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Metal enrichment signatures of the first stars on high- z DLAs

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ABSTRACT

We use numerical N -body hydrodynamical simulations with varying PopIII stellar models to investigate the possibility of detecting first star signatures with observations of high-redshift damped Ly α absorbers (DLAs). The simulations include atomic and molecular cooling, star formation, energy feedback and metal spreading due to the evolution of stars with a range of masses and metallicities. Different initial mass functions (IMFs) and corresponding metal-dependent yields and lifetimes are adopted to model primordial stellar populations. The DLAs in the simulations are selected according to either the local gas temperature (temperature selected) or the host mass (mass selected). We find that 3 per cent (40 per cent) of mass (temperature)-selected high- z ($z \geq 5.5$) DLAs retain signatures of pollution from PopIII stars, independent of the first star model. Such DLAs have low halo mass ($< 10^{9.6} M_{\odot}$), metallicity ($< 10^{-3} Z_{\odot}$) and star formation rate ($< 10^{-1.5} M_{\odot} \text{ yr}^{-1}$). Metal abundance ratios of DLAs imprinted in the spectra of quasi-stellar object can be useful as tools to infer the properties of the polluting stellar generation and to constrain the first star mass ranges. Comparing the abundance ratios derived from our simulations to those observed in DLAs at $z \geq 5$, we find that most of these DLAs are consistent within errors with PopII star dominated enrichment and strongly disfavour the pollution pattern of very massive first stars (i.e. 100–500 M_{\odot}). However, some of them could still result from the pollution of first stars in the mass range [0.1, 100] M_{\odot} . In particular, we find that the abundance ratios from SDSS J1202+3235 are consistent with those expected from PopIII enrichment dominated by massive (but not extreme) first stars.

Key words: stars: Population III – galaxies: abundances – galaxies: high-redshift – early Universe.

1 INTRODUCTION

One of the main challenges in studying primordial star formation is the lack of information on the first generation of structures. The typical environment in which early star formation happens has been largely explored in the literature (e.g. Wise et al. 2012, 2014; Biffi & Maio 2013; Dayal et al. 2013; de Souza et al. 2013, 2014; Salvaterra et al. 2013; Maio et al. 2016, and references therein); however, the chemical details of the first galaxies remain elusive.

While direct observations of the first stars are very unlikely with current and planned instruments,¹ indirect signatures could be exploited to infer the properties of pristine stellar populations. Precious

tools in this respect are represented by the abundance patterns from (carbon-enhanced) extremely metal-poor stars (Spite et al. 2013) to gamma-ray bursts (Wang et al. 2012; Maio & Barkov 2014; Ma et al. 2015, 2017). Indeed, stars with a range of masses and metallicities enrich the surrounding medium with different chemical signatures and on different time-scales (Heger & Woosley 2002, 2010). Therefore, a deep investigation of atomic metal abundances in primordial gas clouds could unveil important properties of stellar populations at high redshift (z).

Techniques based on damped Ly α absorbers (DLAs) have also been proposed to detect the imprint of Population III (PopIII) stellar enrichment in their metal content (Salvadori & Ferrara 2012; Kulkarni et al. 2013, 2014; Maio, Ciardi & Müller 2013a), in particular, in systems with halo masses smaller than $10^9 M_{\odot}$. DLAs are

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¹ Very massive first stars, as predicted by several studies (Abel, Bryan & Norman 2002; Yoshida et al. 2006; Clark et al. 2011; Stacy & Bromm 2014; Ishiyama et al. 2016), might be observable with the next-generation space

telescope JWST (for a review, see Bromm 2013), while primordial neutral gas could be detected by the future SKA mission (Koopmans et al. 2015).

dense gas absorbers with very high neutral hydrogen column density, i.e. $N_{\text{HI}} > 2 \times 10^{20} \text{ cm}^{-2}$ (Wolfe, Gawiser & Prochaska 2005), usually found along the line of sight of powerful quasi-stellar objects (QSOs). Because of the large content of neutral material, they are expected to have a kinetic temperature below 10^4 K (Foltz, Chaffee & Weymann 1986; Smith, Cohen & Bradley 1986; Wolfe et al. 1986) and to host significant amounts of molecules (Noterdaeme et al. 2008). DLA systems are widely detected from $z = 0$ to 7, while only one observed at $z \approx 2.34$ (Cooke et al. 2011) presents a carbon-enriched signature which might be associated with metal-free supernovae of mass $\approx 25 M_{\odot}$ (Kobayashi, Tomonaga & Nomoto 2011). Recently, signatures of heavy elements have been detected in DLAs at $z \geq 5$ (Becker et al. 2012; Morrison et al. 2016). Their metal abundance ratios could be safely derived from ionic column densities. Absolute metallicities cannot be measured directly, although they are expected to be less than 1 per cent solar (Maio et al. 2013a; Keating et al. 2014, 2016; Maio & Tesfari 2015). Compared to local observations, e.g. metal-poor stars in the Milky Way, faint dwarf galaxies and low- z DLAs, the detection of high- z DLAs would provide stronger evidence of the first stars, thanks to the more significant PopIII contribution to the cosmic star formation rate density and metal enrichment at those epochs (Maio et al. 2010).

In this paper, we use 3D numerical hydrodynamical chemistry simulations with different PopIII IMFs to study how gas pollution in primordial gas (in particular DLAs) changes with the properties of the first stars. Despite former numerical studies based on general lower- z DLA properties (Finlator et al. 2013, 2015) or on semi-analytical modelling (Kulkarni et al. 2013, 2014), there has been no attempt yet to investigate this issue by employing detailed 3D numerical chemistry simulations, as we plan to do here.

Throughout this paper, the standard flat Λ CDM cosmological model is adopted with the following parameters: $\Omega_{0,\Lambda} = 0.7$, $\Omega_{0,M} = 0.3$, $\Omega_b = 0.04$, $h = 0.7$, $\sigma_8 = 0.9$ and $n = 1$. The simulations used are described in Section 2, the DLA definition and properties are presented in Section 3, the results for simulated metal abundances of galaxies are shown in Section 4 and the conclusions are discussed in Section 5.

2 SIMULATIONS

In this paper, we employ the code extensively described in Maio et al. (2010, 2011b, 2013b) and the simulations run in Ma et al. (2017). Here, we mention only their main characteristics and we refer the reader to the original papers for more details and tests. The simulations were performed by using an extended version of the GADGET-2 code (Springel 2005) that includes atomic and molecular non-equilibrium chemistry, resonant and fine-structure cooling (according to the implementation of Maio et al. 2007), star formation, stellar evolution and metal enrichment (Tornatore et al. 2007; Maio et al. 2010). All the simulations are started at $z = 100$ and ended at $z = 5.5$, and 20 snapshots from $z = 17$ to 5.5 are saved. We simulate 2×320^3 gas and dark matter particles in a box with the side length of $10 \text{ cMpc } h^{-1}$. The initial mass is $3.39 \times 10^5 M_{\odot}$ for gas particles and $2.2 \times 10^6 M_{\odot}$ for dark matter particles. We apply the same chemistry reactions and corresponding molecular and metal cooling of Maio et al. (2007). Stars are formed when, as a result of cooling, the density of a gas particle reaches 70 cm^{-3} . When this happens, a kinetic wind with velocity 500 km s^{-1} is added to the gas particle to prevent overcooling. The stellar lifetimes we adopted are from Padovani & Matteucci (1993). Stars ending their life as supernovae (SNe) (Woosley & Weaver 1995; Heger & Woosley 2010)

yield and spread abundant metals in the hosting galaxies. As these metals mix quickly within the surrounding medium (de Avillez & Mac Low 2002), in our simulations they are smoothed over the neighbours according to the SPH kernel (Tornatore et al. 2007).

Due to our current lack of definite information about stellar properties in primordial galaxies, in this work we employ simulations for three possible choices of the PopIII IMF: very massive SNe (VMSNe), massive SNe (MSNe) and regular SNe (RSNe). More specifically, in the simulation referred to as VMSNe, the first stars are assumed to be in the mass range $[100, 500] M_{\odot}$ with a Salpeter slope, and contribute to metal pollution and thermal feedback via pair-instability SNe (PISNe) in the mass range $[140, 260] M_{\odot}$ (Heger & Woosley 2002). Both the MSN and RSN models sample a Salpeter IMF over $[0.1, 100] M_{\odot}$ for first stars, but they differ in the SN ranges contributing to metal pollution, which are $[10, 100] M_{\odot}$ (Heger & Woosley 2010) and $[10, 40] M_{\odot}$ (Woosley & Weaver 1995; Heger & Woosley 2002), respectively. PopII/I stars are formed once the local gas has reached a metallicity $Z > 10^{-4} Z_{\odot}$ (in solar units, $Z_{\odot} \approx 0.02$) (Bromm & Loeb 2003; Schneider et al. 2006). For PopII/I stars, we always adopt a Salpeter IMF in the range $[0.1, 100] M_{\odot}$ and include metal yields for AGB stars (van den Hoek & Groenewegen 1997), type Ia SNe (SNIa) (Thielemann et al. 2003) and Type II SNe (SNIa) (Woosley & Weaver 1995). We assume that the energy released from all SNe is 10^{51} erg , with the exception of PISNe, whose released energy is between $\sim 10^{51} - 10^{53} \text{ erg}$ depending on the mass of their progenitors (Heger & Woosley 2002).

Galaxies are identified by applying a friends-of-friends technique. In the following analysis, we only consider the galaxies with a minimum of 200 gas particles in order to prevent numerical artefacts (de Souza et al. 2013; Maio et al. 2013a; Salvaterra et al. 2013). We also exclude metal-free galaxies to focus on the metal enrichment originated by PopIII and PopII/I pollution.

Following the procedure of Ma et al. (2017) to describe the level of metal enrichment from first stars, we define the fraction of metals in a galaxy which is produced by PopIII stars as

$$f_{\text{III}} = \frac{\sum_j m_{Z_j, \text{III}}}{\sum_j m_{Z_j}}, \quad (1)$$

where m_{Z_j} is the mass of metal element Z_j in the galaxy (i.e. except hydrogen and helium) and $m_{Z_j, \text{III}}$ is the mass of metal element Z_j coming from PopIII stars. Then, we classify the selected galaxies into three classes based on the value of f_{III} :

1. ‘PopII-dominated’, if $f_{\text{III}} < 20$ per cent.
2. ‘Intermediate’, if $20 \text{ per cent} < f_{\text{III}} < 60$ per cent.
3. ‘PopIII-dominated’, if $f_{\text{III}} > 60$ per cent.

As a sanity check, we display in Fig. 1 the history of the comoving star formation rate (SFR) densities, $\dot{\rho}_*$, for both PopIII and PopII/I stars, as derived from the three simulations considered. We also overplot observational determinations inferred from data in different bands (Zheng et al. 2012; Coe et al. 2013; Ellis et al. 2013; McLure et al. 2013; Oesch et al. 2013, 2014; Bouwens et al. 2014, 2015; Duncan et al. 2014; Rowan-Robinson et al. 2016). While the simulations have very similar PopII/I SFR, the contribution from PopIII stars differs as a consequence of the different PopIII modelling (Maio et al. 2010). The PopII/I SFR densities are comparable to the one estimated from observations in the infrared (IR) at $z \approx 6$ (Rowan-Robinson et al. 2016), but higher than most of the data points derived from the UV magnitudes of galaxies at $5 < z < 11$ (Bouwens et al. 2014, 2015; Duncan et al. 2014). These latter values may give an incomplete picture, though, since UV determinations

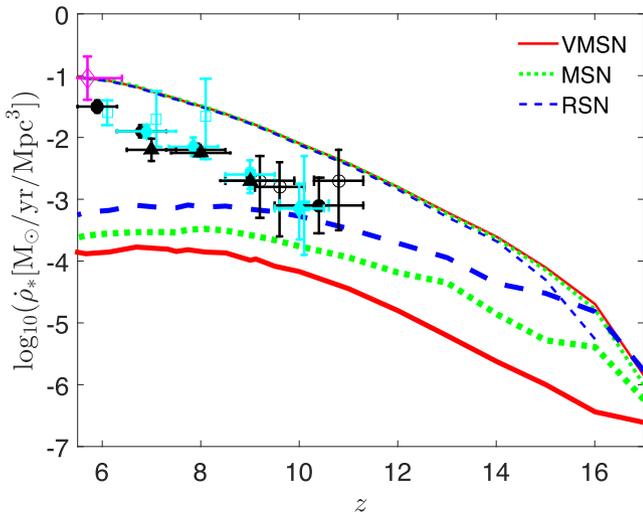


Figure 1. Redshift evolution of the intrinsic SFR comoving densities in models VMSN (solid red lines), MSN (dotted green) and RSN (dashed blue). The lower set of thick lines denote the SFR densities of PopIII stars, while the upper set of thin lines are the SFR densities of PopII/I stars. The data points with error bars show SFR densities derived by UV determinations from Bouwens et al. (2015) (filled black circles); CLASH data from Zheng et al. (2012), Coe et al. (2013) and Bouwens et al. (2014) (open black circles); McLure et al. (2013); Ellis et al. (2013) (filled cyan diamonds); Oesch et al. (2013, 2014) (filled cyan squares); Duncan et al. (2014) (empty cyan squares) and Laporte et al. (2016) (filled black triangles). IR determination (empty magenta diamond) is from Rowan-Robinson et al. (2016). (The colour version is only available in the online journal.)

could underestimate the contribution of embedded star formation (Rowan-Robinson et al. 2016) by up to 1 dex (Kennicutt & Evans 2012; Madau & Dickinson 2014). Despite the lower dust content of early galaxies in comparison to $z < 3$ objects expected by Capak et al. (2015), the critical issues on high- z IR luminosities (Bouwens et al. 2016; Laporte et al. 2016) and the origin of dust grains (Mancini et al. 2015; Bocchio et al. 2016; Ferrara, Viti & Ceccarelli 2016) still persist. Hence, dust extinction might be important already at $z \sim 7$ and could hide the existence of dusty, UV-faint galaxies at early epochs (Salvaterra et al. 2013; Mancini et al. 2016; Mancuso et al. 2016). We also note that the SFR at $z > 10$ might be slightly underestimated, as shown by simulations run at higher resolution (Maio et al. 2010; Campisi et al. 2011).

3 DAMPED LYMAN-ALPHA ABSORBERS

In the following, we first introduce the definition of a DLA in the simulations and then discuss the properties of PopIII/PopII-dominated DLAs as well as their evolution with redshift.

3.1 DLA definition

DLA (sub-DLA) in the observation is defined as a structure with column density of neutral hydrogen $N_{\text{HI}} > 2 \times 10^{20} \text{ cm}^{-2}$ ($10^{19} < N_{\text{HI}} < 2 \times 10^{20} \text{ cm}^{-2}$) along the line of sight of QSOs (Péroux et al. 2003; Wolfe et al. 2005). To correctly evaluate the ionization fraction of neutral hydrogen within the cosmic medium in numerical simulations, the radiative transfer of ionizing photons should be accounted for. Given the ability of our implementation to capture gas physics and chemistry down to very low temperatures, here we prefer to select DLAs in the simulations following the procedure of Maio et al. (2013a) and Maio & Tescari (2015), i.e.

gaseous objects with mass-weighted gas temperature $T \leq 10^4 \text{ K}$ (consistent with Cooke, Pettini & Jorgenson 2015). Here T is calculated as $\sum_k m_k T_k / \sum_k m_k$, where k denotes all the gas particles in one halo, and m_k and T_k are the mass and temperature of the k th gas particle, respectively.

Alternatively, we also consider the description of DLAs in terms of a mass-dependent neutral hydrogen cross-section (as done in Kulkarni et al. 2013, 2014), which is parametrized through a fit calibrated against $z = 3$ observations (see also Pontzen et al. 2008; Font-Ribera et al. 2012):

$$\Sigma(M_{\text{halo}}) = \Sigma_0 \left(\frac{M_{\text{halo}}}{M_0} \right)^2 \left(1 + \frac{M_{\text{halo}}}{M_0} \right)^{\alpha-2}, \quad (2)$$

where $\alpha = 0.2$, $M_0 = 10^{9.5} M_{\odot}$ and $\Sigma_0 = 40 \text{ kpc}^2$ at $z = 3$. This fitting function matches very well the column density distribution function observed at low z (Kulkarni et al. 2013). However, we note that a possible bias may appear for high- z DLAs due to the fact that the fit is calibrated to reproduce data at $z = 3$.

To aid the assessment of the impact of such definitions on our results, we discuss below some general galactic properties. In Fig. 2, we present the distributions of PopII-dominated (red), intermediate (green) and PopIII-dominated (blue) galaxies at three redshifts in terms of mass-weighted temperature T versus dark matter halo mass M_{halo} . As expected, more massive galaxies appear with decreasing redshift. For example, at $z = 5.5$ halo masses are between $10^{8.5}$ and $10^{11.5} M_{\odot}$, with corresponding gas temperatures ranging from 10^3 to several 10^5 K , while at $z = 9.5$ the halo masses are $< 10^{10} M_{\odot}$. The more massive galaxies usually have higher gas temperatures, as the higher-mass objects are statistically more evolved and have deeper potential well, higher SFR and stellar mass.

No matter at which redshift PopII-dominated galaxies show very similar T versus M_{halo} distributions in the three simulations, i.e. their mass and temperature are weakly affected by the first star model adopted.

Also the distributions of PopIII-dominated galaxies are consistent in the three simulations, with halo masses $M_{\text{halo}} < 10^{9.6} M_{\odot}$ and temperatures $T < 10^{4.5} \text{ K}$, lower than those of the PopII-dominated galaxies. No significant redshift evolution is evident here since PopIII metal pollution usually dominates in low-mass galaxies, where gas parcels feature mostly pristine composition, irrespective of redshift.

The halo mass and gas temperature of intermediate galaxies are very close to those of PopIII-dominated galaxies, with a few rare exceptions. The paucity of intermediate-class galaxies is due to the rapidity of metal enrichment, which is quickly dominated by PopII/I star formation (i.e. $f_{\text{III}} < 20$ per cent). Due to the different stellar lifetimes, metal spreading from first PopIII SNe is delayed in the RSN model with respect to the VMSN and MSN models. Therefore, the PopIII SFR reaches higher levels and have more prolonged duration (Ma et al. 2017), delaying the transition from intermediate to PopII-dominated galaxies (i.e. more galaxies retain $f_{\text{III}} > 20$ per cent). For this reason, intermediate galaxies are more abundantly present in the RSN model than in the VMSN and MSN models.

From Fig. 2, it is clear that the two theoretical definitions of DLAs are expected to affect the results on the predicted properties. Indeed, since PopIII-dominated objects have very low gas temperatures, most of them could be DLA candidates when selected with the temperature threshold $T \leq 10^4 \text{ K}$ (in the following referred to as ‘temperature selected’), while PopII-dominated ones, populating the high-mass end of the distributions, would dominate the candidate DLA population when selected according to equation 2 (‘mass

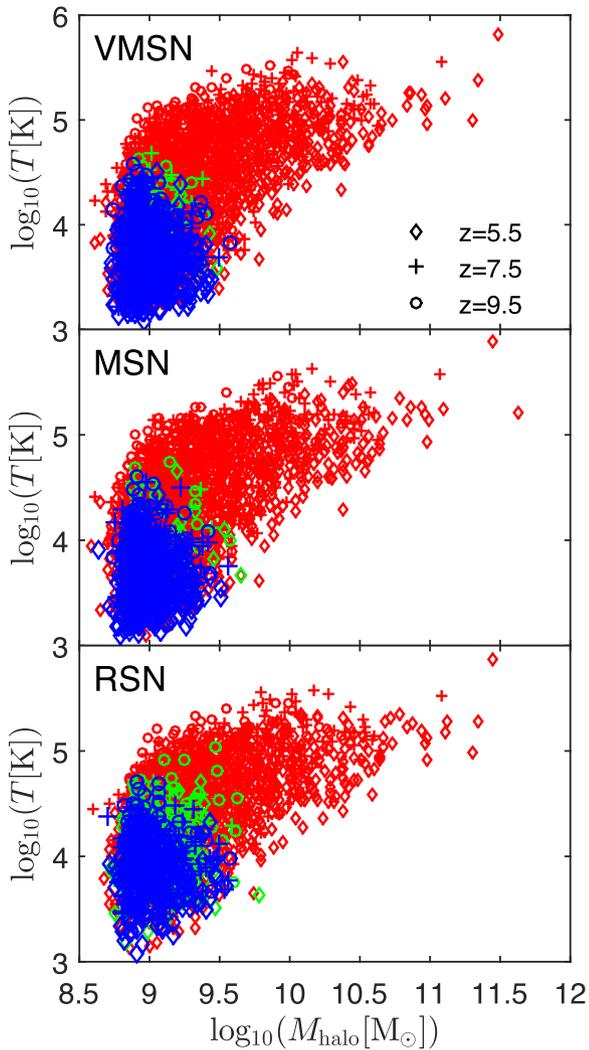


Figure 2. Distributions of mass-weighted temperature (T) of PopII-dominated (small red symbols), intermediate (green) and PopIII-dominated (big blue) galaxies as a function of halo mass (M_{halo}). The results are shown at $z = 9.5$ (circles), 7.5 (crosses) and 5.5 (diamonds). From top to bottom, the panels refer to the VMSN, MSN and RSN models. (The colour version is only available in the online journal.)

selected^{*)}. The differences of these two approaches are also discussed in Morrison et al. (2016). In the following, the evolution and the statistical properties of DLAs selected with both criteria will be shown.

3.2 Evolution of PopII/PopIII-dominated DLAs

In Fig. 3, we show the evolution of the fraction of PopIII-dominated (top panel) and PopII-dominated (bottom panel) DLAs as a function of redshift, according to both DLA definitions. For temperature-selected DLAs (thin lines), the fraction is defined as $N_i(z)/N_{\text{tot}}(z)$, where $N_i(z)$ is the number of class i DLAs and $N_{\text{tot}}(z)$ is the total number of all classes of DLAs at redshift z . For mass-selected DLAs (thick lines) the fraction of class i is computed as $\Delta_i(z)/\Delta_{\text{tot}}(z)$, where $\Delta_i(z)$ is sum of $\Sigma(M_{\text{halo}})$ of class i galaxies and $\Delta_{\text{tot}}(z)$ is the integral of $\Sigma(M_{\text{halo}})$ of all galaxies at redshift z .

Consistently with what mentioned in the previous section, the fraction of PopIII-dominated DLAs decreases with redshift because of the increasing metal contribution from PopII/I stars. As the

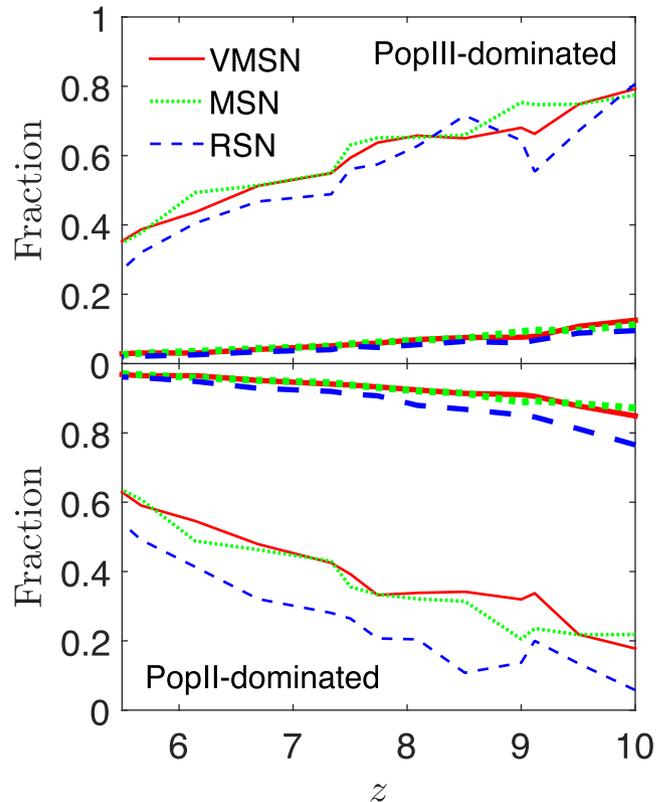


Figure 3. Redshift evolution of the fraction of DLAs which are PopIII-dominated (top) and PopII-dominated (bottom) in models VMSN (solid red lines), MSN (dotted green) and RSN (dashed blue). The thin and thick sets of curves refer to temperature-selected DLAs and mass-selected DLAs, respectively. (The colour version is only available in the online journal.)

cosmic metallicity is hardly affected by the mass range of the first stars (Maio et al. 2010), this fraction is almost independent of the first star models, although the RSN model shows slightly lower PopIII-dominated values than the others, as a result of the delayed metal pollution from smaller-mass first stars that have longer lifetimes. On the other hand, while temperature-selected DLAs have a very high probability to be PopIII-dominated, e.g. 80 per cent at $z = 10$ and 40 per cent at $z = 6$, the corresponding fraction for mass-selected DLAs is only 10 per cent at $z = 10$ and 3 per cent at $z = 6$. The difference is caused by the fact that PopIII-dominated galaxies are likely to have $T \leq 10^4$ K and small halo masses (see Fig. 2). These characteristics are strongly disfavoured when using the fit of equation 2 (which gives a larger weight to typically hotter objects with masses $\gtrsim 10^{9.5} M_{\odot}$).

The fraction of PopII-dominated DLAs is complementary to that of PopIII-dominated ones. It turns to be very high for mass-selected DLAs, with fractions exceeding 80 per cent at $z = 10$ and increasing with decreasing redshift. As just mentioned, this trend is due to the adopted fit in equation 2, which favours more massive PopII-dominated galaxies. On the contrary, temperature-selected PopII-dominated DLAs feature fractions of just 20 per cent at $z = 10$ and 60 per cent at $z = 5.5$, as more massive galaxies are excluded by the temperature selection criterion.

From these results, we also conclude that the fraction of intermediate DLAs is rather low for both definitions. Precisely, it is < 4 per cent in the VMSN and MSN models and only slightly larger in the RSN model (see also Figs 2 and 5).

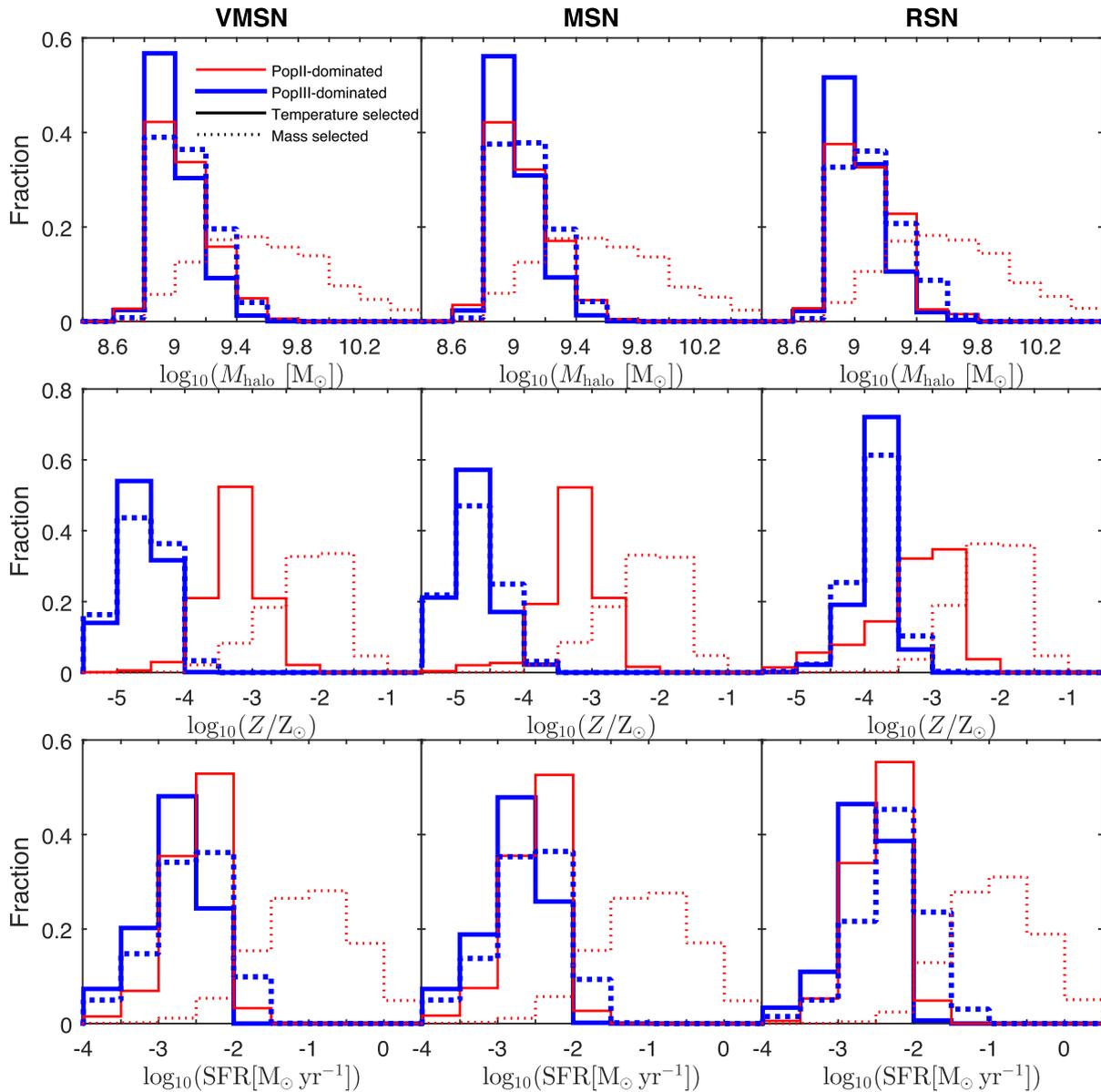


Figure 4. From top to bottom, halo mass (M_{\odot}), metallicity (Z_{\odot}) and SFR ($M_{\odot} \text{ yr}^{-1}$) distributions of PopII-dominated (red thin lines) and PopIII-dominated (blue thick lines) DLAs at $z \geq 5.5$, in base-10 logarithmic scale. Solid lines indicate temperature-selected DLAs, while dotted lines indicate mass-selected DLAs. From left to right, the columns refer to models VMSN, MSN and RSN. (The colour version is only available in the online journal.)

While the fraction of PopIII-dominated DLAs becomes larger at higher redshifts, the total DLA number density is lower at earlier epochs, because less structures are present. Thus, to predict how many high- z DLAs are observable as PopIII-dominated, we calculate the expected cumulative fraction of all DLAs at $z \geq 5.5$ that are PopIII-dominated. For temperature-selected DLAs, it is computed by

$$f_{\text{DLA},T} = \frac{\int_{z>5.5} dz' N_i(z') \Phi(z') \frac{dV}{dz'}}{\int_{z>5.5} dz' N_{\text{tot}}(z') \Phi(z') \frac{dV}{dz'}}, \quad (3)$$

where $\Phi(z') \propto 10^{-0.5 * z'}$ denotes the quasar number at redshift larger than z' (Fan et al. 2001), $dV/dz' = D^2 \frac{dD}{dz'} d\Omega$ indicates the comoving cosmic volume and $D(z')$ is the comoving distance between redshift z' and 0. The cumulative fraction is weighted by the number of background quasars, because the number of observable DLAs also depends on the number of background sources (Fan et al. 2001). As a

result, we get values of 42.5 per cent, 43.6 per cent and 36.5 per cent for the VMSN, MSN and RSN models, respectively. For mass-selected DLAs, this cumulative fraction is expressed as

$$f_{\text{DLA},M} = \frac{\int_{z>5.5} dz' \Delta_i(z') \Phi(z') (1+z')^2 \frac{dD}{dz'}}{\int_{z>5.5} dz' \Delta_{\text{tot}}(z') \Phi(z') (1+z')^2 \frac{dD}{dz'}}, \quad (4)$$

which gives 3.4 per cent, 3.5 per cent and 2.6 per cent for VMSN, MSN and RSN models, respectively.

3.3 DLA properties

In Fig. 4 we show the distributions of halo mass, metallicity and SFR of PopIII-dominated (blue thick lines) and PopII-dominated (red thin lines) DLAs at $z \geq 5.5$. For temperature-selected DLAs

(solid lines) the fraction of class i DLAs in the j th bin is computed by

$$F_{i,j;T} = \frac{\int_{z>5.5} dz' N_{i,j}(z') \Phi(z') \frac{dV}{dz'}}{\int_{z>5.5} dz' N_{i,\text{tot}}(z') \Phi(z') \frac{dV}{dz'}}, \quad (5)$$

where $N_{i,j}(z')$ is the number of class i DLAs in the j th bin at redshift z' , and $N_{i,\text{tot}}(z')$ is the total number of class i DLAs at z' . For mass-selected DLAs (dotted lines), it is calculated by

$$F_{i,j;M} = \frac{\int_{z>5.5} dz' \Delta_{i,j}(z') \Phi(z') (1+z')^2 \frac{dD}{dz'}}{\int_{z>5.5} dz' \Delta_{i,\text{tot}}(z') \Phi(z') (1+z')^2 \frac{dD}{dz'}}, \quad (6)$$

where $\Delta_{i,j}(z')$ is the sum of $\Sigma(M_{\text{halo}})$ of class i DLAs in the j th bin at z' , and $\Delta_{i,\text{tot}}(z')$ is the sum of $\Sigma(M_{\text{halo}})$ of all class i DLAs at z' .

Although no obvious dependence on the first star models, the properties of PopII-dominated DLAs (red lines) feature major differences according to the selection criteria. Specifically, temperature-selected PopII-dominated DLAs (solid red lines) have very low halo masses $M_{\text{halo}} < 10^{9.6} M_{\odot}$, while metallicities are concentrated in the range $(10^{-4.5} \text{ to } 10^{-2}) Z_{\odot}$ and peak around $Z \approx 10^{-3} Z_{\odot}$. Their SFRs are $< 10^{-1.5} M_{\odot} \text{ yr}^{-1}$. Since equation 2 suppresses low-mass structures and imposes a roughly flat trend of neutral hydrogen cross-section for big structures, the mass-selected PopII-dominated DLAs (dotted red lines) have flatter halo mass distributions with $M_{\text{halo}} > 10^{8.8} M_{\odot}$ and a peak at $\approx 10^{9.6} M_{\odot}$. Their metallicities are in the range $(10^{-3.5} \text{ to } 10^{-1}) Z_{\odot}$ and peak at $\approx 10^{-2} Z_{\odot}$, which is 1 dex higher than those of temperature-selected ones. They also show larger SFRs ($> 10^{-2.5} M_{\odot} \text{ yr}^{-1}$), that peak at $\approx 0.1 M_{\odot} \text{ yr}^{-1}$.

As the two DLA definitions select similar samples of PopIII-dominated galaxies (see Fig. 2), the distributions of PopIII-dominated DLAs (blue lines) show similar trends, although mass-selected ones have slightly higher mass, metallicity and SFR. They are typically hosted in haloes with mass $M_{\text{halo}} < 10^{9.6} M_{\odot}$ (also see Fig. 2), consistent with Kulkarni et al. (2013) and Maio & Tescari (2015), and peak at $\approx 10^9 M_{\odot}$. Their typical metallicities are, as expected, lower than those of PopII-dominated DLAs, i.e. $Z < 10^{-4} Z_{\odot}$ in the VMSN and MSN models, while $10^{-5} Z_{\odot} < Z < 10^{-3} Z_{\odot}$ in the RSN model, since the higher small-mass PopIII SFR in the RSN model induces more metal pollution in low-mass galaxies. Their SFRs are $< 10^{-1.5} M_{\odot} \text{ yr}^{-1}$ and peaks around $\approx 10^{-2.5} M_{\odot} \text{ yr}^{-1}$.

Temperature-selected PopII-dominated (solid red lines) DLAs present halo mass and SFR distributions very similar to those of the PopIII-dominated ones (solid blue lines), while they still can be neatly identified through metallicities. In fact, most PopIII-dominated DLAs have $Z < 10^{-3.5} Z_{\odot}$, while only a few temperature-selected PopII-dominated DLAs reach such low values.

4 DLA METAL ABUNDANCES

In this section, we present the results for metal abundance ratios of indicative heavy elements extracted from our simulations. We focus on carbon (C), oxygen (O), silicon (Si) and iron (Fe), which are abundant in galaxies and have been detected in high- z DLAs (Becker et al. 2012; Morrison et al. 2016).

The abundance ratios with respect to the corresponding solar values are defined as $[A/B] = \log_{10}(N_A/N_B) - \log_{10}(N_A/N_B)_{\odot}$, where A and B are two arbitrary species, $N_{A(B)}$ is the number densities of element A(B) and the subscript \odot denotes the solar values from Asplund et al. (2009).

4.1 Abundance ratios

In Fig. 5 we show distributions of abundance ratios at redshift $z = 9.5, 7.5$ and 5.5 , respectively, for 617, 1496 and 2581 metal-enriched galaxies in the VMSN model; 608, 1508 and 2590 in the MSN model and 512, 1332 and 2370 in the RSN model. The RSN model features fewer (by ~ 16 per cent at $z = 9.5$, ~ 11 per cent at $z = 7.5$ and ~ 8 per cent at $z = 5.5$) metal polluted galaxies than the other models due to the longer lifetimes of small-mass first stars that delay the initial phases of cosmic metal enrichment. All these galaxies are possible DLA candidates when selected according to the mass criterion, while only 10 per cent to 30 per cent of them can be DLAs when selected according to their gas temperature. We have verified that this smaller sample of DLAs retains the same characteristics of the metal ratio distributions of the full sample. This is because the metal abundance ratios of galaxies are mainly determined by the yields of the polluting stellar population and thus are weakly dependent on the halo mass or gas temperature. For this reason, the distributions are representative of the DLA population, independent of the selection. In the remaining of the paper we will thus refer to abundance ratios of DLAs rather than of galaxies.

From Fig. 5, PopII-dominated, intermediate and PopIII-dominated DLAs are clearly separated in the panels [Si/O] versus [C/O], [Si/Fe] versus [O/Fe] and [Si/C] versus [C/Fe] for the VMSN and MSN models, while in the RSN model they mostly overlap.

More specifically, PopII-dominated DLAs (red symbols) display abundance ratios very similar in the three models, indicating their common origin (i.e. pollution by the same PopII generation). A linear correlation between [Si/O] and [C/O] is evident, due to the different metal pollution phase in each DLA. The [Si/O] ratios of PopII-dominated DLAs are mostly between -0.2 and 0.2 , while typically [C/O] ratios are in the range $[-0.6, -0.3]$. A few cases at $z = 5.5$ have [C/O] > -0.3 due to pollution by local AGB stars, whose contribution becomes relevant at such epochs as a consequence of AGB lifetimes.² PopII-dominated DLAs also show a roughly linear relation between [Si/Fe] and [O/Fe] and they are preferentially located in regions where [O/Fe] = [0.3, 0.9] and [Si/Fe] = [0.2, 0.7]. Some have very low [O/Fe] and [Si/Fe] ratios at $z = 5.5$, resulting from metal spreading from stars dying as SNeIa (with high iron yields) that become increasingly important with decreasing redshift. In the [Si/C] versus [C/Fe] panels, PopII-dominated DLAs in the three models share the region delimited by [C/Fe] = $[-0.1, 0.3]$ and [Si/C] = [0.2, 0.5]. Some differences appear, though, e.g. DLAs with [Si/C] > 0.5 are present only in the VMSN model as a result of the high [Si/C] yields (≈ 1.14) from PISNe. Then, only the MSN model features some PopII-dominated DLAs with [C/Fe] > 0.3 because of the low iron yields expected for faint supernovae from massive stars (Heger & Woosley 2010).

The distributions of PopIII-dominated DLAs (blue symbols) are visibly different in the three models, and do not show obvious time evolution, because of the leading role of the metal yields for massive stars at these epochs. We note that a PopIII IMF (Hirano et al. 2015) changing with redshift and having strongly variable yields for massive stars might lead to an evolution of the distributions for PopIII-dominated DLAs.

Thanks to the large explosion energy of PISNe and the short lifetime of very massive PopIII stars, PopIII-dominated DLAs in the VMSN model are mostly concentrated in a small region of the

² AGB stars evolve in a few 10^8 yr, hence they can enrich the medium already at $z \sim 6$ by spreading metals via mass loss.

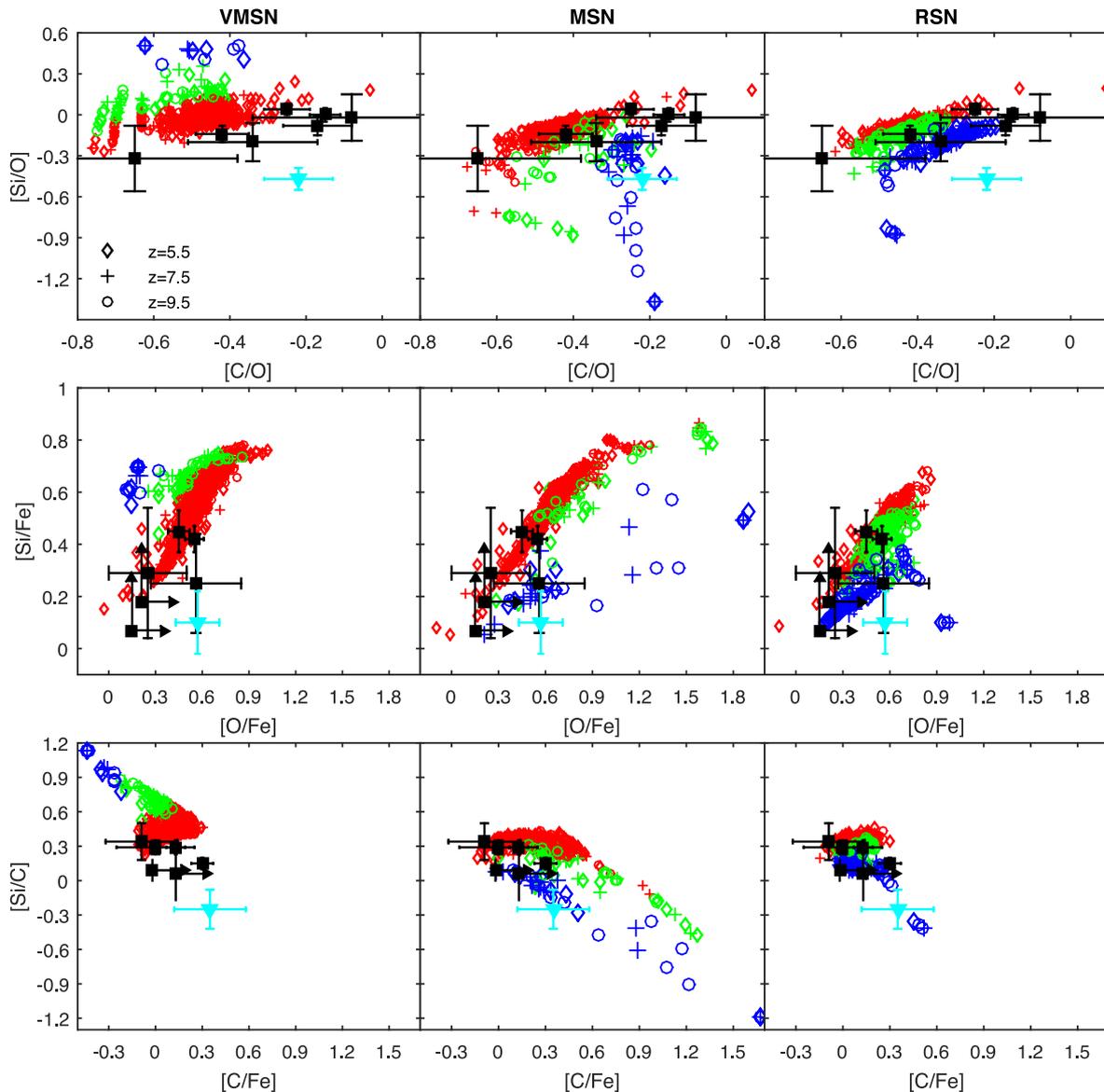


Figure 5. Distributions of abundance ratios $[\text{Si}/\text{O}]$ versus $[\text{C}/\text{O}]$ (top panels), $[\text{Si}/\text{Fe}]$ versus $[\text{O}/\text{Fe}]$ (central) and $[\text{Si}/\text{C}]$ versus $[\text{C}/\text{Fe}]$ (bottom). From left to right the columns refer to models: VMSN, MSN and RSN. The symbols denote PopII-dominated (small red), intermediate (green) and PopIII-dominated (big blue) galaxies at $z = 9.5$ (circles), 7.5 (crosses) and 5.5 (diamonds). Black solid squares with error bars (or arrows to denote lower or upper limits) are DLAs observed in the redshift range $5.3 - 6.2$ (Becker et al. 2012), while the cyan solid triangle refers to a sub-DLA at $z = 5$ reported by Morrison et al. (2016). (The colour version is only available in the online journal.)

$[\text{Si}/\text{O}]$ versus $[\text{C}/\text{O}]$ panel with $[\text{C}/\text{O}] \approx -0.62$ and $[\text{Si}/\text{O}] \approx 0.51$. Similarly, these DLAs typically have $[\text{O}/\text{Fe}] \approx 0.18$ versus $[\text{Si}/\text{Fe}] \approx 0.7$ and $[\text{C}/\text{Fe}] \approx -0.44$ versus $[\text{Si}/\text{C}] \approx 1.14$. Although their $[\text{C}/\text{O}]$ and $[\text{O}/\text{Fe}]$ might be similar to those of PopII-dominated and intermediate DLAs, they would display higher $[\text{Si}/\text{O}]$ and $[\text{Si}/\text{Fe}]$, e.g. $[\text{Si}/\text{O}] > 0.3$.

Since the MSN model includes PopIII metal enrichment both from massive stars ($>40 M_{\odot}$), that have very high carbon and oxygen but low silicon and iron yields when exploding as faint supernovae, and from regular stars ($<40 M_{\odot}$), whose metal yields are very similar to those of PopII stars, a number of PopIII-dominated DLAs have $[\text{Si}/\text{O}] \approx -1.37$ and $[\text{C}/\text{O}] \approx -0.19$, and some are closer to PopII-dominated DLAs, while others are located in between, as a result of the combined contribution of massive and low-mass PopIII stars. The same effects also appear in the distribution of $[\text{Si}/\text{Fe}]$ ver-

sus $[\text{O}/\text{Fe}]$, i.e. some of them have $[\text{O}/\text{Fe}] \approx 1.9$ and $[\text{Si}/\text{Fe}] \approx 0.5$, while others are located around $[\text{O}/\text{Fe}] \approx 0.5$ and $[\text{Si}/\text{Fe}] \approx 0.2$. In the $[\text{Si}/\text{C}]$ versus $[\text{C}/\text{Fe}]$ panel, they are distributed in the range $[\text{C}/\text{Fe}] = [0, 1.68]$ and $[\text{Si}/\text{C}] = [-1.2, 0]$.

The metal yields of pristine SNe with small-mass progenitors are similar to those from PopII/I stars, thus PopIII-dominated DLAs in the RSN model are located near PopII-dominated and intermediate DLAs, although some differences are still visible. For example, given the same $[\text{C}/\text{O}]$ (or $[\text{O}/\text{Fe}]$), they present $[\text{Si}/\text{O}]$ (or $[\text{Si}/\text{Fe}]) \approx 0.2$ dex lower than that of PopII-dominated DLAs, and they also feature a smaller $[\text{Si}/\text{C}]$ (< 0.2).

Intermediate DLAs (green symbols) always locate between PopII-dominated and PopIII-dominated ones. They are relatively few in the VMSN and MSN models when compared to the RSN model, as mentioned earlier. Since massive PopII stars (with mass

Table 1. Abundance ratios of DLAs observed at $z \geq 5$. The first measurement is from Morrison et al. (2016), while the others are from Becker et al. (2012).

QSO	z	[Si/O]	[C/O]	[Si/C]	[C/Fe]	[O/Fe]	[Si/Fe]
SDSS J1202+3235	5.0	-0.47 ± 0.08	-0.22 ± 0.09	-0.25 ± 0.17	0.35 ± 0.23	0.57 ± 0.14	0.10 ± 0.12
SDSS J0231-0728	5.3380	-0.14 ± 0.06	-0.42 ± 0.07	0.29 ± 0.06	0.13 ± 0.06	0.55 ± 0.06	0.42 ± 0.05
SDSS J0818+1722	5.7911	0.00 ± 0.05	-0.15 ± 0.04	0.15 ± 0.05	0.30 ± 0.07	0.45 ± 0.07	0.45 ± 0.08
SDSS J0818+1722	5.8765	-0.08 ± 0.07	-0.17 ± 0.09	0.09 ± 0.09	> -0.02	> 0.15	> 0.07
SDSS J1148+5251	6.0115	0.04 ± 0.04	-0.25 ± 0.06	0.29 ± 0.06	-0.00 ± 0.25	0.25 ± 0.25	0.29 ± 0.25
SDSS J1148+5251	6.1312	-0.32 ± 0.24	-0.65 ± 0.27	0.34 ± 0.16	-0.09 ± 0.23	0.56 ± 0.29	0.25 ± 0.19
SDSS J1148+5251	6.1988	-0.20 ± 0.14	-0.34 ± 0.17	0.14 ± 0.12	–	–	–
SDSS J1148+5251	6.2575	-0.02 ± 0.17	-0.08 ± 0.26	0.06 ± 0.23	> 0.13	> 0.21	> 0.18

$>30 M_{\odot}$) have stellar lifetimes much shorter than low-mass ones, intermediate DLAs retain some hints of their metal pollution by having [C/O] ratios lower, but [Si/Fe] ratios higher than the PopIII-dominated DLAs, which are the typical features of PopII stars with mass $>30 M_{\odot}$.

The unique characteristics shown by PopIII-dominated DLAs in the abundance ratios can be exploited to identify first star signatures and constrain PopIII mass ranges by comparing them to observed DLAs at $z \geq 5$ with high precision abundance ratios measured (Becker et al. 2012; Morrison et al. 2016). These are listed in Table 1 and shown as symbols in Fig. 5.

Most of data points are consistent with PopII-dominated DLAs, although some of them have error bars too large to draw definitive conclusions. Meanwhile, some data points seem to favour the MSN and RSN models, although no observation can clearly discriminate one model over the others.

The DLA at $z = 5.8765$ found along the line-of-sight of QSO SDSS J0818+1722 is very close to the PopIII-dominated DLAs in the RSN model, as indicated by the [Si/O] versus [C/O] panel, as well as by the ratio [Si/C] = 0.09 ± 0.09 . Nevertheless, there are only lower limits for its [Si/Fe] and [O/Fe], which cannot then be used to set more stringent constraints on its class and progenitors.

The DLA at $z = 5.7911$ is also consistent with PopIII-dominated objects in the RSN model in the [Si/O] versus [C/O] panel, while it clearly locates in the region of PopII-dominated DLAs in the [Si/Fe] versus [O/Fe] panel.

All the abundance ratios of the recently observed sub-DLA at $z = 5$ from QSO SDSS J1202+3235 (Morrison et al. 2016) are consistent with PopIII-dominated DLAs in the MSN model, although it does not display the features of faint supernovae, e.g. very low [Si/O] (< -1) or very high [C/Fe] (> 1). Its abundance ratios are also close to those of PopIII-dominated DLAs in the RSN model in the panel [Si/C] versus [C/Fe], but they are not fully consistent with the same model in the panel [Si/O] versus [C/O]. We also note that the metallicity of this sub-DLA ($[O/H] = -2.00 \pm 0.12$) is much higher than the PopIII-dominated DLAs predicted in our simulations.

4.2 Evolution of DLA mean abundance ratios

In Fig. 6, we show the evolution of the average [C/O], [Si/C] and [C/Fe] for DLAs selected under both definitions. As mass-selected DLAs are dominated by massive haloes with metal abundances predominantly coming from PopII stars, their mean abundance ratios (thick lines) are weakly dependent on the first star models and present a mild evolution with redshift (Kulkarni et al. 2013). Instead, the mean abundance ratios of temperature-selected DLAs (thin lines) are dramatically different in the three models and show an evident redshift evolution.

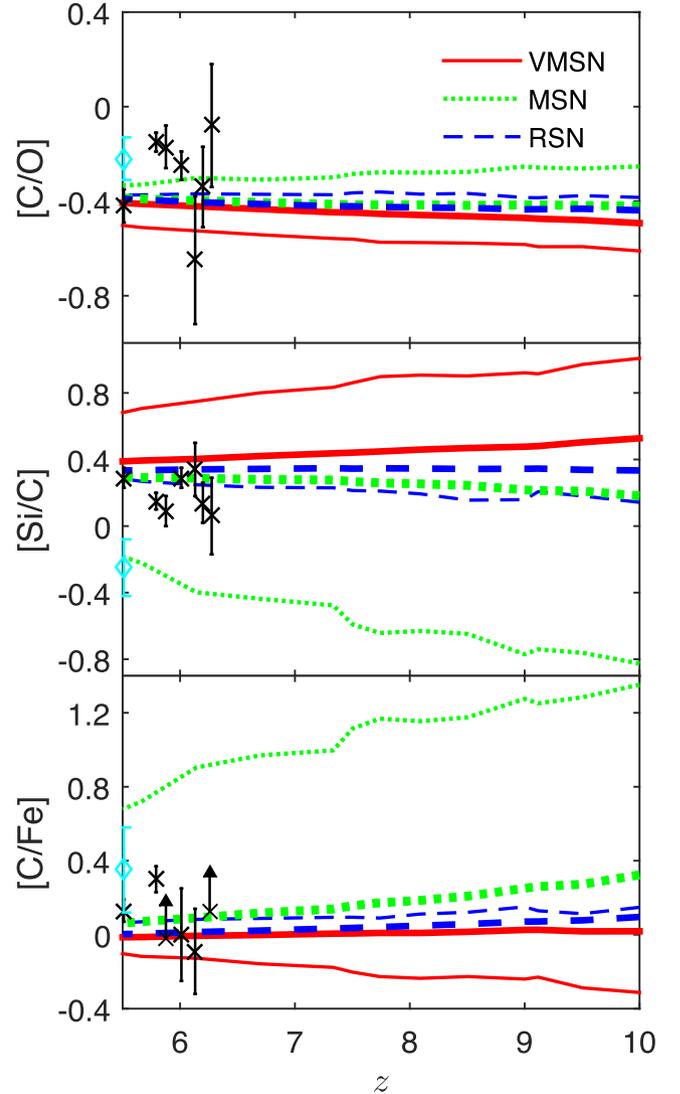


Figure 6. Mean value of [C/O] (top), [Si/C] (middle) and [C/Fe] (bottom) in the VMSN (solid red lines), MSN (dotted green) and RSN (dashed blue) models as a function of redshift. The thin and thick curve sets denote temperature-selected and mass-selected DLAs, respectively. The black crosses with error bars (or arrows) are data for observed DLAs taken from Becker et al. (2012), and the one showed as cyan diamond with error bar is from Morrison et al. (2016). (The colour version is only available in the online journal.)

Specifically, since PopIII stars in all three models have [C/O] yields similar to those of PopII stars, the mean [C/O] ratios of DLAs present differences of only 0.3 dex for the temperature-selected ones and 0.1 dex for the mass-selected ones. In most cases, the mean [C/O] have a little increase with decreasing redshift, reaching ≈ -0.4 at $z = 5.5$, as a result of the increasing contribution from AGB stars. In the VMSN model, this behaviour is also partly caused by the low [C/O] yields of PISNe (< -0.6). PISN pollution is mostly significant at high redshift, so at $z = 10$ the mean [C/O] in the VMSN model is ≈ -0.6 (≈ -0.5) for temperature- (mass-)selected DLAs. However, temperature-selected DLAs in the MSN model show an exception, i.e. their average [C/O] decreases with decreasing redshift. Indeed, massive first stars, that explode as faint supernovae with high [C/O] yields (≈ -0.2), determine a relatively high mean [C/O] ratio at $z = 10$. As redshift decreases, the mean [C/O] is gradually reduced by the growing contribution of PopII stars that have lower [C/O] yields (≈ -0.4).

PopIII-dominated DLAs have very high (low) [Si/C] in the VMSN (MSN) (see also bottom panels in Fig. 5), thus the average [Si/C] values feature the largest differences within the three models, e.g. ≈ 1.8 dex (≈ 0.4 dex) for temperature-selected (mass-selected) DLAs at $z = 10$. As redshift decreases, the contribution from PopIII stars becomes less important, hence the mean [Si/C] ratios in both VMSN and MSN models tend to converge to that of the RSN model, with [Si/C] ≈ 0.3 .

The mean [C/Fe] ratio in the RSN model is ≈ 0 , and slowly decreases with decreasing redshift. This might be caused by SNIa events that start contributing to metal enrichment at low redshift and injecting important amounts of iron. Although mass-selected DLAs in the VMSN model have mean [C/Fe] very similar to that of the RSN model, as a result of their small fraction of PopIII-dominated DLAs (see Fig. 3), the temperature-selected ones in this model show a low mean [C/Fe] at $z = 10$ (≈ -0.3) because of the low [C/Fe] yields of PISNe, while it increases with decreasing redshift and finally converges to the RSN trend. With larger [C/Fe] yields from faint SNe, the average [C/Fe] in the MSN model is the highest, e.g. [C/Fe] ≈ 1.3 (≈ 0.3) for temperature-selected (mass-selected) DLAs at $z = 10$.

[C/O] measurements would not help discriminate the various PopIII models in a solid way, while it is possible to infer interesting conclusions from observational data of [Si/C] and [C/Fe]. The predicted mean metal ratios of mass-selected DLAs (thick lines), for all three models, are consistent with the observed high- z DLAs (see also Kulkarni et al. 2013), while they have clear difficulties in reproducing the observation by Morrison et al. (2016) for [Si/C] and [C/Fe]. Temperature-selected DLAs (thin lines) feature trends similar to the ones obtained by mass selection in the RSN model, while they show higher (lower) mean [Si/C] in the VMSN (MSN) model, as well as much higher [C/Fe] in the MSN model. These latter trends are in tension with [Si/C] and [C/Fe] data by Becker et al. (2012), but explain reasonably well the data of Morrison et al. (2016) through MSN enrichment. We note that the various differences may be caused by a bias in the observations, e.g. PopIII-dominated DLAs with such low metallicities ($Z < 10^{-3.5} Z_{\odot}$) might not be easy to detect (Wolfe et al. 2005).

5 DISCUSSION AND CONCLUSIONS

We use hydrodynamical chemistry simulations with different PopIII IMFs to study the metal signatures of the first stars in high- z ($z \geq 5.5$) DLAs. The simulations include atomic and molecular chemistry, stellar evolution and metal enrichment according to metal-

dependent yields and lifetimes. In all simulations, PopII/I stars are assumed to be distributed according to a Salpeter IMF with mass range $[0.1, 100] M_{\odot}$ and include metal pollution from type II SNe, AGBs and type Ia SNe. Three cases for PopIII stars have been considered: (i) a top-heavy IMF with mass range $[100, 500] M_{\odot}$ and PISN explosions of stars with mass $[140, 260] M_{\odot}$ (VMSN); (ii) a normal Salpeter IMF with mass range $[0.1, 100] M_{\odot}$, which includes massive SNII explosions within $[10, 100] M_{\odot}$ (MSN) and (iii) a normal Salpeter IMF with SNII explosions within $[10, 40] M_{\odot}$ (RSN). Thanks to our detailed implementation, we can retrieve solid results for the chemistry and the pollution history of the primordial Universe, which are mainly led by the first PopIII and PopII/I star formation episodes and by the interplay between different feedback mechanisms.

Although our simulations cover a wide range of PopIII IMFs, other options could be possible for PopIII stars, while even the most extreme scenarios are unlikely to have first stars above $500 M_{\odot}$. Similarly, stellar structure models and physical processes in the stellar cores different from those adopted here, e.g. explosion mechanisms, magnetic fields, rotation and nuclear reaction rates, might affect the theoretical metal yields from supernova explosions and thus the predicted metal abundance ratios for DLAs (see more discussion in Ma et al. 2017), nevertheless the cosmic gas evolution would not be changed significantly. The PopIII SFR, and thus the corresponding metal contribution, may be slightly changed by the adopted criteria for the transition from a PopIII to a PopII/I star formation mode, but the PopII/I SFR is hardly affected (Maio et al. 2010). While significant changes are not expected when varying the adopted underlying cosmological model, effects from warm dark matter (Maio & Viel 2015) can suppress the early SFR, and primordial streaming motions (Tsaliakhovich & Hirata 2010; Maio, Koopmans & Ciardi 2011a) can moderately impact gas evolution and primordial gas collapse at very high z .

In our analysis, we explore the patterns of DLAs extracted from the numerical simulations with two possible selection criteria: the first one selects DLAs as gaseous objects with mass-weighted temperature $T < 10^4$ K (temperature-selected DLAs; Cooke et al. 2015), and the other one is based on the associated neutral hydrogen cross-section (mass-selected DLAs; as in Kulkarni et al. 2013). Our main results can be summarized as follows.

(i) The fraction of DLAs that are PopIII-dominated decreases quickly with redshift, almost independent of the first star models. However, it is dramatically affected by the selection criteria for DLAs, as it can be as high as 40 per cent at $z = 6$ for temperature-selected DLAs, but only 3 per cent for mass-selected ones.

(ii) PopIII-dominated DLAs have very low halo mass ($< 10^{9.6} M_{\odot}$), metallicities ($< 10^{-4} Z_{\odot}$ in the VMSN and MSN models and $< 10^{-3} Z_{\odot}$ in the RSN model) and SFR ($< 0.01 M_{\odot} \text{ yr}^{-1}$), without much dependence on the selection criteria.

(iii) PopII-dominated DLAs, instead, display different properties under the two adopted definitions. While temperature-selected PopII-dominated DLAs show halo mass and SFR distributions very similar to those of PopIII-dominated ones and ≈ 1 dex larger metallicities, mass-selected PopII-dominated DLAs span a larger mass range, i.e. ($10^{8.6}$ to $10^{10.6}$) M_{\odot} , and favour larger metallicities ($> 10^{-3.5} Z_{\odot}$) and SFRs ($> 10^{-2.5} M_{\odot} \text{ yr}^{-1}$).

(iv) No matter which definition we adopt, the PopIII-dominated DLAs could be easily identified through their abundance ratios in the VMSN and MSN models (and possibly also in the RSN model), e.g. by comparing [Si/O] versus [C/O], [Si/Fe] versus [O/Fe] or [Si/C]

versus [C/Fe]. Besides, the average abundance ratios of DLAs (e.g. [C/O], [Si/C] and [C/Fe]) from all the simulations display a visible difference at high redshift, but converge quickly with decreasing redshift.

(v) While most currently observed DLAs at $z \geq 5$ (Becker et al. 2012; Morrison et al. 2016) have metal ratios consistent with a PopII stellar enrichment, the one presented in Morrison et al. (2016) can be explained only with the MSN model and at least one DLA presented in Becker et al. (2012) can be well described by the RSN model. None though supports the VMSN model. Although more precise observations are necessary in the future, these results seem consistent with the conclusions drawn from the most metal-poor DLA currently known (Cooke, Pettini & Steidel 2017), which feature no significant signatures from massive first stars and point towards enrichment from lower-mass stars to explain observed low-metallicity gas patterns.

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REFERENCES

Abel T., Bryan G. L., Norman M. L., 2002, *Science*, 295, 93
 Asplund M., Grevesse N., Sauval A. J., Scott P., 2009, *ARA&A*, 47, 481
 Becker G. D., Sargent W. L. W., Rauch M., Carswell R. F., 2012, *ApJ*, 744, 91
 Biffi V., Maio U., 2013, *MNRAS*, 436, 1621
 Bocchio M., Marassi S., Schneider R., Bianchi S., Limongi M., Chieffi A., 2016, *A&A*, 587, A157
 Bouwens R. J. et al., 2014, *ApJ*, 795, 126
 Bouwens R. J. et al., 2015, *ApJ*, 803, 34
 Bouwens R. J. et al., 2016, *ApJ*, 833, 72
 Bromm V., 2013, *Rep. Prog. Phys.*, 76, 112901
 Bromm V., Loeb A., 2003, *Nature*, 425, 812
 Campisi M. A., Maio U., Salvaterra R., Ciardi B., 2011, *MNRAS*, 416, 2760
 Capak P. L. et al., 2015, *Nature*, 522, 455
 Clark P. C., Glover S. C. O., Smith R. J., Greif T. H., Klessen R. S., Bromm V., 2011, *Science*, 331, 1040
 Coe D. et al., 2013, *ApJ*, 762, 32
 Cooke R., Pettini M., Steidel C. C., Rudie G. C., Nissen P. E., 2011, *MNRAS*, 417, 1534
 Cooke R. J., Pettini M., Jorgenson R. A., 2015, *ApJ*, 800, 12
 Cooke R. J., Pettini M., Steidel C. C., 2017, *MNRAS*,
 Dayal P., Dunlop J. S., Maio U., Ciardi B., 2013, *MNRAS*, 434, 1486
 de Avillez M. A., Mac Low M.-M., 2002, *ApJ*, 581, 1047
 de Souza R. S., Ciardi B., Maio U., Ferrara A., 2013, *MNRAS*, 428, 2109
 de Souza R. S., Maio U., Biffi V., Ciardi B., 2014, *MNRAS*, 440, 240
 Duncan K. et al., 2014, *MNRAS*, 444, 2960
 Ellis R. S. et al., 2013, *ApJ*, 763, L7
 Fan X. et al., 2001, *AJ*, 121, 54
 Ferrara A., Viti S., Ceccarelli C., 2016, *MNRAS*, 463, L112
 Finlator K., Muñoz J. A., Oppenheimer B. D., Oh S. P., Özel F., Davé R., 2013, *MNRAS*, 436, 1818

Finlator K., Thompson R., Huang S., Davé R., Zackrisson E., Oppenheimer B. D., 2015, *MNRAS*, 447, 2526
 Foltz C. B., Chaffee F. H. Jr, Weymann R. J., 1986, *AJ*, 92, 247
 Font-Ribera A. et al., 2012, *J. Cosmology Astropart. Phys.*, 11, 59
 Heger A., Woosley S. E., 2002, *ApJ*, 567, 532
 Heger A., Woosley S. E., 2010, *ApJ*, 724, 341
 Hirano S., Hosokawa T., Yoshida N., Omukai K., Yorke H. W., 2015, *MNRAS*, 448, 568
 Ishiyama T., Sudo K., Yokoi S., Hasegawa K., Tominaga N., Susa H., 2016, *ApJ*, 826, 9
 Keating L. C., Haehnelt M. G., Becker G. D., Bolton J. S., 2014, *MNRAS*, 438, 1820
 Keating L. C., Puchwein E., Haehnelt M. G., Bird S., Bolton J. S., 2016, *MNRAS*, 461, 606
 Kennicutt R. C., Evans N. J., 2012, *ARA&A*, 50, 531
 Kobayashi C., Tominaga N., Nomoto K., 2011, *ApJ*, 730, L14
 Koopmans L. et al., 2015, *Advancing Astrophysics with the Square Kilometre Array (AASKA14)*, p. 1
 Kulkarni G., Rollinde E., Hennawi J. F., Vangioni E., 2013, *ApJ*, 772, 93
 Kulkarni G., Hennawi J. F., Rollinde E., Vangioni E., 2014, *ApJ*, 787, 64
 Laporte N. et al., 2016, *ApJ*, 820, 98
 Ma Q., Maio U., Ciardi B., Salvaterra R., 2015, *MNRAS*, 449, 3006
 Ma Q., Maio U., Ciardi B., Salvaterra R., 2017, *MNRAS*, 466, 1140
 Madau P., Dickinson M., 2014, *ARA&A*, 52, 415
 Maio U., Barkov M. V., 2014, *MNRAS*, 439, 3520
 Maio U., Tescari E., 2015, *MNRAS*, 453, 3798
 Maio U., Viel M., 2015, *MNRAS*, 446, 2760
 Maio U., Dolag K., Ciardi B., Tornatore L., 2007, *MNRAS*, 379, 963
 Maio U., Ciardi B., Dolag K., Tornatore L., Khochfar S., 2010, *MNRAS*, 407, 1003
 Maio U., Koopmans L. V. E., Ciardi B., 2011a, *MNRAS*, 412, L40
 Maio U., Khochfar S., Johnson J. L., Ciardi B., 2011b, *MNRAS*, 414, 1145
 Maio U., Ciardi B., Müller V., 2013a, *MNRAS*, 435, 1443
 Maio U., Dotti M., Petkova M., Perego A., Volonteri M., 2013b, *ApJ*, 767, 37
 Maio U., Petkova M., De Lucia G., Borgani S., 2016, *MNRAS*, 460, 3733
 Mancini M., Schneider R., Graziani L., Valiante R., Dayal P., Maio U., Ciardi B., Hunt L. K., 2015, *MNRAS*, 451, L70
 Mancini M., Schneider R., Graziani L., Valiante R., Dayal P., Maio U., Ciardi B., 2016, *MNRAS*, 462, 3130
 Mancuso C., Lapi A., Shi J., Gonzalez-Nuevo J., Aversa R., Danese L., 2016, *ApJ*, 823, 128
 McLure R. J. et al., 2013, *MNRAS*, 432, 2696
 Morrison S., Kulkarni V. P., Som D., DeMarcy B., Quiert S., Péroux C., 2016, *ApJ*, 830, 158
 Noterdaeme P., Ledoux C., Petitjean P., Srianand R., 2008, *A&A*, 481, 327
 Oesch P. A. et al., 2013, *ApJ*, 773, 75
 Oesch P. A. et al., 2014, *ApJ*, 786, 108
 Padovani P., Matteucci F., 1993, *ApJ*, 416, 26
 Péroux C., McMahon R. G., Storrie-Lombardi L. J., Irwin M. J., 2003, *MNRAS*, 346, 1103
 Pontzen A. et al., 2008, *MNRAS*, 390, 1349
 Rowan-Robinson M. et al., 2016, *MNRAS*, 461, 1100
 Salvadori S., Ferrara A., 2012, *MNRAS*, 421, L29
 Salvaterra R., Maio U., Ciardi B., Campisi M. A., 2013, *MNRAS*, 429, 2718
 Schneider R., Omukai K., Inoue A. K., Ferrara A., 2006, *MNRAS*, 369, 1437
 Smith H. E., Cohen R. D., Bradley S. E., 1986, *ApJ*, 310, 583
 Spite M., Caffau E., Bonifacio P., Spite F., Ludwig H.-G., Plez B., Christlieb N., 2013, *A&A*, 552, A107
 Springel V., 2005, *MNRAS*, 364, 1105
 Stacy A., Bromm V., 2014, *ApJ*, 785, 73
 Thielemann F.-K. et al., 2003, *Nucl. Phys. A*, 718, 139
 Tornatore L., Borgani S., Dolag K., Matteucci F., 2007, *MNRAS*, 382, 1050
 Tselikhovich D., Hirata C., 2010, *Phys. Rev. D*, 82, 083520

van den Hoek L. B., Groenewegen M. A. T., 1997, *A&AS*, 123, 305
Wang F. Y., Bromm V., Greif T. H., Stacy A., Dai Z. G., Loeb A., Cheng
K. S., 2012, *ApJ*, 760, 27
Wise J. H., Turk M. J., Norman M. L., Abel T., 2012, *ApJ*, 745
Wise J. H., Demchenko V. G., Halicek M. T., Norman M. L., Turk M. J.,
Abel T., Smith B. D., 2014, *MNRAS*, 442, 2560
Wolfe A. M., Turnshek D. A., Smith H. E., Cohen R. D., 1986, *ApJS*, 61,
249

Wolfe A. M., Gawiser E., Prochaska J. X., 2005, *ARA&A*, 43, 861
Woosley S. E., Weaver T. A., 1995, *ApJS*, 101, 181
Yoshida N., Omukai K., Hernquist L., Abel T., 2006, *ApJ*, 652, 6
Zheng W. et al., 2012, *Nature*, 489, 406

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