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Neglecting primordial non-Gaussianity threatens future cosmological experiment accuracy

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Future galaxy redshift surveys aim at probing the clustering of the cosmic large-scale structure with unprecedented accuracy, thus complementing cosmic microwave background experiments in the quest to deliver the most precise and accurate picture ever of our Universe. Analyses of such measurements are usually performed within the context of the so-called vanilla Λ CDM model—the six-parameter phenomenological model which, for instance, emerges from best fits against the recent data obtained by the Planck satellite. Here, we show that such an approach is prone to subtle systematics when the Gaussianity of primordial fluctuations is concerned. In particular, we demonstrate that, if we neglect even a tiny amount of primordial non-Gaussianity—fully consistent with current limits—we shall introduce spurious biases in the reconstruction of cosmological parameters. This is a serious issue that must be properly accounted for in view of accurate (as well as precise) cosmology.

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I. INTRODUCTION

The currently accepted standard model for the formation of cosmic structure posits that the Universe underwent an early phase of accelerated expansion (dubbed “inflation” [1]), during which a random field of primordial density fluctuations emerged. Subsequent gravity-driven hierarchical growth of such density fluctuations led to the formation of galaxies, galaxy clusters, and the cosmic large-scale structure (LSS). As inflation is not a fundamental theory, different classes of inflationary models predict different statistical distributions for the primordial density fluctuations (see, e.g., Ref. [2] for a comprehensive review). Naturally, structures accreting from different initial conditions will have different statistical properties. The study of such properties constitutes one of the most powerful probes for understanding the physics of the (mostly unobservable) early Universe.

One of the most general ways to quantify the statistics of primordial density fluctuations is measuring their level of non-Gaussianity. Whilst the simplest slow-roll inflationary model predicts initial conditions that are almost perfectly Gaussian, the relaxation of specific assumptions gives rise to substantial and model-dependent deviations from Gaussianity. A particularly convenient (although not

unique) way to parametrize primordial non-Gaussianity (PNG) is to add a quadratic correction to the original Gaussian Bardeen potential field [3,4],

$$\Phi = \Phi_G + f_{\text{NL}} * (\Phi_G^2 - \langle \Phi_G^2 \rangle). \quad (1)$$

The quantity dubbed f_{NL} , which may be regarded as a free parameter, determines the amplitude of PNG. In the most general case, f_{NL} may depend on both time and scale, and thus we use the $*$ convolution symbol instead of ordinary multiplication; as is often done in the literature, for the sake of simplicity we here assume f_{NL} to be scale independent.

PNG has been studied extensively over the past decade, using data from both the cosmic microwave background (CMB) and the LSS. With respect to the latter, investigations have included cluster number counts, galaxy clustering, cosmic shear, LSS topology, and others (see, for example, Refs. [5–11] and references therein). Recently, analyses of Planck satellite data managed to severely constrain the allowed parameter space of PNG [12]. Henceforth, according to a number of studies, only certain future LSS experiments will be able to provide comparable constraints on f_{NL} , such as those using galaxy redshift surveys [11,13–15] or that are conducted in the radio continuum [16,17], or those using newborn techniques such as neutral hydrogen intensity mapping [18] or cross-correlation with other observables [19–21].

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Given that many LSS cosmological tests keep finding levels of PNG that are consistent with zero (though usually with large error bars; see, e.g., Ref. [22]) and the fact that confidence levels have been dramatically shrunk by Planck data, it is meaningful to ask if PNG could be altogether ignored without significantly affecting constraints on the other cosmological parameters. If this is the case, the analysis of future cosmological data will be significantly simplified. Conversely, PNG should be kept in mind in order to not bias future cosmological constraints. This is the very question we address in this work. Specifically, we investigate whether we would introduce a bias in the best-fit value of other cosmological parameters if we neglected PNG in a universe with a small but nonvanishing value of f_{NL} . Then, we compare this possible bias with the statistical uncertainties predicted for future LSS surveys. Throughout this paper we refer to a Class IV cosmological experiment, of which the most renowned representatives are, e.g., the Square Kilometre Array (SKA) [23] at radio wavelengths, the Dark Energy Survey (DES) [24], the forthcoming European Space Agency Euclid¹ satellite [25,26] and the Large Synoptic Survey Telescope (LSST) [27] at optical frequencies.

As a reference model, we assume a flat Λ CDM universe with total matter density (in units of the critical density) $\Omega_m = 0.315$, baryon fraction $\Omega_b = 0.0487$, dark energy equation-of-state Chevallier-Polarski-Linder parameters $w_0 = -1$ and $w_a = 0$ [28,29], and a dimensionless Hubble constant $h \equiv H_0/(100 \text{ km/s/Mpc}) = 0.673$; the primordial power spectrum is described by its scalar spectral index $n_s = 0.960$ and amplitude $A_s = 2.195 \times 10^{-9}$ [30]. We consider cosmological constraints as expected for galaxy cluster counts, clustering of galaxies, and galaxy clusters, as well as their combination.

II. METHODOLOGY

A. Modeling PNG corrections

The impact of deviations from Gaussianity on the abundance and clustering of the tracers of the underlying dark matter structure have been investigated by many authors, who obtained either theoretical, semianalytic, or fully numerical results. Here, we summarize the most relevant aspects and refer the interested reader to, e.g., Fedeli *et al.* [11] and references therein for additional details. PNG effects mainly concern the mass function and linear bias of dark matter halos. These modifications involve different integrals of the gravitational potential bispectrum, $B^\Phi(\mathbf{k}_1, \mathbf{k}_2, \mathbf{k}_3)$. The bispectrum amplitude depends on both the amplitude A_s of the gravitational potential power spectrum, $P^\Phi(\mathbf{k})$, and on f_{NL} so that

$$B^\Phi(\mathbf{k}_1, \mathbf{k}_2, \mathbf{k}_3) = f_{\text{NL}} A_s^2 \Gamma(\mathbf{k}_1, \mathbf{k}_2, \mathbf{k}_3). \quad (2)$$

The PNG shape is determined by the dependence of $\Gamma(\mathbf{k}_1, \mathbf{k}_2, \mathbf{k}_3)$ upon the three momenta.

Here, we investigate the effect of some bispectrum shapes. Besides the often used local-type PNG—whose bispectrum is maximized for squeezed configurations, where one of the three momenta is much smaller than the other two—we also consider “orthogonal” PNG, so called because its configuration is nearly orthogonal to the local and equilateral shapes (for a review see Ref. [31]). The former is known to have the heaviest impact on the clustering of the LSS, while the latter is nonetheless interesting because it presents degeneracies with other cosmological parameters that are different from those of all other PNG types.

There are a number of prescriptions in the literature for computing PNG deviations to the abundance of dark matter halos. Here, we follow LoVerde *et al.* [32], who used an Edgeworth expansion of the mass density field in order to derive a non-Gaussian generalization of the Press and Schechter [33] mass function, $n_{\text{PS}}(m, z)$. We define a correction factor

$$\mathcal{R}(m, z) = \frac{n_{\text{PS}}(m, z)}{n_{\text{PS}}^{(\text{G})}(m, z)} \quad (3)$$

by means of which one can translate any given Gaussian halo mass function $n^{(\text{G})}(m, z)$, computed according to one’s favorite recipe, to its non-Gaussian counterpart, i.e.,

$$n(m, z) = \mathcal{R}(m, z) n^{(\text{G})}(m, z). \quad (4)$$

In this work, we use a Gaussian Sheth and Tormen [34] mass function.

Moreover, if the initial conditions for structure formation are non-Gaussian, the linear halo bias acquires an additional scale dependence, which can be modeled as [35]

$$b(m, z, k) = b^{(\text{G})}(m, z) + \beta_R(k) \sigma_m^2 [b^{(\text{G})}(m, z) - 1]^2, \quad (5)$$

where $b^{(\text{G})}$ is the Gaussian linear halo bias of Sheth *et al.* [36]. The function $\beta_R(k)$ encodes all the scale dependence of the non-Gaussian halo bias at mass $m = m(R)$, and can be written as [37]

$$\beta_R(k) = \int_0^{+\infty} d\xi \frac{\xi^2 \mathcal{M}_R(\xi)}{8\pi^2 \sigma_m^2 \mathcal{M}_R(k)} \times \int_{-1}^1 d\cos\theta \frac{B^\Phi(\xi, \alpha, k)}{P^\Phi(k)} \mathcal{M}_R(\alpha), \quad (6)$$

with $\alpha = \mathbf{k} + \xi$, where θ is the angle between \mathbf{k} and ξ , σ_m^2 is the mass variance, and

$$\mathcal{M}_R(\mathbf{k}) = \frac{2T(k)W_R(k)k^2}{3H_0^2\Omega_m} \quad (7)$$

¹<http://euclid-ec.org/>.

relates the density fluctuation field smoothed on a scale R to the respective peculiar potential. $T(k)$ is the matter transfer function and $W_R(k)$ is the Fourier transform of a top-hat window function. In the case of a local bispectrum shape, it has been shown that the PNG scale dependence is $\beta_R(k) \propto k^{-2}$ at large scales. For other shapes, the dependence is usually weaker (see Refs. [11,13] for details).

In order to model the impact of PNG and other cosmological parameters on the assembly of the LSS, we make use of the well-established halo model [38,39]. It is a semianalytic framework based on the fundamental assumption that all the objects we are interested in are contained within bound dark matter halos, so that their clustering properties can be simply expressed as a superposition of the object distribution within individual halos and the mutual clustering properties of halos. In this framework, galaxies are distributed within dark matter halos according to some conditional probability distribution, $p(N_g|m)$. (Note that in general this probability distribution would depend also on redshift, whereas for simplicity we ignore this dependence, unless explicitly stated.) Its first and second statistical moments, $\langle N_g|m \rangle$ and $\langle N_g(N_g - 1)|m \rangle$, respectively, represent the average number of galaxies that reside within a dark matter halo of mass m and the variance of that average number. A similar reasoning applies to clusters, except that it is commonly assumed that only one cluster may occupy a given dark matter halo, so that

$$\langle N_c(N_c - 1)|m \rangle = 0, \quad (8)$$

$$\langle N_c|m \rangle = \Theta(m - m_c), \quad (9)$$

where $\Theta(x)$ is the Heaviside step function and m_c is some (possibly redshift-dependent) mass threshold.

Hereafter, we follow Fedeli *et al.* (Ref. [11], Sec. 4) and by means of the halo model and of the quantities introduced above we consistently construct three-dimensional power spectra $P^{XY}(k, z)$, where $X, Y = g, c$ for galaxies and galaxy clusters, respectively. Note that, thanks to this method, we also compute the cross-correlation power spectrum between galaxies and galaxy clusters. Eventually, we calculate number counts of galaxy clusters $N_c(z)$ as well.

Regarding the sources, we consider $H\alpha$ galaxies, which for instance will be selected by a Euclid-like experiment. These are preferentially going to be blue star-forming galaxies, and therefore we model the moments of the galaxy distribution within dark matter halos following semianalytic models of galaxy formation [39], which give

$$\langle N_g|m \rangle = N_{g,0} \Theta(m - m_0) \left(\frac{m}{m_1} \right)^\gamma, \quad (10)$$

where $N_{g,0} = 0.7$, $m_0 = 10^{11} h^{-1} M_\odot$, $m_1 = 4 \times 10^{12} h^{-1} M_\odot$, and $\gamma = 0.8 \Theta(m - m_0)$. Fedeli *et al.* [11] showed that this

choice of parameters produce an effective galaxy bias that is in fair agreement with predictions based on semianalytic galaxy formation models [40,41] for a Class IV survey like Euclid [42]. Moreover, we set

$$\langle N_g(N_g - 1)|m \rangle = \langle N_g|m \rangle^2 f(m), \quad (11)$$

where the function $f(m)$ represents the deviation of the galaxy distribution from a Poissonian and can be modeled as

$$f(m) = \begin{cases} \log^2 \sqrt{\frac{m}{m_0}} & \text{if } m \leq 10^{13} h^{-1} M_\odot, \\ 1 & \text{if } m > 10^{13} h^{-1} M_\odot. \end{cases} \quad (12)$$

Finally, we consider galaxy clusters that will be photometrically selected, and for this reason we choose the minimum cluster halo mass $m_c = m_c(z)$ using the Euclid Red Book photometric selection function [25].

III. RESULTS

A. PNG effects on galaxies and galaxy clusters

Figure 1 illustrates the auto- and cross-correlation power spectra of galaxies and galaxy clusters as a function of scale at $z = 1$ for three values of f_{NL} in the local-shape scenario. Solid curves are for $f_{\text{NL}}^{\text{loc}} = -2.17$. We choose this particular value because, as recently remarked by Camera *et al.* [14], in Λ CDM with slow-roll single-field inflation, galaxy surveys should measure $f_{\text{NL}} \approx -2.17$. This happens because there is a nonlinear general-relativistic correction

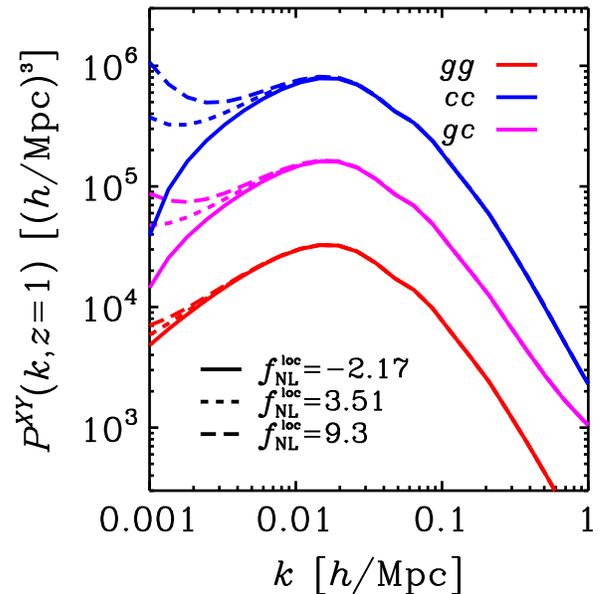


FIG. 1 (color online). Galaxy (red) and galaxy cluster (blue) power spectra and their cross-spectrum (magenta) at $z = 1$ with $f_{\text{NL}}^{\text{loc}} = -2.17, 3.51$ and 9.3 (solid, short- and long-dashed curves, respectively).

on very large scales which mimics a local PNG with $f_{\text{NL}} \approx -5/3$ [13,43]. This correction is derived in the CMB convention because it is based on the primordial Φ . It does not affect CMB measurements of PNG, but it must be added to the local PNG parameter for LSS. The translation from the CMB to the LSS convention (which we adopt here) sets $f_{\text{NL}}^{\text{LSS}} \approx 1.3f_{\text{NL}}^{\text{CMB}}$ (see, e.g., Ref. [11]), which eventually gives $f_{\text{NL}} \approx -2.17$. The other two sets of spectra are for $f_{\text{NL}}^{\text{loc}} = 3.51$ and 9.3 (short- and long-dashed curves, respectively). The former has been chosen because it is the Planck best-fit value (LSS convention) [12], while the latter better shows the PNG departure from the Gaussian prediction still lying within the Planck 1σ bound.

We can extract some useful information from Fig. 1. As expected, galaxy clusters (blue curves) are more biased than galaxies (red curves), so that their power spectrum is larger. Given this, the PNG correction—which is proportional to $[b^{(G)} - 1]^2 k^{-2}$ —kicks in at smaller scales (larger wave numbers) compared to galaxies, as can be seen by looking at the different behavior of the two curves at small k . Besides, we can notice that the cross spectrum between galaxies and galaxy clusters (magenta curve set) is not merely an average of the two progenitors' spectra. Indeed, the three spectra are characterized by different scale dependencies, which means that each observable carries a different piece of information about the clustering of the LSS. For example, there are different shapes at large k , whereby the galaxy 1-halo term carrying information on nonlinear scales starts to become important, but no 1-halo term is present in the cluster power spectrum (since it is commonly assumed that only one cluster is contained inside each dark matter halo).

Opposite to what happens to the power spectra, the effect of PNG on galaxy cluster counts is smaller, since it is integrated over mass and redshift. Therefore, to give a flavor of the non-Gaussian mass function, in Fig. 2 we plot the correction factor of Eq. (3), $\mathcal{R}(m, z)$, at $z = 1$ and for the same f_{NL} values as in Fig. 1.

B. Induced bias on cosmological parameters

To estimate the bias on a set of cosmological parameters $\{\vartheta_\alpha\}$ triggered by neglecting some amount of PNG in the data analysis phase, we follow the Bayesian approach of Heavens *et al.* [44], based on the Fisher information matrix [45]. The basic idea is that if we try to fit a model that does not correctly include all the relevant effects against actual data (PNG in this case), the model likelihood in parameter space will have to shift its peak in order to account for incorrect assumptions. In other words, the true parameter likelihood peaks at a certain point in the full parameter space spanned by $\{\vartheta_\alpha\} \cup \{f_{\text{NL}}\}$; by neglecting PNG, however, we actually look at the $f_{\text{NL}} = 0$ hypersurface, where the likelihood maximum will not in general correspond to its true value. The corresponding shift induced on the other model parameters—what we here call the bias

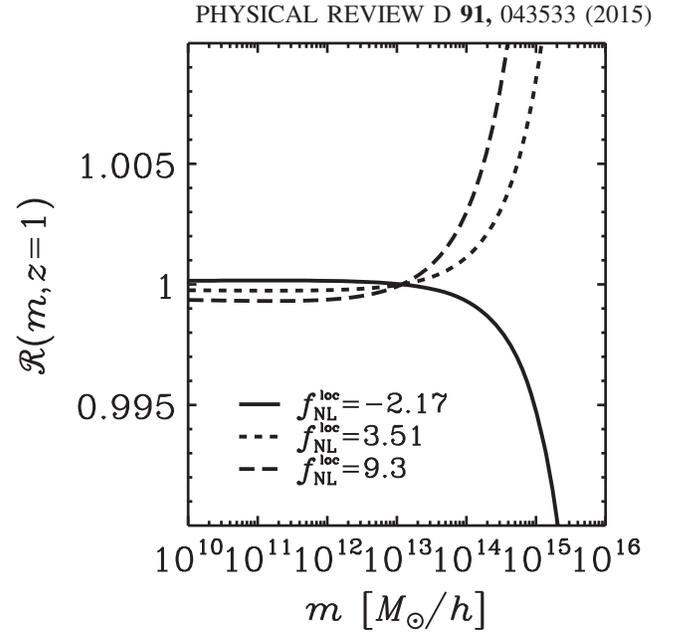


FIG. 2. Ratio of the non-Gaussian Schechter mass function to the Gaussian one as a function of mass at $z = 1$ with $f_{\text{NL}}^{\text{loc}} = -2.17, 3.51,$ and 9.3 (solid, short-, and long-dashed curves, respectively).

$b(\vartheta_\alpha)$ —is directly proportional to $\delta f_{\text{NL}} \equiv f_{\text{NL}}^{\text{true}} - 0$ and may be computed via

$$b(\vartheta_\alpha) = -(\mathbf{F}^{-1})_{\alpha\beta} \hat{\mathbf{F}}_{\beta f_{\text{NL}}} \delta f_{\text{NL}}, \quad (13)$$

where $\mathbf{F}_{\alpha\beta}$ is the Fisher matrix for the wrong parameter set, $\hat{\mathbf{F}}_{\alpha\beta}$ is the true Fisher matrix for the full parameter set

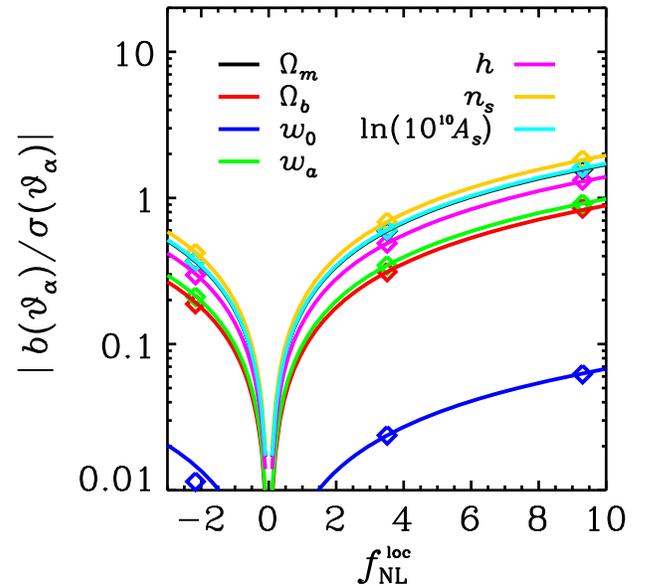


FIG. 3 (color online). Bias on cosmological parameters (in units of the error on the corresponding parameter) induced by neglecting PNG versus $f_{\text{NL}}^{\text{loc}}$ when Fisher matrices for all probes are considered.

TABLE I. Forecast marginal errors σ , correlation parameters r , and (normalized) biases of Λ CDM cosmological parameters ϑ for $f_{\text{NL}}^{\text{loc}} = 9.3$.

Local-type PNG ($f_{\text{NL}}^{\text{loc}} = 9.3$)												
ϑ	TOT			TOT + Planck			PSTOT			PSTOT + Planck		
	σ	r	b/σ	σ	r	b/σ	σ	r	b/σ	σ	r	b/σ
Ω_m	1.2×10^{-3}	5.5×10^{-1}	1.6	9.6×10^{-4}	4.5×10^{-1}	1.5	1.4×10^{-3}	$6. \times 10^{-1}$	9.6×10^{-1}	1.2×10^{-3}	6.7×10^{-1}	1.3
Ω_b	$8. \times 10^{-4}$	$3. \times 10^{-1}$	8.5×10^{-1}	2.7×10^{-4}	$-2. \times 10^{-1}$	-6.8×10^{-1}	8.6×10^{-4}	2.7×10^{-1}	4.4×10^{-1}	2.9×10^{-4}	-4.1×10^{-1}	-8.1×10^{-1}
w_0	2.1×10^{-2}	2.2×10^{-2}	6.2×10^{-2}	7.7×10^{-3}	-4.3×10^{-1}	-1.4	$3. \times 10^{-2}$	-5.4×10^{-1}	-8.6×10^{-1}	9.5×10^{-3}	-6.8×10^{-1}	-1.3
w_a	$1. \times 10^{-1}$	3.2×10^{-1}	9.2×10^{-1}	2.2×10^{-2}	9.6×10^{-2}	3.2×10^{-1}	1.3×10^{-1}	-2.3×10^{-1}	-3.7×10^{-1}	2.2×10^{-2}	-1.1×10^{-2}	-2.2×10^{-2}
h	8.9×10^{-3}	4.6×10^{-1}	1.3	1.9×10^{-3}	1.6×10^{-1}	5.4×10^{-1}	$1. \times 10^{-2}$	4.4×10^{-1}	$7. \times 10^{-1}$	1.9×10^{-3}	1.4×10^{-1}	2.8×10^{-1}
$\ln(10^{10}A_s)$	2.6×10^{-2}	-5.7×10^{-1}	-1.6	7.9×10^{-3}	-3.3×10^{-1}	-1.1	$3. \times 10^{-2}$	-1.8×10^{-1}	-2.8×10^{-1}	$8. \times 10^{-3}$	-2.1×10^{-1}	-4.2×10^{-1}
n_s	5.7×10^{-3}	-6.4×10^{-1}	-1.8	2.7×10^{-3}	-4.7×10^{-1}	-1.6	6.7×10^{-3}	$-6. \times 10^{-1}$	-9.7×10^{-1}	3.1×10^{-3}	-6.4×10^{-1}	-1.2

TABLE II. Same as Table I for $f_{\text{NL}}^{\text{ort}} = -65$.

Orthogonal-type PNG ($f_{\text{NL}}^{\text{ort}} = -65$)												
ϑ	TOT			TOT + Planck			PSTOT			PSTOT + Planck		
	σ	r	b/σ									
Ω_m	1.2×10^{-3}	-5.6×10^{-1}	2.2	$6. \times 10^{-4}$	-2.2×10^{-1}	1.1	1.4×10^{-3}	-6.3×10^{-1}	1.3	6.4×10^{-4}	-4.1×10^{-1}	1.1
Ω_b	$8. \times 10^{-4}$	-2.8×10^{-1}	1.1	2.3×10^{-4}	-4.6×10^{-2}	2.3×10^{-1}	8.6×10^{-4}	-2.8×10^{-1}	$6. \times 10^{-1}$	2.3×10^{-4}	3.2×10^{-2}	-8.9×10^{-2}
w_0	2.1×10^{-2}	1.1×10^{-1}	-4.2×10^{-1}	1.8×10^{-2}	$-1. \times 10^{-1}$	5.1×10^{-1}	2.8×10^{-2}	-4.5×10^{-1}	9.8×10^{-1}	2.4×10^{-2}	$-6. \times 10^{-1}$	1.7
w_a	$1. \times 10^{-1}$	-3.8×10^{-1}	1.5	7.2×10^{-2}	-9.2×10^{-4}	4.6×10^{-3}	1.3×10^{-1}	1.2×10^{-1}	-2.5×10^{-1}	9.3×10^{-2}	5.1×10^{-1}	-1.4
h	8.9×10^{-3}	-4.5×10^{-1}	1.8	1.8×10^{-3}	-2.8×10^{-1}	1.4	$1. \times 10^{-2}$	-4.5×10^{-1}	9.7×10^{-1}	1.9×10^{-3}	-3.6×10^{-1}	9.9×10^{-1}
$\ln(10^{10}A_s)$	2.6×10^{-2}	5.8×10^{-1}	-2.3	1.1×10^{-2}	1.4×10^{-1}	-6.9×10^{-1}	3.1×10^{-2}	2.5×10^{-1}	-5.4×10^{-1}	1.4×10^{-2}	-4.1×10^{-1}	1.1
n_s	5.7×10^{-3}	6.3×10^{-1}	-2.4	$2. \times 10^{-3}$	2.6×10^{-1}	-1.3	6.7×10^{-3}	6.1×10^{-1}	-1.3	2.1×10^{-3}	3.6×10^{-1}	-9.9×10^{-1}

(including f_{NL}), and $\hat{\mathbf{F}}_{\beta f_{\text{NL}}}$ is a vector corresponding to the $\hat{\mathbf{F}}$ matrix line/column relative to f_{NL} .

Details on the Fisher information matrices for galaxy and cluster power spectra and their mutual cross spectrum can be found in Refs. [11,46]. In the following analysis, we consider ten redshift bins of width 0.1 centred from $z = 0.9$ to 1.8. We hold $k_{\text{max}} = 0.3h \text{ Mpc}^{-1}$ fixed to avoid the strongly nonlinear regime, while the smallest wave number, k_{min} , matches the largest available scale in a given redshift bin. [Regarding this point, we notice that general-relativistic corrections on very large scales may affect the results (see, e.g., Refs. [14,47]), but it has been shown that for future surveys (for instance, Euclid) their effect should be negligible [48].] Figure 3 shows $|b(\vartheta_\alpha)/\sigma(\vartheta_\alpha)|$ for the case where we sum the Fisher matrices for all probes, i.e., for galaxies, clusters, and their cross spectrum and cluster number counts. We present the bias in units of the forecast marginal error on the corresponding parameter,

$$\sigma(\vartheta_\alpha) = \sqrt{(\mathbf{F}^{-1})_{\alpha\alpha}}, \quad (14)$$

which allows us to better assess the impact that such a bias will imply. The parameter set which we allow to vary is the full ΛCDM set $\{\vartheta_\alpha\} = \{\Omega_m, \Omega_b, w_0, w_a, h, \ln(10^{10}A_s), n_s\}$, in addition to the PNG parameter f_{NL} . Data points refer to sampled f_{NL} values, whereas solid curves come from the analytic expression of Eq. (13). Clearly, there are parameters that are prone to having their “best-fit” value shifted if we estimate them within the wrong theoretical framework.

This is the case for Ω_m , n_s , and A_s , which are known to be more degenerate with f_{NL} (see the correlation coefficients in Tables I and II and the related discussion). On the other hand, our analysis is in good agreement with the literature, as we do not observe a significant dependence of $\sigma(\vartheta_\alpha)$ on the assumed f_{NL} fiducial value (see, e.g., Refs. [49,50]).

Ultimately, this means that a blatant disregard for PNG (which is somewhat understandable given the stringent Planck limits) may threaten the constraining power of future surveys—if not by worsening their precision, then by undermining their accuracy. Surely, a Fisher-matrix approach does not fully capture the likelihood properties on the whole parameter space. Nevertheless, we want to emphasize that our analysis by no means refers to some extreme situation. On the contrary, the f_{NL} fiducial values here considered are well within Planck 2σ constraints for local-type PNG.

To further stress this point, in Fig. 4 we present 1σ joint marginal contours in the $[\Omega_m, \ln(10^{10}A_s)]$ and (w_0, w_a) planes (left and right panels, respectively), with solid lines for the true error contours from $\hat{\mathbf{F}}$ and dashed contours for when f_{NL} is neglected in the Fisher analysis. This is done for local-type PNG with $f_{\text{NL}}^{\text{loc}} = 9.3$ (left plots in both panels) and for orthogonal-type PNG with $f_{\text{NL}}^{\text{ort}} = -65$ (right plots in both panels), when we consider Fisher matrices for all probes (“TOT,” top plots) or only for the combination of the three auto and cross spectra (“PSTOT,” bottom plots). Black and red ellipses refer to forecasts either ignoring or including current Planck constraints, respectively. It is clear that in both cases, and for all the

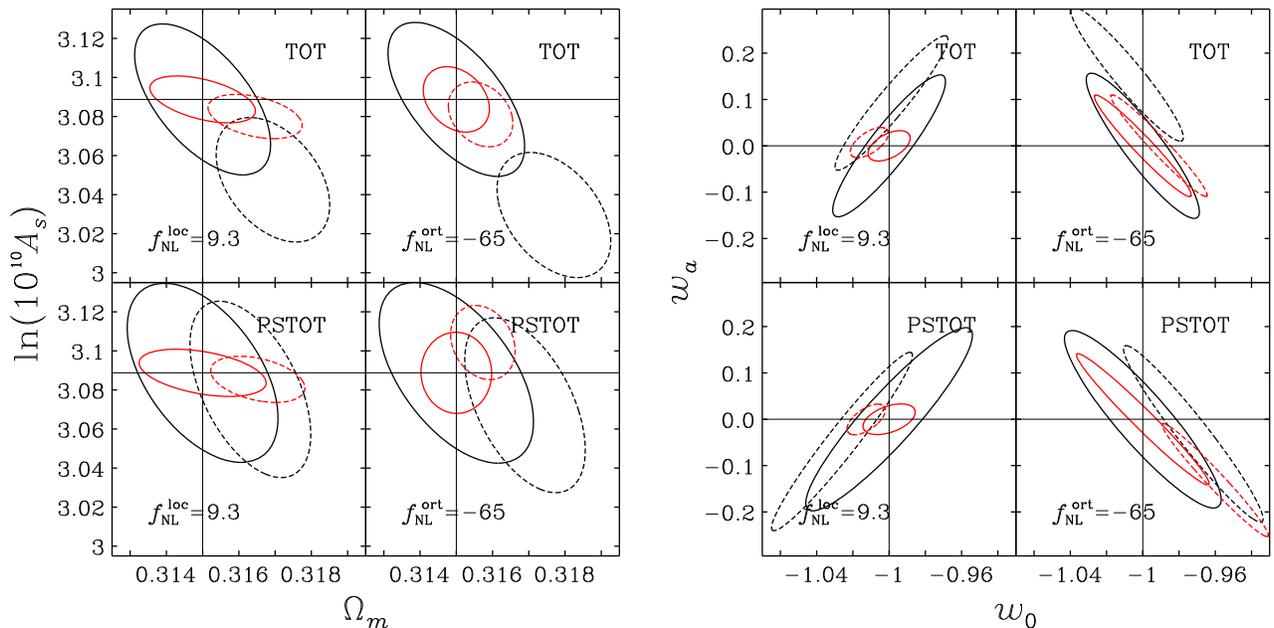


FIG. 4 (color online). Forecast 1σ joint marginal contours in the $[\Omega_m, \ln(10^{10}A_s)]$ and (w_0, w_a) planes (left and right panels, respectively) for local- and orthogonal-type PNG with $f_{\text{NL}}^{\text{loc}} = 9.3$ and $f_{\text{NL}}^{\text{ort}} = -65$. Solid ellipses are the true error contours, while dashed ellipses come from neglecting f_{NL} in the analysis. Black and red colors refer to results without and with Planck priors on ΛCDM parameters, respectively. Top (bottom) panels are for TOT (PSTOT).

configurations and PNG choices considered in this work, some non-negligible shift occurs. A more quantitative insight can be drawn from Tables I and II, where relevant quantities of the Λ CDM cosmological parameters are given, such as forecast marginal errors $\sigma(\vartheta_\alpha)$, f_{NL} correlation parameters

$$r(\vartheta_\alpha, f_{\text{NL}}) = \frac{(\mathbf{F}^{-1})_{\alpha f_{\text{NL}}}}{\sqrt{(\mathbf{F}^{-1})_{\alpha\alpha}(\mathbf{F}^{-1})_{f_{\text{NL}}f_{\text{NL}}}}}, \quad (15)$$

and normalized biases $b(\vartheta_\alpha)/\sigma(\vartheta_\alpha)$.

A major point emerging from this analysis is that—even though orthogonal-type PNG deviations from the Gaussian prediction have a much smaller impact upon the clustering of the LSS compared to PNG with local shape—current constraints on orthogonal PNG are consequently looser. In particular, Planck data [12] agrees with $f_{\text{NL}}^{\text{ort}} = -33 \pm 51$ (LSS convention). That is to say, the value $f_{\text{NL}}^{\text{ort}} = -65$ we assume here is well within Planck 1σ bounds. Nonetheless, if it were the true value and we neglected it, we would miss the true likelihood peak by more than 1σ , which is intolerable for the aims of future cosmological experiments.

IV. CONCLUSIONS

In this paper, we investigated the impact of neglecting PNG when performing parameter reconstruction for an idealized representative of next-generation Class IV cosmological experiments. Specifically, we considered a spectroscopic galaxy redshift survey along the lines of the European Space Agency Euclid satellite. This allowed us to compute galaxy and galaxy cluster three-dimensional power spectra, as well as their cross spectrum and galaxy cluster number counts, in a fully consistent way within the halo model framework.

Hence, we estimated the bias on the reconstruction of standard Λ CDM cosmological parameters induced by disregarding PNG in the analysis. This was done in a Bayesian Fisher-matrix perspective, by considering the Λ CDM vanilla cosmological model as a subspace (in parameter space) of a Λ CDM universe with PNG. In other words, we recover the concordance cosmological model if we restrict the parameter space to the $f_{\text{NL}} = 0$ hypersurface. By doing so, the peak of the parameter likelihood on the hypersurface does not (in general) correspond to its true peak in the full parameter space if f_{NL} is nonzero and it is

not completely uncorrelated to the other cosmological parameters.

Our major results are summarized in Tables I and II and in Fig. 4. In particular, we found that an incorrect treatment of PNG in the data analysis will undermine the experimental accuracy of the reconstruction of some cosmological parameters. For example, the best-fit value of the dark energy parameters w_0 will be biased by more than one standard deviation if local-type PNG is in fact present with a value of f_{NL} consistent with 1σ Planck constraints. This is mainly due to the high precision of oncoming surveys, which will provide us with very tight constraints on the Λ CDM model parameters. Indeed, if on the one hand their expected allowed regions in parameter space will only slightly shrink by neglecting f_{NL} in the analysis (as known in the literature), on the other hand the small but non-negligible degeneracy with f_{NL} will cause a shift of their reconstructed best-fit values. To avoid this, it appears clear that PNG has to be consistently accounted for.

Finally, we emphasize that—although we adopt the specifics of a Euclid-like survey as a reference experiment—our findings should be regarded as potential systematics for the whole class of future, high-precision galaxy surveys, such as DES, LSST, and the SKA.

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