



Publication Year	2017
Acceptance in OA	2020-12-30T13:48:58Z
Title	Probing the interstellar dust towards the Galactic Centre: dust-scattering halo around AX J1745.6-2901
Authors	Jin, Chichuan, PONTI, GABRIELE, Haberl, Frank, Smith, Randall
Publisher's version (DOI)	10.1093/mnras/stx653
Handle	http://hdl.handle.net/20.500.12386/29377
Journal	MONTHLY NOTICES OF THE ROYAL ASTRONOMICAL SOCIETY
Volume	468

Probing the Interstellar Dust towards the Galactic Centre: Dust Scattering Halo around AX J1745.6-2901

Chichuan Jin^{1*}, Gabriele Ponti¹, Frank Haber¹, Randall Smith²

¹Max-Planck-Institut für Extraterrestrische Physik, Giessenbachstrasse, D-85748 Garching, Germany

²Smithsonian Astrophysical Observatory, 60 Garden Street, Cambridge, MA 02138, USA

prepared for MNRAS

ABSTRACT

AX J1745.6-2901 is an X-ray binary located at only 1.45 arcmin from Sgr A^{*}, showcasing a strong X-ray dust scattering halo. We combine *Chandra* and *XMM-Newton* observations to study the halo around this X-ray binary. Our study shows two major thick dust layers along the line of sight (LOS) towards AX J1745.6-2901. The LOS position and N_H of these two layers depend on the dust grain models with different grain size distribution and abundances. But for all the 19 dust grain models considered, dust Layer-1 is consistently found to be within a fractional distance of 0.11 (mean value: 0.05) to AX J1745.6-2901 and contains only (19-34)% (mean value: 26%) of the total LOS dust. The remaining dust is contained in Layer-2, which is distributed from the Earth up to a mean fractional distance of 0.64. A significant separation between the two layers is found for all the dust grain models, with a mean fractional distance of 0.31. Besides, an extended wing component is discovered in the halo, which implies a higher fraction of dust grains with typical sizes $\lesssim 590 \text{ \AA}$ than considered in current dust grain models. Assuming AX J1745.6-2901 is 8 kpc away, dust Layer-2 would be located in the Galactic disk several kpc away from the Galactic Centre (GC). The dust scattering halo biases the observed spectrum of AX J1745.6-2901 severely in both spectral shape and flux, and also introduces a strong dependence on the size of the instrumental point spread function and the source extraction region. We build XSPEC models to account for this spectral bias, which allow us to recover the intrinsic spectrum of AX J1745.6-2901 free from dust scattering opacity. If dust Layer-2 also intervenes along the LOS to Sgr A^{*} and other nearby GC sources, a significant spectral correction for the dust scattering opacity would be necessary for all these GC sources.

Key words: dust, extinction - X-rays: ISM; neutron stars: X-rays: binaries: individual: AX J1745.6-2901

1 INTRODUCTION

The centre of the Milky Way (i.e. the Galactic Centre, hereafter: GC) is one of the most attractive regions for astrophysical studies. It has been observed that thousands of point-like X-ray sources and extended sources are sitting in this region (e.g. Wang, Gotthelf & Lang 2002; Munro et al. 2003, 2005, 2009; Degenaar et al. 2012, 2015; Ponti et al. 2015a, 2016), including Sgr A^{*} which is the super-massive black hole (SMBH) closest to Earth (e.g. Genzel, Eisenhauer & Gillessen 2010). The inner few hundred parsecs contain a large concentration of molecular material (e.g. Morris 1996), which is revealing traces of previous active periods of Sgr A^{*} (see Ponti et al. 2013 for a review). However, the GC region

is also highly extinguished from the optical to the soft X-ray band (e.g. Becklin & Neugebauer 1968; Fritz et al. 2011), elevating the importance of understanding the properties of gas and dust in front of the GC.

Previous studies about the diffuse emission and point sources around the GC LOS often suggest a large hydrogen column density of $N_H \gtrsim 10^{22} \text{ cm}^{-2}$, as measured by absorption models in the X-ray spectral fitting (hereafter: $N_{H,abs}$; e.g. Terrier et al. 2010; Molinari et al. 2011; Degenaar et al. 2012). The large $N_{H,abs}$ also implies a large dust column assuming typical gas-to-dust ratios (e.g. Predehl & Schmitt 1995). Recently, several papers show that dust scattering can have a significant impact on the observed X-ray spectra (Smith, Valencic & Corrales 2016), causing significant bias to $N_{H,abs}$ if the dust scattering opacity is not properly considered in the spectral fitting (Corrales et al. 2016).

* E-mail: chichuan@mpe.mpg.de

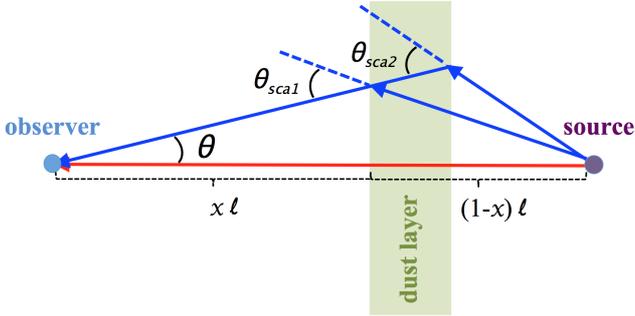


Figure 1. Geometry of the dust scattering by a single thick dust layer. θ is the viewing angle. θ_{sca1} and θ_{sca2} are the two scattering angles at the same viewing angle but different location in the dust layer. ℓ is the absolute distance between the source and the observer, x is the fractional distance from the dust layer to the observer.

Therefore, it is especially important to consider the effects of dust scattering when studying X-ray sources in the GC direction.

1.1 Dust Scattering Halo in X-rays

The idea that dust grains along the LOS can produce an observable halo around X-ray sources, the so called ‘dust scattering halo’, was proposed fifty years ago (Overbeck 1965; Trümper & Schönfelder 1973) and was first observed by the *Einstein* satellite around GX 339-4 (Rolf 1983). The investigation of the dust scattering halo can provide important information about the location of the dust along the LOS as well as the size and composition of the dust grains (e.g. Predehl et al. 1992; Predehl & Schmitt 1995; Xiang, Zhang & Yao 2005; Valencic & Smith 2015; Heinz et al. 2016; Vasilopoulos & Petropoulou 2016).

1.1.1 Basic Principles of the Dust Scattering Theory

We assume a point source that is located at a distance ℓ from the observer characterized by a flux F_X and absorbed by neutral material with a hydrogen column density N_H . With assumptions made in the dust grain models about the dust-to-gas ratio and abundances, this N_H is determined by the total number of dust grains required (hereafter: $N_{H,sca}$). Moreover, we assume that dust grains are distributed along the LOS with a normalized distribution function $f(x)$ (where x ($0 \leq x \leq 1$) is the fractional distance between the observer and the dust grains, as drawn in Fig.1). The intensity of the scattering light at the viewing angle θ , due to a single dust scattering (Mathis & Lee 1991), is given by:

$$I_{sca}^{(1)}(\theta) = F_X N_{H,sca} \int_{E_{min}}^{E_{max}} S(E) \int_0^1 \frac{f(x)}{(1-x)^2} \int_{a_{min}}^{a_{max}} n(a) \times \frac{d\sigma_{sca}(a, x, E, \theta)}{d\Omega} da dx dE \quad (1)$$

where $S(E)$, $n(a)$, $\sigma_{sca}(a, x, E, \theta)$ are the normalized source spectrum, dust grain size distribution and the scattering cross section, respectively. $n(a)$ depends on the dust grain model and contains assumptions about the abundances and

the dust-to-gas ratio. The wavelength of the X-ray photon is much smaller than the typical size range of dust grains in the ISM, and the refraction index of the dust grain is close to unity. Therefore, the Rayleigh-Gans approximation, instead of a full Mie calculation, can be used to simplify the calculation (e.g. Mathis & Lee 1991). So the differential cross section can be written as:

$$\frac{d\sigma_{sca}}{d\Omega} = (1.1 \text{ cm}^2 \text{ sr}^{-1}) \left(\frac{2Z}{M}\right)^2 \left(\frac{\rho}{3 \text{ g cm}^{-3}}\right) \left[\frac{F(E)}{Z}\right]^2 \times \left(\frac{a}{1 \mu\text{m}}\right)^6 \Phi^2(a, x, E, \theta) \quad (2)$$

where $\Phi(a, x, E, \theta)$ is the form factor (Bohren & Huffman 1983) that can be approximated by a Gaussian form for a spherical dust grain with both good accuracy and high computing efficiency:

$$\Phi^2(a, x, E, \theta) = \exp[-0.4575 E(\text{keV})^2 a(\mu\text{m})^2 \theta_{sca}(\text{arcmin})^2] \quad (3)$$

where the scattering angle θ_{sca} can be derived from x and θ from simple geometric calculations:

$$\theta_{sca} = \arctan\left(\frac{x}{1-x} \cdot \tan\theta\right) + \theta \quad (4)$$

The advantage of using the Gaussian approximation is the much higher computing efficiency with good accuracy. Mathis & Lee (1991) compared the scattering profiles between the exact form factor and the Gaussian approximation, and found at most 10 percent difference at all scattering angles. According to the above equations, we can also define a typical dust grain size for a specific θ_{sca} as:

$$a(\mu\text{m}) = \frac{1.4785}{E(\text{keV}) \theta_{sca}(\text{arcmin})} \quad (5)$$

Note that for a dust layer with a non-negligible thickness, the scattering light at a specific viewing angle θ comprises a range of scattering angles ($\theta_{sca1} - \theta_{sca2}$, Fig.1). It can therefore be predicted that the dust scattering from a distinct layer will produce a halo around X-ray sources. Eq.3 indicates that the halo intensity is the strongest in the core and fades out quickly outside a few arcmin, making a point source slightly extended. The typical halo size changes with the photon energy as $\propto \exp(-E^2)$. Therefore, a strong dust scattering halo is most likely to be observed in the soft X-ray band from a bright X-ray source with a significant amount of intervening dust on the LOS.

Finally, the optical depth of the dust scattering τ_{sca} can be estimated from the observation using the fraction of the halo flux (F_{sca}) in the total observed flux (F_{obs}) (Mathis & Lee 1991; Predehl & Schmitt 1995; Valencic et al. 2009):

$$\tau_{sca} = -\ln(1 - F_{sca}/F_{obs}) \quad (6)$$

1.1.2 Variability of the Dust Scattering Halo

A dust scattering halo will show variability if the source is variable (Trümper & Schönfelder 1973; Mao, Ling & Zhang 2014; Heinz et al. 2015; Vasilopoulos & Petropoulou 2016). Time lags are caused by the arrival time differences of scattering photons with different light paths (Xu, McCray & Kelley 1986). Based on simple geometrical calculations, the

time-lag of a scattering photon relative to an un-scattered photon can be derived as:

$$\Delta t(x, \theta, \ell) = \frac{\ell}{c} \left(\sqrt{\left(\frac{x}{\cos\theta}\right)^2 - 2x + 1} + \frac{x}{\cos\theta} - 1 \right) \quad (7)$$

where c is the speed of light. Since dust scattering mainly happens at small angles, $\Delta t(x, \theta, \ell)$ can be simplified with small angle approximations:

$$\Delta t(x, \theta, \ell) \approx (1.21s) \frac{x}{1-x} \theta(\text{arcsec})^2 \ell(\text{kpc}) \quad (8)$$

According to these equations, for a viewing angle of 2 arcsec to a dust grain at $x = 0.9$ from a source at $\ell = 8 \text{ kpc}$, the lag is 350 s. These equations also show that time-lag increases with increasing viewing angles. For a fixed viewing angle, the time-lag increases as the scattering dust is located closer to the source. The lag scales linearly with the absolute source distance.

1.2 Dust Grain Models

The modelling of the dust scattering halo strongly depends on the dust grain model (e.g. Xiang et al. 2005; Valencic & Smith 2015). The classic dust grain model, proposed by Mathis, Rumpl & Nordsieck (1977) (hereafter: MRN77 dust grain), assumes a power law form for the size distribution of graphite and silicate grains with a size range of 0.005-0.25 μm . Later a smaller-size carbonaceous grain population (in the form of Polycyclic Aromatic Hydrocarbon - PAH) was added to explain the infrared (IR) emission features and to reproduce the 2175 \AA hump in the ultraviolet (UV) extinction curve (Li & Draine 2001; Weingartner & Draine 2001 - WD01 dust grains for different R_V (def. $A(V)/E(B-V)$) and carbon abundance; Draine 2003). Zubko, Dwek & Arendt (2004) (hereafter: ZDA04) considered various forms of carbon in the ISM, different abundances, and composite particles encompassing organic refractory material ($\text{C}_{25}\text{H}_{25}\text{O}_5\text{N}$), water ice (H_2O) and voids, and proposed 15 dust grain models which can all reproduce the $R_V = 3.1$ extinction curve. Moreover, Xiang et al. (2011) proposed a different dust grain model (XLNW) where the olivine grains in the BARE-GR-S dust grain model (BARE: without composite grains; GR: graphite; S: solar abundances) of ZDA04 is replaced with a new type of grain comprising a metallic iron core and a troilite/enstatite shell.

Due to the lack of knowledge about the dust grain composition and size distribution in different locations of the ISM, previous studies about the dust scattering halo often adopted various dust grain models and compared the results. It was reported that the BARE- dust grains in ZDA04 were better than COMP- dust grains (with composite grains) in terms of fitting the halo around sources whose LOS has a short overlap with the Galactic plane (Valencic et al. 2009; Xiang et al. 2011; Valencic & Smith 2015). For the GC LOS, Fritz et al. (2011) measured the extinction curve within 1-19 μm and fitted various dust grain models to the IR extinction curve, and reported the necessity of adding composite dust grains containing H_2O to the pure carbonaceous and silicate grains. Therefore, the COMP-AC-S dust grain model (AC: amorphous carbon) of ZDA04 was proposed, for the first time, to be the best dust grain model for the GC LOS.

1.3 AX J1745.6-2901

The GC region is filled with intense and complex diffuse emission, so a bright X-ray source is required to produce a detectable bright dust scattering halo that overwhelms the diffuse emission. Then the intrinsic halo profile can be extracted with enough signal-to-noise (S/N) to study the dust properties along the LOS. AX J1745.6-2901 is an ideal source because it can be very bright in X-rays. It is a low mass X-ray binary (LMXB) consisting of a neutron star and a K0V type companion star (Ponti et al. 2015b). It was first discovered during *ASCA* observations of the GC in 1993 as an eclipsing transient (Kennea & Skinner 1996; Maeda et al. 1996). *Chandra* observations provided precise measurements of its coordinate (17:45:35.64, -29:01:33.90 with a 1σ coordinate error of 0.32'')¹, which is only 1.45 arcmin from Sgr A* (Heinke et al. 2009). Therefore, it has been intensely observed by many X-ray observatories such as *Chandra*, *XMM-Newton*, *Swift* and *Suzaku* during GC monitoring campaigns (Hyodo et al. 2009; Degenaar et al. 2012; Ponti et al. 2015a).

Previous observations revealed periodic eclipsing and irregular dipping in the light curve of AX J1745.6-2901, indicating a high inclination angle of this binary system (Maeda et al. 1996) with an orbital period of $30063.74 \pm 0.14 \text{ s}$ and an eclipse duration of $\sim 1440 \text{ s}$ (Hyodo et al. 2009). Analysing more than 20 years of *XMM-Newton* and *ASCA* observations, Ponti et al. (2017a) measured a long-term decrease of the orbital period of $(4.03 \pm 0.32) \times 10^{-12} \text{ s/s}$ as well as significant jitter.

AX J1745.6-2901 is one of the brightest X-ray transient within a few arcmin of Sgr A* (Degenaar et al. 2012; Ponti et al. 2015b). Its luminosity was observed to reach a few percent Eddington luminosity assuming that the source is inside the GC region (i.e. within a few hundred pc from Sgr A*), which is within the typical luminosity range for ‘atoll’ sources (Gladstone, Done & Gierliński 2007). The $N_{\text{H},\text{abs}}$ towards AX J1745.6-2901 is measured to be $\sim 2.0 \times 10^{23} \text{ cm}^{-2}$ by Ponti et al. (2015b) from fitting the *XMM-Newton* spectra with the XSPEC PHABS absorption model and solar abundances of Anders & Grevesse (1989) (hereafter: AG89). But Paizis et al. (2015) fitted the *Chandra* spectra with the TBABS model and Interstellar Medium (ISM) abundances of Wilms, Allen & McCray (2000) (hereafter: WAM00), and reported $N_{\text{H},\text{abs}} = 3.4 \times 10^{23} \text{ cm}^{-2}$. The difference mainly arise from different assumptions about the abundances. In comparison, the $N_{\text{H},\text{abs}}$ measured from Sgr A*'s X-ray flares is $1.23 \times 10^{23} \text{ cm}^{-2}$ using the WABS model and Anders & Ebihara (1982) solar abundances (Porquet et al. 2008), or $1.5 \times 10^{23} \text{ cm}^{-2}$ using the TBNEW model and WAM00 ISM abundances (Nowak et al. 2012), which is also consistent with $1.6 \times 10^{23} \text{ cm}^{-2}$ measured most recently by Ponti et al.

¹ Degenaar & Wijnands (2009) report a similar *Chandra* position with a 1σ coordinate error of 0.6''. To improve the astrometric accuracy, we choose a long *Chandra* observation (ObsID:03665, 89.9 ks exposure time) where AX J1745.6-2901 is faint (so it is free from photon pile-up) but still have 285 counts for an accurate position measurement. We find three bright stars in the Tycho-2 catalog detected in the soft X-ray band within the *Chandra* field-of-view, and use them to perform a high-precision astrometry correction, which then allows us to reduce the coordinate error by a factor of 2.

Table 1. List of *XMM-Newton* and *Chandra* observations used for this paper. The list of observations during which AX J1745.6-2901 was in quiescence can be found in the Appendix B. ‘Exp’ is the total exposure time and ‘CL-Exp’ is the good exposure time after subtracting the periods of background flares. θ_{off} is the off-axis angle of AX J1745.6-2901. r_{pileup} is the inner circular radius of the PSF affected by photon pileup (Section 2). ‘PFW’: Primary Full Window mode. ‘HETG’: High Energy Transmission Grating.

<i>Chandra</i>							
Obs-ID	Obs-Date	Target	ACIS (Grating)	Exp (ks)	CL-Exp (ks)	θ_{off} (arcmin)	r_{pileup} (arcsec)
9169	2008-05-05	Sgr A*	ACIS-I (None)	27.6	27.6	1.80	2.0
9170	2008-05-06	Sgr A*	ACIS-I (None)	26.8	26.8	1.80	2.0
9171	2008-05-10	Sgr A*	ACIS-I (None)	27.7	27.7	1.80	2.1
9172	2008-05-11	Sgr A*	ACIS-I (None)	27.4	27.4	1.80	2.2
9174	2008-07-25	Sgr A*	ACIS-I (None)	28.8	28.8	1.32	2.0
9173	2008-07-26	Sgr A*	ACIS-I (None)	27.8	27.8	1.32	2.0
17857	2015-08-11	AX J1745-2901	ACIS-S (HETG)	117.2	111.1	0.34	2.4
<i>XMM-Newton</i>							
Obs-ID	Obs-Date	Target	pn-Filter	Exp (ks)	CL-Exp (ks)	θ_{off} (arcmin)	r_{pileup} (arcsec)
0402430701	2007-03-30	Sgr A*	PFW-Medium	34.2	21.3	2.57	20
0402430301	2007-04-01	Sgr A*	PFW-Medium	105.4	56.1	2.55	20
0402430401	2007-04-03	Sgr A*	PFW-Medium	105.8	38.1	2.60	20
0505670101	2008-03-23	Sgr A*	PFW-Medium	105.7	64.5	2.57	20
0724210201	2013-08-30	Sgr A*	PFW-Medium	58.5	53.4	1.94	20
0700980101	2013-09-10	Sgr A*	PFW-Medium	38.7	35.8	1.92	20
0724210501	2013-09-22	Sgr A*	PFW-Medium	43.9	32.2	1.89	20
0743630801	2015-04-01	Sgr A*	PFW-Medium	27.0	23.4	2.56	20

(2017b) using the same absorption model and abundances. These $N_{H,abs}$ measurements are all significantly smaller than that of AX J1745.6-2901. Moreover, Degenaar & Wijmans (2009) reported the brightest X-ray burst from this source reaching a luminosity of $1.3 \times 10^{38} \text{ erg s}^{-1}$ for a GC distance of 8 kpc, approaching the Eddington luminosity of a typical neutron star (e.g. Kuulkers et al. 2003). Therefore, it can be inferred that AX J1745.6-2901 is likely located inside or behind the GC.

This paper presents a detailed analysis of the dust scattering halo around AX J1745.6-2901. We organize the paper as follows. In Section 2 we describe the *Chandra* and *XMM-Newton* observations of AX J1745.6-2901 and data reduction, as well as the steps to extract the dust scattering halo around AX J1745.6-2901. Section 3 describes the point spread function (PSF) of *Chandra* and *XMM-Newton*. Section 4 presents a detailed modelling of the dust scattering halo in order to explore the properties of the dust grains along the LOS. The effect of the dust scattering on the observed spectra is studied in Section 5. Discussions about the dust properties towards AX J1745.6-2901 and more details about the halo profile are all included in Section 6. The final section summarizes our work and provides the main conclusions. Throughout this paper, we adopt a distance of 8 kpc to the GC (Reid et al. 2009; Genzel et al. 2010; Boehle et al. 2016), the uncertainty of which should not affect the results of this work.

2 OBSERVATION AND DATA REDUCTION

AX J1745.6-2901 is one of the brightest transients within a few arcmin from Sgr A*. Being so close, it is in the field of view (FoV) of many observations targeting Sgr A*. In this

work we used archival and new *Chandra* and *XMM-Newton* observations. *Chandra* has high spatial resolution and small PSF, which can produce high resolution halo radial profiles; while *XMM-Newton* provides a large effective area and FoV, which leads to a significant detection of the dust scattering halo up to 10 arcmin. Therefore, combining the data from these two satellites is ideal to study the halo profile.

2.1 *Chandra* Observations

The *Chandra* data website was used to search for observations of AX J1745.6-2901². We selected observations where AX J1745.6-2901 was observed at ≤ 3 arcmin off-axis angle³ because of significant PSF degradation at larger off-axis angles. Among these observations, we chose those where AX J1745.6-2901 was bright enough to show a clear halo above 2 keV ($L_{2-4\text{keV}} \geq 10^{-12} \text{ erg cm}^{-2} \text{ s}^{-1}$). Then we dropped observations with short exposures ($\lesssim 5$ ks), as these observations were too short to constrain the halo profile. Our selection resulted in 6 ACIS-I observations targeting Sgr A* and 1 recent ACIS-S observation targeting AX J1745.6-2901 (Table 1).

In order to measure the detector background and diffuse emission underneath AX J1745.6-2901 we made use of observations where AX J1745.6-2901 was at ≤ 3 arcmin off-axis angle but was in the quiescent state. To ensure that the non-detection of AX J1745.6-2901 was due to its intrinsic weakness rather than the short exposure time, we chose observations with more than 15 ks exposure and found 17 such

² <http://cda.harvard.edu/chaser>

³ the angle between the HRMA optical axis and the source position

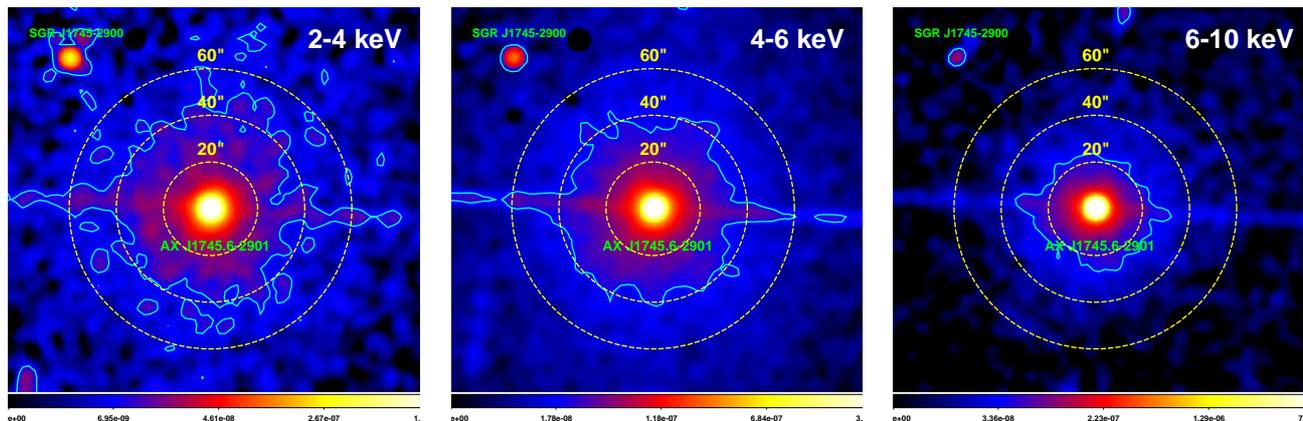


Figure 2. *Chandra* ACIS-S (ObsID:17857) flux images of AX J1745.6-2901 in the 2-4, 4-6 and 6-10 keV bands, showing the shrink of the dust scattering halo towards higher energies. The background based on 12 ACIS-S observations with AX J1745.6-2901 in quiescence has been subtracted (see Section 2). In every panel, the cyan contour indicates the location where the surface brightness is only 0.4% of the value at 2.5 arcsec from the core (inside which more than 1% pile-up is expected). The dashed yellow circles indicate regions with 20, 40 and 60 arcsec radius centred on AX J1745.6-2901. Note that the horizontal structure across AX J1745.6-2901 is due to the readout streak.

observations in the ACIS-I mode (with 713.3 ks clean exposure in total) and 12 observations in the ACIS-S mode (with 702.7 ks clean exposure in total) (Table B1). The long accumulated exposure time allows an accurate determination of both the detector background and the diffuse emission in this region. Moreover, we checked that AX J1745.6-2901 was located in the same type of CCD chip, i.e. Front Illuminated (FI) chip in ACIS-I and Back Illuminated (BI) chip S3 in ACIS-S. This is to ensure the consistency of the detector background among all observations in ACIS-I and ACIS-S, separately.

The *Chandra* software CIAO (v4.8.1) and *Chandra* Calibration Database (CALDB, v4.7.2) were used to analyze all the *Chandra* observations following the standard data reduction steps (see e.g. Clavel et al. 2013 for detailed descriptions). Firstly, the CHANDRA_REPRO script was used to reprocess all the ACIS data. To ensure the astrometry consistency among all the observations, we followed the standard thread to perform the astrometry correction by taking ObsID:14468 (146 ks exposure time) as the reference observation. For every observation, we used the WAVDETECT script to perform the point source detection and all the point sources were excluded before extracting the background light curve, which was then used as the input for the DEFLARE script to remove the background flares from the event file. AX J1745.6-2901 was piled up in all the selected observations, so we used the PILEUP_MAP script in CIAO (v4.8.1) to estimate the size of the 1% pile-up region around AX J1745.6-2901 (Table 1). All data within this pile-up radius were removed from our study. Furthermore, we excluded the data when AX J1745.6-2901 was in the eclipsing phase (detected in 5 observations, ObsID: 9169, 9170, 9173, 9174, 17857) and 300 s after the eclipse egress time. This is because the halo profile was varying due to a delayed response to the eclipsed signal which needs to be studied independently (Jin et al. in prep.).

We used the MERGE_OBS script and REPROJECT_IMAGE_GRID script in the *Chandra* software to create images and exposure maps (including the vignetting effect), re-project the images and combine images from all

observations. The DMIMGCALC script was used to perform background subtraction. Fig.2 shows the final flux images of AX J1745.6-2901 in 2-4, 4-6 and 6-10 keV bands from the ACIS-S observation (ObsID: 17857). It is clear that AX J1745.6-2901 appears much more extended than a point-like source, and the extension has a clear energy-dependence, consistent with the prediction of the dust scattering theory.

To extract the radial profile of AX J1745.6-2901, we first excluded all other point sources including transients in every observations, then an elliptical region of $2' \times 6'$ was used to exclude the region around Sgr A^{*} where the diffuse emission is strongest (e.g. Heard & Warwick 2013a,b; Ponti et al. 2015a). All artificial features such as the readout streak and the HETG arms (in ACIS-S observations) were also masked out by employing box regions of 7 arcsec wide from the edge of the FoV down to 2 arcsec from AX J1745.6-2901. We defined a set of annulus regions, and used the FUNTOOLS in *ds9* to extract the photon counts and exposure in every annulus region from the image and exposure map of every observation⁴. The combined radial profile was obtained by adding counts from all observations and dividing it by the total exposure time in every annulus region. The Poisson counting errors are small because of the accumulated number of counts from all the observations, but there are additional small dispersions in the shape of radial profiles from different observations, which are likely due to the weak dependence of PSF on the small variation of off-axis angles (Table 1), and/or the fluctuation of residual background contamination from e.g. the extended PSF wing of bright sources and diffuse emission in the FoV. Therefore, we used the MPFITEXY routine (Williams, Bureau & Cappellari 2010, Markwardt 2009) to calculate error bars for the combined radial profile, which takes into account both the Poisson error and the small dispersion between different observations (see Tremaine et al. 2002 for the statistical method).

⁴ also see http://cxc.harvard.edu/ciao/threads/radial_profile

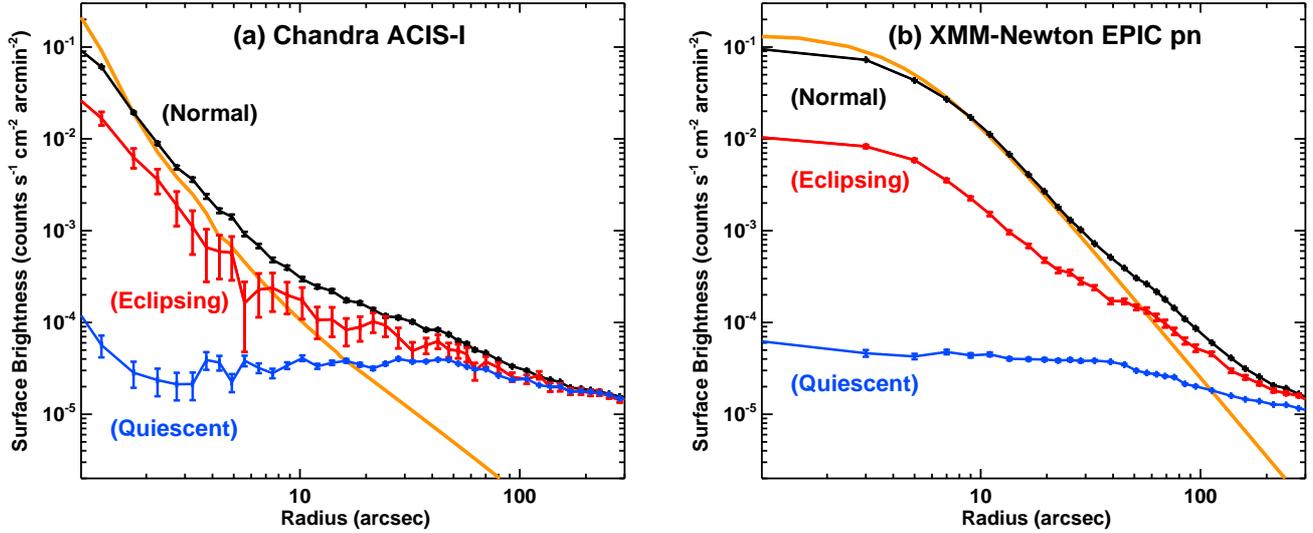


Figure 3. Radial profile of AX J1745.6-2901 in the normal (black) and eclipsing phases (red) in the 4-6 keV band. Panel-a shows the combined radial profile from 6 *Chandra* ACIS-I observations in Table 1. The quiescent profile (blue) is based on a combination of 17 *Chandra* ACIS-I observations where AX J1745.6-2901 was in quiescence. The slight rising of the radial profile within 2 arcsec is due to the weak emission from AX J1745.6-2901 when several observations were combined. Panel-b shows the combined radial profile from 8 *XMM-Newton* observations in Table 1. The emission inside the eclipsing phase comes from the dust scattering. In both panels, the solid orange profile shows the instrumental PSF at 5 keV. The difference between the halo radial profiles of *Chandra* and *XMM-Newton* is caused by their different PSF.

The above procedure was repeated to create the radial profile of AX J1745.6-2901 as well as the radial profile of the underlying diffuse emission and detector background when AX J1745.6-2901 was in quiescence (Fig. 3a, also see Fig. A1 for the energy-dependence of the radial profile). The intrinsic radial profile of AX J1745.6-2901 was derived by subtracting the combined radial profile observed when AX J1745.6-2901 was in quiescence from the one when AX J1745.6-2901 was bright, with the standard error propagations. Note that we did not combine the radial profiles from ACIS-I and ACIS-S observations, because the ACIS-S observation was in the sub-array mode. AX J1745.6-2901 was weakly detected in the combined background radial profile within 2 arcsec due to the enhanced S/N. We also repeated the above procedures to extract the mean radial profile of AX J1745.6-2901 during its eclipsing phase (Fig. 3a red points), which allows us to observe the shape of the dust scattering halo without the PSF of AX J1745.6-2901 itself. However, the light curve of AX J1745.6-2901 in Ponti et al. (2015) shows that the halo is variable during the eclipsing phase due to the time-lag effect (see Section 1.1.2), so in this work we do not model the halo profile in the eclipsing phase.

2.2 *XMM-Newton* Observations

We used the *XMM-Newton* Science Archive (XSA) to search for *XMM-Newton* (Jansen et al. 2001) observations of AX J1745.6-2901 with similar selection criteria as for the *Chandra* observations. We found 9 observations where AX J1745.6-2901 was observed at ≤ 3 arcmin off-axis angle where the PSF was not degraded significantly, and the source was bright enough for a significant halo detection above 2 keV ($L_{2-4\text{keV}} \geq 10^{-11}$ erg cm $^{-2}$ s $^{-1}$). All these observations

were in the same mode and targeted Sgr A * so that AX J1745.6-2901 was at the same off-axis angle. In order to determine the underlying diffuse emission and detector background, we selected another 14 observations targeting Sgr A * with ≥ 15 ks exposure time but with no detection of AX J1745.6-2901 (so it was in quiescence). The accumulated exposure time is 358.3 ks, which can ensure an accurate measurement of the diffuse emission and detector background.

For each of these observations, we used the *XMM-Newton* Science Analysis System (SAS v15.0.0) to process the data and applied the most recent calibrations. We only focused on the EPIC-pn data as pn provides the highest count rate among all three EPIC cameras. The EPCHAIN task was used to reprocess the data. Then we excluded all high background periods detected from the 7-15 keV background light curve, and followed the method in Urban et al. (2011) (also see Leccardi & Molendi 2008) to check and ensure no significant soft proton residuals. We excluded one observation of AX J1745.6-2901 (ObsID: 0504940201), as the clean exposure time was only 6 ks. The remaining 8 observations of AX J1745.6-2901 are listed in Table 1. The 14 observations that caught AX J1745.6-2901 in quiescence are listed in Table B1, with a total of 332.1 ks clean exposure time.

We used the EPATPLOT task to determine the pile-up region, and excluded data within 20 arcsec of AX J1745.6-2901 to safely avoid the pile-up effect. We double-checked this radius by plotting the radial dependence the ratios of single-events and double events in all events for the observation (OBSID: 0743630801) where AX J1745.6-2901 had the highest flux (see Costantini, Freyberg & Predehl 2005). For the observed spectrum of AX J1745.6-2901, the pile-up effect is most obvious above 7 keV, and so we used events within 7-10 keV to perform this test. Fig. 4 shows that, within 20 arc-

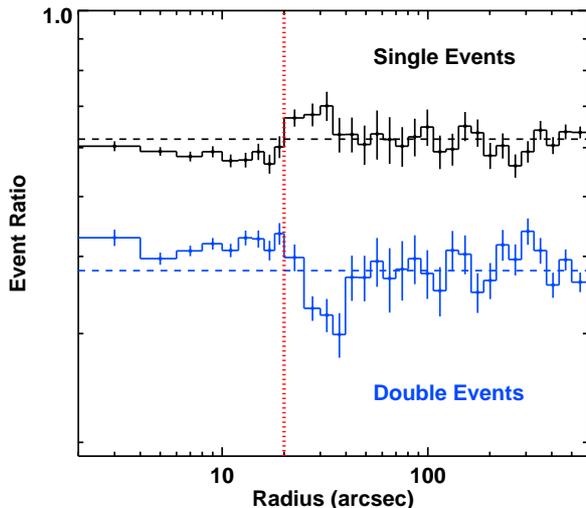


Figure 4. Radius dependence of the ratio between single-events and all events (black), and between double-events and all events (blue) for 7-10 keV in *XMM-Newton* (OBSID: 0743630801). The two horizontal dashed lines indicate the expectation of no pile-up effect. The vertical red dotted line indicates 20 arcsec radius within which the pile-up effect is seen.

sec, the pile-up effect causes single-event’s ratio to be lower than expectation, and double-event’s ratio to be higher than expectation.

For every observation, the 2-10 keV light curve of AX J1745.6-2901 was extracted with the `EVSELECT` task. The light curve was used to search for eclipsing periods and the latter were excluded from the event file. We also carefully excluded some dipping periods. Then we used the Extended Source Analysis Software (ESAS; Snowden & Kuntz 2011) in SAS (v15.0.0) to create and process images from the event file. We performed point source detection using the `CHEESE` task. Then the `PN_SPECTRA` task was used to create the normal and Out-of-Time (OoT) processing images and exposure maps (including the vignetting effect). The OoT image was multiplied by 0.063 and subtracted from the normal image using the `FARITH` task (FTOOLS v6.19). The `PN_BACK` task was used to create the particle background image and then subtracted this from the normal image.

All point sources except AX J1745.6-2901 were excluded in every observation before extracting the radial profile. An elliptical region of $12' \times 16'$ was used to exclude the strong diffuse emission around Sgr A*. We note that there were still residual OoT features in the final image of some observations, which was because AX J1745.6-2901 was piled up so that the correction was not accurate. Therefore, we visually checked all images and, when necessary, we used a rectangular region to mask out the OoT residual from the CCD edge down to 20 arcsec from AX J1745.6-2901. We chose the width of the rectangle up to 250 arcsec depending on the strength of the OoT residual feature. Some weak features near the edge of FoV due to stray light were also visually masked out. The exposure map created by the `PN_BACK` task was not multiplied by the effective area, so we used the SAS task `PSFGEN` to produce the unvignetted Ancillary Response File (ARF), from which we obtained the spectral shape weighted effective area in every energy band. Finally,

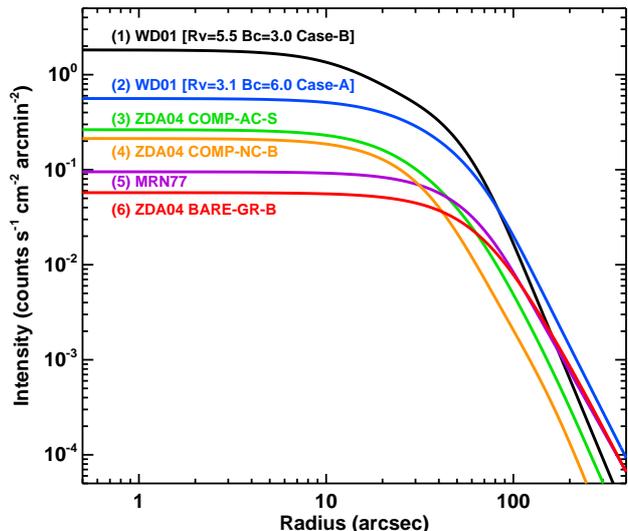


Figure 5. Radial profile comparison for some representative dust grain models. These radial profile models are for a source at 8.15 kpc with a flux of unity and $E = 5$ keV. The light passes through a single smoothly distributed dust layer between 2.5 and 3.5 kpc with a total $N_{H,sca} = 1.5 \times 10^{23} \text{ cm}^{-2}$. No instrumental PSF was convolved. These radial profiles reveal the typical halo shape diversity among the dust grain models in MRN77, ZDA04 and WD01.

we obtain the combined radial profiles from all the observations and performed background subtraction using the same method as for *Chandra* (see Section 2.1 and Fig.3b).

3 PSF OF *Chandra* AND *XMM-Newton*

The image of AX J1745.6-2901 consists of a central point source plus a dust scattering halo, convolved with a 2-dimensional (2D) instrumental PSF. This 2D convolution can significantly smooth the halo profile especially when the PSF is broad such as in *XMM-Newton*.

3.1 PSF of *Chandra* ACIS

The on-axis half energy width of the PSF in the *Chandra* High-Resolution Mirror Assembly (HRMA)/Advanced CCD Imaging Spectrometer (ACIS) is 0.4-0.7 arcsec within 0.3-10 keV (*Chandra* Proposers’ Observatory Guide v18.0⁵), indicating that the halo profile detected by *Chandra* ACIS is heavily smeared by the PSF convolution (Smith, Edgar & Shafer 2002; Smith 2008). Therefore, an accurate PSF is mainly required to disentangle the point source from the dust scattering halo. We used the web-based CHART software to perform the ray-trace simulation which takes parameters directly from real observations (ObsID:09174 for ACIS-I and ObsID:17857 for ACIS-S). The simulation was done for 50 iterations in order to obtain enough counts to extract a high S/N PSF⁶. Then the MARX (v5.3) software

⁵ http://cxc.harvard.edu/proposer/POG/html/chap6.html#tth_sEc6.6

⁶ <http://cxc.harvard.edu/ciao/PSFs/chart2/runchart.html>

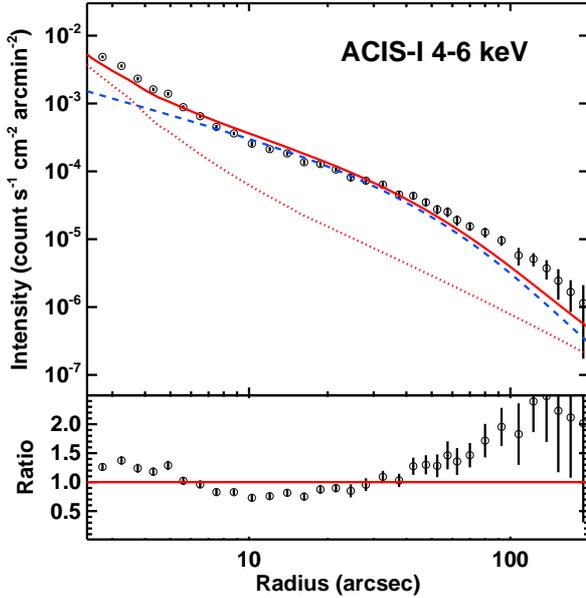


Figure 6. Radial profile observed by *Chandra* ACIS-I in the 4-6 keV band. The model comprises the ACIS-I PSF (red dotted curve) and a single dust layer smoothly distributed from Earth to AX J1745.6-2901 (blue dashed curve), which used the COMPAC-S dust grain model. The lower panel shows the residuals.

and DMMERGE script were used to create a combined pseudo event file from all the 50 iterations. Finally, a high S/N PSF radial profile was extracted from the combined event file. This procedure was repeated for ACIS-I and ACIS-S at 3.3, 5.0 and 7.0 keV separately, because AX J1745.6-2901 was at slightly different off-axis angles during these two observation modes. Note that in the simulations for ACIS-S, there was no typical ‘X’ pattern from the HETG arms, thus the simulated PSF profile is not affected by the gratings.

Meanwhile, Gaetz (2010) performed a detailed analysis of the PSF wing of HRMA using a deep calibration observation of the low mass X-ray binary Her X-1 (ObsID: 3662), which is bright and has low $N_{H,sca}$ in the LOS. This observation was also used in some other works to measure the PSF profile (e.g. Smith, Edgar & Shafer 2002; Xiang, Lee & Nowak 2007; Xiang et al. 2011). Because Her X-1 was heavily piled up in the PSF core, we only used the best-fit King profile in Gaetz (2010) to produce the PSF from 10 arcsec to 500 arcsec. Her X-1 was at a 0.75 arcmin off-axis angle during this observation, while AX J1745.6-2901 was on-axis in the ACIS-S mode and at 1.45 arcmin off-axis in the ACIS-I mode, but the small difference of the off-axis angle should not affect the PSF wing significantly (see the *Chandra* website⁷ and Gaetz 2010).

Since the PSF from the CHART simulation for $E \geq 2$ keV may under-predict the flux in the PSF wing², while the PSF from Gaetz (2010) was affected by severe pile-up within 10 arcsec, we decided to use the core of the CHART PSF and the wing of the Gaetz (2010) PSF by matching them within 10-15 arcsec, so that an accurate PSF profile was derived. This final PSF profile was renormalized and adopted in our halo profile modelling.

⁷ <http://cxc.harvard.edu/ciao/PSFs/chart2/caveats.html>

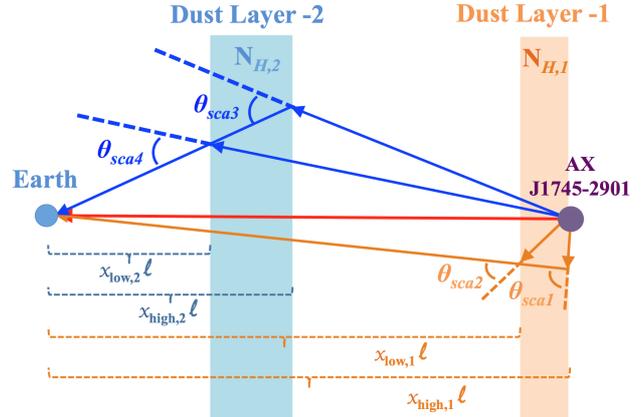


Figure 7. The two dust-layer scenario for the radial profile fitting in Fig.9. $x_{low,1,2}$ and $x_{high,1,2}$ are the fractional distances of the two dust layers from Earth. $N_{H,1}$ and $N_{H,2}$ are the $N_{H,sca}$ in Layer-1 and Layer-2, separately. l is the absolute distance of AX J1745.6-2901 from Earth. $\theta_{sca1,2}$ and $\theta_{sca3,4}$ show the change of dust scattering angle at a specific viewing angle within the two dust layers.

3.2 PSF of *XMM-Newton* EPIC-pn

The on-axis half energy width of EPIC-pn is 16.6 arcsec (*XMM-Newton* Users Handbook⁸), implying a much bigger impact on the halo profile than in *Chandra*. Ghizzardi (2002) studied 110 *XMM-Newton* observations to calibrate the EPIC-pn PSF and provided equations (based on the best-fit King profile) to calculate the 1-dimensional (1D) PSF profile (also see Costantini, Freyberg & Predehl 2005). Meanwhile, the 1D and 2D PSF profiles were modelled separately with a 1D King profile and a 2D ELLBETA model by the *XMM-Newton* calibration team (Read et al. 2011), with the best-fit parameters given in the latest Current Calibration Files (CCF). However, it was noticed that the 2D ELLBETA model requires further calibration in the PSF wing (Schartel, private communication), while the CCF files only provide 1D PSF parameters for a discrete set of off-axis angles and energies, so we decided to adopt the Ghizzardi (2002) PSF which should be accurate enough for our study. Note that the off-axis angle of AX J1745.6-2901 is 1.454-1.678 arcmin in all the EPIC-pn observations, and the PSF changes very little within these small off-axis angles.

4 MODELLING THE DUST SCATTERING HALO OF AX J1745.6-2901

4.1 Radial Profile Fitting of the Dust Scattering Halo

The shape of the dust scattering halo depends on the dust grain model (see Section 1.2). In Fig.5, we compare the halo profiles using some representative dust grain models, including the classic MRN77 dust grain model, two dust grain models from WD01 with [$R_V = 3.1$, $B_c = 6.0$, *Case - A*] (adopted by Valencic & Smith 2015 as the standard WD01

⁸ <http://www.cosmos.esa.int/web/xmm-newton/documentation>

dust grain model) and $[R_V = 5.5, B_c = 3.0, \text{Case - B}]$ (highest intensity in the halo core) where B_c ($\times 10^5$) is the carbon abundance per hydrogen nucleus, and three ZDA04 dust grain models. In this paper, we tried all dust grain models in MRN77, WD01, ZDA04 and XLNW to do the halo profile modelling and compared the results.

A reasonable dust scattering model should reproduce not only the halo radial profile at a specific energy, but also the energy dependence of the halo. Besides, the halo profiles from *Chandra* and *XMM-Newton* should contain the same intrinsic halo shape except for the different instrumental PSF. Therefore, we decided to fit the halo profiles from *Chandra* ACIS-I and *XMM-Newton* in the 2-4, 4-6 and 6-10 keV bands, simultaneously, i.e. six radial profiles in total (see Fig.9). The dust scattering calculation is computationally intensive (Smith, Valencic, Corrales 2016; Corrales & Paerels 2015). In order to increase the computing efficiency, we calculated the spectral weighted effective energy within the 2-4, 4-6 and 6-10 keV bands as 3.3, 5.0 and 7.0 keV and fixed the photon energy at these values in the model. This should not affect the accuracy because the halo shape is not sensitive to a small change of the effective energy. Assuming a rotational symmetry, we performed 2D convolution for the instrumental PSF in order to fit the halo profile observed by various instruments. All the fittings were performed with the SHERPA fitting engine provided in the CIAO (v4.8.2) software⁹.

4.1.1 One Dust-Layer Model

First, we assumed a point source and one foreground dust layer of smooth distribution (as shown in Fig.1). The lower and upper boundary of the dust layer and the $N_{H,sca}$ were all free parameters. The incident flux to the dust layer was linked to the flux of the point source, which was also a free parameter because the fraction of scattering flux in the total observed flux was unknown. Only single scattering was considered (see Section 4.2 for the discussion of multiple scattering). The absolute distance of AX J1745.6-2901 does not affect the halo profile fitting because all distances in the model are fractional. Fig.6 shows an example of the one-layer model fit to the 4-6 keV radial profile of AX J1745.6-2901 observed by *Chandra* ACIS-I, using the COMP-AC-S dust grain model. The best-fit $\chi^2 = 1329$ for 178 degrees of freedom (dof) for all six radial profiles. The fitting also required the dust layer to extend from Earth to the source with $N_{H,sca} = 1.7 \times 10^{23} \text{ cm}^{-2}$. It is clear that this model is too simple to reproduce the halo profile around AX J1745.6-2901. We tried other dust grain models but none could produce a reasonably good fit with reduced $\chi^2_v < 2$. The residuals in Fig.6 clearly indicate that the curvature of the radial profile is more complex than the dust scattering halo from one smooth dust layer regardless of the dust grain model used.

4.1.2 Two Dust-Layers without Wing

Then we adopted a two-layer model, i.e. Layer-1 + Layer-2. Fig.7 shows the setup and free parameters of the two dust

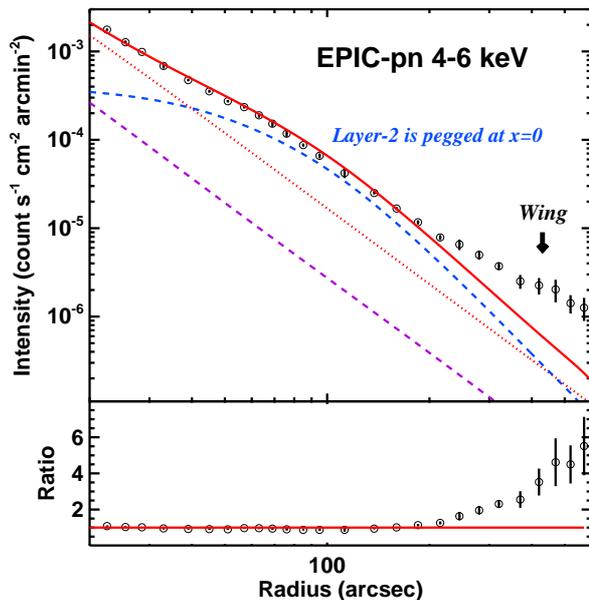


Figure 8. Radial profile observed by *XMM-Newton* EPIC-pn in the 4-6 keV band, overplotted with the EPIC-pn PSF (red dotted curve), dust Layer-1 (magenta dashed curve), Layer-2 whose fractional distance is pegged at 0 and fractional width pegged at the lower limit of 0.01 (blue dashed curve), and the total model (red solid curve). A significant excess flux exists outside 200 arcsec which cannot be fitted by any dust model (black arrow).

layers, including the $N_{H,sca}$, lower and upper boundaries of each layer. Layer-1 was closer to AX J1745.6-2901 and layer-2 was closer to Earth, and so the halo from Layer-2 was more extended than that from Layer-1. The parameters of these two layers were independent from each other. As a first order consideration, we assumed the same dust grain model for the two layers. This two-layer model improved the halo fitting significantly. For example, we found $\chi^2 = 664$ for 175 dof in the case of COMP-AC-S dust grain model. Fig.8 shows that this two-layer model cannot fit the halo wing ≥ 200 arcsec observed by *XMM-Newton*. The discrepancy between the the wing and the best-fit model increases towards large radii up to a factor of 6 at 600 arcsec. The total flux in the detected halo wing (200-600 arcsec) is a factor of 6.0 ± 1.0 higher than in the model. We emphasise that the fractional distance of layer-2 in Fig.8 (blue dashed curve) has already pegged at 0 with the fractional width reaching the lower limit of 0.01, indicating that this wing component cannot be solely explained by the distance of the layer¹⁰. We find that this wing component is related to the X-ray emission from AX J1745.6-2901 and rises from some dust grains with relatively small sizes (see Section 6.3), while scattering from large dust grains still dominates at small radii.

4.1.3 Two Dust-Layers with Wing

Because of the halo wing component detected, we restricted our fitting to the radial profile less than 300 arcsec and added

¹⁰ we chose 0.01 to avoid numerical problem in the halo calculation. Changing it to an even smaller value causes negligible difference to the halo profile.

⁹ <http://cxc.cfa.harvard.edu/ciao>

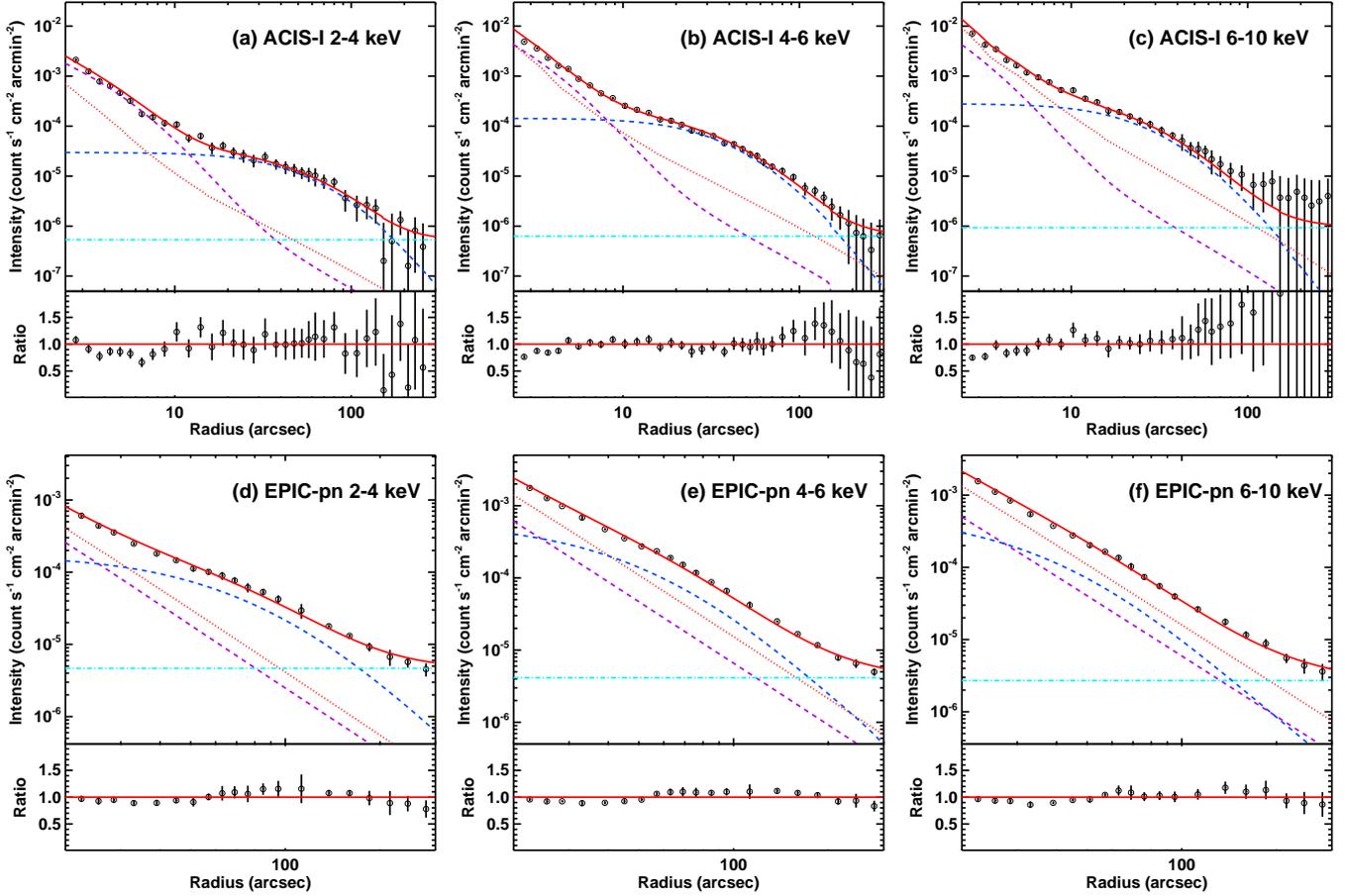


Figure 9. The radial profile fitting for *Chandra* ACIS and *XMM-Newton* EPIC pn in the 2-4, 4-6 and 6-10 keV bands using two dust layers with the COMP-AC-S dust grain model (ZDA04). The background is already subtracted (see Section 2.1). The red solid curve shows the total model. The red dotted curve shows the instrumental PSF. The two dashed curves show the dust scattering from the two dust layers (magenta: Layer-1; blue: Layer-2). The cyan line is a flat component to account for the halo wing dominating outside ~ 200 arcsec, which cannot be reproduced directly by any dust grain model (see Section 6.3). The best-fit parameters can be found in Table 2. The radial profiles before background subtraction can be found in Fig. A1.

a free constant parameter to the model to account for the wing above 200 arcsec. Since this wing component should exist in both *Chandra* and *XMM-Newton* observations, we assumed that it had the same scaling factor to the source flux in both instruments. Therefore, the final two-layer model was Layer-1 + Layer-2 + Wing.

Based on this new two-layer model, we fitted the halo profile with various dust grain models. Fig. 9 shows the halo fitting for the COMP-AC-S dust grain model. 18 out of 19 dust grain models produce $\chi^2_\nu < 2$, indicating a reasonably good fit. The best $\chi^2_\nu = 1.29$ was found by the BARE-GR-B dust grain model (Table 2). The best-fit parameters and their statistical error are listed in Table 2 and 3. It is clear that the statistical errors are much smaller than the dispersion of best-fit values among all the dust grain models, indicating significant systematic uncertainty due to the assumptions made in the dust grain model. Other systematics include the assumptions about the uniform dust distribution inside every layer, the same dust grain model for all the foreground dust, Gaussian approximation for the form factor, calibration systematics especially in the PSF wing.

Among all the 19 dust grain models, the highest

$N_{H,sca} = (29.7 \pm 1.3) \times 10^{22} \text{ cm}^{-2}$ was found for the COMP-AC-B dust grain model, while the lowest $N_{H,sca} = (5.0 \pm 0.2) \times 10^{22} \text{ cm}^{-2}$ was found for the WD01-B dust grain model. This large $N_{H,sca}$ dispersion was also found in previous works (e.g. Smith, Edgar & Shafer 2002; Valencic & Smith 2008; Valencic et al. 2009; Xiang et al. 2011), which is mainly due to different assumptions about the dust-to-gas ratio, abundances and size distribution made in different dust grain models. However, the mean best-fit $N_{H,sca} = 16.9 \times 10^{22} \text{ cm}^{-2}$ and the COMP-AC-S best-fit $N_{H,sca} = 19.2 \times 10^{22} \text{ cm}^{-2}$ are still roughly consistent with $N_{H,abs} \sim 2.0 \times 10^{23} \text{ cm}^{-2}$ using the AG89 solar abundances (Ponti et al. 2015b).

For all dust grain models, the fitting requires dust Layer-1 to be close to AX J1745.6-2901. Depending on the dust grain model used, the fractional distance of Layer-1 is $\lesssim 0.1$ from the source and containing (19-34)% of the total dust along the LOS. Layer-2 is more extended and contains (66-81)% of the intervening dust, with its lower boundary always pegged at 0, indicating that part of this layer is very close to Earth. From Table 2, we see that the upper boundary of layer-2 ($x_{high,2}$) also depends on the dust

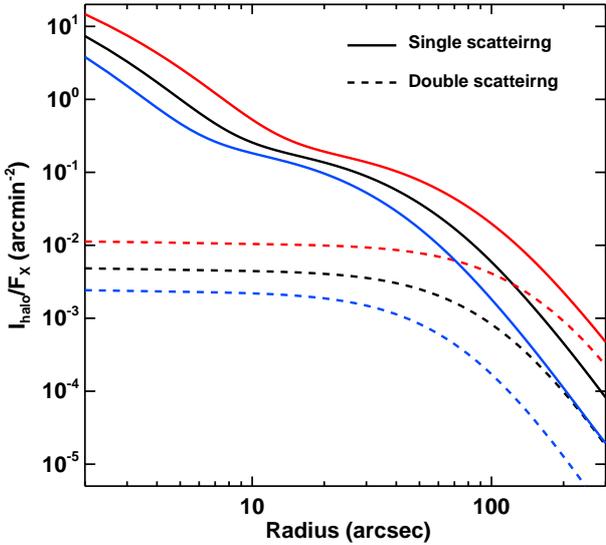


Figure 10. Comparison of the radial profiles from the single scattering (solid lines) and double scattering (dash lines) for 3.3 keV (red, $\tau = 0.82$), 5.0 keV (black, $\tau = 0.55$) and 7.0 keV (blue, $\tau = 0.42$), based on the two-layer model with COMP-AC-S dust grains (Table 3).

grain model, ranging from 0.01 to 0.90. The lowest values are found in the COMP-NC-B ($x_{high,2} = 0.01$) and COMP-NC-GF ($x_{high,2} = 0.11$) dust grain models, but their χ^2 are also the largest. The highest value is found in the BARE-GR-B model ($x_{high,2} = 0.90$). We notice that the boundary of the two layers mainly depend on the size distribution of dust grains in each model. A higher fraction of small dust grains will produce a more extended halo profile, and so $x_{high,2}$ ($x_{low,1}$) will increase (decrease) to enhance the halo intensity at small radii in order to fit the observed halo profile. The mean $x_{high,2}$ is 0.64, which is about half way towards AX J1745.6-2901. The separation between these two layers is quite significant compared to the statistical error, although the value ranges from 0.09 to 0.88 among all the dust grain models (Table 3). Adding more dust layers can improve the fitting statistics but also introduce more parameter degeneracies. To better resolve these two major dust layers, it is necessary to consider the timing properties of the halo due to the variability of AX J1745.6-2901 (e.g. the eclipsing signal), which will be presented in a separate work (Jin et al. in prep.).

4.2 Optical Depth and Multiple Scattering

Although our two-layer single scattering model can provide a reasonably good fit to the halo profile of AX J1745.6-2901 in all three energy bands, it is still necessary to check if multiple scattering is important. The inclusion of higher order scattering will produce a brighter and more extended halo than single scattering (Mathis & Lee 1991). Multiple scattering starts to dominate when the scattering optical depth (τ_{sca}) is significantly bigger than unity (e.g. Mathis & Lee 1991; Xiang, Lee & Nowak 2007). Since AX J1745.6-2901 is an eclipsing source, the total flux and halo flux can be measured directly from inside and outside the eclipsing phase, then τ_{sca} can be calculated using Eq.6.

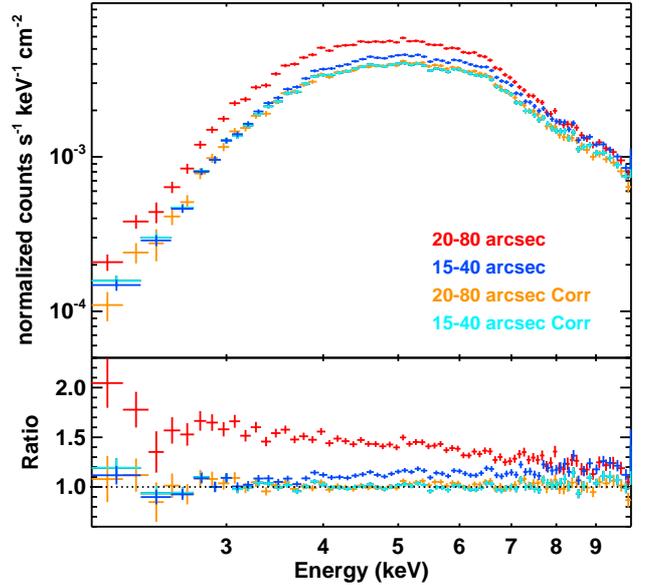


Figure 11. *XMM-Newton* EPIC-pn spectra of AX J1745.6-2901 from different extraction regions (ObsID: 0724210201). The red spectrum is from a 20-80 arcsec annulus region, which is both significantly brighter and softer than the blue spectrum from a 15-40 arcsec annulus region. Both spectra have been corrected for the same EPIC-pn PSF. The spectral difference is essentially caused by the dust scattering halo, which can be corrected for using our dust scattering model (i.e. the consistent orange and cyan spectra, see Section 5.2). Lower panel shows the ratio of the four spectra to the best-fit absorbed power law model to the orange spectrum.

In order to measure the flux accurately, we chose a different *XMM-Newton* observation (ObsID: 0723410301) where AX J1745.6-2901 was not piled up. The same data reduction procedures were applied to this observation. Source light curves were extracted from a circular region of 200 arcsec to include the dust scattering halo in the 2-4, 4-6 and 6-10 keV bands. The remaining halo flux outside 200 arcsec should be negligible (see Fig.9). From background subtracted light curves, F_{sca}/F_{obs} was found to be 0.57, 0.32, 0.16 for the three energy bands, implying $\tau_{sca} = 0.84, 0.39$ and 0.17. We also calculated τ_{sca} from every best-fit two-layer model, which are consistently less than 1 (Table 3). Similar results can be found using common relations between τ_{sca} , $N_{H,abs}$ and the V band extinction A_V (see Section 6.2). These scattering optical depths indicate that single scattering should dominate in the case of AX J1745.6-2901. As a further check, we followed the equations in Mathis & Lee (1991) and Xiang, Lee & Nowak (2007) to calculate the halo profile from double scattering using the best-fit COMP-AC-S two-layer model. Fig.10 shows that the intensity of the double scattering component is indeed dominated by the single scattering component. Therefore, we conclude that it is sufficient to consider only single scattering in this work.

Table 2. Best-fit parameters for all the 19 dust grain models using the *Two Dust-Layers with Wing* model. $x_{high/low,1/2}$ and $N_{H,1/2}$ are the fractional distances and $N_{H,sca}$ of the two dust layers, as shown in Fig.7. The COMP-AC-S dust grain model (indicated by †) was recommended for the GC direction by Fritz et al. (2011). l : the parameter pegs at the lower limit. χ^2/dof (ALL) is for the simultaneous fitting to all the six radial profiles in Fig.9, while χ^2/dof (CXO) is the fitting statistics of the three *Chandra* radial profiles. Error bars are calculated for the 90% confidence range. We emphasise that systematic uncertainties are more important, such as the scatter of best-fit parameters between different dust grain models.

No.	Dust Model	$x_{low,2}$	$x_{high,2}$	$x_{low,1}$	$x_{high,1}$	$N_{H,2}$ (10^{22} cm^{-2})	$N_{H,1}$ (10^{22} cm^{-2})	χ^2/dof (ALL)	χ^2/dof (CXO)
0	MRN77	0^l	0.8949 $^{+0.0138}_{-0.0716}$	0.9783 $^{+0.0010}_{-0.0021}$	0.9976 $^{+0.0008}_{-0.0004}$	13.2 $^{+1.1}_{-1.1}$	6.8 $^{+0.9}_{-0.9}$	199/131	97/83
1	BARE-GR-S	0^l	0.8092 $^{+0.0082}_{-0.0643}$	0.9553 $^{+0.0022}_{-0.0062}$	0.9975 $^{+0.0017}_{-0.0030}$	9.5 $^{+0.5}_{-0.5}$	2.3 $^{+0.3}_{-0.3}$	204/131	126/83
2	BARE-GR-FG	0^l	0.8314 $^{+0.0131}_{-0.0472}$	0.9611 $^{+0.0034}_{-0.0084}$	0.9989 $^{+0.0011}_{-0.0006}$	10.5 $^{+0.8}_{-0.8}$	3.0 $^{+0.5}_{-0.5}$	198/131	107/83
3	BARE-GR-B	0^l	0.9010 $^{+0.0046}_{-0.0421}$	0.9760 $^{+0.0024}_{-0.0087}$	0.9984 $^{+0.0011}_{-0.0003}$	15.1 $^{+0.9}_{-0.9}$	4.8 $^{+0.5}_{-0.5}$	169/131	94/83
4	COMP-GR-S	0^l	0.6242 $^{+0.0035}_{-0.0423}$	0.9403 $^{+0.0102}_{-0.0026}$	0.9985 $^{+0.0014}_{-0.0001}$	12.0 $^{+0.8}_{-0.8}$	5.2 $^{+0.5}_{-0.5}$	230/131	100/83
5	COMP-GR-FG	0^l	0.7465 $^{+0.0117}_{-0.0109}$	0.9523 $^{+0.0076}_{-0.0040}$	0.9976 $^{+0.0011}_{-0.0001}$	11.6 $^{+0.7}_{-0.7}$	4.9 $^{+0.5}_{-0.5}$	229/131	97/83
6	COMP-GR-B	0^l	0.7913 $^{+0.0116}_{-0.0152}$	0.9557 $^{+0.0040}_{-0.0070}$	0.9986 $^{+0.0008}_{-0.0006}$	11.3 $^{+0.8}_{-0.8}$	3.2 $^{+0.5}_{-0.5}$	207/131	113/83
7	BARE-AC-S	0^l	0.8491 $^{+0.0143}_{-0.0193}$	0.9637 $^{+0.0039}_{-0.0088}$	0.9976 $^{+0.0023}_{-0.0005}$	9.6 $^{+0.6}_{-0.6}$	3.2 $^{+0.3}_{-0.3}$	196/131	101/83
8	BARE-AC-FG	0^l	0.8394 $^{+0.0218}_{-0.0283}$	0.9625 $^{+0.0040}_{-0.0089}$	0.9977 $^{+0.0022}_{-0.0001}$	9.3 $^{+0.5}_{-0.5}$	3.1 $^{+0.3}_{-0.3}$	198/131	104/83
9	BARE-AC-B	0^l	0.8783 $^{+0.0172}_{-0.0561}$	0.9665 $^{+0.0017}_{-0.0088}$	0.9982 $^{+0.0009}_{-0.0010}$	11.3 $^{+0.6}_{-0.6}$	3.0 $^{+0.3}_{-0.3}$	184/131	109/83
10	COMP-AC-S†	0^l	0.5143 $^{+0.0092}_{-0.5143}$	0.9389 $^{+0.0060}_{-0.0208}$	0.9978 $^{+0.0009}_{-0.0001}$	13.4 $^{+1.3}_{-1.3}$	5.8 $^{+0.8}_{-0.8}$	224/131	103/83
11	COMP-AC-FG	0^l	0.6737 $^{+0.0067}_{-0.2864}$	0.9474 $^{+0.0040}_{-0.0180}$	0.9977 $^{+0.0009}_{-0.0001}$	12.6 $^{+0.8}_{-0.8}$	5.4 $^{+0.5}_{-0.5}$	239/131	105/83
12	COMP-AC-B	0^l	0.2954 $^{+0.0025}_{-0.2954}$	0.9225 $^{+0.0127}_{-0.0175}$	0.9976 $^{+0.0018}_{-0.0001}$	23.3 $^{+1.1}_{-1.1}$	6.4 $^{+0.7}_{-0.7}$	213/131	137/83
13	COMP-NC-S	0^l	0.3356 $^{+0.0039}_{-0.3356}$	0.9306 $^{+0.0025}_{-0.0358}$	0.9979 $^{+0.0012}_{-0.0002}$	15.5 $^{+1.0}_{-1.0}$	6.5 $^{+0.7}_{-0.7}$	234/131	121/83
14	COMP-NC-FG	0^l	0.1093 $^{+0.0125}_{-0.1093}$	0.9071 $^{+0.0072}_{-0.0234}$	0.9991 $^{+0.0004}_{-0.0002}$	16.2 $^{+1.3}_{-1.3}$	5.4 $^{+0.7}_{-0.7}$	211/131	136/83
15	COMP-NC-B	0^l	0.0100 $^{+0.0273}_{-0.0100}$	0.8873 $^{+0.0161}_{-0.0194}$	0.9990 $^{+0.0005}_{-0.0001}$	17.9 $^{+1.2}_{-1.2}$	5.9 $^{+0.7}_{-0.7}$	271/131	169/83
16	WD01-A	0^l	0.8302 $^{+0.0086}_{-0.1314}$	0.9628 $^{+0.0039}_{-0.0027}$	0.9975 $^{+0.0017}_{-0.0001}$	8.4 $^{+0.5}_{-0.5}$	4.5 $^{+0.3}_{-0.3}$	246/131	97/83
17	WD01-B	0^l	0.4035 $^{+0.0122}_{-0.4035}$	0.9326 $^{+0.0027}_{-0.0243}$	0.9977 $^{+0.0011}_{-0.0001}$	3.6 $^{+0.2}_{-0.2}$	1.4 $^{+0.1}_{-0.1}$	235/131	116/83
18	XLNW	0^l	0.8772 $^{+0.0034}_{-0.0517}$	0.9701 $^{+0.0088}_{-0.0065}$	0.9974 $^{+0.0019}_{-0.0004}$	11.6 $^{+0.6}_{-0.6}$	4.3 $^{+0.4}_{-0.4}$	175/131	88/83
	<i>min. value</i>	0^l	0.0100 ^l	0.8873	0.9974	3.6	1.4	–	–
	<i>max. value</i>	0^l	0.9010	0.9783	0.9991	23.3	6.8	–	–
	<i>mean value</i>	0^l	0.6429	0.9479	0.9980	12.4	4.5	–	–
	<i>median value</i>	0^l	0.7913	0.9523	0.9977	11.6	4.5	–	–

5 IMPACT OF THE DUST SCATTERING FOR SOURCE SPECTRA

5.1 Extraction Region Dependence of the Observed Spectrum

X-ray photons passing through the ISM are exposed to both photoelectric absorption and scattering by dust grains (Draine 2003; Ueda et al. 2010). Recent studies on simulated *Chandra* spectra show that the $N_{H,abs}$ can be overestimated by 25% if the dust scattering opacity is not considered in the spectral fitting (Corrales et al. 2016). Assuming a similar dust-to-gas ratio, the spectral effect of the dust scattering opacity becomes more significant for sources with higher $N_{H,abs}$, such as many sources in the GC.

Dust along the LOS can scatter X-ray photons out of the LOS, which are partially compensated by the photons scattered into the LOS and are subsequently included in the spectral extraction region (see Fig.1 in Smith, Valencic & Corrales 2016). If an infinite spectral extraction region is adopted, the loss of LOS photons will be fully compensated by the extra photons from the entire dust scattering halo, thereby leaving no dust scattering opacity in the observed spectrum. In reality, the spectral extraction region has a limited size. Besides, a small region of the PSF core is sometimes excised in order to avoid the photon pileup. Therefore, the effective dust scattering opacity is often nonzero and so affects the observed spectrum.

Based on our halo profile modelling, we can inves-

tigate this spectral effect for AX J1745.6-2901. We extracted *XMM-Newton* pn spectra from an observation (ObsID:0724210201) with two annulus regions (i.e. 15-40 arcsec and 20-80 arcsec) and applied the PSF correction using SAS task ARFGEN. Background spectra were extracted from an observation (ObsID:0658600201) where AX J1745.6-2901 was in quiescence. Fig.11 shows the two source spectra (background subtracted) differ significantly from each other in terms of both spectral shape and flux. The spectral difference increases towards soft X-rays, which is consistent with the fact that the effect of dust scattering is stronger for lower energy photons. We note that there could also be spectral discrepancies due to an annulus extraction region for point-like sources because of the PSF calibration issue¹¹, but this type of spectral discrepancy is mainly above 4 keV, with ~15% flux uncertainty in the 5-10 keV band, which is much smaller than what we found in AX J1745.6-2901 due to the dust scattering halo.

To quantify the spectral discrepancy, we fit the two spectra within 2-10 keV using a simple absorbed power law model (XSPEC model: TBNEW*POWERLAW), with cross-sections of Verner et al. (1996) and WAM00 ISM abundances. For the same $N_{H,abs}$, the best-fit photon index differs by 0.3 (or 6 σ) between the two spectra, while the flux

¹¹ <http://www.cosmos.esa.int/web/xmm-newton/sas-watchout-epic-spectra-annular-sources>

Table 3. Parameters calculated from the best-fit models in Table 2 with propagated statistical errors. τ is the dust scattering optical depth calculated from each of the best-fit models. $f_{N_{H,2}} = N_{H,2}/(N_{H,1} + N_{H,2}) \times 100\%$, where $N_{H,1}$ and $N_{H,2}$ are the $N_{H,sca}$ in Layer-1 and Layer-2, separately. $f_{N_{H,2}}$ indicates the percentage of dust contained in Layer-2 along the LOS.

No.	Dust Model	$x_{low,1} - x_{high,2}$	$N_{H,2} + N_{H,1}$ (10^{22} cm^{-2})	$f_{N_{H,2}}$ (%)	$\tau(2-4\text{keV})$	$\tau(4-6\text{keV})$	$\tau(6-10\text{keV})$
0	MRN77	0.0834 ^{+0.0716} _{-0.0140}	20.0 ^{+1.4} _{-1.4}	66.0 ^{+3.5} _{-3.5}	0.80	0.53	0.41
1	BARE-GR-S	0.1461 ^{+0.0643} _{-0.0103}	11.8 ^{+0.6} _{-0.6}	80.5 ^{+2.2} _{-2.2}	0.51	0.28	0.18
2	BARE-GR-FG	0.1297 ^{+0.0473} _{-0.0156}	13.5 ^{+0.9} _{-0.9}	77.8 ^{+3.2} _{-3.2}	0.65	0.42	0.32
3	BARE-GR-B	0.0750 ^{+0.0098} _{-0.0098}	19.9 ^{+1.0} _{-1.0}	75.9 ^{+3.2} _{-3.2}	0.66	0.40	0.28
4	COMP-GR-S	0.3161 ^{+0.0960} _{-0.0044}	17.2 ^{+0.9} _{-0.9}	69.8 ^{+2.5} _{-2.5}	0.93	0.65	0.53
5	COMP-GR-FG	0.2058 ^{+0.0222} _{-0.0122}	16.5 ^{+0.9} _{-0.9}	70.3 ^{+2.5} _{-2.5}	0.79	0.51	0.37
6	COMP-GR-B	0.1644 ^{+0.0157} _{-0.0157}	14.5 ^{+0.9} _{-0.9}	77.9 ^{+3.0} _{-3.0}	0.61	0.38	0.28
7	BARE-AC-S	0.1146 ^{+0.0197} _{-0.0168}	12.8 ^{+0.7} _{-0.7}	75.0 ^{+2.1} _{-2.1}	0.63	0.37	0.26
8	BARE-AC-FG	0.1231 ^{+0.0177} _{-0.0235}	12.4 ^{+0.6} _{-0.6}	75.0 ^{+2.1} _{-2.1}	0.62	0.37	0.25
9	BARE-AC-B	0.0882 ^{+0.0561} _{-0.0296}	14.3 ^{+0.7} _{-0.7}	79.0 ^{+1.9} _{-1.9}	0.57	0.33	0.23
10	COMP-AC-S [†]	0.4246 ^{+0.5143} _{-0.0227}	19.2 ^{+1.5} _{-1.5}	69.8 ^{+3.6} _{-3.6}	0.82	0.55	0.43
11	COMP-AC-FG	0.2737 ^{+0.2864} _{-0.0192}	18.0 ^{+0.9} _{-0.9}	70.0 ^{+2.4} _{-2.4}	0.77	0.50	0.37
12	COMP-AC-B	0.6271 ^{+0.2957} _{-0.0177}	29.7 ^{+1.3} _{-1.3}	78.5 ^{+2.0} _{-2.0}	0.55	0.33	0.22
13	COMP-NC-S	0.5950 ^{+0.3356} _{-0.0363}	22.0 ^{+1.2} _{-1.2}	70.5 ^{+2.6} _{-2.6}	0.78	0.53	0.41
14	COMP-NC-FG	0.7978 ^{+0.0265} _{-0.0189}	21.6 ^{+1.3} _{-1.3}	75.0 ^{+2.9} _{-2.9}	0.85	0.63	0.52
15	COMP-NC-B	0.8773 ^{+0.0189} _{-0.0335}	23.8 ^{+1.3} _{-1.3}	75.2 ^{+2.5} _{-2.5}	0.71	0.48	0.38
16	WD01-A	0.1326 ^{+0.1315} _{-0.0254}	12.9 ^{+0.6} _{-0.6}	65.1 ^{+2.0} _{-2.0}	1.07	0.76	0.61
17	WD01-B	0.5291 ^{+0.4035} _{-0.0272}	5.0 ^{+0.2} _{-0.2}	72.0 ^{+1.8} _{-1.8}	0.75	0.51	0.40
18	XLNW	0.0929 ^{+0.0524} _{-0.0073}	15.9 ^{+0.7} _{-0.7}	73.0 ^{+2.1} _{-2.1}	0.71	0.43	0.30
<i>min. value</i>		0.0882	5.0	66.0	0.51	0.28	0.18
<i>max. value</i>		0.8773	29.7	80.5	1.07	0.76	0.61
<i>mean value</i>		0.3051	16.9	74.4	0.73	0.47	0.36
<i>median value</i>		0.1644	16.5	75.0	0.71	0.48	0.37

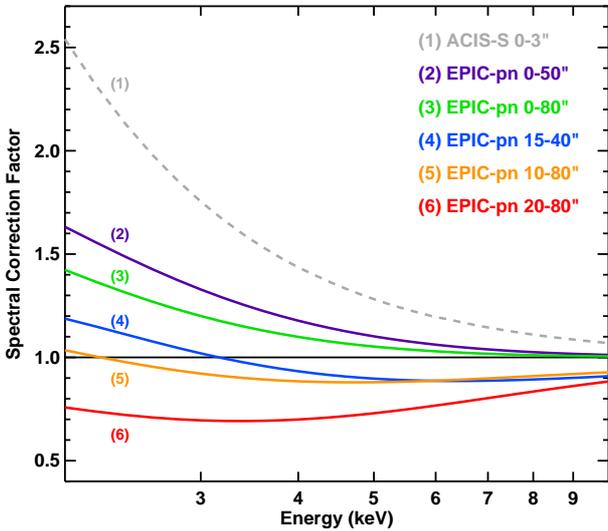


Figure 12. Energy dependent spectral correction factor for the dust scattering opacity (Section 5.2) for various instruments and spectral extraction regions. Note that these curves are only for the dust on the LOS of AX J1745.6-2901.

differs by 30%, 22% and 9% in the 2-4, 4-6 and 6-10 keV bands, respectively (see Table 4), suggesting the spectral dependence on the source extraction region is very significant and so must be addressed properly.

5.2 Spectral Correction for the Dust Scattering halo

Based on the best-fit two-layer model, we can perform corrections on the spectral effect of the dust scattering halo. We assume the PSF correction has been performed on the spectra extracted from different annulus regions, so the spectral correction is essentially the correction for the difference between the observed source radial profile and the radial profile of the instrumental PSF, with the same flux in the source extraction region. Therefore, this correction only requires well-constrained radial profiles of the source at various energies, but does not require the knowledge about intrinsic properties of the intervening dust.

We defined the spectral correction factor as $F_{int,E}/F_{obs,E}$, where $F_{obs,E}$ and $F_{int,E}$ are the flux at energy E before and after the dust scattering correction, and used the best-fit model with the COMP-AC-S dust grain to calculate the correction factor. Fig.12 shows the correction factor as a function of energy for different source extraction regions. A bigger correction factor is found at lower energies with smaller extraction regions, because in this case a bigger halo flux loss is expected. Therefore, for instruments such as *Chandra* ACIS where a small source extraction regions is often used, it is more important to consider the spectral effect of the dust scattering halo (Corrales et al. 2016; Smith, Valencic & Corrales 2016).

Based on the spectral correction factor in Fig.12, we can correct the observed spectra for the dust scattering halo and produce a ‘dust free’ spectrum. Fig.11 shows that the two spectra from different source extraction regions become

Table 4. Comparison of the best-fit parameters for the two spectra in Fig.11 fitted using the TBNEW*POWERLAW model in XSPEC within 2-10 keV. ‘Corr’ indicates whether the AXJDUST model for the dust scattering opacity correction is incorporated or not. Including the AXJDUST model enables us to obtain consistent spectral parameters from the two very different spectra. † this $N_{H,abs}$ does not necessarily reflect the true $N_{H,abs}$ of AX J1745.6-2901, because the absorbed power law model is not a good model for the entire 2-10 keV band (see Section 5.2).

Spectrum	Corr.	† $N_{H,abs}$ (10^{22} cm^{-2})	Γ	Norm	F_{2-4keV}	F_{4-6keV}	$F_{6-10keV}$	χ^2/dof
					$(10^{-12} \text{ erg cm}^{-2} \text{ s}^{-1})$			
Spec (15-40'')	no	40.9 ± 0.6	3.06 ± 0.04	1.65 ± 0.1	15.2 ± 0.1	68.4 ± 0.4	97.8 ± 0.5	1866/1359
Spec (20-80'')	no	40.5 ± 0.4	3.38 ± 0.04	3.46 ± 0.3	21.8 ± 0.1	86.3 ± 0.5	107.3 ± 0.6	1916/1409
Spec (15-40'')	yes	38.8 ± 0.6	3.01 ± 0.04	1.30 ± 0.12	14.1 ± 0.1	60.7 ± 0.3	88.6 ± 0.5	1827/1359
Spec (20-80'')	yes	39.1 ± 0.6	3.06 ± 0.04	1.46 ± 0.13	14.5 ± 0.1	62.3 ± 0.3	89.5 ± 0.5	1942/1409

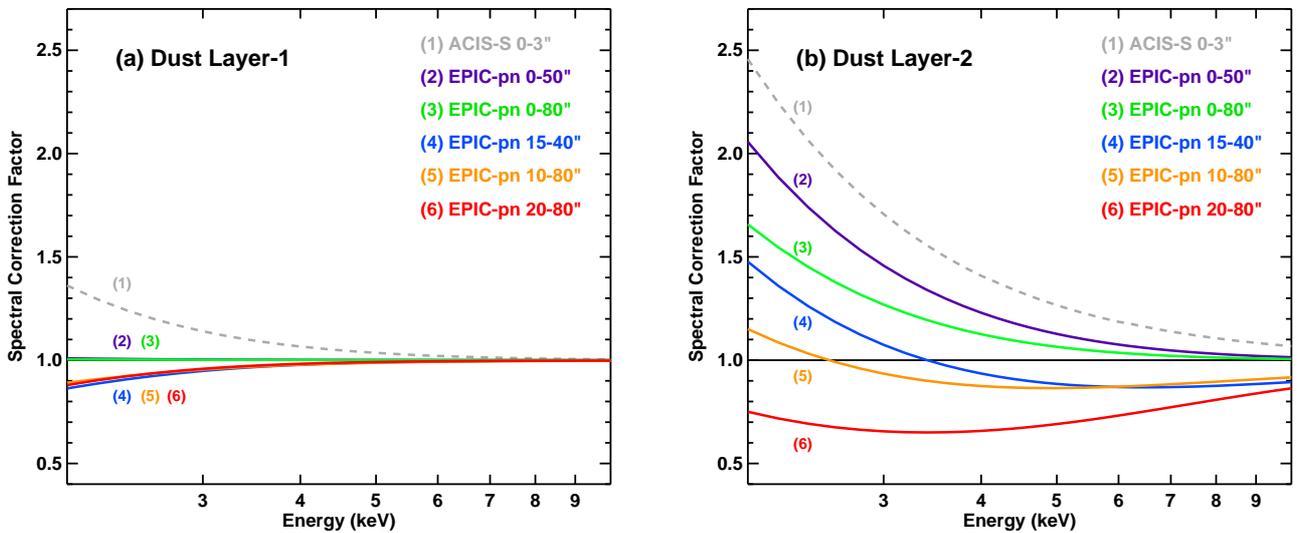


Figure 13. Energy dependent spectral correction factor for the two dust layers for various instruments and spectral extraction regions, based on the best-fit parameters with the COMP-AC-S dust grains (Table 2). Panel-a is for Layer-1 which is local to AX J1745.6-2901; Panel-b is for Layer-2 which is far from AX J1745.6-2901. It is clear that Layer-2 introduces much bigger spectral biases than Layer-1.

consistent with each other after the spectral correction. However, this correction to the spectral file itself is only approximate as it does not incorporate the full response profile of every energy channel in EPIC-pn, and so the most appropriate way is to correct the model rather than the spectral file. Therefore, we built an XSPEC multiplicative model (AXJDUST) for the dust scattering halo around AX J1745.6-2901 for different annulus regions and instruments¹². We found that adding this model component can fully account for the spectral discrepancy in Fig.11 and produce consistent best-fit parameters from the two apparently different spectra. The best-fit parameters also changes significantly after the inclusion of AXJDUST model (see Table 4). Since the absorbed power law model is clearly not good enough for the entire 2-10 keV band, we restrict the fitting to 3-6 keV band and we find a much better fit with $\chi^2/dof = 595/596$, $N_{H,abs} = (3.02 \pm 0.17) \times 10^{23} \text{ cm}^{-2}$ and $\Gamma = 2.02 \pm 0.17$. These parameters are also roughly consistent with those reported by Paizis et al. (2015) using the lower S/N *Chandra* HETG

spectrum and applying the same absorbed power law model with the same cross-sections and abundances.

Furthermore, to assess the contribution from different dust layers, we calculated the spectral correction factor assuming there is only dust layer-1 or layer-2, separately. It is found that the bias is mainly caused by dust layer-2 which is far from AX J1745.6-2901 (Fig.13). This is because dust layer-1 is so close to AX J1745.6-2901 that its halo is very compact and its shape very similar to the instrumental PSF (see Fig.9), and so the PSF correction also corrects for the halo flux loss with small biases. Dust layer-2 is far from AX J1745.6-2901, and so it produces an extended dust scattering halo, which is more extended than the instrumental PSF and is also more difficult to be covered by a small spectral extraction region, thereby requiring a more significant correction factor.

¹² AXJDUST model is available upon request.

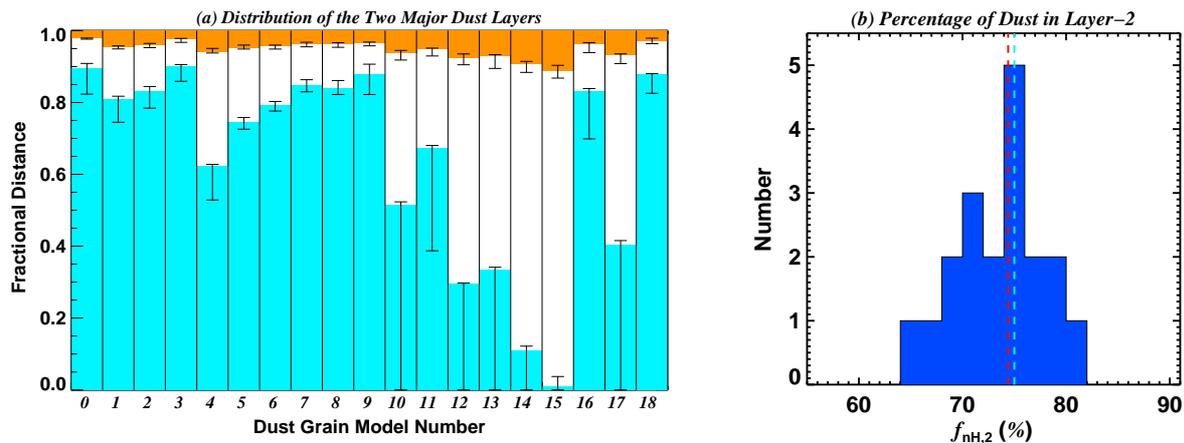


Figure 14. Panel-a: the distribution of the two major dust layers along the LOS, with the orange region being Layer-1 and cyan region being Layer-2. Distance 0 corresponds to the Earth, while distance 1 corresponds to AX J1745.6-2901. Error bars indicate the statistical uncertainty of each layer’s boundary. Significant separation between the two layers can be found in very dust grain model. Panel-b: the distribution of the percentage of LOS dust in dust Layer-2. The red and cyan dash lines indicate the mean value of 74.4% and median value of 75.0%, separately. Dust grain model number and all the data can be found in Table 2, 3.

6 DISCUSSION

6.1 Dust Distribution in the GC Direction

Interstellar dust is typically contained in molecular clouds, whose spatial distribution can be traced by molecular lines (e.g. CO emission, Dame et al. 2001). Previous Galactic CO maps showed a prominent feature due to a massive molecular ring at 4-5 kpc from the GC (the so-called 5 kpc molecular ring, Jackson et al. 2006; Simon et al. 2006), or 3-4 kpc from Earth in the GC direction (Valleé 2014). Later studies showed an enhancement of mass density of the molecular clouds between 3-6 kpc from the GC overlapping with the 5 kpc molecular ring, which also seems to be distributed along the Scutum Arm (Roman-Duval et al. 2010; Dobbs & Burkert 2012; Sato et al. 2014; Heyer & Dame 2015). Previous surveys about the infrared dark clouds (Egan et al. 1998; Carey et al. 1998) found these clouds mainly lie along the Scutum arm between 3-6 kpc from the GC (Simon et al. 2006; Marshall, Joncas & Jones 2009) with typical $N_H \gtrsim 10^{22} \text{ cm}^{-2}$ in a single cloud (Simon et al. 2006; Peretto & Fuller 2010). Therefore, it is indeed possible for the GC LOS to contain a large amount of intervening dust about half way between the Earth and the GC.

Another method to infer the dust distribution is to use the dust extinction map. Using the infrared extinction curve towards the GC, Fritz et al. (2011) reported that most of the extinction is caused by the dust in the Galactic disk rather than in the Nuclear bulge (Mezger et al. 1996). Schultheis et al. (2014) built a 3D dust extinction map for the Galactic bulge and found a peak of extinction at 3 kpc from the GC, implying a dust lane in front of the Galactic bar, but they cannot rule out the possibility that this peak is associated with the 5 kpc molecular ring.

All the above studies consistently show that most of the dust is distributed several kpc from the GC in the Galactic disk. This invalidates the assumption made in Tan & Draine (2004), where the dust scattering halo from Sgr A* was assumed to rise from the intervening dust and gas close to Sgr A*. Interestingly, recent studies of PSR J1745-2900, the magnetar located at only 2.4 arcsec from Sgr A* (Kennea et

al. 2013; Mori et al. 2013; Rea et al. 2013), also show that most of the scattering medium responsible for the broadening of the radio emission from this magnetar is likely associated with the interstellar gas in the Galactic disk, possibly in a nearby spiral arm, rather than close to the GC (Eatough et al. 2013; Shannon & Johnston 2013; Bower et al. 2014; Wucknitz 2015; Sicheneder & Dexter 2016; but see Spitler et al. 2014).

Our study on AX J1745.6-2901 is the first detailed observation and modelling of a dust scattering halo around a source in the GC. The dust properties and distribution of the two major layers show a large dispersion, which is mainly due to the lack of knowledge about the properties of dust grains in the ISM especially along the GC LOS. But there are some robust results that can be derived from all the dust grain models. Firstly, although the separation between the two major layers changes from 0.09 to 0.88 (with mean value: 0.31) for different dust grain models, all the models consistently find a highly significant separation between the two layers (see Fig. 14a), i.e. Layer-1 is always found to be local to AX J1745.6-2901, while Layer-2 is significantly far away and always has its lower boundary reaching 0. Secondly, we find that the percentage of dust that is contained in each layer is tightly constrained and it is observed to be (66-81)% (with mean value: 74.4%) in Layer-2, with a much weaker dependence on the dust grain models than the other parameters (see Fig. 14b). We emphasize that an extra uncertainty of this dust distribution lies in the assumption of the same type of dust grains in different dust layers. It has been reported that the abundances in the GC may be higher than in the Galactic disk (see Section 6.2 for more detailed discussion), and the GC dust grains might have a smaller characteristic size (Hankins et al. 2017). However, there are no strong constraints on these GC parameters, and we are not sure if AX J1745.6-2901 is located inside the GC region or not (see discussions below), thus in this work we do not consider different dust grain models for different dust layers.

The absolute location of these two layers also depends on the absolute distance of AX J1745.6-2901 itself, which is still highly uncertain. From the fact that the $N_{H,abs}$ of

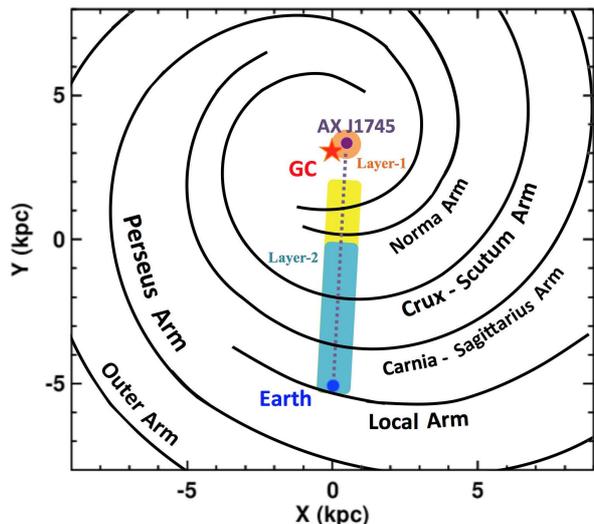


Figure 15. Schematic of the Milky Way spiral arms viewed from the north Galactic pole based on the Milky Way structure reported by Nakanishi & Sofue (2016), Reid et al. (2009) and Caswell & Haynes (1987). The orange and dark green regions indicate the location of dust layer-1 and dust layer-2 based on the average best-fit parameters in Table 2, assuming AX J1745.6-2901 is in the GC. The yellow region shows the highest upper boundary (0.9) of dust layer-2 from the fitting using the BARE-GR-B dust grain model.

AX J1745.6-2901 is significantly larger than that of Sgr A* (fusing the same absorption model and abundances, see Section 1.3), it can be reasonably inferred that AX J1745.6-2901 is NOT closer to the Earth than Sgr A*. Although $N_{H,sca}$ is not directly comparable to the $N_{H,abs}$ as they have different normalizations (see Section 6.2), by assuming the same dust properties in all the layers, we can use the fraction of dust to infer the fraction of $N_{H,abs}$ in each layer. Therefore, adopting $N_{H,abs} = 3 \times 10^{23} \text{ cm}^{-2}$ for AX J1745.6-2901 (see Section 5.2) and $1.6 \times 10^{23} \text{ cm}^{-2}$ for Sgr A* from the TBNEW model and WAM00 abundances (Ponti et al. 2017b), we can estimate $N_{H,abs} = (2.0 - 2.4) \times 10^{23} \text{ cm}^{-2}$ in Layer-2, which is similar to Sgr A*. This is also consistent with the fact that Layer-1 is local to AX J1745.6-2901 and Layer-2 is much farther away, thus it is likely that only Layer-2 is in front of Sgr A*.

If a distance of 8 kpc is adopted for AX J1745.6-2901, the mean fractional distance of Layer-2 would indicate that it is 0-5 kpc from the Earth in the Galactic disk, thus overlaps with the Crux-Scutum Arm, Carina-Sagittarius Arm and the local Arm (see Fig.15; Nakanishi & Sofue 2016; Reid et al. 2009; Caswell & Haynes 1987). This result immediately implies the possibility that the same dust in Layer-2 may also intervene along the LOS of many GC sources. There are two X-ray transients whose LOSs are very close to Sgr A*. Swift J174540.7-290015 was discovered in 2016 at 16 arcsec from Sgr A* (Reynolds et al. 2016), with $N_{H,abs} = 1.7 \times 10^{23} \text{ cm}^{-2}$ (TBNEW model and WAM00 abundances, Ponti et al. 2016). SGR J1745-2900 is located at 2.4 arcsec from Sgr A* (Rea et al. 2013), with $N_{H,abs} = 1.9 \times 10^{23} \text{ cm}^{-2}$ (TBABS model and WAM00 abundances, Coti-Zelati et al. 2015). Although a further correc-

tion for the dust scattering opacity is probably also necessary for these GC sources, their $N_{H,abs}$ are indeed similar to that of Layer-2 (also see Ponti et al. 2017b), which support Layer-2 being a common GC-foreground dust layer in the Galactic disk. A more conclusive study would be to compare the dust scattering halo around many GC sources and see if a similar dust component as Layer-2 exists in their halo profiles.

6.2 Testing Dust-to-Gas Relations in the GC Direction

Since τ_{sca} , $N_{H,sca}$ and $N_{H,abs}$ have all been measured independently in AX J1745.6-2901, it becomes the first source that allows the test of dust-to-gas relations in the GC LOS. However, we would like to emphasize that there are some underlying systematic uncertainties associated with the N_H measurement. Firstly, $N_{H,sca}$ depends heavily on the assumptions made in the dust grain models such as the abundances and the grain size distribution, which affect the halo intensity and profile, which then affect $N_{H,sca}$ during the halo profile fitting. This is why a large dispersion of $N_{H,sca}$ is found among all the dust grain models. Secondly, $N_{H,abs}$ also depends heavily on the assumed abundances during the X-ray spectral fitting. In the case of AX J1745.6-2901, we found $\sim 50\%$ change in $N_{H,abs}$ by simply changing the abundances from the AG89 solar abundances to WAM00 ISM abundances. Moreover, $N_{H,abs}$ depends on the absorption model and intrinsic spectral model used for the spectral fitting. A high S/N and resolution spectra with a good intrinsic model would allow the absorption model to determine $N_{H,abs}$ from various metal absorption edges; while a low S/N spectra or a bad intrinsic model would force the absorption model to trace the spectral curvature (as likely in AX J1745.6-2901 and many other GC sources where the strong extinction leads to a fast decrease of S/N in the soft X-ray band), which would bias $N_{H,abs}$. Therefore, it is not meaningful to do a precise comparison between $N_{H,sca}$ and $N_{H,abs}$. Actually, some of the dust and gas in Layer-1 can be so close to AX J1745.6-2901 that they would absorb X-rays but not *observably* scatter them, and so the *observed* $N_{H,sca}$ should be smaller than the intrinsic N_H . It must also be stressed that it is necessary to provide key informations when an $N_{H,abs}$ is reported, such as the quality of spectra being fitted, models used for the intrinsic spectra and X-ray absorption, and assumptions made for the cross-sections and abundances.

On the other hand, several versions of the dust-to-gas relations between $N_{H,abs}$, A_V and τ_{sca} have been reported so far. For example, Predehl & Schmitt (1995) (PS95) performed X-ray spectral and halo fittings for 25 point sources observed by *ROSAT* and reported $N_{H,abs}/A_V = 1.79 \times 10^{21} \text{ cm}^{-2} \text{ mag}^{-1}$ and $\tau_{sca} = 0.087 \times A_V(\text{mag}) E(\text{keV})^{-2}$. Their absorption model used the Morrison & McCammon (1983) cross-sections, which is equivalent to the XSPEC WABS model. However, Draine & Bone (2004) (DB04) argued that Predehl & Schmitt (1995) may have underestimated the scattering power and so a better relation would be $\tau_{sca} = 0.15 \times A_V(\text{mag}) E(\text{keV})^{-1.8}$. Later Güver & Özel (2009) (LO09) collected 22 supernova remnants which have their $N_{H,abs}$ measured from the high-quality X-ray spectra from *Chandra*, *XMM-Newton* and *Suzaku*, and have in-

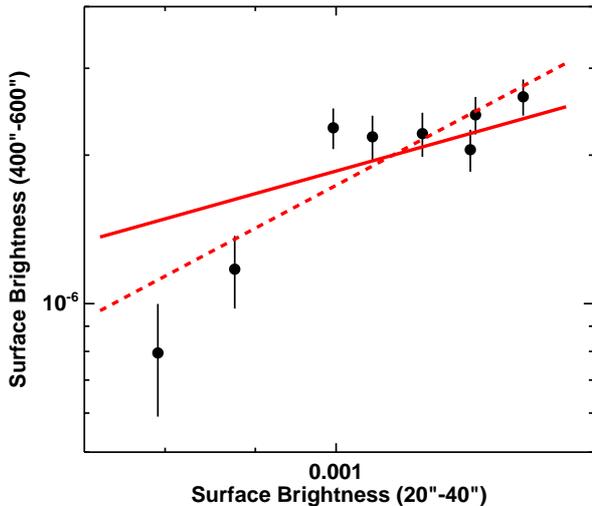


Figure 16. The correlation between the halo wing and halo core intensity in the 4-6 keV band. The unit for the surface brightness is $\text{counts cm}^{-2}\text{s}^{-1}\text{arcmin}^{-2}$. The data points are measured from the eight *XMM-Newton* observations in Table 1. The red solid line is the best-fit line assuming a slope of 1, the red dashed line is the best-fit line with the best-fit slope of 1.9 ± 0.5 .

dependent A_V measurements. They reported $N_{H,abs}/A_V = (2.21 \pm 0.09) \times 10^{21} \text{cm}^{-2} \text{mag}^{-1}$. Their difference from Predehl & Schmitt (1995) was attributed to the narrow band-pass of *ROSAT* and the simple power law model assumed for the intrinsic spectra. However, it is difficult to track down all the absorption models, cross-sections and abundances used to derive the $N_{H,abs}$ of their sample because they were reported by many different papers and authors. Recently, Valencic & Smith (2015) (VS15) analyzed the X-ray dust scattering halo around 35 sources and found $N_{H,abs}/A_V = (2.08 \pm 0.26) \times 10^{21} \text{cm}^{-2} \text{mag}^{-1}$, with $\tau_{sca}(1.0\text{keV})/A_V$ having a typical $\pm 1\sigma$ range of 0.02-0.08 depending on the dust grain model. Their $N_{H,abs}$ was based on the AG89 solar abundances and PHABS model.

In order to use these relations, we adopt $N_{H,abs} \approx 2 \times 10^{23} \text{cm}^{-2}$ for AX J1745.6-2901 also based on the PHABS model and the AG89 solar abundances. It is known that the metallicity near/in the GC is higher than the solar abundances (e.g. Kubryk, Prantzos & Athanassoula 2015 and references therein). If some of the dust and gas in e.g. Layer-1 is indeed close to the GC, the $N_{H,abs}$ from the AG89 abundances would have been over-predicted. However, it is also reported that the AG89 metallicity over-estimated the solar metallicity (e.g. Asplund et al. 2009; Nieva & Przybilla 2012), thus the adoption of AG89 abundances may actually alleviate the potential higher-metallicity problem towards AX J1745.6-2901. We find $A_V = 117, 90, 96 \text{mag}$ from the $N_{H,abs}/A_V$ relations in PS95, LO09 and VS15, respectively. Adopting $A_V \approx 100 \text{mag}$, we derive $\tau_{sca} = 0.80, 0.35$ and 0.18 at 3.3, 5.0 and 7.0 keV from the τ_{sca}/A_V relation in PS95, or $\tau_{sca} = 1.75, 0.83$ and 0.45 from DB04, or $\tau_{sca} = 0.73, 0.32$ and 0.16 from VS15. Therefore, it appears that the dust-to-gas relations in PS95 and VS15 give more consistent τ_{sca} with those measured directly from the eclipsing light curves.

6.3 Excess Flux in the Halo Wing

In this work, we reported, for the first time, the detection of a significant wing of the dust scattering halo at $E \geq 5 \text{keV}$. On the other hand, previous studies about the dust scattering halo mainly focussed on energies below 3 keV (e.g. Predehl & Schmitt 1995; Xiang, Zhang & Yao 2005; Valencic & Smith 2015), and used a free constant to fit the underlying background, hampering the distinguishing of an extended halo wing even if it existed. This is possible in our work because the background is well constrained from observations where AX J1745.6-2901 was in quiescence, which then enabled us to detect this excess flux in the halo wing.

In order to verify that this halo wing is related to the scattering light from AX J1745.6-2901, we tested the correlation between the flux of the halo wing and the flux of the halo core using all the *XMM-Newton* observations in Table 1. Fig.16 shows that a clear correlation is found. The Spearman’s rank correlation coefficient is 0.79 and the p -value is 0.02, and Pearson’s linear correlation coefficient is 0.89. A linear fitting with slope fixed at 1 would produce a χ^2 of 7.1 for 7 dof (red solid line in Fig.16). Freeing the slope in the fitting only improves the χ^2 by 1.0 for 1 extra free parameter, and the best-fit slope is 1.9 ± 0.5 (red dashed line in Fig.16). These results indicate the correlation between the halo wing flux and halo core flux is consistent with a linear correlation, confirming that the wing structure in the radial profile is indeed part of the dust scattering halo around AX J1745.6-2901.

Since smaller dust grains can scatter photons at larger angles more efficiently, this halo wing implies an excess fraction of dust grains of relatively small sizes in the ISM. According to Eq.5, for a viewing angle of 300-600 arcsec at 5 keV, the typical dust grain size would be 300-590 Å, which is still two orders of magnitude bigger than the wavelength of X-ray photons (2.5 Å for 5 keV), and so the Rayleigh-Gans approximation is still valid. However, the outer boundary of the halo wing is still unknown as at larger radii the wing starts to blend with the diffuse and background emission. If the wing is more extended than currently detected, the inferred dust grains can have smaller sizes. Therefore, the halo wing suggests a higher fraction of dust grains with typical sizes of $\lesssim 590 \text{Å}$ than considered in current dust grain models. This size range covers the typical sizes of PAH (3.5-55 Å, ZDA04), and is within the typical sizes of Graphite, ACH2, Olivine, Enstatite, Fe metal and the composite dust grains in ZDA04 (but is one order of magnitude smaller than the largest dust grains). Since we do not know the size distribution and composition of this dust, it is not possible to predict the influence in other aspects such as the IR emission or excitation. Future studies of the dust grain model along the GC LOS should take these halo wing phenomena into account.

6.4 Halo Uniformity

The non-uniformity of dust scattering halos has been noticed before (McCullough, Smith, Valencic 2013; Seward & Smith 2013; Valencic & Smith 2015), which may rise from partially-aligned non-spherical grains (Draine & Allaf-Akbari 2006) or azimuthal variation of the foreground dust column density. Following our previous discussion, if AX J1745.6-2901

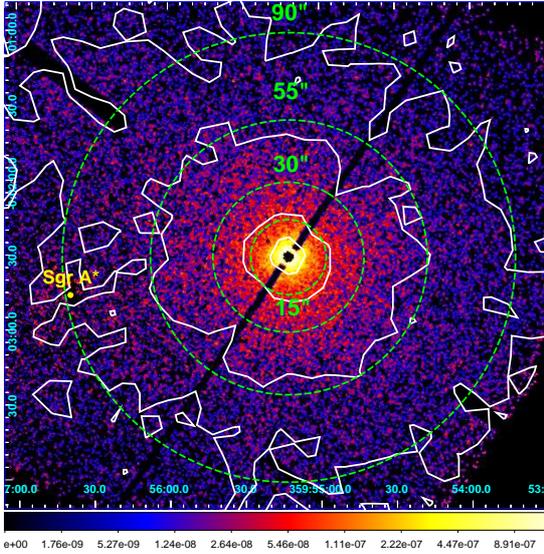


Figure 17. Background subtracted image of AX J1745.6-2901 in the 4-6 keV band using ACIS-S observations. The vertical axis is the Galactic latitude, the horizontal axis is the Galactic longitude. The readout streaks and central 2.5 arcsec are subtracted. The white surface brightness contours are based on the same image smoothed by 15 arcsec, which shows no strong azimuthal variation except for the artificial non-uniformity due to the readout streaks.

is located inside or beyond the GC, then Layer-1 is probably associated with the molecular clouds in the Central Molecular Zone (CMZ), whose column density is known to vary severely along different LOS (e.g. Serabyn & Guesten 1987; Morris & Serabyn 1996; Molinari et al. 2011; Ponti et al. 2013; Langer et al. 2015; Henshaw et al. 2016). If so, the halo produced by Layer-1 might show azimuthal variation. To test this, we generated a contour map on the ACIS-S image of AX J1745.6-2901 in the 4-6 keV band (Fig. 17, based on ObsID: 17857). No obvious azimuthal variation is seen in the image except for the artificial asymmetry due to the subtraction of the readout streaks, suggesting that Layer-1 should not have strong column density variation. The halo outside 20 arcsec is dominated by Layer-2, and we found no clear asymmetry out to the LOS of Sgr A*, either. This is supporting evidence that the same Layer-2 may also intervene along the LOS to Sgr A*. Further study of the non-uniformity would require more photon counts and more careful azimuthal structure analysis (e.g. Seward & Smith 2013), which is beyond the scope of this work.

7 CONCLUSIONS

In this paper, we used a large dataset from *Chandra* and *XMM-Newton* to conduct a detailed analysis of the radial profile of AX J1745.6-2901 whose LOS is ~ 1.45 arcmin away from Sgr A*. Being a bright X-ray source with high column density, this source is surrounded by a strong dust scattering halo across the entire 2-10 keV band.

(1) Based on the halo profile modelling, we identified two major scattering components, namely dust Layer-1 and Layer-2. For all the 19 dust grain models used in this

work, and assuming the same dust grain model along the LOS, we found Layer-1 is produced by the dust close to AX J1745.6-2901 within a fractional distance of $\lesssim 0.11$ and contains (19-34)% (mean value: 26%) of the total LOS dust. The remaining (66-81)% (mean value: 74%) LOS dust is contained in Layer-2, distributing from the Earth up to a fractional distance of 0.01-0.90, with a mean value of 0.64 from all dust grain models. These two layers separate from each other significantly by a fractional distance of 0.08-0.88, with a mean value of 0.31. Assuming that AX J1745.6-2901 is located at 8 kpc in the GC, Layer-2 must be distributed in the Galactic disk and probably associated with the molecular clouds distributed along the spiral arms on the LOS, such as the Crux-Scutum Arm, Carina-Sagittarius Arm and the Local Arm.

(2) In addition to the two dust scattering components, we identified an extended component in the halo wing which cannot be explained by any dust grain model used in this work. This halo wing suggests a higher fraction of dust grains with typical sizes of $\lesssim 590$ Å than considered in current dust grain models.

(3) We also investigated the influence of the dust scattering halo on the observed spectra. We found that the observed spectra of AX J1745.6-2901 can strongly depend on the source extraction region, which is due to the partial inclusion of the dust scattering halo. The more extended halo from Layer-2 has a stronger spectral effect. After applying the spectral correction for the dust scattering towards AX J1745.6-2901, and found $N_{H,abs} = (3.02 \pm 0.17) \times 10^{23} \text{cm}^{-2}$ and $\Gamma = 2.02 \pm 0.17$ by fitting the TBNEW*POWERLAW model to the 3-6 keV spectrum from *XMM-Newton* EPIC-pn, with cross-sections of Verner et al. (1996) and WAM00 ISM abundances.

(4) Since Layer-2 is likely associated with the dust in the Galactic disk several kpc away from AX J1745.6-2901, it may also intervene the LOS of other nearby GC sources such as Sgr A*. We found a rough consistency between the $N_{H,abs}$ of Layer-2, Sgr A* and a couple of nearby transients, so it is possible that Layer-2 also produces dust scattering in these sources. A more conclusive study would be to compare the halo shapes from all these sources. If it is confirmed that the same dust population in Layer-2 also intervenes other nearby LOSs towards the GC, a significant spectral correction for the dust scattering halo would be necessary for many GC sources.

(5) We built XSPEC models to account for the spectral discrepancy caused by the entire dust scattering halo around AX J1745.6-2901 (AXJDUST model), as well as the dust scattering in Layer-2 alone (FGDUST model) so that it can be applicable to the other sources in the GC.. These models can properly correct for different spectral discrepancies arising from different source extraction regions¹³.

¹³ We have created different versions of these dust scattering models for the CCD quality spectra from *XMM-Newton* EPIC, *Chandra* ACIS, *Swift* XRT and *NuSTAR*. These models will be

ACKNOWLEDGEMENTS

We thank Elisa Costantini for helpful discussions and comments. CJ acknowledges valuable discussions with Lynne Valencic about the cosmic abundances. The anonymous referee is appreciated for providing insightful comments and suggestions, which allowed us to improve the paper significantly. We would also like to thank the *XMM-Newton* team for providing important instructions on technical and calibration issues. This project is supported by the Bundesministerium für Wirtschaft und Technologie/Deutsches Zentrum für Luft- und Raumfahrt (BMW/DLR, FKZ 50 OR 1408 and FKZ 50 OR 1604) and the Max Planck Society. This work is based on observations obtained with *XMM-Newton*, an ESA science mission with instruments and contributions directly funded by ESA Member States and NASA. The scientific results reported in this article are also based on observations made by the *Chandra* X-ray Observatory, as well as data obtained from the *Chandra* Data Archive. This research has made use of software provided by the *Chandra* X-ray Center (CXC) in the application packages CIAO, ChIPS, and Sherpa.

REFERENCES

- Anders E., Ebihara M., 1982, *GeCoA*, 46, 2363
 Anders E., Grevesse N., 1989, *GeCoA*, 53, 197
 Asplund M., Grevesse N., Sauval A. J., Scott P., 2009, *ARA&A*, 47, 481
 Becklin E. E., Neugebauer G., 1968, *ApJ*, 151, 145
 Bohren C. F., Huffman D. R., 1983, New York: Wiley, 1983
 Bower G. C., Deller A., Demorest P., et al., 2014, *ApJL*, 780, L2
 Caswell J. L., Haynes R. F., 1987, *A&A*, 171, 261
 Carey S. J., Clark F. O., Egan M. P., et al., 1998, *ApJ*, 508, 721
 Corrales L. R., Paerels F., 2015, *MNRAS*, 453, 1121
 Corrales L. R., García J., Wilms J., Baganoff F., 2016, *MNRAS*, 458, 1345
 Costantini E., Freyberg M. J., Predehl P., 2005, *A&A*, 444, 187
 Dame T. M., Hartmann D., Thaddeus P., 2001, *ApJ*, 547, 792
 Degenaar N., Wijnands R., 2009, *A&A*, 495, 547
 Degenaar N., Wijnands R., Cackett E. M., et al., 2012, *A&A*, 545, A49
 Degenaar N., Wijnands R., Miller J. M., et al., 2015, *Journal of High Energy Astrophysics*, 7, 137
 Coti Zelati F., Rea N., Papitto A., et al., 2015, *MNRAS*, 449, 2685
 Dobbs C. L., Burkert A., 2012, *MNRAS*, 421, 2940
 Draine B. T., 2003, *ARA&A*, 41, 241
 Draine B. T., Allaf-Akbari K., 2006, *ApJ*, 652, 1318
 Draine B. T., Bond N. A., 2004, *ApJ*, 617, 987
 Eatough R. P., Kramer M., Lyne A. G., Keith M. J., 2013, *MNRAS*, 431, 292
 Egan M. P., Shipman R. F., Price S. D., et al., 1998, *ApJL*, 494, L199
 Fritz T. K., Gillessen S., Dodds-Eden K., et al., 2011, *ApJ*, 737, 73
 Gaetz T. J., 2010, Analysis of the Chandra On-Orbit PSF Wings, <http://cxc.harvard.edu/cal/Hrma/PSFWings.html>
 Genzel R., Eisenhauer F., Gillessen S., 2010, *Reviews of Modern Physics*, 82, 3121
 Ghizzardi S., 2002, XMM-Newton Calibration Documentation (XMM-SOC-CAL-TN-0029)
 Gladstone J., Done C., Gierliński M., 2007, *MNRAS*, 378, 13
 Güver T., Özel F., 2009, *MNRAS*, 400, 2050
 Hankins M., Lau R., Morris M., Herter T., 2017, *ApJ*, accepted (arXiv: 1702.04345)
 Heard V., Warwick R. S., 2013a, *MNRAS*, 428, 3462
 Heard V., Warwick R. S., 2013b, *MNRAS*, 434, 1339
 Heinke C. O., Tomsick J. A., Yusef-Zadeh F., Grindlay J. E., 2009, *ApJ*, 701, 1627
 Heinz S., Burton M., Braiding C., et al., 2015, *ApJ*, 806, 265
 Heinz S., Corrales L., Smith R., Brandt W. N., Jonker P. G., Plotkin R. M., Neilsen J., 2016, *ApJ*, 825, 15
 Henshaw J. D., Longmore S. N., Kruijssen J. M. D., et al., 2016, *MNRAS*, 457, 2675
 Heyer M., Dame T. M., 2015, *ARA&A*, 53, 583
 Hyodo Y., Ueda Y., Yuasa T., et al., 2009, *PASJ*, 61, S99
 Jackson J. M., Rathborne J. M., Shah R. Y., et al., 2006, *ApJS*, 163, 145
 Jansen F., Lumb D., Altieri B., et al., 2001, *A&A*, 365, L1
 Kennea J. A., Skinner G. K., 1996, *PASJ*, 48, L117
 Kennea J. A., Burrows D. N., Kouveliotou C., et al., 2013, *ApJL*, 770, L24
 Kubryk M., Prantzos N., Athanassoula E., 2015, *A&A*, 580, A127
 Kuulkers E., den Hartog P. R., in't Zand J. J. M., et al., 2003, *A&A*, 399, 663
 Langer W. D., Goldsmith P. F., Pineda J. L., et al., 2015, *A&A*, 576, A1
 Leccardi A., Molendi S., 2008, *A&A*, 486, 359
 Li A., Draine B. T., 2001, *ApJ*, 554, 778
 Maeda Y., Koyama K., Sakano M., Takeshima T., Yamauchi S., 1996, *PASJ*, 48, 417
 Mao J., Ling Z., Zhang S. N., 2014, *ApJ*, 785, 23
 Markwardt C. B., 2009, *Astronomical Data Analysis Software and Systems XVIII*, 411, 251
 Marshall D. J., Joncas G., Jones A. P., 2009, *ApJ*, 706, 727
 Mathis J. S., Lee C. W., 1991, *ApJ*, 376, 490
 Mathis J. S., Rumpl W., Nordsieck K. H., 1977, *ApJ*, 217, 425
 McCollough M. L., Smith R. K., Valencic L. A., 2013, *ApJ*, 762, 2
 Mezger P. G., Duschl W. J., Zylka R., 1996, *A&ARv*, 7, 289
 Molinari S., Bally J., Noriega-Crespo A., et al., 2011, *ApJL*, 735, L33
 Mori K., Gotthelf E. V., Zhang S., et al., 2013, *ApJL*, 770, L23
 Morris M., Serabyn E., 1996, *ARA&A*, 34, 645
 Morrison R., McCammon D., 1983, *ApJ*, 270, 119
 Munro M. P., Baganoff F. K., Bautz M. W., et al. 2003, *ApJ*, 589, 225
 Munro M. P., Bauer F. E., Baganoff F. K., et al., 2009, *ApJS*, 181, 110
 Munro M. P., Pfahl E., Baganoff F. K., et al., 2005, *ApJL*, 622, L113
 Nakanishi H., Sofue Y., 2016, *PASJ*, 68, 5
 Nieva M. F., Przybilla N., 2012, *A&A*, 539, A143
 Overbeck J. W., 1965, *ApJ*, 141, 864
 Paizis A., et al., 2015, *ApJ*, 808, 34
 Peretto N., Fuller G. A., 2010, *ApJ*, 723, 555
 Ponti G., Morris M. R., Terrier R., et al., 2015a, *MNRAS*, 453, 172
 Ponti G., Bianchi S., Muñoz-Darias T., et al., 2015b, *MNRAS*, 446, 1536
 Ponti G., Jin C., De Marco B., et al., 2016, *MNRAS*, 461, 2688
 Ponti G., De K., Muñoz-Darias T., Stella, L., Nandra K., 2017a, *MNRAS*, 464, 840
 Ponti G., George E., Scaringi S., et al., 2017b, *MNRAS*, accepted (arXiv:1703.03410)
 Ponti G., Morris M. R., Terrier R., Goldwurm A., 2013, *Cosmic Rays in Star-Forming Environments*, 34, 331

made public available. Interested readers can also contact C. Jin directly for these models.

- Predehl P., Schmitt J. H. M. M., Snowden S. L., Truemper, J. 1992, *Science*, 257, 935
- Predehl P., Schmitt J. H. M. M., 1995, *A& A*, 293, 889
- Rea N., Esposito P., Pons J. A., et al., 2013, *ApJL*, 775, L34
- Read A. M., Rosen S. R., Saxton R. D., Ramirez J., 2011, *A& A*, 534, A34
- Reid M. J., Menten K. M., Zheng X. W., et al., 2009, *ApJ*, 700, 137-148
- Reynolds M., Kennea J., Degenaar N., Wijnands R., Miller J., 2016, *The Astron. Telegram*, 8649
- Rolf D. P., 1983, *Nature*, 302, 46
- Roman-Duval J., Jackson J. M., Heyer M., Rathborne J., Simon R., 2010, *ApJ*, 723, 492
- Sato M., Wu Y. W., Immer K., et al., 2014, *ApJ*, 793, 72
- Serabyn E., Guesten R., 1987, *A& A*, 184, 133
- Schultheis M., Chen B. Q., Jiang B. W., et al., 2014, *A& A*, 566, A120
- Seward F. D., Smith R. K., 2013, *ApJ*, 769, 17
- Shannon R. M., Johnston S., 2013, *MNRAS*, 435, L29
- Sicheneder E., Dexter J., 2016, *arXiv:1612.04819*
- Simon R., Rathborne J. M., Shah R. Y., Jackson J. M., Chambers E. T., 2006, *ApJ*, 653, 1325
- Smith R. K., 2008, *ApJ*, 681, 343-349
- Smith R. K., Edgar R. J., Shafer R. A., 2002, *ApJ*, 581, 562
- Smith R. K., Valencic L. A., Corrales L., 2016, *ApJ*, 818, 143
- Snowden S. L., Kuntz K. D., 2011, *Bulletin of the American Astronomical Society*, 43, 344.17
- Soldi S., Clavel M., Goldwurm A., et al., 2014, *The Galactic Center: Feeding and Feedback in a Normal Galactic Nucleus*, 303, 126
- Spitler L. G., Lee K. J., Eatough R. P., et al., 2014, *ApJL*, 780, L3
- Tan J. C., Draine B. T., 2004, *ApJ*, 606, 296
- Tremaine S., Gebhardt K., Bender R., et al., 2002, *ApJ*, 574, 740
- Trümper J., Schönfelder V., 1973, *A& A*, 25, 445
- Ueda Y., Honda K., Takahashi H., et al., 2010, *ApJ*, 713, 257
- Urban O., Werner N., Simionescu A., Allen S. W., Böhringer H., 2011, *MNRAS*, 414, 2101
- Valencic L. A., Smith R. K., 2008, *ApJ*, 672, 984-995
- Valencic L. A., Smith R. K., 2015, *ApJ*, 809, 66
- Valencic L. A., Smith R. K., Dwek E., Graessle D., Dame T. M., 2009, *ApJ*, 692, 502
- Vasilopoulos G., Petropoulou M., 2016, *MNRAS*, 455, 4426
- Vallée J. P., 2014, *AJ*, 148, 5
- Verner D. A., Ferland G. J., Korista K. T., Yakovlev D. G., 1996, *ApJ*, 465, 487
- Wang Q. D., Gotthelf E. V., Lang C. C., 2002, *Nature*, 415, 148
- Weingartner J. C., Draine B. T., 2001, *ApJ*, 548, 296
- Williams M. J., Bureau M., Cappellari M., 2010, *MNRAS*, 409, 1330
- Wilms J., Allen A., McCray R., 2000, *ApJ*, 542, 914
- Wucknitz O., 2015, *arXiv:1501.04510*
- Xiang J., Lee J. C., Nowak M. A., 2007, *ApJ*, 660, 1309
- Xiang J., Zhang S. N., Yao, Y., 2005, *ApJ*, 628, 769
- Xiang J., Lee J. C., Nowak M. A., Wilms, J., 2011, *ApJ*, 738, 78
- Xu Y., McCray R., Kelley R., 1986, *Nature*, 319, 652
- Zhang S., Hailey C. J., Mori K., et al., 2015, *ApJ*, 815, 132
- Zubko V., Dwek E., Arendt R. G., 2004, *ApJS*, 152, 211

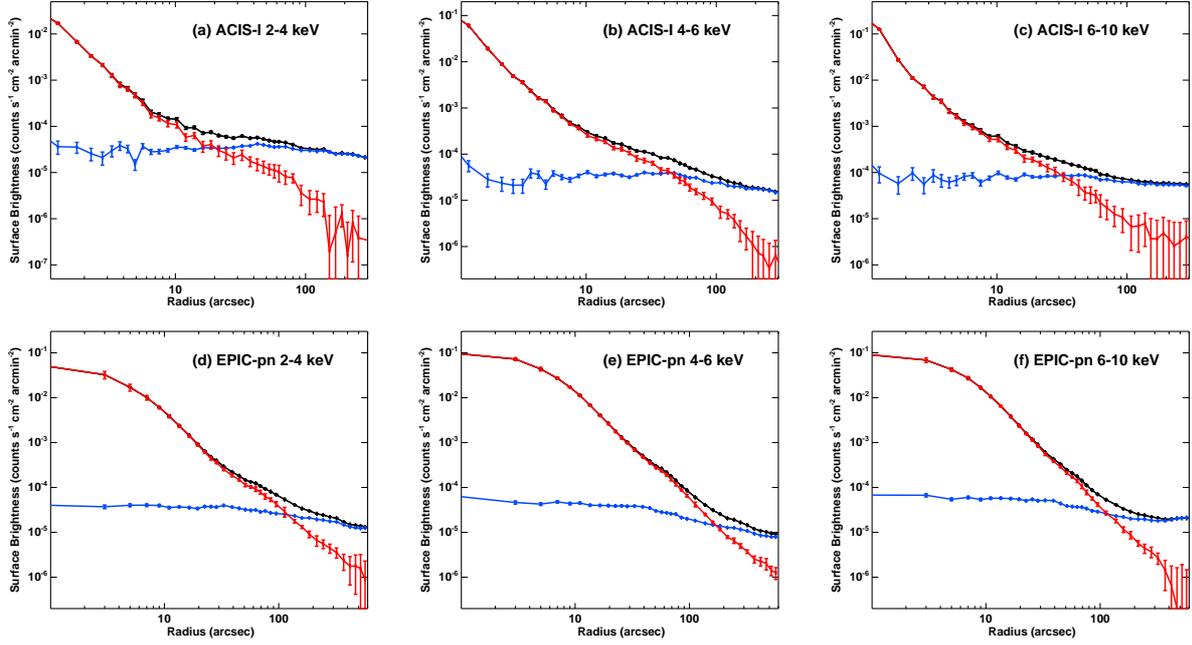


Figure A1. Combined radial profiles of AX J1745.6-2901 in *Chandra* and *XMM-Newton* in three energy bands before background subtraction (black) and after (red). The blue data points show the radial profile of the diffuse emission plus detector background as measured during quiescent periods of AX J1745.6-2901.

APPENDIX A: COMBINED RADIAL PROFILES OF AX J1745.6-2901 AND THE BACKGROUND EMISSION

APPENDIX B: LIST OF BACKGROUND OBSERVATIONS

This paper has been typeset from a $\text{\TeX}/\text{\LaTeX}$ file prepared by the author.

Table B1. List of *XMM-Newton* and *Chandra* observations used to determine the emission underneath the halo of AX J1745.6-2901 when it was in the quiescent state (i.e. detector background plus diffuse emission). ‘Exp’ and ‘CL-Exp’ are the total exposure time and clean exposure time without background flares. θ_{off} is the off-axis angle of AX J1745.6-2901.

<i>Chandra</i>						
ObsID	Start-Time	Target	ACIS-Grating	Exp (ks)	CL-Exp (ks)	θ_{off} (arcmin)
242	1999-09-21	Sgr A*	ACIS-I (None)	45.9	42.2	1.14
2943	2002-05-22	Sgr A*	ACIS-I (None)	37.7	37.7	1.73
3663	2002-05-24	Sgr A*	ACIS-I (None)	38.0	38.0	1.73
4683	2004-07-05	Sgr A*	ACIS-I (None)	49.5	49.5	1.23
4684	2004-07-06	Sgr A*	ACIS-I (None)	49.5	49.5	1.23
5950	2005-07-24	Sgr A*	ACIS-I (None)	48.5	48.5	1.21
5951	2005-07-27	Sgr A*	ACIS-I (None)	44.6	44.6	1.21
5952	2005-07-29	Sgr A*	ACIS-I (None)	45.3	43.5	1.20
5953	2005-07-30	Sgr A*	ACIS-I (None)	45.4	45.4	1.20
5954	2005-08-01	Sgr A*	ACIS-I (None)	17.9	17.9	1.20
6363	2006-07-17	Sgr A*	ACIS-I (None)	29.8	29.8	1.21
10556	2009-05-18	Sgr A*	ACIS-I (None)	112.5	112.5	1.82
11843	2010-05-13	Sgr A*	ACIS-I (None)	78.9	78.9	1.81
13016	2011-03-29	Sgr A*	ACIS-I (None)	17.8	17.8	1.77
13017	2011-03-31	Sgr A*	ACIS-I (None)	17.8	17.8	1.77
14941	2013-04-06	Sgr A*	ACIS-I (None)	19.8	19.8	1.78
14942	2013-04-14	Sgr A*	ACIS-I (None)	19.8	19.8	1.78
combined	–	–	–	718.8	713.3	–
13850	2012-02-10	Sgr A*	ACIS-S (HETG)	59.3	59.0	1.43
14392	2012-02-11	Sgr A*	ACIS-S (HETG)	58.4	58.0	1.43
14394	2012-02-10	Sgr A*	ACIS-S (HETG)	17.8	17.8	1.43
14393	2012-02-11	Sgr A*	ACIS-S (HETG)	41.0	40.9	1.43
13856	2012-03-15	Sgr A*	ACIS-S (HETG)	39.5	39.3	1.43
13857	2012-03-17	Sgr A*	ACIS-S (HETG)	39.0	39.0	1.43
13854	2012-03-20	Sgr A*	ACIS-S (HETG)	22.8	22.5	1.43
14414	2012-03-23	Sgr A*	ACIS-S (HETG)	19.8	19.8	1.43
14427	2012-05-06	Sgr A*	ACIS-S (HETG)	79.0	79.0	1.48
13848	2012-05-09	Sgr A*	ACIS-S (HETG)	96.9	96.9	1.48
13849	2012-05-11	Sgr A*	ACIS-S (HETG)	176.4	176.4	1.48
13846	2012-05-16	Sgr A*	ACIS-S (HETG)	55.5	54.0	1.48
combined	–	–	–	705.5	702.7	–
<i>XMM-Newton</i>						
ObsID	Start-Time	Target	pn-Filter	Exp (ks)	CL-Exp (ks)	θ_{off} (arcmin)
0554750401	2009-04-01	Sgr A*	PFW-Medium	39.9	32.0	2.53
0554750501	2009-04-03	Sgr A*	PFW-Medium	44.3	38.1	2.57
0554750601	2009-04-05	Sgr A*	PFW-Medium	39.1	31.8	2.58
0604300601	2011-03-28	Sgr A*	PFW-Medium	48.8	27.3	2.52
0604300801	2011-04-01	Sgr A*	PFW-Medium	48.8	33.9	2.61
0604300901	2011-04-03	Sgr A*	PFW-Medium	46.9	18.8	2.58
0604301001	2011-04-05	Sgr A*	PFW-Medium	50.7	27.6	2.58
0658600101	2011-08-31	Sgr A*	PFW-Medium	49.9	48.1	2.04
0658600201	2011-09-01	Sgr A*	PFW-Medium	53.2	39.3	1.99
0674600601	2012-03-13	Sgr A*	PFW-Medium	21.5	8.4	2.50
0674600701	2012-03-15	Sgr A*	PFW-Medium	15.9	7.0	2.53
0674601101	2012-03-17	Sgr A*	PFW-Medium	28.0	10.6	2.53
0674600801	2012-03-19	Sgr A*	PFW-Medium	22.9	16.0	2.53
0674601001	2012-03-21	Sgr A*	PFW-Medium	23.9	19.4	2.56
combined	–	–	–	358.3	332.1	–