The distinct build up of dense and normal passive galaxies in VIPERS

ADRIANA GARGIULO & THE VIPERS TEAM

Paris – Galaxy evolution across the cosmic time – 12/06/2017

THE OUTLINE OF THE TALK

- Why studying massive ($M_{star} > 10^{11} M_{sun}$) passive galaxies (MPGs)
- The evolution of the number density and of the stellar population ages of MPGs at $0.5 < z < 1.0$
- The impact of the environment on the mass assembly history of MPGs

WHY STUDYING MASSIVE PASSIVE GALAXIES (MPGs)

Most of the stellar mass in galaxies today resides in massive passive systems (e.g. Renzini 2006).

Still unclear *WHEN & WHERE* the stars in massive galaxies were formed

THE 'MERGER' (EX SITU) HYPOTHESIS

In a ΛCDM Universe the assembly of massive galaxies is dominated by the (dry) accretion of stars formed in other galaxies

(e.g. Naab+ 2009; Qu+ 2017; Rodriguez-Gomez+ 2016; Gabor&Davè 2012; Lackner+ 2012 / e.g. Robertson+ 2006; Cox+ 2006; Pipino+ 2008)

THE IN SITU HYPOTHESIS

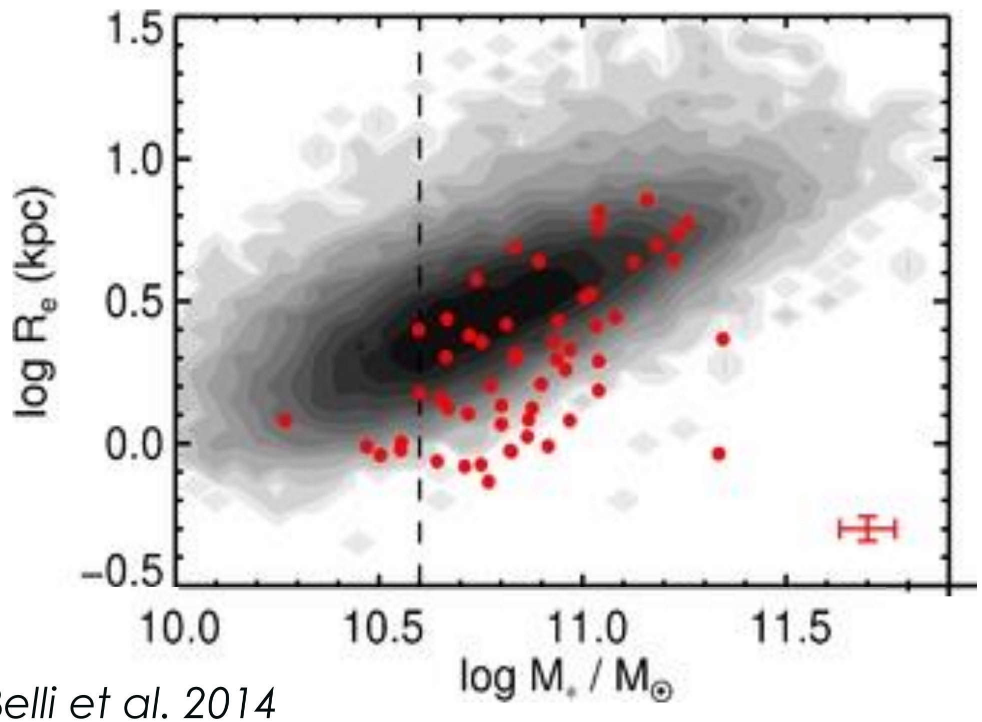
Stars in massime galaxies come from in situ star formation (untill an event – e.g. AGN feedback, halo shock, disc fragmantation – does not stop it)

(Genzel+ 2008, 2011; Förster Schreiber+ 2011; Tacconi+ 2013; Wuyts+ 2013)

STRUCTURAL PROPERTIES OF LOCAL AND HIGH-Z MPGs

THE EX SITU HYPOTHESIS vs. THE IN SITU HYPOTHESIS
Do we need a combination of these?

Evolutionary models have to reproduce MPGs properties as the size-mass relation and its evolution with time



Belli et al. 2014

At fixed stellar mass

- local MPGs have dimensions that vary up to an order of magnitude;
- high-z (z ~ 2) MPGs are smaller than local MPGs by a factor ~ 5

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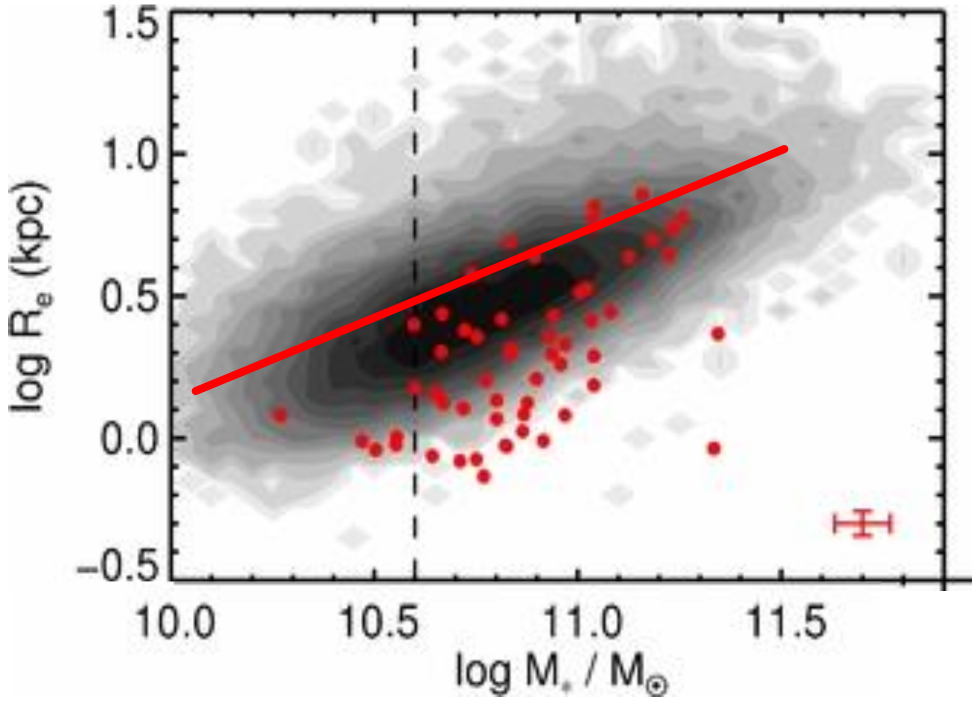
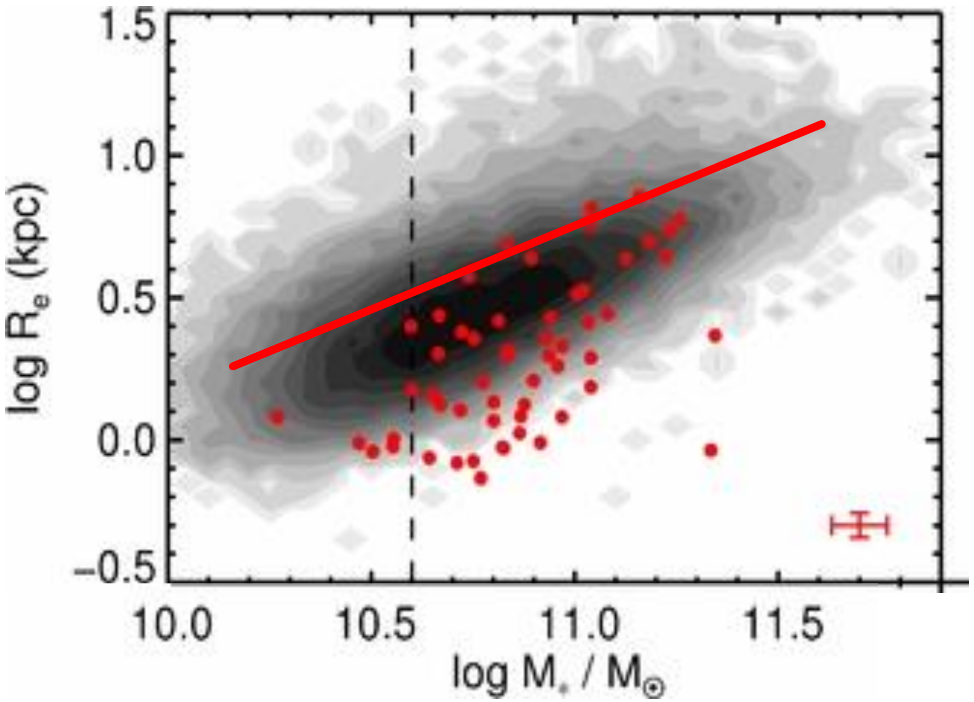


Given the mean stellar mass density
$\Sigma = M_* / (2 \pi R_e^2)$

on average
high – z MPGs were denser than local MPGs

THE OPEN QUESTIONS



- Why, at fixed stellar mass, the typical dimension of a MPGs varies up to an order of magnitude?
- Have dense and less dense local MPGs different stellar mass assembly history?

In which way was build up the population of local MPGs?

TO FIND THE ANSWER WE STUDIED:

1. the evolution of the number density of MPGs as a function of Σ
2. the evolution of the stellar population ages of MPGs as a function of Σ
3. the correlation of Σ and the local environment

USING VIPERS

VIMOS PUBLIC EXTRAGALACTIC REDSHIFT SURVEY

VIPERS IN A NUTSHELL

ESO LARGE PROGRAM (PI: L. Guzzo)

SPECTRA :

LRR grism (R = 200) → [5500 – 9500] A → $\Delta z = 0.00047(1 + z)$

TARGET SAMPLE :

- i(AB) < 22.5 on the CFHTLS Wide W1 and W4 fields
- ugri colour pre-selection → z > 0.5

VOLUME :

$5 \times 10^7\ h^{-3}\ Mpc^3$ - comparable to 2DF but at redshift ~ 1

AREA :

24 sq. dg (~ 16 without gaps) – 40% sampling rate

PHOTOMETRY :

NUV, u, g, r, i, z, K ++ (Moutard et al. 2016)

SURVEY STATUS AS OF 30/11/2015

EFFECTIVE TARGETS	MEASURED REDSHIFTS	STELLAR CONTAMINATION	COVERED AREA
93252	88901	2265 (2.5 %)	100.0 %

Now available at http://vipers.inaf.it/

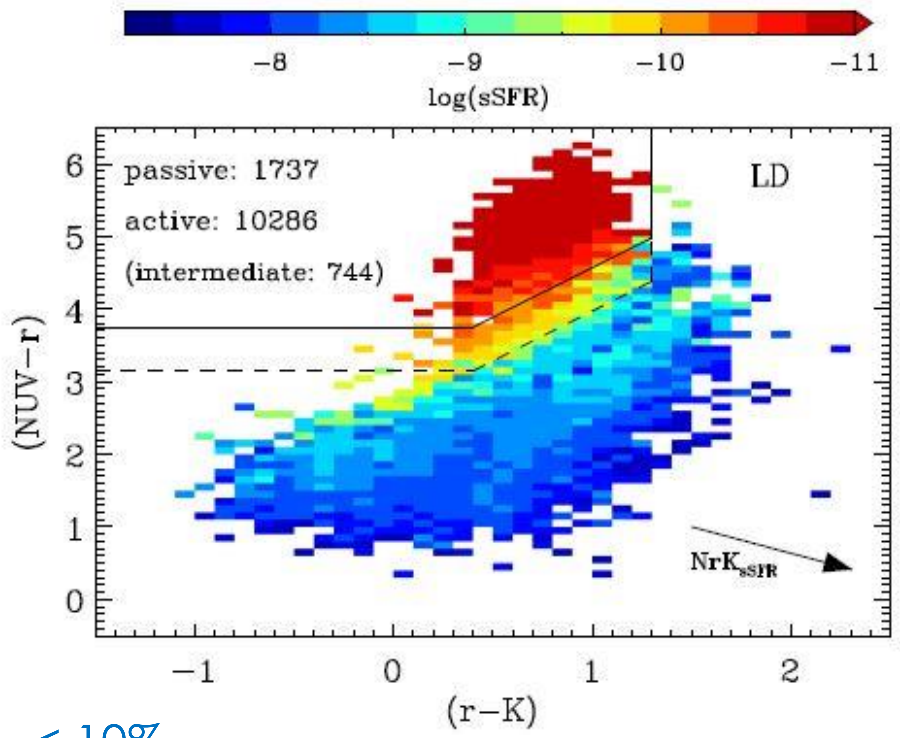
0.5 < z < 1.2

THE VIPERS SAMPLE OF MPGs

The selection criteria:

- $2 \leq z_{flag} \leq 9.0$
- $0.5 \leq z \leq 1.0$
- (NUV – r) vs (r –K)
- $M_{star} \geq 10^{11}$ M_{sun} (Cha IMF)

~2000 MPGs at 0.5 < z < 1.0 with z-spec



For structural parameters in VIPERS see Krywult+ 2016, $\delta R_e < 10\%$

Davidzon et al. 2013, 2016

To derive Σ: Re in i band for galaxies at $z \geq 0.8$ / Re r band for galaxies at $z < 0.8$ } ~ U band rest frame over the whole redshift range

High-Σ MPGs → $\Sigma > 2000 M_{sun}\ pc^{-2}$ --- Intermediate- Σ MPGs → $1000 < \Sigma \leq 2000 M_{sun}\ pc^{-2}$

Low- Σ MPGs → $\Sigma \leq 1000 M_{sun}\ pc^{-2}$

THE EVOLUTION OF THE NUMBER DENSITY OF MPGs AS A FUNCTION OF Z AND Σ

The evolution of the number densities depends on Σ :
the lower the Σ, the faster the evolution

From z = 1.0 → z = 0.5:

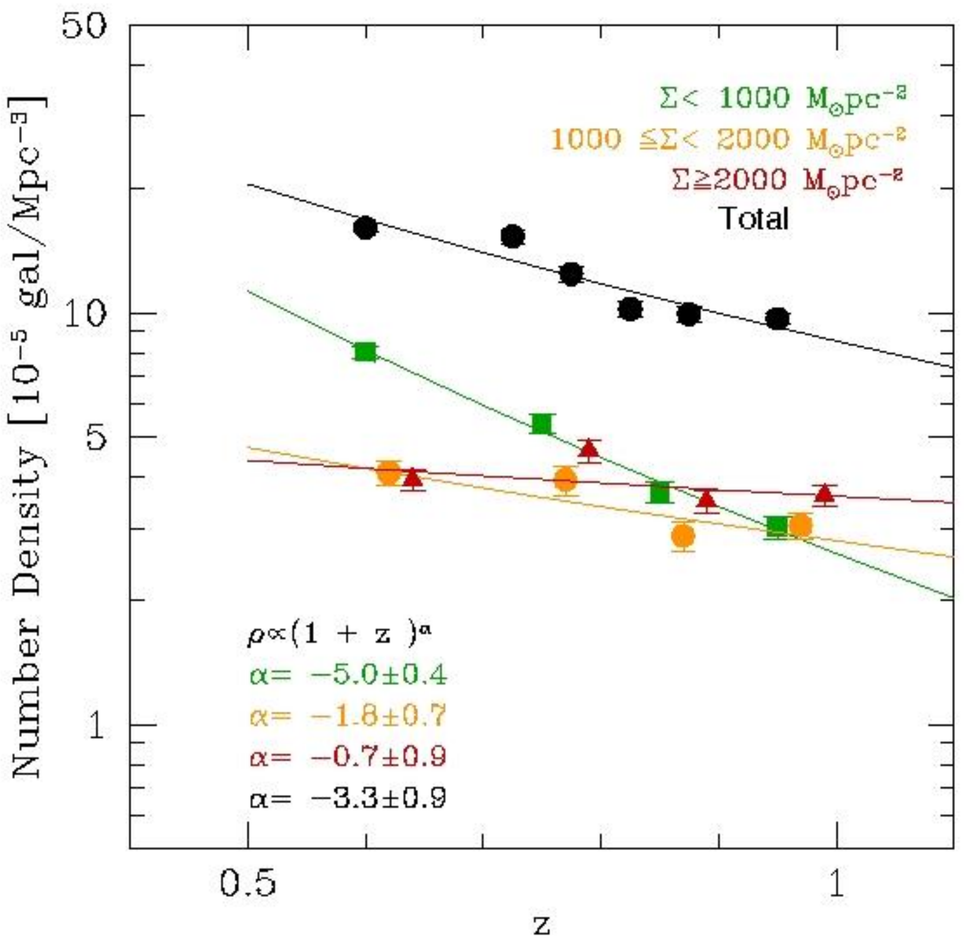
Increase factor:

Total : ~ 2.5

High Σ : ~ 1.2

Int Σ : ~ 1.7

Low Σ : ~ 4.2



Number densities fully corrected for incompleten
Errors take into account the Poisson fluctuations and the error on Re

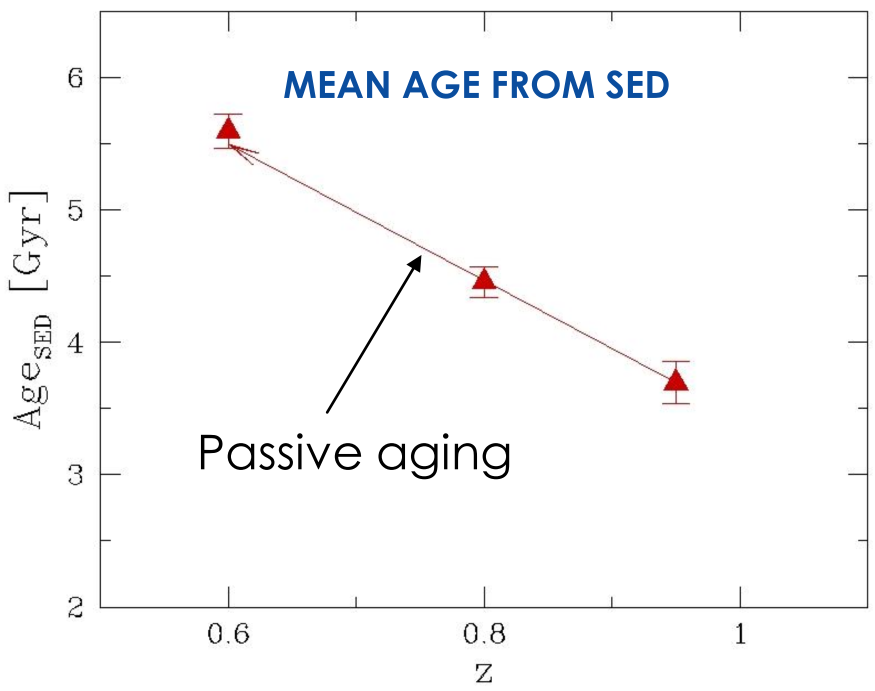
Gargiulo et al. 2017, in press

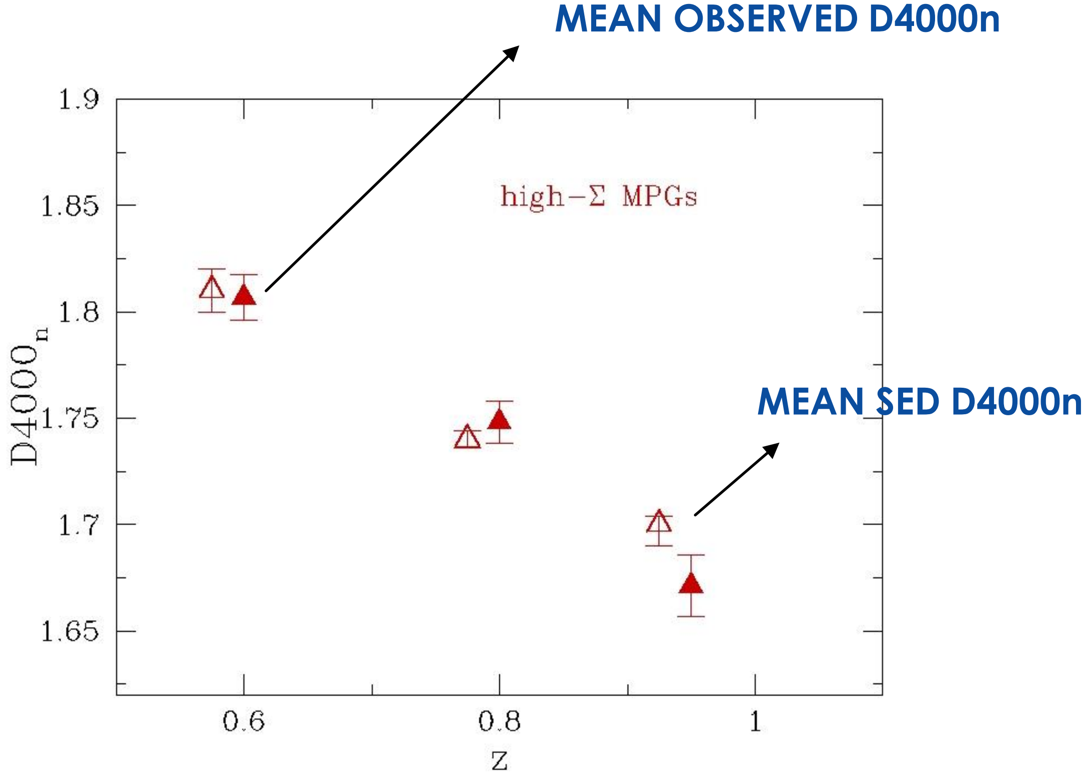
THE EVOLUTION OF THE STELLAR POPULATION AGES OF MPGs AS A FUNCTION OF Z AND Σ

The approach:

1. ages derived from the SED fitting → Mean Age/Z/Tau (z, Σ);
2. Mean Age/Z/Tau (z, Σ) + BC03 models → $D4000_{SED}$ (z, Σ);
3. $D4000_{SED}$ (z, Σ) vs $D4000_{obs}$(z, Σ)

Compact MPGs

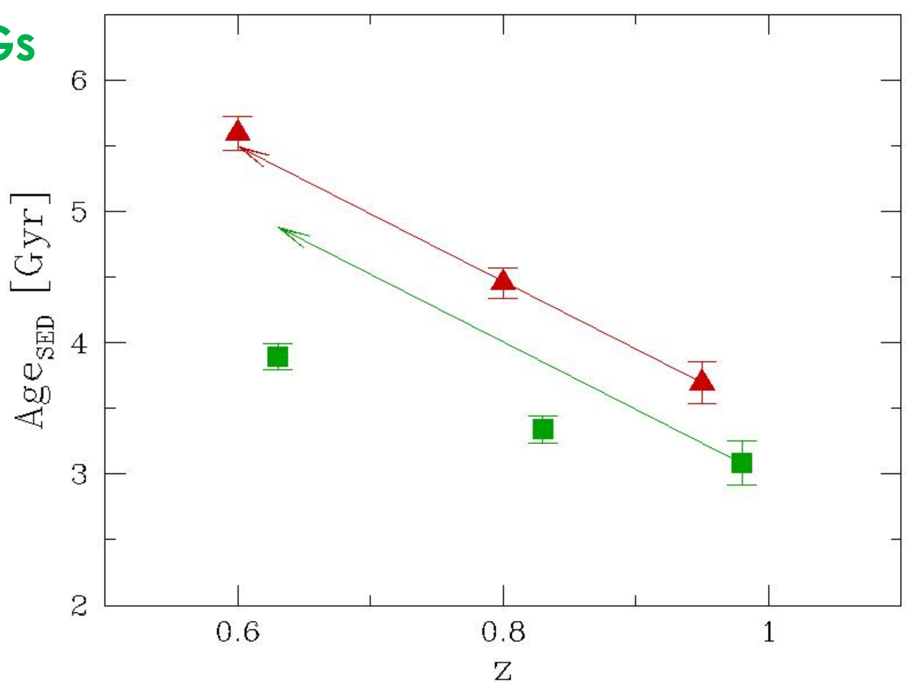


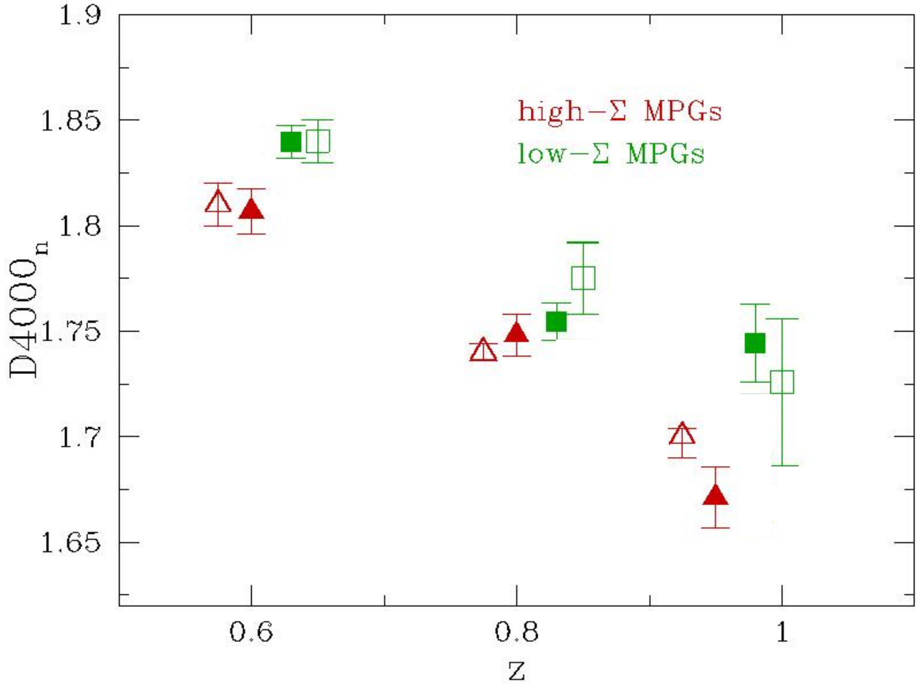


The evolution both of the number density and of the mean age of dense MPGs show that they passively evolve

THE EVOLUTION OF THE STELLAR POPULATION AGES OF MPGs AS A FUNCTION OF Z AND Σ

Less dense MPGs





Dense MPGs are older than less dense ones

(see also, e.g. , Poggianti+ 2013, Saracco+ 2010, Williams+ 2016, Fagioli+ 2016 at smaller M_{star})

The evolution of the number density and of the mean age of less dense MPGs show that a significant fraction of NEW and YOUNGER MPGs should appear at later epoch

CONCLUSIONS - 1

From redshift 1.0 to 0.5
the population of MPGs (mainly) grows bottom – up:

on top of the population of denser MPGs already in place at $z \sim 1.0$,
new, younger, and larger MPGs appear at lower z

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Where do these new MPGs come from?

THE PROGENITORS OF (LESS DENSE) MPGs

The increase in number density of MPGs at z < 0.8 is totally accounted for by the decrease in number density of active massive galaxies

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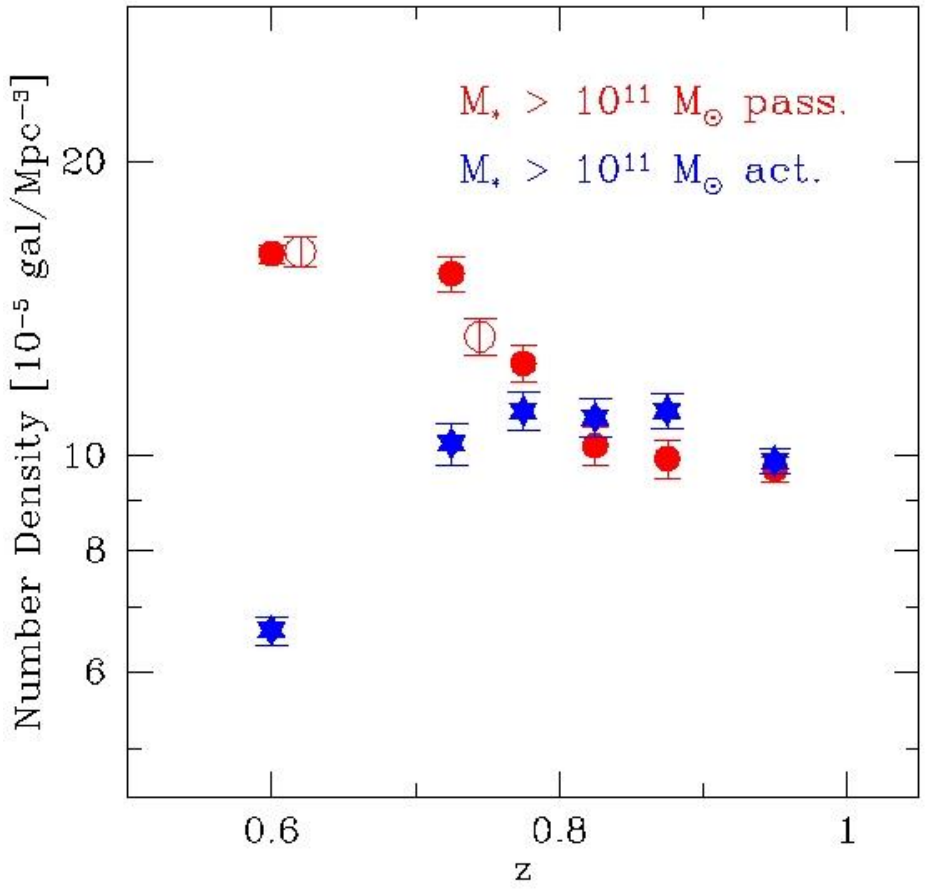
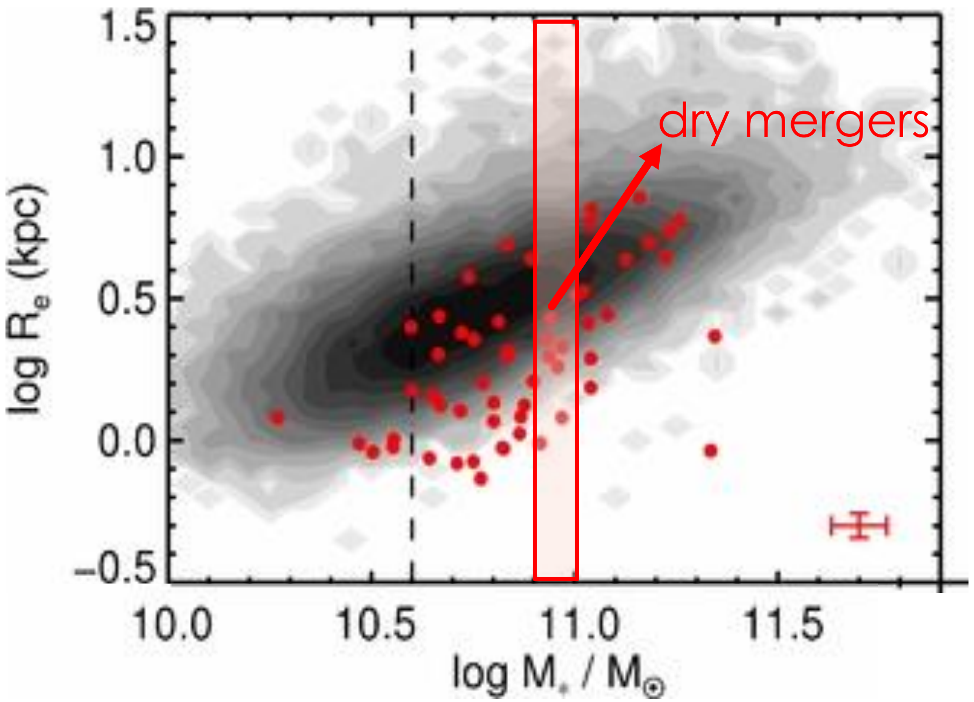
Active massive galaxies most likely progenitors of MPGs

(see, e.g. Lilly & Carollo 2016).

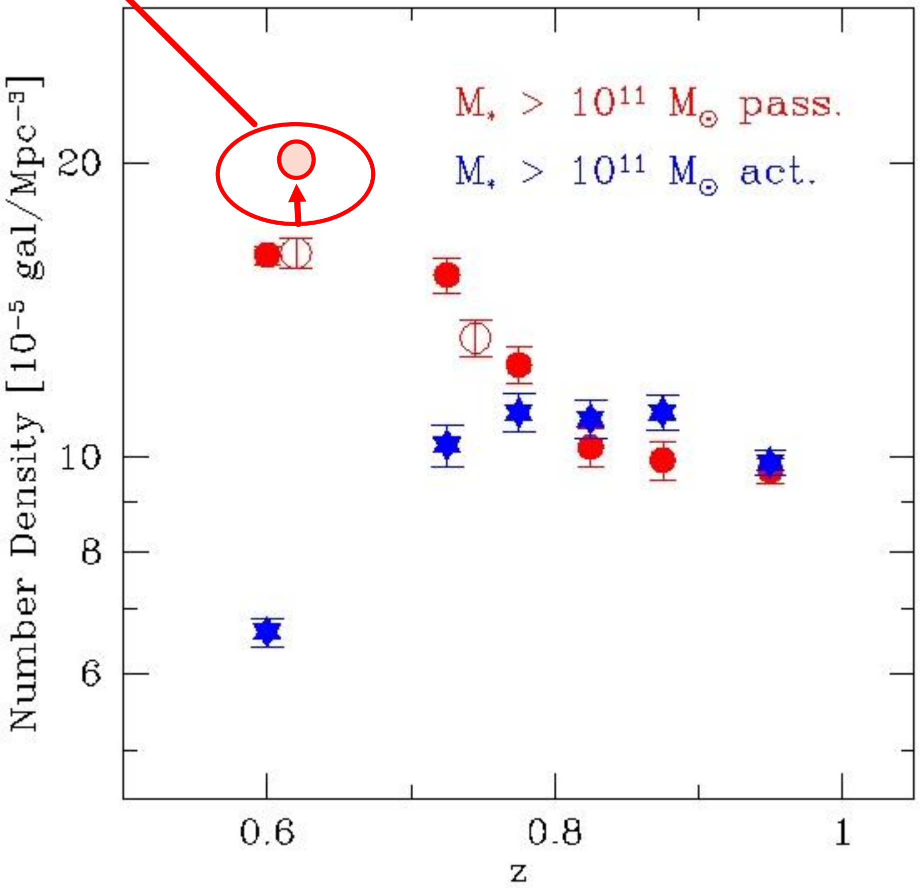


THE PROGENITORS OF (LESS DENSE) MPGs

Any other channel of MPGs formation different from the quenching of active massive galaxies will result in an overabundance of the number of MPGs



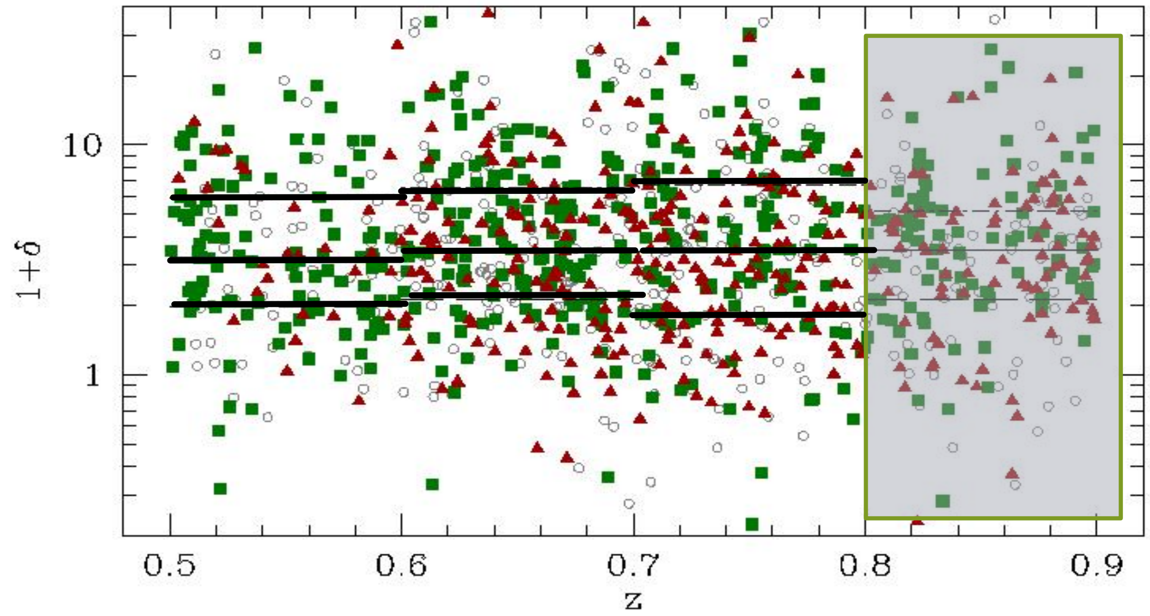
Increase in the number density due to other MPGs formation channels, e.g. dry mergers



THE IMPACT OF THE ENVIRONMENT ON Σ

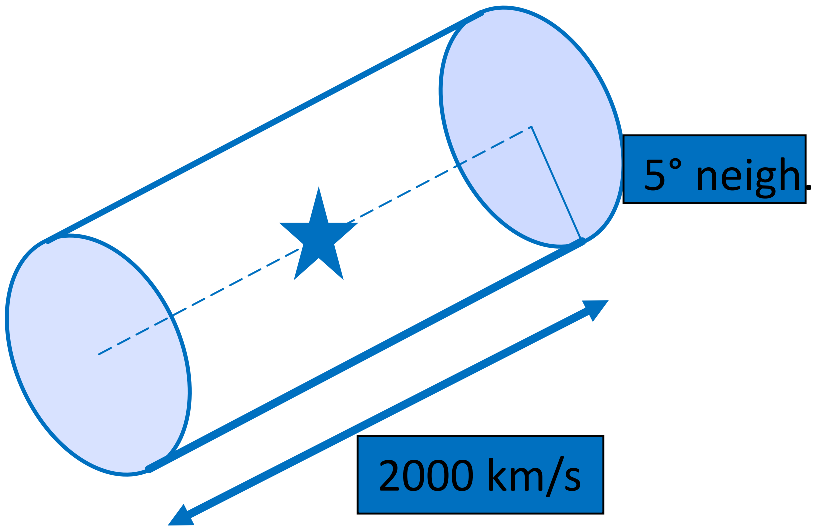
Dry mergers increase the galaxy size + merger activity enhanced in higher density regions

Larger galaxies in denser environment



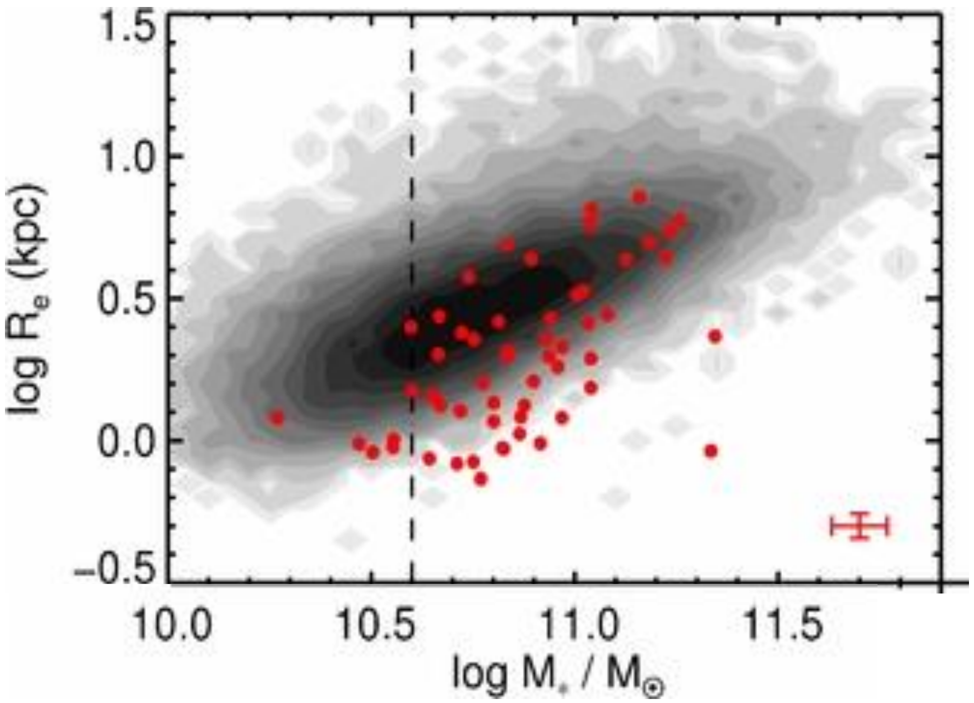
$\delta = [\rho(ra, dec, z)/\rho(z)] - 1$

Cucciati et al. 2014, 2017



THE IMPACT OF THE STELLAR MASS DISTRIBUTION ON THE δ vs. Σ CORRELATION

Positive corralation bewteen M_{star} and R_e



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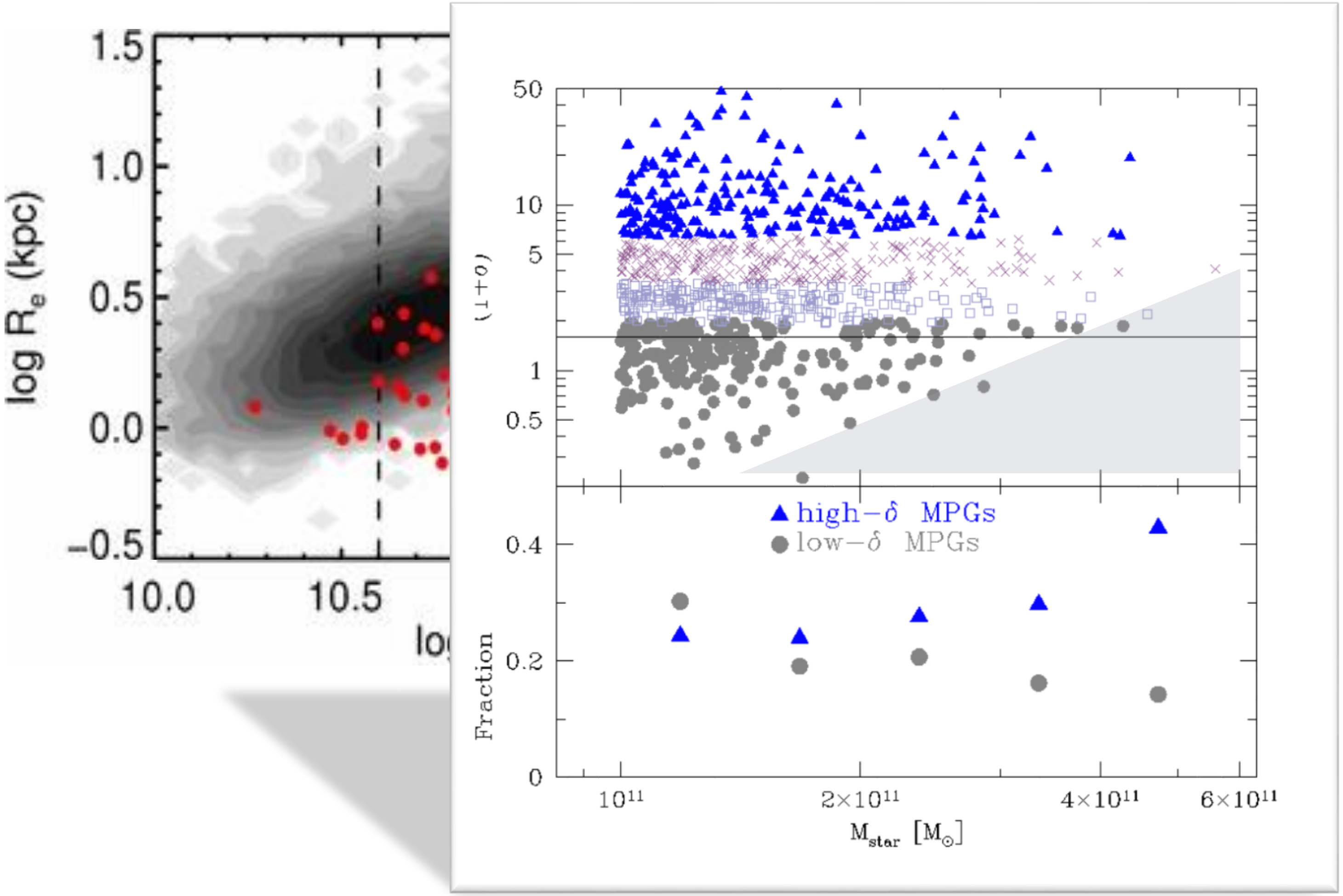
Positive corralation bewteen M_{star} and R_e & δ

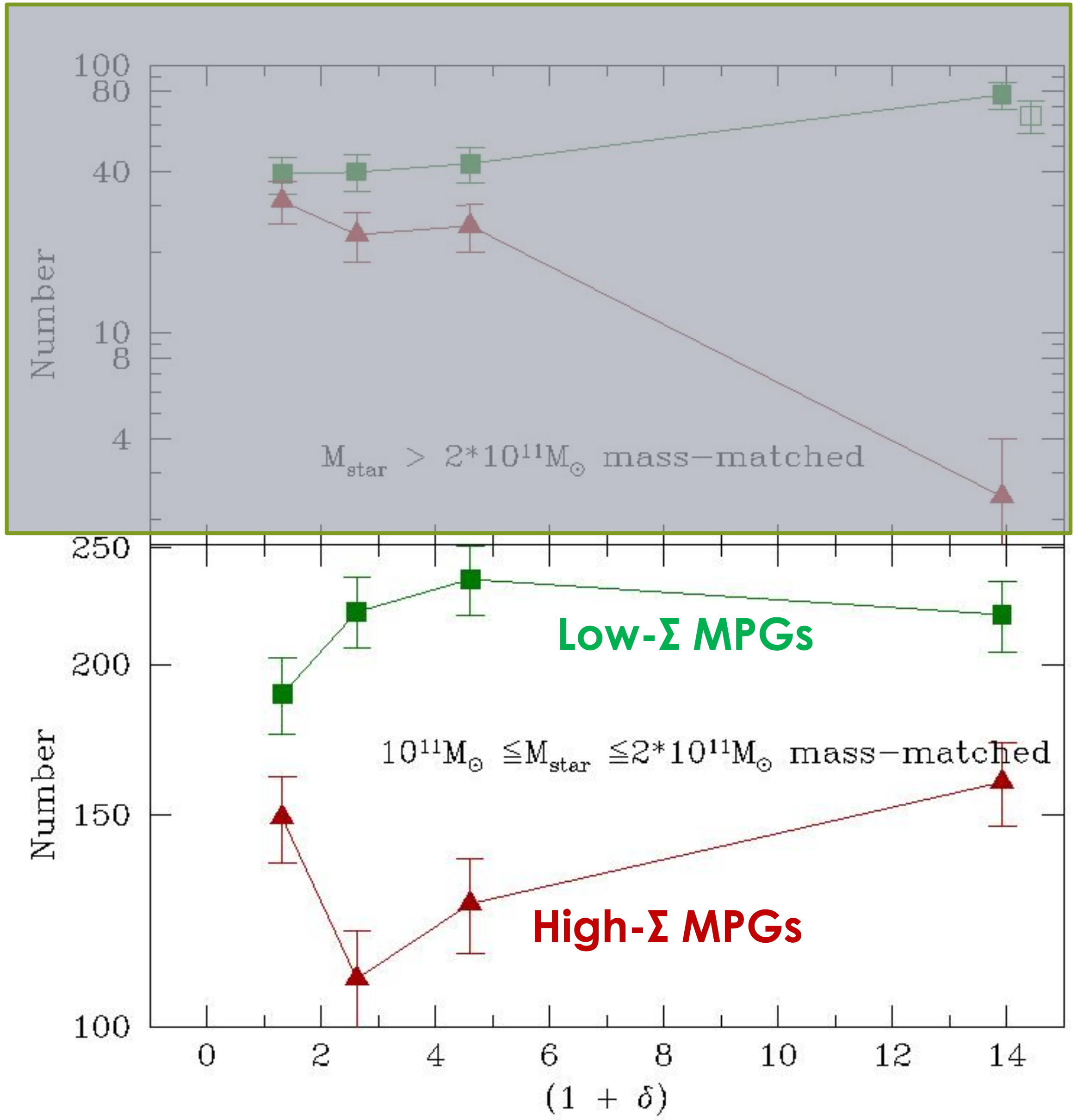


Spurious positive correlation between δ and Σ

Mass-matched samples

Gargiulo et al. 2017, in preparation

(PRELIMINARY) RESULTS



At $M_{star} < 2*10^{11}M_{sun}$ we do not find significative evidence of a trend between δ and Σ

Gargiulo et al. 2017, in preparation

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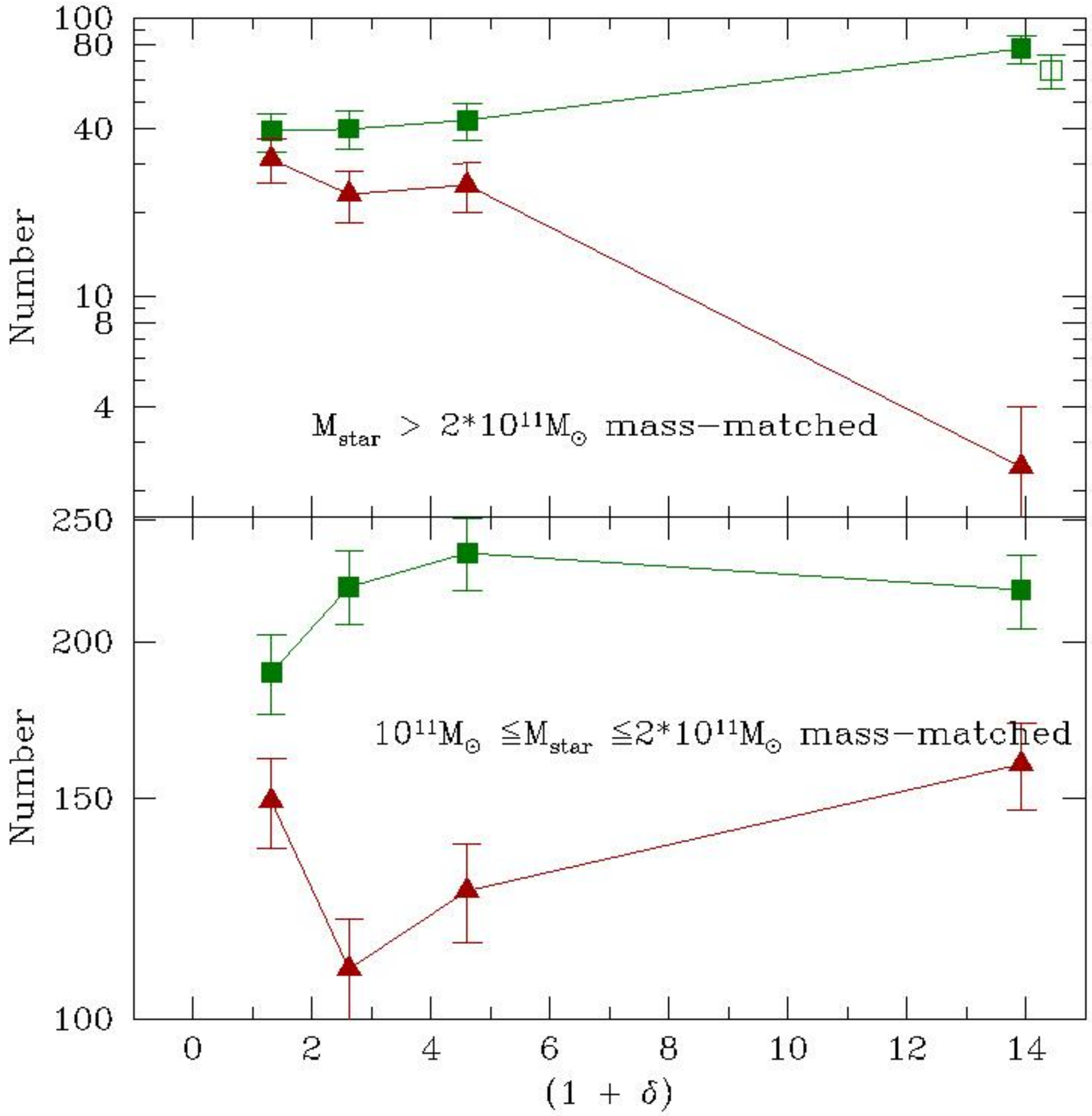

Lack(abundance) of compact(less dense) verymassive PGs in the densest regions:

1. High density regions prevent(favor) the formation of compact(less dense) MPGs with $M_{star} > 2*10^{11}M_{sun}$

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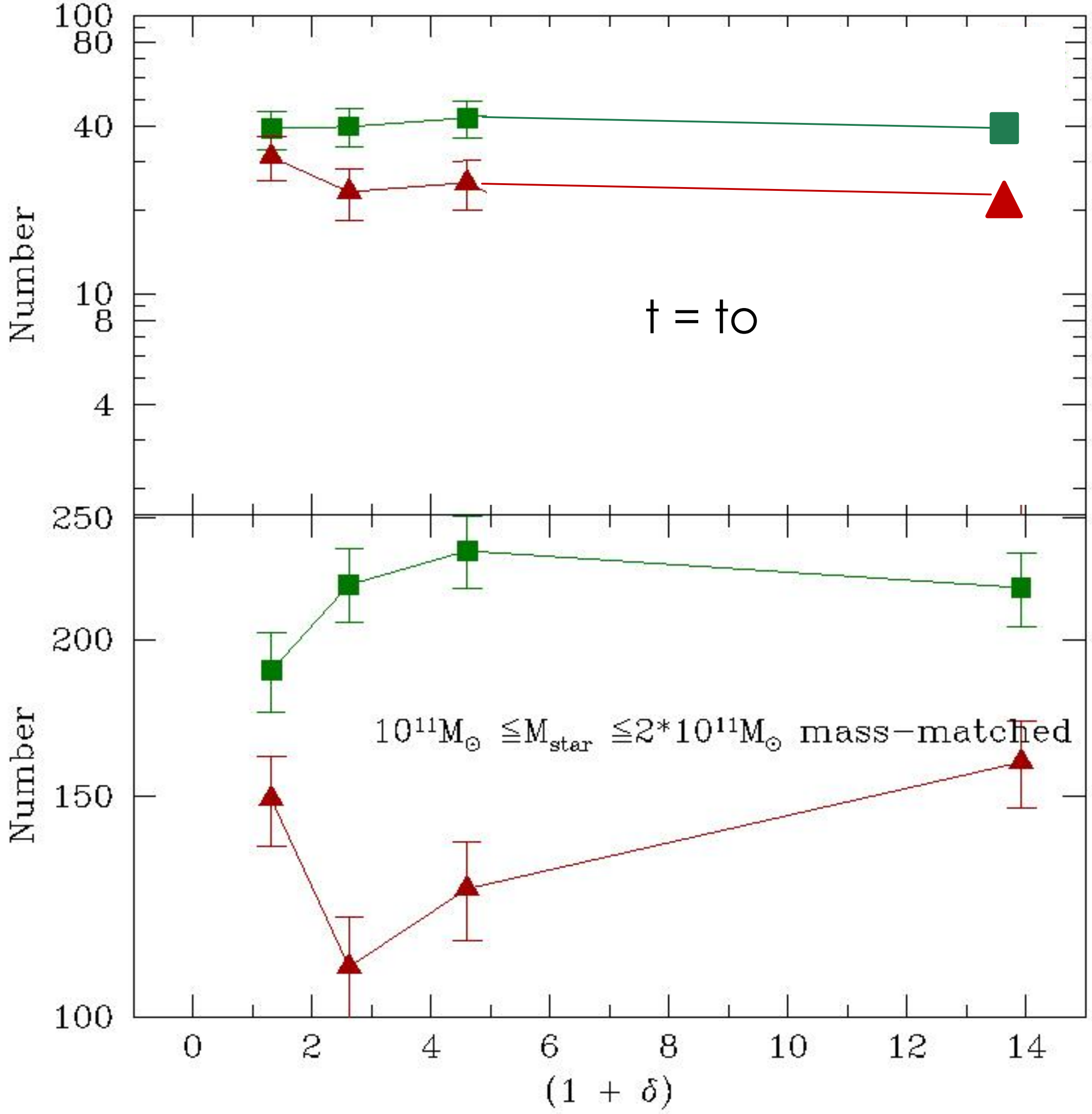

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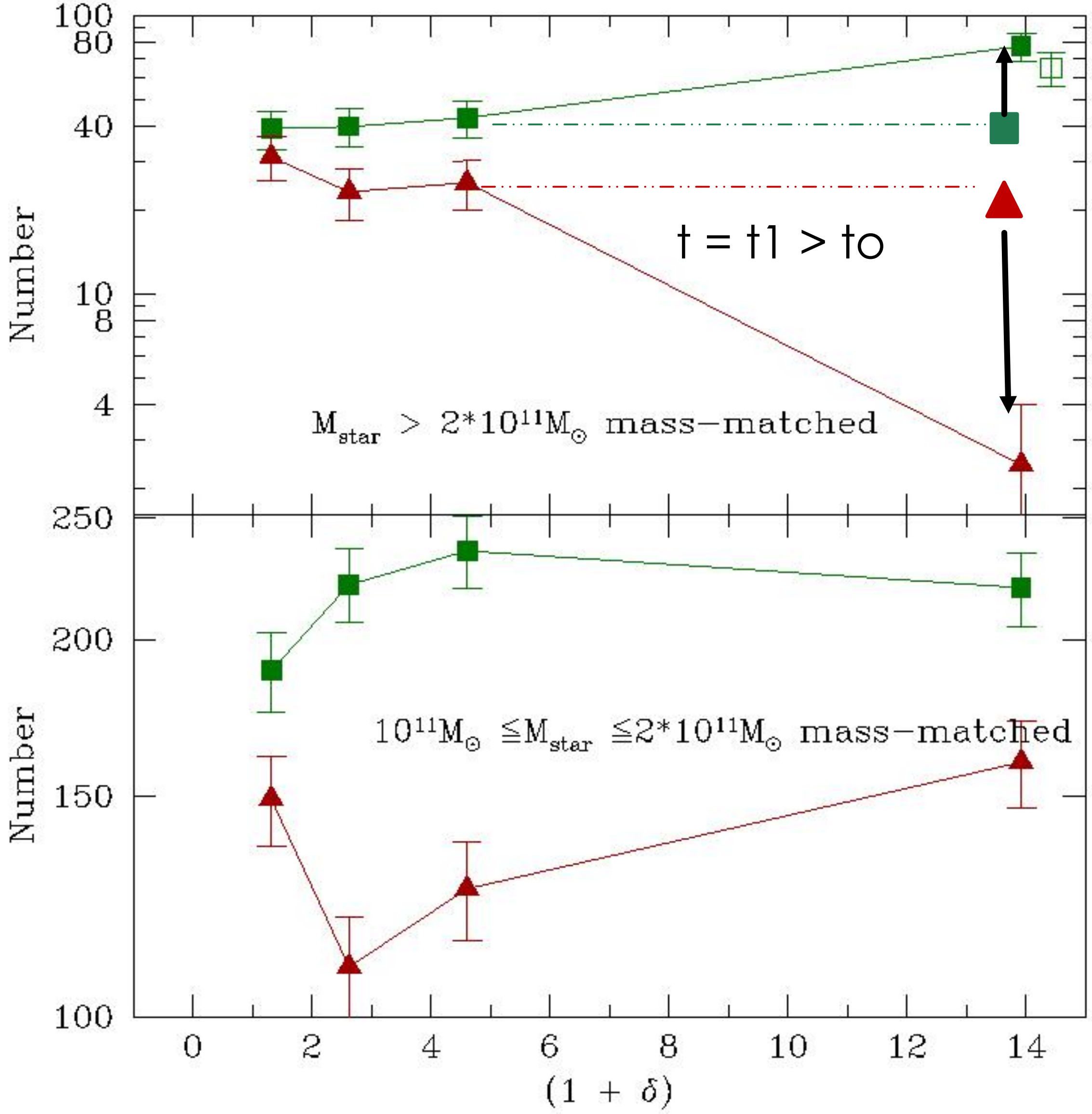

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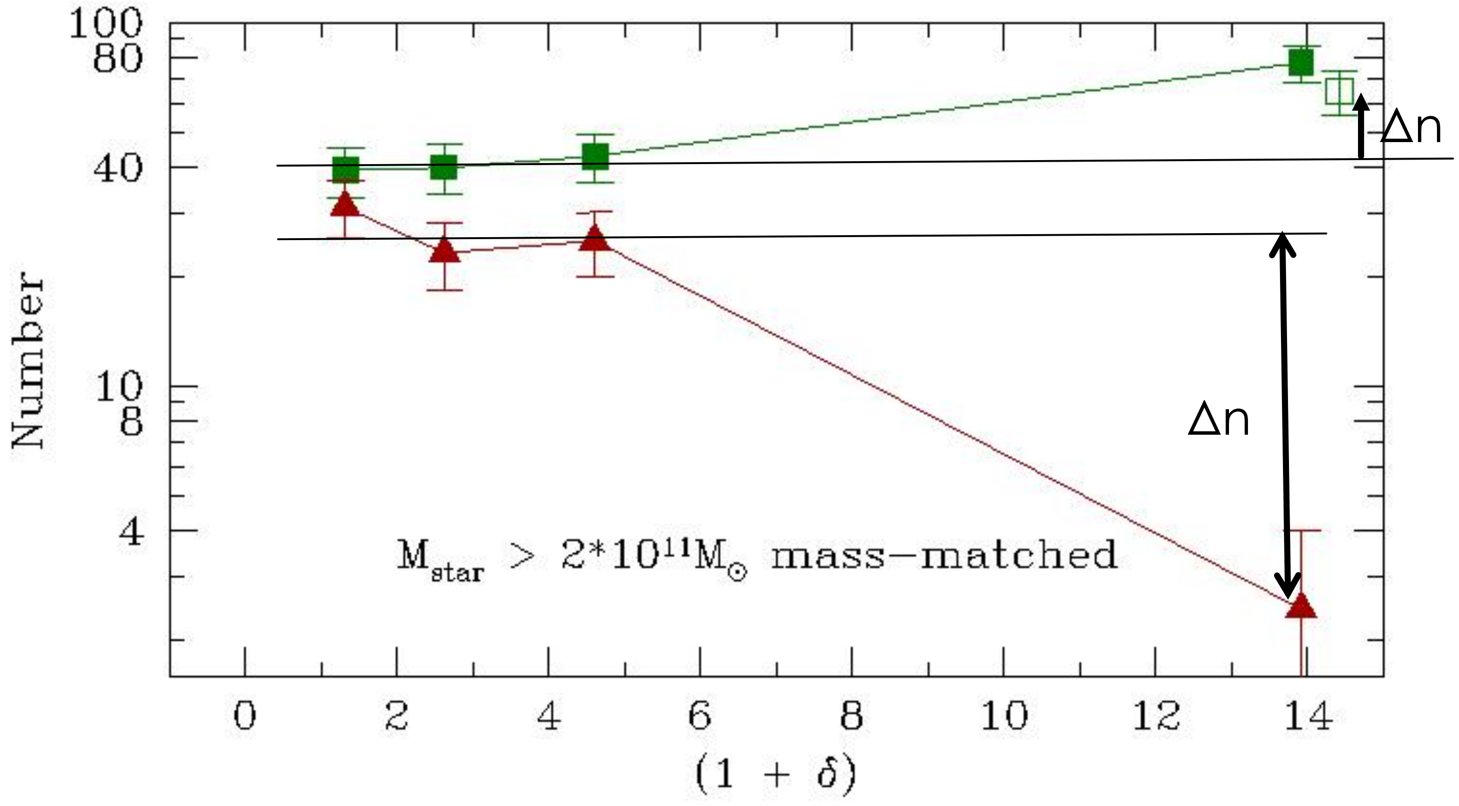

The drop in the number of compact MPGs in the highest desnity regions

=

to the increase in the number of less dense MPGs in the highest density regions

BCGs 'creation'?

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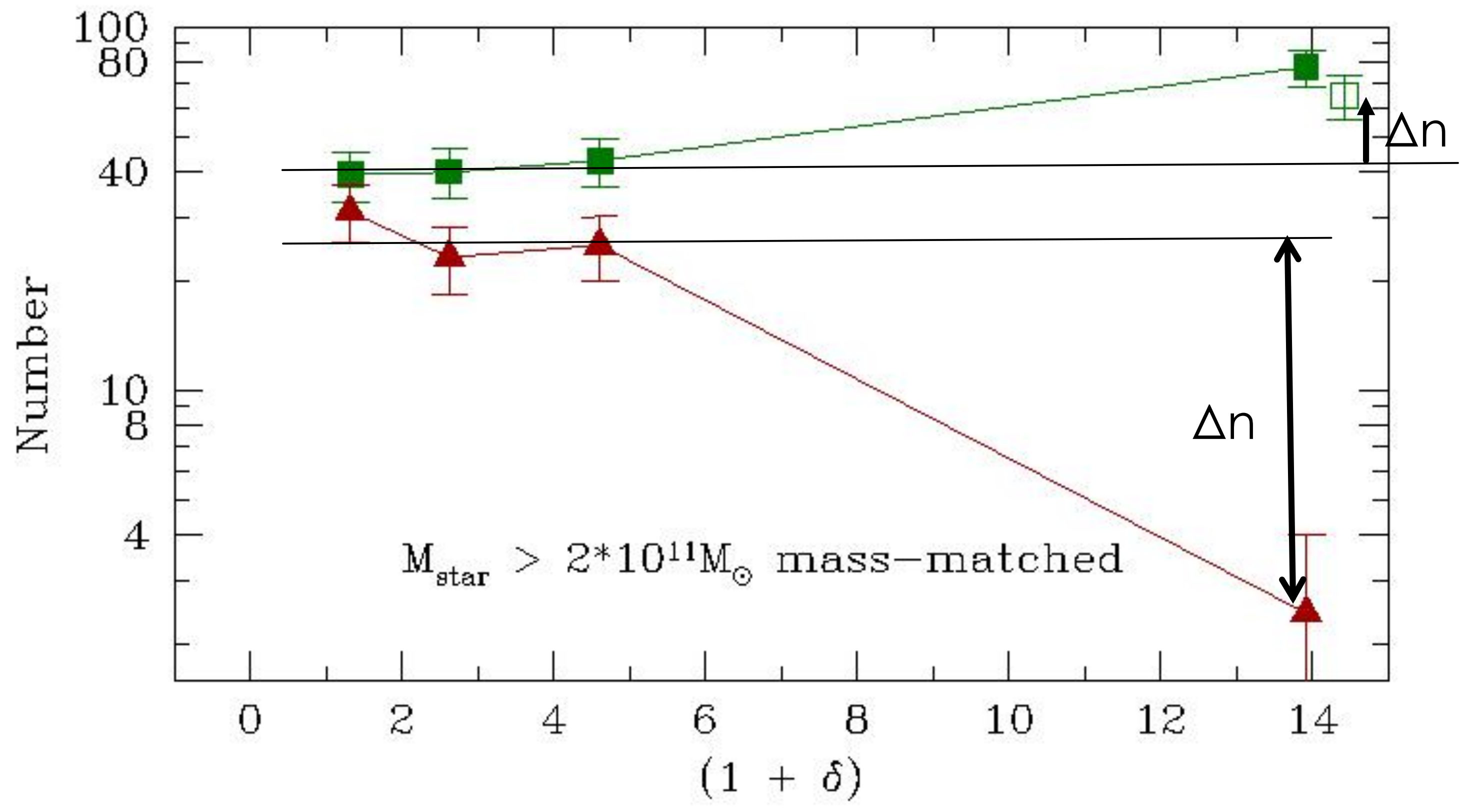

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Thank you!