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LETTER TO THE EDITOR

Deuterated methanol on Solar System scale around the HH212 protostar

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ABSTRACT

Context. Deuterium fractionation is a precious tool to understand the chemical evolution during the process leading to the formation of a Sun-like planetary system.

Aims. Methanol is thought to be mainly formed during the prestellar phase and its deuterated form keeps memory of the conditions at that epoch. Thanks to the unique combination of high angular resolution and sensitivity provided by ALMA, we wish to measure methanol deuteration in the planet formation region around a Class 0 protostar and to understand its origin.

Methods. We mapped both the $^{13}\text{CH}_3\text{OH}$ and CH_2DOH distribution in the inner regions (~ 100 au) of the HH212 system in Orion B. To this end, we used ALMA Cycle 1 and Cycle 4 observations in Band 7 with angular resolution down to $\sim 0''.15$.

Results. We detected 6 lines of $^{13}\text{CH}_3\text{OH}$ and 13 lines of CH_2DOH with upper level energies up to 438 K in temperature units. We derived a rotational temperature of (171 ± 52) K and column densities of $7 \times 10^{16} \text{ cm}^{-2}$ ($^{13}\text{CH}_3\text{OH}$) and $1 \times 10^{17} \text{ cm}^{-2}$ (CH_2DOH), respectively. Consequently, the D/H ratio is $(2.4 \pm 0.4) \times 10^{-2}$, a value lower by an order of magnitude with respect to what was previously measured using single dish telescopes toward protostars located in Perseus. Our findings are consistent with the higher dust temperatures in Orion B with respect to that derived for the Perseus cloud. The emission is tracing a rotating structure extending up to 45 au from the jet axis and elongated by 90 au along the jet axis. So far, the origin of the observed emission appears to be related with the accretion disk. Only higher spatial resolution measurements however, will be able to disentangle between different possible scenarios: disk wind, disk atmosphere, or accretion shocks.

Key words. Stars: formation – ISM: abundances – ISM: molecules – ISM: individual objects: HH212

1. Introduction

Molecular deuteration is a powerful diagnostic tool to study the past history of the gas associated with the formation of a proto-Sun and its protoplanetary system (see e.g. Ceccarelli et al. 2014, and references therein). In the prestellar core phase, the low temperatures and the resulting CO freeze-out enhance the abundance of the deuterated molecules. These molecules are stored in the icy grain mantles and then released into the gas phase in the protostellar stage (Ceccarelli et al. 2007; Caselli et al. 2008). This is the case of methanol (CH_3OH), formed on the grain surfaces (e.g. Tielens 1983; Rimola et al. 2014) and then released into the gas phase because dust mantles are thermally evaporated in the inner warm regions around low-mass protostars (e.g. Ceccarelli

et al. 2000, 2007; Parise et al. 2002, 2004, 2006) and/or sputtered by shocks (Codella et al. 2012; Fontani et al. 2014).

Deuterated methanol has been analysed towards Class 0 protostars (Parise et al. 2006): the measured D/H of 0.4–0.6 indicates that CH_2DOH can be almost as abundant as CH_3OH . In the only one more evolved source studied so far, the same ratio appears to be smaller ($\text{D/H} \sim 1.7 \times 10^{-3}$), suggesting a decrease of the deuterated species due to evolutionary effects (Bianchi et al. 2017). However, all the existing observations have been performed with single-dish telescopes (angular resolution $\geq 10''$) which do not disentangle the different components associated with a protostellar system (jet, high-velocity shocks, slower accretion shocks, inner envelope). Sub-arcsecond resolution observations are then crucial to measure the D/H ratio within 50–100 au from the protostar, i.e. the region hosting the protoplanetary disks. Dust grains and ices may also be affected by the accretion

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shock, near the centrifugal barrier, so we do not know what the deuterium fractionation is at small scales.

The HH212 star-forming region in Orion ($d = 450$ pc) is an ideal laboratory to study the jet-disk system. The HH212-mm Class 0 protostar is driving a bright, extended, and bipolar molecular jet extensively observed using IRAM-NOEMA, SMA, and ALMA (e.g. Lee et al. 2006, 2007, 2008, 2015, 2016, 2017a, 2017b; Codella et al. 2007; Cabrit et al. 2007, 2012). A molecular disk was revealed in HCO^+ , C^{17}O , and SO emission, using the first ALMA cycles with angular resolutions of $\sim 0''.6$. Velocity gradients have been detected along the equatorial plane consistently with a rotating disk around a $0.2\text{--}0.3 M_\odot$ protostar (Lee et al. 2014; Codella et al. 2014; Podio et al. 2015). Indeed, a disk with a radius of 60 au has been imaged by Lee et al. (2017a) observing dust continuum emission at $850 \mu\text{m}$. Finally, Codella et al. (2016) and Leurini et al. (2016) suggested that HDO and CH_3OH emission is likely associated with outflowing gas, and possibly with a disk wind. To conclude, the HH212 region is the unique Class 0 protostellar region where a bright jet, a compact rotating disk, and signatures of a disk wind have been revealed. In this Letter, we exploit ALMA Band 7 observations with an angular resolution down to $0''.15$ to obtain the first measurement of methanol deuteration in the disk formation region.

2. Observations

HH212 was observed in Band 7 with ALMA using 34 12-m antennas between 15 June and 19 July 2014 during the Cycle 1 phase. Further observations of HH212 were performed in Band 7 with ALMA using 44 12-m antennas between 6 October and 26 November 2016 during the Cycle 4 phase. The maximum baselines for Cycle 1 and 4 were 650 m and 3 km, respectively.

In Cycle 1 the spectral windows between 337.1–338.9 GHz and 348.4–350.7 GHz were observed using spectral channels of 488 kHz ($0.42\text{--}0.43 \text{ km s}^{-1}$), subsequently smoothed to 1.0 km s^{-1} . Calibration was carried out following standard procedures, using quasars J0607-0834, J0541-0541, J0423-013, and Ganymede. The continuum-subtracted images have clean-beam FWHMs from $0''.41 \times 0''.33$ to $0''.52 \times 0''.34$ ($\text{PA} = -63^\circ$), and an rms noise level of $5\text{--}6 \text{ mJy beam}^{-1}$ in 1.0 km s^{-1} channels. For Cycle 4 spectral units of 977 kHz (0.87 km s^{-1}) were used to observe the spectral windows between 334.1–336.0 GHz. Calibration was carried out following standard procedures, using quasars J0510+1800, J0552+0313, J0541-0211 and J0552-3627. The continuum-subtracted images have a typical clean-beam FWHM of $0''.15 \times 0''.12$ ($\text{PA} = -88^\circ$), and an rms noise level of 1 mJy beam^{-1} in 0.87 km s^{-1} channels. Spectral line imaging was achieved with the CASA package, while data analysis was performed using the GILDAS¹ package. Positions are given with respect to the MM1 protostar continuum peak located at $\alpha(\text{J2000}) = 05^h 43^m 51\rlap{.}^s 41$, $\delta(\text{J2000}) = -01^\circ 02' 53\rlap{.}'' 17$ (Lee et al. 2014). Lines were identified using spectroscopic parameters extracted from the Jet Propulsion Laboratory (JPL², Pickett et al. 1998) and Cologne Database for Molecular Spectroscopy (CDMS³; Müller et al. 2001, 2005) molecular databases.

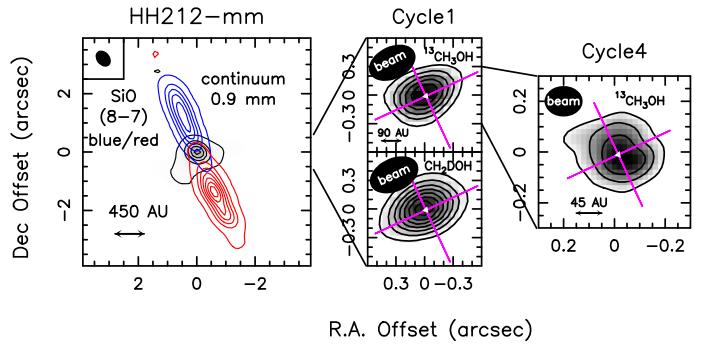


Fig. 1: *Left Panel:* The HH212 protostellar system as observed by ALMA-Band 7 (Codella et al. 2014). Blue/red contours plot the blue/red-shifted $\text{SiO}(8\text{--}7)$ jet overlaid on the continuum at 0.9 mm (black contours). Positions are with respect to the coordinates reported in Sect. 2. The filled ellipse shows the synthesised beam (HPBW): $0''.61 \times 0''.45$. *Central Panels:* Zoom-in of the central region as observed by ALMA-Band 7 Cycle 1: $^{13}\text{CH}_3\text{OH}(13_{0,13}\text{--}12_{1,12})\text{E}$ and $\text{CH}_2\text{DOH}(9_{0,9}\text{--}8_{1,8})\text{e0}$ integrated over $\pm 4 \text{ km s}^{-1}$ with respect to the $v_{\text{sys}} = +1.7 \text{ km s}^{-1}$, (black contours and grey scale). First contours and steps are 5σ for $^{13}\text{CH}_3\text{OH}$ ($15 \text{ mJy beam}^{-1} \text{ km s}^{-1}$) and CH_2DOH ($35 \text{ mJy beam}^{-1} \text{ km s}^{-1}$). The HPBWs are: $0''.52 \times 0''.34$ ($\text{PA} = -64^\circ$) for $^{13}\text{CH}_3\text{OH}$ and $0''.54 \times 0''.35$ ($\text{PA} = -65^\circ$) for CH_2DOH . *Right Panel:* Zoom-in of the central region, as observed by ALMA-Band 7 Cycle 4, showing the $^{13}\text{CH}_3\text{OH}(12_{1,11}\text{--}12_{0,12})\text{A}$ emission integrated over 10 km s^{-1} around v_{sys} (black contours and grey scale). First contours and steps are 5σ ($20 \text{ mJy beam}^{-1} \text{ km s}^{-1}$) and 3σ , respectively. The HPBW is $0''.15 \times 0''.12$ ($\text{PA} = -88^\circ$).

3. Results and Discussion

3.1. Line spectra and maps

In Cycle 1 data, we detected 5 lines of $^{13}\text{CH}_3\text{OH}$ and 12 lines of CH_2DOH covering excitation energies, E_{up} , from 17 to 438 K (see Table A.1). In addition, in the Cycle 4 dataset we revealed one transition of $^{13}\text{CH}_3\text{OH}$ ($E_{\text{up}} = 193 \text{ K}$), and one of CH_2DOH ($E_{\text{up}} = 327 \text{ K}$). Examples of the profiles of the detected lines are shown in Fig. A.1: the emission peaks at velocities between $\approx +1 \text{ km s}^{-1}$ and $+3 \text{ km s}^{-1}$, consistently with the systemic velocity of $+1.7 \text{ km s}^{-1}$ (Lee et al. 2014), once considered the spectral resolution of 1 km s^{-1} . The profiles, fit using the GILDAS package, are Gaussian with typical Full Width at Half Maximum (FWHM) $\sim 5 \text{ km s}^{-1}$. The spectral parameters of the detected lines are presented in Table A.1. Figure 1 introduces the HH212 system: the 0.9 mm continuum traces the envelope hosting the protostar. The bipolar blue- and red-shifted lobes are traced by SiO (Codella et al. 2007; Lee et al. 2014). The central panels, covering a zoom-in of the central protostellar region ($\lesssim 0''.5$, i.e. 225 au), show the Cycle 1 data, specifically examples of the emission maps of $^{13}\text{CH}_3\text{OH}$ and CH_2DOH . The emission is spatially unresolved with a size smaller than the Cycle 1 beam ($0''.5 \times 0''.3$) corresponding to a size of 225×135 au. However, the definitely higher angular resolution provided by the Cycle 4 dataset ($0''.15 \times 0''.12$ corresponding to a radius of 34×27 au) provides spatially resolved images of both $^{13}\text{CH}_3\text{OH}$ and CH_2DOH ⁴ emission lines. We used the $^{13}\text{CH}_3\text{OH}$ emission as revealed in the Cycle 4 image to show in the right panel of Fig. 1 a further zoom-in, sampling the inner 100 au.

¹ <http://www.iram.fr/IRAMFR/GILDAS>

² <https://spec.jpl.nasa.gov/>

³ <http://www.astro.uni-koeln.de/cdms/>

⁴ The CH_2DOH line is blended with the $\text{SO}_2(8_{2,6}\text{--}7_{1,7})$ line (Tab. A.1).

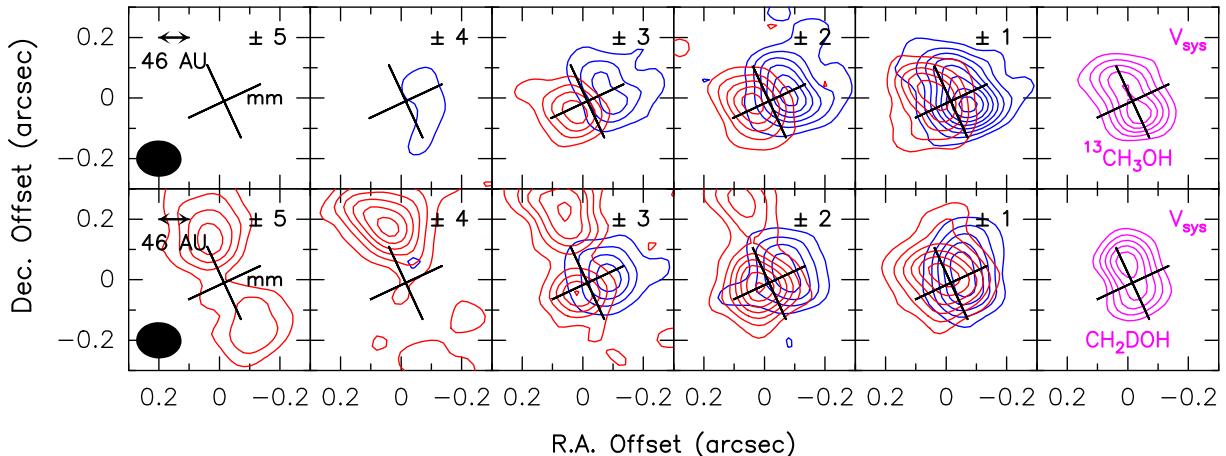


Fig. 2: Channel maps of the $^{13}\text{CH}_3\text{OH}(12_{1,11}-12_{0,12})\text{A}$ (upper panels) and $\text{CH}_2\text{DOH}(16_{2,15}-16_{1,15})\text{o1}$ (lower panels) blue- and redshifted (continuum subtracted) emissions observed during ALMA-Cycle 4 towards the HH212-mm protostar. Each panel shows the emission integrated over a velocity interval of 1 km s^{-1} shifted with respect to the systemic velocity (see the magenta channel, sampling the velocity between $+1.5 \text{ km s}^{-1}$ and $+2.5 \text{ km s}^{-1}$) by the value given in the upper-right corner. The black cross (inclined in order to point the SiO jet direction and consequently the equatorial plane, see Fig. 1) indicates the position of the protostar. The ellipse in the top left panel shows the ALMA synthesised beam (HPBW): $0''.15 \times 0''.12$ ($\text{PA} = -88^\circ$). First contours and steps correspond to 3σ ($2.7 \text{ mJy beam}^{-1} \text{ km s}^{-1}$). Note that for CH_2DOH at velocities larger than $+1.0 \text{ km s}^{-1}$ a further emission appears along the jet axes to the north and to the south with respect to the CH_2DOH emission. This is due to the $\text{SO}_2(8_{2,6}-7_{1,7})$ emission line (see Tab. 1 and Fig. 2), which confirms to be a good tracer of outflowing material (Podio et al. 2015). Given that the velocity scale has been derived by first fixing the CH_2DOH frequency, the SO_2 emission looks artificially red, but it is, obviously, blue-shifted (see Fig. A.1).

3.2. Methanol deuteration

We used the $^{13}\text{CH}_3\text{OH}$ lines detected in Cycle 1 to perform a rotation diagram analysis, assuming LTE conditions and optically thin lines. In order to exclude opacity effects, we use the $^{13}\text{CH}_3\text{OH}$ lines to derive the CH_3OH column density as the main isotopologue lines observed by Lurani et al. (2016) are moderately optically thick ($\tau < 0.4$). We also verified that the CH_2DOH transitions are indeed optically thin using a Population Diagram analysis (Goldsmith & Langer 1999) in which opacities are self-consistently computed. We used $^{13}\text{CH}_3\text{OH}$ and CH_2DOH lines extracted from the same ALMA dataset in order to have the same $u-v$ coverage and to minimise calibration effects. Moreover, the Cycle 0 (beam $\approx 0''.6$) observations of CH_3OH (Lurani et al. 2016) trace a much warmer component dominated by lines with E_{up} up to 747 K with respect to the emission lines analysed in this paper. A source size of $0''.19 \pm 0''.02$ has been derived from the $^{13}\text{CH}_3\text{OH}$ Cycle 4 map integrated over the whole emission range and then used to correct the Cycle 1 observed values. This value is in good agreement with $\sim 0''.2$, as derived from CH_3OH (Lurani et al. 2016). Figure A.2 shows the derived rotational temperature and column density which are $T_{\text{rot}} = (171 \pm 52)$ K and $N_{\text{tot}} = (6.5 \pm 2.1) \times 10^{16} \text{ cm}^{-2}$, respectively. Conservatively, we used only Cycle 1 data to perform the fit; however, the $^{13}\text{CH}_3\text{OH}$ Cycle 4 line is in good agreement with Cycle 1 observations. In the lower panel of Fig. A.2 we derive the column density of CH_2DOH , assuming the same rotational temperature and size derived for $^{13}\text{CH}_3\text{OH}$. We excluded from the fitting the two CH_2DOH transitions with $E_{\text{up}} > 300$ K to exclude any possible contamination from non-thermal excitation processes or the occurrence of a component with a different, higher, excitation condition. However the fitting obtained with the low-energies transitions (continuous line) is in agreement with the two excluded transitions, as well as with the Cycle 4 transition, as shown by the dotted line. The derived column density for CH_2DOH is then

$N_{\text{tot}} = (1.1 \pm 0.2) \times 10^{17} \text{ cm}^{-2}$. Note that an LVG analysis (see Ceccarelli et al. 2003 for further details) of the $^{13}\text{CH}_3\text{OH}$ emission indicates a source size of $0''.1-0''.2$, T_{kin} larger than 100 K and densities larger than 10^6 cm^{-3} in agreement with the HDO results reported by Codella et al. (2016).

The column densities derived from the rotation diagrams have been used to derive the methanol D/H for the first time in the inner 100 au around a low-mass protostar. Assuming a $^{12}\text{C}/^{13}\text{C}$ ratio of 70 (Milam et al. 2005) at the galactocentric distance of HH212 (~ 8.3 kpc), the D/H is consequently $(2.4 \pm 0.4) \times 10^{-2}$. This value is in agreement with the upper limit of 0.27 derived by Lee et al. (2017b) using CH_2DOH and three lines of CH_3OH . In addition, our measurement is lower than previous ones of CH_2DOH in other Class 0 objects performed with the IRAM-30m single dish, which indicate $\text{D/H} \approx 40-60\%$ (Parise et al. 2006). Taking into account that, as reported by Belloche et al. (2016), the column densities of CH_2DOH derived in Parise et al. (2006) were overestimated by a factor ~ 2 because of a problem in the spectroscopic parameters, the HH212 deuteration is still one order of magnitude lower. Which are the reasons of this difference? To start with, the single-dish measurements by Parise et al. (2006) are sampling regions ≥ 2000 au: the D/H is calculated assuming that the main isotopologue and the deuterated species come from the same emitting source, but this cannot be verified with low angular resolutions. This issue is overcome with high angular resolution observations, like those presented here that allow us to directly image the emitting region. Interestingly, Jørgensen et al. (2016) report a level of deuteration for glycolaldehyde of $\sim 5\%$ in IRAS16293 in Ophiuchus on 50 au scale. That said, a possible explanation for the lower methanol deuteration measured for HH212-mm could be related to different physical conditions, during the formation of methanol on dust mantles during the prestellar phase. Indeed, a larger gas temperature reduces the atomic gas D/H ratio landing on the grain surfaces, reducing, consequently, the deuteration of methanol. All

the sources observed by Parise et al. (2006) are located in the Perseus star-forming region, which could have experienced different conditions with respect to the Orion B region where HH212 lies. Specifically, while in Perseus the dust temperature is about ~ 12 K (Zari et al. 2016), HH212 is located about one degree north of the high-mass star forming region NGC 2024, and the dust temperature here is ≥ 16 K (Lombardi et al. 2014). As shown e.g. by Taquet et al. (2012, 2014), deuteration (including that of methanol), once the volume density is fixed, decreases as temperature increases. The models indicate that D/H can decrease by up to one order of magnitude by increasing the temperature from 10 K to 20 K. Interestingly, we note that according to Fuente et al. (2014) the methanol deuteration in the hot core in the intermediate mass star forming region NGC7129, FIRS 2 is $\sim 2\%$. However in hot cores associated with massive star forming regions in Orion (Peng et al. 2012), the methanol D/H is between 0.8×10^{-3} and 1.3×10^{-3} , lower by one order of magnitude than the values reported here.

3.3. The emitting region of deuterated methanol

Figure 2 shows the channel maps of the two transitions observed during ALMA Cycle 4, i.e. at the highest spatial resolution presented here: $^{13}\text{CH}_3\text{OH}$ (12_{1,11}–12_{0,12})A and CH_2DOH (16_{2,15}–16_{1,15})o1 (see Table A.1). For both lines the rightmost panels show the spatial distributions of the emission close to systemic velocity. The rest of the panels are for the blue- and red-shifted velocities, imaged up to ± 4 km s $^{-1}$. Note that for CH_2DOH the red-shifted emission is partially contaminated by the SO_2 (8_{2,6}–7_{1,7}) emission (see Table A.1 and Fig. A.1).

The spatial distribution of the methanol isotopologues at the systemic velocity is elongated along the jet axis ($\text{PA} = 22^\circ$). A fit in the image plane gives for $^{13}\text{CH}_3\text{OH}$ a beam deconvolved FWHM size of $0''.18(0''.02) \times 0''.12(0''.02)$ at $\text{PA} = 33^\circ$. For CH_2DOH we derived $0''.20(0''.02) \times 0''.10(0''.02)$ and $\text{PA} = 19^\circ$. These values correspond to FWHM sizes of 81×54 au ($^{13}\text{CH}_3\text{OH}$) and 90×45 au (CH_2DOH). Moving towards the blue- and red-shifted emission we resolve a clear velocity gradient parallel to the equatorial plane (see also Fig. A.3) with shifts of $\pm 0''.1 = 45$ au, in the same sense as the HH212 disk rotation. The typical (beam deconvolved) size of the blue- and red-shifted emission at ± 2 km s $^{-1}$ from systemic is $0''.22(0''.04) \times 0''.14(0''.04)$ and $0''.18(0''.02) \times 0''.10(0''.03)$, respectively, similar to the size at V_{sys} . Rotation was already noted in CH_3OH by Lurini et al. (2016) from unresolved maps at lower resolution ($\simeq 0''.6$). The authors further noted that emission centroids (fitted in $u-v$ plane) moved away from the source, and eventually away from the mid plane, at higher velocity, speculating that the methanol emission is not dominated by a Keplerian disk or the rotating-infalling cavity, and possibly associated with a disk wind. With the present images we reach higher angular resolution which allows us to refine the picture: the systemic velocity looks arising mainly from two peaks at $\pm 0''.05 = 20$ au above and below the disk plane (of which the centroids in Lurini et al. 2016 only traced the barycenter). In addition, channel maps at ± 1 km s $^{-1}$ suggest red-shifted emission peaking in the north and blue-shifted emission peaking in the south, which is not the sense we would naively expect for a wind from HH212-mm. Since the jet axis is so close to the plane of the sky⁵, this might perhaps still be explained by non-axisymmetric structure in a disk wind, but synthetic predictions would be needed to test this. Alternatively, our data suggest that the emission could arise from accretion shocks occurring

at the centrifugal barrier associated with the accretion disk, and heating the gas at temperatures around 100 K (Sakai et al. 2014; Oya et al. 2016). Indeed, the radial extent ~ 50 au of the observed red- and blue-shifted emission is consistent with the dust disk radius ~ 60 au, recently resolved by Lee et al. (2017a) at 850 μm with ALMA. In this case, we would observe a thicker accretion shock near the centrifugal barrier in HH212, as recently seen in C_2H in L1527 (Sakai et al. 2017). We may also have a contribution from a warm disk atmosphere, emitting in the northern/southern portions of the accretion disk, thicker than the dust atmosphere.

4. Conclusions

The ALMA Cycle 1 and 4 observations allow us to measure methanol deuteration in the inner 50 au of the jet-disk system associated with the Class 0 protostar HH212, in Orion B. The deuteration is $\simeq 2 \times 10^{-2}$, a value lower than what previously measured using single-dish towards Class 0 protostars in Perseus. Although we cannot exclude that single-dish observations are mixing different gas components with different D/H values, our findings are consistent with a higher dust temperature in Orion B with respect to the Perseus cloud. This confirms the diagnostic value of molecular deuteration to recover the physical conditions during the pre-collapse phase. The emission is confined in a rotating structure which extends at ± 45 au from the equatorial plane and is elongated along the jet axis. Disk wind, disk atmosphere, and accretion shocks could explain the observed images. Higher spatial resolution maps will be necessary to distinguish between these possibilities.

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⁵ The system inclination is 4° to the plane of sky; Claussen et al. 1998

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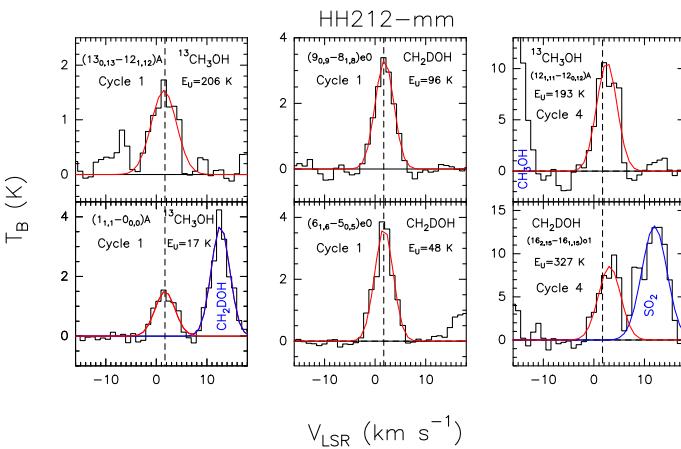


Fig. A.1: Examples of line profiles in T_B scale (not corrected for the beam dilution): species and transitions are reported. The vertical dashed line stands for the systemic V_{LSR} velocity ($+ 1.7 \text{ km s}^{-1}$, Lee et al. 2014). Red and blue curves are for the Gaussian fit. In the spectrum of Cycle 1 ($1_{1,1}-0_{0,0}$)A $^{13}\text{CH}_3\text{OH}$ transition, a CH_2DOH doublet containing the ($5_{4,2}-5_{3,2}$)e1 and the ($5_{4,1}-5_{3,3}$)e1 transitions is also shown (see Table A.1). In Cycle 4 spectra the ($7_{1,7}-6_{1,6}$)A CH_3OH and the ($8_{2,6}-7_{1,7}$) SO_2 transitions are also shown.

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Appendix A: Additional material

Table A.1 lists all the methanol isotopologue lines observed towards HH212-mm during ALMA Cycle 1 and Cycle 4 operations. Fig. A.1 shows examples of line profiles in T_B scale, while Fig. A.2 shows the rotation diagrams for $^{13}\text{CH}_3\text{OH}$ and CH_2DOH . Fig. A.3 shows the $^{13}\text{CH}_3\text{OH}(12_{1,11}-12_{0,12})\text{A}$ spectrum extracted at $\pm 0''.06$ from the protostar, in the equatorial plane direction.

Table A.1: Emission lines detected in ALMA Cycle 1 and 4 measurements towards HH212-mm.

Transition	ν^a (GHz)	E_{up}^a (K)	$S\mu^2a$ (D 2)	rms (K)	F_{int}^b (K km s $^{-1}$)
$^{13}\text{CH}_3\text{OH}$ Cycle 1					
$13_{0,13}-12_{1,12}\text{A}$	338.75995	206	13	0.29	9.6(1.1)
$13_{4,9}-14_{3,11}\text{E}$	347.78840	303	4	0.36	4.0(1.4)
$11_{0,11}-10_{1,9}\text{E}$	348.10019	162	5	0.84	7.5(3.0)
$11_{1,1}-0_{0,0}\text{A}$	350.10312	17	2	1.16	7.4(3.0)
$8_{1,7}-7_{2,5}\text{E}$	350.42158	103	2	0.93	8.7(2.6)
CH_2DOH Cycle 1					
$9_{0,9}-8_{1,8}\text{e0}$	337.34866	96	6	0.15	15.3(0.5)
$9_{1,8}-8_{2,6}\text{o1}^c$	338.46254	120	2	0.26	10.7(2.1) ^c
$13_{1,12}-12_{0,12}\text{e0}$	338.86898	202	2	0.13	9.7(0.4)
$6_{1,6}-5_{0,5}\text{e0}$	338.95711	48	5	0.29	16.7(0.9)
$18_{4,15}-18_{3,15}\text{e1}$	347.76728	438	10	0.09	6.6(0.3)
$18_{4,14}-18_{3,16}\text{e1}$	347.95281	438	10	0.18	6.4(0.6)
$4_{1,3}-4_{0,4}\text{e1}$	348.16076	38	4	0.15	18.1(0.5)
$7_{4,4}-7_{3,4}\text{e1}$	349.95168	132	3	0.16	25.3(0.5)
$7_{4,3}-7_{3,5}\text{e1}$	349.95272	132	3		
$6_{4,3}-6_{3,3}\text{e1}$	350.02735	117	2		
$6_{4,2}-6_{3,4}\text{e1}$	350.02777	117	2		
$5_{4,2}-5_{3,2}\text{e1}$	350.09024	104	2		
$5_{4,1}-5_{3,3}\text{e1}$	350.09038	104	2		
$6_{2,5}-5_{1,5}\text{e1}^c$	350.45387	72	4	0.63	17.0(3.9) ^c
$5_{1,4}-5_{0,5}\text{e1}$	350.63207	49	5	0.60	21.0(2.0)
$^{13}\text{CH}_3\text{OH}$ Cycle 4					
$12_{1,11}-12_{0,12}\text{A}$	335.56021	193	23	0.81	61.7(2.7)
CH_2DOH Cycle 4					
$16_{2,15}-16_{1,15}\text{o1}^d$	334.68395	327	8	0.67	51.1(2.9) ^d

Notes. ^a Frequencies and spectroscopic parameters are extracted from the Jet Propulsion Laboratory molecular database (JPL, Pickett et al. 1998) and from the Cologne Database for Molecular Spectroscopy (CDMS, Müller et al. 2001). Upper level energies refer to the ground state of each symmetry. ^b Gaussian fit. ^c The line transition ($9_{1,8}-8_{2,6}$)o1 is partially blended with the ($7_{-5,2}-6_{-5,1}$)E1 CH_3OH transition. The line ($6_{2,5}-5_{1,5}$)e1 is partially blended with the ($18_{1,17}-17_{1,16}$)E CH_3CHO transition. ^d The CH_2DOH ($16_{2,15}-16_{1,15}$) o1 line ($E_u = 327 \text{ K}$, $S_{ij}\mu^2 = 8 \text{ D}^2$) is partially blended with the SO_2 (8 – 7) transition.

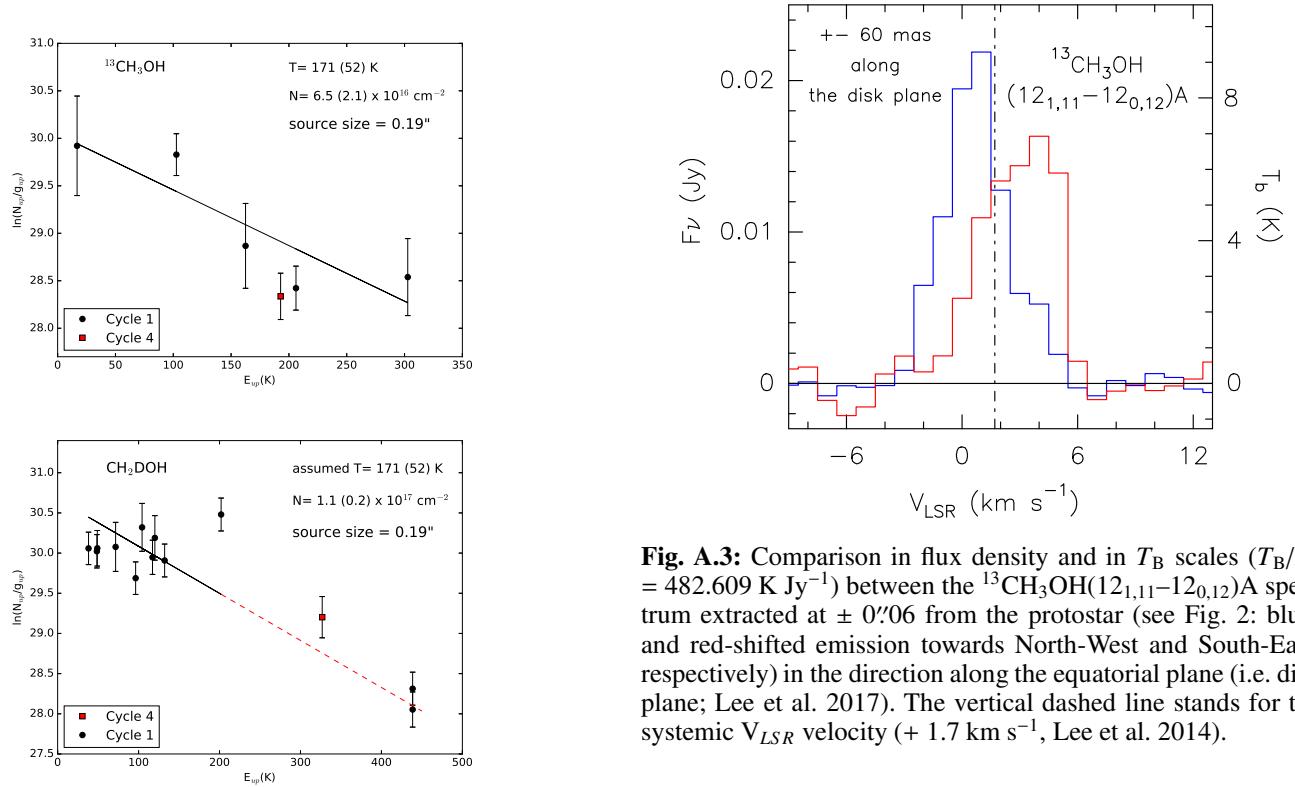


Fig. A.2: Rotation diagrams for $^{13}\text{CH}_3\text{OH}$ (upper panel) and CH_2DOH (lower panel). An emitting region size of $0''.19$ is assumed (see text). The parameters N_u , g_u , and E_{up} are, respectively, the column density, the degeneracy and the energy of the upper level. The derived value of the rotational temperature is reported for $^{13}\text{CH}_3\text{OH}$. For CH_2DOH the T_{rot} derived from the ^{13}C -isotopologues is assumed to derive the column density. The two CH_2DOH transitions with excitation energies higher than 400 K as well as the Cycle 4 point are excluded from the fit although they are in good agreement with it.

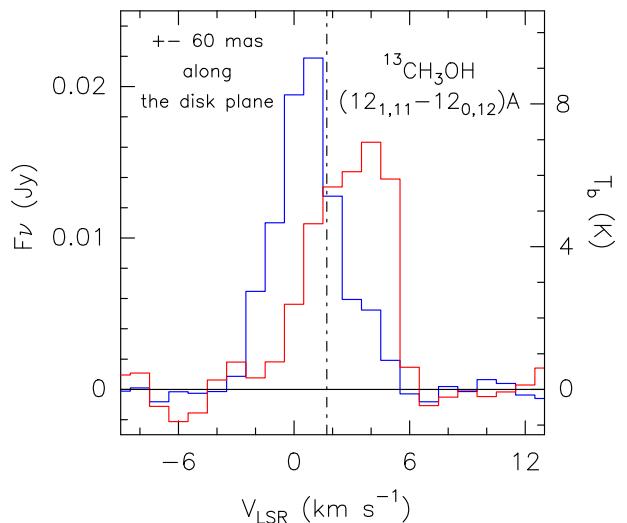


Fig. A.3: Comparison in flux density and in T_B scales ($T_B/F_\nu = 482.609 \text{ K Jy}^{-1}$) between the $^{13}\text{CH}_3\text{OH}(12_{1,11}-12_{0,12})\text{A}$ spectrum extracted at $\pm 0''.06$ from the protostar (see Fig. 2: blue- and red-shifted emission towards North-West and South-East, respectively) in the direction along the equatorial plane (i.e. disk plane; Lee et al. 2017). The vertical dashed line stands for the systemic V_{LSR} velocity ($+ 1.7 \text{ km s}^{-1}$, Lee et al. 2014).