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# Uncertainties on near-core mixing in red-clump stars: effects on the period spacing and on the luminosity of the AGB bump

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## ABSTRACT

Low-mass stars in the He-core-burning (HeCB) phase play a major role in stellar, galactic, and extragalactic astrophysics. The ability to predict accurately the properties of these stars, however, depends on our understanding of convection, which remains one of the key open questions in stellar modelling. We argue that the combination of the luminosity of the AGB bump (AGBb) and the period spacing of gravity modes ( $\Delta\Pi_1$ ) during the HeCB phase provides us with a decisive test to discriminate between competing models of these stars. We use the Modules for Experiments in Stellar Astrophysics (MESA), a Bag of Stellar Tracks and Isochrones (BaSTI), and Padova & Trieste Stellar Evolution Code (PARSEC) stellar evolution codes to model a typical giant star observed by *Kepler*. We explore how various near-core-mixing scenarios affect the predictions of the above-mentioned constraints, and we find that  $\Delta\Pi_1$  depends strongly on the prescription adopted. Moreover we show that the detailed behaviour of  $\Delta\Pi_1$  shows the signature of sharp variations in the Brunt–Väisälä frequency, which could potentially give additional information about near-core features. We find evidence for the AGBb among *Kepler* targets, and a first comparison with observations shows that, even if standard models are able to reproduce the luminosity distribution, no standard model can account for satisfactorily the period spacing of HeCB stars. Our analysis allows us to outline a candidate model to describe simultaneously the two observed distributions: a model with a moderate overshooting region characterized by an adiabatic thermal stratification. This prescription will be tested in the future on cluster stars, to limit possible observational biases.

**Key words:** asteroseismology – stars: evolution – stars: interiors – stars: late-type – stars: low-mass.

## 1 INTRODUCTION

The helium-core-burning (HeCB) phase of low-mass stars affects many aspects of their subsequent evolution, and has therefore been the subject of numerous studies in the literature. In particular, the red clump (RC) is a well-known feature in the Hertzsprung–Russell diagram – as well as the colour–magnitude diagram – of simple and composite stellar populations, and is associated with the low-mass,

metal-rich stars in the HeCB phase (Cannon 1970). The RC plays a key role in many fields of astrophysics: its luminosity, for instance, can be used as a distance and age indicator of clusters and nearby galaxies, while the observed chemical composition of its members is useful to investigate the chemical evolution of galaxies (e.g. see Girardi & Salaris 2001; Catelan 2009; Nidever et al. 2014).

Current models of the internal structure and evolution of such stars, however, still suffer from systematic uncertainties which are due in most cases to our limited understanding of the physical processes in stellar conditions (energy transport and nuclear processes). In particular, predictions of stellar lifetimes in the HeCB phase are

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strongly dependent on the (poorly constrained) amount of mixing applied beyond boundaries of convective regions and on the adopted definition (and correct implementation) of such boundaries (e.g. see Castellani, Giannone & Renzini 1971a; Chiosi 2007; Gabriel et al. 2014; Bressan et al. 2015). Crucially, our ability to test models has been limited so far by the lack of observational constraints which are specific to the internal structure of evolved stars.

Asteroseismology of thousands of red giants observed by CoRoT (Baglin et al. 2006) and *Kepler* (Borucki et al. 2010) has changed the situation. We can now use the pulsation frequencies to place tight constraints not only on the fundamental stellar properties, but also to probe their internal structure (see e.g. Chaplin & Miglio 2013 and references therein). In particular, as presented in Montalbán et al. (2013), the frequencies of oscillation modes detected in HeCB stars are sensitive diagnostics of the chemical and thermal stratification of the energy-generating core, providing us with a novel and independent constraint, which is specific to the core structure of HeCB stars. In this paper we argue that seismic constraints, when used in conjunction with well-known classical indicators [in particular the luminosity of the AGB bump (AGBb), see below], provide us with a decisive test to discriminate competing models of different near-core mixing schemes adopted during the He-burning phase. We also show how current standard models (see Section 3) from different stellar codes do not match these constraints simultaneously.

The paper is organized as follows: we start by reviewing currently available indicators of core mixing in HeCB stars (Section 2). We describe in Section 3 the set of stellar models we use, while in Section 4 we explore the impact of several commonly adopted prescriptions for near-core mixing on the observables, with emphasis on asteroseismic constraints (Section 5). We then compare in Section 6 our predictions with asteroseismic constraints as published in recent papers (Mosser et al. 2014; Pinsonneault et al. 2014). Finally, we present a summary and future prospects in Section 7.

## 2 OBSERVATIONAL CONSTRAINTS ON HE-CORE-BURNING MODELS

One of the main observables used to constrain the mass of the fully mixed core during the HeCB phase is the  $R_2$  ratio (Buonanno, Corsi & Fusi Pecci 1985).  $R_2$  is defined as the ratio between the number of early asymptotic giant branch (eAGB) horizontal branch (HB) stars in simple stellar populations (chemically homogeneous and coeval stars) and is directly connected to the lifetime of the two phases ( $R_2 \sim \tau_{\text{AGB}}/\tau_{\text{HB}}$ ). The value of  $R_2$  is affected by the core mixing during the HeCB phase, as discussed in, e.g. Bressan, Bertelli & Chiosi (1986), Caputo et al. (1989). The use of  $R_2$  has hitherto been limited to stars in clusters, due to the small number of field stars with accurate distance available, and due to the inherent complication of dealing with an ensemble of stars with a spread in age and chemical composition (which also hampers robust inferences on their evolutionary state).

Another important observable related to the HeCB and AGB evolution of low- and intermediate-mass stars is the luminosity of the AGBb (see e.g. the review by Catelan 2007). Similarly to the bump in RGB, the AGBb is a feature of the eAGB that consists in three passages of the evolutionary track in a small interval of luminosity and causes a local increment of stars in the luminosity distribution of a stellar population. The occurrence of the AGBb is connected to the formation of the helium-burning shell (see e.g. Cassisi & Salaris 2013, paragraph 5.2). The luminosity at which it occurs depends on the location in mass of the He shell at its ignition, hence it is determined by the maximum extension of the

mixed core during the HeCB phase. While the AGBb had been highlighted in stellar evolutionary tracks a few decades ago (see e.g. Caputo, Castellani & Wood 1978), its first identification as a distinct observational feature in galaxies was reported by Gallart (1998). The AGBb is nowadays observed in a large number clusters and nearby galaxies (e.g. see Alcock et al. 1997; Ferraro et al. 1999; Beccari et al. 2006; Dalcanton et al. 2012). A further important property of the AGBb is the weak dependence of the luminosity ratio between the RC and the AGBb on the metallicity and initial helium abundance (Castellani, Chieffi & Pulone 1991; Bono et al. 1995).

Asteroseismology of white dwarfs may also provide us with observational constraints to test models in the HeCB and AGB phase. The C/O profile of a white dwarf at the beginning of the cooling sequence corresponds to that of the stellar core at the end of the AGB, hence it is largely determined by the combined action of chemical mixing and nuclear burning during the HeCB phase. Using pulsation modes detected in WDs, Metcalfe, Salaris & Winget (2002) found a discrepancy in the central oxygen-to-carbon ratio between stellar models and the value inferred from seismic data. They ascribed this difference to an underestimation of  $^{12}\text{C}(\alpha,\gamma)^{16}\text{O}$  cross-section; however, as pointed out by Straniero et al. (2003), the final C/O ratio in models also depends on the amount of mixing applied in the convective regions beyond the Schwarzschild border, and on the adopted definition of such boundaries.

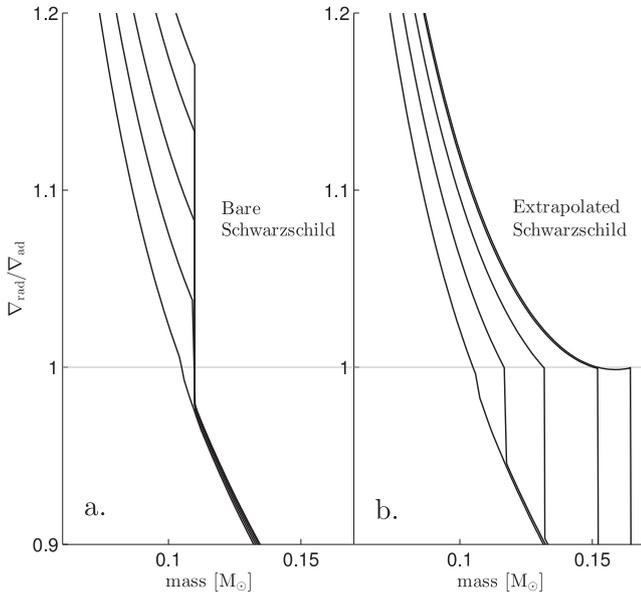
While providing valuable information, currently available tests cannot be used to discriminate competing models. In this paper we propose a way forward, which is based on the combination of said constraints and the more direct diagnostics of the conditions in the core provided by non-radial modes observed in HeCB stars.

## 3 STELLAR MODELS

In our exploratory analysis we consider models with  $M = 1.5 M_{\odot}$  and solar chemical composition. This set of parameters is both typical of the sample of RC giants observed by *Kepler* (see e.g. Mosser et al. 2014; Pinsonneault et al. 2014), and similar to those of giants in the cluster NGC 6819 (see Basu et al. 2011; Miglio et al. 2012; Sandquist et al. 2013). Models in HeCB phase are computed using three different stellar structure and evolution codes [Modules for Experiments in Stellar Astrophysics (MESA), a Bag of Stellar Tracks and Isochrones (BaSTI), and PAdova & TRieste Stellar Evolution Code (PARSEC)] and several assumptions about near-core mixing.

A first set of models presented in this study is computed using the MESA code (Paxton et al. 2011). The choice of parameters and of the relevant physics is the following:

- (i) the initial mass is  $1.5 M_{\odot}$ , and no mass loss is considered during the RGB;
- (ii) the mass fraction of heavy elements ( $Z = 0.01756$ ), the initial helium abundance ( $Y = 0.26556$ ), and the mixing-length parameter ( $\alpha_{\text{mlt}} = 1.69278$ ) are those resulting from a solar-calibrated model without diffusion;
- (iii) the partition of heavy elements is taken from Grevesse & Noels (1993);
- (iv) the nuclear reaction rates are taken from Angulo et al. (1999);
- (v) during the He-burning stage one has to properly account for the C and O abundance changes in the stellar matter and their effect on radiative opacities, hence we use opacity tables accounting for various levels of C and O enhancements (Type 2 opacity tables; Iglesias & Rogers 1996);



**Figure 1.** Ratio between the radiative and adiabatic temperature gradients as a function of mass in MESA HeCB models computed using the BS (panel a) and ES (panel b) scheme. The latter leads to a self-consistent definition of the convective boundary (following the Schwarzschild criterion), as the adiabatic and radiative gradients are equal at the boundary of the core. From bottom to top the lines are for a sequence of models at the stages of  $Y_c = 0.93, 0.87, 0.81, 0.74,$  and  $0.70$ .

(vi) the evolution includes the pre-main-sequence phase and is computed until the first thermal pulse along the AGB (TP-AGB);

(vii) in models computed with overshooting, the overshooting scheme adopted is a step function in which the parameter  $\alpha_{\text{ovHe}}$  indicates the extension of the mixed region from the Schwarzschild border in units of the local pressure scale height. The overshooting region is instantaneously mixed (Maeder 1975);

(viii) we consider two scenarios for the temperature gradient ( $\nabla_T$ ) in the overshooting region, according to the definition given in Zahn (1991):  $\nabla_T = \nabla_r$  (overshooting) or  $\nabla_T = \nabla_{\text{ad}}$  (penetrative convection), where  $\nabla_r$  and  $\nabla_{\text{ad}}$  follow the usual notation for the radiative and adiabatic temperature gradient;

(ix) during the main-sequence phase we adopt an overshooting parameter  $\alpha_{\text{ovH}} = 0.2$  (Nevertheless, the choice of  $\alpha_{\text{ovH}}$  does not impact the structure of a low-mass HeCB star).

We compute the following MESA models with different schemes for convective mixing:

**BS.** The ‘bare Schwarzschild’ (BS) model (to follow the notation in Straniero et al. 2003). In the BS model, the Schwarzschild criterion is applied on the radiative side of the convective boundary which implies that the convective core is not allowed to grow, leading to a discontinuity in the chemical composition (hence in the radiative gradient) at the border of the convective core (Fig. 1 panel a).

A recent paper by Gabriel et al. (2014) showed how this implementation of the Schwarzschild criterion leads to an inconsistent location of the convective border, as the convective luminosity is non-zero there ( $\nabla_r > \nabla_{\text{ad}}$  at the inner side of the convective boundary). The book of Schwarzschild (1958) and the analysis presented in Castellani et al. (1971a) also led to similar conclusions. The latter consider the BS convective border as in an unstable equilibrium in the sense that by extending outwards the convective core

by an arbitrarily small quantity,  $\nabla_r$  at the new border may be larger than  $\nabla_{\text{ad}}$ , hence, the region is convectively unstable according to the Schwarzschild criterion. To find a stable border, the convective core must be extended until radiative and adiabatic gradient become equal (induced overshooting). However, as the HeCB proceeds ( $Y_c \lesssim 0.69$ ) this scheme generates a local minimum in  $\nabla_r$  within the mixed core. The addition of more radiative layers surrounding the increasingly larger convective core will decrease  $\nabla_r$  to the value of  $\nabla_{\text{ad}}$  at the location of the minimum. This creates a separate convective region in layers beyond the location of this minimum, where  $\nabla_r$  is still larger than  $\nabla_{\text{ad}}$ . The treatment of this external convective region is problematic. Full mixing between layers inside the minimum of  $\nabla_r$  and the external convective shell cannot happen, because otherwise the minimum of  $\nabla_r$  would decrease below the local value of  $\nabla_{\text{ad}}$  with the contradiction of a fully mixed region where  $\nabla_r$  is however locally lower than  $\nabla_{\text{ad}}$ . A solution to this problem is the formation of a partially mixed – *semiconvective*– region between the minimum of  $\nabla_r$  and the outer radiative zone (Castellani, Giannone & Renzini 1971b). This is usually treated with dedicated algorithms that allow for partial chemical mixing to satisfy  $\nabla_r = \nabla_{\text{ad}}$  in this region, with the consequence of creating a smooth gradient of chemical composition before the sharp discontinuity due to the HeCB. Although we discourage to adopt the BS model, we use this mixing scheme in order to compare our models with those described in the literature (e.g. Straniero et al. 2003; Kippenhahn, Weigert & Weiss 2013).

**HOV.** High overshooting model with  $\alpha_{\text{ovHe}} = 1$  (HOV).

**HPC.** Penetrative Convection model with a high overshooting parameter  $\alpha_{\text{ovHe}} = 1$  (HPC).

For low and moderate values of  $\alpha_{\text{ovHe}}$ , overshooting models (as well as penetrative convection) might develop a semiconvective region similar to the case presented by Castellani et al. (1971b). For high values of the overshooting parameter (e.g. HOV and HPC) the extra mixed region becomes large enough to prevent the formation of a semiconvective region (see e.g. Bressan et al. 1986; Straniero et al. 2003). Moreover, in these cases the determination of the fully mixed region requires the application of the Schwarzschild criterion in layers where there is no chemical composition gradient/discontinuity, which greatly simplifies the numerical implementation of the convective-instability criterion.

In order to extend our study of the mixing-schemes commonly adopted, we also consider stellar models from other evolution codes (BaSTI and PARSEC). To compute those models we used, where possible, the same physical inputs adopted in MESA.

**BaSTI-SC.** We have also computed models with the BaSTI code (Pietrinferni et al. 2004, 2006, 2013). The BaSTI model has  $M = 1.5 M_{\odot}$ , computed for  $Z = 0.0176$ ,  $Y = 0.266$ , and  $\alpha_{\text{mlt}} = 1.69$  with Grevesse & Noels (1993) heavy element partition. The input physics relevant to this analysis is the same as in the MESA calculations but for the  $^{12}\text{C}(\alpha, \gamma)^{16}\text{O}$  reaction rate that comes from Kunz et al. (2002). Core mixing during the HeCB stage, induced overshooting and semiconvection have been taken into account by adopting the numerical scheme first introduced by Castellani et al. (1985) and previously described. Indeed, a semiconvective region starts to appear when the central abundance by mass of He is  $\sim 0.7$ . The occurrence of breathing pulses – which appears when the central abundance of He drops below of 0.12 – is inhibited by imposing that the abundance of helium in the core is not allowed to increase at each time-step. The evolution starts from the pre-main sequence to the RGB-tip, and is resumed at the start of the HeCB after the core

electron degeneracy has been removed (the helium-flash evolution is not computed), and 3 per cent of carbon has been produced during the He-flash. No core convective overshoot during the central H-burning stage and no mass loss during the RGB phase are taken into account.

**PARSEC-LOV.** We also computed models using PARSEC (Bressan et al. 2012, 2013) for  $M = 1.5 M_{\odot}$ ,  $Z = 0.017$ , and  $Y = 0.279$ . We briefly summarize below the main input physics adopted for the current models. The nuclear reaction rates and corresponding  $Q$ -values are the recommended values in the Joint Institute for Nuclear Astrophysics (JINA) reaclib data base (Cyburt et al. 2010). The high-temperature opacities,  $4.2 \leq \log(T/K) \leq 8.7$ , are provided by the Opacity Project At Livermore (OPAL) team (Iglesias & Rogers 1996, and references therein) and the low-temperature opacities,  $3.2 \leq \log(T/K) \leq 4.1$ , are from  $\text{\AA}ESOPUS^1$  tool (Marigo & Aringer 2009). The equation of state is computed with the FreeEOS code (A.W. Irwin<sup>2</sup>). The heavy element partition is from Caffau et al. (2011). The MLT parameter,  $\alpha_{\text{mlt}} = 1.74$ , is calibrated on the solar model accounting for element diffusion. The standard mixing scheme in PARSEC considers an overshooting parameter of  $\Lambda_c = 0.5$  across the formal Schwarzschild border, which means about  $0.25H_p$  above it, and a radiative thermal stratification of the extra-mixed region. Similar to BaSTI, the evolution starts from the pre-main sequence, stops at the helium flash and restarts after the core electron degeneracy has been removed after the flash. The amount of carbon consumed to remove the electron degeneracy is computed from the variation of the gravitational binding energy of the core during the flash. During the HeCB phase, besides accounting for core overshooting, the code may deal with residual semi-convective instabilities using the Schwarzschild criterion, and suppresses possible breathing pulses of convection.

**PARSEC-MPC.** We use PARSEC to compute a modified track in which we extend the mixed He core by penetrative convection (adiabatic stratification) with a moderate value of overshooting ( $\Lambda_c = 1.0$ , i.e.  $\alpha_{\text{ovHe}} \sim 0.5$ ).

**ES.** Finally, we include in the analysis a model obtained with MESA in which we include an external routine to implement a similar prescription as in BaSTI-SC. This routine consists in the following steps:

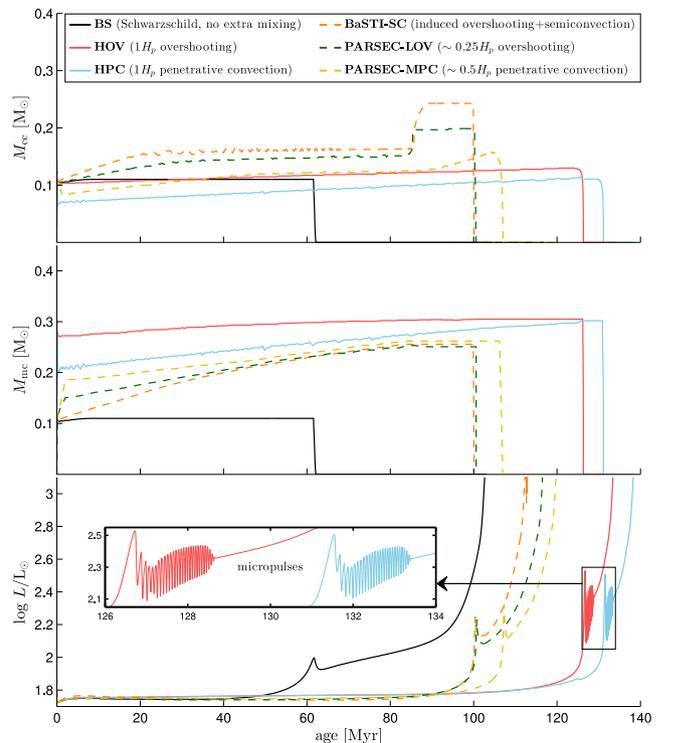
- (1) at the beginning of each time-step we set the position of the convective-core boundary and fully mix the convective region according to the Schwarzschild criterion as implemented in MESA.
- (2) we let the code calculate the burning during the time-step.
- (3) at the end of the time-step, we check whether the radiative gradient (with the new composition determined by the burning) at the convective border is higher than the adiabatic gradient.
- (4a) if this is the case, we restart the time-step from point 1, but extending the core boundary by one mesh.
- (4b) if not, the equality of the gradients has been achieved and the code can continue to the next time-step (point 1).

As a result of this algorithm the convective core extends naturally (Fig. 1 panel b). We call this model Extrapolated Schwarzschild (ES). The internal structure and seismic predictions of ES (see Section 5.1 and Figs 3 and 4) are remarkably similar to the BaSTI-SC. Differently from BaSTI and PARSEC, however, a routine to treat semiconvection is still to be developed in MESA, hence we stop the evolution of the ES model at  $Y_c \sim 0.69$ .

#### 4 PREDICTED STELLAR PROPERTIES OF MODELS IN THE HECB AND AGB PHASE

We now compare the properties of the series of models presented in Section 3, with particular emphasis on those which can be tested via a direct comparisons to observations. We start by comparing predictions of non-seismic observables (see Section 2), while in Section 5 we focus on seismic diagnostics.

The main consequence of applying different mixing schemes in the HeCB phase is to vary the core mass undergoing convective mixing. We refer to *convective – core mass* ( $M_{\text{cc}}$ ) as the core mass in which  $\nabla_r \geq \nabla_{\text{ad}}$  (formal Schwarzschild core), while the *mixed – core mass* ( $M_{\text{mc}}$ ) includes also the extra-mixing region (e.g. convective core + overshooting and/or semiconvection). The mixed-core mass also indicates the location in the stellar structure of the discontinuity in the chemical composition due to HeCB. In the ES models (in the domain we were able to explore) and BS models,  $M_{\text{cc}}$  corresponds to  $M_{\text{mc}}$  since no extra mixing is introduced. This is also true for the BaSTI-SC model when  $Y_c > 0.7$ , i.e. where semiconvection has not appeared yet, and during the last stages of the HeCB, where the convective core grows rapidly and the size of the semiconvective region is reduced (see Fig. 2, upper panel). In HOV, HPC, PARSEC-LOV, and PARSEC-MPC,  $M_{\text{cc}}$  and  $M_{\text{mc}}$  are distinct from the beginning to the end of the HeCB phase (Fig. 2, upper and middle panels). Although penetrative convection models are expected to have smaller  $M_{\text{cc}}$  and  $M_{\text{mc}}$  than overshooting models (see Godart 2007; Noels et al. 2010, in the case of massive main-sequence stars),  $\max(M_{\text{mc}})$  in HOV and HPC is very similar, providing a similar  $L_{\text{AGBb}}$  (see Fig. 2, upper and lower panels).



**Figure 2.** Convective-core (upper panel), mixed-core mass (middle panel) and total luminosity (lower panel) as a function of time for the BS, HOV, HPC, BaSTI-SC, PARSEC-LOV, and PARSEC-MPC models from the start of the HeCB phase up to the first AGB-TP. Both HPC and HOV models show micropulses.

<sup>1</sup> <http://stev.oapd.inaf.it/aesopus>

<sup>2</sup> <http://freeeos.sourceforge.net/>

**Table 1.** Non-seismic properties of 1.5- $M_{\odot}$  models computed adopting different mixing schemes.

	MESA BS	MESA HOV	MESA HPC	MESA ES	BaSTI SC	PARSEC LOV	PARSEC MPC
Time-span HeCB phase ( $\tau_{\text{HeCB}}$ ) [Myr]	60.4	126.2	131.1	–	100.2	100.6	107.0
$R_2$ ratio ( $\frac{\tau_{\text{AGB}}}{\tau_{\text{HeCB}}}$ )	0.71	0.06	0.06	–	0.12	0.15	0.12
$M_{\text{cc}}$ at $Y_c \sim 0.7$ [ $M_{\odot}$ ]	0.100	0.113	0.084	0.162	0.165	0.140	0.107
$M_{\text{mc}}$ at $Y_c \sim 0.7$ [ $M_{\odot}$ ]	0.100	0.292	0.242	0.162	0.165	0.187	0.206
Maximum extension $M_{\text{cc}}$ [ $M_{\odot}$ ]	0.100	0.130	0.114	–	0.243	0.199	0.158
Maximum extension $M_{\text{mc}}$ [ $M_{\odot}$ ]	0.100	0.305	0.302	–	0.243	0.251	0.262

#### 4.1 Luminosity and duration of the HeCB and eAGB phases

An obvious effect of increasing  $M_{\text{mc}}$  is to increase the duration of the HeCB phase ( $\tau_{\text{HeCB}}$ ). In models with larger  $M_{\text{mc}}$  more fuel is available for the triple- $\alpha$  and  $^{12}\text{C}(\alpha, \gamma)^{16}\text{O}$  nuclear reactions, increasing the time necessary to deplete all the helium in the core, i.e. the lifetime of the HeCB phase. In HOV and HPC models, the HeCB lifetime increases by about 40 per cent compared to the BS model, while in BaSTI-SC and PARSEC-LOV by about 20 per cent. However, if we look at the lifespans of the single phases, the increase in duration of the HeCB phase corresponds to a (non-linearly proportional) decrease of the duration of the AGB phase, since the formation of the He shell takes place closer to the H shell. The BS model has shorter lifetime with respect to the HPC and HOV models; however, it has a longer AGB phase. The BaSTI-SC, PARSEC-LOV, and PARSEC-MPC tracks have a behaviour which is in between the BS and the HPC/HOV models. In general, we expect that models with increasing  $M_{\text{mc}}$  produce a longer HeCB phase and a less populated early-AGB, combined with a decreasing  $R_2$  factor from BS models to the HOV and HPC models. These results are reported in Table 1.

The stellar luminosity is also affected by the mixing scheme adopted during HeCB (Fig. 2, lower panel), and  $L_{\text{AGBb}}$  increases when extra mixing is added, while it remains almost unchanged when comparing HOV and HPC tracks (Table 1). The maximum extension of  $M_{\text{mc}}$  (which is very similar in HOV and HPC models) corresponds to the inner border of the He-shell at its ignition and it determines the  $L_{\text{AGBb}}$ .

Models with larger cores (HOV and HPC) show *micropulses* (Mazzitelli & D’Antona 1986), i.e. a series of secular instabilities that may occur during the formation of the He-burning shell. Micropulses appear after the maximum in luminosity of the AGBb if the nuclear reactions in the core stop before the shell is ignited in ‘thin’ conditions (see Schwarzschild & Härm 1965). To support this statement we have looked at the central helium abundance and the nuclear energy generation rate at the start of the AGBb. The values found indicate that in the HPC and HOV models the contribution of HeCB to the luminosity is negligible, since helium is almost completely depleted. On the other hand, in the BS model core-He burning still contributes significantly to the luminosity. We note that for micropulses to be resolved by a stellar evolution code, a small time-step is needed in the numerical simulations: each pulse lasts for about  $5 \times 10^4$  yr and the duration of the entire phenomenon is about 2 million years.

The study of micropulses goes beyond the purpose of this paper and for detailed explanations we refer the reader to Mazzitelli & D’Antona (1986) and the more recent work by Gautschy & Althaus (2007). Also notable are the papers by Bono et al. (1997) and Sweigart et al. (2000) where the connection between ‘gravonuclear instabilities’ and micropulses is investigated.

## 5 ASTEROSEISMIC SIGNATURES

### 5.1 Asymptotic gravity-mode period spacing

We have shown that there are no significant differences between models with similar  $M_{\text{mc}}$  when considering the luminosity as a function of time (see e.g. HOV and HPC in Fig. 2). These models, however, have very distinct seismic properties. Montalbán et al. (2013) showed that extending the adiabatically stratified central region leads to a larger value of the gravity-mode period spacing. The reason why the asymptotic period spacing of gravity modes ( $\Delta\Pi_1$ ) during the HeCB phase is a sensitive probe of the temperature stratification of near-core regions is directly related to the behaviour of the Brunt–Väisälä frequency  $N$ , and its relation with  $\Delta\Pi_1$ .

In the stellar interior,  $N$  depends on the local temperature and chemical composition gradients:

$$N^2 = \frac{g^2 \rho}{P} \frac{\chi_T}{\chi_\rho} \left( \nabla_{\text{ad}} - \nabla_T - \frac{\chi_\mu}{\chi_T} \nabla_\mu \right), \quad (1)$$

where  $\nabla_\mu = \text{dln } \mu / \text{dln } P$ ,  $\chi_T = (\partial \ln P / \partial \ln T)_{\rho, \mu}$ ,  $\chi_\rho = (\partial \ln P / \partial \ln \rho)_{T, \mu}$ , and  $\chi_\mu = (\partial \ln P / \partial \ln \mu)_{T, \rho}$ . In a fully mixed region,  $\nabla_\mu$  is null, therefore the  $N^2$  profile is directly proportional to the difference between  $\nabla_T$  and  $\nabla_{\text{ad}}$  (Fig. 4a).

In the extra-mixed region of models with overshooting (e.g. HOV and PARSEC-LOV)  $\nabla_T < \nabla_{\text{ad}}$ , therefore

$$N^2 \propto \nabla_{\text{ad}} - \nabla_T = \nabla_{\text{ad}} - \nabla_T > 0,$$

while in the corresponding region of a penetrative convection model (e.g. HPC and PARSEC-MPC)  $\nabla_T = \nabla_{\text{ad}}$ , hence

$$N^2 \propto \nabla_{\text{ad}} - \nabla_T = 0.$$

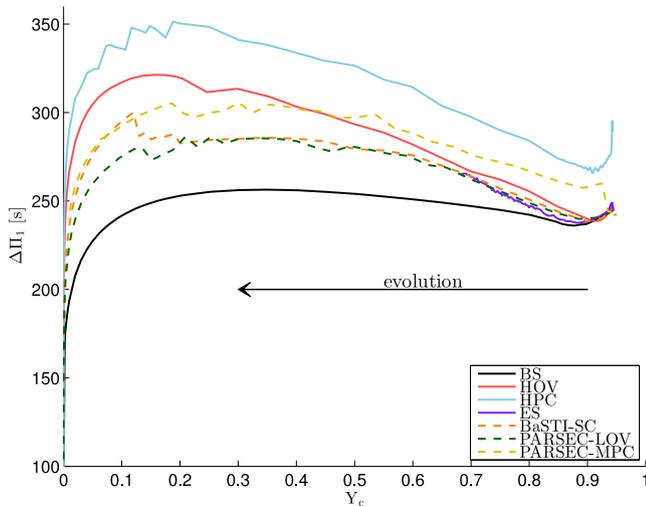
The asymptotic period spacing of gravity modes is related to the Brunt–Väisälä frequency according to the relation (Tassoul 1980):

$$\Delta\Pi_l = \frac{2\pi^2}{\sqrt{l(l+1)}} \left( \int_{r_1}^{r_2} N \frac{dr}{r} \right)^{-1}, \quad (2)$$

where  $r_2$  and  $r_1$  are the boundaries of the g-mode cavity (i.e. where  $\omega^2 = N^2$ ). Consequently, HeCB penetrative-convection models have higher values of  $\Delta\Pi_1$  compared to overshooting models (Fig. 3).

A similar effect can be found if we compare models with increased mixed-core size. BS models, in fact, have lower  $\Delta\Pi_1$  compared to all other models, followed by models with induced overshooting (BaSTI-SC and ES), then high overshooting model, and finally high penetrative convection models (see Fig. 4a).

Independently from the convective-mixing scheme adopted (BS, HOV and HPC), the period spacing of models in the HeCB is higher compared to that on the RGB at the same luminosity ( $\Delta\Pi_{1\text{RGB}} \sim 60 - 50$  s), while after the early-AGB phase  $\Delta\Pi_1$  decreases to similar or smaller values (Montalbán & Noels 2013).



**Figure 3.** Period spacing against the central helium mass fraction during the HeCB phase for the model computed. BS models present the smallest period spacing, while the penetrative convection models have the largest values. Moreover, penetrative convective models start with larger period spacing.

It is interesting to notice that if models present multiple gravity-mode cavities, we do not expect them to show a regular period spacing. This is the case of HOV and HPC models in the post-HeCB phase during the micropulses. In fact, the He shell can experience convection within each pulse, in conjunction with the maximum of the nuclear energy generation. Stars presenting this scenario (if any at all exists) may be missed by analysis based upon looking for a simple pattern in  $\Delta\Pi_1$  (e.g. Mosser et al. 2012).

## 5.2 Period spacing of numerically computed adiabatic frequencies

From the observed frequency spectra we can estimate  $\Delta\Pi_1$  based on the detection of modes that have relatively small inertias, i.e. gravity modes that have a significant coupling with the low-inertia acoustic modes (see e.g. Christensen-Dalsgaard 2012 and reference therein). The frequencies of these mixed gravito-acoustic modes are expected from theory to follow a relatively simple pattern (see e.g. Unno et al. 1989; Mosser et al. 2012), which can be fitted to the observation to estimate  $\Delta\Pi_1$ , provided that a sufficient number of modes are detected, and that the analytical approximation for the expected pattern of mixed modes is accurate (see e.g. Beck et al. 2011; Mosser et al. 2012; Montalbán et al. 2013; Jiang & Christensen-Dalsgaard 2014).

It is thus crucial not only to make predictions of RC  $\Delta\Pi_1$  using the approximated expression in equation (2), but also to compute the spectrum of individual modes, which may show interesting departures from the expected approximated relation/pattern, and that eventually can be compared with the detailed observed frequency spectrum.

In Figs 4(b)–(f) we present the properties of adiabatic pulsation modes computed with GYRE (Townsend & Teitler 2013), and compare frequency spectra of models with the same central helium abundance ( $Y_c \simeq 0.7$ ), but computed with different convective-mixing schemes. In the upper panel of each figure, it is possible to see how the inertia ( $E$ ) of each mode varies in a frequency interval around the frequency of maximum oscillations power ( $\nu_{\max}$ ). Figs 4(b)–(f) also show that the characteristic asymptotic behaviour

of the modes (the constant frequency separation for the low-inertia, pressure-dominated modes and the constant period spacing for the high-inertia, gravity-dominated modes) is a good representation of the detailed, numerically computed frequency spectrum. Moreover, the asymptotic value of the period spacing clearly reflects the differences in the Brunt–Väisälä frequency near the core (Fig. 4a and equation 2), with the BS model having the lowest  $\Delta\Pi_1$ , and the HPC model the highest.

## 5.3 Signatures of sharp-structure variations in the period spacing

As evinced from Fig. 4(a) and, more clearly, from Fig. 5,  $N$  may have sharp changes due e.g. to chemical composition gradients and/or in the temperature gradient in radiative regions. Whether such glitches have a significant impact of the period spacing depends on their location, their sharpness, and to the typical local wavelength of the gravity modes of interest.

As described in the literature (see e.g. Berthomieu & Provost 1988; Brassard et al. 1992; Miglio et al. 2008), the signature of a sharp feature in the Brunt–Väisälä frequency is a periodic component in the periods of oscillations, and therefore in the period spacing, with a periodicity in terms of the radial order  $n$  given by:

$$\Delta n \simeq \frac{\Pi_R}{\Pi_{\text{glitch}}}, \quad (3)$$

where the total buoyancy radius is defined as:

$$\Pi_R^{-1} = \int_{r_0}^R \frac{N}{r'} dr', \quad (4)$$

and local buoyancy radius is

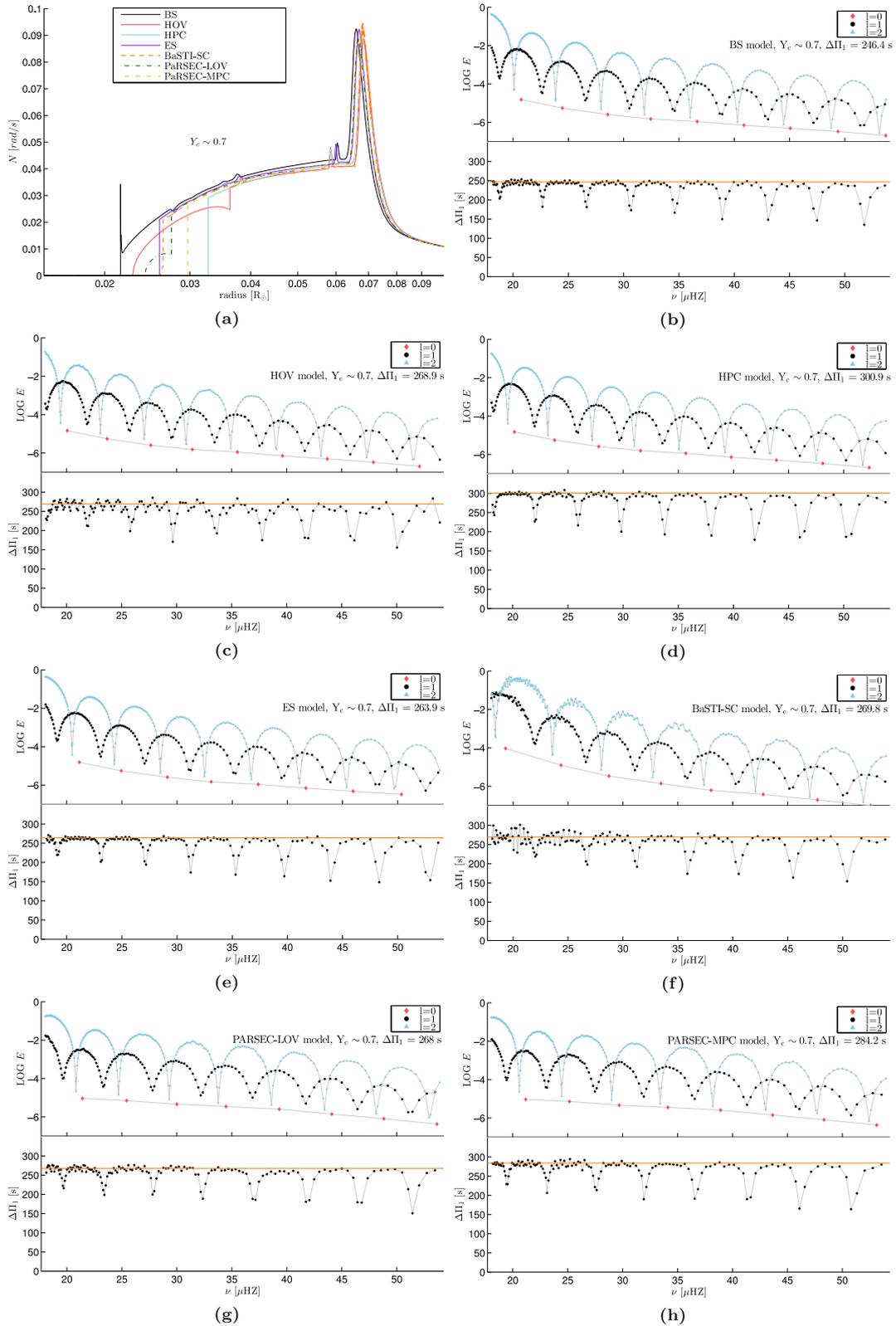
$$\Pi_r^{-1} = \int_{r_0}^r \frac{N}{r'} dr', \quad (5)$$

with  $r_0$  and  $R$  being the inner and outer boundary of the g-mode propagation region. The periodicity of the components in terms of radial order is therefore an indicator of the location of the glitch, expressed in terms of its normalized buoyancy radius ( $\Pi_R/\Pi_{\text{glitch}}$ ).

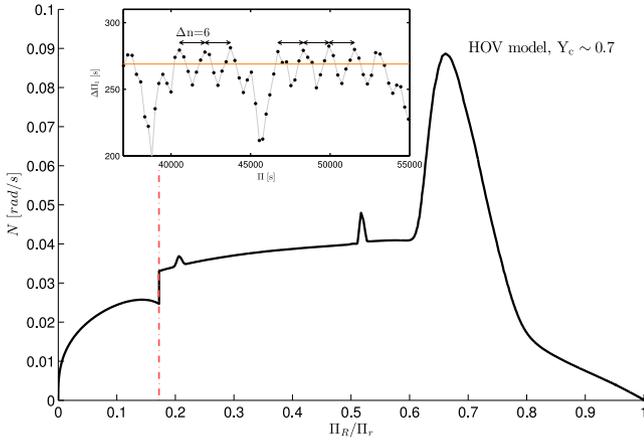
In the models we are focusing on, we notice two main glitches. A prominent, yet smooth, glitch due to the H-burning shell and associated with  $\nabla_\mu$  (see e.g. the glitch located at  $\Pi_R/\Pi_r \simeq 0.7$  in Fig. 5). In models at the beginning of the HeCB phase, in which the H-burning shell is still very thin, this glitch is sharper and may give rise to significant departures from a smooth g-mode period spacing (see Bildsten et al., in preparation). In models during most of the HeCB phase, however, this glitch does not appear to give rise to significant deviations from the asymptotic  $\Delta\Pi_1$  expected for high-order g modes.

More interestingly, in the HOV model (see Figs 4c and 5) we notice a sharp variation in  $N$  which can be well described by a step function.<sup>3</sup> Given the location of this glitch  $\Pi_R/\Pi_r \simeq 0.17$  (see Fig. 5) we expect a periodic deviation from the asymptotic  $\Delta\Pi_1$  with a periodicity of  $\Delta n \simeq 6$ , where  $n$  is the radial order of gravity modes (see e.g. Miglio et al. 2008). This corresponds well (at least in the regions of pure g modes, e.g. in the range  $4.7\text{--}5.3 \times 10^4$  s) to the periodicity of the component (see inset of Fig. 5).

<sup>3</sup> At the boundary of the fully mixed radiative (overshoot) region, the discontinuity in opacity, due to the difference between carbon-rich mixed layers and He-rich surrounding layers, leads to a discontinuity in  $\nabla_T$  and thus in  $N$ .



**Figure 4.** (a) Brunt-Väisälä frequency in the stellar interior of models with  $Y_c \simeq 0.7$  and different convective-mixing scheme. b-h: oscillation mode properties for the different convective schemes. Upper panel: mode inertia as a function of the frequency for modes with angular degree  $\ell = 0, 1, 2$ . Lower panel: period spacing of numerically computed dipolar-mode frequencies (dots) compared with the asymptotic value (solid orange line).



**Figure 5.** Brunt–Väisälä frequency as a function of the normalized buoyancy radius  $\Pi_r/\Pi_R$  (see text) in an HOV model with  $Y_c \simeq 0.7$ . In the rectangle,  $\ell = 1$  period spacing as a function of period for the same model. The period spacing of high-order g modes (e.g. in the period range 4.7–5.3  $10^4$  s) is well described by the superposition of the asymptotic  $\Delta\Pi_1$  (orange line) and a component with periodicity  $\Delta n \simeq 6$ . This periodicity indicates (see equation 3) a sharp-structure variation located at  $\Pi_R/\Pi_r \simeq 1/\Delta n \simeq 0.17$ , which corresponds well with the position of glitch in  $N$ .

While the full description of these glitches is beyond the scope of the present paper, we note that departures from the simple description of  $\Delta\Pi_1$  expected from the interaction between high-order g modes and an acoustic mode (e.g. see Unno et al. 1989; Mosser et al. 2012; Jiang & Christensen-Dalsgaard 2014; Cunha et al. 2015) provide additional, potentially very sensitive, probes of sharp-structure variations in near-core regions during the HeCB phase.

## 6 FIRST COMPARISON WITH OBSERVATIONS: AGB BUMP AND PERIOD SPACING

### 6.1 The AGB bump in *Kepler* red giants

A catalogue of seismic ( $\Delta\nu$  and  $\nu_{\max}$ ) and spectroscopic ( $[M/H]$  and  $T_{\text{eff}}$ ) constraints for  $\sim 1600$  *Kepler* giants was recently published by the APOKASC collaboration (Pinsonneault et al. 2014).

We estimate stellar masses and luminosities of these stars by using the so-called direct method, i.e. by combining  $T_{\text{eff}}$  with the seismic radii (estimated using  $\Delta\nu$  and  $\nu_{\max}$  through scaling relations). This method is known to lead to less precise estimates of  $M$  and  $R$  than so-called ‘grid-based’ approaches, yet it is less dependent on stellar evolutionary tracks (e.g. see Chaplin & Miglio 2013, and references therein) and completely independent of bolometric corrections. We select stars from the APOKASC catalogue in a range of mass  $M = 1.3\text{--}1.7 M_{\odot}$  and metallicity  $[M/H] = -0.4$  to 0.4. The luminosity function of such stars displays a peak that is spread over about four bins and has a maximum at  $\log L/L_{\odot} \sim 2.2$ : we interpret this peak as a strong candidate for the AGBb (Fig. 6). This statement is supported by the fact that 97.5 per cent of the stars have errors smaller than the bin size. We also calculated 1000 realizations of the observed sample, assuming Gaussian errors on  $\Delta\nu$ ,  $\nu_{\max}$ ,  $T_{\text{eff}}$  (taken from the APOKASC catalogue) and found that the properties of the peak in the luminosity function are not significantly affected.

As a word of caution we would like to stress that the APOKASC catalogue may, however, be affected by target selection biases (Pinsonneault et al. 2014). Although the maximum in the observed luminosity distribution does not appear to be significantly affected

by widening the metallicity range, we notice that, if we extend the range of masses down to  $1 M_{\odot}$ , the position of the peak is lowered by 1 bin (0.07 dex).

To check whether the AGBb is a feature we expect to be able to detect in a composite stellar population we use the TRILEGAL code (Girardi et al. 2012) to simulate the galactic population expected in *Kepler* field. The stellar models used are based on Padova tracks (Bressan et al. 2012), with an overshooting parameter of  $\Lambda_c = 0.5$  (same mixing scheme adopted for PARSEC-LOV model). We apply to the synthetic sample the same selection in mass and metallicity as in the observed sample. The synthetic population also displays a well-defined peak with a luminosity compatible with that of the candidate AGBb in the APOKASC catalogue (Fig. 7).

From the simulations we estimate that about half of the red-giant stars in that peak belong to the RGB. This means that in the case of the full APOKASC catalogue (considering the entire range of mass and metallicity) we expect about 20–40 AGBb stars.

In order to make a first comparison between our models and the observations, we generate histograms of the luminosity based on each of our tracks. Results of this comparison are shown in Fig. 8. For the sake of clarity in the figure we have omitted RGB models. All the models are able to reproduce reasonably well the position of the RC, taking into account also the fact that our models are representative of a single-mass, single-metallicity population only, and that we have not added the effect of observational uncertainties when building the synthetic luminosity function.

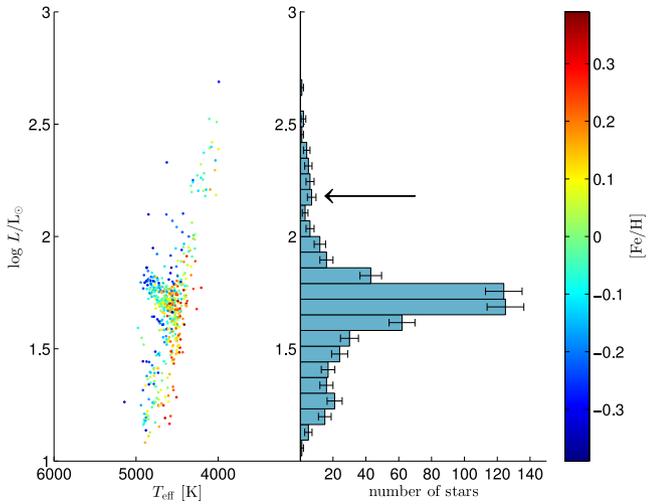
The luminosity of the AGBb predicted by BaSTI-SC, and PARSEC-MPC, is remarkably similar, and in good agreement with the candidate AGBb luminosity as detected in the observations. While PARSEC-LOV is another acceptable model, interestingly, the AGBb predicted by the BS, HOV, and HPC models is in clear disagreement with the observations, being either too faint (BS) or too bright (HOV, HPC).

### 6.2 Period spacing of *Kepler* RC stars

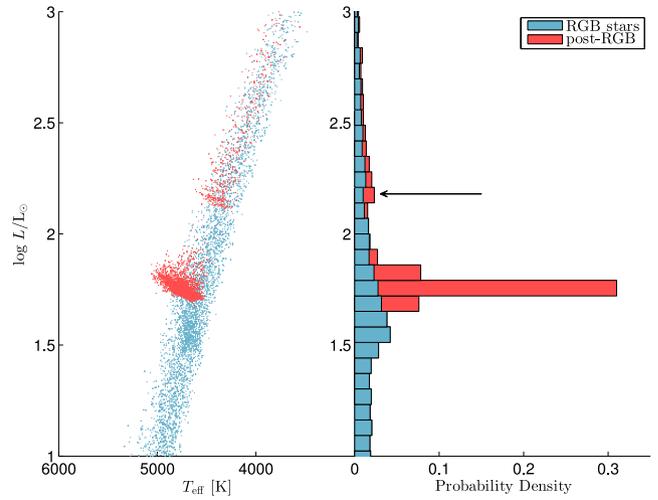
Mosser et al. (2014) recently published a catalogue that contains a large sample of red giants in the *Kepler* field for which g-mode asymptotic period spacings were inferred by modelling the interaction between pressure and gravity modes (see Unno et al. 1989; Mosser et al. 2012). Thanks to this catalogue we are able to compare the observed period spacing with theoretical predictions. We select RC stars (i.e. stars with  $\Delta\Pi_1 > 175$  s) in a range of mass between  $1.3 M_{\odot}$  and  $1.7 M_{\odot}$ , and compare the distribution of observed  $\Delta\Pi_1$  with predictions from our models (see Fig. 9).

None of the models considered seems to describe satisfactorily the entire observed distribution. The BS model can be ruled out, since  $\max \Delta\Pi_1$  is 80–90 s smaller than the maximum in the observed data. BaSTI-SC, which was one of our best candidate models based on the comparison with the AGBb luminosity, shows a main peak in the  $\Delta\Pi_1$  distribution about 20–30 s below the observed one and cannot describe period spacings higher than 300s. PARSEC-LOV has a very similar behaviour to BaSTI-SC and has a maximum  $\Delta\Pi_1 \simeq 285$  s. The HPC model is able to reproduce a large domain of the observed distribution, although its  $\Delta\Pi_1$  appears shifted to higher values with respect to the observations. Finally, although PARSEC-MPC and HOV cannot reproduce the highest observed  $\Delta\Pi_1$ , their main peaks agree with the observed one within 1–2 bins resolution.

In the observed distribution of  $\Delta\Pi_1$  (see Fig. 9) we notice, in addition to the main peak, a secondary peak at  $\Delta\Pi_1 \sim 253$  s, which still belongs to the RC population. Interestingly all our models,



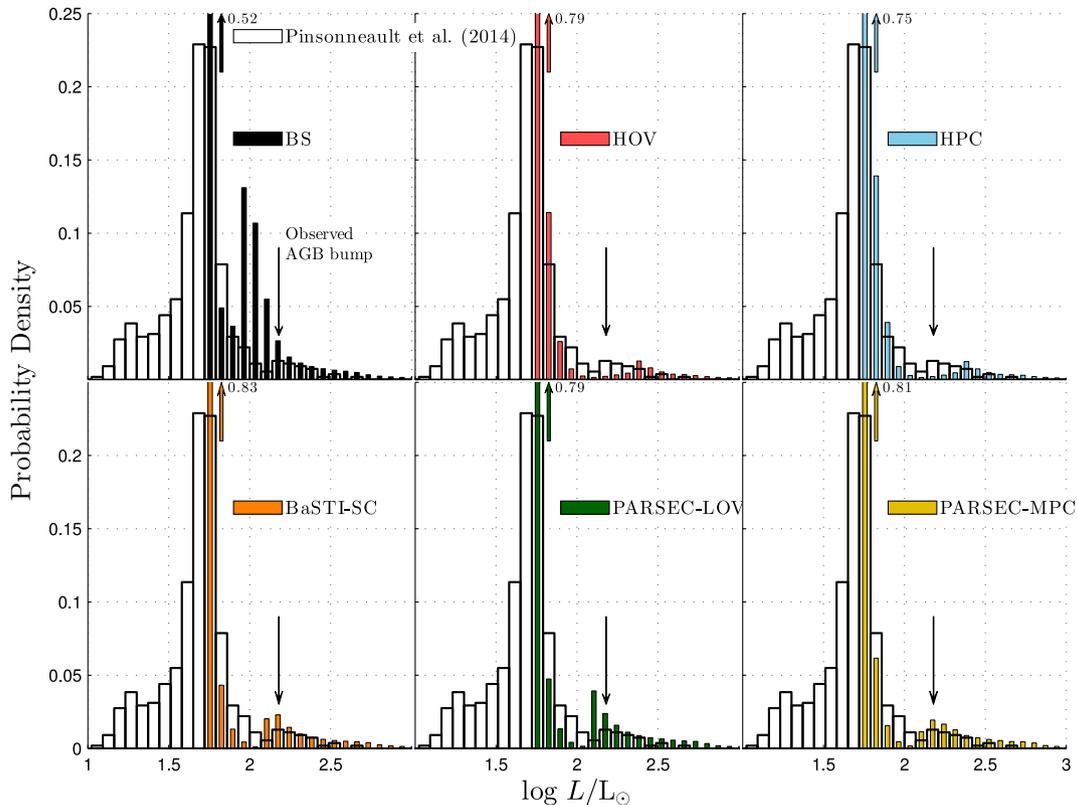
**Figure 6.** Hertzsprung–Russell diagram and luminosity distribution of the APOKASC catalogue (Pinsonneault et al. 2014). The sample is selected in a range of mass  $M = 1.3\text{--}1.7 M_{\odot}$  and metallicity  $[M/H] = -0.4\text{--}0.4$ . The AGBb can be identified in the right panel at around  $\log L/L_{\odot} \sim 2.2$  (black arrow) while the major peak corresponds to the RC.



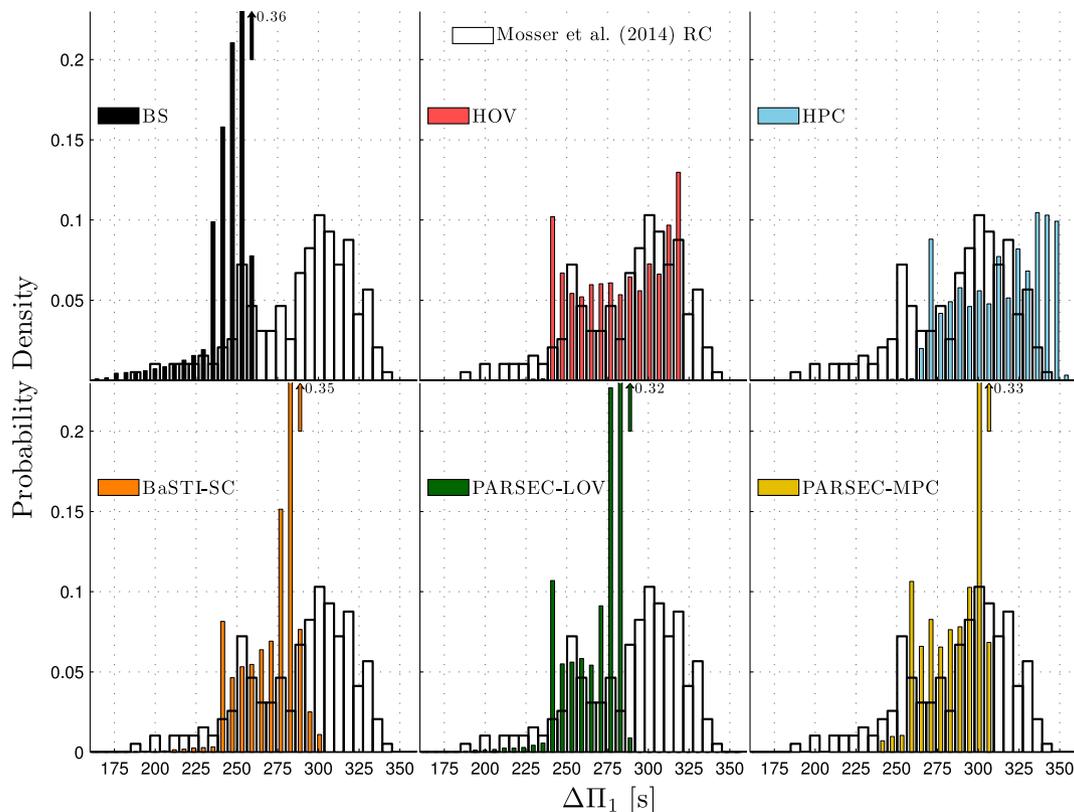
**Figure 7.** Hertzsprung–Russell diagram (left panel) and luminosity density distribution (stacked histogram, right panel) of stars in the TRILEGAL simulation of the *Kepler* field for the same masses and metallicities as in Fig. 6. The AGBb is visible in the right panel around  $\log L/L_{\odot} \sim 2.2$  (black arrow).

but BS, show the presence of this second structure. The reason behind it can be deduced simply following the theoretical tracks in the  $Y_c - \Delta\Pi_1$  diagram (see Fig. 3). The tracks cross three times a narrow interval of period spacing: the first two times at the very beginning of the HeCB phase, i.e. where  $\Delta\Pi_1$  decreases until a local minimum and then starts to rise, while the third passage happens during the rapid decrease that follows the absolute

maximum (at the end of the HeCB). Since time and  $Y_c$  are, in first approximation, linearly related during the HeCB, the third passage is very quick compared to the first and the second, therefore the peak is almost entirely populated by stars at the beginning of the phase. In the BS model the main and the second peak are not resolved since  $\Delta\Pi_1$  stays nearly constant for most of the HeCB phase.



**Figure 8.** Comparison between the observed luminosity density distribution of stars in APOKASC public catalogue selected with the same criteria as in Fig. 6, and the predictions from the models described in the paper. In the models, RGB stars are omitted.



**Figure 9.** Comparison between the observed period-spacing distribution of stars classified by Mosser et al. (2014) as RC and AGB and the predictions from the models described in the paper. We have selected stars with mass  $1.3 \leq M/M_{\odot} \leq 1.7$ .

At this stage we are, however, limited to a qualitative comparison between observed and theoretical distributions of  $\Delta\Pi_1$ . Such limitation arises from the fact that observational biases at the target selection stage, and in the determination of the  $\Delta\Pi_1$  from the power spectra, have not been fully explored yet. Moreover, although we have chosen models with a mass and metallicity representative of the stars observed by *Kepler* ( $[M/H] = -0.07 \pm 0.24$ ), and selected stars according to the mass, our synthetic population is rather simplistic. As a first test of the effect of changing the mass and metallicity, we consider HOV models with different mass ( $M = 1 M_{\odot}$ ) or metallicity ( $Z = 0.007$ ). We find that while the  $\Delta\Pi_1$  distribution depends little on the mass, the effect of reducing significantly the metallicity is to extend to the range of  $\Delta\Pi_1$  by  $\sim 20$  s. This effect needs to be taken into account when making quantitative comparisons between the observed and theoretically predicted distributions.

While these limitations can be partly mitigated by a more realistic synthetic population, and by a thorough examination of selection biases, it is likely that the most robust inference will be possible when applying our test to simple stellar populations, i.e. to red giants in the old-open clusters NGC 6791 and NGC 6819.

## 7 SUMMARY AND FUTURE PROSPECTS

The ability to accurately predict the properties of HeCB stars depends on our understanding of convection, which remains one of the key-open questions in stellar modelling (e.g. see Castellani et al. 1971a; Chiosi 2007; Salaris 2007; Bressan et al. 2015). Crucially, stringent tests of models have been limited so far by the lack of ob-

servational constraints specific to the internal structure of evolved stars.

In this work we propose a way forward. We argue that the combination of two observational constraints, i.e. the luminosity of the AGBb and the RC period spacing of gravity modes, provides us with a decisive test to discriminate between competing models of HeCB low-mass stars.

We have computed a series of stellar models with various prescriptions for the transport of chemicals and for the thermal stratification of near-core regions, and using different evolution codes (MESA, BaSTI, and PARSEC). First, we used these models to make predictions about the duration of the RC and eAGB phases, the luminosity of the AGBb, and C/O ratios in WDs. A summary of the model characteristics can be found in Table 1. We then focused on the prediction of seismic observables (see Section 5). We found that the asymptotic period spacing of gravity modes depends strongly on the prescription adopted (with differences up to about 40 per cent, when comparing the BS and HOV model).

We complemented this analysis by a numerical computation of adiabatic oscillation frequencies. This allowed us to confirm that the asymptotic approximation (equation 2) is a good representation of the period spacing of gravity-dominated modes. Moreover (see Section 5.3), the detailed behaviour of the period spacing of  $g$  modes shows the seismic signature of sharp variations in the Brunt–Väisälä frequency, which could potentially give additional information about near-core features (localized chemical composition gradients and near-discontinuities in the temperature gradient).

We then presented (Section 6) a first comparison between our predictions and the observational constraints obtained from the analysis of *Kepler* light curves (Mosser et al. 2014; Pinsonneault et al. 2014).

We found evidence for the AGBb among *Kepler* targets, which allowed us to make a first combined analysis of classical (AGBb luminosity) and seismic ( $RC \Delta\Pi_1$ ) constraints. Our main conclusion is that, while standard models (BaSTI-SC, PARSEC-LOV) are able to reproduce the luminosity of the AGBb, they cannot describe satisfactorily the distribution of the observed period spacing of RC stars while models with high overshooting (HOV), although giving a much better description of the observed  $RC \Delta\Pi_1$  distribution, fail to reproduce the AGBb luminosity. We then suggest a candidate model to describe simultaneously the two observed distributions: a model with a moderate overshooting region in which we apply an adiabatic thermal stratification. This prescription (which we have tested using PARSEC; see PARSEC-MPC model) gives indeed a better description of the observations.

At this stage of the analysis we are, however, prevented from drawing any further quantitative conclusions. To achieve the latter, we will follow two complementary approaches. On the one hand, we will couple our models with TRILEGAL, to generate synthetic stellar populations which can be quantitatively compared with the observed composite disc population. We will also investigate in detail possible observational biases, both in the target selection and in the detection of the period spacing from oscillations spectra. On the other hand, to limit/quantify such biases, we will test our models considering stars in the clusters NGC 6791 and NGC 6819, in which oscillations were detected in  $\sim 30$  HeCB giants (e.g. see Stello et al. 2011; Corsaro et al. 2012; Miglio et al. 2012).

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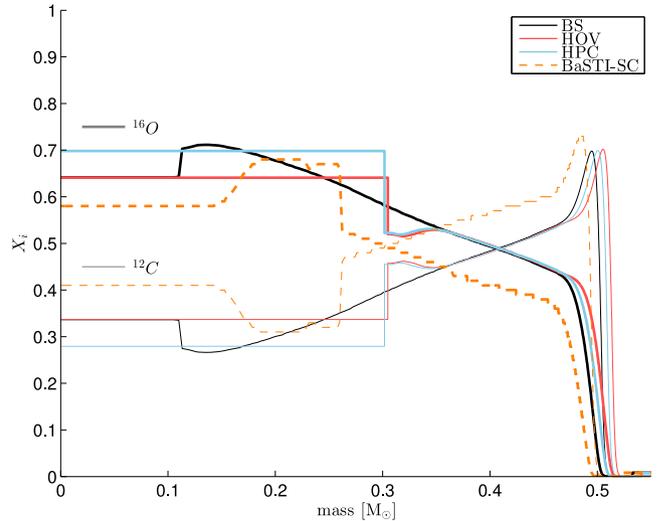
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## APPENDIX A: CARBON AND OXYGEN CONTENT

As mentioned in Section 2, an additional consequence of adopting different mixing schemes in the HeCB phase is the resulting chemical profile of the C-O core, which affects the chemical profile of a WD at the beginning of the cooling sequence.



**Figure A1.** Carbon and oxygen profile of  $1.5\text{-}M_{\odot}$  models at the first AGB-TP.

To provide the reader with a complete set of predictions, we show in Fig. A1 the chemical-composition profile of AGB models at the first thermal pulse. The results are consistent with what Straniero et al. (2003) found in models of a  $3\text{-}M_{\odot}$  star. They showed that the maximum of  $X_{\text{O}}$  for a BS and induced overshooting (plus semi-convection) models lies outside the former convective core while in models with high overshooting it corresponds to the value found in the former mixed core. Moreover, in our HPC model the  $X_{\text{O}}$  is higher than the HOV model, mostly due to the larger duration of the He-burning core phase (Fig. A1, left panel) in which  $^{12}\text{C}(\alpha,\gamma)^{16}\text{O}$  reaction has more time to convert carbon into oxygen.

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