



Publication Year	2015
Acceptance in OA@INAF	2020-03-13T22:52:12Z
Title	Lower mass normalization of the stellar initial mass function for dense massive early-type galaxies at $z \sim 1.4$
Authors	GARGIULO, ADRIANA; SARACCO, Paolo; LONGHETTI, Marcella; Tamburri, S.; Lonoce, I.; et al.
DOI	10.1051/0004-6361/201424235
Handle	http://hdl.handle.net/20.500.12386/23229
Journal	ASTRONOMY & ASTROPHYSICS
Number	573

Lower mass normalization of the stellar initial mass function for dense massive early-type galaxies at $z \sim 1.4$

A. Gargiulo¹, P. Saracco¹, M. Longhetti¹, S. Tamburri^{1,2}, I. Lonoce^{1,2}, and F. Ciocca^{1,2}

¹ INAF – Osservatorio Astronomico di Brera, via Brera 28, 20121 Milan, Italy
e-mail: adriana.gargiulo@brera.inaf.it

² Dipartimento di Scienza e Alta Tecnologia, Università degli Studi dell'Insubria, via Valleggio 11, 22100 Como, Italy

Received 19 May 2014 / Accepted 16 October 2014

ABSTRACT

Aims. This paper aims at understanding whether the normalization of the stellar initial mass function (IMF) of massive galaxies varies with cosmic time and/or with mean stellar mass density $\Sigma = M_*/(2\pi R_e^2)$.

Methods. We have tackled this question by taking advantage of a spectroscopic sample of 18 dense ($\Sigma > 2500 M_\odot \text{pc}^{-2}$) massive early-type galaxies (ETGs) that we collected at $1.2 \lesssim z \lesssim 1.6$. Each galaxy in the sample was selected in order to have available: i) a high-resolution deep HST-F160W image to visually classify it as an ETG; ii) an accurate velocity dispersion estimate; iii) stellar mass derived through the fit of multiband photometry; and iv) structural parameters (i.e. effective radius R_e and Sersic index n) derived in the F160W-band. We have constrained the mass-normalization of the IMF of dense high- z ETGs by comparing the true stellar masses of the ETGs in the sample (M_{true}) derived through virial theorem, hence IMF independent, with those inferred through the fit of the photometry which assume a reference IMF (M_{ref}). Adopting the virial estimator as proxy of the true stellar mass, we have implicitly assumed that these systems have zero dark matter. However, recent dynamical analysis of massive local ETGs have shown that the dark matter fraction within R_e in dense ETGs is negligible ($<5\text{--}10\%$) and simulations of dissipationless mergers of spheroidal galaxies have shown that this fraction decreases going back with time. Accurate dynamical models of local ETGs performed by the ATLAS^{3D} team have shown that the virial estimator is prone to underestimating or overestimating the total masses. We have considered this, and based on the results of ATLAS^{3D} we have shown that for dense ETGs the mean value of total masses derived through the virial estimator with a non-homologous virial coefficient and Sersic- R_e are perfectly in agreement with the mean value of those derived through more sophisticated dynamical models, although, of course, the estimates show higher uncertainties.

Results. Tracing the variation of the parameter $\Gamma = M_{\text{true}}/M_{\text{ref}}$ with velocity dispersion σ_e , we have found that, on average, dense ETGs at $\langle z \rangle = 1.4$ follow the same IMF- σ_e trend of typical local ETGs, but with a lower mass-normalization. The observed lower normalization could be evidence of i) an evolution of the IMF with time or ii) a correlation with Σ . To discriminate between the two possibilities, we have compared the IMF- σ_e trend that we have found for high- z dense ETGs with that of local ETGs with similar mean stellar mass density and velocity dispersion and we have found that the IMF of massive dense ETGs does not depend on redshift. The similarity between the IMF- σ_e trends observed both in dense high- z and low- z ETGs over 9 Gyr of evolution and their lower mass-normalization with respect to the mean value of local ETGs suggests that, independently of formation redshift, the physical conditions which characterized the formation of a dense spheroid *on average* lead to a mass spectrum of newly formed stars with a higher ratio of high- to low-mass stars with respect to the IMF of normal local ETGs. In the direction of our findings, recent hydrodynamical simulations show that the higher star-formation rate that should have characterized the early stage of star formation of dense ETGs is expected to inhibit the formation of low-mass stars. Hence, compact ETGs should have higher ratio of high- to low-mass stars than normal spheroids, as we observe.

Key words. galaxies: elliptical and lenticular, cD – galaxies: evolution – galaxies: formation – galaxies: high-redshift – galaxies: stellar content

1. Introduction

The stellar initial mass function (IMF) is the mass distribution of a stellar generation at birth; fixing the relative number of low- and high-mass stars formed in a burst of star formation, it determines the properties of the whole stellar system, i.e. its spectral energy distribution (SED), stellar mass, stellar mass-to-light ratio (M/L), chemical enrichment, and how these properties change with time. In fact, the most widely adopted approaches to constrain the assembly history of galaxies pass through the analysis of their stellar populations properties (e.g. stellar mass, age, star-formation histories) over the cosmic time. The properties of these stellar populations are derived, at different redshifts, by comparing observations with stellar population synthesis (SPS) models which strictly depend on the IMF. Consequently, the knowledge of the IMF is a key ingredient for our understanding

of the stellar mass accretion over the cosmic time. Nowadays, the local direct star counts have shown that the IMF is approximately constant throughout the disk of the Milky Way (MW) and can be described by two declining power laws $dN/dm \propto m^{-s}$, one with slope $s \approx 2.35$ for stars with mass $m \gtrsim 1 M_\odot$, as originally suggested by Salpeter (1955), and the other, flatter, for lower masses (e.g. Kroupa 2001; Chabrier 2003). This form, measured in the MW disk, is then assumed to be universal, i.e. invariant both throughout the wide population of galaxies, and across the cosmic time. Nonetheless, theoretical arguments cast strong doubts on IMF universality (see Bastian et al. 2010; Kroupa et al. 2013, for complete reviews).

The Jeans mass M_J is dependent on the temperature T and the density ρ of the gas ($M_J \propto T^{3/2} \rho^{-1/2}$) (e.g. Larson 1998, 2005; Bonnell et al. 2006). The temperature of a gas is related to its metallicity Z : at fixed heating rate, the gas cooling efficiency

is reduced in absence of metals, and consequently the gravitational collapse of low-mass stars is inhibited in metal poor gas (e.g. Larson 2005; Bate & Bonnell 2005; Bate 2005; Bonnell et al. 2006). In addition to this, there is evidence that the turbulence level of the interstellar medium plays a major role in shaping the IMF (e.g. Padoan & Nordlund 2002; Hennebelle & Chabrier 2008; Hopkins 2013). These theoretical considerations show that the mass spectrum of the stars formed in a burst of star formation is deeply connected with the physical properties of the gas cloud, and hence raise doubts about the validity of the current IMF universality assumption.

From an observational point of view, the IMF universality has recently been contrasted by indirect emerging evidence in external galaxies. The study of spectral features known to be sensitive to the presence of low-mass stars (e.g. the Na I doublet, the Ca II triplet) in massive local early-type galaxies (ETGs) have shown that the strength of these gravity-sensitive features varies with ETG velocity dispersion (σ) and the Mg/Fe ratio in the direction of progressively bottom-heavy IMF, i.e. IMF with a lower ratio of high- to low-mass stars than the MW IMF, with the increase of the two parameters (e.g. Conroy & van Dokkum 2012; Spiniello et al. 2012; Ferreras et al. 2013; La Barbera et al. 2013).

A similar IMF- σ trend has been found through a different (dynamical) approach based on the direct comparison of the *true* stellar mass-to-light ratio ($M/L_{\star,\text{true}}$) with that inferred through, e.g. the fit of the spectral energy distribution adopting SPS models which assume a fixed IMF as reference ($M/L_{\star,\text{ref}}$). The $M/L_{\star,\text{true}}$ has been derived by different authors through, e.g. detailed 2D axisymmetric dynamical models (Cappellari et al. 2012), simpler spherical dynamical models (Tortora et al. 2013), or through a joint lensing and dynamical analysis of gravitational lenses (Treu et al. 2010). The two independent approaches have been shown to achieve a similar trend with velocity dispersion (Tortora et al. 2013; Conroy et al. 2013). This consolidates the robustness of the IMF- σ trend in local massive ETGs, although caveats and limits of both methods need to be fully understood yet (Smith 2014).

Nonetheless, the most relevant aspect for galaxy evolution, i.e. the assumption of a universality across the cosmic time, is still largely observationally unexplored. Theoretically, the higher temperature of intergalactic medium, which is known to increase with redshift as $\approx 2.73(1+z)$ [K], should contrast the formation of low-mass stars at high redshift and hence point toward a top-heavier IMF in early epoch with respect to that of the local Universe (e.g. Larson 1998, 2005). Similarly, the formation of low-mass stars should be inhibited by the absence of metals in the primordial gas (e.g. Larson 2005), as well as by the higher star-formation rate per unit volume observed in the early universe (e.g. Larson 2005; Dabringhausen et al. 2010; Papadopoulos 2010; Papadopoulos et al. 2011; Narayanan & Davé 2012).

Actually, in a recent analysis, Shetty & Cappellari (2014) have provided the first attempt to constrain the IMF in massive ($M_{\star} > 10^{11} M_{\odot}$) ETGs at $z \sim 0.8$ using axisymmetric dynamical models. They find that, on average, the IMF of $z \sim 0.8$ ETGs is similar to those of local ones with similar stellar mass. This result is consistent with the evidences that, on average, the stellar content of ETGs passively evolves from $z \sim 0.8-1$ down to zero (e.g. di Serego Alighieri et al. 2005; van der Wel et al. 2005; Jørgensen & Chiboucas 2013; Choi et al. 2014). In fact, if a variation of IMF in massive ETGs occurs with time, it will be mostly detectable through the analysis of the stellar content of ETGs at $z > 1$. Actually, the number density of spheroidal

massive galaxies increases by a factor ~ 10 in the range $1 < z < 2$ and is almost flat at lower redshift (Ilbert et al. 2013). This, coupled with the many evidences of passive evolution of the stellar content of massive ETGs in the last 7–8 Gyr, implies that, on average, the mean properties of the stellar content of massive ETGs observed at $z < 1$ are not dissimilar to those observed in local ETGs. In contrast, the population of ETGs $z \sim 1.5$ provides us with information only on those spheroidal galaxies that assembled their stellar content in the first 3–4 Gyr of the age of the Universe. Consequently, they retain information on the mass spectrum of stars in the early Universe.

A piece of evidence concerning the stellar content of spheroidal galaxies up to $z \sim 1.5$, is that at fixed stellar mass M_{\star} , ETGs show a remarkable spread in mean stellar mass density ($\Sigma = M_{\star}/2\pi R_c^2$). Although conclusive results have not been achieved yet, a few recent works (e.g. Läsker et al. 2013; Smith & Lucey 2013) and theoretical expectations have suggested that the hotter dynamical status of a system could have a fundamental role in shaping the IMF (e.g. Dabringhausen et al. 2008, 2010; Papadopoulos 2010; Dabringhausen et al. 2012; Marks et al. 2012).

In this paper, taking advantage of a sample of ~ 18 dense ($\Sigma > 2500 M_{\odot} \text{pc}^{-2}$) ETGs we have collected at $1.26 < z < 1.6$ (see Sect. 2 for more details), we address the following questions: i) Does dense high- z ETGs follow the same IMF- σ trend observed for local ETGs? ii) Does the IMF of ETGs depend on mean stellar mass density and/or on redshift? In Sect. 3 we derive the upper limit of the IMF- σ trend of dense ETGs at $\langle z \rangle = 1.4$. We infer the trend in dynamical masses derived via the virial estimator. Although this procedure is less accurate than the more sophisticated 2D axisymmetric model, in Sect. 3 we discuss in detail how the source of uncertainties could affect our conclusions. In Sect. 4 we present our discussion, and we summarize our results in Sect. 5.

Throughout the paper we adopt standard cosmology with $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_m = 0.3$ and $\Omega_{\Lambda} = 0.7$.

2. The sample of dense ETGs at $\langle z \rangle = 1.4$

The dataset we have adopted in this paper is a collection from the literature of all the 18 galaxies morphologically confirmed as ETGs at $1.26 < z < 1.6$, having stellar mass density $\Sigma > 2500 M_{\odot} \text{pc}^{-2}$, and both kinematic (i.e. velocity dispersion σ_e) and structural (i.e. effective radius R_e and stellar mass M_{\star}) properties available. For all the galaxies in the sample, the morphological classification was performed on the basis of a visual inspection of the galaxies carried out on the HST-NIC2 and WFC3 images in the *F160W* filter. We classified as ETGs those galaxies having a regular shape with no sign of disk and no irregular or structured residuals resulting from the profile fitting. We selected only those ETGs with structural parameters derived by fitting the HST images with a point spread function-convolved 2D Sersic profile and with stellar masses derived with the Chabrier IMF and Bruzual & Charlot (2003) models. We consider that stellar masses have a typical error of $\sim 20\%$, which take into account the uncertainties related to models and photometry, but not on IMF, that for the purpose of our study we consider fixed. In Table 1 we list all the properties of the 18 galaxies relevant for this analysis, i.e. the effective radii, the filter in which they were derived, the stellar masses, as well as their velocity dispersions within the effective radius. We rescaled the measured velocity dispersions to R_e aperture following Cappellari et al. (2006). For 14 ETGs in the sample the values of both structural and kinematical parameters have already been published and we refer the reader

Table 1. Total sample of $\langle z \rangle = 1.4$ galaxies.

Object	z	R_e (kpc)	n	Camera-Filter	σ (km s ⁻¹)	σ_e (km s ⁻¹)	$\log M_\star$ (M_\odot)
S2F1-511 ¹	1.267	2.09 ± 0.07	3.3	HST/NIC2- <i>F160W</i>	269	281 ± 23	11.05
S2F1-633 ¹	1.297	2.59 ± 0.11	4.1	HST/NIC2- <i>F160W</i>	434	447 ± 27	11.48
S2F1-527 ¹	1.331	1.57 ± 0.24	3.1	HST/NIC2- <i>F160W</i>	226	240 ± 26	11.14
S2F1-389 ¹	1.406	2.10 ± 0.18	4.5	HST/NIC2- <i>F160W</i>	224	234 ± 45	11.14
2239 ²	1.415	2.16 ± 0.43^3	2.2 ³	HST/ACS- <i>F850LP</i>	111 ± 35	116 ± 37	10.64 ³
2470 ²	1.415	1.81 ± 0.36^3	4.2 ³	HST/ACS- <i>F850LP</i>	141 ± 26	150 ± 28	10.94 ³
A17300 ⁴	1.423	2.9 ± 0.51^5	5.3	HST/WFC3- <i>F160W</i>	265 ± 7	272 ± 7	11.26
A21129 ⁴	1.583	1.5 ± 0.27^5	5.0	HST/WFC3- <i>F160W</i>	260 ± 9	278 ± 10	11.18
C21434 ⁴	1.522	1.9 ± 0.33^5	3.1	HST/WFC3- <i>F160W</i>	218 ± 16	230 ± 17	11.20
C20866 ⁴	1.522	2.4 ± 0.42^5	3.0	HST/WFC3- <i>F160W</i>	272 ± 23	282 ± 23	11.30
32915 ⁶	1.261	1.33 ± 0.13	6.3	HST/WFC3- <i>F160W</i>		271 ± 17	10.88
2341 ⁶	1.266	1.21 ± 0.12	3.8	HST/WFC3- <i>F160W</i>		196 ± 27	10.87
29059 ⁶	1.278	1.62 ± 0.16	4.3	HST/WFC3- <i>F160W</i>		211 ± 16	10.90
2337 ⁶	1.327	1.54 ± 0.15	3.5	HST/WFC3- <i>F160W</i>		284 ± 20	11.04
14758 ⁶	1.331	0.83 ± 0.08	2.2	HST/WFC3- <i>F160W</i>		165 ± 16	10.71
5020 ⁶	1.415	2.07 ± 0.20	4.6	HST/WFC3- <i>F160W</i>		180 ± 54	10.83
13880 ⁶	1.432	0.87 ± 0.08	2.6	HST/WFC3- <i>F160W</i>		179 ± 70	10.64
S2F1-142 ⁷	1.386	3.04 ± 0.12	3.5	HST/NIC2- <i>F160W</i>	340^{120}_{-60}	347^{120}_{-60}	11.54

Notes. Column 1: ID, Column 2: redshift, Column 3: effective radius, Column 4: Sersic index, Column 5: filter adopted to derive the surface brightness parameters, Column 6: measured velocity dispersion, Column 7: velocity dispersion corrected to R_e aperture following Cappellari et al. (2006), Column 8: Stellar mass. References for the velocity dispersion values are listed near the ID value. When a data of the galaxy comes from a paper different from the one where the σ is published, the new reference is specified near its value. ¹Gargiulo et al. (2014); ²Cappellari et al. (2009); ³Cimatti et al. (2008); ⁴Bezanson et al. (2013); ⁵van de Sande et al. (2013); ⁶Belli et al. (2014); ⁷Longhetti et al. (2014).

to original papers for all the details (see references in Table 1), while for the remaining four (S2F1-511, S2F1-633, S2F1-527, S2F1-389) the estimates of velocity dispersions and the updated values of structural parameters will be presented in a forthcoming paper (Gargiulo et al. 2014). In the following paragraph we briefly describe the method used to derive the physical parameters of these four galaxies (hereafter our high- z sample galaxies).

2.1. Our high- z sample galaxies

We derived the velocity dispersions of the four galaxies at $z \sim 1.3$ from spectra obtained with the ESO VLT-FORS2 spectrograph. The OG590 filter and the GRIS-600 z grism we chose for the observations, allowed us to cover the wavelength range $0.7 < \lambda < 1.0 \mu\text{m}$ with a sampling of 1.6 \AA per pixel. We fixed the width of the slit to $1''$, which provides a resolution $R \simeq 1400$, corresponding to a $FWHM \simeq 6.5 \text{ \AA}$ at 9000 \AA . The total effective integration time per galaxy is ~ 8.5 h which results in a signal-to-noise per pixel $\sim [8-10]$. After standard reduction and calibration, we derived the velocity dispersions best-fitting the observed spectra with a set of template stars from Munari et al. (2005) using the software pPXF (Cappellari & Emsellem 2004). With an extensive set of simulations we tested the reliability of our estimates, given the wavelength region covered and the signal-to-noise of the spectra and found that fitting the observed spectra with a Gauss+Hermite polynomial over the whole range of λ , with a penalty BIAS $\simeq 0.1$ provides reliable estimates of σ . The effective radii were derived by Longhetti et al. (2007) fitting the HST-NIC2/*F160W* images with a 2D-psf convolved Sersic profile with GALFIT (Peng et al. 2002). We updated the values they found taking in consideration the more robust estimate of the redshift based on the FORS2 spectra. Finally, stellar masses of the four ETGs were already presented in Saracco et al. (2009). However, following the same procedure adopted by the authors,

we re-estimated the stellar masses to take into account the update value of the redshift and the now available near-IR photometry in the WISE *RSR-W1* filter ($\lambda_{\text{eff}} \simeq 3.4 \mu\text{m}$).

For 16 ETGs in the total sample their effective radii were derived fitting with a Sersic profile the *WFC3/F160W*-band images while for two of them structural parameters were derived on ACS/F850LP-band images. Since high- z ETGs have significant colour gradients (e.g. Gargiulo et al. 2011, 2012; Guo et al. 2011; Cassata et al. 2011), in order not to affect the relative distribution of the effective radii of our galaxies (hence of their virial masses, see Sect. 3.2), we converted the radii of these two galaxies to those in the filter *F160W* scaling the fitted values of a factor ~ 0.8 (Cassata et al. 2011; Gargiulo et al. 2012). At redshift $\langle z \rangle = 1.4$ the HST/*F160W* filter samples approximately the *R* band restframe. For all the galaxies in the sample, the stellar masses M_\star were derived fitting their observed SEDs with the SPS models by Bruzual & Charlot (2003) and the Chabrier IMF. Figure 1 shows the distribution of our sample of ETGs at $\langle z \rangle = 1.4$ in the plane defined by effective radius and stellar mass. As a reference, we plot also a complete sample of ETGs in the redshift range $1.2 < z < 1.6$ (Tamburri et al. 2014). The sample was extracted from the GOODS-MUSIC catalogue (Grazian et al. 2006; Santini et al. 2009) selecting all the galaxies (1302) with $K_{s,AB} \leq 22$. The spectroscopic redshift completeness is ~ 76 per cent. The morphological classification was performed by two authors, independently, and following the same criteria adopted for our spectroscopic sample. This results in a final sample of ~ 196 ETGs in the redshift range $0.6 \leq z \leq 2.5$ from which we selected those at $1.26 < z < 1.6$ shown in Fig. 1. In this redshift range the photometric sample is complete down to $M_\star \simeq 2 \times 10^{10} M_\odot$. By comparing the spectroscopic sample at $\langle z \rangle \sim 1.4$ with the photometrically complete one, it is evident that the first is skewed towards denser systems. This is principally because in almost all the cases the spectroscopic

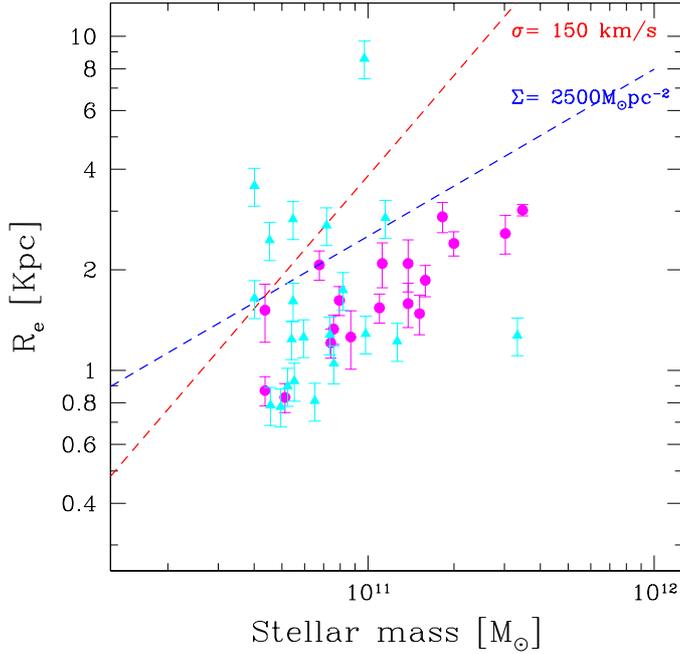


Fig. 1. Distribution of the ETGs of the total high- z sample in the $R_e - M_*$ plane (magenta points). Cyan triangles are a complete sample of ETGs selected from the GOODS-South field in the redshift range $1.2 < z < 1.6$ (Tamburri et al. 2014) for reference. Blue(red) lines indicate lines of constant $\Sigma(\sigma)$. We note that the sample of ETGs at $\langle z \rangle = 1.4$ with velocity dispersions available is skewed toward more compact systems.

targets were identified through a colour selection criteria which is shown to preferentially select compact ETGs (e.g. Saracco et al. 2009; Szomoru et al. 2011). We note that, in comparison with the spectroscopic sample, the complete photometric sample does not include galaxies with $M_* \gtrsim 10^{11} M_\odot$. Actually, the number density of such massive ETGs at $z \sim 1.4$ is extremely low ($\sim 10^{-5} \text{ Mpc}^{-3}$, Ilbert et al. 2010). Considering the relatively small area of GOODS-South field ($\sim 140 \text{ arcmin}^2$), from which the photometric sample was taken, we expect to find ~ 1.5 galaxies, as we observe. Instead, our high- z spectroscopic sample is a collection of spheroids usually identified as spectroscopic targets among the most massive and red galaxies in all the deep currently available surveys. This increases, with respect to photometric sample, their occurrence in the range of higher stellar masses.

3. IMF- σ trend as a function of time and mean stellar mass density

3.1. IMF- σ trend in local universe

Local studies have shown that IMF in spheroidal galaxies varies with their velocity dispersions. These studies compare the stellar mass-to-light ratio $M/L_{\star,\text{ref}}$ derived through, e.g. the fit of the spectral energy distribution assuming a fixed reference IMF, with the *true* stellar mass-to-light ratio $M/L_{\star,\text{true}}$, derived through either dynamical models or spectral features analysis, and investigate how their ratio $\Gamma = (M/L)_{\star,\text{true}}/(M/L)_{\star,\text{ref}}$ varies with the physical properties of local ETGs.

Usually, this IMF- σ trend is presented in a form similar to

$$\log \Gamma = \log[(M/L)_{\star,\text{true}}/(M/L)_{\star,\text{ref}}] = \alpha \log \sigma_e + \beta. \quad (1)$$

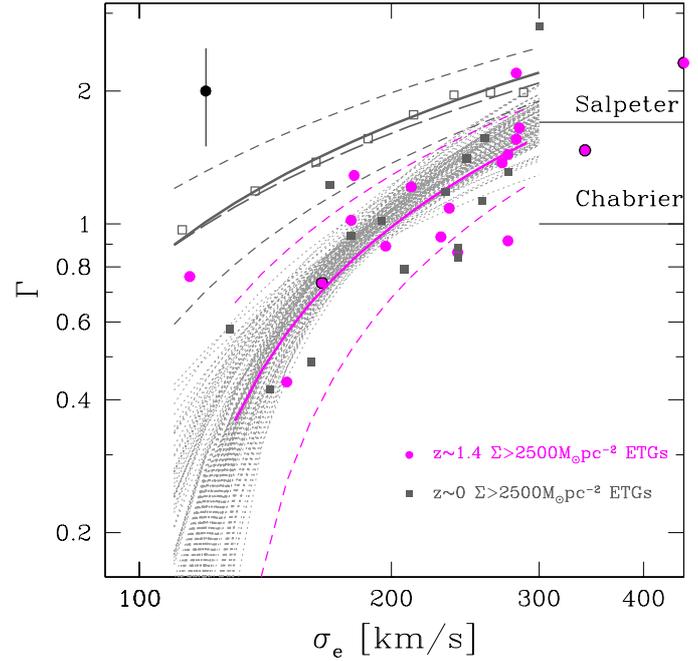


Fig. 2. Trend of Γ parameter (see Eq. (1)) as a function of velocity dispersion, for normal (mean stellar mass density $\Sigma < 2500 M_\odot \text{ pc}^{-2}$) local ETGs (dark grey solid line) and for the whole local ETG population (long-dashed dark grey line). Dashed grey lines are the scatter around the best-fit relation of normal local ETGs, while grey open squares are the average values of the Γ s of local normal ETGs in bin of $\sim 50 \text{ km s}^{-1}$. Magenta points are the upper limits of Γ we have derived for high- z ETGs in our spectroscopic sample at $\langle z \rangle = 1.4$, and magenta solid line is their best-fit relation (magenta dashed lines indicate the rms around the relation). The black small-dotted lines are the best-fit relations of 100 subsamples of local ETGs, randomly extracted from the local sample by Tortora et al. (2013), in order to have the same Σ and σ_e distribution of our spectroscopic sample at $\langle z \rangle = 1.4$ (filled dark grey squares are an example of one local subsample). In this procedure we excluded three galaxies of the high- z sample (identified with magenta points with a black contour) since the local sample is missing ETGs with similar Σ and σ_e . The black point indicates the typical error on Γ s of our spectroscopic sample at $\langle z \rangle = 1.4$.

For the purpose of our analysis, as an example of $\Gamma - \sigma_e$ trend in local ETGs, we have derived from the dataset of Tortora et al. (2013) the best-fit relation of Γ with σ_e for normal ($\Sigma \leq 2500 M_\odot \text{ pc}^{-2}$) local ETGs with stellar mass and velocity dispersion comparable to that of our high- z sample (i.e. $\text{Log} M_* > 10.64$ and $\sigma_e \gtrsim 110 \text{ km s}^{-1}$; we have already stated in the Introduction that all the trends found in local Universe are similar). Specifically, Tortora et al. (2013) derived, assuming different dark matter (DM) profiles, the Γ values for ~ 4500 giant ETGs drawn from the SPIDER project (La Barbera et al. 2010) in the redshift range of $z = 0.05 - 0.1$. The SPIDER sample is 95% complete at a stellar mass $M_* = 3 \times 10^{10} M_\odot$, which corresponds to $\sigma \sim 160 \text{ km s}^{-1}$. In Fig. 2 we have reported the trend found assuming in the derivation of Γ a Navarro, Frenk and White profile (NFW, Navarro et al. 1996) for the DM component. Clearly, the effect of reducing the DM fraction is to increase the overall stellar mass normalization (for a clear picture of the effect of the degree of contraction on mass normalization see Dutton et al. 2011). In this figure, ETGs with $\Gamma = 1$ have true stellar masses in agreement with those derived with the reference IMF, in this case the one by Chabrier, while ETGs with $\Gamma > 1$ have stellar masses greater than that inferred with the Chabrier IMF, i.e. their IMF normalization is more massive.

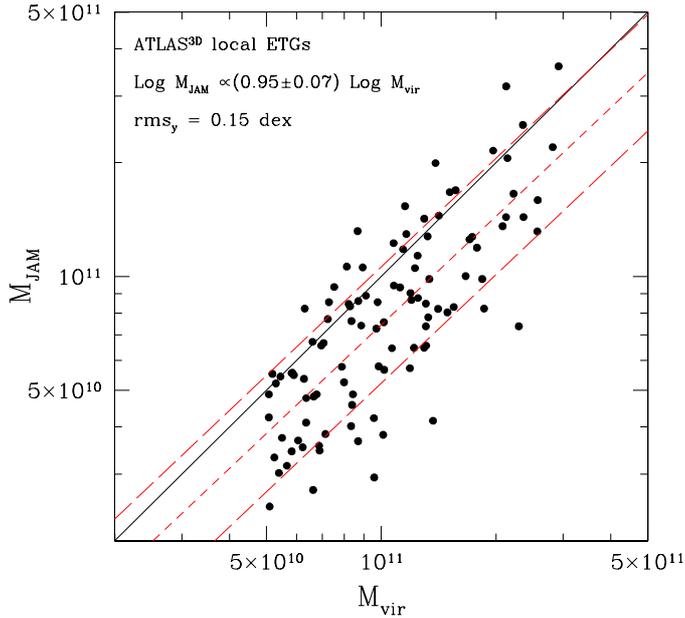


Fig. 3. Dynamical mass derived with 2D axisymmetric models, M_{JAM} , versus virial mass, M_{vir} , derived as in Eq. (2) for the ETGs of the local ATLAS^{3D} sample. M_{JAM} was derived as stated in Cappellari et al. (2013b) in the caption of their Table 1. Although, on average, Eq. (2) overestimates the dynamical masses (black solid line is the 1:1 correlation), the offset (~ 0.2 dex) is not mass dependent.

3.2. IMF- σ trend for dense ETGs at $\langle z \rangle = 1.4$

Following the same procedure adopted in the local Universe we have investigated the IMF- σ_e trend for the dense ETGs in our sample at $\langle z \rangle = 1.4$.

After defining M_{tot} the total mass of the galaxy, and $M_{\star, \text{ref}}$ the stellar mass given in Table 1, we derived the upper limit of the Γ parameter for the ETGs in our sample as $\Gamma = M_{\text{tot}}/M_{\star, \text{ref}}$. Assuming the total mass to be a proxy of the true stellar, we have implicitly assumed that these systems have zero DM.

An important step in our analysis consists in deriving reliable *total* masses for ETGs in our sample. We addressed this point adopting the virial mass estimator, so for each galaxy in the sample

$$M_{\text{tot}} = M_{\text{vir}} = (\beta R_e \sigma_e^2)/G, \quad (2)$$

where G is the gravitational constant and β is equal to

$$\beta = 8.87 - 0.831 \times n + 0.0241 \times n^2 \quad (3)$$

following Cappellari et al. (2006).

Typical errors on our Γ estimates are ~ 0.5 . Detailed estimates of dynamical masses based on more sophisticated 2D axisymmetric dynamical models of local ETGs have shown that the virial Eq. (2) could overestimate the total masses (Cappellari et al. 2013b). In addition to the uncertainties on the *true* value of virial mass for each ETG, to which we will return later, we point out that the *relative* values of M_{vir} are not expected to change. In Fig. 3, the total dynamical mass of local ATLAS^{3D} ETGs (Cappellari et al. 2011) derived with accurate 2D axisymmetric models (M_{JAM}) is given versus their virial mass M_{vir} derived as in Eq. (2). The two mass estimates were derived from the ATLAS^{3D} database as follows. The M_{JAM} masses, following Cappellari et al. (2013b) (see caption of their Table 1), were derived as $(M/L)_{\text{JAM}} \times L$, where $(M/L)_{\text{JAM}}$ and L are the mass-to-light ratio obtained via dynamical models and the galaxy total luminosity, found in Cols. 6 and 15 of their Table 1, respectively. For the estimate of M_{vir} for ATLAS^{3D} ETGs, we adopted

for R_e and n those derived from the 1D Sersic fit of the observed light profiles of the galaxy ($R_{e, \text{1D}}$ and n_{1D} , Cols. 3 and 4 of the Table C.1 in Krajnović et al. 2013). Although the procedure they adopted to derive the R_e is different from the 2D approach of GALFIT, the authors show that, at least for the Sersic index, the two procedures return similar results, with a difference of just 0.08 between the median values of the two distributions. For what concerns the σ , we referred to the stellar velocity dispersion measured by co-adding all SAURON spectra contained within the physical region of 1 kpc (Col. 4 of Table 1 in Cappellari et al. 2013b). We corrected this measure to those within $R_{e, \text{1D}}$ following Cappellari et al. (2006). In Fig. 3 we plot the ATLAS^{3D} ETGs with $10.7 \lesssim \text{Log} M_{\text{vir}} \lesssim 11.5$ in order to approximately match the mass range of our spectroscopic sample. In this mass range, on average, we observe that Eq. (2) overestimates the dynamical masses, but the offset (~ 0.2 dex) is not mass dependent. The values of virial mass estimated via Eq. (2), therefore, are *relatively* reliable and hence, so are the Γ parameters of our high- z ETGs. Based on these considerations, and looking at the distribution of high- z dense ETGs in the Γ - σ_e plane, the first result we have found is that the Γ of high redshift dense ETGs show a *trend* with their σ_e (Spearman’s rank coefficient 0.84, probability 6.9×10^{-5}) and that this trend is very similar to that observed in the local Universe, that is, their IMF becomes more massive with increasing velocity dispersion. In particular, in the range of velocity dispersion $\sim [120-300]$ km s⁻¹ we have found $\Gamma = 2.99 \text{Log} \sigma_e - 5.22$ with a scatter of 0.28.

The possible uncertainties on the true value of M_{vir} do not affect this result, since the *trend* is dependent on relative values. Thus, within the approximation of our analysis, the trend of IMF with σ_e is also observed in high- z ETGs.

The second aspect we want to investigate is whether the IMF of ETGs depends on redshift and/or on mean stellar mass density. To address this second question we have compared the *normalization* of the IMF- σ_e trend we have found for dense ETGs at $z \sim 1.4$ with that observed in local universe.

Although we are making the extreme assumption of no dark matter, and thus the Γ we have derived for high- z ETGs are upper limits, the normalization of the IMF- σ_e trend for dense high- z ETGs is at 1σ lower than that observed for normal local ETGs. In particular, on average, at fixed velocity dispersion, the Γ s of high- z dense ETGs are a factor of ~ 2 lower than the average value observed for normal local ETGs¹. This could be ascribed to an evolution of IMF with redshift or to a dependence of IMF on mean stellar mass density. To discriminate between the two possibilities, we have compared the IMF- σ_e trend of high- z dense ETGs with that of equally dense local ETGs. In fact, if the IMF of massive dense ETGs depends on redshift, we should note that the average IMF- σ_e trend of local dense ETGs differ from the one we observe at high- z . Specifically, to address the topic, we randomly extracted from the sample of Tortora et al. (2013), 100 subsamples of local ETGs with $\text{Log}[M_{\star}/M_{\odot}] > 10.64$ (as the high- z sample) and the same distribution in Σ and σ_e of our

¹ Since the Γ of a given galaxy increases for decreasing amount of dark matter in the model, it could be more intuitive to expect that the IMF- σ_e trend we have derived had a higher mass-normalization than that of typical local ETGs. However, by definition, the Γ of a galaxy is $\Gamma = M_{\text{true}}/M_{\text{ref}} = \frac{R_e \beta (1-f_{\text{DM}}) \sigma_e^2}{GM_{\text{ref}}}$ where f_{DM} is the galaxy dark matter fraction. Thus, if we consider two different galaxies with the same velocity dispersion σ_e , the values of their Γ s are not simply related to their dark matter fractions, but also to their M_{ref}/R_e ratio, hence to their mean stellar mass density. Thus, the position occupied by a galaxy in the IMF- σ plane does not only depend on its f_{DM} , but also on its R_e/M_{\star} ratio.

spectroscopic high- z sample. We checked the goodness of the local subsamples in reproducing the high- z sample with a KS test. In the random extraction, we retained a local subsample if the probability that its distribution of Σ and σ_e and that of the high- z sample are extracted from the same parent distribution is higher than 60%. In this procedure we excluded three galaxies of the high- z sample since the local sample is missing ETGs with similar Σ and σ_e . For each of these 100 subsamples we derived the best fitting Γ - σ relation in the range $\sim[120\text{--}300]$ km s $^{-1}$. The resulting relations are shown in Fig. 2. Actually, the comparison shows that the IMF trend of dense high- z ETGs is consistent with that of local ETGs with similar velocity dispersion and mean stellar mass density.

However, we notice that our estimates of Γ are based on two assumptions, namely that on the DM and on the Eq. (2) as a reliable estimator of total mass. Although these assumptions do not affect the relative values of total masses (as we have shown above), they could have a significant impact on their *absolute* values, and hence on the normalization of the IMF- σ_e trend we observe. With regard to our assumption of zero dark matter fraction, we note that dense ETGs are expected to be stellar matter dominated (e.g. Conroy et al. 2013), and, on average, the dark matter fraction within the effective radius in massive dense local ETGs is found to be $<10\%$ (Tortora, priv. comm.). Similar results are derived also from ATLAS 3D survey. Specifically, we checked the dark matter fraction f_{DM} of ATLAS 3D ETGs (taken from Cappellari et al. 2013a) with $10.7 \lesssim \text{Log}M_{vir} \lesssim 11.5$ and $\Sigma > 2500 M_{\odot} \text{pc}^{-2}$ (for a detailed description of Σ estimates for ATLAS 3D ETGs see below) and found that $f_{DM} = 0.11 \pm 0.11$. Moreover, simulations of dissipationless mergers of spheroidal galaxies have shown that this fraction can only decrease going back with time, pointing toward an even lower fraction of dark matter in high- z spheroidal galaxies (Hilz et al. 2013).

However, the values of Γ we have found for high- z systems could be affected by our uncertainties on true values of M_{tot} , hence of M_{vir} . In the following using the results obtained through 2D dynamical models of local ATLAS 3D ETGs we show that Eq. (2) provides reliable absolute estimates of total masses for dense ETGs at $z \sim 1.4$. Cappellari et al. (2013b) in their Figs. 13 and 14, by comparing the estimates of total mass of local ETGs derived through accurate 2D dynamical models (M_{JAM}) with those inferred from Eq. (2), clearly showed that the reliability of total masses estimated using the virial theorem is closely related to the choice of virial coefficient β and to the accuracy of R_e estimates, which depends on the quality of the data and on the technique adopted to derive them. Following a similar approach, analogous results have been found by Shetty & Cappellari (2014) for massive galaxies at $z \sim 1$. As far as the robustness of size measurements of high- z ETGs, many works (e.g. Szomoru et al. 2012; Davari et al. 2014) have addressed the topic, showing that the typical approach adopted at high- z , i.e. 2D Sersic-fit of the light profile out to infinity on deep space-based images, is able to provide accurate R_e estimates (uncertainties of $\sim 15\%$). In the case in which radii are derived through this approach (i.e. through a Sersic-fit of the light profile out to infinity), Cappellari et al. (2013b) showed that the best approximation of dynamical masses is recovered assuming in the virial equation the non-homologous β coefficient of Eq. (3) rather than a constant virial coefficient (see the two right panels in their Fig. 14). Actually, in these conditions, M_{vir} shows systematic offset of ~ 0.2 dex with M_{JAM} with a scatter 0.15 dex as we have shown in Fig. 3. To further exploit the origin of this offset we compared for ATLAS 3D ETGs the value of their β coefficient, derived following Eq. (3), with those inferred from their

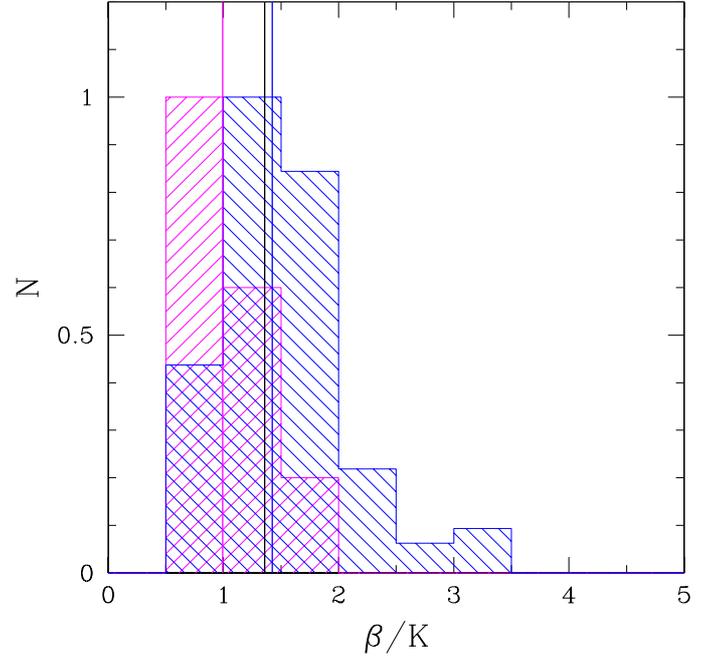


Fig. 4. Distribution of the ratio β/K for ATLAS 3D ETGs with $\Sigma \geq 2500 M_{\odot} \text{pc}^2$ (magenta histogram) and for those with $\Sigma < 2500 M_{\odot} \text{pc}^2$ (blue histogram). Coloured solid lines are the median values of the two relative distributions, while black one refers to the total sample. We have included in the plot only local ETGs with stellar masses in the same range of dense high- z ETGs in our sample ($10.7 \lesssim \text{Log}M_{vir}/M_{\odot} \lesssim 11.5$). The β parameter was derived through Eq. (3), while the K coefficient through Eq. (4).

2D-dynamical models (K) as

$$K = (M_{JAM} * G)/(R_e \sigma_e^2), \quad (4)$$

where M_{JAM} was derived as described above, R_e is $R_{e,1D}$ (Table C.1 from Krajnović et al. 2013) and σ_e is the velocity dispersion within 1 kpc (Table 1 Cappellari et al. 2013b) corrected to the $R_{e,1D}$ aperture.

In Fig. 4 we plot the distribution of the ratio β/K for ATLAS 3D ETGs with $\Sigma \geq 2500 M_{\odot} \text{pc}^2$ and for those with $\Sigma < 2500 M_{\odot} \text{pc}^2$. To estimate the mean stellar mass density of ATLAS 3D ETGs we adopted for R_e the $R_{e,1D}$ values and derived their stellar masses multiplying the stellar mass-to-light ratio (Table 1 from Cappellari et al. 2013b) by the total luminosity L (Table 1 from Cappellari et al. 2013b). Since the ATLAS 3D stellar mass-to-light ratio refers to the Salpeter IMF, we converted the stellar masses to the Chabrier IMF following Longhetti & Saracco (2009).

We have included in the plot of Fig. 4 only local ETGs with $10.7 \lesssim \text{Log}M_{vir}/M_{\odot} \lesssim 11.5$. Although, on average, for the total sample the β coefficient overestimates the virial masses of a factor ~ 1.35 , this is principally due to normal ETGs (i.e. ETGs with $\Sigma < 2500 M_{\odot} \text{pc}^2$). In fact, a clear trend of the β parameter with the compactness of the system emerges: while the majority of normal ETGs shows a β/K ratio >1 (median value $1.42_{-0.36}^{+0.37}$, where errors indicate the 25th and 75th percentile), for dense ETGs (~ 20 ETGs on a total of 103 ETGs in the figure) the virial masses estimated by means of Eq. (2) are perfectly in agreement with those inferred through more sophisticated dynamical models (median value of $\beta/K = 0.99_{-0.11}^{+0.09}$). This result could be expected given the emerging picture on the variation of ETGs structure in the $R_e - M_{\star}$ plane. Actually, Cappellari et al. (2013a)

in their schematic summary of Fig. 14 clearly showed that in the mass range $3 \times 10^{10} < M_{\star} < 2 \times 10^{11}$ ETGs are dominated by the fast rotators, with very few slow rotators. However, at fixed stellar mass, the relevance of the bulge steadily increases with the mean stellar mass density. We argue, therefore, that the assumption of spherical isotropic models in the derivation of the β coefficient could fit better for dense systems than for normal ETGs, as can be seen in Fig. 4. Clearly, this result could not be necessarily valid for high- z dense ETGs which could be structurally and dynamically different from local ones. However, it has been found that both the kinematical and structural properties of dense high- z ETGs are very similar to those of equally dense local ETGs, and their β parameters should be similar as well. (e.g. Saracco et al. 2014; Trujillo et al. 2014; Ferré-Mateu et al. 2012; van der Wel et al. 2012; Buitrago et al. 2013). In this context, very interestingly, the mean value of the β parameter we have derived for our dense high- z ETGs is 7.0 ± 1.7 , and perfectly agrees with the mean value of massive ATLAS^{3D} dense ETGs, 7.5 ± 1.7 .

To conclude, we have found that dense high- z ETGs i) follow the same IMF- σ trend of local ETGs; ii) have a lower mass normalization with respect to the average IMF- σ trend of normal local ETGs population; and iii) their mass-normalization is consistent with that of local ETGs with similar velocity dispersion and mean stellar mass density. In fact, neither our assumption about the dark matter fraction or the estimate of dynamical masses through virial estimator affect this result. Our finding of a lower mass normalization for dense high- z ETGs is in agreement with the recent result obtained through lens modeling of a $z = 1.62$ compact ($R_e \simeq 1.3$ kpc, and $M_{\star} \simeq 2 \times 10^{11} M_{\odot}$) early-type gravitational lens galaxy (Wong et al. 2014). Although the constraints are broad, the authors have found that the stellar mass of this galaxy is more consistent with that derived with a Chabrier IMF than with a Salpeter IMF.

In term of stellar content, these findings imply that at fixed velocity dispersion and at any redshift, the IMF of dense ETGs has a higher ratio of high- to low-mass stars, with respect to the IMF of typical local ETGs. In particular, assuming that the IMF is modelled as a single power law $dN/dm \propto m^{-s}$ in the range of stellar masses $[0.1-100] M_{\odot}$, we found that its slope varies from $s = \sim 1.5$ for high- z dense ETGs with velocity dispersion ~ 150 km s⁻¹ to $s = 2.35$ (hence the Salpeter) for ETGs with velocity dispersion ~ 300 km s⁻¹.

4. The IMF in massive spheroidal galaxies: Does it depend on when or on how ETGs assembled their stellar mass?

The similarity between the IMF trend of high- z and low- z dense ETGs over 9 Gyr of evolution and their lower normalization with respect to that of a typical local ETG suggest that i) the IMF in massive dense ETGs does not depend on the *epoch* at which they accreted their stellar mass and that ii) the physical conditions responsible for the higher mean stellar mass density of dense ETGs, at any redshift, favour a mass spectrum of new stars with a higher ratio of high- to low-mass stars with respect to that of a normal local ETGs.

Our current knowledge of dense ETGs shows that the compactness of a spheroidal galaxy is closely related to its star-formation rate (SFR). Local studies have shown that the FP residuals of local ETGs are anticorrelated both with galaxy age (e.g. Forbes et al. 1998; Reda et al. 2005; Gargiulo et al. 2009; Graves et al. 2009) and with the α -element abundance ratio,

α/Fe (e.g. Gargiulo et al. 2009; Graves et al. 2009) in trends whereby ETGs more compact than expected from the FP relation have stellar populations systematically older and with higher abundances than average. However, a multiple regression analysis of FP residuals showed that α/Fe is the driving parameter, since there is *no* age correlation with compactness at fixed α/Fe , but a very strong α/Fe correlation at fixed age (e.g. Gargiulo et al. 2009). It has been long realized that the stellar α/Fe ratio is a robust indicator of the timescale τ of the star formation (e.g. Tinsley 1980; Greggio & Renzini 1983; Matteucci 1996), with higher abundance for shorter timescale. Thus, local results show that the compactness of a spheroidal galaxy is not the imprint of *when* it had assembled its stellar mass, but mostly of *how*, and more specifically, of the *duration* of its mass assembly history.

These findings show that the higher the SFR at which a galaxy assembles its stellar mass, the higher its compactness. Recent hydrodynamical simulations have shown that, in a galaxy, star formation at a rate greater than few $M_{\odot} \text{ yr}^{-1}$ leads to an increase of the Jeans mass of the system (Narayanan & Davé 2012), and thus inhibits the formation of low-mass stars. This result agrees with our findings. With respect to normal ETGs, the higher SFR that characterizes the formation of dense ETGs, independently of z , should lead to a lack of low-mass stars with the consequent increase in the ratio of high- to low-mass stars, as we observe.

To summarize, on the basis of our results we have shown that high- z dense ETGs follow the same IMF- σ trend of local dense ETGs. This implies that the IMF of massive dense ETGs do not depend on time. Independently of z , the observed lower mass-normalization is a characteristic of ETGs with higher mean stellar mass density. If this result is extendible to all ETGs, that is, if the mass normalization is linked to the mean stellar mass density of the galaxies, then, at any redshift, the IMF of massive ETGs will have a higher or lower ratio of high- to low-mass stars according to the peculiar physical conditions in which a galaxy assembled its stellar mass.

Starting from these results, coupled with the recent findings on normal and dense ETGs, in the following we discuss the implication that our findings can have on the evolution with time of the *mean* IMF of ETGs population. Recent evidence suggests that the first spheroids formed in the early Universe, were dense (e.g. Saracco et al. 2010; Cassata et al. 2013; van Dokkum & Conroy 2010). At later epochs, compact ETGs continued to appear, together with normal ETGs, at least down to $z \sim 1$. In fact, the first studies of the number density of dense and normal massive quiescent galaxies have shown that dense spheroids were numerically dominant at $z > 1$ while normal ETGs are more common at $z < 1$ (Cassata et al. 2013). If the mass normalization of IMF is related to the mean stellar mass density, with densest systems those with a lower mass-normalization as we have found, the evolution of number density of dense and normal ETGs should imply a trend of the *average* IMF of massive ETGs population with z , in the direction of a higher mass-normalization with decreasing redshift. On this basis, consistent with the many theoretical predictions on the evolution of Jeans mass with redshift (see Introduction), our result implicitly suggests that the mean IMF of ETGs population in the early universe had a higher abundance of high-mass stars than the mean IMF of local ETG (assuming a power-law IMF). Clearly, this proposed picture needs to be further constrained by observations aimed at investigating the IMF in complete sample of massive ETGs at different epochs.

5. Summary and conclusions

We have used a spectroscopic sample of dense ($\Sigma > 2500 M_{\odot} \text{pc}^{-2}$) ETGs we collected at $z = 1.4$ to investigate the variation of the IMF of massive ETGs as a function both of Σ and z . The results of our analysis are the following:

- Massive dense ETGs at $z = 1.4$ follow the same IMF- σ_e trend of typical local ETGs, but with a lower mass-normalization. If massive dense ETGs were the first spheroids to form, as recent findings suggest, this result shows that the *average* normalization of the IMF of the first massive ETGs is lower than that of a Chabrier IMF. This lower normalization could be an evidence of an evolution of the IMF with time or of a correlation with Σ
- We find that the IMF- σ_e trend of dense ETGs at $z = 1.4$ is consistent with those of local ETGs with similar velocity dispersion and mean stellar mass density. This result is an evidence that the IMF, at least for dense ETGs, is not time dependent.

Irrespective of the formation redshift, at fixed velocity dispersion, the physical conditions which characterized the formation of dense ETGs lead to a mass spectrum of newly formed stars with a higher ratio of high- to low-mass stars than the IMF of typical ETGs of similar velocity dispersion. In fact, hydrodynamical simulations have shown that the Jeans Mass increases with the star-formation rate (SFR). Thus, given the higher SFR that should characterize the early phases of star formation of dense ETGs, we suggest that the lower mass-normalization we observe in dense ETGs is the result of a lack of low-mass stars.

On this basis, we suggest that the choice of the proper IMF for an individual massive ETG is not dependent on the time at which it assembled its stellar mass, but mostly on the efficiency of its star formation, hence on Σ .

Acknowledgements. The authors warmly thank the referee for his/her helpful comments and suggestions, and C. Tortora for the scientific discussions and for providing us with local data. This work has received financial support from Prin-INAF 1.05.09.01.05 and is based on observations made with ESO Very Large Telescope under programme ID 085.A-0135A.

References

- Bastian, N., Covey, K. R., & Meyer, M. R. 2010, *ARA&A*, 48, 339
- Bate, M. R. 2005, *MNRAS*, 363, 363
- Bate, M. R., & Bonnell, I. A. 2005, *MNRAS*, 356, 1201
- Belli, S., Newman, A. B., & Ellis, R. S. 2014, *ApJ*, 783, 117
- Bezanson, R., van Dokkum, P., van de Sande, J., Franx, M., & Kriek, M. 2013, *ApJ*, 764, L8
- Bonnell, I. A., Clarke, C. J., & Bate, M. R. 2006, *MNRAS*, 368, 1296
- Bruzual, G., & Charlot, S. 2003, *MNRAS*, 344, 1000
- Buitrago, F., Trujillo, I., Conselice, C. J., & Häußler, B. 2013, *MNRAS*, 428, 1460
- Cappellari, M., & Emsellem, E. 2004, *PASP*, 116, 138
- Cappellari, M., Bacon, R., Bureau, M., et al. 2006, *MNRAS*, 366, 1126
- Cappellari, M., di Serego Alighieri, S., Cimatti, A., et al. 2009, *ApJ*, 704, L34
- Cappellari, M., Emsellem, E., Krajnović, D., et al. 2011, *MNRAS*, 413, 813
- Cappellari, M., McDermid, R. M., Alatalo, K., et al. 2012, *Nature*, 484, 485
- Cappellari, M., McDermid, R. M., Alatalo, K., et al. 2013a, *MNRAS*, 432, 1862
- Cappellari, M., Scott, N., Alatalo, K., et al. 2013b, *MNRAS*, 432, 1709
- Cassata, P., Giavalisco, M., Guo, Y., et al. 2011, *ApJ*, 743, 96
- Cassata, P., Giavalisco, M., Williams, C. C., et al. 2013, *ApJ*, 775, 106
- Chabrier, G. 2003, *PASP*, 115, 763
- Choi, J., Conroy, C., Moustakas, J., et al. 2014, *ApJ*, 792, 95
- Cimatti, A., Cassata, P., Pozzetti, L., et al. 2008, *A&A*, 482, 21
- Conroy, C., & van Dokkum, P. G. 2012, *ApJ*, 760, 71
- Conroy, C., Dutton, A. A., Graves, G. J., Mendel, J. T., & van Dokkum, P. G. 2013, *ApJ*, 776, L26
- Dabringhausen, J., Hilker, M., & Kroupa, P. 2008, *MNRAS*, 386, 864
- Dabringhausen, J., Fellhauer, M., & Kroupa, P. 2010, *MNRAS*, 403, 1054
- Dabringhausen, J., Kroupa, P., Pflamm-Altenburg, J., & Mieske, S. 2012, *ApJ*, 747, 72
- Davari, R., Ho, L. C., Peng, C. Y., & Huang, S. 2014, *ApJ*, 787, 69
- di Serego Alighieri, S., Vernet, J., Cimatti, A., et al. 2005, *A&A*, 442, 125
- Dutton, A. A., Conroy, C., van den Bosch, F. C., et al. 2011, *MNRAS*, 416, 322
- Ferré-Mateu, A., Vazdekis, A., Trujillo, I., et al. 2012, *MNRAS*, 423, 632
- Ferreras, I., La Barbera, F., de la Rosa, I. G., et al. 2013, *MNRAS*, 429, L15
- Forbes, D. A., Ponman, T. J., & Brown, R. J. N. 1998, *ApJ*, 508, L43
- Gargiulo, A., Haines, C. P., Merluzzi, P., et al. 2009, *MNRAS*, 397, 75
- Gargiulo, A., Saracco, P., & Longhetti, M. 2011, *MNRAS*, 412, 1804
- Gargiulo, A., Saracco, P., Longhetti, M., La Barbera, F., & Tamburri, S. 2012, *MNRAS*, 425, 2698
- Gargiulo, A., Saracco, P., Longhetti, M., et al. 2014, *A&A*, submitted
- Graves, G. J., Faber, S. M., & Schiavon, R. P. 2009, *ApJ*, 698, 1590
- Grazian, A., Fontana, A., de Santis, C., et al. 2006, *A&A*, 449, 951
- Greggio, L., & Renzini, A. 1983, *A&A*, 118, 217
- Guo, Y., Giavalisco, M., Cassata, P., et al. 2011, *ApJ*, 735, 18
- Hennebelle, P., & Chabrier, G. 2008, *ApJ*, 684, 395
- Hilz, M., Naab, T., & Ostriker, J. P. 2013, *MNRAS*, 429, 2924
- Hopkins, P. F. 2013, *MNRAS*, 433, 170
- Ilbert, O., Salvato, M., Le Floc'h, E., et al. 2010, *ApJ*, 709, 644
- Ilbert, O., McCracken, H. J., Le Fèvre, O., et al. 2013, *A&A*, 556, A55
- Jørgensen, I., & Chiboucas, K. 2013, *AJ*, 145, 77
- Krajnović, D., Alatalo, K., Blitz, L., et al. 2013, *MNRAS*, 432, 1768
- Kroupa, P. 2001, *MNRAS*, 322, 231
- Kroupa, P., Weidner, C., Pflamm-Altenburg, J., et al. 2013, *The Stellar and Sub-Stellar Initial Mass Function of Simple and Composite Populations*, eds. T. D. Oswalt & G. Gilmore, 115
- La Barbera, F., de Carvalho, R. R., de La Rosa, I. G., et al. 2010, *MNRAS*, 408, 1313
- La Barbera, F., Ferreras, I., Vazdekis, A., et al. 2013, *MNRAS*, 433, 3017
- Larson, R. B. 1998, *MNRAS*, 301, 569
- Larson, R. B. 2005, *MNRAS*, 359, 211
- Läscher, R., van den Bosch, R. C. E., van de Ven, G., et al. 2013, *MNRAS*, 434, L31
- Longhetti, M., & Saracco, P. 2009, *MNRAS*, 394, 774
- Longhetti, M., Saracco, P., Severgnini, P., et al. 2007, *MNRAS*, 374, 614
- Longhetti, M., Saracco, P., Gargiulo, A., Tamburri, S., & Lonoce, I. 2014, *MNRAS*, 439, 3962
- Marks, M., Kroupa, P., Dabringhausen, J., & Pawlowski, M. S. 2012, *MNRAS*, 422, 2246
- Matteucci, F. 1996, *Fund. Cosm. Phys.*, 17, 283
- Munari, U., Sordo, R., Castelli, F., & Zwitter, T. 2005, *A&A*, 442, 1127
- Narayanan, D., & Davé, R. 2012, *MNRAS*, 423, 3601
- Navarro, J. F., Frenk, C. S., & White, S. D. M. 1996, *ApJ*, 462, 563
- Padoan, P., & Nordlund, Å. 2002, *ApJ*, 576, 870
- Papadopoulos, P. P. 2010, *ApJ*, 720, 226
- Papadopoulos, P. P., Thi, W.-F., Miniati, F., & Viti, S. 2011, *MNRAS*, 414, 1705
- Peng, C. Y., Ho, L. C., Impey, C. D., & Rix, H. 2002, *AJ*, 124, 266
- Reda, F. M., Forbes, D. A., & Hau, G. K. T. 2005, *MNRAS*, 360, 693
- Salpeter, E. E. 1955, *ApJ*, 121, 161
- Santini, P., Fontana, A., Grazian, A., et al. 2009, *A&A*, 504, 751
- Saracco, P., Longhetti, M., & Andreon, S. 2009, *MNRAS*, 392, 718
- Saracco, P., Longhetti, M., & Gargiulo, A. 2010, *MNRAS*, L115
- Saracco, P., Casati, A., Gargiulo, A., et al. 2014, *A&A*, 567, A94
- Shetty, S., & Cappellari, M. 2014, *ApJ*, 786, L10
- Smith, R. J. 2014, *MNRAS*, 443, L69
- Smith, R. J., & Lucey, J. R. 2013, *MNRAS*, 434, 1964
- Spiniello, C., Trager, S. C., Koopmans, L. V. E., & Chen, Y. P. 2012, *ApJ*, 753, L32
- Szomoru, D., Franx, M., Bouwens, R. J., et al. 2011, *ApJ*, 735, L22
- Szomoru, D., Franx, M., & van Dokkum, P. G. 2012, *ApJ*, 749, 121
- Tamburri, S., Saracco, P., Longhetti, M., et al. 2014, *A&A*, 570, A102
- Tinsley, B. M. 1980, *Fund. Cosm. Phys.*, 5, 287
- Tortora, C., Romanowsky, A. J., & Napolitano, N. R. 2013, *ApJ*, 765, 8
- Treu, T., Auger, M. W., Koopmans, L. V. E., et al. 2010, *ApJ*, 709, 1195
- Trujillo, I., Ferré-Mateu, A., Balcells, M., Vazdekis, A., & Sánchez-Blázquez, P. 2014, *ApJ*, 780, L20
- van de Sande, J., Kriek, M., Franx, M., et al. 2013, *ApJ*, 771, 85
- van der Wel, A., Franx, M., van Dokkum, P. G., et al. 2005, *ApJ*, 631, 145
- van der Wel, A., Bell, E. F., Häußler, B., et al. 2012, *ApJS*, 203, 24
- van Dokkum, P. G., & Conroy, C. 2010, *Nature*, 468, 940
- Wong, K. C., Tran, K.-V. H., Suyu, S. H., et al. 2014, *ApJ*, 789, L31