



Publication Year	2016
Acceptance in OA@INAF	2020-04-29T14:39:27Z
Title	The PEP survey: evidence for intense star-forming activity in the majority of radio-selected AGN at $z > 1$
Authors	MAGLIOCCHETTI, MANUELA; Lutz, D.; SANTINI, Paola; Salvato, M.; Popesso, P.; et al.
DOI	10.1093/mnras/stv2645
Handle	http://hdl.handle.net/20.500.12386/24334
Journal	MONTHLY NOTICES OF THE ROYAL ASTRONOMICAL SOCIETY
Number	456

The PEP survey: evidence for intense star-forming activity in the majority of radio-selected AGN at $z \gtrsim 1$

M. Magliocchetti,¹★ D. Lutz,²★ P. Santini,³ M. Salvato,² P. Popesso,⁴ S. Berta² and F. Pozzi⁵

¹INAF-IAPS, Via Fosso del Cavaliere 100, I-00133 Roma, Italy

²Max Planck Institut für extraterrestrische Physik (MPE), Postfach 1312, D-85741, Garching, Germany

³INAF-Osservatorio Astronomico di Roma, Via di Frascati 33, I-00040 Monte Porzio Catone, Italy

⁴Excellence Cluster, Boltzmannstr. 2, D-85748, Garching, Germany

⁵Dipartimento di Astronomia, Università di Bologna, Via Ranzani 1, I-40127, Bologna, Italy

Accepted 2015 November 9. Received 2015 November 9; in original form 2015 July 29

ABSTRACT

In order to investigate the far-infrared (FIR) properties of radio-active active galactic nuclei (AGN), we have considered three different fields where both radio and FIR observations are the deepest to date: GOODS-South, GOODS-North and the Lockman Hole. Out of a total of 92 radio-selected AGN, ~ 64 per cent are found to have a counterpart in *Herschel* maps. The percentage is maximum in the GOODS-North (72 per cent) and minimum (~ 50 per cent) in the Lockman Hole, where FIR observations are shallower. Our study shows that in all cases FIR emission is associated with star-forming activity within the host galaxy. Such an activity can even be extremely intense, with star-forming rates as high as $\sim 10^3$ – 10^4 M_{\odot} yr $^{-1}$. AGN activity does not inhibit star formation in the host galaxy, just as on-site star formation does not seem to affect AGN properties, at least those detected at radio wavelengths and for $z \gtrsim 1$. Given the very high rate of FIR detections, we stress that this refers to the majority of the sample: most radio-active AGN are associated with intense episodes of star formation. However, the two processes proceed independently within the same galaxy, at all redshifts but in the local universe, where powerful enough radio activity reaches the necessary strength to switch off the on-site star formation. Our data also show that for $z \gtrsim 1$ the hosts of radio-selected star-forming galaxies and AGN are indistinguishable from each other in terms of both mass and IR luminosity distributions. The two populations only differentiate in the very local universe, whereby the few AGN which are still FIR-active are found in galaxies with much higher masses and luminosities.

Key words: methods: observational – galaxies: active – galaxies: evolution – galaxies: starburst – infrared: galaxies – radio continuum: galaxies.

1 INTRODUCTION

Radio sources constitute an unequally powerful tool to investigate the origin and evolution of our Universe. In fact, radio emission is not obscured or attenuated by dust or gas, so deep radio surveys offer a unique opportunity to detect and study sources up to the highest redshifts. It is following this rationale that very ambitious programmes like the planned Square Kilometre Array (SKA; Carilli et al. 2004) facility, a gigantic radio telescope with a total collective area of 1 km², or its precursors ASKAP (Australian SKA pathfinder; Johnston et al. 2007) and MeerKAT (Jonas 2009), will soon see their first light. When applied to radio sources, the expected advantages in

sensitivity, field of view, frequency range and spectral resolution will yield substantial progress in many research fields, from cosmology to astrophysics.

Extragalactic radio sources are, however, a mixed bag of astrophysical objects: amongst the most powerful ones we find radio-loud QSOs and FRII (Fanaroff & Riley 1974) galaxies while, moving to weaker sources, the dominant populations become FRI, low-excitation galaxies and star-forming galaxies, whose contribution to the total radio counts become predominant at the sub-mJy level (see e.g. Magliocchetti et al. 2000; Prandoni et al. 2001). Discerning amongst these objects and their relative weight is an impossible task to carry out only on the basis of radio-continuum data. This is why increasingly more effort was recently put on photometric and spectroscopic follow-up studies of radio sources detected in deep-field surveys (amongst the many: Schinnerer et al. 2004, 2007,

* E-mail: manuela@iaps.inaf.it (MM); lutz@mpe.mpg.de (DL)

2010, for the COSMOS field; Morrison et al. 2010, for GOODS-N; McAlpine, Jarvis & Bonfield 2013, for the VIDEO-XMM3 field; Bondi et al. 2003, for the VVDS field; Simpson et al. 2012, for the *Subaru/XMM-Newton* Deep Field; Mao et al. 2011, for the Extended *Chandra* Deep Field). Investigations of the optical properties of radio sources can then disentangle the different populations and also allow the determination of a number of quantities which influence the source behaviour like e.g. the black hole mass in optical spectra dominated by broad emission lines (e.g. Metcalf & Magliocchetti 2006).

Following the launch of the *Spitzer* satellite, in recent years radio sources have also been investigated at mid-infrared (MIR) wavelengths (e.g. Appleton et al. 2004; Boyle et al. 2007; Magliocchetti, Andreani & Zwaan 2008; Garn et al. 2009; Leipski et al. 2009; De Breuck et al. 2010; Norris et al. 2011 amongst the many). These studies proved to be extremely useful for two main reasons: they allowed investigations of the properties and evolution of the sub-population of star-forming galaxies up to $z \sim 2$, while on the other hand they provided interesting hints on the central engine responsible for radio (and MIR) active galactic nuclei (AGN) emission.

It is only in these very last few years that the radio community has started to realize the importance of multi-wavelength studies which also include far-infrared (FIR) information. This growing interest is primarily due to the advent and launch of the *Herschel* satellite, which for the first time has observed galaxies at FIR wavelengths up to very large cosmological ($z \sim 4$) distances (Lutz 2014). Earlier FIR missions such as *IRAS*, ISO and the 70 μm channel of the *Spitzer* satellite only probed the relatively local ($z \lesssim 1$) universe, and since AGN-powered radio galaxies in the local universe are almost always hosted by massive red galaxies with little or no ongoing star formation activity (e.g. Magliocchetti et al. 2002, 2004), and since FIR emission in galaxies is almost entirely due to processes connected with ongoing star formation, it follows that any pre-*Herschel* interest in a joint radio-FIR analysis was limited to reassessing the precision of the tight radio–FIR correlation found by early *IRAS*-based studies of local star-forming galaxies (Condon et al. 1982).

However, when moving to higher redshifts, galaxies undergo dramatic changes. For instance, it is now well established that both the cosmic star-forming activity and AGN output steadily increase with look-back time at least up to redshifts $z \sim 2$ (e.g. Merlini, Rudnick & Di Matteo 2004; Gruppioni et al. 2013), and many galaxies are observed to host both an AGN and ongoing star formation (e.g. Alexander et al. 2008). It is then legitimate to wonder whether radio galaxies also follow such a trend and if in the early universe they are also associated with ongoing stellar production. First attempts in this direction were presented by the work of Seymour et al. (2011) which analyses the FIR emission as observed by the SPIRE instrument (Griffin et al. 2010) on board of the *Herschel* satellite (Pilbratt et al. 2010) for a sample of very powerful, $L_{1.4\text{GHz}} \geq 10^{25} \text{ W Hz}^{-1}$ radio sources selected in the extragalactic *Spitzer* First Look Survey field, and by the work of Del Moro et al. (2013) who investigate hidden AGN activity in a sample of GOODS-North sources with deep radio, infrared and X-ray data.

Magliocchetti et al. (2014) pushed the Seymour et al. (2011) results further and investigated, for the first time, the FIR properties of radio-selected AGN of *all* radio luminosities and at *all* redshifts $z \lesssim 3.5$. This was done by using the very deep radio catalogue obtained at 1.4 GHz in the COSMOS field by Schinnerer et al. (2004, 2007, 2010) and Bondi et al. (2008). AGN were selected solely on the basis of their radio-luminosity and FIR information came from the Photoconductor Array Camera and Spectrometer

(PACS) Evolutionary Probe (PEP) survey (Lutz et al. 2011). Results from this analysis indicate that the probability for a radio-selected AGN to be detected at FIR wavelengths is both a function of radio power and redshift. Powerful sources were found more likely to be FIR emitters at earlier epochs because of two distinct effects: (1) at all radio luminosities, FIR activity monotonically increases with look-back time and (2) in the earlier universe, radio activity of AGN origin is increasingly less effective at inhibiting FIR emission. Magliocchetti et al. (2014) also found that the observed FIR (and MIR) emission was due to processes which were indistinguishable from those which power star-forming galaxies.

The above results, although interesting and obtained on a relatively large area, however need the complementarity of deep enough FIR observations which can match the depth of the radio data. As a matter of fact, only ~ 36 per cent of the sources identified in the COSMOS field as radio-emitting AGN possessed a counterpart in the PEP catalogues either at 100 μm or at 160 μm . The present work then overcomes the above limitations and analyses the FIR properties of radio-selected AGN on three fields where FIR observations are the deepest-to-date: GOODS-North, GOODS-South and the Lockman Hole (LH). FIR data come from both the PEP survey (Lutz et al. 2011) and the *Herschel*-GOODS-Survey (Elbaz et al. 2011) and, in the best-observed areas such as the two GOODS fields, ensure limiting FIR fluxes which are deeper by as much as a factor of ~ 6 at 100 μm and ~ 5 at 160 μm than the PEP observations on COSMOS. Throughout the work we will assume a Λ cold dark matter cosmology with $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ ($h = 0.7$), $\Omega_0 = 0.3$, $\Omega_\Lambda = 0.7$ and $\sigma_8^m = 0.8$. In all cases, masses for the sources under examination are calculated by using the Bruzual & Charlot (2003) templates and adopting a Salpeter (1955) initial mass function (IMF).

2 THE FIELDS

2.1 Lockman Hole

The radio sources used in this work have been taken from Ibar et al. (2009). For consistency with the other surveys considered in this paper, we focused on the catalogue of objects observed with the Very Large Array (VLA) at 1.4 GHz. The sensitivity of these observations was $\sigma \sim 6 \mu\text{Jy beam}^{-1}$, with a resolution (full width at half maximum – FWHM) of ~ 5 arcsec. Ibar et al. (2009) provide a catalogue of 1303 unique sources (1346 including sub-components of composite radio objects), selected to have a peak-to-local noise (PNR) greater than 5. As these objects were also observed at 610 MHz with the Giant Metrewave Radio Telescope (GMRT), the overwhelming majority of them is also endowed with an estimate of the radio spectral index $\alpha_{1400 - 610}$ (hereafter just α for clarity, where α is defined as $F_v \propto v^{-\alpha}$, with F -radio flux). Ibar et al. (2009) also provide a classification based on the radio source appearance, and distinguish them into single (S), double (D), triple (T) and Extended (M), plus some more complicated morphological cases such as single source with one close neighbour (SD), single source with multiple close neighbours (ST) and so on. Fig. 1 shows the radio sources as taken from the Ibar et al. (2009) 1.4 GHz selection as (blue) crosses.

Optical, near-IR and mid-IR information for these sources, as well as their redshift determination, can be obtained from the work of Fotopoulou et al. (2012) which provides very precise broad-band photometry and photometric redshifts for about 188 000 sources located in $\sim 0.5 \text{ deg}^2$ of the LH area (green small dots in Fig. 1). Special care was taken for objects identified as AGN. Note that the

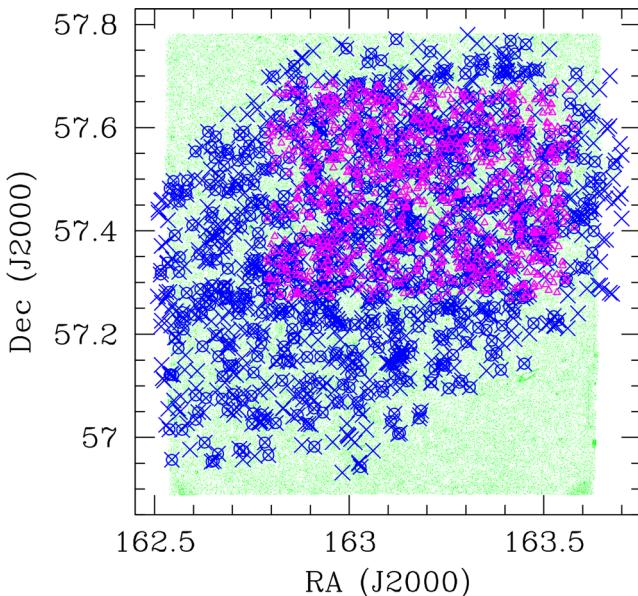


Figure 1. Lockman Hole. (Blue) crosses represent radio sources selected at 1.4 GHz by Ibar et al. (2009), (green) small dots show the area covered by optical and near-IR surveys, while (magenta) open triangles illustrate the sources imaged by the PEP survey. Crosses surrounded by open circles indicate radio sources with a redshift (either spectroscopic or photometric) determination from the catalogue of Fotopoulou et al. (2012).

Fotopoulou et al. (2012) catalogue also includes 388 X-ray detected sources (the overwhelming majority of which are AGN) identified in the deep *XMM-Newton* observations by Brunner et al. (2008), available for an area of 0.2 deg^2 .

Optical identifications for the Ibar et al. (2009) radio sources were then obtained from the Fotopoulou et al. (2012) catalogue for a maximum tolerance radius of 0.9 arcsec. This value was chosen to maximize the number of real associations, while reducing the percentage of spurious matches to roughly 6 per cent. The total number of radio sources on the area covered by the broad-band observations of Fotopoulou et al. (2012) is 1309. The number of sources with a photometric ID obtained in the way described above is 723, corresponding to 55.7 per cent of the original parent sample.

The fraction of radio sources in the LH endowed with either a spectroscopic or photometric redshift determination as a function of integrated 1.4 GHz radio flux is shown in Fig. 2 by the red circles, where the associated error bars correspond to Poissonian estimates. From the figure, one can notice that such a fraction is constant within the errors, and equal to approximately 55 per cent of the parent sample throughout the whole flux range probed by the Ibar et al. (2009) survey. This indicates that no radio-flux bias affects the optical identifications of the radio sources under consideration.

A smaller area, corresponding to roughly $24 \times 24 \text{ arcmin}^2$, has also been observed by the PACS instrument (Poglitsch et al. 2010) onboard of the *Herschel* satellite (Pilbratt et al. 2010) as a part of the PEP (Lutz et al. 2011) Survey, aimed at studying the properties and cosmological evolution of the infrared population up to redshifts $z \sim 3\text{--}4$. 1328 sources were observed in this field at either 100 or 160 μm by the PEP survey and the blind (i.e. source extraction performed without relying on priors) catalogues include objects down to a 3σ confidence level, respectively, of ~ 3.6 and $\sim 7.5 \text{ mJy}$. These are shown in Fig. 1 by the (magenta) open triangles. 698 radio sources from the Ibar et al. (2009) catalogue reside within the area

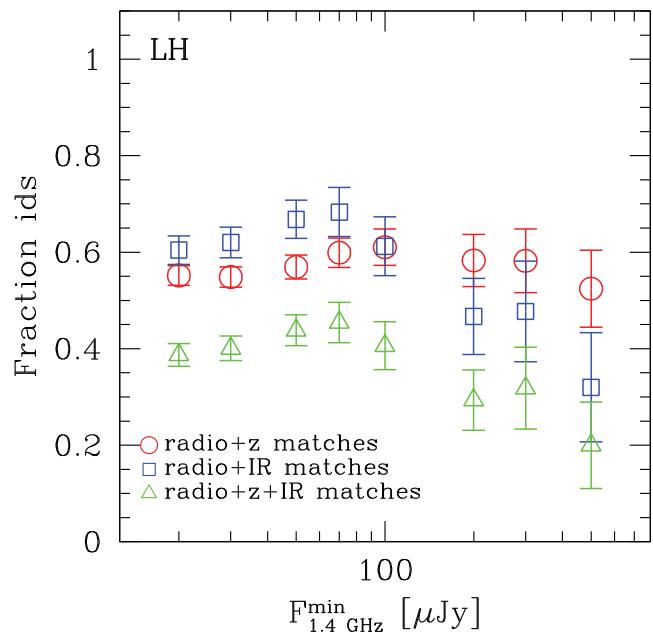


Figure 2. Fraction of radio sources in the LH endowed with a redshift determination (red circles), with an FIR counterpart from the PEP catalogue (blue squares) and with both an FIR ID and a redshift determination (green triangles) as a function of integrated 1.4 GHz flux. Error bars correspond to Poissonian estimates.

covered by the PEP observations, corresponding to 53.8 per cent of the original sample.

In order to provide the radio sources with an FIR counterpart, we followed the approach presented in Magliocchetti et al. (2014) and cross-matched the radio and FIR catalogues with a maximum coherence radius of 4 arcsec. This choice ensures a negligible contribution from spurious associations (~ 4 per cent). The number of radio sources with an FIR counterpart at either 100 μm or 160 μm is 422, corresponding to 60 per cent of the parent radio catalogue. The blue squares in Fig. 2 represent the trend for FIR identifications as a function of 1.4 GHz radio flux. This percentage is roughly constant up to about 100 μJy and then slowly starts declining beyond that value. The green triangles in Fig. 2 represent the fraction of radio sources which are endowed with both a redshift determination and an FIR counterpart. Such a fraction is roughly equal to 40 per cent at the lowest fluxes and then also starts declining beyond radio fluxes $F_{1.4 \text{ GHz}} \gtrsim 100 \mu\text{Jy}$.

2.2 GOODS-South

Deep, 1.4 GHz, radio observations of the whole Extended *Chandra* Deep Field South (ECDFS) are presented in Miller et al. (2013). This survey covers about a third of a square degree and images 883 radio sources – out of which 17 are likely multiple-component objects – at the 5σ level down to a peak rms sensitivity of 6 μJy .

A smaller area of the ECDFS, corresponding to what is called the GOODS-Southern Field (hereafter GOODS-S), is one of the best studied regions in cosmology. Thanks to the very deep and exquisite observations taken at all wavelengths, Santini et al. (2009) were able to provide a high-quality photometric catalogue which covers all available wavelengths from 0.3 to 24 μm . The GOODS-MUSIC catalogue (Santini et al. 2009) also includes high-precision photometric redshifts for about 15 000 objects and, whenever possible, spectroscopic redshifts for about 3000 of them.

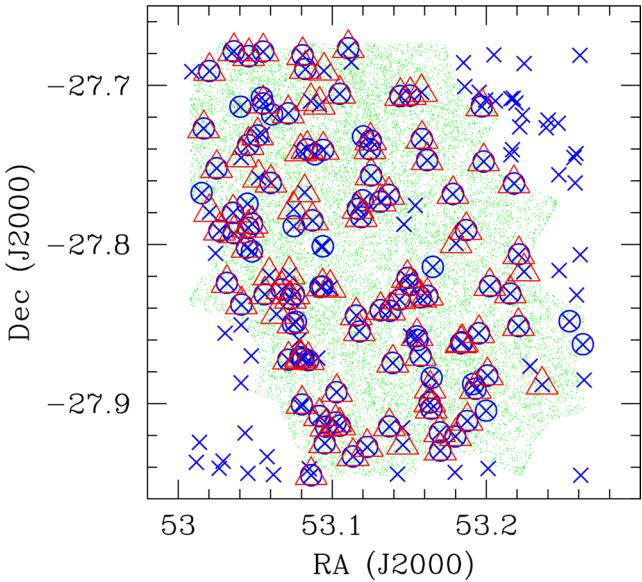


Figure 3. GOODS-South. (Blue) crosses represent radio sources selected at 1.4 GHz by Miller et al. (2013), (green) small dots show the area covered by optical and near-IR surveys, while (red) open triangles illustrate the sources imaged by the PEP and the GOODS-*Herschel* surveys. Crosses surrounded by open circles indicate radio sources with a redshift (either spectroscopic or photometric) determination from the catalogue of Santini et al. (2009).

We then cross-correlated the Miller et al. (2013) radio catalogue with the GOODS-MUSIC one in order to provide the 1.4 GHz-selected radio sources with a redshift determination. 142 radio objects out of the 883 included in the Miller et al. (2013) catalogue fall within the region covered by the GOODS-MUSIC observations. The distribution of the Miller et al. (2013) 1.4 GHz-selected sources on the GOODS-S field is shown in Fig. 3 by the (blue) crosses. The green-shaded area instead shows the distribution of sources from the GOODS-MUSIC catalogue.

Once again with the aim of maximizing the number of sources with a redshift determination, while at the same time minimizing the number of spurious matches, we chose the value of 0.8 arcsec as the maximum radius to be considered for radio-to-optical associations. This procedure leads to 111 photometrically identified radio sources, corresponding to ~ 78 per cent of the parent catalogue. The percentage of expected spurious associations is limited to a mere 3.5 per cent.

The fraction of radio sources in the GOODS-S endowed with a redshift determination as a function of integrated 1.4 GHz radio flux is shown in Fig. 4 by the red circles, where the associated error bars correspond to Poissonian estimates. From the figure, one can notice that such a fraction is constant within the errors, and equal approximately to 80 per cent of the parent sample throughout the whole flux range probed by the Miller et al. (2013) survey. As it was in the case of the LH, the above trend proves that no radio-flux bias affects the optical identifications of the radio sources in exam.

FIR information for the GOODS-S radio sources can be obtained through the PEP (Lutz et al. 2011) and PEP/GOODS-H catalogues (Magnelli et al. 2013). This last one combines observations of the GOODS fields from the PEP and GOODS-*Herschel* (Elbaz et al. 2011) key programmes. This catalogue reaches 3σ depth of 0.9, 0.6 and 1.3 mJy, respectively, at 70, 100 and 160 μ m and includes 1635 sources. As the original PEP catalogue for GOODS-S covers a wider area than that of PEP/GOODS-H, we have identified the FIR counterparts in two steps, first by using the PEP catalogue, and

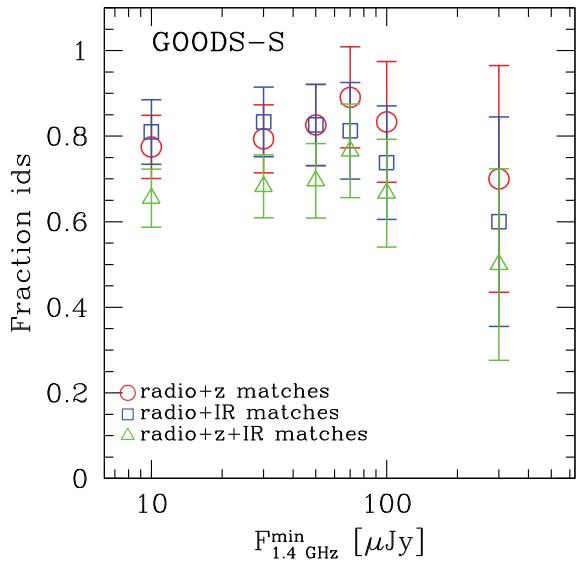


Figure 4. Fraction of radio sources in the GOODS-S endowed with a redshift determination (red circles), with an FIR counterpart from the PEP catalogue (blue squares) and with both an FIR ID and a redshift determination (green triangles) as a function of integrated 1.4 GHz flux. Error bars correspond to Poissonian estimates.

then the PEP/GOODS-H one. As already specified in the previous section, the maximum tolerance radius chosen to associate a radio source with an FIR counterpart was 4 arcsec. This procedure leads to 115, 1.4 GHz-selected, radio sources with an FIR counterpart (corresponding to ~ 81 per cent of the parent catalogue), out of which 110 come from the PEP catalogues either at 100 or at 160 μ m and five from the combined PEP/GOODS-H catalogue. The expected fraction of random associations is less than 2 per cent.

The trend for the fraction of FIR identifications as a function of integrated radio flux is presented in Fig. 4 by the blue squares, while that for sources with both an FIR counterpart and a redshift determination by the green triangles. Errors correspond to Poissonian estimates. Note that in both cases the fractions tend to stay constant, equal to the values of ~ 80 and ~ 70 per cent up to radio fluxes $F_{1.4 \text{ GHz}} \sim 100 \mu\text{Jy}$; then they start to monotonically decline.

2.3 GOODS-North

Deep, 1.4 GHz, radio observations of the GOODS-North field (hereafter GOODS-N) are presented in Morrison et al. (2010). These authors provide a catalogue of 1230 discrete radio sources, imaged over an area of $40 \times 40 \text{ arcmin}^2$. The catalogue comprises all sources above a 5σ detection threshold of $\sim 20 \mu\text{Jy}$ at the field centre.

The GOODS-N, likewise the GOODS-S, is provided with extremely deep observations at all wavelengths. This enables to derive photometric redshifts for the majority of the sources residing there. Furthermore, many spectroscopic redshifts are also available, and the PEP collaboration (Berta et al. 2011¹) provides a list of sources observed in the optical and NIR bands, the majority of which are listed with a redshift, stellar mass and age determination. We then cross-correlated the GOODS-N radio catalogue with that

¹ The multiwavelength GOODS-N catalogue is available on the PEP web page at http://www.mpe.mpg.de/ir/Research/PEP/public_data_releases.php

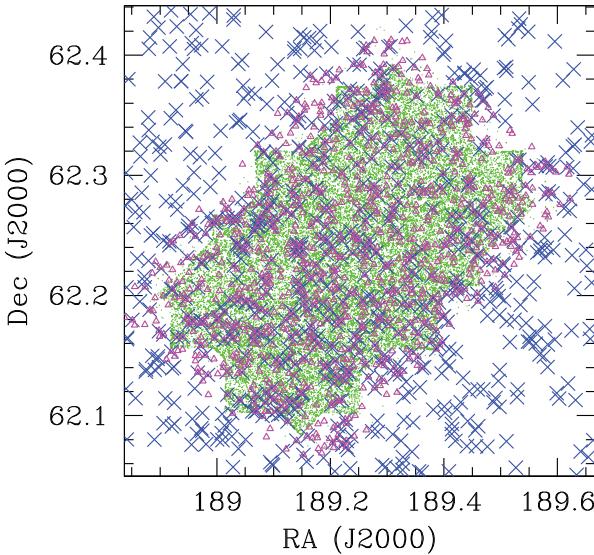


Figure 5. GOODS-North. (Blue) crosses represent radio sources selected at 1.4 GHz by Morrison et al. (2010), (green) small dots show the area covered by optical and near-IR surveys, while (magenta) open triangles illustrate the sources imaged by the PEP and the GOODS-*Herschel* surveys.

of optically selected sources presented above. The number of radio objects which fall within the area observed at the other wavebands is 401. Out of these, 295 have a counterpart in the photo- z catalogue, out to a maximum radius of 0.7 arcsec. 267 of them, corresponding to ~ 67 per cent of the parent radio catalogue, are also endowed with a redshift determination. The distribution of the Morrison et al. (2010) 1.4 GHz-selected sources on the GOODS-N field is shown by Fig. 5 by the (blue) crosses. The green-shaded area instead shows the distribution of optical/near-IR sources from the Berta et al. (2011) catalogue.

As it was for GOODS-S, FIR information for the GOODS-N radio sources was obtained through the PEP (Lutz et al. 2011) and PEP/GOODS-H (Magnelli et al. 2013) catalogues. Once again, the maximum tolerance radius chosen to associate a radio source with an FIR counterpart was 4 arcsec. This procedure leads to 315, 1.4 GHz-selected, radio sources with an FIR counterpart either at 100 μm or at 160 μm . This corresponds to ~ 79 per cent of the parent catalogue, in good agreement with that found in the GOODS-South. Also in this case, the expected fraction of random associations is less than 2 per cent.

The trend for the fraction of FIR identifications as a function of integrated radio flux is presented in Fig. 6 by the blue squares, while that for sources with a redshift determination by the red circles and that for sources with both an FIR counterpart and a redshift determination by the green triangles. In all cases errors correspond to Poissonian estimates. Note that, once again, while the fraction of sources with a redshift determination tends not to depend on flux at all radio fluxes, the fraction of radio sources with an FIR counterpart stays constant, equal to the values of ~ 80 per cent, only up to radio fluxes $F_{1.4\text{ GHz}} \sim 100 \mu\text{Jy}$. Beyond this value such a ratio starts to monotonically decline. The same trend is also reflected in the behaviour of the fraction of radio sources with both a redshift determination and an FIR counterpart.

Table 1 presents a summary of the properties of the 1.4 GHz-selected radio sources on all the three fields considered in this work.

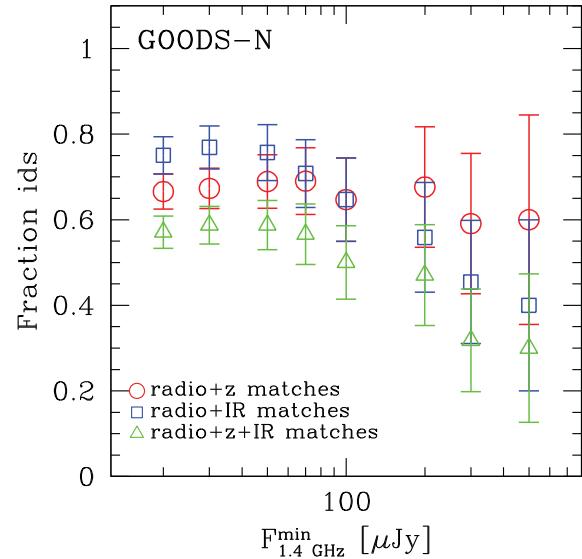


Figure 6. Fraction of radio sources in the GOODS-N endowed with a redshift determination (red circles), with an FIR counterpart from the PEP catalogue (blue squares) and with both an FIR ID and a redshift determination (green triangles) as a function of integrated 1.4 GHz flux. Error bars correspond to Poissonian estimates.

Table 1. Properties of the analysed samples. The various columns represent the three different fields considered in this work. The first row indicates the total number of radio sources selected at 1.4 GHz on regions of the sky covered by deep *Herschel*/PACS observations, N_{TOT} , the second one the number of radio sources with either a spectroscopic or a photometric redshift determination, N_z , the third row the number of radio sources with an FIR counterpart either at 100 μm or at 160 μm , N_{IR} , and the last row the number of radio sources with both an FIR counterpart and a redshift determination, $N_{\text{IR}+z}$.

	LH	GOODS-S	GOODS-N
N_{TOT}	698	142	401
N_z	402	114	267
N_{IR}	422	115	315
$N_{\text{IR}+z}$	260	87	229

3 AGN SELECTION VIA RADIO LUMINOSITY

A tricky issue in the present work is the distinction between radio emission produced by the presence of an AGN and that originating from star formation processes.

The method we adopt here was already introduced in Magliocchetti et al. (2014) and is based on the results of McAlpine et al. (2013) who used the optical and near-infrared spectral energy distributions (SED) of a sample of 942 radio sources (out of 1054 objects selected at 1.4 GHz, with a completeness level of 91 per cent) from the VIDEO-XMM3 field to distinguish between star-forming and AGN-powered galaxies and derive their redshifts. These authors provide luminosity functions for the two classes of sources up to redshifts ~ 2.5 and find different evolutionary behaviours, with star-forming objects evolving in luminosity in a much stronger way [$\propto (1+z)^{2.5}$] than radio-selected AGN [$\propto (1+z)^{1.2}$]. Investigations of their results show that the radio luminosity P_{cross} beyond which AGN-powered galaxies become the dominant radio population scales with redshift roughly as

$$\log_{10} P_{\text{cross}}(z) = \log_{10} P_{0,\text{cross}} + z, \quad (1)$$

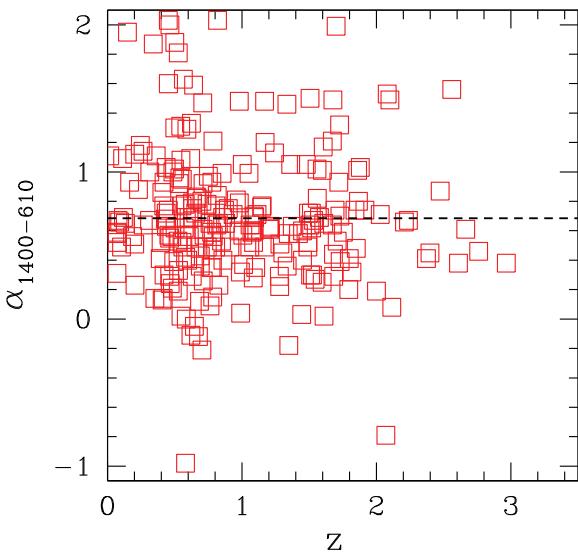


Figure 7. Distribution of radio spectral indices between 1.4 GHz and 610 MHz as a function of redshift for the radio sources in the LH Field with measured redshifts. The horizontal dashed line represents the average of the distribution $\langle \alpha \rangle = 0.685$.

at least up to $z \sim 1.8$. $P_{0,\text{cross}} = 10^{21.7} [\text{W Hz}^{-1} \text{ sr}^{-1}]$ is the value which is found in the local universe and which roughly coincides with the break in the radio luminosity function of star-forming galaxies (cf. Magliocchetti et al. 2002; Mauch & Sadler 2007). Beyond this value, their luminosity function steeply declines, and the contribution of star-forming galaxies to the total radio population is drastically reduced to a negligible percentage. The same trend is true at higher redshifts, and since the radio luminosity function of star-forming galaxies drops off in a much steeper way than that of AGN at all z , we expect the chances of contamination in the two populations to be quite low.

Radio luminosities for our three samples of radio-selected sources endowed with a redshift estimate have been calculated according to the relation:

$$P_{1.4\text{GHz}} = F_{1.4\text{GHz}} D^2 (1+z)^{3+\alpha}, \quad (2)$$

where the result is in $[\text{W Hz}^{-1} \text{ sr}^{-1}]$ units, D is the angular diameter distance and α is the spectral index of the radio emission ($F(\nu) \propto \nu^{-\alpha}$).

As the LH region has been imaged by Ibar et al. (2009) both at 1.4 GHz and at 610 MHz, most of the radio sources in this field have α estimates. The same is not true for either the GOODS-N or the GOODS-S fields. In these latter cases, we then adopted the average value $\alpha = 0.7$ found for similar surveys (e.g. Randall et al. 2012 and references therein) both for star-forming galaxies and for AGN emission. Such an assumption should not be too far from the truth as (1) radio sources in the GOODS samples are faint, therefore the chances of finding a large number of bright, flat spectrum AGN are rather thin and (2) recent results report values $\alpha \simeq 0.7$ also for star-forming galaxies at $z \simeq 2$ (Ibar et al. 2010), similar to that found locally for the same population (Condon 1992). Furthermore, the value of $\alpha = 0.7$ is in perfect agreement with the average value found for sources in the LH ($\langle \alpha \rangle = 0.685$), independent of flux and redshift, as clearly indicated in Fig. 7.

We then distinguished between AGN-powered galaxies and star-forming galaxies by means of equation (1) for $z \leq 1.8$ and by fixing $\log_{10} P_{\text{cross}}(z) = 23.5 [\text{W Hz}^{-1} \text{ sr}^{-1}]$ at higher redshifts (cf. McAlpine et al. 2013). This procedure identifies 357 star-forming

galaxies and 45 AGN (corresponding to 11 per cent of the total radio population) in the LH, 89 star-forming galaxies and 15 AGN (corresponding to 13 per cent of the total radio population) in the GOODS-S, 235 SF and 32 AGN (corresponding to 12 per cent of the total radio population) in the GOODS-N. Note that due to the adopted selection criteria and thanks to the great depths of the considered surveys, *all* the AGN samples are *complete* with respect to radio selection at all probed redshifts, i.e. the considered samples include *all* radio-emitting AGN endowed with a redshift determination.

The AGN distributions as a function of their 1.4 GHz fluxes are presented in the four top panels of Fig. 8 by the solid black lines, where the bottom-right plot refers to the combined GOODS-N and GOODS-S fields, added together because of the very similar depths of the radio, FIR and multi-wavelength surveys. The red, dashed lines represent the distributions of radio-emitting AGN which also possess an FIR counterpart. The histograms in the four bottom panels show the ratio between the full radio-selected AGN population and the sub-class of FIR emitters. As it is possible to notice, in all cases but GOODS-S, where the clarity of the results is hampered by the paucity of sources, there is a systematic trend for the fractions of radio-selected AGN which are also FIR emitters to roughly stay constant at the lowest radio fluxes probed by the different surveys and then possibly decline above $F_{1.4\text{GHz}} \sim 0.5-1 \text{ mJy}$ in the GOODS fields (statistical significance $\sim 2\sigma$) and $F_{1.4\text{GHz}} \sim 1 \text{ mJy}$ in the LH (statistical significance above the 3σ level). Note that even if the general trend is the same in all fields, the success rate for FIR identifications indeed varies from field to field, being lower (~ 50 per cent) in the LH where FIR observations are shallower, and higher (~ 80 per cent) in the two GOODS fields which present the deepest FIR observations obtained so far.

The AGN distributions as a function of redshift are presented in the four top panels of Fig. 9 by the solid black lines, where again the bottom-right plot refers to the combined GOODS-N and GOODS-S fields. The red dashed lines represent the distributions of radio-emitting AGN which also possess an FIR counterpart. The histograms in the four bottom panels show the ratio between the full radio-selected AGN population and the sub-class of FIR emitters. As it is possible to notice, at variance with the former case, there is a clear indication for the fractions of radio-selected AGN which are also FIR emitters to remain constant with look-back time, up to the highest redshifts and in all the fields probed by our analysis. Note that once again the success rate for FIR identifications varies from field to field, being lower (~ 50 per cent) in the LH and higher (~ 80 per cent) in the two GOODS fields.

Some interesting pieces of information on the general properties of radio-emitting AGN can then be gathered through investigation of Figs 8 and 9. In all considered fields their redshift distributions, which we remind, due to the exquisite depths of the radio surveys considered in this work, refer to *complete AGN samples selected in radio luminosity*, present a prominent peak at around $z \sim 1$ and then a possible secondary bump for $z \sim 2$. This is in agreement with the results found by Magliocchetti et al. (2014) for the COSMOS field. Furthermore, the redshift distribution of AGN which also present FIR emission closely follows that of their parent population. Again, this is true for all the considered fields, as well as it was in COSMOS (Magliocchetti et al. 2014). It follows that, in all cases, the fraction of FIR emitters stays constant throughout the whole redshift range up to the earliest epochs probed by our analysis ($z \sim 3-4$), i.e. *there is no preferred age for radio-emitting AGN to be associated with FIR emission*. Indeed, the only appreciable difference amongst the different fields is the value of such a fraction, which

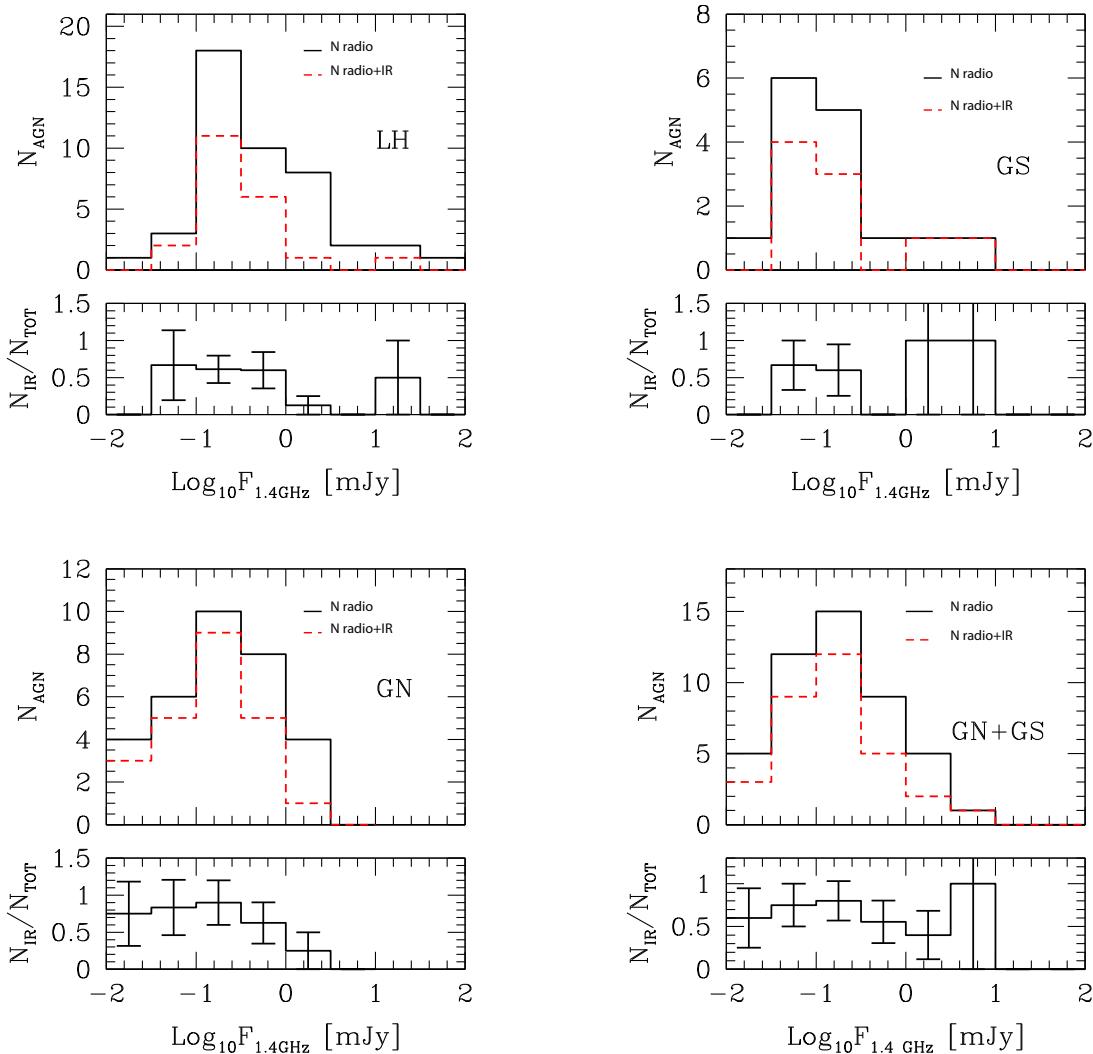


Figure 8. Top panels: distribution of radio-emitting AGN as a function of 1.4 GHz flux. Clockwise, the panels show sources in the LH, in the GOODS-S (GS), in the combined GOODS-N and GOODS-S fields (GN+GS), and in the GOODS-N (GN). The solid black lines represent the whole population of radio-emitting AGN selected on the basis of their radio luminosity, while red dashed lines show the subsets which also have counterparts at FIR wavelengths. Bottom panels: fraction of radio-emitting AGN with an FIR counterpart as a function of radio flux. Error bars correspond to 1σ Poissonian estimates.

is lower and equal to approximately 50 per cent in the LH, while reaches values of about 70–80 per cent in the two GOODS fields where, we remind, FIR observations are the deepest ones available to date.

A point which is worth stressing is the extremely high number of radio-emitting AGN which are found to possess a counterpart in *Herschel* maps. This can be appreciated both in Table 2 and in Figs 8 and 9. In fact, as already mentioned, one goes from the ~ 50 per cent of the LH, to the ~ 80 per cent of the GOODS fields. These figures have to be compared with the ~ 40 per cent found in COSMOS (Magliocchetti et al. 2014), where the different percentages are only due to the different depths of FIR observations on the different fields. Allowing for the relatively large uncertainties due to the small areas covered, these results imply that the overwhelming majority (66 ± 12 per cent, percentage which rises to ~ 72 per cent in the GOODS-N) of radio-emitting AGN in the GOODS-S and GOODS-N are imaged at FIR wavelengths down to fluxes 0.6 mJy at 100 μ m and 1.3 mJy at 160 μ m. This is a rather surprising result, since – at least locally – AGN selected at radio wavelengths have always been associated with passive galaxies, with little or no ongoing stellar activity.

Another remarkable point that emerges from our analysis is that deeper FIR observations do not pick up any new population nor they do single out any different property of the considered population. In fact, the only effect of increased depth is that of evenly increasing the number of sources which are also FIR emitters, independent of look-back time, radio flux and (as it will be more clear later on) many other physical parameters which characterize the sources under examination.

4 PROPERTIES OF THE INNER ENGINE

Important information on the properties of the inner engine of radio-emitting AGN and assessment on whether there are any differences in the presence of FIR emission can be gathered by investigation of the distributions of radio luminosities and radio spectral indices for radio-emitting AGN both in the presence and in absence of FIR emission.

The top panels of Fig. 10 illustrate the distribution of AGN as a function of radio luminosity in the four fields considered here (LH, GOODS-S, GOODS-N and the combined GOODS-S and GOODS-N). Once again, the solid lines represent the full population of

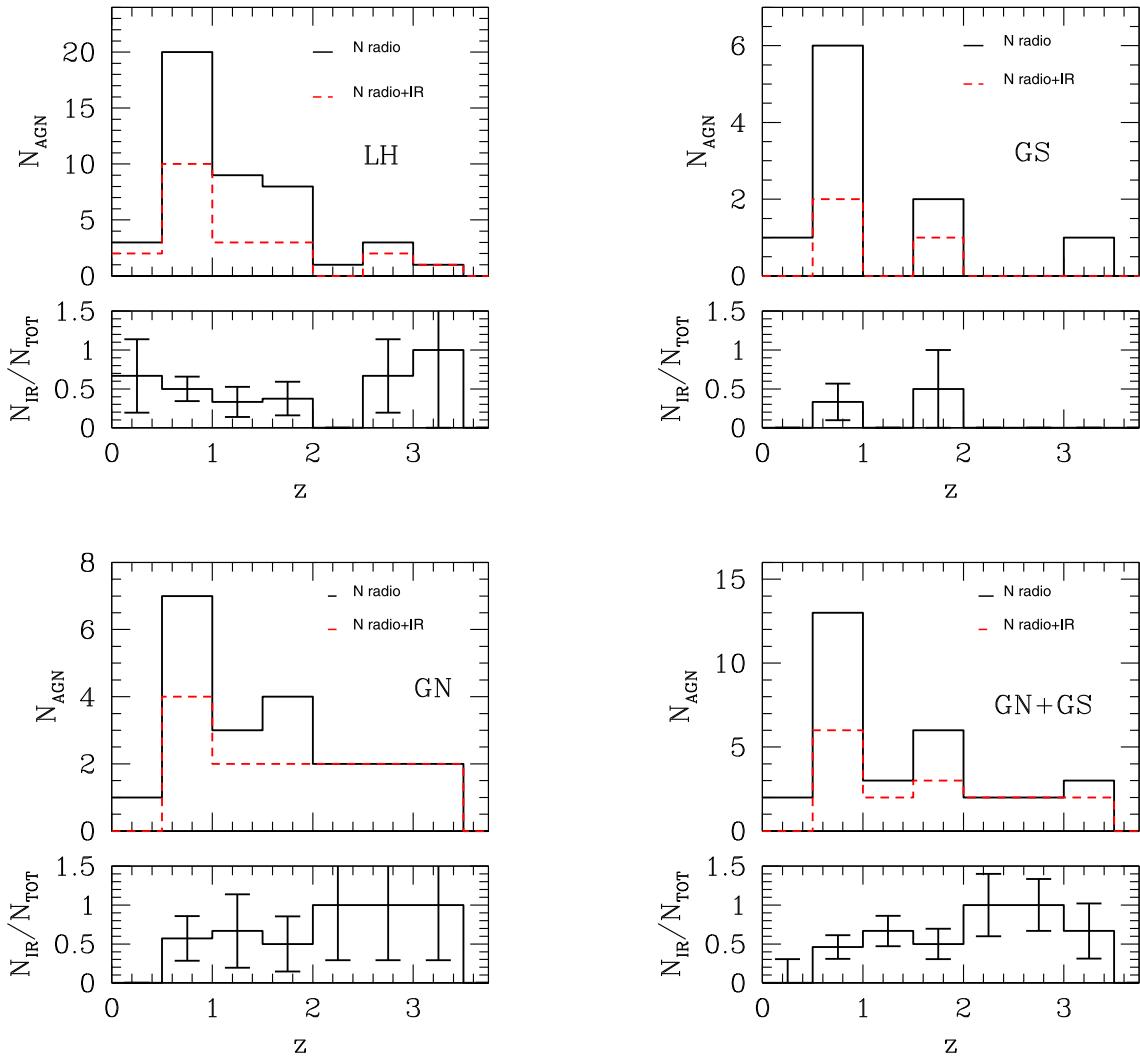


Figure 9. Top panels: distribution of radio-emitting AGN as a function of redshift. Clockwise, the panels show sources in the LH, in the GOODS-S (GS), in the combined GOODS-N and GOODS-S fields (GN+GS), and in the GOODS-N (GN). The solid lines indicate the whole population of radio-emitting AGN selected on the basis of their radio luminosity, while red dashed lines show the subsets which also have counterparts at FIR wavelengths. Bottom panels: fraction of radio-emitting AGN with an FIR counterpart as a function of redshift. Error bars correspond to Poissonian estimates.

Table 2. Properties of the analysed samples. The various columns represent the three different fields considered in this work. The first and second rows respectively indicate the total number of radio-emitting AGN, N_{AGN} , and star-forming galaxies, N_{SF} , selected on the basis of their radio luminosity. The numbers in brackets are for those sources which are also FIR emitters.

	LH	GOODS-S	GOODS-N
N_{AGN}	45 (21)	15 (8)	32 (23)
N_{SF}	357 (238)	89 (79)	235 (206)

radio-emitting AGN, while the dashed ones show the subsets which also have counterparts at FIR wavelengths. The bottom panels instead present the distribution of the fraction of FIR emitters as a function of radio luminosity.

In all the considered fields, there is a systematic drop of sources beyond luminosities $P_{1.4\text{GHz}} \gtrsim 10^{24} \text{ W Hz}^{-1} \text{ sr}^{-1}$. This is due to the fact that very bright radio-selected AGN are extremely rare, and finding them in small fields like those considered here is rather unlikely. Once again, the relative fraction of FIR emitters varies according to the depth of FIR observations on the field, and goes

from the ~ 50 per cent of the LH to the $\gtrsim 70$ per cent of the GOODS-N.

Some interesting information on the behaviour of such sources can be found when splitting the samples into different redshift bins. This is done in Fig. 11 for both the LH and the combined GOODS field. Despite the relative small number of sources, one can clearly see that the number of $P_{1.4\text{GHz}} \gtrsim 10^{23} \text{ W Hz}^{-1} \text{ sr}^{-1}$ radio-emitting AGN with an FIR counterpart steadily increases as one moves back to the earlier universe. In more detail, for the combined GOODS fields we find that *all sources* with $P_{1.4\text{GHz}} \gtrsim 10^{24} \text{ W Hz}^{-1} \text{ sr}^{-1}$ and $z \gtrsim 1$ have a counterpart at FIR wavelengths and, at the same redshifts, between 70 and 80 per cent of them are imaged in *Herschel* maps in the luminosity range $10^{23} \text{ W Hz}^{-1} \text{ sr}^{-1} \lesssim P_{1.4\text{GHz}} \lesssim 10^{24} \text{ W Hz}^{-1} \text{ sr}^{-1}$. The case for the LH is substantially similar, even though the fraction of FIR emitters is not as high as in the GOODS fields as in this latter case FIR observations are shallower.

A better way to visualize the above trend is provided by Fig. 12 which shows the fraction of FIR emitting AGN as a function of radio luminosity in the same three different redshift bins as in Fig. 11. To increase the statistical significance of the results, we

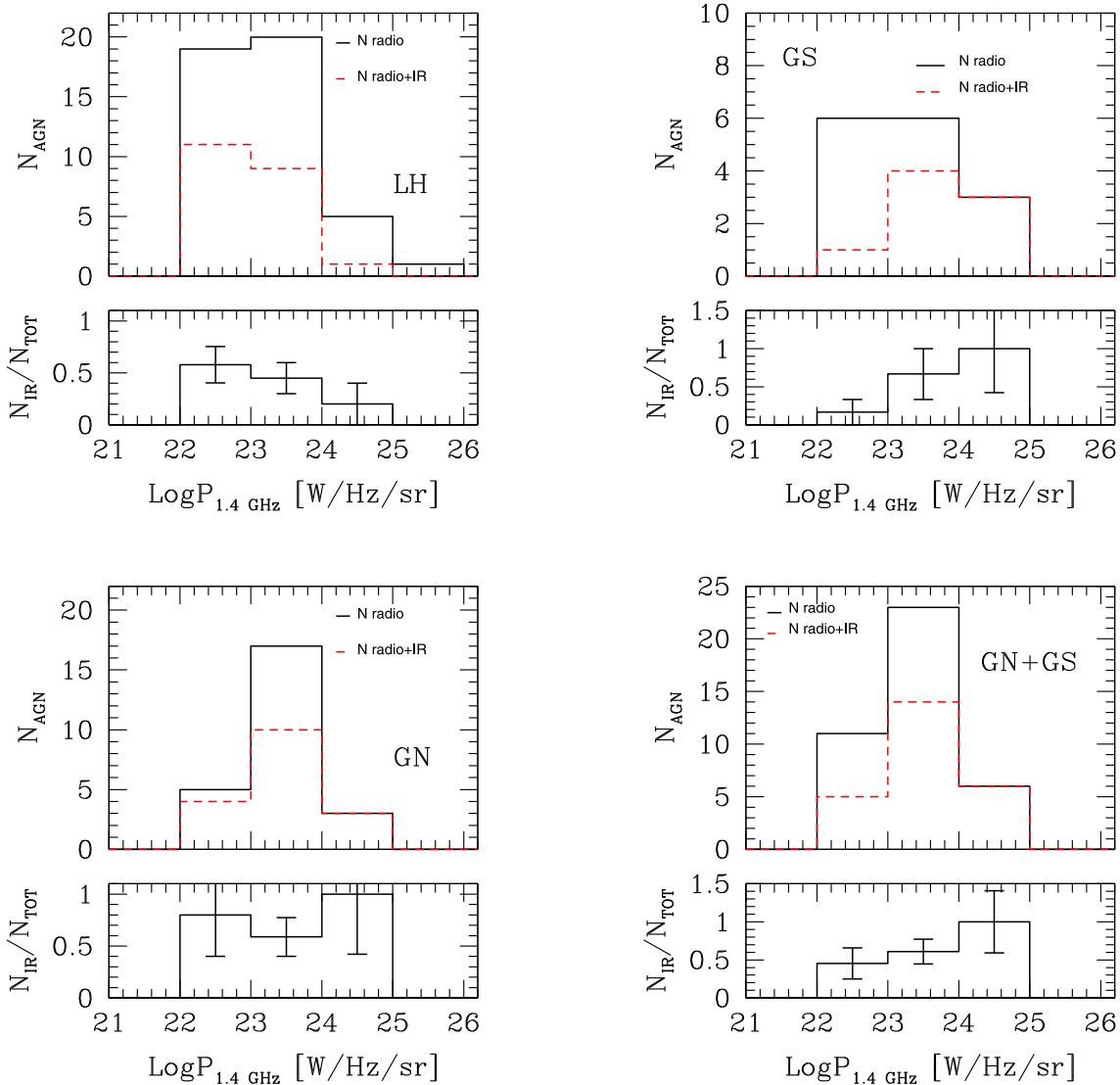


Figure 10. Top panels: distribution of radio-emitting AGN as a function of radio luminosity $P_{1.4\text{ GHz}}$. Clockwise, the panels show sources in the LH, in the GOODS-S (GS), in the combined GOODS-N and GOODS-S fields (GN+GS), and in the GOODS-N (GN). The solid lines indicate the whole population of radio-emitting AGN selected on the basis of their radio luminosity, while red dashed lines show the subsets which also have counterparts at FIR wavelengths. Bottom panels: fraction of radio-emitting AGN with an FIR counterpart. Error bars correspond to Poissonian estimates.

have plotted this quantity for all the three fields combined together. In this case, one can appreciate that for redshifts greater than ~ 1 there is a substantial constancy with radio luminosity of the fraction of AGN which are also FIR emitters. In other words, *radio-selected AGN at redshifts $\gtrsim 1$ have all the same chances of being FIR emitters independent of their radio luminosity*. However, this is not true at lower redshifts. In fact, the plot shows that in the more local universe there is a ($\sim 2\sigma$) trend for FIR emitters to be mostly associated with low-luminosity sources. Such a probability then decreases for brighter AGN. We further stress that, despite the different depth of FIR observations, the same behaviours are observed both in the LH and in the combined GOODS fields, therefore strengthening the significance of our statement.

What one can conclude from this analysis is that chances for an AGN which is active at radio wavelengths to also emit at FIR wavelengths is maximum at high ($z \gtrsim 1$) redshifts. For deep enough FIR observations such as those performed on the GOODS fields, this probability reaches the value of ~ 100 per cent, independent of

radio luminosity. This implies that *almost all* radio-selected AGN at $z \gtrsim 1$ are also the site of FIR emission. The situation changes in the more local, $z \lesssim 1$, universe, as in this latter case one finds that the relative fraction of FIR-active AGN is lower and, most importantly, decreases as a function of radio luminosity. In other words, one has that in the local universe, mostly low-luminosity AGN are found to be associated with FIR emission. Only very few higher-power, $P_{1.4\text{ GHz}} \gtrsim 10^{23} \text{ W Hz}^{-1} \text{ sr}^{-1}$, sources instead appear on FIR maps. We note that these conclusions, although based on a relatively small sample of sources, agree and further extend to fainter FIR emission those obtained by Magliocchetti et al. (2014) for the COSMOS field which, although covering a wider area, suffered of the limits due to relatively shallow FIR observations. They also agree with the local results which find radio-emitting AGN to reside in early-type galaxies, with little or no on-site ongoing star-forming activity. Our data, however, indicate that this is only true in the local universe, while in the more distant one FIR emission is extremely common amongst radio-active AGN.

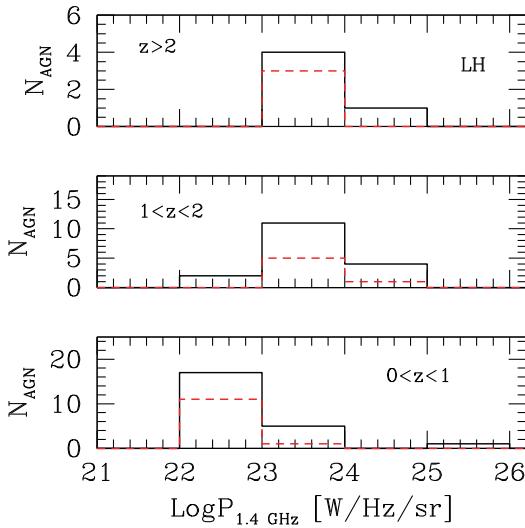


Figure 11. Distribution of radio-emitting AGN as a function of radio luminosity $P_{1.4 \text{ GHz}}$ in three different redshift bins. The left-hand panel shows sources in the LH, while that on the right-hand side sources in the combined GOODS-N and GOODS-S fields (GN+GS). The solid lines indicate the whole population of radio-emitting AGN selected on the basis of their radio luminosity, while red dashed lines show the subsets which also have counterparts at FIR wavelengths.

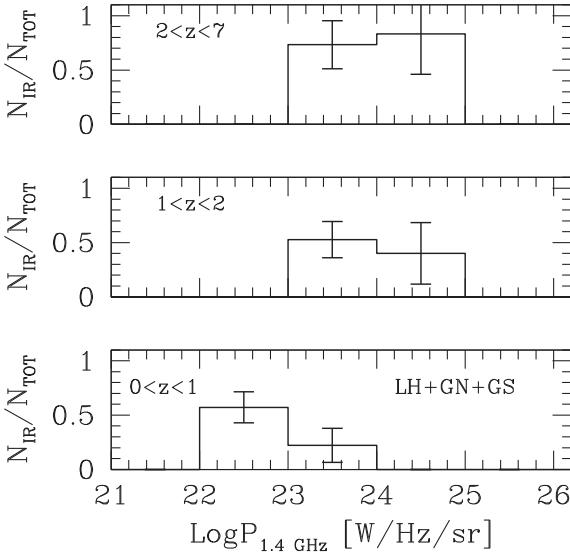
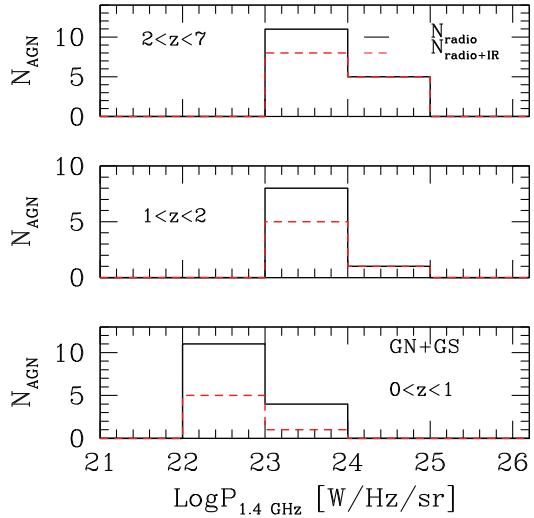


Figure 12. Distribution of the fraction of FIR emitting AGN as a function of radio luminosity in three different redshift intervals. Error bars correspond to Poissonian estimates. Data refer to all fields considered in this work.

Further information on the engine of radio-active AGN can be obtained by investigating the spectral indices of their radio emission. As already discussed in Section 2.1, the LH has been observed by Ibar et al. (2009) both at 1400 and 600 MHz. Most of the observed sources (41 out of 45) are endowed with solid estimates of the quantity $\alpha_{1400 - 600}$. The trends for this quantity as a function of both redshift and radio luminosity are shown in Fig. 13. Squares represent the entire population, and filled ones are for the subset of FIR emitters. The horizontal dotted line shows the average $\langle \alpha \rangle = 0.58 \pm 0.42$ value obtained for the entire population, and the dashed one represents that of $\langle \alpha \rangle = 0.68 \pm 0.29$ derived for FIR emitters. As it is clear also from these visual representations, there is no specific trend in the distribution of α values due to the eventual presence of FIR emission within the galaxies host of a radio-active AGN.



All the distributions look identical, irrespective of the fact that the radio-active AGN also show emission at FIR wavelengths. This strongly suggests that, at least at radio wavelengths, the central engine responsible for the AGN phenomenon is oblivious to the FIR activity which is ongoing within the same host galaxy or, in other words, that FIR processes do not influence the activity of the central engine, at least that detected at radio wavelengths.

5 PROPERTIES OF THE HOST GALAXY

The results obtained in Section 4 have clearly shown that FIR activity is a very common process amongst radio-active AGN. In fact, up to $\gtrsim 70$ per cent of the radio-AGN population is found to show up in the *Herschel* maps. We have also shown that, except maybe for the most local ($z \lesssim 0.5$; cf. Fig. 9) universe, the probability for a radio-active AGN to also emit at FIR wavelengths is independent of redshift. This implies that there is no preferred epoch for a galaxy to be the site of both radio-AGN and FIR activities. Furthermore, we have found that, at high enough ($z \gtrsim 1$) redshifts, such a probability does not even depend on the power of radio emission from the central black hole. Radio luminosity of AGN origin only starts becoming important in the low-redshift universe, whereby the chances of finding FIR emission within the host galaxy are maximum for low-luminosity sources, while they monotonically decrease for increasing radio luminosities. Perhaps more importantly, our study has proven that there is no difference between the (radio) AGN emission associated with galaxies with ongoing FIR activity and that in those galaxies which are FIR-quiet.

All the above conclusions strongly suggest that, at least for what concerns radio emission, the central engine ignores the existence of FIR activity within the same host galaxy, both in terms of its cosmological evolution, of its inner properties and, except for in the very local universe, in terms of its radiative power. In other words, FIR activity does not influence radio-AGN activity and – except in the very local universe – the two processes seem to proceed independently of each other.

But what happens at the host galaxy level? Is there any difference in the properties of galaxies which only host a radio-active AGN or are the sites of both radio-AGN and FIR activity? In order to answer

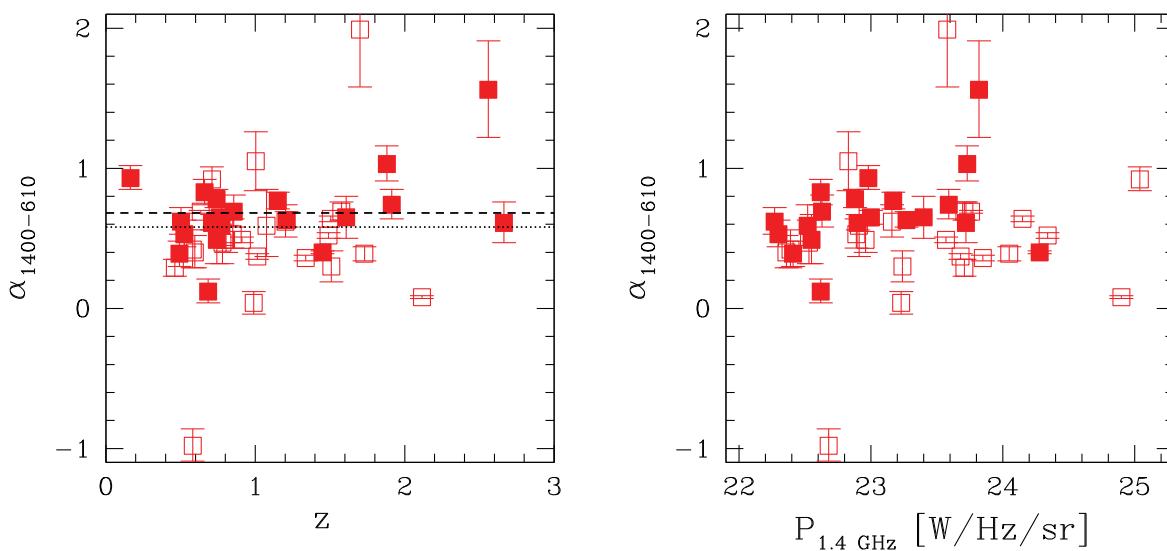


Figure 13. Radio spectral indices between 1.4 GHz and 600 MHz for the population of radio-emitting AGN as a function of redshift (left-hand plot) and radio luminosity (right-hand plot). The filled squares indicate the sub-population of FIR emitters. In the left hand plot, the dotted line shows the average value of $\alpha_{1400} - 610 \simeq 0.58$ obtained for the whole population, and the dashed one that of $\alpha_{1400} - 610 \simeq 0.68$ derived for FIR emitters. Data refer to the LH.

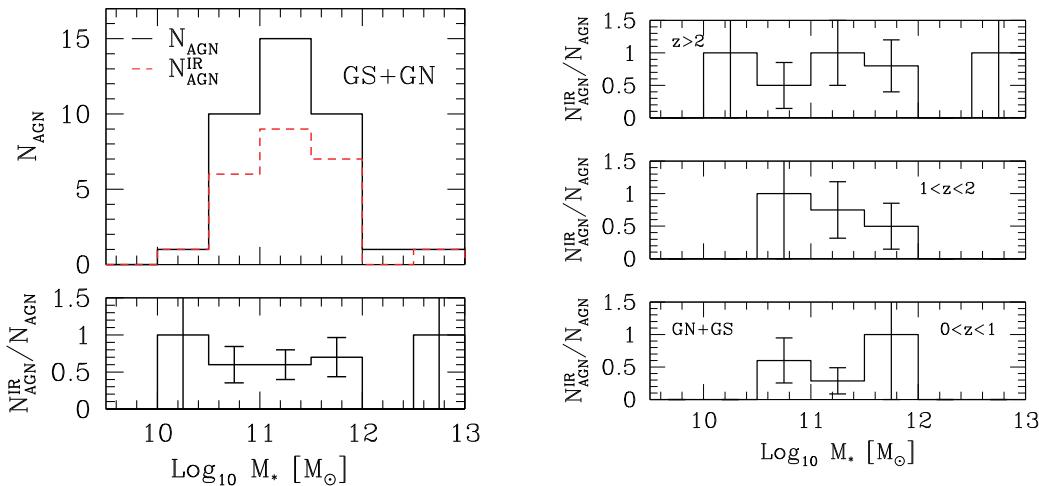


Figure 14. Left-hand panel: stellar mass distribution for the galaxies hosts of radio-active AGN activity. The solid line shows the whole galaxy population, while the dashed one represents the sub-class of FIR emitters. The bottom panel illustrates the ratio between these two quantities. Right-hand panel: fraction of radio-active galaxies which are also FIR emitters as a function of stellar mass in three different redshift bins. In all cases error bars correspond to Poissonian estimates. Data refer to the combined GOODS-N and GOODS-S fields.

these questions, we have then looked into the distributions of stellar masses and ages for galaxies hosts of the AGN phenomenon only and of both AGN and FIR activity. This has only been possible for the GOODS-N and GOODS-S fields, as the LH does not have estimates for either masses or ages of the galaxies within the field.

The distribution of stellar masses for galaxies in the combined GOODS-S and GOODS-N fields host of the radio-active AGN selected in our work is presented in the top-left panel of Fig. 14. The solid histogram refers to the whole population, while the dashed histogram indicates the sub-class of FIR emitters. The bottom-left panel shows the ratio between these two quantities.

Two features can be clearly appreciated from these plots: the first one is that AGN hosts are very massive galaxies. Indeed, the mass distribution in both fields peaks in the range $M_* = 10^{11}\text{--}10^{11.5} M_\odot$, and more than 95 per cent of the objects have masses larger than $10^{10.5} M_\odot$. Furthermore, a non-negligible fraction of

these sources (~ 26 per cent) possesses stellar masses in the range $[10^{11.5}\text{--}10^{12}] M_\odot$.

The second feature that can be appreciated is that there is no difference in the mass distribution of FIR-active and FIR-inactive AGN. As clearly shown in the bottom-left panel of Fig. 14, the two distributions are identical. Not only that, such a constancy is seen throughout the whole redshift range probed by our data (cf. right-hand plot of Fig. 14), implying that there was never any difference between the masses of the hosts of AGN phenomena alone and the masses of the hosts of both AGN and FIR activity. This finding is somehow in contradiction with the results of Magliocchetti et al. (2014) on the COSMOS field, which indicated that FIR emitters preferred somehow slightly lower-mass galaxies ($\log_{10}[M_*^{\text{AGN}}/M_\odot] - \log_{10}[M_*^{\text{AGN+FIR}}/M_\odot] \sim 0.5$). However, the discrepancy can be explained by recalling that COSMOS is covered by much shallower FIR observations than the two GOODS

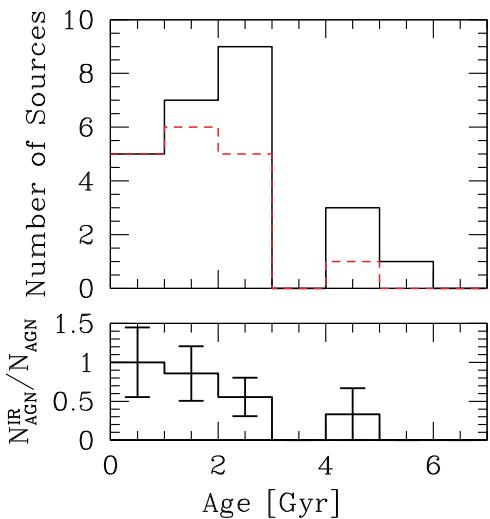


Figure 15. Distribution of ages for galaxies in the GOODS-N hosting a radio-active AGN. The solid histogram in the top panel shows the trend for the entire AGN population, while the dashed one that for the subset of FIR emitters. The bottom panel presents the ratio between these two quantities. Error bars correspond to Poissonian estimates.

fields and therefore the information one can gather from its analysis is possibly affected by biases due to incompleteness. Indeed, only 40 per cent of the radio-active AGN in COSMOS were found to have a counterpart in *Herschel* maps, while this percentage rises to $\gtrsim 70$ per cent in the GOODS fields.

Some more information on these galaxies can be obtained, albeit only in the case of the GOODS-N, from investigations of the distribution of their ages, τ . This is presented in Fig. 15, which shows as a solid histogram the distribution of ages for the entire AGN population, and with the dashed one that for the sub-population of FIR emitters. Although the statistical significance of the result ($\sim 1\sigma$) is hampered by the small number of sources, Fig. 15 suggests that even in this case the presence of FIR emission does not alter the age properties of the galaxies under exam, at least up to about 2–3 Gyr. For galaxies older than this value, instead there seems to be a deficit of FIR emitters. The bulk of the distribution, both in the presence and absence of FIR emission, is found for relatively young, $\tau \lesssim 3$ Gyr sources. Beyond this value, the number of AGN drops dramatically, as only ~ 15 per cent of them is older than this value.

On the basis of what is found in this section, we can then conclude that neither the age nor the mass distribution of galaxies host of a radio-AGN phenomenon varies in the presence of FIR emission.

6 FIR PROPERTIES

Magliocchetti et al. (2014) have proven that FIR emission from the hosts of radio AGN was entirely to be attributed to star-forming processes within the galaxy itself. This is also true in the present case. Fig. 16 reports the distribution of FIR, $F_{100\mu\text{m}}/F_{160\mu\text{m}}$, colours as a function of redshift and for different ranges of radio luminosity. The figure refers to all the three considered fields, whereby small symbols are for GOODS-S, intermediate-sized ones for GOODS-N and large symbols for the LH. The dashed and dotted lines represent the SED of two well-studied templates for star-forming galaxies such as M82 (dashed line) and Arp220 (dotted line).

It is clear that irrespective of cosmic age and radio luminosity, and except for very few outliers, the FIR colour distribution of

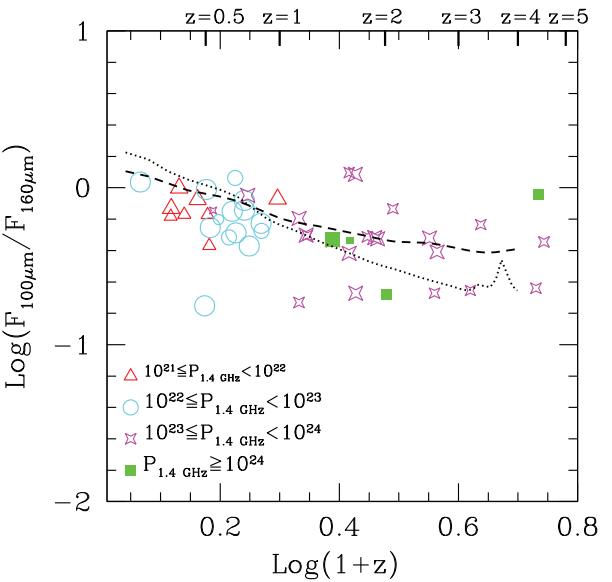


Figure 16. FIR colour distribution as a function of redshift for those radio-active AGN which also emit at FIR wavelengths. As illustrated by the plot itself, sources are colour-coded according to their radio luminosity. Small symbols are for AGN within the GOODS-S, intermediate-sized ones for those within the GOODS-N and large symbols for AGN within the LH. The dashed line represents the SED of M82, while the dotted one the SED of Arp220.

radio-active AGN perfectly follows those of typical moderate-to-intense star-forming galaxies. The radio excess due to the presence of a radio-active AGN is only visible, and not in all cases, in the distributions of the quantities $q_{100} = \log [F_{100\mu\text{m}}/F_{1.4\text{GHz}}]$ and $q_{160} = \log [F_{160\mu\text{m}}/F_{1.4\text{GHz}}]$ represented in Fig. 17.

So, we have shown that FIR emission from radio-selected AGN originates from star-forming processes within the host galaxy. But how bright are these sources at FIR wavelengths? And what are their star-forming rates (SFRs)? In order to provide such answers, we have then estimated the total IR luminosities, L_{IR} , of those radio-active AGN which showed a counterpart in *Herschel* maps, and subsequently calculated their SFR according to the standard relation (Kennicutt 1998, which holds for a Salpeter IMF and for stellar masses in the range $\sim 0.1\text{--}100 M_{\odot}$): $\text{SFR} [\text{M}_{\odot} \text{ yr}^{-1}] = 1.8 \times 10^{-10} L_{\text{IR}}/L_{\odot}$.

These are shown in Fig. 18, where the left-hand plot refers to the LH, while the right-hand one is for sources in the combined GOODS-N+GOODS-S field. The bottom x -axes report the bolometric luminosities, while SFRs are shown on the top x -axes. Note that the three objects in the combined GOODS-N+GOODS-S field endowed with bolometric luminosities above $10^{14} L_{\odot}$ have estimated redshifts between $z \simeq 4.5$ and $z \simeq 7$. Until such redshifts are spectroscopically confirmed, caution should be taken in interpreting results connected to these galaxies.

Our data show how radio-active AGN are extremely luminous sources: their bolometric luminosities extend up to $\sim 10^{13}\text{--}10^{14} L_{\odot}$, which in turn correspond to extremely high, $\text{SFR} \gtrsim 10^3 M_{\odot} \text{ yr}^{-1}$, SFRs. These results imply that galaxy hosts of radio-AGN activity are indeed also the sites of extremely intense star-forming emission. AGN activity does not seem to inhibit star formation within the host galaxy, just as FIR activity does not seem to affect radio-AGN luminosity, at least in the relatively distant ($z \gtrsim 1$; cf. Section 4) universe. In other words, during most of the lifetime of such sources, AGN and star formation do not seem to influence

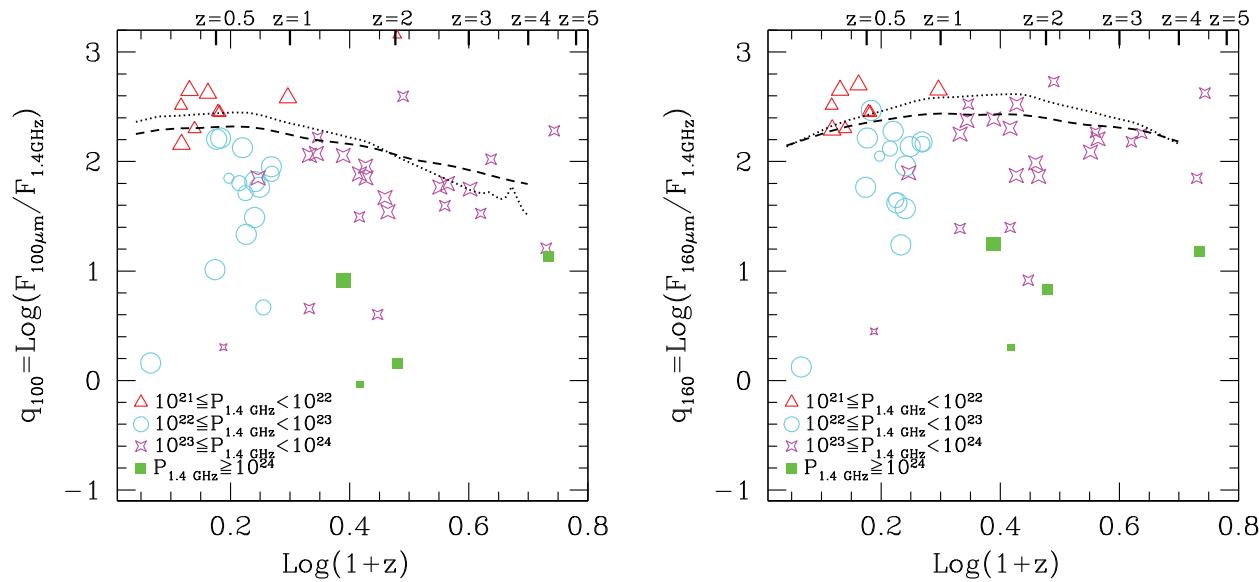


Figure 17. Distribution of the quantities $q_{100} = \log [F_{1.4\text{GHz}}/F_{100\mu\text{m}}]$ (left-hand plot) and $q_{160} = \log [F_{1.4\text{GHz}}/F_{160\mu\text{m}}]$ (right-hand plot) as a function of redshift for those radio-active AGN which also emit at FIR wavelengths. Sources are colour-coded according to their radio luminosity. Small symbols are for AGN within the GOODS-S, intermediate-sized ones for those within the GOODS-N and large symbols for AGN within the LH. The dashed lines represent the SED of M82, while the dotted ones those of Arp220.

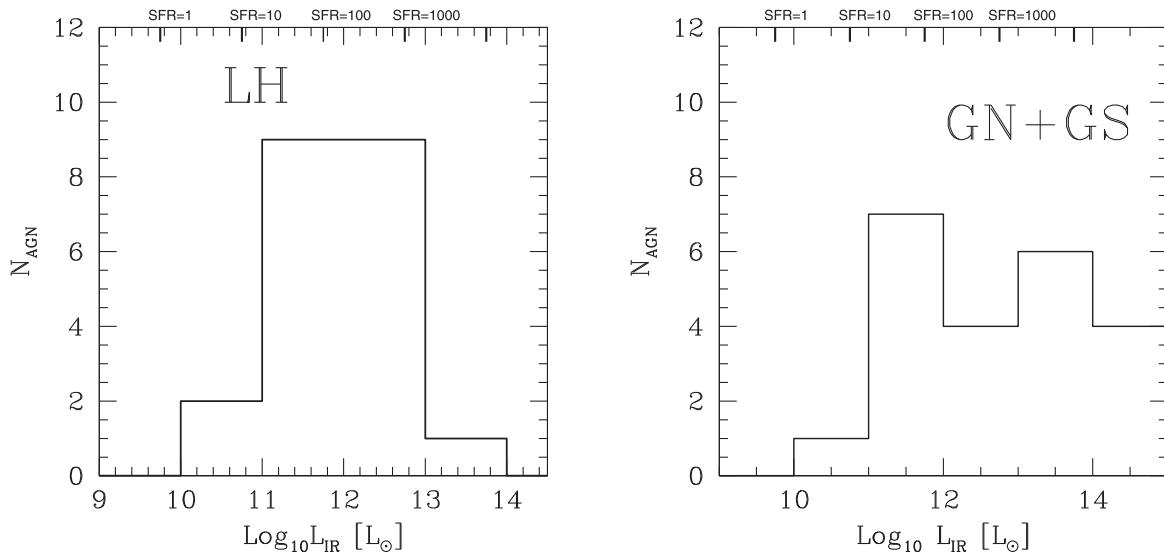


Figure 18. Distribution of bolometric luminosities (L_{IR} , bottom x -axes) and SFR (top x -axes, expressed in $M_{\odot}\text{ yr}^{-1}$ units) for radio-active AGN selected within the LH (left-hand panel), and the combined GOODS-S and GOODS-N fields (right-hand panel).

the activity of each other. Furthermore, as we learned in Section 5, also the physical properties of the galaxies host of a radio-active AGN are only – if any – very minimally affected by the presence of a star-forming event within the same galaxy. Given the extremely high rate of associations which, we re-iterate, approaches $\gtrsim 70$ per cent in deep fields such as the GOODS-N, we stress that the above conclusions do not simply refer to a limited sub-sample of radio-active AGN but rather include the overwhelming majority of the considered sources: most of radio-active AGN are associated with intense episodes of star formation. However, the two processes proceed independently within the same galaxy, at all redshifts except in the most local universe where possibly intense radio activity acquires enough power to switch off the on-site star formation.

7 AGN CONTRIBUTION TO GALAXY EVOLUTION IN RADIO-SELECTED OBJECTS

In the previous sections we have shown that there is substantially no difference between the global properties of radio-selected AGN and those of the sub-population of sources which are also associated with episodes of star formation within their hosts. These findings suggest that, at least in galaxies hosting a radio-active AGN event, star formation is not the main trigger responsible for shaping the formation and evolution of such sources. What we are left with is then the presence of the AGN. Is this really the main factor which influences galaxy evolution?

In order to investigate this issue in more detail, we have compared the distributions of masses and infrared luminosities/star

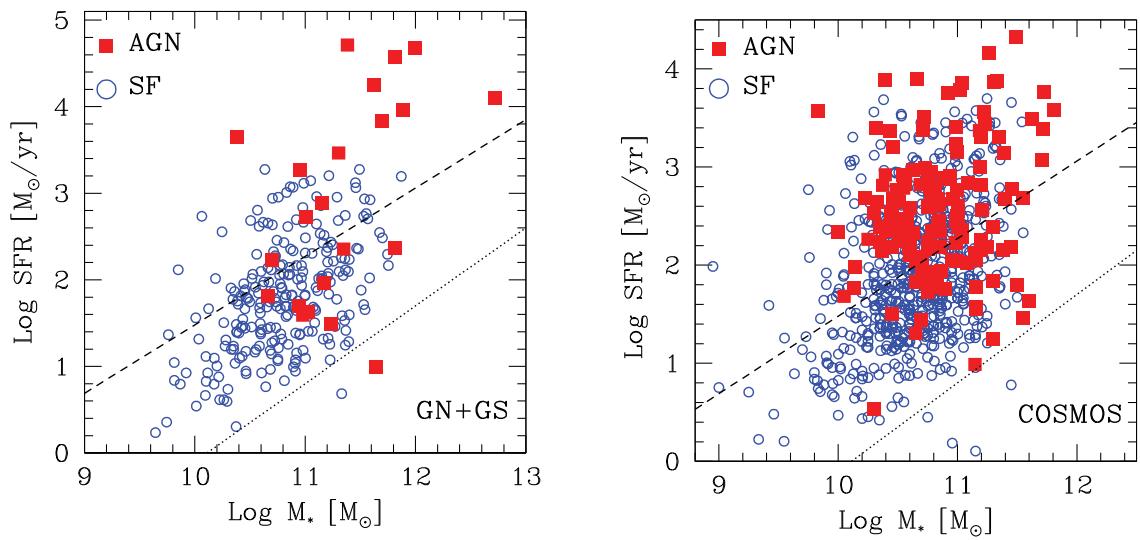


Figure 19. Distributions of stellar masses versus SFRs for radio-selected sources at all redshifts which are also FIR-detected. Filled squares represent AGN, while open circles show star-forming galaxies. The left-hand plot refers to the combined GOODS-N+GOODS-S fields, while the right-hand plot illustrates the case for COSMOS. The dashed lines indicate the relation obtained for main-sequence galaxies at $z \sim 2$ by Rodighiero et al. (2011), while the dotted lines that derived for local galaxies by Brinchmann et al. (2004) and Peng et al. (2010).

formation rates for the hosts of radio-selected AGN with those of radio-selected star-forming galaxies, derived from the samples considered in the present work on the basis of their radio luminosity as explained in Section 3. This is shown in Fig. 19, where the filled squares refer to AGN, while the open circles to star-forming galaxies. The left-hand side plot shows the results obtained for the combined GOODS-N and GOODS-S fields. The right-hand side plot instead refers to COSMOS, as the LH does not have estimates for the masses of the galaxies which reside there (cf. Section 5). Objects in the COSMOS field belong to the sample presented by Magliocchetti et al. (2014) and have been selected in the same way as sources in the present work. In order to guide the eye, the dashed lines in both plots represent the so-called main sequence for star-forming galaxies obtained by Rodighiero et al. (2011) for $z \sim 2$ galaxies, while the dotted line shows that derived in the local universe by Brinchmann et al. (2004) and Peng et al. (2010).

Fig. 19 clearly shows that the distribution of radio-selected sources which contain a central radio-active AGN is markedly different from that of those galaxies which do not. First of all, AGN are much more FIR-bright objects. The overwhelming majority of these sources indeed exhibit SFRs above $\sim 100 M_{\odot} \text{ yr}^{-1}$, while star-forming galaxies are also found down to very low IR luminosities (of the order of a few $10^{10} L_{\odot}$). At the same time, the luminosity distribution of radio-selected AGN extends up to extremely bright luminosities, both in the case of COSMOS and in the case of GOODS-S+GOODS-S fields, while that of star-forming galaxies stops at around $10^3 M_{\odot} \text{ yr}^{-1}$ (with slightly higher values observed in COSMOS which, we remind, includes on average brighter FIR sources due to the lower sensitivity of the *Herschel* observations both at $100 \mu\text{m}$ and $160 \mu\text{m}$).

In addition, our data clearly show that star-forming galaxies are much smaller objects. Indeed, while AGN hosts have masses always $\gtrsim 10^{10.5} M_{\odot}$, star-forming galaxies can be found for stellar masses as low as $\sim 10^9 M_{\odot}$. On the other hand, while star-forming galaxies disappear beyond $\sim 10^{11.5} M_{\odot}$, a non-negligible fraction of radio-selected AGN can still be found beyond that value. This result is better illustrated in Fig. 20 which indeed shows that, irrespective of the considered field, there is a trend for AGN to be increasingly

more numerous within the population of radio-selected sources as the stellar mass of the hosts increases (cf. bottom panels of both plots).

So, are these differences between radio-selected AGN and star-forming galaxies inbuilt in the process of their formation or do these two populations differentiate at a certain point of their evolution? The answer can be found in the plots presented in Figs 21 and 22. These show that a part from a number of outliers which mostly represent objects set at very high ($z \gtrsim 4$) redshifts and mainly lie in the combined GOODS-S+GOODS-N fields (see also the discussion in Section 6), the distributions of radio-selected AGN and star-forming galaxies are almost indistinguishable one from the other down to $z \sim 1$. The two populations only differentiate in the most local universe, whereby AGN are endowed with much larger masses and luminosities than star-forming galaxies.

These results are in agreement and extend the findings of the former sections which indicate that for redshifts above ~ 1 there is no difference between radio-selected AGN and the sub-population of those which are also active at FIR wavelengths. Here we have seen that, except for a few cases which mostly reside in the very early universe, at the same high redshifts there is also no difference between FIR-active, radio-selected, AGN and FIR-active, radio-selected, star-forming galaxies. In all cases, differences appear only in the local universe: radio-active AGN start progressively becoming FIR-quiet the more as the higher is their radio luminosity (cf. Section 4). At the same time, AGN are found associated with much higher galaxy masses than star-forming galaxies. Interestingly enough, the few AGN which still emit at FIR wavelengths in the local universe exhibit higher FIR luminosities than star-forming galaxies (cf. left-hand plots of Figs 21 and 22).

8 CONCLUSIONS

With the aim of providing an analysis as exhaustive as possible of the FIR properties of radio-selected AGN, we have considered three very well studied fields such as the LH, the GOODS-N and

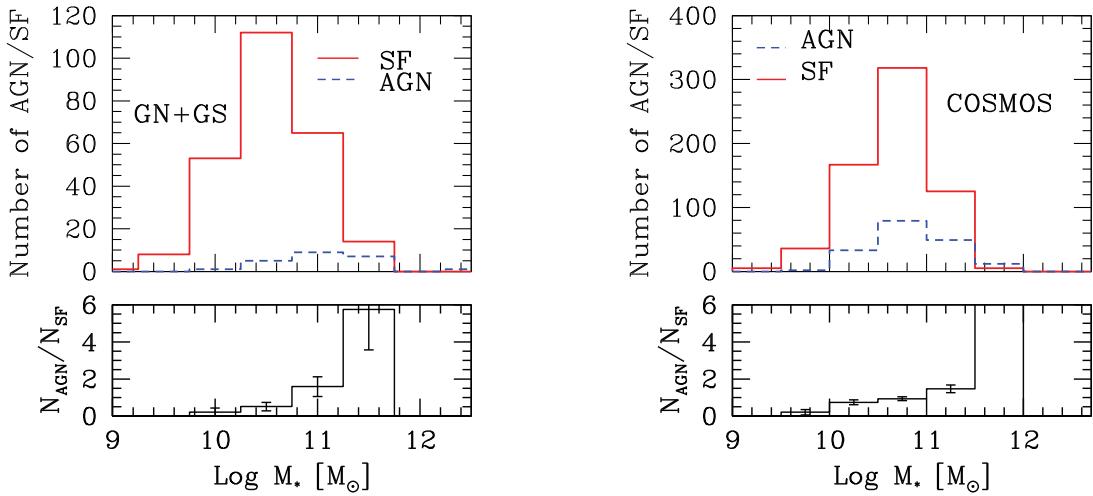


Figure 20. Mass distribution for sources in the combined GOODS-N and GOODS-S fields (left-hand plot) and COSMOS (right-hand plot). The solid histograms represent the class of radio-selected star-forming galaxies, while the dashed ones that of radio-selected AGN. The bottom panels show the ratio between the two distributions, normalized by the total number of AGN and SF in each field. Error bars correspond to Poissonian estimates.

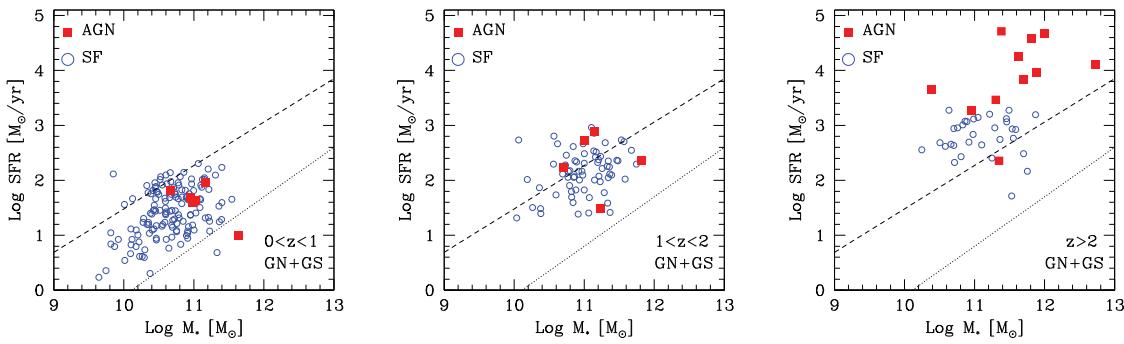


Figure 21. Distributions of stellar masses versus SFRs for radio-selected sources which are also FIR detected in the combined GOODS-N+GOODS-S fields. Filled squares represent AGN, while open circles show star-forming galaxies. Different panels refer to different redshift intervals. The dashed and dotted lines are as in Fig. 19.

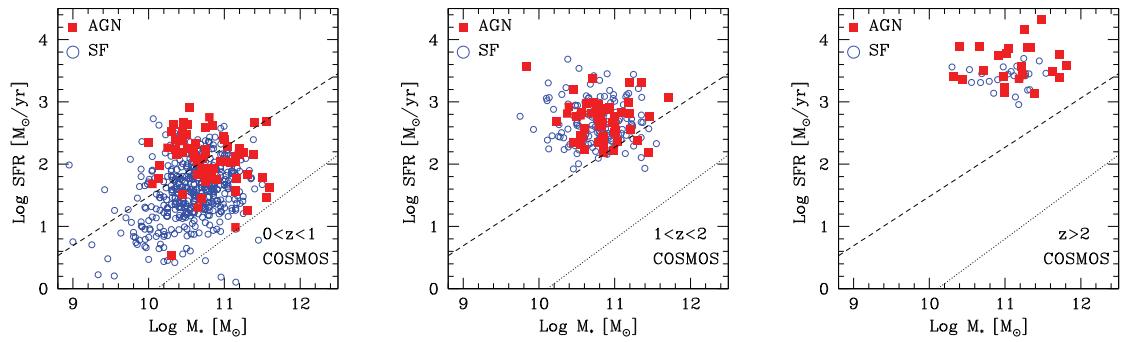


Figure 22. Distributions of stellar masses versus SFRs for radio-selected sources which are also FIR-detected in the COSMOS field. Filled squares represent AGN, while open circles show star-forming galaxies. Different panels refer to different redshift intervals. The dashed and dotted lines are as in Fig. 19.

the GOODS-S. These fields are indeed provided with the deepest radio and FIR observations available up-to-date and have also been observed at a large number of other wavelengths, allowing the determination of spectroscopic and photometric redshifts for the overwhelming majority of the sources residing there.

For the primary radio selection, we used data from Ibar et al. (2009), Miller et al. (2013) and Morrison et al. (2010). In all cases, these reach depths of $\sim 20 \mu\text{Jy}$ at 1.4 GHz.

In order to distinguish between radio emission due to star formation ongoing in the host galaxy and radio emission of AGN origin, we used the method already presented in Magliocchetti et al. (2014) which is based on the radio luminosity of the sources. This marks as an AGN all those objects with 1.4 GHz luminosities brighter than a chosen threshold, which roughly corresponds to the break in the radio luminosity function of star-forming galaxies and which positively evolves with look-back time. The number of

radio-emitting AGN found on the three considered fields with the above method is 92: 45 in the LH, 15 in the GOODS-S and 32 in the GOODS-N.

By then relying on the very deep observations provided by the *Herschel* mission at 100 μm and 160 μm , we used the PEP (Lutz et al. 2011) and GOODS-*Herschel* (Elbaz et al. 2011) data to look for FIR counterparts to radio-selected AGN. Thanks to the extreme depth of FIR observations, we found that the large majority of radio-emitting AGN are indeed FIR emitters. The fraction of sources with ongoing FIR activity reaches values as high as $\gtrsim 70$ per cent in the case of GOODS-N and ~ 50 per cent in the LH, where FIR observations are shallower. We found that possibly except for the very local, $z \lesssim 0.5$, universe the chances for FIR emission in a radio-active AGN do not depend on the redshift of the sources, implying that FIR activity is a common event throughout the whole lifetime evolution of these objects.

As shown in Magliocchetti et al. (2014), we confirm that FIR emission in radio-active AGN is entirely due to star-forming activity within the host galaxy. The bolometric luminosities of these sources are extremely high: the peak of their distribution is observed at around $L_{\text{IR}} = 10^{12.5} L_{\odot}$, with AGN reaching L_{IR} values as high as $10^{14} L_{\odot}$. Converted into SFRs, this implies sources with average $\text{SFR} \sim 10^2 M_{\odot} \text{ yr}^{-1}$ which can even reach values up to $10^4 M_{\odot} \text{ yr}^{-1}$. Recalling that the large majority of radio-active AGN are also found to be FIR emitters, we can then draw the conclusion that most of the galaxies which host a radio-active AGN are also the hosts of extremely intense star-forming activity. This finding confirms and extends to the radio-active case all recent results obtained on the very high star-forming activity observed in the hosts of AGN selected in different ways and at various wavelengths (e.g. Hatziminaoglou et al. 2010; Barthel et al. 2012; Santini et al. 2012; Rosario et al. 2013; Drouart et al. 2014; Podigachoski et al. 2015, just to mention a few).

Perhaps more importantly, the data show that neither the spectral slope of the radio emission nor the radio-luminosity of the sources under examination varies in the presence of FIR emission. This is true for the spectral slope at all redshifts, and suggests that the central engine responsible for the AGN phenomenon is oblivious, at least at radio wavelengths, to the presence of ongoing star formation within the same galaxy. In other words, FIR activity does not seem to influence radio-AGN activity. This is also true for radio luminosities, at least for redshifts larger than $z \sim 1$: indeed we find that, irrespective of radio luminosity, in the more distant universe virtually all radio-emitting AGN cohabit with a star-forming event. Once again, this suggests that radio luminosity, and therefore AGN activity at radio wavelengths, is not affected by the nearby star-forming processes. The situation varies in the more local universe, as there is a deficit of FIR emitters amongst bright radio AGN. This is in accordance with the general view that powerful radio-loud AGN reside in old galaxies with little or no ongoing star-forming activity. Our results then show that this lack of star-forming activity within the hosts of radio-active AGN is only true in the nearby universe, and possibly implies that at some late stage of their evolution, and only for weak enough FIR emission, powerful enough AGN are capable to swipe off the gas in the surrounding star-forming regions and consequently switch off the on-site star formation. Note that our results are in excellent agreement with those of Best et al. (2014), which also show a clear positive evolution of radio-loud AGN which are in the 'radiative-mode' (i.e. associated with cold gas and therefore to star formation) with respect to those which are in the so-called 'jet-mode' (radiatively inefficient and generally

associated with passive galaxies) between $z = 0$ and $z = 1$. Another possible interpretation of the data is provided by recent studies of the mass function of galaxies which show that there is a transition at around $z \sim 1$ between a mass function dominated by late-type galaxies to one dominated by early-type ones (e.g. Pozzetti et al. 2010; Ilbert et al. 2013). The radio-AGN behaviour might then only mirror this more general trend derived for the whole galaxy population.

Another relevant result of our analysis is that also the general properties of the galaxies host of radio-selected AGN, such as their stellar mass or their age, do not vary in the presence of ongoing FIR activity. This means that the hosts of radio-active AGN all belong to the same population, irrespective of whether or not there is ongoing star-forming activity, i.e. that star formation is not the dominant ingredient in shaping a galaxy which also hosts a radio-active AGN.

We stress once again that the above results rely on an extremely high success rate of radio-to-FIR associations and therefore that the conclusions we can draw with the present work refer to the majority of radio-active AGN. For this class of sources no effect of star-forming origin is observed on AGN activity. All the considered properties both of the host galaxy and of the AGN itself remain unaltered in the presence of FIR emission, showing that star formation does not influence (at least radio) AGN activity or shape the physical properties of the hosting galaxies. On the other hand, most of such AGN hosts (virtually the totality for redshifts $z \gtrsim 1$) are also found to be the sites of even extremely intense star-forming activity. This indicates that also star formation activity is not affected by the presence of an AGN, in all cases except for radio-luminous sources in the local, $z \lesssim 1$, universe. We can then conclude that, except for the case of relatively local sources and high radio luminosities, no interaction is observed between the processes of AGN emission and star formation, as they seem to proceed independently of each other. Note that no clear evidence for the dominance of AGN-feedback in shaping galaxy evolution has also been found by Pozzi et al. (2015), who use a backward approach capable of reproducing the galaxy K band and the far-IR luminosity functions across the redshift range $0 < z < 3$.

Finally, we have shown that, except for a few cases of sources mainly residing at very high ($z \gtrsim 4$) redshifts, galaxies hosting radio-selected star-forming events and those hosting radio-selected AGN are indistinguishable from each other in terms of both mass and IR luminosity (or SFR) at all redshifts larger than ~ 1 . Differences only start appearing in the local universe, where radio-active AGN start progressively becoming FIR-quiet, the more as the higher is their radio luminosity, and at the same time are found associated with much larger galaxies than those which host a star-forming event.

Why differences in the various radio-selected populations considered here only appear in the local universe is still unclear and will be investigated in a forthcoming paper.

ACKNOWLEDGEMENTS

We wish to thank the referee for his/her careful reading of the manuscript and constructive suggestions which helped us in shaping and improving the paper. PACS has been developed by a consortium of institutes led by MPE (Germany) and including UVIE (Austria); KU Leuven, CSL, IMEC (Belgium); CEA, LAM (France); MPIA (Germany); INAF- IFSI/OAA/OAP/OAT, LENS, SISSA (Italy); IAC (Spain). This development has been supported by the funding agencies BMVIT (Austria), ESA-PRODEX (Belgium), CEA/CNES

(France), DLR (Germany), ASI/INAF (Italy), and CICYT/MCYT (Spain).

REFERENCES

- Alexander D. M. et al., 2008, *ApJ*, 135, 1968
 Appleton P. N. et al., 2004, *ApJS*, 154, 147
 Barthel P., Haas M., Leipski C., Wilkes B., 2012, *ApJ*, 757, L26
 Berta S. et al., 2011, *A&A*, 532, 49
 Best P. N., Ker L. M., Simpson C., Rigby E. E., Sabater J., 2014, *MNRAS*, 445, 955
 Bondi M. et al., 2003, *A&A*, 403, 857
 Bondi M., Ciliegi P., Schinnerer E., Smolcic V., Jahke K., Carilli C., Zamorani G., 2008, *ApJ*, 681, 1135
 Boyle B. J., Cornwell T. J., Middleberg E., Norris R. P., Appleton P. N., Smail I., 2007, *MNRAS*, 385, 1143
 Brinchmann J., Charlot S., White S. D. M., Tremonti C., Kauffmann G., Heckman T., Brinkman J., 2004, *MNRAS*, 351, 1151
 Brunner H., Cappelluti N., Hasinger G., Barcons X., Fabian A. C., Mainieri V., Szokoly G., 2008, *A&A*, 479, 283
 Bruzual G., Charlot S., 2003, *MNRAS*, 344, 1000
 Carilli C. L., Furlanetto S., Briggs F., Jarvis M., Rawlings S., Falcke H., 2004, *New Astron. Rev.*, 48, 1029
 Condon J. J., 1992, *ARA&A*, 30, 575
 Condon J. J., Condon M. A., Gisler G., Pushell J., 1982, *ApJ*, 252, 102
 De Breuck C. et al., 2010, *ApJ*, 725, 36
 Del Moro A. et al., 2013, *A&A*, 549, A59
 Drouart G. et al., 2014, *A&A*, 566, A53
 Elbaz D. et al., 2011, *A&A*, 533, A119
 Fanaroff B. L., Riley J. M., 1974, *MNRAS*, 167, 31
 Fotopoulou S. et al., 2012, *ApJS*, 198, 1
 Garn T., Green D. A., Riley J. M., Alezander P., 2009, *MNRAS*, 397, 1101
 Griffin M. J. et al., 2010, *A&A*, 518, L3
 Gruppioni C. et al., 2013, *MNRAS*, 432, 23
 Hatziminaoglou E. et al., 2010, *A&A*, 518L, 33
 Ibar E., Ivison R. J., Biggs A. D., Lal D. V., Best P. N., Green D. A., 2009, *MNRAS*, 397, 281
 Ibar E., Ivison R. J., Best P. N., Coppin K., Pope A., Smail I., Dunlop J. S., 2010, *MNRAS*, 401, L53
 Ilbert O. et al., 2013, *A&A*, 556, 55
 Johnston S. et al., 2007, *PASA*, 24, 174
 Jonas J. L., 2009, *IEEE Proc. 97*, 1522
 Kennicut R. C., Jr, 1998, *ARA&A*, 36, 189
 Leipski C., Antonucci R., Ogle P., Whysong D., 2009, 701, 891
 Lutz D., 2014, *ARA&A*, 52, 373
 Lutz D. et al., 2011, *A&A*, 532, 90
 Magliocchetti M. (the 2dFGRS Team) et al., 2002, *MNRAS*, 333, 100
 Magliocchetti M. (the 2dFGRS Team) et al., 2004, *MNRAS*, 350, 1485
 Magliocchetti M., Maddox S. J., Wall J. V., Benn C. R., Cotter G., 2000, *MNRAS*, 318, 1047
 Magliocchetti M., Andreani P., Zwaan M. A., 2008, *MNRAS*, 383, 479
 Magliocchetti M. et al., 2014, *MNRAS*, 442, 682
 Magnelli B. et al., 2013, *A&A*, 553, A132
 Mao M. Y., Huynh M. T., Norris R. P., Dickinson M., Frayer D., Helou G., Monkiewicz J. A., 2011, *ApJ*, 731, 79
 Mauch T., Sadler E. M., 2007, *MNRAS*, 375, 931
 McAlpine K., Jarvis M. J., Bonfield D. G., 2013, *MNRAS*, 436, 1084
 Merloni A., Rudnick G., Di Matteo T., 2004, *MNRAS*, 354, L37
 Metcalf R. B., Magliocchetti M., 2006, *MNRAS*, 365, 101
 Miller N. A. et al., 2013, *ApJS*, 205, 13
 Morrison G. E., Owen F. N., Dickinson M., Ivison R. J., Ibar E., 2010, *ApJS*, 188, 178
 Norris R. P. et al., 2011, *ApJ*, 736, 55
 Peng Y.J. et al., 2010, *ApJ*, 721, 193
 Pilbratt G. et al., 2010, *A&A*, 518, L1
 Podigachoski P. et al., 2015, *A&A*, 575, A80
 Poglitsch A. et al., 2010, *A&A*, 518, L2
 Pozzetti L. et al., 2010, *A&A*, 517, 11
 Pozzi F. et al., 2015, *ApJ*, 803, 35
 Prandoni I., Gregorini L., Parma P., de Ruiter H. R., Vettolani G., Zanichelli A., Wieringa M. H., Ekers R. D., 2001, *A&A*, 369, 787
 Randall K. E., Hopkins A. M., Norris R. P., Zinn P.-C., Middleberg E., Mao M. Y., Sharp R. G., 2012, *MNRAS*, 421, 1644
 Rodighiero G. et al., 2011, *ApJ*, 739, L40
 Rosario D. et al., 2013, *A&A*, 560, 72
 Salpeter E. E., 1955, *ApJ*, 121, 161
 Santini P. et al., 2009, *A&A*, 504, 751
 Santini P. et al., 2012, *A&A*, 540, 109
 Schinnerer E. et al., 2004, *AJ*, 128, 1974
 Schinnerer E. et al., 2007, *ApJS*, 172, 46
 Schinnerer E. et al., 2010, *ApJS*, 188, 384
 Seymour N. et al., 2011, *MNRAS*, 413, 1777
 Simpson C. et al., 2012, *MNRAS*, 421, 3060

This paper has been typeset from a *T_EX/L_AT_EX* file prepared by the author.