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# The Compact, $\sim 1$ kpc Host Galaxy of a Quasar at a Redshift of 7.1

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## Abstract

We present Atacama Large Millimeter/submillimeter Array (ALMA) observations of the [C II] fine-structure line and the underlying far-infrared (FIR) dust continuum emission in J1120+0641, the most distant quasar currently known ( $z = 7.1$ ). We also present observations targeting the CO(2–1), CO(7–6), and [C I] 369  $\mu\text{m}$  lines in the same source obtained at the Very Large Array and Plateau de Bure Interferometer. We find a [C II] line flux of  $F_{\text{[C II]}} = 1.11 \pm 0.10 \text{ Jy km s}^{-1}$  and a continuum flux density of  $S_{227 \text{ GHz}} = 0.53 \pm 0.04 \text{ mJy beam}^{-1}$ , consistent with previous unresolved measurements. No other source is detected in continuum or [C II] emission in the field covered by ALMA ( $\sim 25''$ ). At the resolution of our ALMA observations ( $0''.23$ , or 1.2 kpc, a factor of  $\sim 70$  smaller beam area compared to previous measurements), we find that the majority of the emission is very compact: a high fraction ( $\sim 80\%$ ) of the total line and continuum flux is associated with a region 1–1.5 kpc in diameter. The remaining  $\sim 20\%$  of the emission is distributed over a larger area with radius  $\lesssim 4$  kpc. The [C II] emission does not exhibit ordered motion on kiloparsec scales: applying the virial theorem yields an upper limit on the dynamical mass of the host galaxy of  $(4.3 \pm 0.9) \times 10^{10} M_{\odot}$ , only  $\sim 20 \times$  higher than the central black hole (BH). The other targeted lines (CO(2–1), CO(7–6), and [C I]) are not detected, but the limits of the line ratios with respect to the [C II] emission imply that the heating in the quasar host is dominated by star formation, and not by the accreting BH. The star formation rate (SFR) implied by the FIR continuum is  $105\text{--}340 M_{\odot} \text{ yr}^{-1}$ , with a resulting SFR surface density of  $\sim 100\text{--}350 M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$ , well below the value for Eddington-accretion-limited star formation.

**Key words:** cosmology: observations – galaxies: active – galaxies: high-redshift – galaxies: individual (ULAS J112001.48+064124.3) – galaxies: ISM

## 1. Introduction

Luminous quasars are powered by accreting, supermassive black holes (BHs). Quasars and thus supermassive BHs have been found at high redshifts,  $z \sim 7$ , when the universe was less than a billion years old (e.g., Mortlock et al. 2011; Venemans et al. 2013). These early, supermassive BHs are generally hosted by massive galaxies that form stars at a high rate. Indeed, locally there is a relation between the mass of the bulge and the mass of the BH in its center (see, e.g., Kormendy & Ho 2013, for a review). An important question in astronomy is how the first BHs formed and whether the BH and hosting galaxy coevolved.

Over the past 15 yr, numerous surveys have established a sample of  $\sim 100$  quasars at  $z > 6$  (e.g., Fan et al. 2006; Jiang et al. 2009, 2015; Mortlock et al. 2009; Willott et al. 2010b; Venemans et al. 2013, 2015; Carnall et al. 2015; Reed et al. 2015; Bañados et al. 2016; Matsuoka et al. 2016). The most luminous of these have BHs with masses in excess of  $10^9 M_{\odot}$  (e.g., Jiang et al. 2007; Kurk et al. 2007; Willott et al. 2010a; De Rosa et al. 2011, 2014; Venemans et al. 2015; Wu et al. 2015). As the accreting BH generally dominates the

emission at rest-frame UV and optical wavelengths, observations at (sub)millimeter wavelength are needed to study the galaxies hosting these BHs. Several groups have targeted and detected  $z \gtrsim 6$  quasars with millimeter facilities such as the IRAM Plateau de Bure Interferometer (PdBI) and Atacama Large Millimeter/submillimeter Array (ALMA) (e.g., Bertoldi et al. 2003; Maiolino et al. 2005; Wang et al. 2008, 2013; Willott et al. 2013, 2015; Venemans et al. 2016). These data show that rapid BH growth is, in some cases, accompanied by very high star formation rates (SFRs) of up to  $\sim 1000 M_{\odot} \text{ yr}^{-1}$ .

In this paper we investigate the host galaxy of the most distant quasar currently known, ULAS J112001.48+064124.3 (hereafter J1120+0641), at a redshift of  $z = 7.085$  (Mortlock et al. 2011). The quasar is powered by a BH with a mass of  $(2.4 \pm 0.2) \times 10^9 M_{\odot}$  (Mortlock et al. 2011; De Rosa et al. 2014) and is accreting close to the Eddington limit (Mortlock et al. 2011; De Rosa et al. 2014; Barnett et al. 2015). The quasar host galaxy has previously been detected with the IRAM PdBI in [C II] and the dust continuum (Venemans et al. 2012). In these data the host galaxy was unresolved in a  $\sim 2''$  beam and the dynamical mass and the morphology of the line-emitting gas could not be constrained. Here we present high spatial resolution ( $0''.23$ , or 1.2 kpc) observations with ALMA

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(Section 2.1), decreasing the beam area by a factor of  $\sim 70$ . We also present observations with the PdBI (Section 2.2) and the NRAO Karl G. Jansky Very Large Array (VLA; Section 2.3) targeting the neutral and molecular gas lines CO(7–6), [C I] 369  $\mu\text{m}$ , and CO(2–1). Our results are detailed in Section 3, followed by a discussion about the implications of our findings in Section 4. A summary is presented in Section 5.

Throughout this paper, we adopt a concordance cosmology with  $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ ,  $\Omega_M = 0.3$ , and  $\Omega_\Lambda = 0.7$ , leading to a spatial scale of 5.2 proper kpc per arcsec at  $z = 7.085$ . SFRs are calculated assuming a Kroupa (2001) initial mass function.

## 2. Observations

### 2.1. ALMA Cycle 1 Observations

The host galaxy of J1120+0641 was observed with ALMA between 2014 and 2015 June (program 2012.1.00882.S). In 2014 the antennas were in a compact configuration (33–36 antennas with baselines between 20 and 650 m), and in 2015 they were in a more extended configuration, with baselines between 34 and 1574 m and a total of 38–47 antennas. J1120+0641 was observed on 2014 June 9 for 33.5 minutes (on source), on 2014 June 10 for 26.8 minutes, and on 2014 June 11 for 33.5 minutes. The combined 2014 data reached an rms noise of 0.39 mJy beam $^{-1}$  in a 20 MHz bin, and the beam size was  $0''.62 \times 0''.51$  (natural weighting). On 2015 June 26 and 27, the quasar was observed for 33.4 minutes (on source) during both days. The beam size of the combined 2015 data was  $0''.25 \times 0''.24$  (using natural weighting), and the rms noise per 20 MHz bin is 0.21 mJy beam $^{-1}$ . The full data set reached a noise of 0.15 mJy beam $^{-1}$  (20 MHz) $^{-1}$  and has a beam size of  $0''.31 \times 0''.29$  using natural weighting. Using a weighting factor of robust = 0.5 results in a slightly higher rms noise of 0.17 mJy beam $^{-1}$  per 20 MHz bin and a beam size of  $0''.23 \times 0''.22$ . This corresponds to 1.2 kpc at the redshift of the quasar.

The ALMA observations covered the redshifted [C II] line at  $\nu_{\text{obs}} = 235.07 \text{ GHz}$  with two overlapping bandpasses of 1.875 GHz each. The overlap was 20%, which resulted in frequency coverage of 3.375 GHz around the [C II] line. Two additional bandpasses of 1.875 GHz each were placed around an observed frequency of 220 GHz. The data were reduced using Common Astronomy Software Applications (CASA; McMullin et al. 2007), following standard reduction steps.

### 2.2. PdBI 3 mm Observations

The CO(7–6) ( $\nu_{\text{rest}} = 806.652 \text{ GHz}$ ) and [C I] ( $\nu_{\text{rest}} = 809.344 \text{ GHz}$ ) emission lines from J1120+0641 (redshifted to observed frequency around  $\nu_{\text{obs}} \sim 100 \text{ GHz}$ ) were targeted by the PdBI between 2011 July 4 and 2012 May 3. The observations were carried out with five to six antennas. The antennas were in the most compact (D) configuration, providing a beam with a size of  $4''.7 \times 4''.2$ . The WideX correlator that was used provided a continuous frequency coverage of 3.6 GHz and covered both CO(7–6) and [C I] emission lines in a single setup. The data were reduced using the Grenoble Image and Line Data Analysis System software package. The total time on source was 18.3 hr (six-antenna equivalent), resulting in an rms noise of 0.33 mJy beam $^{-1}$  per 20 MHz bin. The continuum rms noise, measured in an image that was created by averaging all channels that do not cover the

emission lines (resulting in a continuum bandwidth of 2.7 GHz; see Section 3.3), is 29  $\mu\text{Jy}$ .

### 2.3. VLA Observations

We observed CO(2–1) in J1120+0641 (rest frequency  $\nu_{\text{rest}} = 230.5424 \text{ GHz}$ , redshifted to  $\nu_{\text{obs}} = 28.5145 \text{ GHz}$ ) as part of VLA project 11A-285. The observations were taken on 2013 January 12 and 13 in the DnC configuration (consisting of 27 antennas). We utilized the Ka-band receivers with the A/C IF pair tuned to 32.008 GHz and the B/D IF pair tuned to 27.810 GHz. The observations were taken in full polarization mode with 16 128 MHz spectral windows (eight spectral windows per IF) and 64 2 MHz channels per spectral window.

The observations consisted of a single pointing centered on the quasar host galaxy. We used fast-switching phase calibration (Carilli & Holdaway 1999) on VLA calibrator J1058+0133 with a 4-minute cycle time, and the same source served as our bandpass calibrator. The bright source 3C 286 served as the flux calibrator for all observations. The observing time, including overhead, was 10 hr. The total time on source was 6.3 hr.

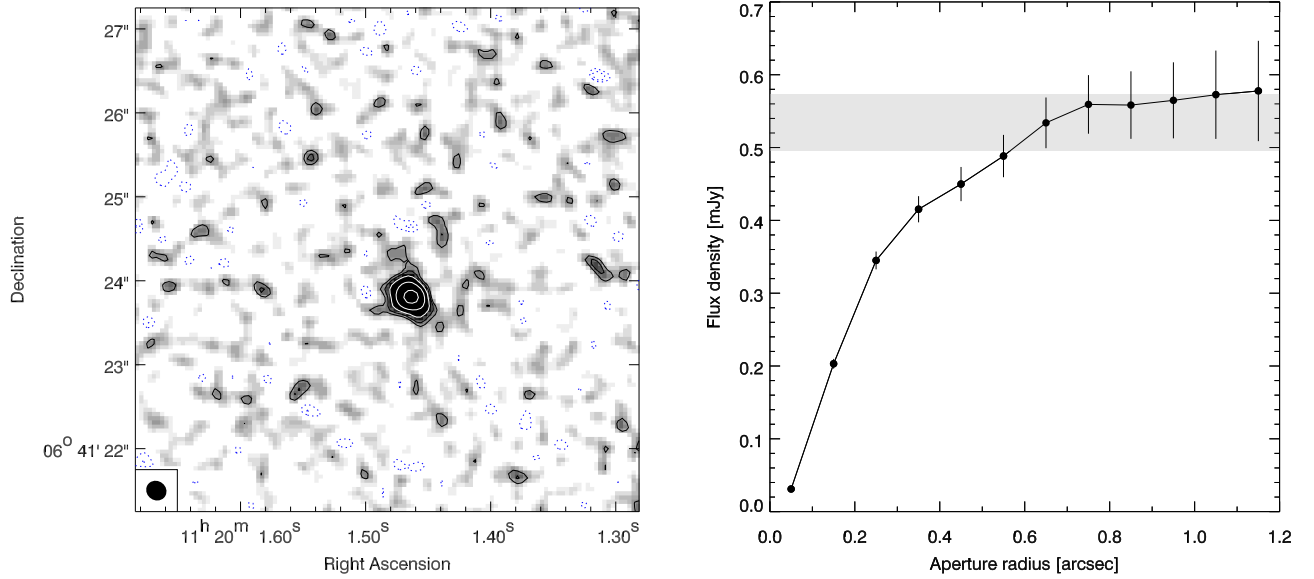
The data were reduced using the CASA package. We imaged the calibrated data using the clean algorithm with a cell size of  $0''.5$  and natural weighting, resulting in a synthesized beam of  $2''.3 \times 1''.3$ . The final cube has an rms of 15  $\mu\text{Jy beam}^{-1}$  in a 76 MHz window (800 km s $^{-1}$ ) centered on the frequency of the expected CO line. The continuum rms is 4  $\mu\text{Jy beam}^{-1}$ .

## 3. Results

### 3.1. FIR Continuum

In Figure 1 we show a 227 GHz (observed wavelength  $\lambda_{\text{obs}} = 1.3 \text{ mm}$ ) continuum map of the field that was created by averaging all channels at least 0.75 GHz away from the expected [C II] emission line. As a consequence, the reconstructed continuum emission is dominated by ALMA’s lower-sideband observations (at frequencies  $\sim 10 \text{ GHz}$  below the redshifted [C II] line). In the whole field (with a half-power beam width of  $24''.8$ ) only the quasar is detected (at signal-to-noise ratio  $[S/N] > 5$ , or  $S_{227 \text{ GHz}} > 55 \mu\text{Jy beam}^{-1}$ ). The position (R.A. =  $11^{\text{h}}20^{\text{m}}01^{\text{s}}.465$ ; decl. =  $+06^\circ 41'23''.810$ ) is approximately  $0''.5$  to the south of the position from the UK Infrared Telescope (UKIRT) Infrared Deep Sky Survey (UKIDSS) published in Mortlock et al. (2011). This difference could be due to a systematic difference in the absolute astrometric calibration between UKIDSS and ALMA. It is not possible to verify this, as there are no other sources in the field within the ALMA field-of-view beam, but similar offsets between ALMA and optical/near-infrared images have been reported (e.g., in the Hubble Ultra Deep Field; see Aravena et al. 2016; Rujopakarn et al. 2016). No offset from the UKIDSS position was seen in the PdBI data (Venemans et al. 2012). Within the uncertainties of the PdBI observations, the earlier millimeter continuum position was identical to that of the UKIDSS position.

The peak flux density of the quasar host is  $S_{227 \text{ GHz}} = 0.26 \pm 0.01 \text{ mJy beam}^{-1}$ . To measure the total continuum flux density, we tapered the emission to  $1''.9$ . This resolution is similar to the beam of the original [C II] observations ( $2''.0 \times 1''.7$ ; Venemans et al. 2012). In this map, the peak flux density of the host is  $S_{227 \text{ GHz}} = 0.53 \pm 0.04 \text{ mJy beam}^{-1}$ . A similar value was



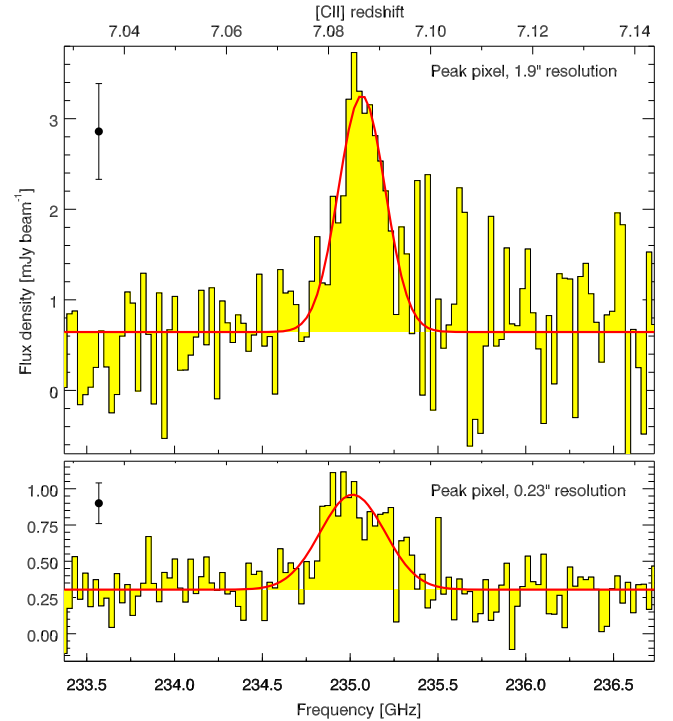
**Figure 1.** Left: 227 GHz ( $\lambda_{\text{obs}} = 1.3$  mm) continuum map of the field of J1120+0641. To create the map, all channels at least 0.75 GHz away from the position of the [C II] emission line ( $\nu_{\text{[C II]}} \approx 235$  GHz; see Figure 2) were averaged. The beam of  $0''.25 \times 0''.23$  is shown in the bottom left corner. The  $1\sigma$  rms noise in this map is  $11 \mu\text{Jy beam}^{-1}$ . The blue dashed contours are  $-3\sigma$  and  $-2\sigma$ ; the black solid contours are  $+2\sigma$  and  $+3\sigma$ ; the white solid contours are  $+5\sigma$ ,  $+10\sigma$ , and  $+20\sigma$ . Only the central quasar is detected ( $S/N \sim 26$ ). Right: estimated flux density of the central continuum source as a function of aperture radius. The flux density reaches a roughly constant value for apertures with radius  $> 0''.7$ , and the value is consistent within the errors with the peak flux density when observed with a beam of  $1''.9$  (the gray band).

derived when performing aperture photometry (see Figure 1). For apertures larger than  $\sim 0''.7$  the flux density is roughly constant and we measure a total flux density of  $S_{227 \text{ GHz}} = 0.56 \pm 0.04$  mJy (see Section 3.4). Note that due to the shape of the far-infrared (FIR) continuum, the continuum flux density is higher around the [C II] line (Figure 2; see also, e.g., Section 4.2 and Venemans et al. 2016). From the spectrum of the [C II] line (around an observed frequency of 235 GHz; Figure 2) we measure  $S_{235 \text{ GHz}} = 0.64 \pm 0.08$  mJy  $\text{beam}^{-1}$ , which is consistent with the published value of  $S_{235 \text{ GHz}} = 0.61 \pm 0.16$  mJy  $\text{beam}^{-1}$  (Venemans et al. 2012).

We also created a map of the 100 GHz ( $\lambda_{\text{obs}} \approx 3$  mm) continuum emission from the PdBI data, using the channels that are expected to be line-free. At the position of the 227 GHz continuum source we obtain a tentative  $3\sigma$  detection (flux density of  $S_{100 \text{ GHz}} = 86 \pm 29 \mu\text{Jy}$ ). Based on the 1 mm continuum detection in the ALMA data, assuming an intrinsic dust temperature of 30–50 K (see Sections 4.1 and 4.2), an emissivity index of  $\beta = 1.6$ , and taking the cosmic microwave background (CMB) into account (see, e.g., da Cunha et al. 2013; Venemans et al. 2016, Section 4.1), the expected flux density at 100 GHz is 42–55  $\mu\text{Jy}$ , consistent with our low-S/N measurement. The nondetection of the continuum in the VLA data ( $S_{30 \text{ GHz}} < 11 \mu\text{Jy}$ , Table 1) is consistent with J1120+0641 being radio-quiet (Momjian et al. 2014).

### 3.2. [C II] Emission Line

The [C II] emission line is detected at high S/N in the ALMA data. To get an estimate of the total line flux, we tapered the data cube to a beam of  $1''.9$ . We show the spectrum from the peak pixel in this cube, together with a Gaussian fit, in Figure 2 (top). The redshift of the [C II] line is  $z_{\text{[C II]}} = 7.0851 \pm 0.0005$ , the peak flux density is  $f_p = 2.60 \pm 0.25$  mJy  $\text{beam}^{-1}$ , and the FWHM is  $\text{FWHM}_{\text{[C II]}} = 400 \pm 45 \text{ km s}^{-1}$  (see also Table 1). The line flux of  $F_{\text{[C II]}} = 1.11 \pm 0.10 \text{ Jy km s}^{-1}$  is consistent



**Figure 2.** Top: spectrum extracted from the peak pixel in the ALMA data cube tapered to a spatial resolution of  $1''.9$ . The bins are 30 MHz, which corresponds to  $\sim 38 \text{ km s}^{-1}$ . The typical  $1\sigma$  uncertainty per bin of  $0.52 \text{ mJy beam}^{-1}$  is plotted in the top left corner. The red line is a Gaussian+constant fit to the spectrum. Bottom: same as above, but this spectrum is extracted from the peak pixel in the data cube with the full spatial resolution ( $0''.23$ ). The rms noise is  $0.14 \text{ mJy beam}^{-1}$  per 30 MHz bin.

with the value of  $F_{\text{[C II]}} = 1.03 \pm 0.14 \text{ Jy km s}^{-1}$  published by Venemans et al. (2012), while the line width measured in the ALMA data is a factor of  $1.69 \pm 0.32$  larger than the earlier value that was based on lower-S/N data. The [C II] rest-frame



**Table 1**  
Observed Properties of J1120+0641

R.A. (J2000)	11 <sup>h</sup> 20 <sup>m</sup> 01 <sup>s</sup> .465
Decl. (J2000)	+06° 41' 23" 810
$z_{[\text{C II}]}$	$7.0851 \pm 0.0005$
$F_{[\text{C II}]}$ (Jy km s <sup>-1</sup> )	$1.11 \pm 0.10$
$\text{FWHM}_{[\text{C II}]}$ (km s <sup>-1</sup> )	$400 \pm 45$
$S_{227 \text{ GHz}}$ (mJy)	$0.53 \pm 0.04$
$S_{100 \text{ GHz}}$ (mJy)	$0.086 \pm 0.029$
$S_{30 \text{ GHz}}$ (mJy)	$< 0.011$
$\text{EW}_{[\text{C II}]}$ ( $\mu\text{m}$ )	$0.90 \pm 0.14$
$F_{\text{CO}(2-1)}$ (Jy km s <sup>-1</sup> )	$< 0.034$
$F_{\text{CO}(7-6)}$ (Jy km s <sup>-1</sup> )	$< 0.20$
$F_{[\text{C I}]}$ (Jy km s <sup>-1</sup> )	$< 0.20$
Size continuum <sup>a</sup> (arcsec <sup>2</sup> )	$(0.23 \pm 0.03) \times (0.16 \pm 0.03)$
Size continuum <sup>a</sup> (kpc <sup>2</sup> )	$(1.24 \pm 0.14) \times (0.83 \pm 0.14)$
Size [C II] emission <sup>a</sup> (arcsec <sup>2</sup> )	$(0.31 \pm 0.05) \times (0.27 \pm 0.05)$
Size [C II] emission <sup>a</sup> (kpc <sup>2</sup> )	$(1.65 \pm 0.29) \times (1.44 \pm 0.26)$

**Note.**

<sup>a</sup> The sizes listed here are diameters and are derived for the central component that contains around 80% of the total flux density (see Section 3.4).

equivalent width is  $\text{EW}_{[\text{C II}]} = 0.90 \pm 0.14 \mu\text{m}$ . This is only  $\sim 30\%$  lower than that of local starburst galaxies (e.g., Díaz-Santos et al. 2013; Sargsyan et al. 2014).

Also shown in Figure 2 (bottom) is the spectrum of the peak pixel in the high-resolution data cube (with a beam of  $0''.23 \times 0''.22$ ). The [C II] line is significantly fainter in the center, with a peak flux density  $f_p = 0.66 \pm 0.06 \text{ mJy beam}^{-1}$  and a flux of  $F_{[\text{C II}]} = 0.39 \pm 0.04 \text{ Jy km s}^{-1}$ , which is 35% of the total line flux. This means that the [C II] line is spatially resolved in our data (see Section 3.4). The peak of the [C II] emission coincides with the peak of the continuum emission (Figure 3). The line width is broader in the center, with an  $\text{FWHM}_{[\text{C II}]} = 555 \pm 60 \text{ km s}^{-1}$ . We will come back to this in Section 4.4.

### 3.3. Limits on CO and [C I] Emission

We searched for CO(7–6) and [C I] emission in the PdBI 3 mm data. After creating a continuum-subtracted data cube using the CASA task “uvcontsub,” we averaged the channels where we expected the CO(7–6) and [C I] lines, based on the [C II] redshift, over  $400 \text{ km s}^{-1}$  (the [C II] line FWHM; see Table 1). In the resulting maps (not shown here) no significant ( $> 3\sigma$ ) line emission was detected at the location of the quasar host galaxy. For the CO(7–6) and [C I] lines we derived  $3\sigma$  upper limits on the line strength of  $F_{\text{CO}(7-6), [\text{C I}]} < 0.20 \text{ Jy km s}^{-1}$ .<sup>13</sup>

Similarly, we derived an upper limit on the CO(2–1) line from the VLA data. Averaging the data over  $400 \text{ km s}^{-1}$  resulted in a map with no significant emission at the position of the quasar host. The rms of the image was  $0.021 \text{ mJy}$ , and the  $3\sigma$  upper limit on the CO(2–1) line was  $F_{\text{CO}(2-1)} < 0.034 \text{ Jy km s}^{-1}$ . We will discuss the implications of these nondetections in Section 4.1.

<sup>13</sup> These upper limits are corrected to account for flux in the outer line wings that are not included when averaging over the FWHM of a Gaussian line.

### 3.4. Size and Structure of the Emission Regions

**Continuum:** to estimate the size of the continuum-emitting region, we fitted a 2D Gaussian to the continuum map using the CASA task “imfit.” The quasar host galaxy is marginally resolved, and we derive a deconvolved size with an FWHM of  $(0.23 \pm 0.03) \times (0.16 \pm 0.03) \text{ arcsec}^2$ , or  $(1.24 \pm 0.14) \times (0.83 \pm 0.14) \text{ kpc}^2$ . The integrated flux density of this central component is  $0.43 \pm 0.03 \text{ mJy}$ , which is  $\sim 80\%$  of the peak measured in the tapered continuum image (Section 3.1). The remaining 20% of the total continuum flux density comes from a more extended region, with a size  $\lesssim 0''.6$  in radius ( $\lesssim 3 \text{ kpc}$ ; see Figure 1).

**[C II] line:** in order to measure the size of the [C II]-emitting region, we first created a continuum-subtracted data cube: we fitted a first-order polynomial to the channels at least  $0.75 \text{ GHz}$  away from the line center and subtracted this continuum from the data using the CASA task “uvcontsub.” A map of the [C II] emission was produced by averaging the cube over  $600 \text{ km s}^{-1}$  ( $1.5 \times \text{FWHM}_{[\text{C II}]}$ ) around the peak of the [C II] emission. The result is shown in Figure 3. From a 2D Gaussian fit to this map we obtained a deconvolved size of the [C II]-emitting region of  $(0.31 \pm 0.05) \times (0.27 \pm 0.05) \text{ arcsec}^2$  in diameter, which corresponds to  $(1.65 \pm 0.29) \times (1.44 \pm 0.26) \text{ kpc}^2$ . The area of the resolved [C II]-emitting region of  $1.9 \pm 0.5 \text{ kpc}^2$  is larger than the area of the continuum region ( $0.8 \pm 0.2 \text{ kpc}^2$ ). A larger extent of [C II] emission compared to that of the continuum emission has also been reported in other  $z \gtrsim 6$  quasar host galaxies (e.g., Wang et al. 2013; Venemans et al. 2016). The total flux density of the resolved component is  $1.35 \pm 0.15 \text{ mJy}$ , which corresponds to a flux of  $0.81 \pm 0.09 \text{ Jy km s}^{-1}$ .

We also performed aperture photometry on the [C II] image. We recovered all the flux measured in the tapered spectrum within a radius of  $0''.8$  ( $\sim 4.3 \text{ kpc}$ ). At larger aperture radii we tentatively detected additional flux. We estimated that up to 20% of additional line flux might be present at scales up to  $\sim 7 \text{ kpc}$  from the quasar, although the significance is low ( $1\sigma$ ).

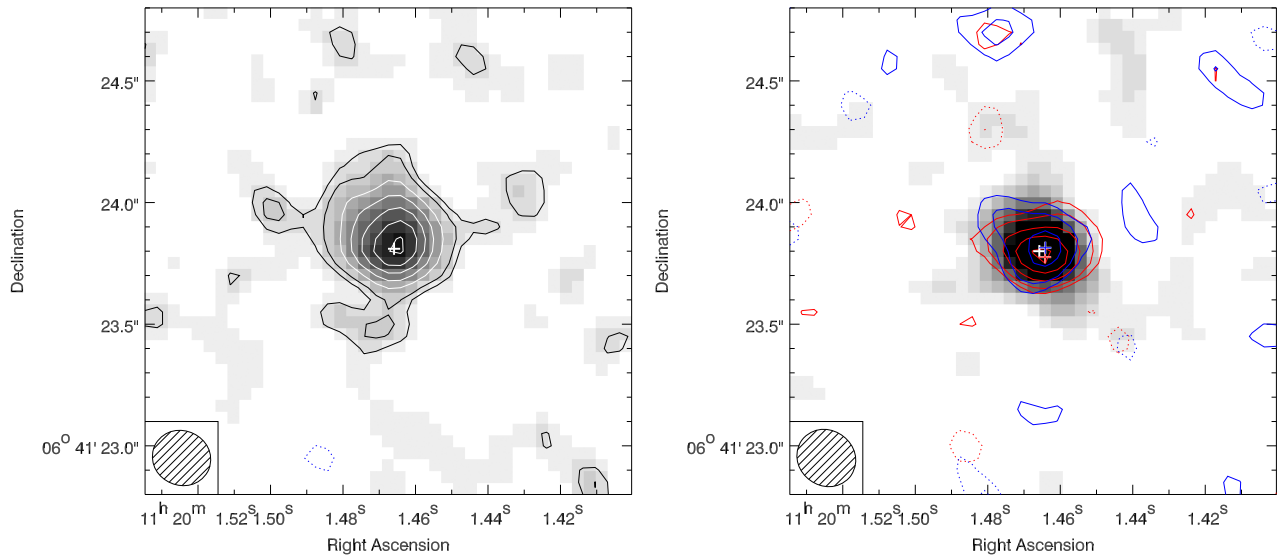
Bright [C II] emission in  $z \gtrsim 6$  quasar hosts often shows indications of rotation (e.g., Wang et al. 2013; Willott et al. 2013; Venemans et al. 2016). To investigate whether the gas in J1120+0641 displays ordered motion, we separately mapped the blue and red sides of the emission line; see Figure 3. The blue- and redshifted emission peaks coincide with the continuum emission. It therefore appears that the gas traced by [C II] emission does not show rotation on scales of  $\gtrsim 1 \text{ kpc}$ . We will discuss the implications of this in Section 4.4.

A summary of our results, described in Sections 3.1–3.4, is listed in Table 1.

## 4. Discussion

### 4.1. Origin of the Heating Radiation

We now compare our (limits on the) emission-line ratios to models to constrain the physical parameters of the emitting gas (see, e.g., Kaufman et al. 1999; Meijerink & Spaans 2005; Meijerink et al. 2007). In particular, the line ratio  $[\text{C II}]/[\text{C I}]$  can be used to determine the dominant source of radiation, the hard X-ray radiation of the accreting, supermassive BH (an X-ray-dominated region, or XDR), or UV radiation from hot stars (a photon-dominated region, or PDR) (see Figure 4). To calculate the intrinsic line ratios, we first need to determine the intrinsic luminosity of the emission lines. Because we are



**Figure 3.** Left: map of the [C II] emission in J1120+0641 (shown with gray scale and contours), averaged over  $1.5 \times \text{FWHM}_{[\text{C II}]}$  (600 km s<sup>-1</sup>, or 470 MHz). The 1 $\sigma$  rms noise in this map is 42  $\mu\text{Jy}$ . The blue dashed contours are  $-3\sigma$  and  $-2\sigma$ ; the black solid contours are  $+2\sigma$  and  $+3\sigma$ ; the white solid contours are  $[5, 7, 9, 11, 13] \times \sigma$ . The size of the beam is shown in the bottom left corner. The white plus sign shows where the continuum emission of the quasar host peaks. Right: the red and blue sides of the emission line are shown in contours, averaged over 265 km s<sup>-1</sup> centered at  $+265$  km s<sup>-1</sup> (red contours) and  $-265$  km s<sup>-1</sup> (blue contours) from the line peak. Contour levels are  $-2\sigma, +2\sigma, +3\sigma, +5\sigma, +7\sigma$ , and  $+9\sigma$ , with  $\sigma \sim 57 \mu\text{Jy}$ . The gray scale is a representation of the continuum map. The white, red, and blue plus signs indicate the peak of the total, redshifted, and blueshifted [C II] emission, respectively.

measuring the flux of the lines against the CMB, the intrinsic luminosity will depend on the excitation temperature (da Cunha et al. 2013):  $F_{\text{in}} = F_{\text{obs}} / \{1 - B_{\nu}[T_{\text{CMB}}(z = 7.0851)] / B_{\nu}[T_{\text{ex}}]\}$ , where  $F_{\text{obs}}$  and  $F_{\text{in}}$  are the observed and intrinsic line flux,  $B_{\nu}$  the Planck function at the rest-frame frequency  $\nu$  of the line,  $T_{\text{CMB}}(z = 7.0851)$  the temperature of the CMB at redshift  $z = 7.0851$  ( $\approx 22$  K), and  $T_{\text{ex}}$  the excitation temperature.

If collisions dominate the excitation, then the excitation temperature is set by the kinetic temperature of the gas. In this paper we further assume thermodynamic equilibrium between the dust and the gas, i.e.,  $T_{\text{gas}} = T_{\text{dust}}$ . This assumption is motivated by the study of dust and CO emission in the host galaxy of quasar J1148+5251 at  $z = 6.42$  in which  $T_{\text{ex}} \approx T_{\text{dust}}$  (Beelen et al. 2006; Riechers et al. 2009; Stefan et al. 2015). We also assume that the dust has a constant temperature throughout the host galaxy. To explore the range of luminosities and line ratios in J1120+0641, we derive intrinsic luminosities of the emission lines for various temperatures that are found in the literature (Table 2). Several studies of  $z \sim 6$  quasar host galaxies (e.g., Wang et al. 2013; Willott et al. 2013, 2015) implement a dust temperature of  $T_{\text{d}} = 47$  K as derived by Beelen et al. (2006) for distant luminous quasars. A study by Priddey & McMahon (2001) found an average dust temperature in quasars of  $T_{\text{d}} = 41$  K. We also computed the line luminosities in the case of a lower temperature of  $T_{\text{d}} = 30$  K (Walter et al. 2011; Venemans et al. 2016). Alternatively, the gas temperature could be much higher than that of the dust,  $T_{\text{gas}} \gg 100$  K (e.g., Contursi et al. 2013), and the effect of the CMB becomes negligible (the “no CMB” column in Table 2).

We can compare the limits on the [C II]/[C I] ratio in J1120+0641 (Table 2) to those from PDR and XDR models (Figure 4). In PDR models, the [C II]/[C I] ratio covers a large range from  $\sim 3$  to 50, depending on the density and strength of the UV radiation field (Kaufman et al. 1999; Meijerink et al. 2007). On the other hand, in XDRs the [C II]/[C I] ratio

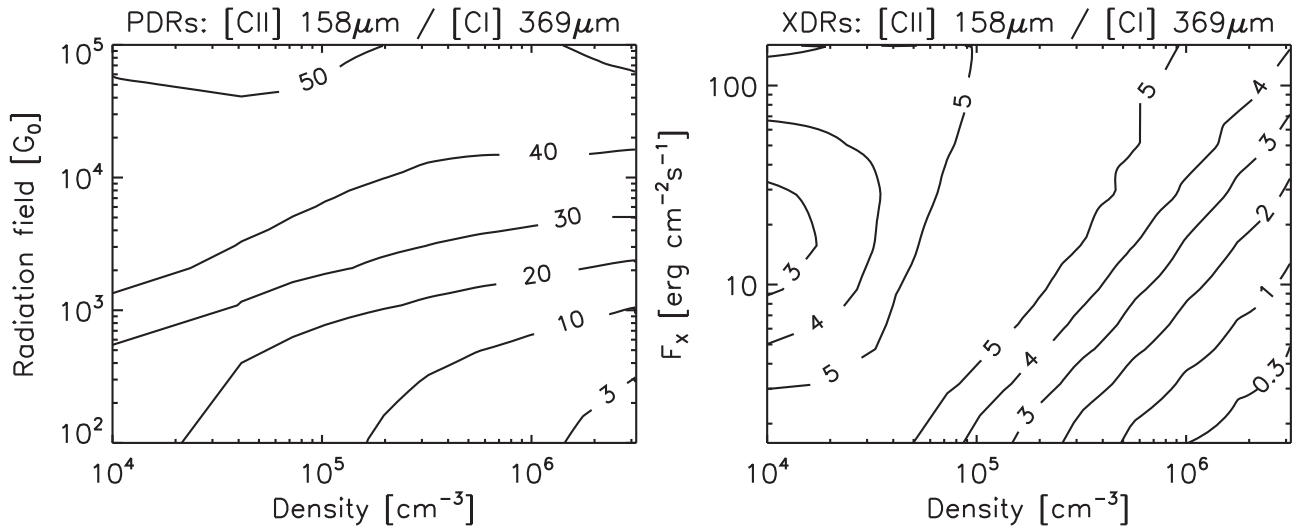
is generally lower than in a PDR, with a maximum of around  $\sim 6$  (Meijerink et al. 2007). We measure a lower limit on the [C II]/[C I] line ratio of  $\sim 8.7$ . Therefore, based on the XDR models, we can conclude that the radiation illuminating the gas is dominated by hard X-ray radiation from the accreting BH, but instead should mainly come from hot stars.

#### 4.2. Infrared Luminosity and SFRs

To compute the FIR luminosity, we have to assume a shape of the dust emission. The cold dust spectral energy distribution (SED) of distant quasar host galaxies is often parameterized as an optically thin modified blackbody (e.g., Priddey & McMahon 2001; Beelen et al. 2006; Leipski et al. 2014) with a dust temperature  $T_{\text{d}}$  and emissivity index  $\beta$ . Adopting  $T_{\text{d}} = 47$  K and  $\beta = 1.6$  (Beelen et al. 2006), integrating the dust SED from 42.5 to 122.5  $\mu\text{m}$ , and taking the CMB into account, we derive an FIR luminosity of  $L_{\text{FIR}} = (1.5 \pm 0.1) \times 10^{12} L_{\odot}$ . For  $T_{\text{d}} = 41$  K and  $\beta = 1.95$  (Priddey & McMahon 2001), we derive a similar value for the FIR luminosity. A lower dust temperature of 30 K (Venemans et al. 2016) and  $\beta = 1.6$  results in  $L_{\text{FIR}} = (5.6 \pm 0.4) \times 10^{11} L_{\odot}$  (see Table 2 for a summary).

In Section 4.1 we concluded that the gas is predominantly heated by UV radiation from stars. If that also applies to the dust, then we can use the infrared luminosity to constrain the SFR of the host galaxy. This is supported by Barnett et al. (2015), who analyzed the full SED of J1120+0641 and concluded that the emission around 235 GHz in the rest frame is dominated by a cool dust component and not by the accreting BH.

To estimate the SFR from the continuum detection, we first integrated the modified blackbody from 8 to 1000  $\mu\text{m}$  to obtain the total infrared luminosity  $L_{\text{TIR}}$ . Depending on the parameters of the modified blackbody, we derive total infrared luminosities ranging from  $L_{\text{FIR}} = (7.7 \pm 0.6) \times 10^{11} L_{\odot}$  to



**Figure 4.** Left: [C II]/[C I] line ratio as a function of density and radiation field in the case of a PDR, adapted from Meijerink & Spaans (2005) and Meijerink et al. (2007). Right: same line ratio, but this time for an XDR, also adapted from Meijerink & Spaans (2005) and Meijerink et al. (2007). The maximum ratio expected in an XDR is  $\sim 6$ , independent of (realistic) radiation strength and density. According to these models, our measured lower limit of  $\sim 8.7$  (Table 2) excludes that the X-ray radiation of the quasar dominates the gas heating.

$L_{\text{FIR}} = (2.1 \pm 0.2) \times 10^{12} L_{\odot}$  (Table 2). We then applied the local scaling relation between the total infrared luminosity and SFR from Murphy et al. (2011):  $\text{SFR}_{\text{TIR}} = L_{\text{TIR}} / 6.7 \times 10^9$  with  $\text{SFR}_{\text{TIR}}$  in units of  $M_{\odot} \text{ yr}^{-1}$  and  $L_{\text{TIR}}$  in units of  $L_{\odot}$ . We estimate an SFR of  $105\text{--}340 M_{\odot} \text{ yr}^{-1}$  (Table 2), where the main uncertainty is the shape of the dust SED. In Section 3.4 we concluded that  $\sim 80\%$  of the continuum emission originates from a region that measures  $1.2 \text{ kpc} \times 0.8 \text{ kpc}$  in diameter, or  $0.8 \text{ kpc}^2$ . This means that the star formation rate density (SFRD) is  $\sim 100\text{--}350 M_{\odot} \text{ yr}^{-1} \text{ kpc}^2$ . The lower limit is an order of magnitude smaller than the SFRD derived for the more FIR-luminous bright quasar J1148+5251 at  $z = 6.42$  (Walter et al. 2009), in which the SFRD approaches the Eddington limit for star formation.

Alternatively, we can compute the SFR from the luminosity of the [C II] line. Applying the [C II] SFR conversion for high-redshift sources from De Looze et al. (2014), we derive an  $\text{SFR}_{[\text{C II}]} = 70\text{--}700 M_{\odot} \text{ yr}^{-1}$  (Table 2). Within the large uncertainties, this value is consistent with the one derived from the FIR continuum emission.

#### 4.3. Dust and Gas Mass

Following Venemans et al. (2012), we derived the dust mass from the FIR luminosity assuming a temperature and a dust mass opacity coefficient:  $M_{\text{dust}} \sim S_{\nu} / [\kappa_{\lambda} \times B_{\nu}(T_{\text{d}})]$  (e.g., Hildebrand 1983) with  $S_{\nu}$  the continuum flux density at rest-frame frequency  $\nu$  and the dust mass opacity  $\kappa_{\lambda} = 0.77(850 \mu\text{m}/\lambda)^{\beta} \text{ cm}^2 \text{ g}^{-1}$  (Dunne et al. 2000). For the range of temperatures and emissivity indices considered in this paper, our best estimate for the dust mass in J1120+0641 is  $(0.8\text{--}4) \times 10^8 M_{\odot}$  (Table 2).

A limit on the molecular gas mass can be derived from the upper limit on the CO luminosity  $L'_{\text{CO}(2-1)}$  (in units of  $\text{K km s}^{-1} \text{ pc}^2$ ) using  $L'_{\text{CO}(1-0)} \approx L'_{\text{CO}(2-1)}$  (e.g., Carilli & Walter 2013) and applying a conversion factor of  $\alpha = M_{\text{gas}} / L'_{\text{CO}(1-0)} = 0.8 M_{\odot} (\text{K km s}^{-1} \text{ pc}^2)^{-1}$  as found for nearby ULIRGs (e.g., Downes & Solomon 1998). Taking the effects of the CMB into account, our VLA upper limit on the CO(2–1) line flux results in an upper limit on the CO luminosity of  $L'_{\text{CO}(1-0)} < 4 \times 10^{10} \text{ K km s}^{-1} \text{ pc}^2$

(see Table 2) and a gas mass limit of  $M_{\text{gas}} < 3 \times 10^{10} M_{\odot}$ . Alternatively, we can use the PdBI upper limit on the CO(7–6) line flux to derive  $L'_{\text{CO}(1-0)}$ . Applying a CO excitation ladder that is similar to that in quasar J1146+5251 at  $z = 6.42$  (e.g., Stefan et al. 2015), we derive an upper limit of  $M_{\text{gas}} \lesssim 2 \times 10^{10} M_{\odot}$ .

If we combine these upper limits with our estimates of the dust mass, we derive gas-to-dust mass ratio limits of  $< 80\text{--}250$  (strongly depending on the dust and gas temperature; see Table 2), which are consistent with locally derived values of  $\sim 100$  (e.g., Draine et al. 2007; Sandstrom et al. 2013).

#### 4.4. Dynamical Mass Estimate

The dynamical mass of  $z > 6$  quasar hosts has often been computed by assuming that the gas is rotating in an inclined disk (e.g., Walter et al. 2003; Wang et al. 2013; Willott et al. 2015; Venemans et al. 2016). This approach was motivated by the detection of velocity gradients in the [C II]-emitting gas. From Figure 3 it is clear that in J1120+0641 there is no evidence for rotation on scales of  $\gtrsim 1 \text{ kpc}$ . Instead, we here use the virial theorem to estimate a dynamical mass of the quasar host galaxy:  $M_{\text{dyn}} = 3R\sigma^2/2G$ , with  $R$  the radius of the line-emitting region,  $\sigma$  the velocity dispersion of the gas, and  $G$  the gravitational constant. In Section 3.4 we measure a maximum radius of  $4.3 \text{ kpc}$ , and from the tapered spectrum in Figure 2 we obtain a velocity dispersion of  $\sigma = 169 \pm 18 \text{ km s}^{-1}$ , which results in an upper limit on the dynamical mass of  $M_{\text{dyn}} = (4.3 \pm 0.9) \times 10^{10} M_{\odot}$ , which is similar to dynamical masses derived for other  $z \gtrsim 6$  quasar host galaxies (e.g., Walter et al. 2009; Wang et al. 2013; Willott et al. 2015; Venemans et al. 2016). Intriguingly, the BH with a mass of <sup>14</sup>  $M_{\text{BH}} = (2.4 \pm 0.2) \times 10^9 M_{\odot}$  (De Rosa et al. 2014) already contains  $\sim 6\%$  of this dynamical mass. This fraction is about  $10\times$  higher than the mass ratio of BHs and bulges in local early-type galaxies ( $M_{\text{BH}}/M_{\text{bulge}} = 0.49\%$ ; Kormendy & Ho 2013). High BH-to-dynamical mass ratios are also found in other  $z \gtrsim 6$  quasar host galaxies, which have BH-to-

<sup>14</sup> The uncertainty quoted here does not include the systematic uncertainty of a factor of  $\sim 3$  in the method applied by De Rosa et al. (2014) to derive the mass of BH.

**Table 2**  
Derived Properties of the Host Galaxy of Quasar J1120+0641 at  $z = 7.0851$  as a Function of Temperature  $T_d$  and Emissivity Index  $\beta$

Parameter	$T_d = 47 \text{ K}, \beta = 1.6, \text{ no CMB}$	$T_d = 47 \text{ K}, \beta = 1.6$	$T_d = 41 \text{ K}, \beta = 1.95$	$T_d = 30 \text{ K}, \beta = 1.6$
$L_{\text{FIR}} (L_\odot)$	$(1.3 \pm 0.1) \times 10^{12}$	$(1.5 \pm 0.1) \times 10^{12}$	$(1.5 \pm 0.1) \times 10^{12}$	$(5.6 \pm 0.4) \times 10^{11}$
$L_{\text{TIR}} (L_\odot)$	$(1.9 \pm 0.1) \times 10^{12}$	$(2.1 \pm 0.2) \times 10^{12}$	$(1.9 \pm 0.1) \times 10^{12}$	$(7.7 \pm 0.6) \times 10^{11}$
$M_{\text{dust}} (M_\odot)$	$(8.6 \pm 0.6) \times 10^7$	$(9.6 \pm 0.7) \times 10^7$	$(7.7 \pm 0.6) \times 10^7$	$(4.2 \pm 0.3) \times 10^8$
$L_{[\text{C II}]} (L_\odot)$	$(1.3 \pm 0.1) \times 10^9$	$(1.5 \pm 0.1) \times 10^9$	$(1.6 \pm 0.1) \times 10^9$	$(2.0 \pm 0.2) \times 10^9$
$L_{[\text{C I}]} (L_\odot)$	$< 1.0 \times 10^8$	$< 1.4 \times 10^8$	$< 1.5 \times 10^8$	$< 2.3 \times 10^8$
$L_{\text{CO}(2-1)} (L_\odot)$	$< 5.0 \times 10^6$	$< 8.4 \times 10^6$	$< 9.4 \times 10^6$	$< 1.6 \times 10^7$
$L_{\text{CO}(7-6)} (L_\odot)$	$< 1.0 \times 10^8$	$< 1.4 \times 10^8$	$< 1.5 \times 10^8$	$< 2.3 \times 10^8$
$L_{[\text{C II}]} / L_{[\text{C I}]}$	$> 13.0$	$> 10.6$	$> 10.1$	$> 8.7$
$L'_{\text{CO}(1-0)} (K \text{ km s}^{-1} \text{ pc}^2)$	$< 1.3 \times 10^{10}$	$< 2.1 \times 10^{10}$	$< 2.4 \times 10^{10}$	$< 4.0 \times 10^{10}$
$L_{[\text{C II}]} / L_{\text{CO}(1-0)}^a$	$> 2200$	$> 1400$	$> 1300$	$> 1000$
$M_{\text{gas}} / M_{\text{dust}}$	$< 120$	$< 180$	$< 250$	$< 80$
$\text{SFR}_{\text{TIR}} (M_\odot \text{ yr}^{-1})$	$280 \pm 20$	$315 \pm 25$	$290 \pm 20$	$115 \pm 10$
$\text{SFR}_{[\text{C II}]} (M_\odot \text{ yr}^{-1})$	70–440	80–500	85–525	110–700

**Note.**

<sup>a</sup> Derived from the  $3\sigma$  limit on the CO(2–1) emission and assuming that the molecular gas is thermalized ( $L'_{\text{CO}(2-1)} = L'_{\text{CO}(1-0)}$ ). Taking instead the limit on the CO (7–6) emission and adopting a CO spectral line energy distribution similar to that observed in the quasar J1148+5251 at  $z = 6.42$  gives upper limits on the CO(1–0) emission that are a factor of 1.0–1.4 smaller.

dynamical mass ratios up to  $M_{\text{BH}}/M_{\text{dyn}} \lesssim 25\%$  and an average of  $\langle M_{\text{BH}}/M_{\text{dyn}} \rangle \approx 2\%$  (Venemans et al. 2016, and references therein). The high ratios disagree with some simulations of high-redshift quasar host galaxies (e.g., Valiante et al. 2014). Various solutions for this discrepancy have been proposed in the literature. As nearly all high-redshift quasars are selected from flux-limited, optical/near-infrared surveys, by design only the most massive BHs with the highest accretion rates are selected. Also, due to scatter in the correlation between BH and galaxy mass, the massive BHs of  $z \gtrsim 6$  quasars are preferentially found in galaxies that can be an order of magnitude less massive than expected based on the correlation itself (e.g., Willott et al. 2005). Alternatively, FIR emission lines only trace the gas in the inner regions of a galaxy, and using these lines could underestimate the true dynamical mass of the quasar host (e.g., Valiante et al. 2014).

We can also estimate the dynamical mass of the central, unresolved emission. The gas within the central beam with an FWHM of  $0''.23$  (1.2 kpc) has a velocity dispersion of  $\sigma = 235 \pm 25 \text{ km s}^{-1}$ . Setting the radius of this region to  $R = 0.5 \times \text{FWHM}_{\text{beam}}$ , we derive a mass of  $M_{\text{dyn,center}} = (1.2 \pm 0.2) \times 10^{10} M_\odot$ , which is only  $\sim 5$  times more than the mass of the BH.

We can compare the dynamical mass with the mass of the molecular gas in the quasar host galaxy. Assuming a gas-to-dust mass ratio of 100, the gas mass is  $(0.8\text{--}4) \times 10^{10} M_\odot$ , which is 20%–95% of the dynamical mass. In Section 3.4 we showed that a high fraction ( $\sim 80\%$ ) of the dust and [C II] emission resides in a compact region with a diameter  $< 1.5 \text{ kpc}$ . If this is also the case for the molecular gas, then there is not much room for a massive stellar component in the central  $\sim 1\text{--}1.5 \text{ kpc}$  of the host galaxy, which raises interesting questions regarding the origin of the detected dust emission. Due to the large uncertainties in the molecular gas mass, BH mass, and dynamical mass, we cannot put any strong constraints on the stellar mass in the quasar host galaxy.

## 5. Summary

We present ALMA, PdBI, and VLA observations targeting the dust emission and [C II], [C I], and two CO lines in the host galaxy of quasar J1120+0641 at  $z = 7.1$ . The ALMA observations of the [C II] line and the underlying continuum greatly increase the spatial resolution compared to earlier measurements with the PdBI (factor of 70 in beam area). Our main findings can be summarized as follows.

1. Within the field of view mapped by ALMA ( $\sim 25''$ ) the quasar is the only detected source.
2. The dust continuum and [C II] emission regions are very compact and only marginally resolved in the ALMA data. The majority of the emission (80%) is associated with a compact region of size  $1.2 \times 0.8 \text{ kpc}^2$  in diameter.
3. The nondetection of the [C I] line indicates that the heating in the quasar host galaxy is dominated by star formation (PDR), and not by the accreting supermassive BH (XDR).
4. From the limits on the CO(2–1) and CO(7–6) lines we derived upper limits on the molecular gas mass of  $M_{\text{gas}} < 4 \times 10^{10} M_\odot$ . This is consistent with a gas-to-dust mass ratio of  $\sim 100$  that is measured in the local universe.
5. We estimate the SFR in the quasar host using both the FIR continuum and the [C II] line measurement. Both methods give consistent results ( $\text{SFR}_{\text{FIR}} \sim 105\text{--}340 M_\odot \text{ yr}^{-1}$ ), with a resulting SFR surface density of  $\sim 100\text{--}350 M_\odot \text{ yr}^{-1} \text{ kpc}^{-2}$ , well below the value for Eddington-accretion-limited star formation (Walter et al. 2009).
6. Surprisingly, the compact [C II] emission does not exhibit ordered motion on kiloparsec scales: applying the virial theorem yields a dynamical mass of the host galaxy of  $(4.3 \pm 0.9) \times 10^{10} M_\odot$ , only  $\sim 20\times$  higher than that of the central supermassive BH. In the very central region, the dynamical mass of the host is only 5 times that of the central BH. In this region, the mass of the BH and that of the implied dust and gas are able to explain the dynamical



mass. In other words, there is not much room for a massive stellar component in the very central region. However, we note that the large uncertainties in both the molecular gas and dust mass, as well as the dynamical mass, make it unfeasible to put tight constraints on the stellar mass.

The ALMA observations presented here start to spatially resolve the host galaxy of the most distant quasar known. With the recent commissioning of even longer baselines at ALMA, even higher-resolution imaging of this quasar host is now possible that will start to spatially resolve the sphere of influence of the central supermassive BH.

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*Facilities:* ALMA, IRAM:Interferometer, EVLA.

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