



<b>Publication Year</b>	2018
<b>Acceptance in OA</b>	2020-09-29T09:14:52Z
<b>Title</b>	SN 2015ba: a Type IIP supernova with a long plateau
<b>Authors</b>	Dastidar, Raya, Misra, Kuntal, Hosseinzadeh, G., PASTORELLO, Andrea, Pumo, M. L., Valenti, S., McCully, C., TOMASELLA, Lina, Arcavi, I., ELIAS DE LA ROSA, NANCY DEL CARMEN, Singh, Mridweeka, Gangopadhyay, Anjasha, Howell, D. A., Morales-Garoffolo, Antonia, ZAMPIERI, Luca, Kumar, Brijesh, TURATTO, Massimo, BENETTI, Stefano, TARTAGLIA, Leonardo, Ochner, P., Sahu, D. K., Anupama, G. C., Pandey, S. B.
<b>Publisher's version (DOI)</b>	10.1093/mnras/sty1634
<b>Handle</b>	<a href="http://hdl.handle.net/20.500.12386/27512">http://hdl.handle.net/20.500.12386/27512</a>
<b>Journal</b>	MONTHLY NOTICES OF THE ROYAL ASTRONOMICAL SOCIETY
<b>Volume</b>	479

# SN 2015ba: a Type IIP supernova with a long plateau

Raya Dastidar,<sup>1,2★</sup> Kuntal Misra,<sup>1</sup> G. Hosseinzadeh,<sup>3,4</sup> A. Pastorello,<sup>5</sup> M. L. Pumo,<sup>5,6,7</sup> S. Valenti,<sup>8</sup> C. McCully,<sup>3,4</sup> L. Tomasella,<sup>5</sup> I. Arcavi,<sup>3,4†</sup> N. Elias-Rosa,<sup>5</sup> Mridweeka Singh,<sup>1,9</sup> Anjasha Gangopadhyay,<sup>1,9</sup> D. A. Howell,<sup>3,4</sup> Antonia Morales-Garoffolo,<sup>10</sup> L. Zampieri,<sup>5</sup> Brijesh Kumar,<sup>1</sup> M. Turatto,<sup>5</sup> S. Benetti,<sup>5</sup> L. Tartaglia,<sup>11</sup> P. Ochner,<sup>5</sup> D. K. Sahu,<sup>12</sup> G. C. Anupama<sup>12</sup> and S. B. Pandey<sup>1</sup>

<sup>1</sup>Aryabhata Research Institute of observational sciences, Manora Peak, Nainital 263 001, India

<sup>2</sup>Department of Physics & Astrophysics, University of Delhi, Delhi 110 007, India

<sup>3</sup>Las Cumbres Observatory, 6740 Cortona Dr, Suite 102, Goleta, CA 93117-5575, USA

<sup>4</sup>Department of Physics, University of California, Santa Barbara, CA 93106-9530, USA

<sup>5</sup>INAF Osservatorio Astronomico di Padova, Vicolo dell'Osservatorio 5, I-35122 Padova, Italy

<sup>6</sup>Università degli Studi di Catania, DIEEI and DFA, Via Santa Sofia 64, I-95123 Catania, Italy

<sup>7</sup>INFN-Laboratori Nazionali del Sud, Via Santa Sofia 62, I-95123 Catania, Italy

<sup>8</sup>Department of Physics, University of California, 1 Shields Ave, Davis, CA 95616-5270, USA

<sup>9</sup>Pt.Ravi Shankar Shukla University, Raipur 492 010, India

<sup>10</sup>Department of Applied Physics, University of Cádiz, Campus of Puerto Real, E-11510 Cádiz, Spain

<sup>11</sup>Department of Astronomy, The Oskar Klein Centre, AlbaNova, SE-106 91 Stockholm, Sweden

<sup>12</sup>Indian Institute of Astrophysics, Koramangala, Bangalore 560 034, India

Accepted 2018 June 13. Received 2018 June 13; in original form 2018 April 20

## ABSTRACT

We present optical photometry and spectroscopy from about a week after explosion to  $\sim 272$  d of an atypical Type IIP supernova, SN 2015ba, which exploded in the edge-on galaxy IC 1029. SN 2015ba is a luminous event with an absolute *V*-band magnitude of  $-17.1 \pm 0.2$  mag at 50 d since explosion and has a long plateau lasting for  $\sim 123$  d. The distance to the SN is estimated to be  $34.8 \pm 0.7$  Mpc using the expanding photosphere and standard candle methods. High-velocity H Balmer components constant with time are observed in the late-plateau phase spectra of SN 2015ba, which suggests a possible role of circumstellar interaction at these phases. Both hydrodynamical and analytical modelling suggest a massive progenitor of SN 2015ba with a pre-explosion mass of 24–26  $M_{\odot}$ . However, the nebular spectra of SN 2015ba exhibit insignificant levels of oxygen, which is otherwise expected from a massive progenitor. This might be suggestive of the non-monotonical link between O-core masses and the zero-age main sequence mass of pre-supernova stars and/or uncertainties in the mixing scenario in the ejecta of supernovae.

**Key words:** techniques: photometric – techniques: spectroscopic – supernovae: general – supernovae: individual: SN 2015ba – galaxies: individual: IC 1029.

## 1 INTRODUCTION

Type IIP Supernovae (SNe IIP, hereafter) originate from precursor stars that had retained a substantial amount of their hydrogen layers (greater than  $\sim 3\text{--}5 M_{\odot}$ , e.g. Smartt 2009; Pumo & Zampieri 2013) before exploding as core-collapse SNe (CCSNe), and are characterized by a plateau in the light curve typically lasting  $\sim 100$  d (e.g. SNe 1999em, 2004et, 2014cx; Nadyozhin 2003; Arcavi et al. 2012; Huang et al. 2016). During this phase, the *VRI*-band luminosity

remains constant to within  $\leq 0.5$  mag (Li et al. 2011; Faran et al. 2014b) and P-Cygni features in the spectra become conspicuous. The plateau phase is ascribed to the recombination wave that moves inward in mass coordinates, while the photosphere stays at roughly the same radius and temperature, resulting in the release of internal energy at a roughly constant rate. This phase ends when all of the hydrogen envelope has recombined, and is followed by a sudden 2–5 mag drop in brightness in  $< 30$  d (Olivares et al. 2010). A recent study of SNe II by Valenti et al. (2016) suggests a 1.0–2.6 mag drop in magnitude in luminous SNe IIP (e.g. 1.6 mag in SN 2004et) from the plateau to the radioactive tail, while their less luminous, lower velocity cousins exhibit a larger drop ranging between 3 and 5 mag

\* E-mail: rdastidr@gmail.com

† Einstein Fellow.

(e.g. 3.83 mag in SN 2005cs). The SN then enters the nebular phase, with the light curve being powered by the decay of  $^{56}\text{Co}$  to  $^{56}\text{Fe}$ , which is characterized by an exponential decay.

Early theoretical works of Grassberg, Imshennik & Nadyozhin (1971), Chevalier (1976), Falk & Arnett (1977), and Arnett (1980) suggest that SNe IIP stem from red supergiant (RSG) stars. At present, there are about a dozen confirmed cases of RSGs as SN progenitors detected from direct imaging. The first discovered was the 8–9  $M_{\odot}$  RSG precursor of SN 2003gd (Van Dyk, Li & Filippenko 2003; Smartt et al. 2004; Hendry et al. 2005), while the latest was the  $12.5 \pm 1.2 M_{\odot}$  progenitor to SN 2012aw (Fraser 2016). According to modern stellar evolution models (Heger et al. 2003; Eldridge & Tout 2004), all stars with initial masses in the range  $8 \lesssim M \lesssim 25\text{--}30 M_{\odot}$  are expected to end up as RSGs, subsequently exploding as CCSNe. However, the direct detection of progenitors in the archival pre-explosion images constrained the SN IIP precursor masses close to the theoretical low-mass limit of the core-collapse process viz. 8.5 (e.g. SN 2003gd) to 16.5  $M_{\odot}$  (e.g. SNe 2009kr, 2012ec; Smartt 2009, 2015, and references therein). This has also been inferred from nucleosynthetic yields from analysis of nebular spectra (Dessart, Livne & Waldman 2010; Jerkstrand et al. 2014), while hydrodynamical models yield larger upper limits (Utrobin & Chugai 2008, 2009; Bersten, Benvenuto & Hamuy 2011; Pumo & Zampieri 2011; Pumo et al. 2017).

Large-scale synoptic surveys are discovering large number of SNe events every year, out of which SNe IIP constitute nearly 50 percent of all CCSNe (Li et al. 2011; Graur et al. 2017). A number of studies have been undertaken to constrain the physical properties of these explosions such as the ejected mass, explosion energy, and pre-supernova radius by modelling the observed SN properties, such as their light curves and spectra. A detailed study of SNe IIP has also been done by Hamuy (2003), suggesting that these events exhibit a wide range of plateau luminosities ( $-15 < M_V < -18$ ), ejecta masses (14–56  $M_{\odot}$ ), kinetic energies ( $0.6\text{--}5.5 \times 10^{51}$  erg),  $^{56}\text{Ni}$  masses (0.0016–0.26  $M_{\odot}$ ), and ejecta velocities (2000–8000  $\text{km s}^{-1}$ ). A continuity in the observable properties of SNe is seen from faint, low-velocity, nickel-poor events such as SN 1997D (Turatto et al. 1998; Zampieri et al. 2003; Pastorello et al. 2004) to bright, high-velocity (HV), nickel-rich objects like SN 1992am (Hamuy 2003). This was further confirmed with the analysis of samples of low and standard luminosity SNe IIP by several authors (Pastorello et al. 2004; Anderson et al. 2014a; Faran et al. 2014a; Spiro et al. 2014; Pumo et al. 2017; Tomasella et al. 2018).

The early time spectra of SNe IIP in most cases show a blue continuum, with few broad P-Cygni features. A blue notch in the absorption component of H Balmer lines has been observed in few SNe II, which is attributed to Si II/Ba II  $\lambda 6497$  when observed at early phases, and to HV H I when observed at late phases, possibly arising from weak circumstellar interaction (Baron et al. 2000; Leonard et al. 2002a; Chugai, Chevalier & Utrobin 2007; Valenti et al. 2014; Gutiérrez et al. 2017). The spectra obtained at late times are also a useful diagnostic to probe the core of the progenitor star. Maguire et al. (2012) and Jerkstrand et al. (2012, 2014) compared their spectral synthesis models with the nebular spectra of SNe II, and estimated progenitor masses that are in good agreement with the mass estimates from direct detection of progenitor. Nevertheless, for the recent event SN 2015bs, which exploded in a low-metallicity environment, the comparison to nebular phase spectral models yielded a zero-age main sequence (ZAMS) mass (17–25  $M_{\odot}$ ; Anderson et al. 2018) that is much higher than that deduced for normal Type II SN using similar methods. Moreover, the recent analysis of 91 nebular spectra of 38 SNe IIP by Silver-

man et al. (2017), indicates the presence of primordial oxygen in the hydrogen layer of massive progenitors of SNe IIP, issuing caveats in the estimation of progenitor masses deduced from the forbidden line fluxes of oxygen.

Although, the observed properties of SNe IIP seem to follow a continuous trend, a handful of peculiar events display rare characteristics in their light curves or spectra. SN 2016X, showcased a long rise time to the *R*-band peak ( $12.6 \pm 0.5$  d; Huang et al. 2018) as compared to normal SNe IIP ( $\sim 7$  d). SN 2009ib, an intermediate luminosity SN IIP, had a longer plateau phase, more massive ejecta and more  $^{56}\text{Ni}$ -rich material compared to intermediate luminosity SNe (Takáts et al. 2015). SN 2009bw, an SN IIP with interaction signatures, recorded the fastest drop from photospheric to the nebular phase ( $\sim 2.2$  mag in 13 d; Inserra et al. 2012). Despite the bright peak ( $M_V = -18$  mag) and plateau magnitudes, the radioactive tail of SN 2007od is fainter than most SNe IIP (Inserra et al. 2011). The explosion energy, ejecta mass, and  $^{56}\text{Ni}$  mass of SN 2009js is found to be similar to subluminescent SNe, while exhibiting a significantly higher luminosity, comparable to those of intermediate luminosity events (Gandhi et al. 2013). The early spectra of SN 2013am featured relatively narrow P-Cygni profiles and prominent Ca II and Ba II lines, similar to the low-luminosity and low-velocity SN 2005cs, while with a shorter plateau phase and brighter light-curve tail, the photometric properties of this event deviated from those of SN 2005cs (Zhang et al. 2014; Tomasella et al. 2018). These controversial properties of SNe IIP motivated the study of each of them in greater detail.

In this paper, we present the detailed analysis of SN 2015ba, a relatively bright SN IIP, which exploded in the nearly edge-on galaxy IC 1029. The paper is structured as follows: Section 2 presents the data and a short overview of the reduction procedure; in Section 3, the distance, explosion epoch, and the reddening towards SN 2015ba is extensively discussed. The light and colour curves are illustrated in Section 4. The distance to SN 2015ba is estimated using the expanding photosphere method (EPM) and standard candle method (SCM) in Section 5. The temperature and radius evolution is discussed in Section 6. The spectral evolution,  $\text{SN}^{++}$  modelling, and line velocities are discussed in Section 7. The bolometric light curve,  $^{56}\text{Ni}$  mass yield, and the progenitor properties derived using analytical and hydrodynamical modelling are given in Section 8. Finally, a short summary of the work is presented in Section 9.

## 2 SN 2015BA: DATA AND REDUCTION

SN 2015ba was discovered by Koichi Itagaki (Teppo-cho, Yamagata, Japan) on 2015 November 28.8071 UT (JD = 2457355.31) in the galaxy IC 1029 at an unfiltered magnitude of 16.7 mag (CBET 4209<sup>1</sup>). The precise location of the SN was given by T. Yusa and Osaki to be around 19°E, 43°S of the centre of IC 1029. Benetti et al. (2015) obtained an optical spectrum on 2015 December 2.13 UT showing weak and shallow P-Cygni profiles of the Balmer lines superimposed on a blue continuum. This is consistent with an SN II spectrum at early phase. The expansion velocities reported from the measurement of the position of the minima of the H  $\alpha$  and H  $\beta$  Balmer lines were around 8900 and 8400  $\text{km s}^{-1}$ , respectively. The details of SN 2015ba and its host galaxy IC 1029 are given in Table 1.

Our observing campaign of SN 2015ba was triggered 3 d after the discovery, using instruments equipped with broadband *BVR*

<sup>1</sup>[www.cbat.eps.harvard.edu/iau/cbet/004200/CBET004209.txt](http://www.cbat.eps.harvard.edu/iau/cbet/004200/CBET004209.txt)

**Table 1.** Basic information on SN 2015ba and the host galaxy IC 1029. The host galaxy parameters are taken from NED.

Host galaxy	IC 1029
Galaxy type	SABb
Redshift	0.007 949
Major diameter	3.34 arcmin
Minor diameter	0.61 arcmin
Helio. radial velocity	$2383 \pm 7 \text{ km s}^{-1}$
RA(J2000.0)	$14^{\text{h}}32^{\text{m}}29^{\text{s}}.19$
Dec.(J2000.0)	$+49^{\text{d}}53^{\text{m}}34^{\text{s}}.5$
Distance modulus	$32.73 \pm 0.04 \text{ mag}$
Total extinction $E(B - V)$	$0.46 \pm 0.05 \text{ mag}$
SN type	IIP
Offset from nucleus	$19''.2 \text{ E}, 41''.7 \text{ S}$
Date of discovery	2457355.3 (JD)
Estimated date of explosion	2457349.7 (JD)

**Table 2.** The distances to the host galaxy of SN 2015ba determined in this paper. The mean of the values is also shown.

Method	$D$ (Mpc)	$\mu$ (mag)	References
Virgo-infall	36.4 (0.9)	32.80 (0.05)	Makarov et al. (2014)
EPM	36.6 (1.9)	32.82 (0.11)	This paper
SCM	30.1 (1.4)	32.39 (0.10)	This paper
Weighted mean	34.8 (0.7)	32.73 (0.04)	

and *ugriz* filters listed in Table A1. High cadence photometric data were obtained up to 266 d from discovery, beyond which the SN was below the detection limit of 1–2 m class telescopes. The images were pre-processed and reduced as discussed in Appendix A and night to night zero-points calculated using the local standards (see Fig. A2 and Table A2) were applied to obtain the final SN magnitudes as listed in Table A3. The spectroscopic monitoring was conducted at 27 epochs from a number of facilities as listed in Table A1 and the log of spectroscopic observations is presented in Table B1.

### 3 PARAMETERS OF SN 2015BA

#### 3.1 Distance

Using the recessional velocity  $v_{\text{vir}} = 2664 \pm 3 \text{ km s}^{-1}$  given in HyperLeda (Makarov et al. 2014)<sup>2</sup> and Hubble constant  $H_0 = 73.24 \pm 1.74 \text{ km s}^{-1} \text{ Mpc}^{-1}$  (Riess et al. 2016), we obtain a Virgo infall distance of about  $36.4 \pm 0.9 \text{ Mpc}$ . Further, we implement the EPM and the SCM to estimate the distance from early photometric and spectroscopic observations. This will be discussed in detail in Section 5. The EPM and SCM provides a distance of  $36.6 \pm 1.9 \text{ Mpc}$  and  $30.1 \pm 1.4 \text{ Mpc}$ , respectively. We take the weighted mean of distances estimated using the three methods (listed in Table 2), and we hereafter adopt a distance of  $34.8 \pm 0.7 \text{ Mpc}$ , corresponding to a distance modulus  $\mu = 32.73 \pm 0.04 \text{ mag}$ .

#### 3.2 Explosion epoch

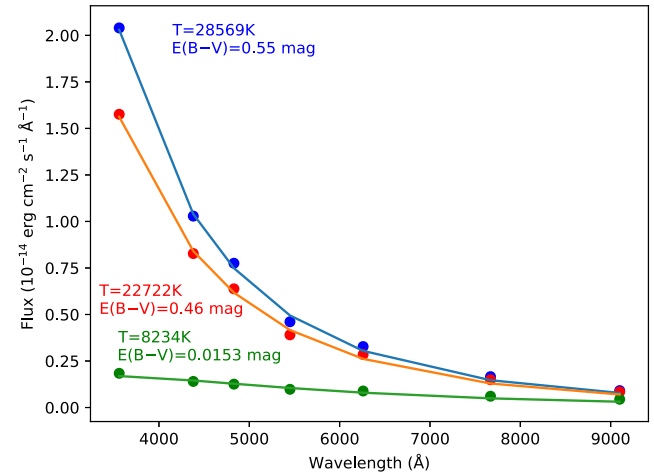
We use the SNID code (Blondin & Tonry 2007) to constrain the explosion epoch of SN 2015ba, which has been tested in the works of Anderson et al. (2014b) and Gutiérrez et al. (2017). This code matches the SN spectrum to a library of spectral templates, and

<sup>2</sup><http://leda.univ-lyon1.fr/>

**Table 3.** The best three matches to the spectrum obtained on 2015 December 4 along with the rlap parameter to estimate the age of SN 2015ba.

SN	rlap <sup>a</sup>	Age since $V_{\text{max}}$ (d)	Age since explosion (d)
SN 2004et	10.2	−1.9	$14.2 \pm 3$
SN 2004et	9.2	−2.9	$13.2 \pm 3$
SN 1999em	9.1	0.2	$10.1 \pm 5$
Mean			$12.5 \pm 7.3$

<sup>a</sup>Quality of fit.



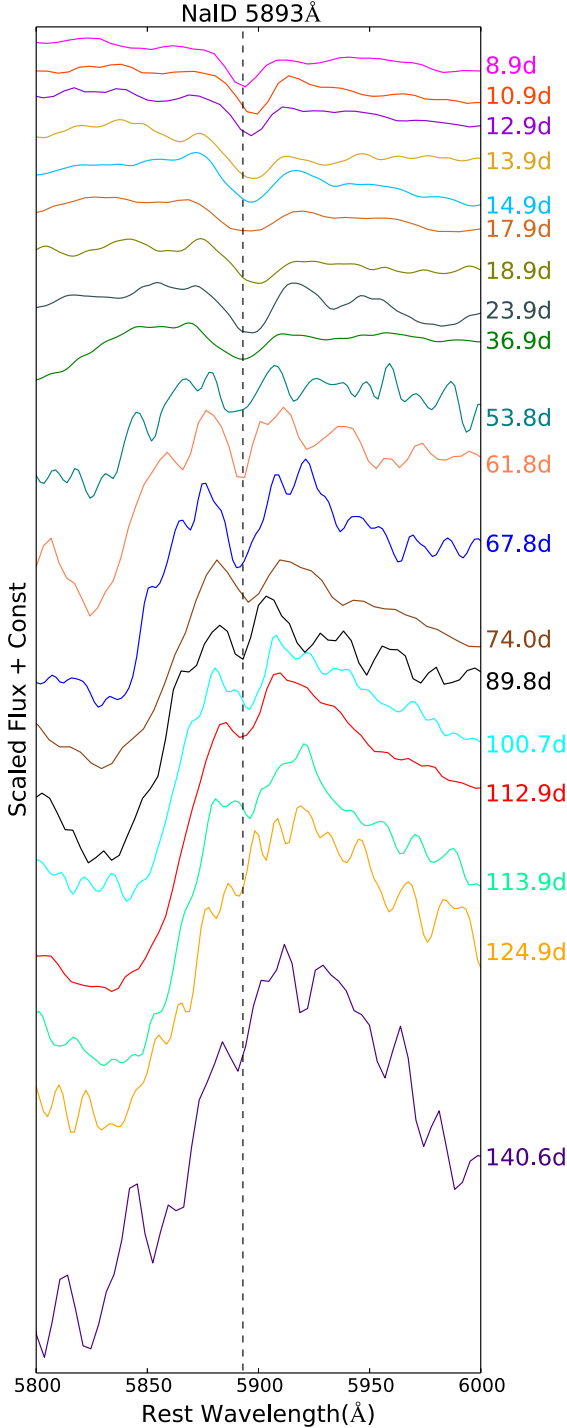
**Figure 1.** The SED of SN 2015ba at 10 d with varying reddening is fitted by a blackbody function to derive the best-fitting temperature. We used  $E(B - V) = 0.0153 \text{ mag}$ , corresponding to Galactic extinction,  $E(B - V) = 0.46 \text{ mag}$  corresponding to Galactic plus host galaxy extinction (as discussed in the text) and  $E(B - V) = 0.55 \text{ mag}$  for which the best-fitting temperature becomes unphysical. The filled circles denote the fluxes in *uBgriz* corrected for the corresponding reddening and the continuous lines are the best-fitting blackbody models.

uses cross-correlation technique to constrain the explosion epoch. Since most spectral lines lie at the blue wavelengths, we perform our fits in the wavelength range 3500–6000 Å. The quality of the fit is represented by the ‘rlap’ parameter, with higher value implying a better correlation. From the best three matches, we derive a mean value of  $12.5 \pm 7.3 \text{ d}$  from explosion for the spectrum obtained on 2015 December 4 (JD 2457360.67), which provides 2015 November 21.5 UT (JD 2457348.2  $\pm$  7.3) as the explosion date. The best three matches with the rlap parameter are listed in Table 3.

The EPM also gives an approximate explosion epoch (see Section 5 for details), as 2015 November 23.2 UT (JD 2457349.7  $\pm$  1.0), which is consistent with that estimated using SNID. Therefore, we adopt throughout the paper 2015 November 23 (JD 2457349.7  $\pm$  1.0) as the most reliable explosion epoch ( $t_0$ ) of SN 2015ba.

#### 3.3 Extinction

In order to gain insights on the true nature of an event and to derive physical parameters from photometry, it is important to estimate the Galactic and the host extinction in the SN direction. The Galactic extinction in the direction of IC 1029 corresponds to a colour excess of  $E(B - V) = 0.0153 \pm 0.0003 \text{ mag}$  (Schlafly & Finkbeiner 2011). There are various methods for a crude estimation of host galaxy extinction. In our case, we inspect three methods: (i) fitting a blackbody function to the early spectral energy distribution (SED)



**Figure 2.** The narrow Na I D absorption dip due to host galaxy, IC 1029. The spectra on 8.9, 12.9, 14.9, 36.9, and 74.0 d with SNR > 50 at 0.6  $\mu\text{m}$  are used to estimate the host galaxy extinction.

of the SN, (ii) estimating from the equivalent width (EW) of Na I D line, and (iii) using the ‘colour method’ suggested by Olivares et al. (2010).

The comparison of the early spectral shape of a classical SN to a blackbody function can provide an upper bound of the colour excess, as discussed by Eastman, Schmidt & Kirshner (1996). With increasing extinction values, the blackbody spectra will start to

**Table 4.** The EW of the narrow Na I D absorption feature due to host galaxy from the spectra of SN 2015ba.

Date (yyyy-mm-dd)	Phase <sup>a</sup> (d)	SNR <sup>b</sup>	EW (Å)
2015-12-01	8.9	126.3	1.8 (0.2)
2015-12-05	12.9	80.6	2.1 (0.4)
2015-12-08	14.9	86.0	1.9 (0.1)
2015-12-30	36.9	55.8	1.9 (0.7)
2016-02-04	74.0	71.9	1.6 (0.7)
Weighted mean			1.88 (0.06)

Notes. <sup>a</sup>since explosion  $t_0 = 2457349.7$  JD.

<sup>b</sup>At 0.6  $\mu\text{m}$ .

deviate from the continuum, establishing an upper limit to the reddening. More recently, theoretical modelling of Dessart & Hillier (2006) and Bersten et al. (2011) suggest that the maximum likely temperature in the early phases of an SN II is 25–30 kK. We fit a Planck function to the SED constructed from the magnitudes of SN 2015ba obtained on 2015 December 3.21 UT (+10 d since explosion) by varying the amount of reddening. We used the reddening law of Cardelli, Clayton & Mathis (1989) with a total-to-selective extinction ratio ( $R_V$ ) of 3.1 to deredden the observed fluxes. The best-fitting temperature corresponding to each value of the colour excess is shown in Fig. 1. This exercise sets an upper limit of  $E(B - V)$  to 0.55 mag above which the temperature becomes unphysical.

The presence of a narrow Na I D line in the spectra of SN 2015ba at the redshift of the host galaxy (Fig. 2) is a possible indication of extinction within the host. Several empirical relations exist to correlate the EW of Na I D to the colour excess  $E(B - V)$ . The EW of the Na I D lines in five spectra with good signal-to-noise ratio are tabulated in Table 4. However, most of these relations show a large scatter and are unreliable with low-resolution spectra (Poznanski et al. 2011). Munari & Zwitter (1997) showed that the lines saturate for EW of Na I D  $\gtrsim 0.6$  Å, hence leading to an underestimate of the colour excess. Nevertheless, we have listed all the existing relations in Table 5 and calculated the corresponding  $E(B - V)$  values.

The third method implemented in the calculations of the colour excess  $E(B - V)$  is the ‘colour method’. This predicts that all SNe IIP should reach the same intrinsic colour towards the end of the plateau phase. It assumes that the opacity of an SN IIP is dominated by electron scattering, and will attain the recombination temperature of hydrogen at the end of plateau phase. Hence, the difference in the observed and the intrinsic colours may be attributed to reddening due to dust along the line of sight. Using a library of SN II spectra, Olivares et al. (2010) computed a conversion factor between  $E(V - I)$  and  $A_V$  and suggested the following prescription:

$$A_V(V - I) = 2.518[(V - I) - 0.656] \quad (1)$$

$$\sigma(A_V) = 2.518\sqrt{\sigma_{(V-I)}^2 + 0.0053^2 + 0.0059^2}. \quad (2)$$

We calculate the weighted mean of  $(V - I)$  colours (corrected for Galactic extinction) on 112.3, 115.3, and 119.3 d from the explosion corresponding to the end of the plateau phase, which results in  $(V - I) = 1.20 \pm 0.01$  mag. The estimated  $A_{V(\text{host})}$  is  $1.37 \pm 0.20$  and  $E(B - V)_{\text{host}} = 0.44 \pm 0.06$  mag.

The reddening estimate using the EW and the colour method are consistent with each other within the errors, and are listed in Table 5. We, therefore, compute the weighted mean of the host galaxy reddening as  $E(B - V)_{\text{host}} = 0.44 \pm 0.05$  mag. The total

**Table 5.** The host galaxy reddening derived for SN 2015ba.

Method	Formula	$E(B - V)_{\text{host}}$ mag	Reference
Na I D	$E(B - V) = 0.25EW$	$0.48 \pm 0.20$	Barbon et al. (1990)
Na I D	$E(B - V) = 0.16EW - 0.01$	$0.30 \pm 0.20$	Turatto, Benetti & Cappellaro (2003)
Na I D	$E(B - V) = 0.43EW - 0.08$	$0.74 \pm 0.30$	Poznanski et al. (2011)
Colour curve	$A_V(V - I) = 2.518[(V - I) - 0.656]$	$0.44 \pm 0.06$	Olivares et al. (2010)
Weighted mean		$0.44 (0.05)$	

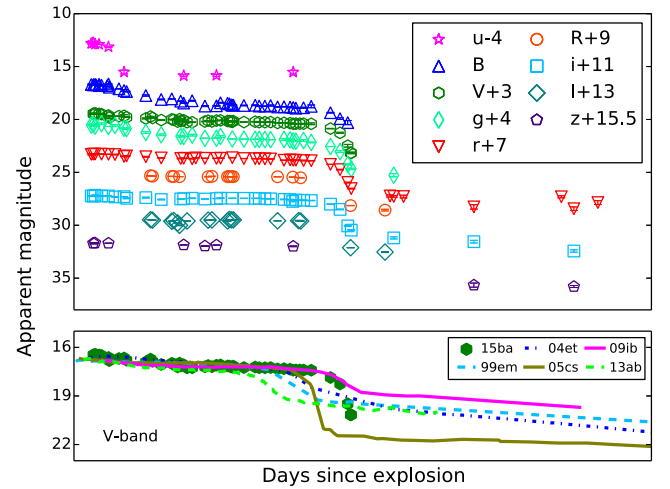
$E(B - V)$  due to host and Milky Way is  $0.46 \pm 0.05$  mag, which is adopted throughout the paper.

#### 4 TEMPORAL EVOLUTION OF SN 2015BA LIGHT CURVE

A careful study of the light-curve properties is essential for the characterization of the event and constrain the properties of the explosion. By comparing the light-curve parameters such as the mid-plateau absolute magnitude in  $V$  band with those of other archetypal SNe IIP, one can subclassify the events as subluminous, normal, or overluminous. Moreover, the steepness of the plateau indicates the extent of the thermalization of the ejecta, with a flatter plateau indicating efficient thermalization and higher  $^{56}\text{Ni}$  mass yield. Hydrodynamical modelling of Bersten et al. (2011) showed that extensive mixing of  $^{56}\text{Ni}$  is required to produce flatter plateaus. For the comparison, we have selected a sample of SNe IIP with variable luminosities, plateau lengths, and plateau slopes. SN 2005cs (Pastorello et al. 2009) is a low-luminosity ( $M_V = -14.83$  mag) Type IIP event while DLT16am (Tartaglia et al. 2018) is one of the brightest ( $M_V = -17.76$  mag) SNe IIP; the plateau length of SN 2007od is  $\sim 25$  d in  $R$  band (Inserra et al. 2011) while the plateau length of SN 2009ib is 118 d in  $V$  band (Takáts et al. 2015); the  $^{56}\text{Ni}$  mass yield of SN 2005cs is  $< 0.003 M_\odot$  (Pastorello et al. 2009) while for SN 2004et  $M(^{56}\text{Ni}) = 0.06 M_\odot$  (Sahu et al. 2006; Misra et al. 2007; Maguire et al. 2010). Further, we include the bright SNe 2007od and 2009bw, the normal luminosity SNe 1999em, 1999gi, 2004dj, 2012aw, 2013ab, 2014cx, and 2016X, and the intermediate luminosity SNe 2008in, 2009N in the sample. Our sample includes a total of 16 events. The parameters of the comparison sample are listed in Table 6.

##### 4.1 Main light-curve features

SN 2015ba was discovered  $\sim 5.6$  d after explosion ( $t_0$  taken from Section 3.2), caught dropping off from the peak and settling onto the plateau. The  $BVRI$  and  $ugriz$  light curves ranging from 8 to 272 d since explosion are shown in the top panel of Fig. 3. We have limited number of measurements at late epochs, and observations were obtained only with the  $griz$  filters. We converted the  $g$ -band data in the tail to  $V$  and  $B$  band using the relations (1) and (6) given in Jordi, Grebel & Ammon (2006). In the bottom panel of Fig. 3, we show the  $V$ -band light curve of SN 2015ba along with other SNe IIP. While the plateau length of SN 2015ba ( $\sim 123$  d in  $V$  band) is similar to that of SN 2009ib ( $\sim 118$  d in  $V$  band; Takáts et al. 2015), the magnitude drop from the plateau to the nebular phase is larger for SN 2015ba ( $\sim 3$  mag). The decline rate in the  $r$  band of SN 2015ba in the first 50 d is  $\sim 0.5$  mag, consistent with the criterion suggested by Li et al. (2011), that the drop in  $R$ -band magnitude in the first 50 d for SNe IIP is  $\leq 0.5$  mag. The  $B$ -band decline rate of



**Figure 3.** Top panel: broadband  $uBVgrRiZ$  light curves of SN 2015ba shifted arbitrarily for clarity. Bottom panel: the  $V$ -band light curve of SN 2015ba compared with other SNe IIP.

SN 2015ba is  $2.9 \text{ mag } 100 \text{ d}^{-1}$  that is also typical of SNe IIP (Patat et al. 1994). The mid-plateau absolute magnitude of SN 2015ba in  $V$  band ( $M_V^p$ ) is  $-17.1 \pm 0.2$  mag, well within the magnitude limits of normal Type IIP SNe ( $-18 < M_V^p < -15$ ; Patat et al. 1994). While, all these numbers suggest that SN 2015ba belong to the plateau subgroup of Type II events, the decline rate in the  $V$  band for SN 2015ba in the first 50 d is  $0.69 \pm 0.05$  mag and is higher than that used by Faran et al. (2014a) to filter out Type IIP events from their sample.

Further, we carried out a comparison of the plateau lengths of SNe IIP sample by fitting the analytic function provided in Valenti et al. (2016), as given below:

$$y(t) = \frac{-a_0}{1 + e^{(t-t_{\text{PT}})/w_0}} + (p_0 \times t) + m_0, \quad (3)$$

where  $t$  is the time from explosion in days,  $t_{\text{PT}}$  is the time in days from explosion to the transition point between the end of the plateau phase and start of the radioactive tail,  $a_0$  is the depth of the drop from plateau to radioactive tail,  $w_0$  is the slope of the drop and  $p_0$  constrains the slope before and after the drop. We fitted this function to the  $V$ -band light curve of SN 2015ba and the best-fitting value of  $a_0$  is  $3.01 \pm 0.09$  mag in  $\sim 20$  d, which corresponds to the upper limit of the range of 2–3 mag drop for typical SNe IIP (Olivares et al. 2010), while this exceeds the range (1–2.6 mag) suggested by Valenti et al. (2016). The same exercise is done for the comparison sample, and the best-fitting values are listed in Table 7. The  $t_{\text{PT}}$  of our sample range from 95 d in SN 2016X to  $\sim 141$  d in SN 2015ba. The drop in magnitude from plateau to nebular phase is maximum

**Table 6.** Parameters of the SNe IIP sample.

Supernova	Parent Galaxy	Distance <sup>a</sup> (Mpc)	$A_V^{\text{tot}}$ (mag)	$M^V$ (mag)	$t_p$ (d)	$E$ ( $10^{51}$ ) (erg)	$R$ ( $R_\odot$ )	$M_{\text{ej}}$ ( $M_\odot$ )	$M_{\text{ms}}$ ( $M_\odot$ )	$^{56}\text{Ni}$ ( $M_\odot$ )	Ref.
1987A	LMC	0.05	0.60	–	40	1.3	40	15	20	0.075	1
1999em	NGC 1637	11.7 (0.1)	0.31	–16.71	95	$1.2^{+0.6}_{-0.3}$	$249^{+243}_{-150}$	$27^{+14}_{-8}$	$12 \pm 1$	$0.042^{+0.027}_{-0.019}$	2,3,4
1999gi	NGC 3184	11.1	0.65	–15.68	95	$1.5^{+0.7}_{-0.5}$	$81^{+110}_{-51}$	$43^{+24}_{-14}$	$15^{+5}_{-3}$	$0.018^{+0.013}_{-0.009}$	5
2004dj	NGC 2403	3.5 (0.3)	0.22	–15.86	$100 \pm 20$	$0.86^{+0.89}_{-0.49}$	$155^{+150}_{-75}$	$19^{+20}_{-10}$	$>20$	$0.02 \pm 0.01$	6
2004et	NGC 6946	5.4 (1.0)	1.27	–17.04	$110 \pm 10$	$0.98 \pm 0.25$	$530 \pm 280$	$16 \pm 5$	$\sim 20$	$0.06 \pm 0.03$	7,8,9
2005cs	M51	7.1 (1.2)	0.34	–14.83	90–120	0.3	100	8–13	10–15	0.003–0.004	10,11
2007od	UGC 12846	25.3 (0.8)	0.12	–17.64	25	0.5	670	5–7	7–9	0.02	12
2008in	NGC 4303/M61	12.6 (1.0)	0.3	–15.25	98	0.54	126	16.7	$>20$	0.015	13
2009bw	UGC 2890	20.2 (0.6)	0.96	–17.24	100	0.3	510–1000	8.3–12	10–14	0.022	14
2009ib	NGC 1559	19.5 (2.8)	0.496	–15.77	118	0.55	400	15	16.5–17	$0.046 \pm 0.015$	15
2009N	NGC 4487	19.8 (1.1)	0.403	–15.52	110	0.48	287	11.5	13–13.5	$0.020 \pm 0.004$	16
2012aw	NGC 3351	9.9 (0.1)	0.23	–16.67	$\sim 90$	1–2	337	9–14	11–16	$0.06 \pm 0.01$	17
2013ab	NGC 5669	23.9 (0.9)	0.14	–16.69	80	0.35	$\sim 600$	$\sim 7$	$\sim 9$	0.064	18
2014cx	NGC 337	18.5 (3.7)	0.3	–16.56	$\sim 100$	0.4	640	8	10	$0.056 \pm 0.008$	19
DLT16am	NGC 1532	20.3 (4.1)	6.04 (0.46)	–17.76	$\sim 80$	–	–	–	–	$0.208 \pm 0.044$	20
2016X	UGC 08041	15.6 (1.3)	0.05 (0.21)	–16.36	$\sim 90$	–	$930 \pm 70$	–	18.5–19.7	$0.034 \pm 0.006$	21
2015ba	IC 1029	34.8 (0.7)	1.42 (0.16)	–17.00	$\sim 123$	1.6	690	24	–	$0.032 \pm 0.006$	This paper

Notes. <sup>a</sup> In the  $H_0 = 73.24 \text{ km s}^{-1} \text{ Mpc}^{-1}$  scale.

References: (1) Hamuy & Suntzeff (1990), (2) Hamuy et al. (2001), (3) Leonard et al. (2002a), (4) Leonard et al. (2003), (5) Leonard et al. (2002b), (6) Vinkó et al. (2006), (7) Sahu et al. (2006), (8) Misra et al. (2007), (9) Maguire et al. (2010), (10) Pastorello et al. (2006), (11) Pastorello et al. (2009), (12) Inserra et al. (2011), (13) Roy et al. (2011), (14) Inserra et al. (2012), (15) Takáts et al. (2015), (16) Takáts et al. (2014), (17) Bose et al. (2013); Dall’Ora et al. (2014), (18) Bose et al. (2015a), (19) Huang et al. (2016), (20) Tartaglia et al. (2018), (21) Huang et al. (2018).

**Table 7.** Best-fitting values of  $t_{\text{PT}}$ ,  $a_0$ , and  $w_0$  using (Valenti et al. 2016) expression.

SN	$t_0$ JD (2400000+)	$t_0 - t_{\text{PT}}$ (d)	$a_0$ (mag)	$w_0$	References
SN 1999em	51477.0 (1.0)	$116.8 \pm 1.8$	$1.84 \pm 0.03$	$4.1 \pm 0.6$	1,2
SN 1999gi	51518.3 (3.0)	$126.5 \pm 0.6$	$2.2 \pm 0.1$	$5.1 \pm 0.6$	3
SN 2004et	53270.5 (0.2)	$122.9 \pm 0.6$	$1.55 \pm 0.03$	$3.6 \pm 0.8$	4,5,6
SN 2005cs	53549.0 (0.5)	$125.4 \pm 0.2$	$4.07 \pm 0.05$	$2.2 \pm 0.2$	7
SN 2008in	54825.6 (0.8)	$107.8 \pm 0.3$	$1.93 \pm 0.03$	$2.2 \pm 0.3$	10
SN 2009N	54848.1 (1.2)	$108.1 \pm 0.2$	$1.87 \pm 0.03$	$2.6 \pm 0.19$	11
SN 2009ib	55041.8 (3.1)	$139.7 \pm 0.3$	$1.61 \pm 0.04$	$5.4 \pm 0.3$	12
SN 2013ab	56340.0 (1.0)	$102.1 \pm 0.2$	$1.52 \pm 0.03$	$3.6 \pm 0.2$	14
SN 2014cx	56902.4 (0.5)	$109.5 \pm 0.4$	$1.67 \pm 0.04$	$6.0 \pm 0.4$	15
SN 2016X	57406.4 (0.6)	$95.0 \pm 0.3$	$1.33 \pm 0.04$	$2.28 \pm 0.33$	16
SN 2015ba	57349.7 (1.0)	$140.7 \pm 0.2$	$3.01 \pm 0.09$	$2.3 \pm 0.2$	This paper

Notes. References: (1) Hamuy et al. (2001), (2) Leonard et al. (2002a), (3) Leonard et al. (2002b), (4) Sahu et al. (2006), (5) Misra et al. (2007), (6) Maguire et al. (2010), (7) Pastorello et al. (2006), (8) Pastorello et al. (2009), (9) Inserra et al. (2011), (10) Roy et al. (2011), (11) Takáts et al. (2014), (12) Takáts et al. (2015), (13) Gandhi et al. (2013), (14) Bose et al. (2015a), (15) Huang et al. (2016), (16) Huang et al. (2018).

for SN 2005cs ( $4.07 \pm 0.05$  mag) and minimum for SN 2016X ( $1.33 \pm 0.04$  mag).

We reproduced the plot between absolute peak  $V$  magnitude and  $a_0$  from Valenti et al. (2016) along with two additional SNe IIP (SNe 2016X and 2015ba; marked in red) from our sample (Fig. 4). The three objects in Valenti et al. (2016) sample that shows high  $a_0$  are SNe 2005cs, 2013bu, and ASASSN-14ha, which are also low-luminosity SNe IIP. As noted by Valenti et al. (2016), the two parameters ( $a_0$  and  $V_{\text{magnitude}_{\text{max}}}$ ) seem to correlate; however, the correlation is largely affected by the significant drop of the three faint SNe and the small drop of SN 1979C.

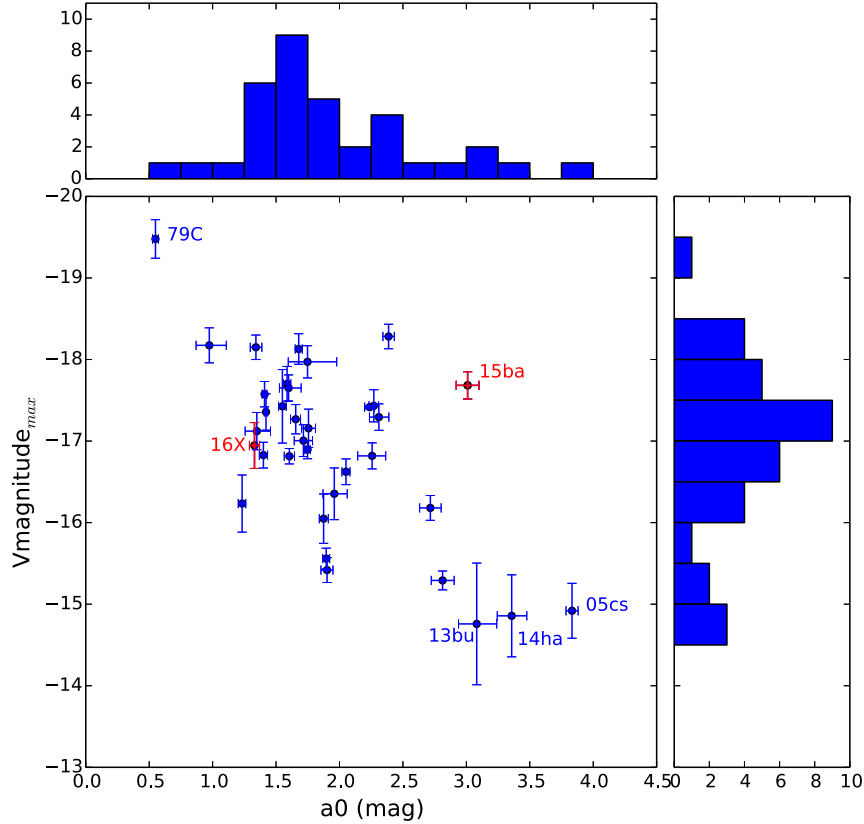
## 4.2 Colour curves

The temporal evolution of reddening-corrected broadband colours provide important clues to the dynamics of SN ejecta. The reddening-corrected colour evolution  $[(B - V)_0, (V - R)_0]$ , and

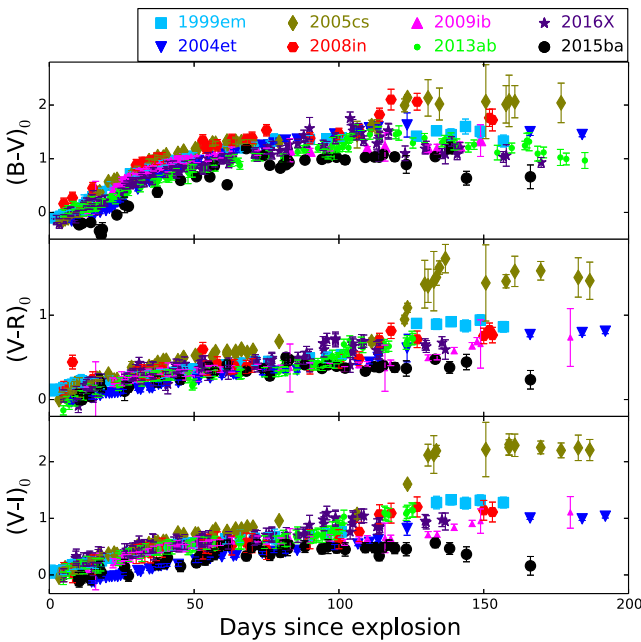
$[(V - I)_0]$  of SN 2015ba is shown in Fig. 5, with those of other well-studied Type IIP SNe. The  $(B - V)_0$  colour gradually becomes redder by about  $\sim 1$  mag in the first  $\sim 50$  d, as a consequence of the expansion and cooling of the ejecta. The  $(V - R)_0$  and  $(V - I)_0$  colours evolve more slowly, with  $<0.5$  mag in 50 d. In the plateau phase, the colours show negligible evolution with time, signifying a nearly constant temperature during this phase. After  $\sim 144$  d, when the SN enters the nebular phase, the colours become bluer. The overall evolution is similar to typical SNe IIP with the colours being marginally bluer than those of the comparison SNe IIP.

## 5 DISTANCE

There are several redshift independent distance estimates listed in NED for the host galaxy IC 1029. However, the values are inconsistent and vary over a large range (31.1–105 Mpc). For this reason, we carry out an independent estimate of the distance to SN 2015ba



**Figure 4.** Plot of absolute peak magnitude in V band and the depth of drop from plateau to nebular phase ( $a_0$ ). The sample of SNe from Valenti et al. (2016) are marked in blue, while the newly added SNe are marked in red.



**Figure 5.** Comparisons of the  $(B - V)_0$ ,  $(V - R)_0$ , and  $(V - I)_0$  colour curves of SN 2015ba with other SNe IIP.

using the EPM and SCM prescriptions. The EPM is a variant of the Baade–Wesslink method to estimate SN distances (Kirshner & Kwan 1974). It compares the linear and angular radius of homogeneously expanding optically thick SN ejecta to compute the SN

distance. The SCM (Hamuy & Pinto 2002) is based on the correlation between the SN brightness and the expansion velocity in the middle of the plateau. It needs less input data, but requires calibration via SNe with well-established distances. We discuss these methods in details below.

### 5.1 Expanding photosphere method

The early phase of an SN, when the ejecta is fully ionized and electron scattering dominates the total opacity at the photosphere, can be approximated to be radiating as a diluted blackbody. The angular radius of the expanding ejecta at any time  $t$  can then be approximated as

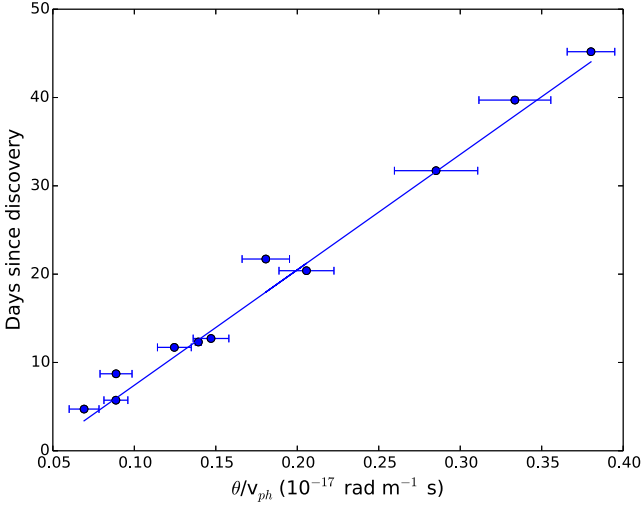
$$\theta = \frac{R}{D} = \sqrt{\frac{f_\lambda 10^{0.4A_\lambda}}{\zeta_\lambda^2(T_c) \pi B_\lambda(T_c)}} \quad (4)$$

where  $B_\lambda$  is the Planck function at colour temperature  $T_c$ ,  $f_\lambda$  is the flux density received at Earth,  $A_\lambda$  is the extinction, and  $\zeta_\lambda(T_c)$  is the colour temperature dependent ‘dilution factor’ and  $R = v(t - t_0)$ , where  $(t - t_0)$  is the time since explosion.

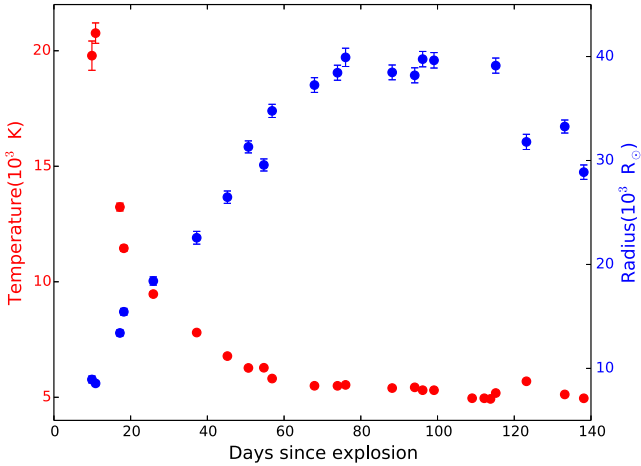
Recasting this equation in terms of broadband photometry by integrating over the filter response function, we get

$$m_\lambda = -5 \log(\zeta_\lambda \theta) + A_\lambda + b_\lambda, \quad (5)$$

where  $b_\lambda$  is the convolution of the filter response function with the blackbody model  $\pi B_\lambda$ .



**Figure 6.** Distance determination with the EPM for SN 2015ba, using Dessart & Hillier (2005b) prescription for the dilution factors and the *BVI* observations. The distance and the explosion epoch are obtained from the best fit.



**Figure 7.** Evolution of the temperature and the radius of SN 2015ba.

Hamuy et al. (2001) computed  $b_\lambda(T)$  by fitting a polynomial in the temperature range of 4000–25 000 K, as

$$b_\lambda(T) = \sum_{i=1}^5 C_i(\lambda) \left( \frac{10^4 \text{K}}{T} \right)^i. \quad (6)$$

The coefficients  $C_i$ 's are listed in table 13 of Hamuy et al. (2001).

Minimizing the quantity,

$$\epsilon = \sum_{\lambda \in S} [m_\lambda + 5 \log(\zeta_\lambda \theta) - A_\lambda - b_\lambda]^2 \quad (7)$$

we get  $\zeta_\lambda \theta$  and  $T_c$  simultaneously. The errors in  $\zeta_\lambda \theta$  and  $T_c$  are estimated by randomly mixing the magnitudes with uncertainties drawn from their normal distribution, thereby generating a sample of 1000 data points. The standard deviation of the result gives the error for these quantities.

The dilution factors  $\zeta$  can be expressed as a function of  $T_c$  from atmospheric model calculations. Dessart & Hillier (2005b) used second-order polynomial fits to the distribution of dilution factors

versus blackbody colour temperatures,

$$\zeta_S(T_S) = \sum_{i=0}^2 a_{S,i} \left( \frac{10^4 \text{K}}{T_S} \right)^i. \quad (8)$$

We estimated  $\theta$  adopting the coefficients from Dessart & Hillier (2005b) and the values obtained are listed in Table 8. The photospheric velocity  $v_{\text{ph}}$  is obtained from the spectral modelling of He I 5876Å line up to 13 d after discovery and Fe II 5169Å line up to 40 d after discovery. We generate a sample of 1000 data points from the normal distribution of the uncertainty in the photospheric velocity. To each of this data set, cubic spline interpolations is performed to account for the missing velocity values for the corresponding epoch of photometric data. From the standard deviation of these values, we obtain the errors. We perform a linear fit to  $t$  versus  $\theta/v_{\text{ph}}$  to estimate the distance from the slope and the time of outburst  $t_0$  from the y-intercept, following the expression:

$$t = D(\theta/v_{\text{ph}}) + t_0. \quad (9)$$

The linear fit to the data is shown in Fig. 6 and the distance and explosion epoch estimates from the fit are  $36.6 \pm 1.9$  Mpc and  $5.6 \pm 1.0$  d before the discovery date (2015 November 28.8 UT, JD 2457355.3), respectively.

## 5.2 Standard candle method

The SCM has several versions (Nugent et al. 2006; Poznanski et al. 2009; Olivares et al. 2010) with different SNe samples, but, yielding consistent results. We will discuss two versions and use their average value as the SCM distance estimate. The value of Hubble constant adopted below is  $H_0 = 73.24 \pm 1.74$  km s<sup>-1</sup> Mpc<sup>-1</sup> (Riess et al. 2016).

(i) In the version of Poznanski et al. (2009), the measured brightness and expansion velocity of a sample of 34 SNe with known distances were used to calibrate the equation

$$\begin{aligned} \mathcal{M}_I - \alpha \cdot \log \left( \frac{v_{\text{Fe}}(50 \text{d})}{5000} \right) + R_I((V - I) - (V - I)_0) - m_I \\ = -5 \cdot \log(H_0 D), \end{aligned} \quad (10)$$

where  $\mathcal{M}_I = -1.615 \pm 0.08$  mag,  $\alpha = 4.4 \pm 0.6$ ,  $R_I = 0.8 \pm 0.3$ , and  $(V - I)_0 = 0.53$  mag. Note that the  $R_I$  value used here differs from the Cardelli et al. (1989) value ( $R_I = 1.36$ ). We measure the *I*-band magnitude of SN 2015ba on day +50 as  $m_I = 16.07 \pm 0.04$  mag,  $(V - I) = 0.954 \pm 0.01$  mag, and  $V_{\text{FeII}} = 3332 \pm 55$  km s<sup>-1</sup> and obtain a distance  $D = 28.8 \pm 2.6$  Mpc.

(ii) Olivares et al. (2010) used the magnitudes and velocities measured 30 d before the middle of the transition phase ( $t_{\text{PT}}$ ) in the relation of Hamuy & Pinto (2002)

$$m + \alpha \cdot \log \left( \frac{v_{\text{Fe}}}{5000} \right) - \beta(V - I) = 5 \cdot \log(H_0 D) + z_p \quad (11)$$

and calibrated the formula using a sample of 37 nearby SNe IIP. The coefficients  $\alpha$ ,  $\beta$ , and  $z_p$  are different for *V* and *I* bands. We determine  $t_{\text{PT}} - t_0 = 140.7$  d after explosion (see Section 4.1). 30 d before this date, we measure the values  $v(110.7 \text{d}) = 2252 \pm 54$  km s<sup>-1</sup>,  $m_V = 17.31 \pm 0.02$  mag and  $m_I = 16.16 \pm 0.02$  mag. We estimate the distances to be  $29.9 \pm 2.5$  Mpc and  $31.2 \pm 2.3$  Mpc using equation (11) for *V* and *I* bands, respectively. However, note that no K-correction has been applied in this case.

The SCM distance estimates using the above two methods are consistent with each other. The weighted SCM distance is  $30.1 \pm 1.4$  Mpc. We list the distance estimates obtained via recessional veloc-

**Table 8.** The derived parameters of SN 2015ba needed to apply EPM: the angular size ( $\theta$ ), photospheric temperature ( $T$ ), dilution factor [ $\zeta_\lambda(T)$ ], and the interpolated photospheric velocity ( $v_{\text{ph}}$ ).

$t^a$ (d)	$\theta$ ( $10^8$ km Mpc $^{-1}$ )	$T$ (K)	$\zeta_\lambda(T)$	$v_{\text{ph}}$ (km s $^{-1}$ )
4.71	1.870 (0.247)	23019 (2794)	0.519 (0.007)	8757 (56)
5.72	2.330 (0.192)	18766 (1350)	0.509 (0.003)	8512 (47)
8.72	2.149 (0.237)	19641 (1876)	0.511 (0.005)	7841 (38)
11.70	2.787 (0.232)	15115 (1906)	0.503 (0.001)	7250 (43)
12.33	3.069 (0.037)	13404 (169)	0.504 (0.0004)	7138 (44)
12.72	3.208 (0.239)	12840 (1232)	0.506 (0.004)	7067 (45)
20.42	3.750 (0.306)	10367 (703)	0.527 (0.01)	5909 (52)
21.71	3.204 (0.256)	11671 (786)	0.512 (0.006)	5746 (53)
31.72	4.181 (0.368)	8550 (434)	0.572 (0.02)	4750 (82)
39.72	4.248 (0.258)	7648 (252)	0.616 (0.02)	4127 (108)
45.20	4.346 (0.063)	7238 (508)	0.645 (0.004)	3701 (132)

Note. <sup>a</sup>Since discovery,  $t_0 = 2457355.3$  JD.

**Table 9.** The line velocities of Fe II  $\lambda 5169$  and He I  $\lambda 5876$  estimated with SYN++ modelling of the observed spectra of SN 2015ba.

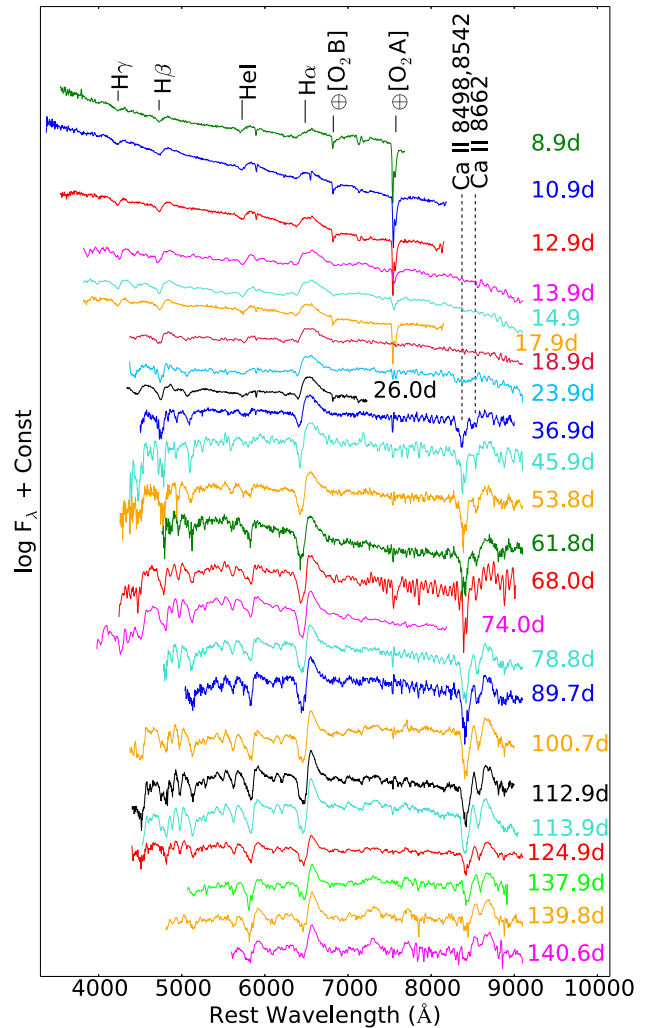
Date (yyyy-mm-dd)	Phase <sup>a</sup> (days)	$v(\text{He I})$ $10^3$ km s $^{-1}$	$v(\text{Fe II})$ $10^3$ km s $^{-1}$
2015-12-02	8.9	$9.2 \pm 0.1$	–
2015-12-04	10.9	$8.4 \pm 0.1$	–
2015-12-06	12.9	$8.4 \pm 0.1$	–
2015-12-07	13.9	$7.7 \pm 0.2$	–
2015-12-08	14.9	$7.8 \pm 0.1$	–
2015-12-11	17.9	$7.2 \pm 0.1$	–
2015-12-12	18.9	$6.3 \pm 0.1$	–
2015-12-17	23.9	–	$6.7 \pm 0.1$
2015-12-19	26.0	–	$6.0 \pm 0.1$
2015-12-30	36.9	–	$4.5 \pm 0.1$
2016-01-08	45.9	–	$4.5 \pm 0.2$
2016-01-16	53.8	–	$3.6 \pm 0.3$
2016-01-24	61.8	–	$2.4 \pm 0.4$
2016-01-29	68.0	–	$2.6 \pm 0.2$

Note. <sup>a</sup>Relative to the date of explosion,  $t_0 = 2457349.7$  JD.

ity, EPM, and SCM in Table 2. The weighted mean of these distances is  $34.8 \pm 0.7$  Mpc and has been adopted throughout the paper.

## 6 TEMPERATURE AND RADIUS

We also determine the photospheric temperature and the radius evolution by constructing SEDs with the *BVI* photometric fluxes at each epoch and fitting it with a Planck function until  $\sim 140$  d. We have also applied the dilution factor correction (Dessart & Hillier 2005b) corresponding to the *BVI* filter set to the blackbody radii to estimate the photospheric radii. As shown in Fig. 7, the temperature drops from 20 000 K at 10 d to 6300 K at 50 d since explosion due to rapid adiabatic cooling, and thereafter it declines very slowly to 5300 K at 100 d since explosion. The slow decline of the late phase is due to the low opacity of the ejecta that favours cooling through photon energy diffusion. The photospheric radius shows a constant growth up to  $\sim 70$  d and then remains nearly constant until 115 d. During the plateau phase, due to the declining electron density, the photosphere remains contiguous to the receding recombination front on top of the expanding envelope, leaving the photospheric radius nearly constant. Afterwards, the radius starts to decrease up to 140 d, beyond which the SED cannot be compared to a blackbody any longer.



**Figure 8.** Doppler-corrected spectra of SN 2015ba from 8.9 to 140.6 d since explosion. The symbol  $\oplus$  indicates the position of the most prominent telluric absorption bands.

## 7 SPECTRAL ANALYSIS

Fig. 8 presents the spectral evolution of SN 2015ba from  $\sim 9$  d to  $\sim 141$  d after explosion, with a preliminary identification of absorp-

tion features. This is necessary to get an idea of the overall structure of the envelope with the revelation of deeper layers at later times. Prominent features of H are visible throughout the evolution. We attempt a detailed line identification at two phases, 8.9 and 53.8 d spectrum using the `SYN++` (Thomas, Nugent & Meza 2011) spectral code.

`SYN++` is an evolution of `SYNOW`, which uses the Sobolev approximation to produce synthetic spectra of SNe during the photospheric phase. `SYN++` assumes that spectral lines are formed via resonance scattering above a sharp photosphere and the ejecta are homologously expanding. The location of the photosphere is expressed in velocity coordinates as  $v_{\text{ph}}$  (in  $\text{km s}^{-1}$ ). With the optical depths for each species, line strengths are computed assuming Boltzmann excitation (i.e. local thermodynamic equilibrium – LTE) using a specified excitation temperature  $T_{\text{exc}}$  (in K). Non-LTE effects are partially accounted for by allowing different  $T_{\text{exc}}$  values for each species, which can be different from the photospheric temperature  $T_{\text{phot}}$ . The latter is used only for computing the blackbody radiation emitted by the photosphere.

The early spectrum mostly exhibits broad H and He I lines superposed on a blue continuum, which are fairly well reproduced in the synthetic spectrum. There is a weak but conspicuous absorption feature at  $\sim 4560 \text{ \AA}$ , marked by ‘?’ and identified from the modelling as He II 4686. With time, multiple metal lines gain prominence, resulting in line blendings. The `SYN++` modelling aids in identifying most of the blended features in the plateau phase spectrum at 53.8 d (Fig. 9). The P-Cygni profiles to the blue side of H $\alpha$  profile are identified as Sc II and Ba II  $\lambda 6142$ . However, there is a discrepancy between the observed and the modelled spectra at shorter wavelengths, which could perhaps improve with the addition of other species such as Ti II lines. Although, the `SYN++` fit reproduces most of the features, it fails to reproduce the emission component of H $\alpha$  line. The net emission from H $\alpha$  line is a characteristic feature of SN II spectra and is not reproducible in such simplified models without assuming an ad hoc net line emission (as reported in Jeffery & Branch 1990).

### 7.1 Early phase features

The early spectra (8.9–12.9 d) are characterized by a few broad P-Cygni profiles superimposed on a blue continuum. The most prominent are the H and He I  $\lambda 5876$  lines. H $\alpha$  has a weaker absorption component as compared to other H I lines. The 8.9 d spectrum modelled using `SYN++` suggests the absorption feature at  $\sim 4560 \text{ \AA}$  arises from the excitation of He. We compared the 8.9 d spectra of SN 2015ba with coeval epoch spectra of SNe 1999em, 2007od, 2012aw, and 2013ab in Fig. 10. The P-Cygni profiles are quite well developed in the spectra of SNe 1999em, 2007od, and 2012aw, while those of SNe 2013ab and 2015ba have a more featureless continuum. This is possibly due to temperature difference in the initial phases among the SNe. The spectrum from 12.9 to 18.9 d marks the transition from the early to the plateau phase. The He I feature completely disappears after 18.9 d, and later on metal lines start to develop.

### 7.2 Plateau phase

As the SN expands and cools, the opacity drops and the continuum becomes redder. Metal lines develop, the emission component of H $\alpha$  becomes stronger, while the other H I Balmer lines attenuate. In place of He I 5876, Na I D line develops and become stronger with

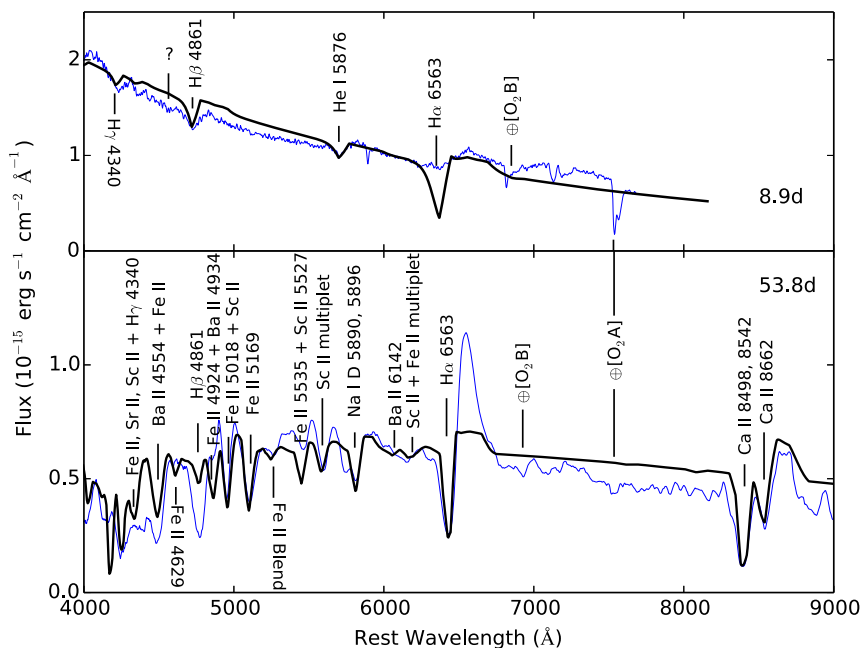
time. The spectra evolves very slowly during the plateau (36.9–124.9 d, Fig. 8). Fe II  $\lambda 5169$  is feebly visible in the 18.9 d spectrum and becomes prominent after 23.9 d. Other lines of Fe II (Fe II  $\lambda\lambda$  5018, 4924) appear in the 23.9 d spectrum, but are weak. These lines become prominent from the 36.9 d spectrum. Ba II lines appear in the 26 d spectrum. The Ca II IR triplet ( $\lambda\lambda$  8498, 8542, 8662) becomes visible from day 23.9. Thereafter, a number of P-Cygni profiles develop in the spectra. We compare the mid-plateau phase spectrum of SN 2015ba with those of other SNe IIP in Fig. 11. The spectra of SN 2015ba shows remarkable similarity in features with other SNe IIP, only the strength of the Ba II, Sc II, and Fe II lines are weaker in SNe 2015ba, 2013ab, and 2009bw.

### 7.3 Nebular phase

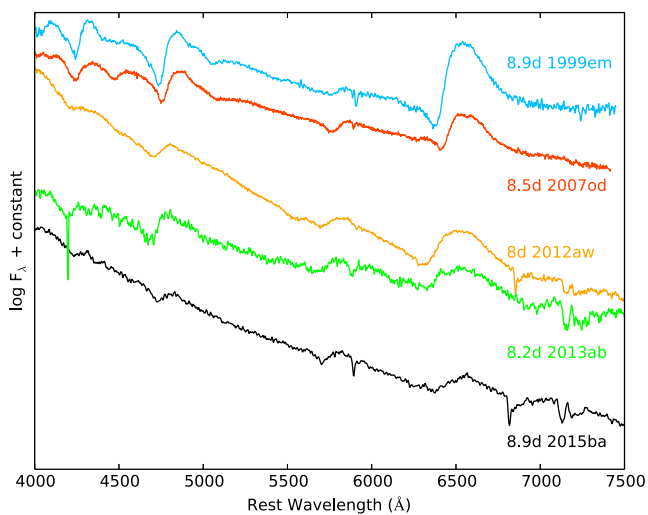
In the nebular phase, owing to the optically thin ejecta, the absorption component of the P-Cygni profiles vanishes. The 253.2 and 272.2 d spectra obtained with the 10.4 m GTC, shown in Fig. 12, mainly exhibit H, Na I D, and Ca II infrared (IR) triplet features along with some nebular lines. The H $\alpha$  line is the most prominent, while the other H Balmer lines, though detectable, are much weaker. The Balmer lines exhibit double-peaked profiles in the 253.2 d spectrum, with the red peak at the rest wavelength. The double peaked feature in H $\alpha$  may possibly be attributed to noise in the spectra, but a similar feature is detected for H $\beta$ . This supports that it is an intrinsic feature. Jerkstrand et al. (2015) suggested that double peaked profiles may arise from dust in the circumstellar medium (CSM) at very late epochs ( $\gtrsim 400$  d). However, at the epochs ranging from 200 to 300 d, Jerkstrand et al. (2015) state that dust will have negligible effect on the optical spectra of SNe IIP. We, therefore, disfavour an effect of dust, and suggest this to be an artefact due to background line contamination. The appearance of [O III] 5007 and [S II] 6716, 6731 emission lines from the host galaxy are clear signatures that the background is contaminating the SN spectra at late phases. Forbidden lines of [Ca II] 7291, 7324 and the K I  $\lambda\lambda$  7665, 7699 also becomes prominent in the nebular spectra. The plethora of permitted and forbidden emission lines of iron from various multiplets generates several blends. The comparison of the nebular phase spectrum of SN 2015ba at 272.2 d with late-time spectra of other SNe IIP is shown in Fig. 13. Most of the features in the nebular spectra in the comparison SNe IIP are similar; however, the forbidden [O I] 6300, 6364 doublet is conspicuous in the comparison spectra and is absent in SNe 2015ba and 2009bw. Also, the absorption components in SN 2015ba is weaker with respect to the comparison sample.

### 7.4 Evolution of the line profiles

The evolution of H $\alpha$ , H $\beta$ , Na I D, and Ba II  $\lambda 6142$  lines from 14.9 to 140.6 d is shown in Fig. 14 in the velocity domain and centered at the rest wavelengths of these lines. The H $\alpha$  feature is broad and weak at early phases and becomes stronger and narrower at late phases. The profiles of H $\beta$ , Na I D, and Ba II  $\lambda 6142$  are hardly discernible at early phases and only gain prominence at later phases. The He I 5876 line is visible at early phases at the position of Na I D line, and gradually attenuates while the Na I D line emerges. A blueshifted emission profile is seen for H $\alpha$  as well as H $\beta$  line, which gradually shifts to redder wavelengths, reaching zero velocity at the end of plateau phase. As pointed out by Chugai (1988) and later by Jeffery & Branch (1990), the blueshifted emission peak results from the diffuse reflection by the photosphere of the resonance photons emitted in the SN atmosphere that are not absorbed because of



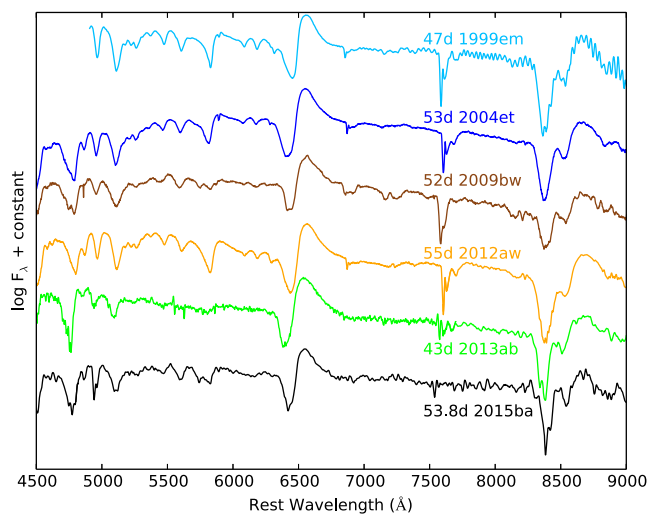
**Figure 9.** SYN++ modelling of the 8.9 and 53.8 d spectra of SN 2015ba. Model spectra are shown with thick solid (black) lines, while the observed ones are in thin (blue) solid line.



**Figure 10.** Comparison of an early spectrum (+8.9 d) of SN 2015ba with other archetypal Type IIP SNe at similar phases.

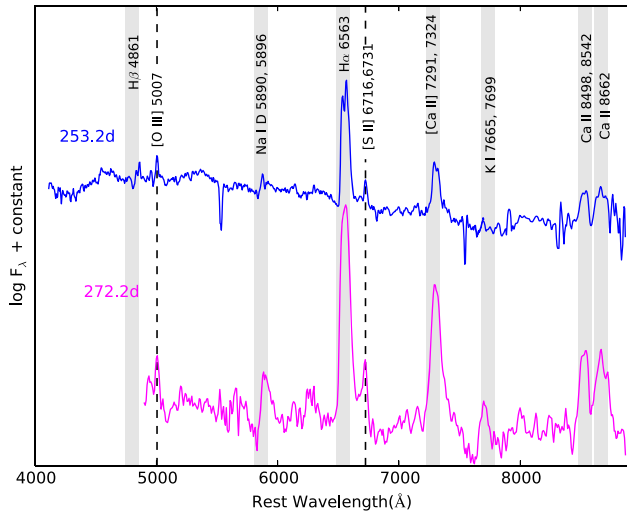
the dominant electron scattering opacity at early phases. However, Anderson et al. (2014a) carried out an analysis of a sample of SNe IIP and linked the blueshift to the steep density structure of the ejecta, which gives rise to strong occultation effects in the receding part of the ejecta, shifting the emission peak towards the blue. Such blueshifted emission lines have been successfully modelled by Dessart & Hillier (2005a) using the non-LTE model CMFGEN, taking into account the steep density profile of ejecta. Since  $H\alpha$  is the strongest emission line in SN II spectra, the velocity offset is more conspicuous in this line as compared to other emission lines, although the peaks of other emission lines are also blueshifted by a smaller amount.

We also observe a non-evolving, narrow absorption component to the blue side of  $H\alpha$  lines from the 89.7 to 140.6 d spectra of SN



**Figure 11.** Comparison of a mid-plateau spectrum (53.8 d) of SN 2015ba with other archetypal SNe IIP at similar phases.

2015ba (marked with dash-dotted lines in Fig. 14). This component is most conspicuous in  $H\alpha$  trough at  $\sim 6420$  Å and is detectable in the blue wings of  $H\beta$  at  $\sim 4750$  Å in the 124.9 and 137.9 d spectra. This component has been observed at the mid-plateau epochs in several SNe II (Baron et al. 2000; Leonard et al. 2002a; Chugai et al. 2007; Inserra et al. 2012; Valenti et al. 2014; Bose et al. 2015b; Gutiérrez et al. 2017) and has been attributed to either Si II 6355 Å or HV hydrogen lines as a result of weak interaction between the SN ejecta and the RSG wind. However, this feature does not seem to evolve with time in SN 2015ba like other species in the shell, hence it is not arising from Si II. Chugai et al. (2007) suggested that the blue shoulder in the absorption profile of  $H\alpha$  lines originate from the excitation of the cool, dense shell formed at the interface of the SN ejecta and CSM. Considering this feature to be HV components of



**Figure 12.** Doppler-corrected nebular phase spectra of SN 2015ba. The emission lines from the galaxy are shown with dashed lines.

H $\alpha$  and H $\beta$ , both of these are observed at about 6800 km s<sup>-1</sup> from the rest wavelength of these lines, which supports the identification of these components as HV Balmer lines. We are also able to model this feature in the 137.9 d spectrum with a HV component of H I in the SYN++ modelling as shown in Fig. 15, which further strengthens our assumption.

### 7.5 Velocity evolution

The photospheric velocity of the ejecta traces the velocity of propagation of the recombination wave that in turn depends on the physical conditions of the ejecta such as the temperature, density, and extent of mixing of elements in the ejecta. Dessart & Hillier (2005b) described the photosphere as the layer where the continuum optical depth is  $\sim 2/3$ . However, there is no single prescribed line that can accurately represent the true photospheric layer and its velocity. We used Fe II 5169 line as a proxy for the photospheric velocity in the plateau phase (Dessart & Hillier 2005b) while at early phases, He I 5876 line is used as suggested in Takáts & Vinkó (2012). The velocity is determined from the shift in the absorption minima of these lines or through spectral modelling. The velocities estimated from spectral modelling is listed in Table 9. We also compute the radial velocity from the absorption minima of H $\alpha$ , H $\beta$  and Sc II lines (see Fig. 16). These lines form in the outer layers of the envelope, exhibiting higher expansion velocities than the Fe II lines. However, we also note the drop in velocity of H $\beta$  at later phases, indicative of the presence of hydrogen in the deeper ejecta layers.

The photospheric velocity evolution of SN 2015ba is compared to other SNe IIP in Fig. 17. The velocities derived from the absorption minima of He I/Fe II line have been used for all SNe to maintain consistency. While the velocity profile of SN 2015ba is similar to other SNe IIP, the velocities are significantly higher than SNe 1987A and 2005cs. The velocities of SN 2015ba are, however, more similar to SNe 1999em and 1999gi and are about  $\sim 500$  km s<sup>-1</sup> lower than SNe 2004et, 2012aw, and 2013ab.

## 8 EXPLOSION PROPERTIES

### 8.1 Bolometric light curve

The bolometric light curve, with contributions ranging from the ultraviolet (UV) to the IR domains, is a useful tool to derive the

physical parameters of the explosion such as the synthesized <sup>56</sup>Ni mass, and the mass and kinetic energy of the ejecta. However, a well-sampled data set over all wavelengths is usually rare. Therefore, we construct a pseudo-bolometric light curve by integrating the fluxes in *BVri* bands only.

We convert the dereddened *BVri* magnitudes of SN 2015ba to fluxes using the zero points adopted from Bessell, Castelli & Plez (1998) and Fukugita et al. (1996). We integrate over the filter bandpass using these monochromatic fluxes implementing trapezoidal integration rule, and for each epoch we obtain the pseudo-bolometric flux. Adopting a distance of  $34.8 \pm 0.7$  Mpc, we compute the pseudo-bolometric luminosity ( $L_{BVri}$ ).

The pseudo-bolometric light curve of SN 2015ba is compared with those of the SNe sample in Fig. 18. We generate the pseudo-bolometric light curves of all comparison SNe using the method described above for consistency. The luminosity of SN 2015ba in the plateau phase is comparable with that of the luminous SN 2004et. However, the constant luminosity phase is extended for SN 2015ba, possibly indicating a massive hydrogen envelope. Due to limited multiband data in the radioactive tail phase, we are unable to estimate the pseudo-bolometric luminosities for SN 2015ba at late phases.

We construct the true bolometric light curve, using the direct integration module in SuperBoL (Lusk & Baron 2017) that is based on the method given in Bersten & Hamuy (2009). With this method, the IR contribution is obtained by fitting a blackbody spectrum to the observed fluxes redward of 5000 Å (as suggested in Lyman, Bersier & James 2014) and then integrating the best-fitting blackbody curve from the longest observed wavelength to  $\lambda = \infty$ . The missing UV flux is obtained by integrating the blackbody function from the shortest observed flux to  $f_{\lambda} = 0$  at  $\lambda = 2000$  Å. We cannot adopt this method to estimate the UV contribution to the bolometric light curve in the nebular phase since this method is only valid during the optically thick phases (up to  $\sim 115$  d in SN 2015ba) when the SED of the SN can be approximated to a blackbody. So, we estimate the UV contribution at late phases with a linear function from the effective wavelength of the shortest observed filter to  $f_{\lambda} = 0$  at  $\lambda = 2000$  Å. The UV (2000–3500 Å) and IR contribution (8200–21 900 Å) in SN 2015ba, thus obtained, is  $\sim 36$  per cent and  $\sim 10$  per cent, respectively, on day 11. The UV contribution falls to  $\sim 10$  per cent and the IR contribution rises to 40 per cent by day 120.

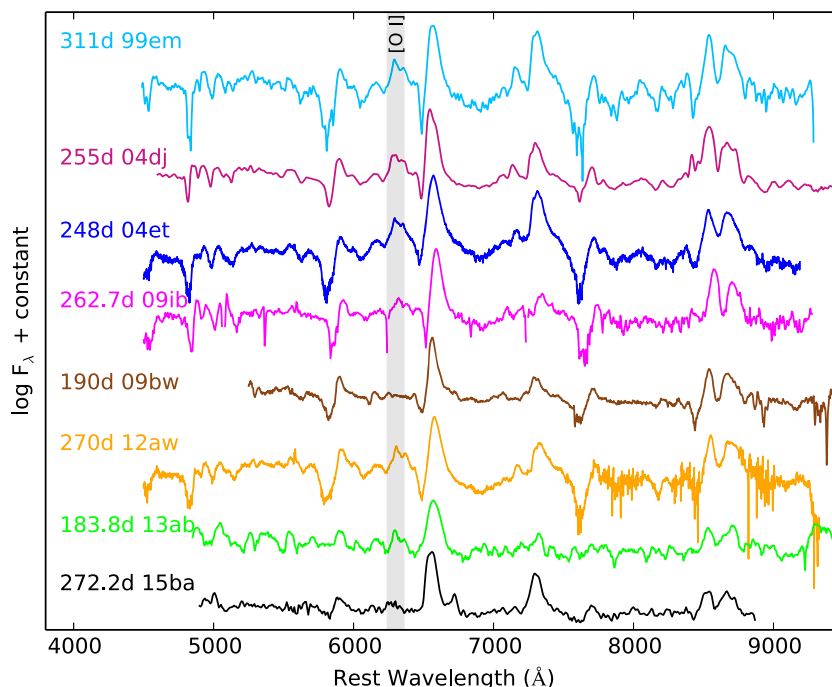
The contributions in the UV and IR are added to the pseudo-bolometric luminosities in order to obtain the true bolometric luminosities. The true bolometric luminosities are used in the analytical and hydrodynamical modelling discussed later in Sections 8.3 and 8.4, respectively.

### 8.2 <sup>56</sup>Ni mass

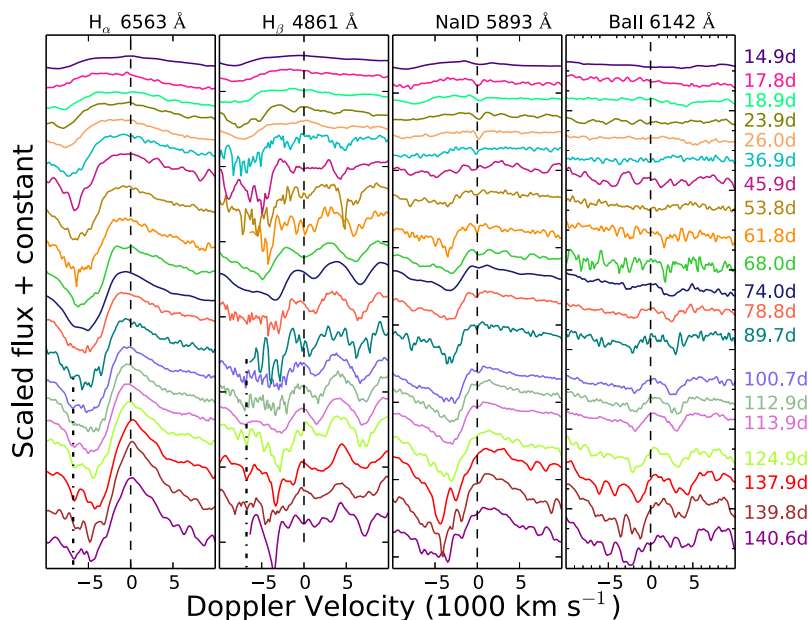
The nebular phase light curve is thought to be powered by the radioactive decay of <sup>56</sup>Co to <sup>56</sup>Fe resulting in the emission of gamma rays and positrons. As the ejecta in the nebular phase are still opaque to gamma rays and the <sup>56</sup>Co is the radioactive decay product of <sup>56</sup>Ni, the luminosity is proportional to the <sup>56</sup>Ni nucleosynthesized at the time of shock breakout. We determine the mass of <sup>56</sup>Ni using the bolometric luminosity by adopting two empirical methods discussed below.

#### 8.2.1 <sup>56</sup>Ni mass from tail luminosity

Assuming that the deposition of  $\gamma$ -photons and positrons emitted from radioactive decay of <sup>56</sup>Co results in the thermalization of



**Figure 13.** Comparison of the nebular spectrum (272.2 d) of SN 2015ba with those of other archetypal Type IIP SNe.



**Figure 14.** Evolution of H  $\alpha$ , H  $\beta$ , Na I D, and Ba II 6142 Å emission line profiles during the photospheric phase. The HV feature conspicuous from the 89.7 d spectrum up to the 140.6 d spectrum is shown with dash-dotted lines.

energy in the ejecta,  $^{56}\text{Ni}$  mass can be independently estimated from the tail luminosity, as described by Hamuy (2003):

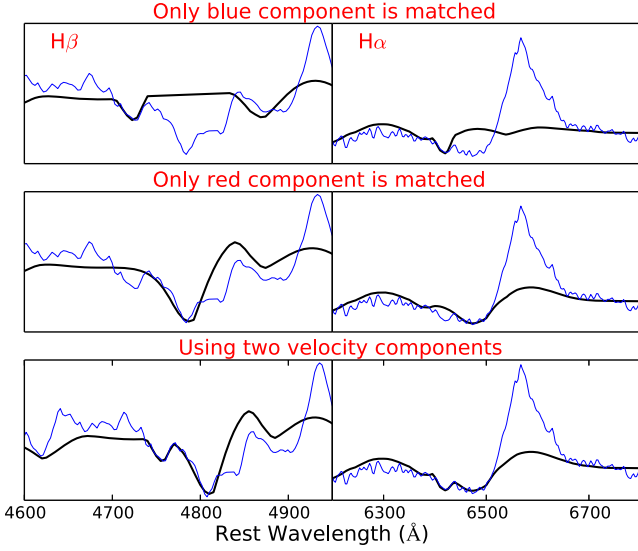
$$M_{\text{Ni}} = 7.866 \times 10^{-44} \times L_t \exp \left[ \frac{(t - t_0)/(1+z) - 6.1}{111.26} \right] M_{\odot}, \quad (12)$$

where  $t_0$  is the explosion time, 6.1 d is the half-life time of  $^{56}\text{Ni}$  and 111.26 d is the e-folding time of the  $^{56}\text{Co}$  decay. The tail luminosity  $L_t$  is computed at the only epoch available at the tail ( $\sim 166$  d) from the V-band magnitude corrected for distance, extinction, and bolometric correction factor of  $0.26 \pm 0.06$  mag during the nebular

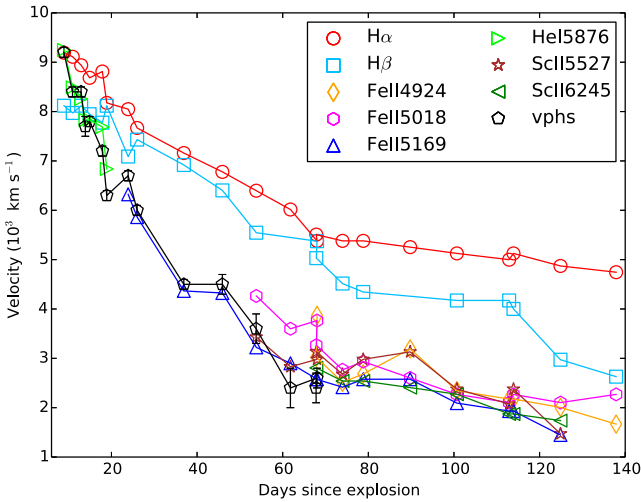
phase (Hamuy 2003) using the following equation:

$$\log(L_t) = -0.4[V_t - A_V(\text{tot}) + \text{BC}] + 2 \log D - 3.256, \quad (13)$$

where  $D$  is the distance in cm and BC is the bolometric correction in the nebular phase. The value of  $L_t$  obtained with equation (13) is  $9.9 \pm 1.8 \times 10^{40}$  erg s $^{-1}$ , which corresponds to  $^{56}\text{Ni}$  mass of  $0.032 \pm 0.006 M_{\odot}$ .



**Figure 15.** SYN++ modelling of the 137.9 d spectrum showing fits to the HV notch in H $\alpha$  (right-hand panel) and H $\beta$  (left-hand panel). A single HV component (6800 km s $^{-1}$ ) is used to match the blue dip (top panel), a single low-velocity component (3100 km s $^{-1}$  for H $\beta$  and 4500 km s $^{-1}$  for H $\alpha$ ) is used to match the red dip (middle panel) and two velocity components are used to fit both the absorption dips (bottom panel).



**Figure 16.** Line velocity evolution of H $\alpha$ , H $\beta$ , He I, Fe II, and Sc II lines. The velocities are estimated from the absorption minima. The SYN++ modelled photospheric velocities ( $v_{\text{phs}}$ ) follow roughly the He I velocities at early phases and Fe II/Sc II velocities at late phases.

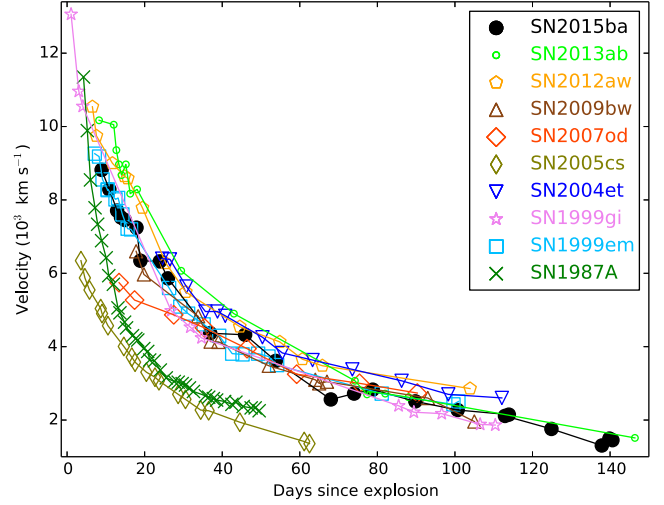
### 8.2.2 $^{56}\text{Ni}$ mass from nebular spectra

We also deduce the  $^{56}\text{Ni}$  mass from full width at half-maximum (FWHM) of the H $\alpha$  emission line in the late-time spectra of SN 2015ba using the relation from Maguire et al. (2012):

$$M(^{56}\text{Ni}) = A \times 10^{B \cdot \text{FWHM}_{\text{corr}}} M_{\odot}, \quad (14)$$

where  $A = 1.81^{+1.05}_{-0.68} \times 10^{-3}$ ,  $B = 0.0233 \pm 0.0041$ , and  $\text{FWHM}_{\text{corr}}$  is the FWHM of line profiles corrected for the spectral resolution of the instruments used to obtain the spectrum given by

$$\text{FWHM}_{\text{corr}} = \sqrt{\text{FWHM}_{\text{obs}}^2 - \text{FWHM}_{\text{instr}}^2}. \quad (15)$$



**Figure 17.** The evolution of photospheric velocity ( $v_{\text{ph}}$ ) of SN 2015ba is compared with those of other well-studied SNe. The  $v_{\text{ph}}$  plotted here are absorption trough velocities (from He I at early phases and Fe II  $\lambda 5169$  at late phases).

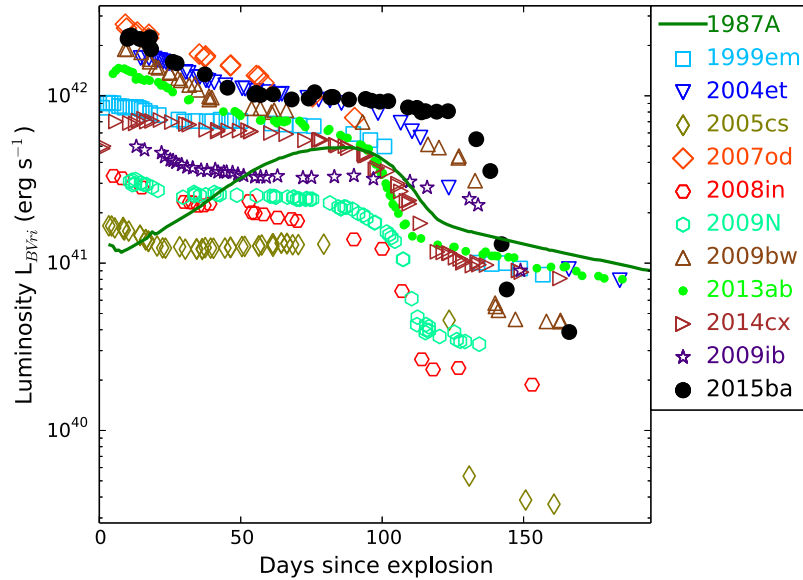
We obtain the FWHM by fitting a Gaussian function to the emission lines in the spectra at 253 and 272 d. For the instrumental FWHM ( $\text{FWHM}_{\text{instr}}$ ), we measured the FWHM of [O I] 6300 Å sky emission line, which was found to be  $18.17 \pm 1.71$  and  $27.94 \pm 1.48$  Å on 253 and 272 d, respectively. The FWHM of H $\alpha$  emission line is  $60.68 \pm 1.83$  Å and  $65.68 \pm 2.16$  Å and using equation (15), we obtain the corrected FWHM to be  $57.89 \pm 2.50$  Å and  $59.44 \pm 2.62$  Å on 253 and 272 d, respectively. Substituting the corrected FWHM in equation (14), we obtain the average  $^{56}\text{Ni}$  mass from the two late-time spectra to be  $0.042^{+0.026}_{-0.020} M_{\odot}$ .

We find that the  $^{56}\text{Ni}$  mass obtained using the above two methods yields consistent results within the errors. The weighted mean of ejected  $^{56}\text{Ni}$  is  $0.032 \pm 0.006 M_{\odot}$ .

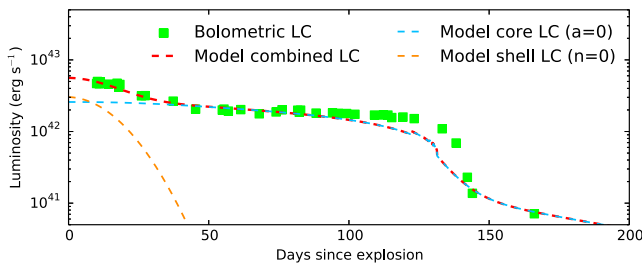
### 8.3 Analytical modelling

The physical parameters of SN 2015ba such as the explosion energy, ejected mass, and initial radius can be estimated using the two-component semi-analytical modelling of Nagy et al. (2014) and Nagy & Vinkó (2016) incorporating a dense inner core and an extended low-mass envelope. This technique uses the radiative diffusion model to include the effect of the recombination front in the ejecta that was originally developed by Arnett & Fu (1989). The main assumptions in the extended model of Nagy et al. (2014) is a homologously expanding spherical ejecta and constant Thomson scattering opacity in a given ionized layer. The photon diffusion equation in the ejecta is then solved considering the recombination energy, the radioactive energy from the decay of  $^{56}\text{Ni}$  and  $^{56}\text{Co}$  and the effect of gamma ray leakage. We fit this model to the true bolometric light curve (see Section 8.1) and derive the best-fitting parameters. One should note, however, that independent values of opacity, ejected mass ( $M_{\text{ej}}$ ), and kinetic energy ( $E_{\text{kin}}$ ) cannot be estimated using this technique; only the degenerate combinations of  $M_{\text{ej}}\kappa$  or  $E_{\text{kin}}\kappa$  can be constrained.

We used a constant density profile for the core and the shell. The best-fitting ejecta mass, progenitor radius, and explosion energy are  $22 M_{\odot}$ ,  $4.8 \times 10^{13}$  cm ( $\sim 690 R_{\odot}$ ), and 2.3 foe, respectively, and the total mass of the progenitor, assuming 1.3–2.0  $M_{\odot}$  remnant mass,



**Figure 18.** *BVri* pseudo-bolometric light curve of SN 2015ba shown along with those of a sample of Type IIP SNe. *r/R* and *i/I* bands were used for the comparison sample depending on the availability of observations in these filters.



**Figure 19.** The best-fitting analytical model LC of SN 2015ba using Nagy & Vinkó (2016). The derived model parameters from the fitting are  $E = 2.3$  foe,  $R = 4.8 \times 10^{13}$  cm, and  $M_{\text{ej}} = 22 M_{\odot}$ .

**Table 10.** The best-fitting core and shell parameters for true bolometric light curve of SN 2015ba using Nagy & Vinkó (2016).

Parameter	Core	Shell	Remarks
$R_0$ (cm)	$4.8 \times 10^{13}$	$2.6 \times 10^{13}$	Initial radius of ejecta
$T_{\text{rec}}$ (K)	9500	–	Recombination temperature
$M_{\text{ej}}$ ( $M_{\odot}$ )	22	0.4	Ejecta mass
$E_{\text{th}}$ (foe)	0.8	0.09	Initial thermal energy
$E_{\text{kin}}$ (foe)	1.4	0.05	Initial kinetic energy
$M_{\text{Ni}}$ ( $M_{\odot}$ )	0.032	–	Initial $^{56}\text{Ni}$ mass
$\kappa$ ( $\text{cm}^2 \text{g}^{-1}$ )	0.14	0.4	Opacity
$A_{\text{g}}$ ( $\text{day}^2$ )	3e4	1e10	Gamma-ray leakage

is  $24 M_{\odot}$ . The mass of  $^{56}\text{Ni}$  estimated from the fit is  $0.032 M_{\odot}$ , which is consistent with the value derived from tail luminosity.

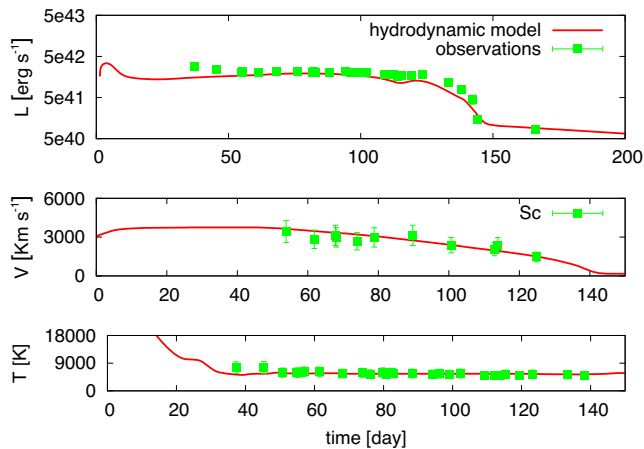
The parameters of the outer shell can also be estimated using this model, like the radius of the H-envelope ( $R = 2.6 \times 10^{13}$  cm). Although the poor sampling of the early light curve makes this value not reliable enough for consideration, yet we consider this value as a lower limit for this explosion. The data along with the best-fitting model is shown in Fig. 19. The parameters of the shell and core are listed in Table 10.

#### 8.4 Hydrodynamical modelling

We derive the SN progenitor’s physical properties at the time of the explosion (namely, the ejected mass  $M_{\text{ej}}$ , the progenitor radius  $R$ , and the total – kinetic plus thermal – energy  $E$ ) by means of a well-tested radiation-hydrodynamical modelling procedure already applied to many other observed SNe (e.g. SNe 2007od, 2009bw, 2009E, 2012A, 2012aw, 2012ec, 2013ab, 2013am, 2014cx, and OGLE-2014-SN-073; see Inserra et al. 2011, 2012; Pastorello et al. 2012; Tomasella et al. 2013, 2018; Dall’Ora et al. 2014; Barbarino et al. 2015; Bose et al. 2015a; Huang et al. 2016; Terreran et al. 2017, respectively).

A complete description of this procedure can be found in Pumo et al. (2017), here we recall that it includes the hydrodynamical modelling of all the main SN observables (i.e. bolometric light curve, evolution of line velocities, and the temperature at the photosphere) so as to constrain the SN progenitor’s physical properties through a simultaneous  $\chi^2$  fit of these observables against model calculations. Two codes are employed for computing the models. The first one is the semi-analytic code described in Zampieri et al. (2003), which solves the energy balance equation for an ejecta of constant density in homologous expansion. The second one is the general-relativistic, radiation-hydrodynamics Lagrangian code presented in Pumo, Zampieri & Turatto (2010) and Pumo & Zampieri (2011), which was specifically tailored to simulate the evolution of the physical properties of SN ejecta and the behaviour of the main SN observables up to the nebular stage, taking into account both the gravitational effects of the compact remnant and the heating effects due to the decay of the radioactive isotopes synthesized during the explosion. The semi-analytic code is used for a preparatory analysis aimed at individuating the parameter space describing the SN progenitor at the time of the explosion and, consequently, to guide the more realistic, but time consuming simulations performed with the general-relativistic, radiation-hydrodynamics Lagrangian code.

Based on the adopted explosion epochs (Section 3.2), bolometric luminosities (Section 8.1), and nickel masses (Section 8.2), we find the best-fitting model shown in Fig. 20. It has  $E = 1.6$  foe,  $R = 4.8 \times 10^{13}$  cm ( $\sim 690 R_{\odot}$ ), and  $M_{\text{ej}} = 24 M_{\odot}$ . Considering the  $2\sigma$



**Figure 20.** Comparison of the evolution of the main observables of SN 2015ba with the best-fitting model computed with the general-relativistic, radiation-hydrodynamics code. The best-fitting model parameters are  $E = 1.6$  foe,  $R = 4.8 \times 10^{13}$  cm, and  $M_{\text{ej}} = 24 M_{\odot}$ . Top, middle, and bottom panels show the bolometric light curve, the photospheric velocity, and the photospheric temperature as a function of time. To estimate the photosphere velocity from observations, we use the minima of the profile of the Sc lines.

confidence intervals for one parameter based on the  $\chi^2$  distributions produced by the semi-analytical models, we estimate that the error due to the  $\chi^2$  fitting procedure is about 15–20 per cent for  $M_{\text{ej}}$  and  $R$  and 30–40 per cent for  $E$  (see Pumo et al. 2017, for details on the procedure used for the error estimate).

The values of the best-fitting model parameters are consistent with the explosion of a relatively massive RSG star. In particular, adding the mass of the compact remnant ( $\sim 1.3\text{--}2 M_{\odot}$ ) to that of the ejected material, we obtain a total stellar mass at the time of the explosion of  $\sim 25.3\text{--}26 M_{\odot}$ . This value seems to indicate a progenitor mass clearly higher than the observational limit of  $16.5 \pm 1.5 M_{\odot}$  of the Type IIP events that raised the so-called ‘RSG problem’ (Smartt et al. 2009). High progenitor mass has also been estimated for SN 2012aw from the hydrodynamical modelling ( $\sim 22 M_{\odot}$ ; Dall’Ora et al. 2014). Moreover, the inferred progenitor’s mass is fully consistent (within the errors) with the maximum mass for an SN IIP progenitor found more recently by Walmswell & Eldridge (2012) although the dependability of such higher limit values is strongly debated (e.g. Kochanek, Khan & Dai 2012). We also note that the moderately small amount of oxygen and  $^{56}\text{Ni}$  found in the ejecta of SN 2015ba should be not in contrast with a relatively massive progenitor scenario given (1) the uncertainties in modelling the mixing processes of the stellar inner layers that may lead to a not clear connection between the progenitor mass and the nucleosynthesis efficiency (e.g. Jerkstrand et al. 2014), and (2) the non-monotonical link among the pre-supernova structure, the explosion properties, and the progenitor mass (Woosley, Heger & Weaver 2002; Sukhbold & Woosley 2014; Ertl et al. 2016; Sukhbold et al. 2016). Finally we point out that, although the progenitor masses estimated from the hydrodynamical modelling are generally higher to be consistent with those determined from direct observations of SN progenitors, the code used here gives lower progenitor masses (compared to other hydrodynamical codes) that are often consistent with mass estimates from the direct progenitor detection method and, more in general, from other independent methods such as the modelling of the observed nebular spectra (e.g. Tomasella et al. 2013; Pumo et al. 2017). Here, the inferred progenitor’s mass is in good agreement

with that found using the independent analytical modelling reported in Section 8.3.

## 9 SUMMARY

The photometric and spectroscopic data and analysis of SN 2015ba in IC 1029 is presented in this paper. This is a relatively bright SN IIP with a strikingly long plateau ( $\sim 123$  d). The absolute  $V$ -band magnitude at mid-plateau of SN 2015ba is similar to SN 2004et ( $M_V^p = -17.1$  mag; Sahu et al. 2006), while significantly lower ( $\sim 0.7$  mag) than the brightest SNe 2007od and DLT16am ( $-17.70$  and  $-17.73$  mag, respectively; Inserra et al. 2011; Tartaglia et al. 2018). The mean distance to the SN is estimated to be  $34.8 \pm 0.7$  Mpc. The  $^{56}\text{Ni}$  mass obtained from the tail of the light curve is  $0.032 \pm 0.006 M_{\odot}$ . Compared to its brightness, the  $^{56}\text{Ni}$  production in SN 2015ba is much lower than SN 2004et ( $0.06 M_{\odot}$ ). However, brighter events with comparatively lower  $^{56}\text{Ni}$  masses have been reported for the cases of SNe 2007od and 2009bw (Inserra et al. 2011, 2012). The drop in magnitude from the plateau to the radioactive tail is also relatively larger and steeper ( $\sim 3$  mag in 20 d in  $V$  band) than normal SNe IIP; however, smaller than the subluminous SNe IIP (e.g. SN 2005cs,  $\sim 3.8$  mag in 20 d in  $V$  band; Pastorello et al. 2009). We mapped the temperature evolution by fitting a blackbody to the SED obtained from the photometric fluxes of SN 2015ba, resulting in temperature as high as  $\sim 20\,000$  K at early epochs, falling to  $\sim 6300$  K at 50 d and finally settling to  $\sim 4800$  K at late epochs.

The early spectra of SN 2015ba exhibit a blue continuum with some broad Balmer P-Cygni profiles. The initial velocities of  $H\alpha$  and  $H\beta$  are  $\sim 9200$  km  $\text{s}^{-1}$  and  $\sim 8100$  km  $\text{s}^{-1}$ , respectively. The emergence of weak CSM interaction in SN 2015ba has been observed in the late photospheric stages evident from the appearance of HV features blueward of  $H\text{I}$  lines. These non-evolving, narrow features at similar epochs have been reported in SNe 1999em, 2004dj, and 2009bw (Chugai et al. 2007; Inserra et al. 2012).

The analytical and hydrodynamical modelling yield parameters of the progenitor and the explosion. While both analytical and hydrodynamical modelling gives consistent values of ejecta mass  $M_{\text{ej}}$ ,  $^{56}\text{Ni}$  mass, and pre-explosion radius, the inferred values of explosion energy from the analytical modelling (2.3 foe) is higher than the hydrodynamical modelling (1.6 foe). The mass of the progenitor of SN 2015ba estimated from the modelling ( $\sim 26 M_{\odot}$ ) significantly overshoots the observational limit of  $16.5 \pm 1.5 M_{\odot}$  of SNe IIP progenitors, which gave rise to the ‘RSG problem’ (Smartt et al. 2009). Nevertheless, massive RSG progenitors ( $20 < M_{\text{ZAMS}} < 30$ ) can give rise to plateaus as long as 140 d, which is consistent with SN 2015ba (Sukhbold et al. 2016). Also, we note the complete absence of [O I] 6300,6364 doublet in the nebular spectra of SN 2015ba, which is commonly found in massive progenitors. SN 2009bw is another SN IIP, coming from a relatively massive progenitor, which exhibited oxygen deprived nebular spectra. However, the non-monotonic variation of O-core masses with the ZAMS mass in massive pre-supernova stars ( $\sim 9 M_{\odot}$ ; Woosley et al. 2002; Sukhbold & Woosley 2014; Ertl et al. 2016; Sukhbold et al. 2016) and uncertainties in the mixing scenario may give rise to such an explosion. While hydrodynamical models have been known to produce higher estimates of progenitor mass, the code used in this work gives lower progenitor masses that are often found to be consistent with masses derived from other prevalent methods, like direct progenitor detection in pre-explosion images or modelling of nucleosynthesis yields in the nebular spectra (e.g. Tomasella et al. 2013; Pumo et al. 2017).

## ACKNOWLEDGEMENTS

Based on observations made with: Copernico 1.82 m Telescope operated by INAF Osservatorio Astronomico di Padova at Asiago, Italy, the Gran Telescopio Canarias (GTC) operated on the island of La Palma at the Spanish Observatorio del Roque de los Muchachos of the Instituto de Astrofísica de Canarias, the Nordic Optical Telescope (NOT), operated by the NOT Scientific Association at the Spanish Observatorio del Roque de los Muchachos of the Instituto de Astrofísica de Canarias. This work also makes use of observations with Las Cumbres Observatory (LCO) from the Supernova Key Project. We thank the observing staff and observing assistants of all the telescopes used in this work for their support during observations of SN 2015ba. We also thank S. Taubenberger and A. Floers for their support. We acknowledge Wiermann Interactive Supernova data REpository (WiSeREP) (<http://wiserep.weizmann.ac.il>; Yaron & Gal-Yam 2012). This research has made use of the NASA/IPAC Extragalactic Database (NED) that is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration. We acknowledge the usage of the HyperLeda database (<http://leda.univ-lyon1.fr>). DAH, GH, and CM are supported by NSF grant AST-1313484. Support for IA was provided by NASA through the Einstein Fellowship Program, grant PF6-170148. AP, LT, SB, PO, and MT are partially supported by the PRIN-INAF 2017 ‘Towards the SKA and CTA era: discovery, localization, and physics of transient sources’ (PI Giroletti). AMG acknowledges financial support by the University of Cádiz grant PR2017-64. SBP and KM acknowledges BRICS grant DST/IMRCD/BRICS/Pilotcall/ProFCheap/2017(G) for this work.

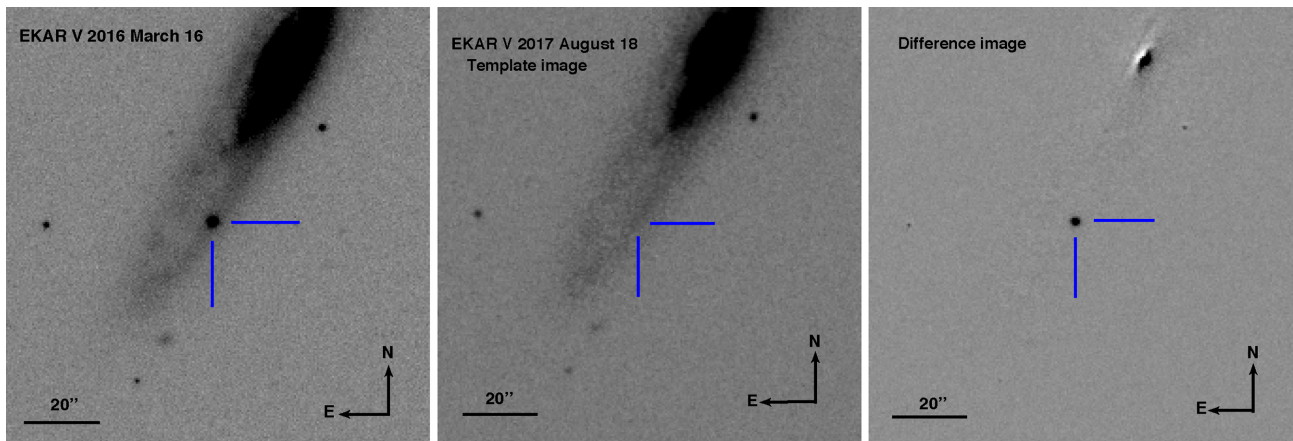
## REFERENCES

- Anderson J. P. et al., 2014a, *MNRAS*, 441, 671  
 Anderson J. P. et al., 2014b, *ApJ*, 786, 67  
 Anderson J. P. et al., 2018, *Nat. Astron.*, preprint ([arxiv:1805.04434](https://arxiv.org/abs/1805.04434)),  
 Arcavi I. et al., 2012, *ApJ*, 756, L30  
 Arnett W. D., 1980, *ApJ*, 237, 541  
 Arnett W. D., Fu A., 1989, *ApJ*, 340, 396  
 Barbarino C. et al., 2015, *MNRAS*, 448, 2312  
 Barbon R., Benetti S., Rosino L., Cappellaro E., Turatto M., 1990, *A&A*, 237, 79  
 Baron E. et al., 2000, *ApJ*, 545, 444  
 Benetti S. et al., 2015, *Astron. Telegram*, 8353  
 Bersten M. C., Hamuy M., 2009, *ApJ*, 701, 200  
 Bersten M. C., Benvenuto O., Hamuy M., 2011, *ApJ*, 729, 61  
 Bessell M. S., Castelli F., Plez B., 1998, *A&A*, 333, 231  
 Blondin S., Tonry J. L., 2007, *ApJ*, 666, 1024  
 Bose S. et al., 2013, *MNRAS*, 433, 1871  
 Bose S. et al., 2015a, *MNRAS*, 450, 2373  
 Bose S. et al., 2015b, *ApJ*, 806, 160  
 Cardelli J. A., Clayton G. C., Mathis J. S., 1989, *ApJ*, 345, 245  
 Chevalier R. A., 1976, *ApJ*, 207, 872  
 Chugai N. N., 1988, *Sov. Astron. Lett.*, 14, 334  
 Chugai N. N., Chevalier R. A., Utrobin V. P., 2007, *ApJ*, 662, 1136  
 Dall’Ora M. et al., 2014, *ApJ*, 787, 139  
 Dessart L., Hillier D. J., 2005a, *A&A*, 437, 667  
 Dessart L., Hillier D. J., 2005b, *A&A*, 439, 671  
 Dessart L., Hillier D. J., 2006, *A&A*, 447, 691  
 Dessart L., Livne E., Waldman R., 2010, *MNRAS*, 405, 2113  
 Eastman R. G., Schmidt B. P., Kirshner R., 1996, *ApJ*, 466, 911  
 Eldridge J. J., Tout C. A., 2004, *MNRAS*, 353, 87  
 Ertl T., Janka H.-T., Woosley S. E., Sukhbold T., Ugliano M., 2016, *ApJ*, 818, 124  
 Falk S. W., Arnett W. D., 1977, *ApJS*, 33, 515  
 Faran T. et al., 2014a, *MNRAS*, 442, 844  
 Faran T. et al., 2014b, *MNRAS*, 445, 554  
 Fraser M., 2016, *MNRAS*, 456, L16  
 Fukugita M., Ichikawa T., Gunn J. E., Doi M., Shimasaku K., Schneider D. P., 1996, *AJ*, 111, 1748  
 Gandhi P. et al., 2013, *ApJ*, 767, 166  
 Grassberg E. K., Imshennik V. S., Nadyozhin D. K., 1971, *Ap&SS*, 10, 28  
 Graur O., Bianco F. B., Modjaz M., Shivvers I., Filippenko A. V., Li W., Smith N., 2017, *ApJ*, 837, 121  
 Gutiérrez C. P. et al., 2017, *ApJ*, 850, 89  
 Hamuy M. et al., 2001, *ApJ*, 558, 615  
 Hamuy M., 2003, *ApJ*, 582, 905  
 Hamuy M., Pinto P. A., 2002, *ApJ*, 566, L63  
 Hamuy M., Suntzeff N. B., 1990, *AJ*, 99, 1146  
 Heger A., Fryer C. L., Woosley S. E., Langer N., Hartmann D. H., 2003, *ApJ*, 591, 288  
 Hendry M. A. et al., 2005, *MNRAS*, 359, 906  
 Huang F. et al., 2016, *ApJ*, 832, 139  
 Huang F. et al., 2018, *MNRAS*, 475, 3959  
 Inserra C. et al., 2011, *MNRAS*, 417, 261  
 Inserra C. et al., 2012, *MNRAS*, 422, 1122  
 Jeffery D. J., Branch D., 1990, in Wheeler J. C., Piran T., Weinberg S., eds, *Supernovae, Jerusalem Winter School for Theoretical Physics*. World Scientific Publishing Co., Singapore, p. 149  
 Jerkstrand A. et al., 2015, *MNRAS*, 448, 2482  
 Jerkstrand A., Fransson C., Maguire K., Smartt S., Ergon M., Spyromilio J., 2012, *A&A*, 546, A28  
 Jerkstrand A., Smartt S. J., Fraser M., Fransson C., Sollerman J., Taddia F., Kotak R., 2014, *MNRAS*, 439, 3694  
 Jordi K., Grebel E. K., Ammon K., 2006, *A&A*, 460, 339  
 Kirshner R. P., Kwan J., 1974, *ApJ*, 193, 27  
 Kochanek C. S., Khan R., Dai X., 2012, *ApJ*, 759, 20  
 Landolt A. U., 2009, *AJ*, 137, 4186  
 Leonard D. C. et al., 2002a, *PASP*, 114, 35  
 Leonard D. C. et al., 2002b, *AJ*, 124, 2490  
 Leonard D. C., Kanbur S. M., Ngeow C. C., Tanvir N. R., 2003, *ApJ*, 594, 247  
 Li W. et al., 2011, *MNRAS*, 412, 1441  
 Lusk J. A., Baron E., 2017, *PASP*, 129, 044202  
 Lyman J. D., Bersier D., James P. A., 2014, *MNRAS*, 437, 3848  
 Maguire K. et al., 2010, *MNRAS*, 404, 981  
 Maguire K. et al., 2012, *MNRAS*, 420, 3451  
 Makarov D., Prugniel P., Terekhova N., Courtois H., Vauglin I., 2014, *A&A*, 570, A13  
 Misra K., Pooley D., Chandra P., Bhattacharya D., Ray A. K., Sagar R., Lewin W. H. G., 2007, *MNRAS*, 381, 280  
 Munari U., Zwitter T., 1997, *A&A*, 318, 269  
 Nadyozhin D. K., 2003, *MNRAS*, 346, 97  
 Nagy A. P., Vinkó J., 2016, *A&A*, 589, A53  
 Nagy A. P., Ordasi A., Vinkó J., Wheeler J. C., 2014, *A&A*, 571, A77  
 Nugent P. et al., 2006, *ApJ*, 645, 841  
 Olivares E. F. et al., 2010, *ApJ*, 715, 833  
 Pastorello A. et al., 2004, *MNRAS*, 347, 74  
 Pastorello A. et al., 2006, *MNRAS*, 370, 1752  
 Pastorello A. et al., 2009, *MNRAS*, 394, 2266  
 Pastorello A. et al., 2012, *A&A*, 537, A141  
 Patat F., Barbon R., Cappellaro E., Turatto M., 1994, *A&A*, 282, 731  
 Poznanski D. et al., 2009, *ApJ*, 694, 1067  
 Poznanski D., Ganeshalingam M., Silverman J. M., Filippenko A. V., 2011, *MNRAS*, 415, L81  
 Pumo M. L., Zampieri L., 2011, *ApJ*, 741, 41  
 Pumo M. L., Zampieri L., 2013, *MNRAS*, 434, 3445  
 Pumo M. L., Zampieri L., Turatto M., 2010, *Mem. Soc. Astron. Ital.*, 14, 123  
 Pumo M. L., Zampieri L., Spiro S., Pastorello A., Benetti S., Cappellaro E., Manicò G., Turatto M., 2017, *MNRAS*, 464, 3013  
 Riess A. G. et al., 2016, *ApJ*, 826, 56  
 Roy R. et al., 2011, *ApJ*, 736, 76

- Sahu D. K., Anupama G. C., Srividya S., Muneer S., 2006, *MNRAS*, 372, 1315
- Schlafly E. F., Finkbeiner D. P., 2011, *ApJ*, 737, 103
- Silverman J. M. et al., 2017, *MNRAS*, 467, 369
- Smartt S. J., 2009, *ARA&A*, 47, 63
- Smartt S. J., 2015, *Publ. Astron. Soc. Aust.*, 32, e016
- Smartt S. J., Maund J. R., Hendry M. A., Tout C. A., Gilmore G. F., Mattila S., Benn C. R., 2004, *Science*, 303, 499
- Smartt S. J., Eldridge J. J., Crockett R. M., Maund J. R., 2009, *MNRAS*, 395, 1409
- Spiro S. et al., 2014, *MNRAS*, 439, 2873
- Stalin C. S., Hegde M., Sahu D. K., Parihar P. S., Anupama G. C., Bhatt B. C., Prabhu T. P., 2008, *Bull. Astron. Soc. India*, 36, 111
- Stetson P. B., 1987, *PASP*, 99, 191
- Stetson P. B., 1992, *J. R. Astron. Soc. Can.*, 86, 71
- Sukhbold T., Woosley S. E., 2014, *ApJ*, 783, 10
- Sukhbold T., Ertl T., Woosley S. E., Brown J. M., Janka H.-T., 2016, *ApJ*, 821, 38
- Takáts K. et al., 2014, *MNRAS*, 438, 368
- Takáts K. et al., 2015, *MNRAS*, 450, 3137
- Takáts K., Vinkó J., 2012, *MNRAS*, 419, 2783
- Tartaglia L. et al., 2018, *ApJ*, 853, 62
- Terreran G. et al., 2017, *Nat. Astron.*, 1, 713
- Thomas R. C., Nugent P. E., Meza J. C., 2011, *PASP*, 123, 237
- Tomasella L. et al., 2013, *MNRAS*, 434, 1636
- Tomasella L. et al., 2018, *MNRAS*, 475, 1937
- Turatto M. et al., 1998, *ApJ*, 498, L129
- Turatto M., Benetti S., Cappellaro E., 2003, in Hillebrandt W., Leibundgut B., eds, *From Twilight to Highlight: The Physics of Supernovae*. Springer-Verlag, Berlin, p. 200
- Utrobin V. P., Chugai N. N., 2008, *A&A*, 491, 507
- Utrobin V. P., Chugai N. N., 2009, *A&A*, 506, 829
- Valenti S. et al., 2014, *MNRAS*, 438, L101
- Valenti S. et al., 2016, *MNRAS*, 459, 3939
- van Dokkum P. G., 2001, *PASP*, 113, 1420
- Van Dyk S. D., Li W., Filippenko A. V., 2003, *PASP*, 115, 1289
- Vinkó J. et al., 2006, *MNRAS*, 369, 1780
- Walmswell J. J., Eldridge J. J., 2012, *MNRAS*, 419, 2054
- Woosley S. E., Heger A., Weaver T. A., 2002, *Rev. Mod. Phys.*, 74, 1015
- Yaron O., Gal-Yam A., 2012, *PASP*, 124, 668
- Zampieri L., Pastorello A., Turatto M., Cappellaro E., Benetti S., Altavilla G., Mazzali P., Hamuy M., 2003, *MNRAS*, 338, 711
- Zhang J. et al., 2014, *ApJ*, 797, 5

## APPENDIX A: PHOTOMETRY

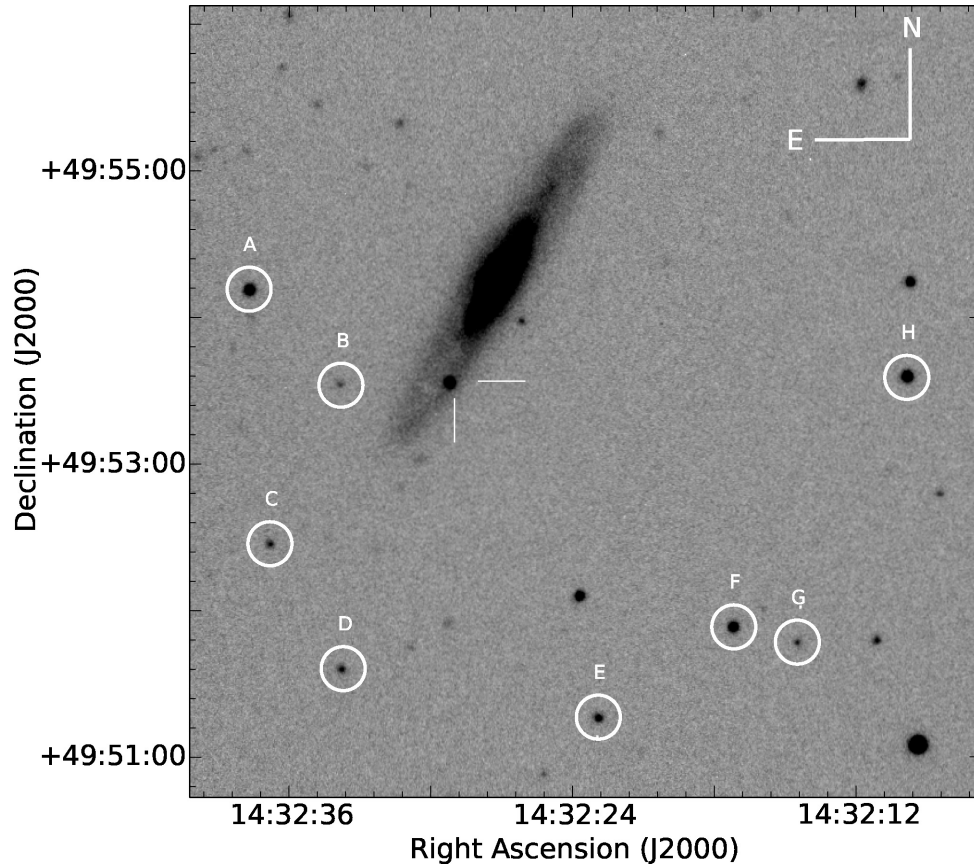
The images were pre-processed to eliminate detector signatures including overscan correction, bias subtraction, flat-field correction,



**Figure A1.** The left-hand panel taken with the 1.82 m EKAR telescope in V band on 2016 March 16 shows the presence of SN 2015ba. The middle panel is acquired with the EKAR telescope in V band on 2017 August 18 (~634 d since explosion) and shows the disappearance of the SN. The right-hand panel shows the subtracted image.

**Table A1.** Summary of the instruments used for the follow-up campaign of SN 2015ba.

Telescope	Location	Instrument	Pixel scale (arcsec pixel <sup>-1</sup> )	Imaging bands	Dispersers/ Grisms	
1 m LCO	McDonald Observatory, USA	SBIG/Sinistro	0.464/0.389	<i>BVgri</i>	–	1
1.04 m Sampurnanand Telescope (ST)	ARIES Observatory, India	Tek 1k×1k	0.53	<i>VRI</i>	–	2
1.3 m Devasthal Fast Optical Telescope (DFOT)	ARIES Observatory, India	Andor 512 × 512	0.64	<i>BVRI</i>	–	3
1.82 m Copernico Telescope	Asiago, Mount Ekar (Italy)	AFOSC	0.48	<i>uBVgriz</i>	Gr04, VPH7	4
2 m LCO Faulkes Telescope North (FTN)	Haleakala Observatory, USA	Spectral, FLOYDS	0.34	<i>BVgri</i>	Cross disperser	5
2.01 m Himalayan Chandra Telescope (HCT)	Indian Astronomical Observatory, Hanle, India	HFOSC	0.17	<i>BVRI</i>	–	6
2.56 m Nordic Optical Telescope (NOT)	Roque de los Muchachos, La Palma, Canaris (Spain)	ALFOSC	0.19	<i>uBVgriz</i>	Gr4	7
10.4 m Gran Telescopio CANARIS (GTC)	Roque de los Muchachos, La Palma, Canaris (Spain)	OSIRIS	0.25	<i>r</i>	R300B, R300R	8



**Figure A2.** SN 2015ba in IC 1029 along with eight local standard stars. This V-band image was taken on 2015 December 15 with a 1 m LCO telescope.

**Table A2.** Coordinates and photometry of the local sequence reference stars in the *BVRI* and in the *ugriz* systems.

ID	$\alpha_{J2000.0}$ (hh:mm:ss)	$\delta_{J2000.0}$ (dd:mm:ss)	<i>B</i> mag	<i>V</i> mag	<i>R</i> mag	<i>I</i> mag	<i>u</i> mag	<i>g</i> mag	<i>r</i> mag	<i>i</i> mag	<i>z</i> mag
A	14:32:37.5	+49:54:12.4	17.690 (.060)	17.060 (.040)	16.560 (.005)	16.245 (.007)	18.390 (.010)	17.215 (.004)	16.781 (.004)	16.650 (.005)	16.600 (.009)
B	14:32:34.0	+49:53:32.0	20.300 (.400)	19.410 (.060)	18.460 (.010)	17.850 (.010)	22.900 (.400)	20.030 (.020)	18.780 (.010)	18.291 (.009)	17.990 (.020)
C	14:32:36.8	+49:52:23.6	19.800 (.100)	18.840 (.080)	18.137 (.011)	17.691 (.007)	21.220 (.090)	19.210 (.010)	18.368 (.008)	18.100 (.008)	17.910 (.020)
D	14:32:33.7	+49:51:31.3	20.000 (.100)	18.680 (.060)	17.683 (.005)	16.967 (.006)	21.800 (.200)	19.370 (.010)	18.017 (.007)	17.413 (.006)	17.090 (.010)
E	14:32:22.8	+49:51:15.2	18.860 (.060)	18.200 (.060)	17.644 (.006)	17.277 (.007)	19.650 (.030)	18.396 (.007)	17.893 (.007)	17.723 (.007)	17.700 (.020)
F	14:32:17.1	+49:51:52.5	17.918 (.100)	17.400 (.040)	16.949 (.003)	16.632 (.005)	18.510 (.010)	17.488 (.005)	17.171 (.005)	17.068 (.006)	17.060 (.010)
G	14:32:14.3	+49:51:47.5	19.900 (.200)	19.100 (.100)	18.508 (.008)	18.110 (.010)	21.200 (.100)	19.390 (.010)	18.740 (.010)	18.510 (.010)	18.400 (.030)
H	14:32:09.9	+49:53:34.2	17.630 (.060)	16.960 (.030)	16.435 (.004)	16.094 (.003)	18.460 (.020)	17.133 (.004)	16.673 (.004)	16.490 (.005)	16.466 (.008)

and trimming, following standard procedures in IRAF.<sup>3</sup> Cosmic hits were detected and removed using the laplacian detection algorithm L.A.Cosmic (van Dokkum 2001). When multiple frames were taken on the same night, the frames were median-combined in order to improve the signal-to-noise ratio. Point spread function (PSF) fitting photometry was performed on the images using DAOPHOT II (Stetson 1987) to derive the SN magnitudes. The mean FWHM of the frame was used as the first aperture and a 8 pixel wide sky background annulus was used, with an inner radius of four times the FWHM.

<sup>3</sup>IRAF stands for Image Reduction and Analysis Facility distributed by the National Optical Astronomy Observatories that is operated by the Association of Universities for research in Astronomy, Inc., under cooperative agreement with the National Science Foundation.

Considering the location of the SN, we performed template subtraction to eliminate any possible contamination of the host galaxy flux in the SN magnitudes. The template images used for subtraction were obtained on 2017 February 26, May 05, and August 18 for the different instrumental setups listed in Table A1. For template subtraction, the images were first registered astrometrically to the template image and subsequently subtracted using High Order Transform of PSF AND Template Subtraction (HOTPANTS).<sup>4</sup> We used the PSF constructed using the unsubtracted SN frames to derive the SN magnitudes from the subtracted frames. Fig. A1 displays 2016 March 16 image before and after performing template subtraction.

<sup>4</sup><https://github.com/acbecker/hotpants>

**Table A3.** Optical photometry of SN 2015ba.

UT date (yyyy-mm-dd)	JD 2457000+	Phase <sup>a</sup> (d)	<i>B</i> (mag)	<i>V</i> (mag)	<i>R</i> (mag)	<i>I</i> (mag)	Tel
2015-12-02.18	358.68	8.98	16.74 ± 0.06	–	–	–	4
2015-12-03.21	359.71	10.01	16.76 ± 0.02	16.51 ± 0.06	–	–	4
2015-12-03.52	360.02	10.32	16.69 ± 0.02	16.47 ± 0.02	–	–	1
2015-12-04.16	360.66	10.96	16.74 ± 0.02	16.48 ± 0.02	–	–	4
2015-12-04.52	361.02	11.32	16.69 ± 0.02	16.44 ± 0.02	–	–	1
2015-12-06.24	362.74	13.04	16.82 ± 0.01	16.45 ± 0.02	–	–	4
2015-12-07.52	364.02	14.32	16.77 ± 0.02	16.50 ± 0.02	–	–	1
2015-12-10.50	367.00	17.3	16.80 ± 0.07	16.68 ± 0.07	–	–	1
2015-12-11.24	367.63	17.93	16.96 ± 0.01	16.61 ± 0.01	–	–	4
2015-12-11.53	368.03	18.33	16.95 ± 0.07	16.81 ± 0.08	–	–	1
2015-12-16.53	373.03	23.33	17.09 ± 0.02	16.68 ± 0.02	–	–	1
2015-12-19.13	375.71	26.01	17.34 ± 0.09	16.74 ± 0.10	–	–	1
2015-12-20.52	377.02	27.32	17.36 ± 0.04	16.79 ± 0.02	–	–	1
2015-12-30.52	387.02	37.32	17.73 ± 0.05	16.90 ± 0.04	–	–	1
2016-01-01.96	389.47	39.77	–	16.91 ± 0.02	16.32 ± 0.01	15.98 ± 0.01	2
2016-01-02.97	390.48	40.78	–	17.06 ± 0.02	16.36 ± 0.01	16.01 ± 0.01	2
2016-01-07.52	395.02	45.32	18.11 ± 0.03	17.05 ± 0.02	–	–	1
2016-01-13.02	400.50	50.8	18.20 ± 0.01	17.07 ± 0.01	16.38 ± 0.01	16.12 ± 0.01	3
2016-01-14.02	401.53	51.83	–	17.12 ± 0.01	16.39 ± 0.01	16.05 ± 0.01	3
2016-01-16.03	403.53	53.83	–	17.29 ± 0.03	16.40 ± 0.01	16.28 ± 0.01	3
2016-01-17.03	404.53	54.83	17.58 ± 0.01	17.24 ± 0.01	16.41 ± 0.01	16.50 ± 0.03	3
2016-01-17.64	405.14	55.44	18.27 ± 0.02	17.15 ± 0.02	–	–	5
2016-01-19.21	406.68	56.98	18.48 ± 0.01	17.16 ± 0.01	–	–	4
2016-01-20.94	408.45	58.75	–	17.34 ± 0.04	16.38 ± 0.01	16.10 ± 0.01	2
2016-01-23.62	411.12	61.42	18.23 ± 0.04	17.26 ± 0.03	–	–	5
2016-01-30.22	417.72	68.02	18.82 ± 0.05	17.17 ± 0.02	–	–	7
2016-01-31.89	419.40	69.7	–	17.24 ± 0.01	16.41 ± 0.01	16.02 ± 0.01	2
2016-02-02.87	421.38	71.68	–	17.28 ± 0.03	16.41 ± 0.01	16.00 ± 0.01	2
2016-02-05.21	423.72	74.02	18.71 ± 0.04	17.19 ± 0.01	–	–	4
2016-02-07.33	425.83	76.13	18.42 ± 0.03	17.08 ± 0.02	–	–	1
2016-02-10.98	429.48	79.78	18.52 ± 0.01	17.23 ± 0.01	16.41 ± 0.01	16.10 ± 0.01	3
2016-02-11.91	430.41	80.71	18.58 ± 0.01	17.23 ± 0.01	16.41 ± 0.01	15.96 ± 0.01	3
2016-02-12.93	431.43	81.73	18.62 ± 0.01	17.28 ± 0.01	16.41 ± 0.01	15.99 ± 0.01	3
2016-02-13.97	432.47	82.77	18.64 ± 0.01	17.25 ± 0.01	16.41 ± 0.01	15.99 ± 0.01	3
2016-02-13.47	431.97	82.27	18.66 ± 0.02	17.22 ± 0.02	–	–	1
2016-02-19.48	437.98	88.28	18.67 ± 0.03	17.24 ± 0.02	–	–	1
2016-02-25.35	443.85	94.15	18.65 ± 0.03	17.21 ± 0.02	–	–	1
2016-02-27.42	445.92	96.22	18.73 ± 0.03	17.25 ± 0.01	–	–	1
2016-03-01.38	448.88	99.18	18.74 ± 0.02	17.26 ± 0.01	–	–	1
2016-03-04.35	451.85	102.15	18.68 ± 0.05	17.24 ± 0.02	–	–	1
2016-03-07.96	455.46	105.66	18.82 ± 0.01	17.33 ± 0.01	16.42 ± 0.01	16.08 ± 0.01	3
2016-03-11.32	458.82	109.12	18.79 ± 0.02	17.31 ± 0.01	–	–	1
2016-03-14.48	461.98	112.28	18.83 ± 0.02	17.34 ± 0.02	–	–	1
2016-03-15.88	463.37	113.67	–	–	16.45 ± 0.02	–	2
2016-03-16.21	463.60	113.9	18.90 ± 0.02	17.40 ± 0.01	–	–	7
2016-03-17.48	464.98	115.28	18.93 ± 0.02	17.40 ± 0.01	–	–	1
2016-03-18.89	466.42	116.72	–	17.37 ± 0.02	–	16.06 ± 0.02	2
2016-03-19.96	467.45	117.75	–	–	16.51 ± 0.01	16.10 ± 0.02	2
2016-03-21.47	468.97	119.27	18.89 ± 0.04	17.40 ± 0.03	–	–	1
2016-03-25.47	472.97	123.27	18.70 ± 0.10	17.41 ± 0.08	–	–	1
2016-04-04.46	482.96	133.26	19.36 ± 0.03	17.86 ± 0.02	–	–	1
2016-04-09.42	487.92	138.22	19.92 ± 0.07	18.26 ± 0.05	–	–	1
2016-04-13.43	491.93	142.23	20.29 ± 0.10	19.40 ± 0.10	–	–	1
2016-04-14.91	493.41	143.61	–	20.17 ± 0.02	19.14 ± 0.01	18.62 ± 0.01	3
2016-04-15.23	493.73	144.03	–	20.16 ± 0.07	–	–	1
2016-05-02.81	511.31	161.61	–	–	19.56 ± 0.08	19.03 ± 0.04	2

We observed three Landolt standard fields PG 918, PG 942, and PG 1525 (Landolt 2009) on 2017 February 26 under photometric night conditions (FWHM seeing in  $V \sim 2$  arcsec) with the 2.01 m Himalayan Chandra Telescope (HCT) at Indian Astronomical Observatory (IAO), Hanle, to convert the instrumental magnitudes in

$BVRI$  to standard magnitudes. The  $V$ -band magnitudes of the stars in the standard field lies in the range 12.27–16.4 mag and  $B - V$  colour varies between  $-0.271$  and  $1.109$  mag. These fields were observed at three altitudes covering airmass from 1.15 to 1.58. With the aid of the instrumental and standard magnitudes of the Landolt field

Table A3 – continued

UT date (yyyy-mm-dd)	JD 2457000+	Phase <sup>a</sup> (d)	<i>u</i> (mag)	<i>g</i> (mag)	<i>r</i> (mag)	<i>i</i> (mag)	<i>z</i> (mag)	Tel
2015-12-02.18	358.68	8.98	16.82 ± 0.06	16.53 ± 0.05	16.29 ± 0.04	16.23 ± 0.01	16.20 ± 0.04	4
2015-12-03.21	359.70	10	16.74 ± 0.02	16.48 ± 0.05	16.27 ± 0.05	16.24 ± 0.04	16.20 ± 0.06	4
2015-12-04.16	360.66	10.96	16.87 ± 0.01	16.63 ± 0.02	16.27 ± 0.02	16.16 ± 0.03	16.17 ± 0.03	4
2015-12-06.24	362.74	13.04	16.90 ± 0.01	–	–	–	–	4
2015-12-07.51	364.01	14.31	–	16.90 ± 0.01	16.27 ± 0.03	16.23 ± 0.04	–	1
2015-12-11.24	367.62	17.92	17.13 ± 0.06	16.663 ± 0.005	16.29 ± 0.01	16.23 ± 0.01	16.20 ± 0.01	4
2015-12-14.50	371.00	21.3	–	16.55 ± 0.07	16.31 ± 0.03	16.32 ± 0.04	–	1
2015-12-19.13	375.73	26.03	19.50 ± 0.10	16.82 ± 0.01	16.49 ± 0.02	16.30 ± 0.01	18.60 ± 0.10	4
2015-12-20.53	377.03	27.33	–	16.87 ± 0.08	16.44 ± 0.03	16.44 ± 0.04	–	1
2015-12-30.53	387.03	37.33	–	17.22 ± 0.08	16.52 ± 0.03	16.38 ± 0.04	–	1
2016-01-07.53	395.03	45.33	–	17.45 ± 0.07	16.63 ± 0.02	16.56 ± 0.03	–	1
2016-01-17.64	405.14	55.44	–	17.58 ± 0.07	16.68 ± 0.02	16.51 ± 0.03	–	5
2016-01-19.21	406.69	56.99	19.86 ± 0.08	17.61 ± 0.01	16.67 ± 0.01	16.49 ± 0.01	16.36 ± 0.01	4
2016-01-23.62	411.12	61.42	–	17.48 ± 0.08	16.71 ± 0.02	16.44 ± 0.02	–	5
2016-01-25.47	412.97	63.27	–	–	16.63 ± 0.01	–	–	1
2016-01-30.22	417.72	68.02	–	17.78 ± 0.01	16.67 ± 0.01	16.52 ± 0.01	16.46 ± 0.04	4
2016-02-05.21	423.72	74.02	19.7 ± 0.1	17.77 ± 0.02	16.68 ± 0.01	16.48 ± 0.01	16.38 ± 0.01	4
2016-02-07.34	425.84	76.14	–	17.60 ± 0.08	16.66 ± 0.02	16.43 ± 0.03	–	1
2016-02-13.48	431.98	82.28	–	17.84 ± 0.01	16.62 ± 0.01	16.45 ± 0.01	–	1
2016-02-19.48	437.98	88.28	–	17.86 ± 0.01	16.69 ± 0.01	16.46 ± 0.01	–	1
2016-02-25.36	443.86	94.16	–	17.87 ± 0.01	16.70 ± 0.01	16.46 ± 0.01	–	1
2016-02-27.44	445.94	96.24	–	17.89 ± 0.01	16.71 ± 0.01	16.46 ± 0.01	–	1
2016-03-01.39	448.89	99.19	–	17.89 ± 0.01	16.72 ± 0.01	16.49 ± 0.01	–	1
2016-03-04.36	451.86	102.16	–	17.90 ± 0.01	16.73 ± 0.01	16.50 ± 0.01	–	1
2016-03-11.33	458.83	109.13	–	17.89 ± 0.07	16.85 ± 0.02	16.55 ± 0.02	–	1
2016-03-14.50	462.00	112.30	–	17.92 ± 0.07	16.82 ± 0.02	16.57 ± 0.02	–	1
2016-03-16.21	463.60	113.90	19.50 ± 0.10	18.03 ± 0.01	16.86 ± 0.01	16.74 ± 0.02	16.49 ± 0.03	4
2016-03-17.49	464.99	115.29	–	18.05 ± 0.01	16.85 ± 0.01	16.62 ± 0.01	–	1
2016-03-21.49	468.99	119.29	–	17.99 ± 0.08	16.88 ± 0.03	16.64 ± 0.03	–	1
2016-03-25.48	472.98	123.28	–	18.17 ± 0.03	16.90 ± 0.01	16.66 ± 0.01	–	1
2016-04-04.48	482.98	133.28	–	18.56 ± 0.01	17.24 ± 0.01	17.01 ± 0.01	–	1
2016-04-09.44	487.94	138.24	–	19.04 ± 0.08	17.74 ± 0.02	17.50 ± 0.03	–	1
2016-04-13.44	491.94	142.24	–	20.25 ± 0.04	18.95 ± 0.02	19.06 ± 0.04	–	1
2016-04-15.25	493.75	144.05	–	20.67 ± 0.06	19.53 ± 0.03	19.46 ± 0.05	–	1
2016-05-05.40	513.90	164.20	–	–	20.22 ± 0.09	–	–	1
2016-05-07.33	515.83	166.13	–	21.20 ± 0.10	20.27 ± 0.08	20.19 ± 0.09	–	1
2016-05-12.42	520.92	171.22	–	–	20.20 ± 0.10	–	–	1
2016-06-17.96	557.46	207.76	–	–	21.20 ± 0.20	20.50 ± 0.10	20.10 ± 0.10	4
2016-08-02.90	602.90	253.20	–	–	20.30 ± 0.10	–	–	8
2016-08-08.92	609.42	259.72	–	–	21.40 ± 0.10	21.40 ± 0.10	20.29 ± 0.09	4
2016-08-21.87	621.87	272.17	–	–	20.80 ± 0.10	–	–	8

<sup>a</sup>Since the explosion epoch  $t_0 = 2457349.7$  JD.

stars, the zero points and colour coefficients of the transformation equations were fitted using the least square linear regression technique as described in Stetson (1992) and the values thus obtained are given below:

$$b = B + (1.632 \pm 0.001) + (0.022 \pm 0.003)(B - V) + k_b \cdot X$$

$$v = V + (1.001 \pm 0.001) + (-0.057 \pm 0.001)(B - V) + k_v \cdot X$$

$$r = R + (1.018 \pm 0.004) + (0.03 \pm 0.03)(V - R) + k_r \cdot X$$

$$i = I + (1.264 \pm 0.003) + (0.065 \pm 0.004)(V - I) + k_i \cdot X,$$

where  $X$  is the airmass and the site extinction coefficients  $k_\lambda$  are 0.21, 0.12, 0.08, and 0.05 mag per unit airmass for the  $B$ ,  $V$ ,  $R$ , and  $I$  bands, respectively (Stalin et al. 2008). The rms scatter between the observed and standard magnitudes of the Landolt field stars are

$\sim 0.08$  mag in  $B$ ,  $\sim 0.04$  mag in  $V$  and  $I$ , and  $\sim 0.03$  mag in  $R$ . We used these coefficients to generate a photometric sequence of eight non-variable stars in the SN field. Observations of standard fields and subsequent calibrations were also performed with the LCO telescopes on multiple photometric nights. We picked the eight local standards chosen before and calibrated their magnitudes in  $ugriz$  and  $BV$  filters. Since, the  $B$  and  $V$  bands are in common to ARIES, LCO, and ASIAGO data, to maintain uniformity we preferred to use LCO secondary standard magnitudes for all  $BV$  bands data. Similarly, to calibrate the  $ugriz$  data obtained from LCO and ASIAGO, we used the LCO secondary standard  $ugriz$  magnitudes. The eight local standards are marked in Fig. A2, and their magnitudes are tabulated in Table A2. Night to night zero points were estimated and applied for obtaining the final SN magnitudes. The combined PSF fitting and photometric calibration errors propagated in quadrature are listed in Table A3.

**APPENDIX B: SPECTROSCOPY**

We performed the initial pre-processing steps on each frame, like bias subtraction, flat-fielding, trimming, and eliminating the cosmic hits from the frames using the L.A. Cosmic routine (van Dokkum 2001). One-dimensional spectra were extracted using the APALL task, and then wavelength and flux calibrated using arc lamps and spectrophotometric standard-star spectra, respectively, observed at a similar airmass during the same night. Night sky emission lines were used to check the wavelength calibration, and shifts were

applied when necessary. On the nights when standard star observations were unavailable, we used the sensitivity function, generated using STANDARD and SENSFUNC tasks, of the standard star observed in a close by night for flux calibration. The calibrated spectra were compared with the photometry and, when necessary, multiplied by a constant factor to correct for atmospheric transparency, bad-seeing conditions, and slit losses to match the spectroscopic and photometric continuum flux. Finally, each spectra were de-redshifted to the heliocentric frame using the *DOPCOR* task.

**Table B1.** Log of the spectroscopic observations.

Date	Phase <sup>a</sup> (d)	Grism	Spectral range (Å)	Resolution	Instrument	Telescope
2015-12-02	8.91	Gr04	3360–7740	311	AFOSC	EKAR
2015-12-04	10.95	Gr04	3360–7740	311	AFOSC	EKAR
2015-12-06	12.93	Gr04	3360–7740	311	AFOSC	EKAR
2015-12-07	13.92	–	3200–10 000	400–700	FLOYDS	FTN
2015-12-08	14.90	–	3200–10 000	400–700	FLOYDS	FTN
2015-12-11	17.88	Gr04	3360–7740	311	AFOSC	EKAR
2015-12-12	18.90	–	3200–10 000	400–700	FLOYDS	FTN
2015-12-17	23.90	–	3200–10 000	400–700	FLOYDS	FTN
2015-12-19	26.00	VPH7	3200–7000	311	AFOSC	EKAR
2015-12-30	36.91	–	3200–10 000	400–700	FLOYDS	FTN
2016-01-08	45.91	–	3200–10 000	400–700	FLOYDS	FTN
2016-01-16	53.83	–	3200–10 000	400–700	FLOYDS	FTN
2016-01-24	61.81	–	5400–10 000	400–700	FLOYDS	FTN
2016-01-30	68.02	Gr04	3200–9600	311	ALFOSC	NOT
2016-01-30	67.85	–	3200–10 000	400–700	FLOYDS	FTN
2016-02-05	73.97	Gr04	3360–7740	311	AFOSC	EKAR
2016-02-10	78.81	–	3200–10 000	400–700	FLOYDS	FTN
2016-02-21	89.77	–	5400–10 000	400–700	FLOYDS	FTN
2016-03-03	100.68	–	3200–10 000	400–700	FLOYDS	FTN
2016-03-15	112.89	–	3200–10 000	400–700	FLOYDS	FTN
2016-03-16	113.90	Gr04	3200–9600	311	ALFOSC	NOT
2016-03-27	124.88	–	3200–10 000	400–700	FLOYDS	FTN
2016-04-09	137.88	–	3200–10 000	400–700	FLOYDS	FTN
2016-04-11	139.83	–	3200–10 000	400–700	FLOYDS	FTN
2016-04-12	140.61	–	5400–10 000	400–700	FLOYDS	FTN
2016-08-02	253.23	R300B	3600–7200	360	OSIRIS	GTC
2016-08-21	272.16	R300R	4800–10 000	348	OSIRIS	GTC

<sup>a</sup>Since explosion epoch  $t_0 = 2457349.7$  JD (2015 November 23).