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<b>Authors</b>	Wu, Y. W.; Reid, M. J.; Sakai, N.; Dame, T. M.; Menten, K. M.; et al.
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# Trigonometric Parallaxes of Star-forming Regions beyond the Tangent Point of the Sagittarius Spiral Arm

Y. W. Wu (吴元伟)<sup>1,2</sup> , M. J. Reid<sup>3</sup> , N. Sakai (坂井伸行)<sup>2</sup>, T. M. Dame<sup>3</sup> , K. M. Menten<sup>4</sup> , A. Brunthaler<sup>4</sup> , Y. Xu (徐焯)<sup>5</sup>, J. J. Li (李晶晶)<sup>5</sup>, B. Ho (胡波)<sup>5</sup> , B. Zhang (张波)<sup>6</sup> , K. L. J. Rygl<sup>7</sup>, and X. W. Zheng (郑兴武)<sup>8</sup>

<sup>1</sup> National Time Service Center, Key Laboratory of Precise Positioning and Timing Technology, Chinese Academy of Sciences, Xi'an 710600, People's Republic of China

<sup>2</sup> National Astronomical Observatory of Japan, 2-21-1 Osawa, Mitaka, Tokyo 181-8588, Japan

<sup>3</sup> Harvard-Smithsonian Center for Astrophysics, 60 Garden Street, Cambridge, MA 02138, USA

<sup>4</sup> Max-Planck-Institut für Radioastronomie, Auf dem Hügel 69, D-53121 Bonn, Germany

<sup>5</sup> Purple Mountain Observatory, Chinese Academy of Sciences, Nanjing 210008, People's Republic of China

<sup>6</sup> Shanghai Astronomical Observatory, Chinese Academy of Sciences, Shanghai 200030, People's Republic of China

<sup>7</sup> Italian ALMA Regional Centre, INAF—Istituto di Radioastronomia, Via P. Gobetti 101, I-40129 Bologna, Italy

<sup>8</sup> School of Astronomy and Space Sciences of Nanjing University, Nanjing 210093, People's Republic of China

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## Abstract

As part of the BeSSeL Survey, we report trigonometric parallaxes and proper motions of molecular maser sources associated with 13 distant high-mass star-forming regions in the Sagittarius spiral arm of the Milky Way. In particular, we obtain improved parallax distance estimates for three well-studied regions:  $1.9^{+0.1}_{-0.1}$  kpc for M17,  $5.3^{+1.3}_{-0.9}$  kpc for W51, and  $7.9^{+0.9}_{-0.7}$  kpc for GAL 045.5+00.0. Peculiar motions for all but one source are less than  $20 \text{ km s}^{-1}$ . We fit a log-periodic spiral to the locations and estimate an average pitch angle of  $7.2 \pm 1.9$ . We find that the section of the arm beyond the tangent point in the first quadrant of the Milky Way appears 15 pc below the IAU-defined Galactic plane.

*Key words:* astrometry – Galaxy: kinematics and dynamics – Galaxy: structure – masers – stars: formation

## 1. Introduction

Trigonometric parallaxes for molecular masers associated with high-mass star-forming regions (HMSFRs) from the Bar and Spiral Structure Legacy (BeSSeL) Survey and the VLBI Exploration of Radio Astrometry (VERA) project are accurately tracing the spiral structure of the Milky Way. As of 2015, over 100 parallaxes had been published (Reid et al. 2014; Honma et al. 2015). These contained 18 parallaxes for sources in the Sagittarius spiral arm, most of which were located closer to the Sun than the arm's tangent point at  $l \approx 49^\circ$  and a distance of  $\approx 5$  kpc (Wu et al. 2014). Beyond the tangent point, only four parallax measurements had been made, and those tended to have large uncertainties ( $>1$  kpc). With the aim of better constraining the structure of the “far” side of the Sagittarius arm (Sgr Far, i.e., beyond the tangent point), we measured parallaxes of 13 additional maser sources. In Section 2 we describe the observations and data analysis. Parallaxes and proper motions are presented in Section 3. We discuss how sources were assigned to the Sgr Far arm in Section 4. The geometry and kinematics of the arm are discussed in Section 5 and 6, and we summarize our conclusions in Section 7.

## 2. Observations and Data Analysis

Between 2012 and 2015, we conducted parallax observations of 11 6.7 GHz methanol and two 22 GHz water masers in the Sag Far arm segment using the National Radio Astronomy Observatory's<sup>9</sup> Very Long Baseline Array (VLBA). Generally,

three masers nearby on the sky were observed within the same 7 hr track. The dates of the observations are listed in Table 1, and coordinates of masers and background quasars are presented in Table 2.

The observations used the same equipment setup and calibration procedures described by Xu et al. (2016). The data were correlated with the DiFX<sup>10</sup> software correlator (Deller et al. 2007) at the VLBA correlation facility in Socorro, NM. Continuum emissions were recorded with four adjacent dual-circular polarization intermediate-frequency (IF) bands of 16 MHz and correlated with 32 channels per IF. The maser-line emissions were set at band center of the second or third IF and were correlated with 2000 and 4000 spectral channels for BR149 and BR198 observations, respectively, yielding velocity channel spacings of  $0.36$  and  $0.18 \text{ km s}^{-1}$  for the 6.7 GHz masers and  $0.11 \text{ km s}^{-1}$  for the 22 GHz masers.

Generally, the channel with the strongest maser emission was selected as the interferometer phase reference used to calibrate both the maser and background quasar data. For G041.15–00.20 and G041.22–00.19, whose background quasars (J1907+0907, J1919+0619) were strong enough to fringe-fit delays with precision better than 100 ps, we performed an additional calibration to remove residual (assumed dispersive) delays before phase referencing. After calibration, images of continuum emission of the quasars and cubes of maser-line emission were produced with AIPS task IMAGR. The task JMFIT was used to fit 2D Gaussian brightness distributions to compact emission peaks, yielding positions of maser spots and background quasars.

<sup>9</sup> The National Radio Astronomy Observatory is a facility of the National Science Foundation operated under cooperative agreement by Associated Universities, Inc.

<sup>10</sup> DiFX, a software correlator for VLBI, is developed as part of the Australian Major National Research Facilities Programme by the Swinburne University of Technology and operated under licence.

**Table 1**  
Observations

Source	Project	Maser	Epoch 1	Epoch 2	Epoch 3	Epoch 4	Epoch 5	Epoch 6
G032.74–00.07	BR198N	H <sub>2</sub> O	2014 Apr 11	2014 Jul 21	2014 Sep 19	2014 Oct 29	2014 Nov 30	2015 Apr 14
G035.79–00.17	BR149N	CH <sub>3</sub> OH	2012 Sep 29	2013 Mar 22	2013 Apr 30	2013 Oct 27	2013 Dec 31	2014 Oct 20
G037.47–00.10	BR198P	CH <sub>3</sub> OH	2014 Apr 13	2014 Oct 1	2014 Nov 6	2015 Apr 30	...	...
G038.03–00.30	BR198P	CH <sub>3</sub> OH	2014 Apr 13	2014 Oct 1	2014 Nov 6	2015 Apr 30	...	...
G041.15–00.20	BR198Q	CH <sub>3</sub> OH	2014 Apr 17	2014 Oct 18	2014 Nov 7	2015 Apr 25	...	...
G041.22–00.19	BR198R	CH <sub>3</sub> OH	2014 Apr 19	2014 Oct 9	2014 Nov 8	2015 May 2	...	...
G043.03–00.45	BR198R	CH <sub>3</sub> OH	2014 Apr 19	2014 Oct 9	2014 Nov 8	2015 May 2	...	...
G043.89–00.78	BR149O	CH <sub>3</sub> OH	2012 Nov 8	2013 Mar 24	2013 Apr 28	2013 Oct 29	...	...
G045.49+00.12	BR198R	CH <sub>3</sub> OH	2014 Apr 19	2014 Oct 9	2014 Nov 8	2015 May 2	...	...
G045.80–00.35	BR149P	CH <sub>3</sub> OH	2012 Nov 10	2013 Mar 26	2013 Apr 29	2013 Oct 30	...	...
G049.34+00.41	BR198T	CH <sub>3</sub> OH	2014 Apr 20	2014 Oct 30	2014 Nov 10	2015 May 4	...	...
G052.10+01.04	BR198S	H <sub>2</sub> O	2014 Apr 18	2014 Jul 25	2014 Sep 22	2014 Nov 9	2014 Dec 18	2015 May 3

**Table 2**  
First-epoch Source Information

Source	R.A.(J2000) (h m s)	Decl.(J2000) (°,′,″)	R.A. <sub>sep</sub> (deg)	Decl. <sub>sep</sub> (deg)	Beam Size (mas × mas at deg)	$F_{\text{peak}}$ (Jy beam <sup>-1</sup> )	$V_{\text{peak}}$ (km s <sup>-1</sup> )
G032.74–00.07	18:51:21.8627	–00:12:06.3359	...	...	1.28 × 0.82 at –23	9.18	+33.0
J1857–0048	18:57:51.35860	–00:48:21.9496	0.45	0.02	2.42 × 0.90 at –46	0.009	...
J1853–0010	18:53:10.26920	–00:10:50.7400	1.62	–0.60	2.54 × 0.91 at –46	0.008	...
J1848+0138	18:48:21.81035	+01:38:26.6322	–0.75	1.84	1.09 × 0.79 at –16	0.036	...
G035.79–00.17	18:57:16.89050	+02:27:58.0070	...	...	3.17 × 1.20 at –2	4.07	+61.00
1855+0251	18:55:35.43640	+02:51:19.5629	–0.42	0.39	3.03 × 1.25 at 1	0.043	...
1903+0145	19:03:53.06320	+01:45:26.3108	1.65	–0.71	3.04 × 1.24 at 0	0.072	...
1848+0138	18:48:21.81030	+01:38:26.6322	–2.23	–0.83	3.06 × 1.25 at 0	0.025	...
1907+0127	19:07:11.99610	+01:27:08.9644	2.48	–1.01	3.08 × 1.28 at 2	0.086	...
G037.47–00.10	19:00:07.14300	+03:59:52.9750	...	...	4.52 × 2.68 at 12	2.87	+62.00
J1855+0251	18:55:35.43640	+02:51:19.5630	–1.13	–1.14	5.66 × 2.39 at 25	0.065	...
J1858+0313	18:58:02.35227	+03:13:16.3116	–0.52	–0.78	4.71 × 2.85 at 18	0.137	...
J1856+0610	18:56:31.83880	+06:10:16.7650	–0.89	2.17	4.57 × 2.82 at 0	0.147	...
J1903+0145	19:03:53.06327	+01:45:26.3113	0.94	–2.24	4.29 × 2.52 at –2	0.052	...
G038.03–00.30	19:01:50.46760	+04:24:18.8999	...	...	3.76 × 2.18 at 4	3.10	+56.00
J1855+0251	18:55:35.43640	+02:51:19.5630	–1.56	–1.55	4.15 × 2.47 at 16	0.060	...
J1856+0610	18:56:31.83880	+06:10:16.7650	–1.32	1.77	3.64 × 2.08 at –2	0.138	...
J1912+0518	19:12:54.25770	+05:18:00.4220	2.76	0.89	3.73 × 2.24 at –3	0.047	...
J1858+0313	18:58:02.35227	+03:13:16.3116	–0.95	–1.18	3.74 × 1.95 at 5	0.092	...
G041.15–00.20	19:07:14.37300	+07:13:17.6500	...	...	3.11 × 1.37 at 1	3.26	+56.0
J1905+0652	19:05:21.21048	+06:52:10.7803	–0.47	–0.35	3.23 × 1.47 at 0	0.092	...
J1907+0907	19:07:41.96340	+09:07:12.3970	0.11	1.90	3.19 × 1.46 at 0	0.186	...
J1856+0610	18:56:31.83880	+06:10:16.7650	–2.66	–1.05	3.21 × 1.46 at 0	0.030	...
J1912+0518	19:12:54.25770	+05:18:00.4220	1.40	–1.92	3.32 × 1.47 at –3	0.035	...
G041.22–00.19	19:07:21.3720	+07:17:08.100	...	...	3.31 × 1.15 at –2	1.89	+55.38
J1905+0652	19:05:21.2104	+06:52:10.7803	–0.50	–0.42	2.73 × 0.98 at 0	0.031	...
J1919+0619	19:19:17.3502	+06:19:42.7700	2.96	–0.96	3.22 × 0.85 at 16	0.158	...
J1856+0610	18:56:31.8388	+06:10:16.7650	–2.68	–1.11	3.08 × 1.17 at 0	0.089	...
J1907+0907	19:07:41.9634	+09:07:12.397	0.09	1.83	3.04 × 1.16 at 0	0.077	...
G043.03–00.45	19:11:38.99000	+08:46:30.8000	...	...	3.31 × 1.15 at –2	1.55	+54.82
J1907+0907	19:07:41.96340	+09:07:12.3970	–0.98	0.34	2.85 × 1.17 at –1	0.073	...
J1913+0932	19:13:24.02539	+09:32:45.3791	0.43	0.77	2.94 × 1.12 at –3	0.008	...
J1905+0952	19:05:39.89890	+09:52:08.4070	–1.48	1.09	2.86 × 1.18 at 0	0.063	...
J1922+0841	19:22:18.63370	+08:41:57.3730	2.63	–0.08	3.09 × 1.23 at –4	0.041	...
G043.89–00.78	19:14:26.39050	+09:22:36.5660	...	...	3.70 × 1.37 at –9	3.85	+48.00
J1913+0932	19:13:24.02539	+09:32:45.3791	–0.26	0.17	3.45 × 1.36 at –9	0.009	...
J1907+0907	19:07:41.96336	+09:07:12.3968	–1.66	–0.26	3.53 × 1.33 at –10	0.096	...
J1922+0841	19:22:18.63360	+08:41:57.3780	1.94	–0.68	3.55 × 1.44 at –8	0.054	...
J1905+0952	19:05:39.89889	+09:52:08.4071	–2.16	0.49	3.35 × 1.37 at –8	0.075	...

**Table 2**  
(Continued)

Source	R.A.(J2000) (h m s)	Decl.(J2000) ( $^{\circ}$ , $'$ , $''$ )	R.A. <sub>sep</sub> (deg)	Decl. <sub>sep</sub> (deg)	Beam Size (mas $\times$ mas at deg)	$F_{\text{peak}}$ (Jy beam $^{-1}$ )	$V_{\text{peak}}$ (km s $^{-1}$ )
G045.49+00.12	19 14 11.35530	+11:13:06.3700	...	...	$3.51 \times 1.20$ at $-6$	3.10	+57.36
J1908+1201	19:08:40.32064	+12:01:58.8609	-1.35	0.81	$3.14 \times 1.34$ at $-4$	0.063	...
J1913+1307	19:13:14.00638	+13:07:47.3307	-0.23	1.91	$3.07 \times 1.21$ at $-5$	0.087	...
J1905+0952	19:05:39.89890	+09:52:08.4070	-2.09	-1.35	$3.09 \times 1.22$ at $-3$	0.076	...
J1925+1227	19:25:40.81710	+12:27:38.0870	2.82	1.24	$3.18 \times 1.28$ at $-4$	0.041	...
G045.80-00.35	19 16 31.07930	+11:16:11.9880	...	...	$3.55 \times 1.34$ at $-3$	2.23	+60.00
J1908+1201	19:08:40.32064	+12:01:58.8609	-1.92	0.76	$3.17 \times 1.30$ at $-2$	0.078	...
J1913+0932	19:13:24.02539	+09:32:45.3791	-0.76	-1.7	$3.17 \times 1.30$ at $-2$	0.155	...
J1925+1227	19:25:40.81708	+12:27:38.0856	2.25	1.19	$3.17 \times 1.30$ at $-2$	0.113	...
J1913+1307	19:13:14.00638	+13:07:47.3307	-0.81	1.86	$3.17 \times 1.30$ at $-2$	0.079	...
G049.34+00.41	19:20:32.4500	++14:45:45.400	...	...	$6.39 \times 1.28$ at $-19$	1.772	+67.60
J1922+1504	19:22:33.27280	++15:04:47.5370	0.49	0.32	$7.17 \times 1.12$ at $-20$	0.021	...
J1922+1530	19:22:34.69918	++15:30:10.0314	0.49	0.74	$6.24 \times 1.74$ at $-17$	0.080	...
J1917+1405	19:17:18.06392	++14:05:09.7704	-0.78	-0.68	$6.62 \times 1.34$ at $-18$	0.048	...
J1924+1540	19:24:39.45574	++15:40:43.9405	1.00	0.92	$6.18 \times 1.64$ at $-19$	0.257	...
G052.10+01.04	19:23:37.3295	+17:29:10.495	...	...	$0.93 \times 0.54$ at $-13$	6.330	-32.30
J1921+1625	19:21:57.38244	+16:25:01.9231	-0.40	-1.07	$1.01 \times 0.58$ at $-9$	0.041	...
J1927+1847	19:27:32.31207	+18:47:07.9048	0.93	1.30	$0.97 \times 0.57$ at $-11$	0.016	...
J1928+1859	19:28:42.71875	+18:59:24.5627	1.21	1.50	$0.94 \times 0.60$ at $-11$	0.032	...
J1924+1540	19:24:39.45590	+15:40:43.9420	0.25	-1.81	$1.01 \times 0.57$ at $-10$	0.267	...

**Note.** The  $R.A._{\text{sep}}$  is  $\Delta R.A.\cos(\text{decl.})$ .

### 3. Parallax Estimation

For 22 GHz water masers, fitting parallax and proper motion of the masers (relative to background quasars) is straightforward and has been discussed in detail in Reid et al. (2009a). However, at the lower observing frequency of the 6.7 GHz methanol masers, the effects of unmodeled ionospheric “wedges” above individual antennas can result in systematic position shifts across the sky (relative to different quasars), complicating the parallax estimation. Reid et al. (2017) discuss this problem in detail and present a strategy to minimize the effects on parallax estimates. This involves a two-step method of, first, using the relative position data for multiple quasars to generate an “artificial quasar” at the location of the target maser and, second, fitting parallax and proper-motion parameters to the artificial quasar data. While this technique improves parallax accuracy, compared to averaging parallax results for each quasar, with our typical four-epoch data it yielded poor estimates of the parallax *uncertainty*.

Since the publication of Reid et al. (2017), we developed a better approach that preserves the accuracy of the artificial quasar method, while improving the *uncertainty* estimates. In a single step, we modeled the positional data of a maser spot relative to multiple quasars at epoch  $t$  as the sum of the maser’s parallax and proper motion and a planar “tilt,” owing to ionospheric wedges, of the quasar positions about the maser position:

$$\Delta\theta_{s,q}^x(t) = \Pi^x(t) + (\Delta\theta_s^x - \Delta\theta_q^x) + \mu_s^x \delta t + S_x^x(t)\Theta_q^x + S_y^x(t)\Theta_q^y. \quad (1)$$

In Equation (1),  $\Pi^x(t)$  is the  $x$ -component of the parallax shift;  $\Delta\theta_s^x$  and  $\Delta\theta_q^x$  are constant offsets of maser spot,  $s$ , and QSO,  $q$ ,

from the position used in correlation;  $\mu_s^x \delta t$  is the  $x$ -component of maser spot position shifted by proper motion;  $S_x^x(t)\Theta_q^x$  is the  $x$ -position shift owing to an ionospheric wedge (“slope” in mas deg $^{-1}$ ) in the  $x$ -direction times the  $x$ -component of the separation between the maser and QSO  $q$  (in deg); and  $S_y^x(t)\Theta_q^y$  is the  $x$ -position shift owing to an ionospheric wedge in the  $y$ -direction times the  $y$ -component of the separation between the maser and QSO  $q$  at epoch  $t$ . The offset of one maser spot was set to zero and held constant, since one cannot solve for all  $\Delta\theta^x$  terms with relative position information. There is an analogous equation that holds for the  $y$ -coordinate.

We used a Markov chain Monte Carlo approach to generate probability distribution functions for all parameters, varying all parameters simultaneously, and then fitted the marginalized distribution functions to obtain the parallaxes and proper-motion values given in Table 3. Typical 6.7 GHz maser parallax uncertainties are  $\pm 0.02$  mas, which correspond to  $\pm 1.3$  kpc uncertainty at a distance of 8 kpc, which is characteristic of our sources past the tangent point of the Sagittarius arm. In Figure 1, we show the parallax and proper motion fitting results.

Regarding G052.10+1.04, Oh et al. (2010) reported a parallax of  $0.251 \pm 0.036$  mas measured at 22 GHz by the VERA project, whereas our 6.7 GHz data yield a parallax of  $0.162 \pm 0.013$  mas. Formally, the parallax difference has  $2.3\sigma$  tension. However, we note that Oh et al. (2010) fitted the parallax for G052.10+1.04 by combining four maser spots, implicitly assuming that the position differences of these maser spots with respect to their reference QSO are uncorrelated. This gives an optimistic parallax uncertainty, since atmospheric delay residuals between the maser and the QSO, which are

**Table 3**  
Parallaxes and Proper Motions of HMSFRs in the Sgr Arm Segment

Source name	Maser	Parallax (mas)	Distance (kpc)	$\mu_x$ (mas yr <sup>-1</sup> )	$\mu_y$ (mas yr <sup>-1</sup> )	$V_{\text{LSR}}$ (km s <sup>-1</sup> )	Reference
G351.44+00.65	12 GHz	0.752 ± 0.069	1.33 <sup>+0.13</sup> <sub>-0.11</sub>	0.31 ± 0.58	-2.17 ± 0.90	-8 ± 3	2, 11
G011.49-01.48	12 GHz	0.800 ± 0.033	1.25 <sup>+0.05</sup> <sub>-0.05</sub>	1.42 ± 0.52	-0.60 ± 0.65	11 ± 3	2
G014.33-00.64	22 GHz	0.893 ± 0.101	1.12 <sup>+0.14</sup> <sub>-0.11</sub>	0.95 ± 1.50	-2.40 ± 1.30	22 ± 5	4
G014.63-00.57	22 GHz	0.546 ± 0.022	1.83 <sup>+0.08</sup> <sub>-0.07</sub>	0.22 ± 1.20	-2.07 ± 1.20	19 ± 5	2
G015.03-00.67	12 GHz	0.499 ± 0.026	2.00 <sup>+0.11</sup> <sub>-0.10</sub>	0.68 ± 0.53	-1.42 ± 0.54	22 ± 3	5, 12
G034.39+00.22	22 GHz	0.643 ± 0.049	1.56 <sup>+0.13</sup> <sub>-0.11</sub>	-0.90 ± 1.00	-2.75 ± 2.00	57 ± 5	6
G035.02+00.34	22 GHz	0.430 ± 0.040	2.33 <sup>+0.24</sup> <sub>-0.20</sub>	-0.92 ± 0.90	-3.61 ± 0.90	52 ± 5	2
G035.19-00.74	12 GHz	0.456 ± 0.045	2.19 <sup>+0.24</sup> <sub>-0.20</sub>	-0.18 ± 0.50	-3.63 ± 0.50	30 ± 7	7
G035.20-01.73	12 GHz	0.412 ± 0.014	2.43 <sup>+0.09</sup> <sub>-0.08</sub>	-0.68 ± 0.44	-3.60 ± 0.44	43 ± 5	7, 13
G037.42+01.51	12 and 22 GHz	0.532 ± 0.021	1.88 <sup>+0.08</sup> <sub>-0.07</sub>	-0.45 ± 0.35	-3.69 ± 0.39	41 ± 3	2
G049.48-00.36	12 GHz	0.195 ± 0.071	5.13 <sup>+2.94</sup> <sub>-1.37</sub>	-2.49 ± 0.14	-5.51 ± 0.16	56 ± 3	8
G049.48-00.38	22 GHz	0.185 ± 0.010	5.41 <sup>+0.31</sup> <sub>-0.28</sub>	-2.64 ± 0.20	-5.11 ± 0.20	58 ± 4	9
<hr/>							
G032.74-00.07	22 GHz	0.126 ± 0.016	7.94 <sup>+1.15</sup> <sub>-0.89</sub>	-3.15 ± 0.27	-6.10 ± 0.29	37 ± 10	1
G035.79-00.17	6.7 GHz	0.113 ± 0.013	8.85 <sup>+1.15</sup> <sub>-0.91</sub>	-2.96 ± 0.12	-6.23 ± 0.14	61 ± 5	1
G037.47-00.10	6.7 GHz	0.088 ± 0.030	11.36 <sup>+5.88</sup> <sub>-2.89</sub>	-2.63 ± 0.07	-6.19 ± 0.15	58 ± 3	1
G038.03-00.30	6.7 GHz	0.095 ± 0.022	10.53 <sup>+3.17</sup> <sub>-1.98</sub>	-3.01 ± 0.06	-6.20 ± 0.11	60 ± 3	1
G041.15-00.20	6.7 GHz	0.125 ± 0.018	8.00 <sup>+1.35</sup> <sub>-1.01</sub>	-2.79 ± 0.14	-5.85 ± 0.16	60 ± 3	1
G041.22-00.19	6.7 GHz	0.113 ± 0.022	8.85 <sup>+2.14</sup> <sub>-1.44</sub>	-2.82 ± 0.13	-5.89 ± 0.16	59 ± 5	1
G043.03-00.45	6.7 GHz	0.130 ± 0.019	7.69 <sup>+1.32</sup> <sub>-0.98</sub>	-3.03 ± 0.15	-6.56 ± 0.20	56 ± 3	1
G043.79-00.12	22 GHz	0.166 ± 0.010	6.02 <sup>+0.39</sup> <sub>-0.34</sub>	-3.02 ± 0.36	-6.20 ± 0.36	44 ± 10	2
G043.89-00.78	6.7 and 22 GHz	0.134 ± 0.013	7.41 <sup>+0.79</sup> <sub>-0.65</sub>	-3.11 ± 0.15	-5.35 ± 0.21	50 ± 3	1, 2
G045.07+00.13	22 GHz	0.129 ± 0.007	7.75 <sup>+0.45</sup> <sub>-0.40</sub>	-3.21 ± 0.26	-6.11 ± 0.26	59 ± 5	2
G045.45+00.06	22 GHz	0.119 ± 0.017	8.40 <sup>+1.40</sup> <sub>-1.05</sub>	-2.34 ± 0.38	-6.00 ± 0.54	55 ± 7	2
G045.49+00.12	6.7 GHz	0.144 ± 0.024	6.94 <sup>+1.39</sup> <sub>-0.99</sub>	-2.62 ± 0.17	-5.61 ± 0.16	57 ± 3	1
G045.80-00.35	6.7 GHz	0.137 ± 0.023	7.30 <sup>+1.47</sup> <sub>-1.05</sub>	-2.52 ± 0.17	-6.08 ± 0.27	64 ± 5	1
G048.99-00.29	22 GHz	0.178 ± 0.017	5.62 <sup>+0.59</sup> <sub>-0.49</sub>	-2.20 ± 0.48	-5.84 ± 0.63	67 ± 10	3
G049.19-00.33	22 GHz	0.197 ± 0.008	5.08 <sup>+0.22</sup> <sub>-0.20</sub>	-3.08 ± 0.40	-5.50 ± 0.40	67 ± 5	2, 3
G049.34+00.41	6.7 GHz	0.241 ± 0.031	4.15 <sup>+0.61</sup> <sub>-0.47</sub>	-2.36 ± 0.26	-5.59 ± 0.33	68 ± 5	1
G049.59-00.24	6.7 GHz	0.218 ± 0.009	4.59 <sup>+0.20</sup> <sub>-0.18</sub>	-2.25 ± 0.24	-6.12 ± 0.25	63 ± 5	1
G052.10+01.04	22 GHz	0.165 ± 0.013	6.06 <sup>+0.52</sup> <sub>-0.44</sub>	-2.77 ± 1.40	-5.85 ± 1.40	42 ± 40	1, 10

**Note.** Column (1): source name. Column (2): maser transition; 6.7 and 12.2 GHz are CH<sub>3</sub>OH and 22 GHz are H<sub>2</sub>O masers. Columns (3) and (4): trigonometric parallaxes and distances. Columns (5), (6), and (7): proper-motion components in the eastward ( $\mu_x = \mu_\alpha \cos\delta$ ) and northward directions ( $\mu_y = \mu_\delta$ ) and the estimated LSR velocity of the central exciting star. Column (8): references. The dividing line within the table separates Sgr Far sources (below line) from Sgr Near sources (above line).

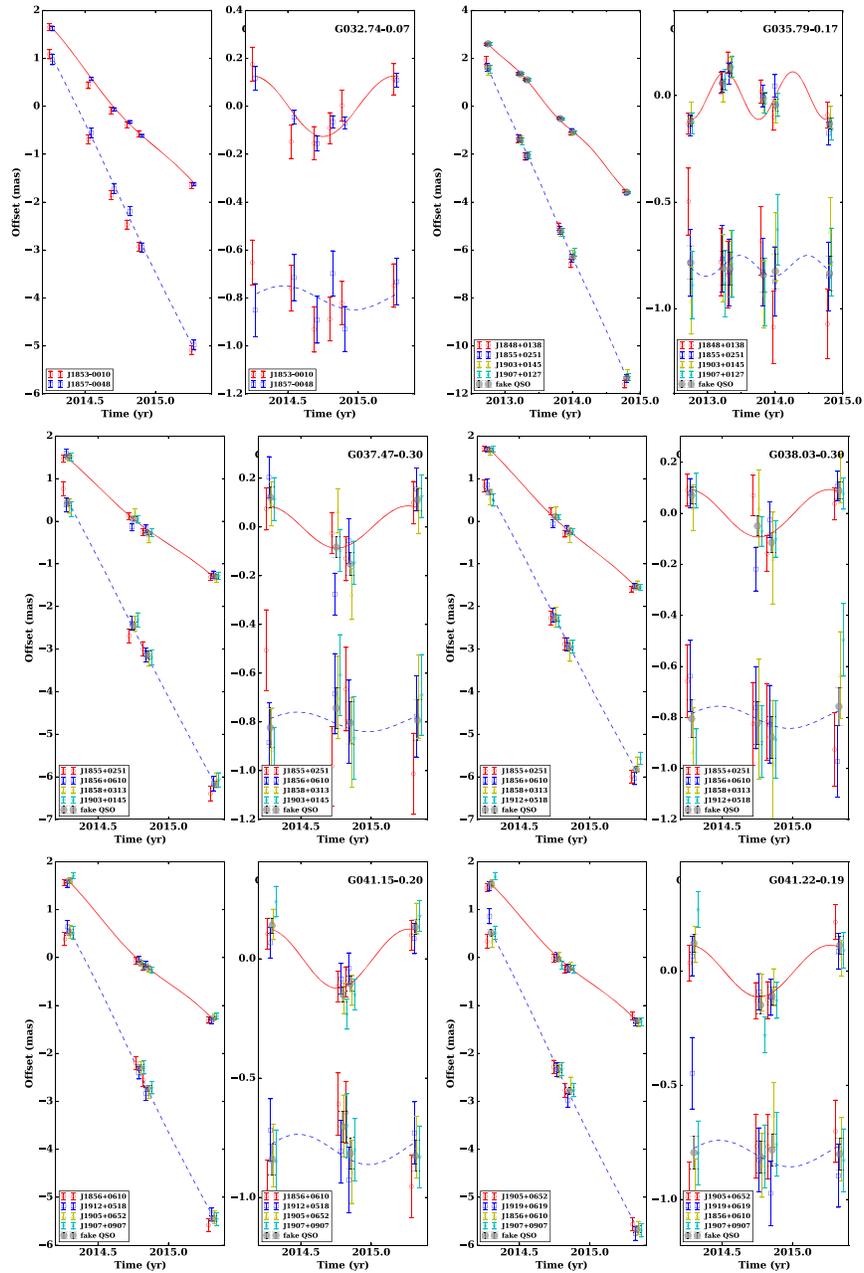
**References.** (1) This paper; (2) Wu et al. 2014; (3) Nagayama et al. 2015; (4) Sato et al. 2010a; (5) Xu et al. 2011; (6) Kurayama et al. 2011; (7) Zhang et al. 2009; (8) Xu et al. 2009; (9) Sato et al. 2010b; (10) Oh et al. 2010; (11) Chibueze et al. 2014; (12) Chibueze et al. 2016; (13) K. L. Rygl et al. 2018, in preparation.

common to all maser spots, usually dominate systematic uncertainties. BeSSeL Survey results generally allow for correlated parallaxes among maser spots, as we conservatively inflate the formal uncertainty by  $\sqrt{N}$ , where  $N$  is the number of spots used for parallax fitting. Doing this for the VERA parallax would give an uncertainty of  $\pm 0.072$  mas, making these two parallaxes statistically consistent at the  $1.2\sigma$  level. Then, combining VERA and BeSSeL measurements yields a parallax of  $0.165 \pm 0.013$  mas.

For G043.89-00.78, our 6.7 GHz data yield a parallax of  $0.144 \pm 0.017$  mas, and combining this with the 22 GHz parallax of  $0.121 \pm 0.020$  mas of Wu et al. (2014) yields a variance-weighted average parallax of  $0.134 \pm 0.013$  mas. Rather than averaging the proper motions from the two maser species, we adopt the motion of the 6.7 GHz methanol masers as the best estimate of the exciting star's motion. We do this

because 22 GHz water masers trace outflows with speeds of typically tens of kilometers per second, whereas methanol masers are generally within  $5 \text{ km s}^{-1}$  of the exciting star.

We note that in our sample masers with similar coordinates, distances, and radial velocities are likely in the same giant molecular cloud. For example, G014.63-00.57 and G015.03-00.67 are associated with the large star-forming region M17 (see Figure 2); G049.48-00.36, G049.48-00.38, G048.99-00.29, and G049.19-00.33 appear to be associated with W51 (see Figure 3); and G045.45+00.06, G045.07+00.13, and G045.49+00.12 may be associated with the H II region GAL 045.5+00.0 (see Figure 4). By combining parallax measurements of the associated sources, we can improve the distance estimates to the larger star-forming regions, yielding averaged parallaxes of  $0.524 \pm 0.024$  mas for M17,  $0.189 \pm 0.037$  mas for W51, and  $0.126 \pm 0.013$  mas for GAL 045.5+00.0.



**Figure 1.** Parallax and proper-motion data and fits. There are two panels for each source: the left panel plots position vs. time, and the right panel has the fitted proper motion removed to better show the parallax signatures. For methanol masers, the gray circles indicate data for “artificial quasar” at the position of the maser.

#### 4. Assigning Sources to Arm Segments

In Figure 5, we show the locations of HMSFRs with measured parallaxes superposed on a CO longitude–LSR velocity ( $l$ – $v$ ) plot. The assignment of a maser source to a specific spiral arm or arm segment can usually be done via such a plot, by comparing the Galactic coordinates and LSR velocity of the source with continuous traces of CO (or H I) emission previously associated with an arm. However, ( $l$ – $v$ ) information alone does not always allow a definitive assignment. Even for sources with direct parallax distances, the distance uncertainty does not always guarantee a spiral arm segment assignment. Instead, distance can best be assigned on a probabilistic basis, and when doing this, it is important that all possible information is utilized.

For the BeSSeL Survey, we evaluate arm segment assignment using our Bayesian distance estimator (Reid et al. 2016), which considers all distance-related information, including a new feature that incorporates a proper-motion “kinematic distance” probability density function (pdf; M. J. Reid et al. 2018, in preparation). For example, consider G032.74–00.07, a water maser at  $v = 37 \text{ km s}^{-1}$ , which in ( $l$ ,  $v$ ) space lies between the Sgr Near and Sgr Far arm segments<sup>11</sup> (as seen in Figure 5). The distance pdf for this source is shown in Figure 6. As expected, the probabilities of matching a spiral arm segment,  $\text{Prob}_{\text{SA}}$ , are

<sup>11</sup> We define “near” and “far” for the Sgr arm as relative to the tangent point. While this is equivalent to kinematic near/far convention, it does not imply that we have used absorption information to make this distinction.

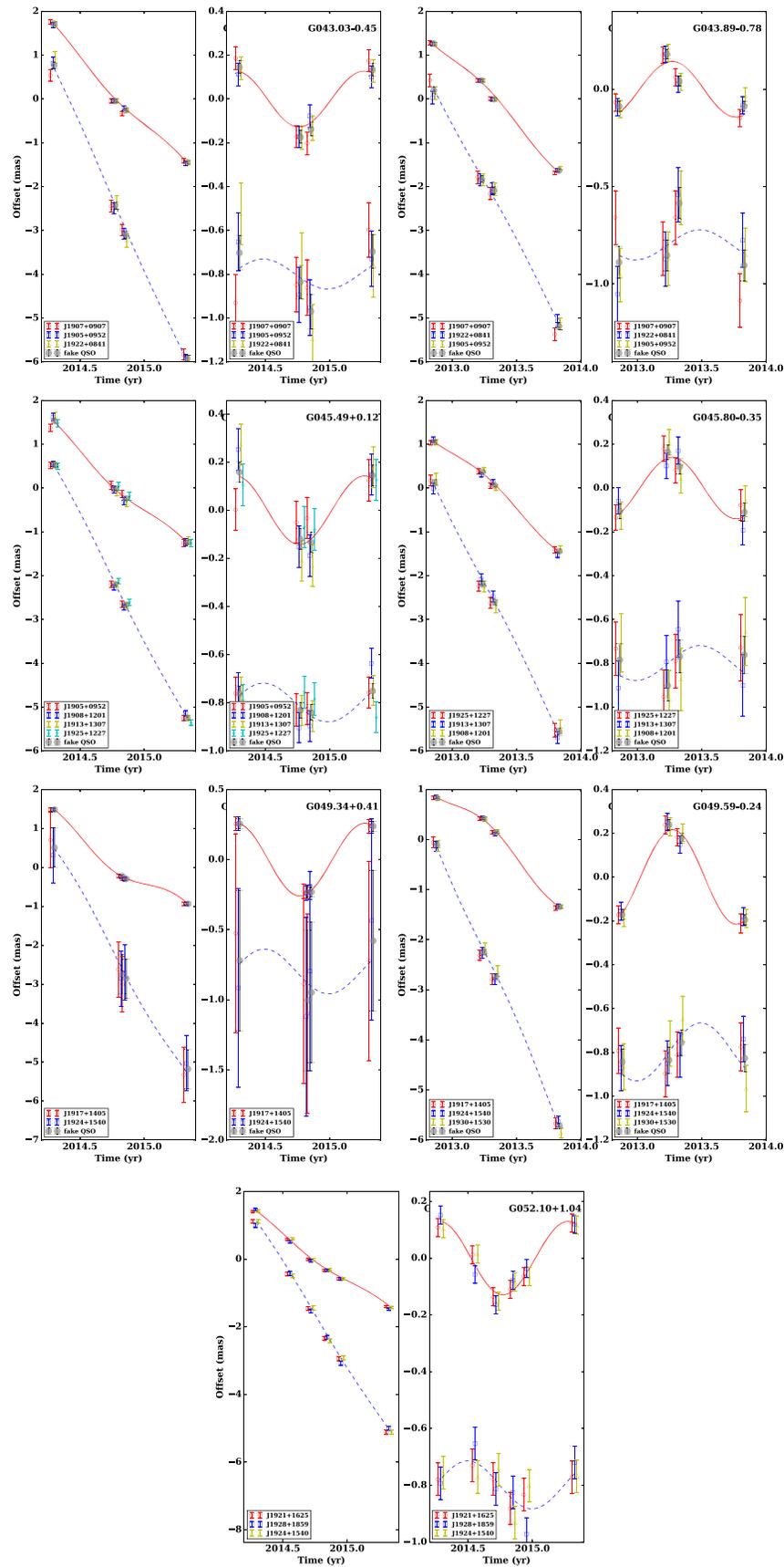
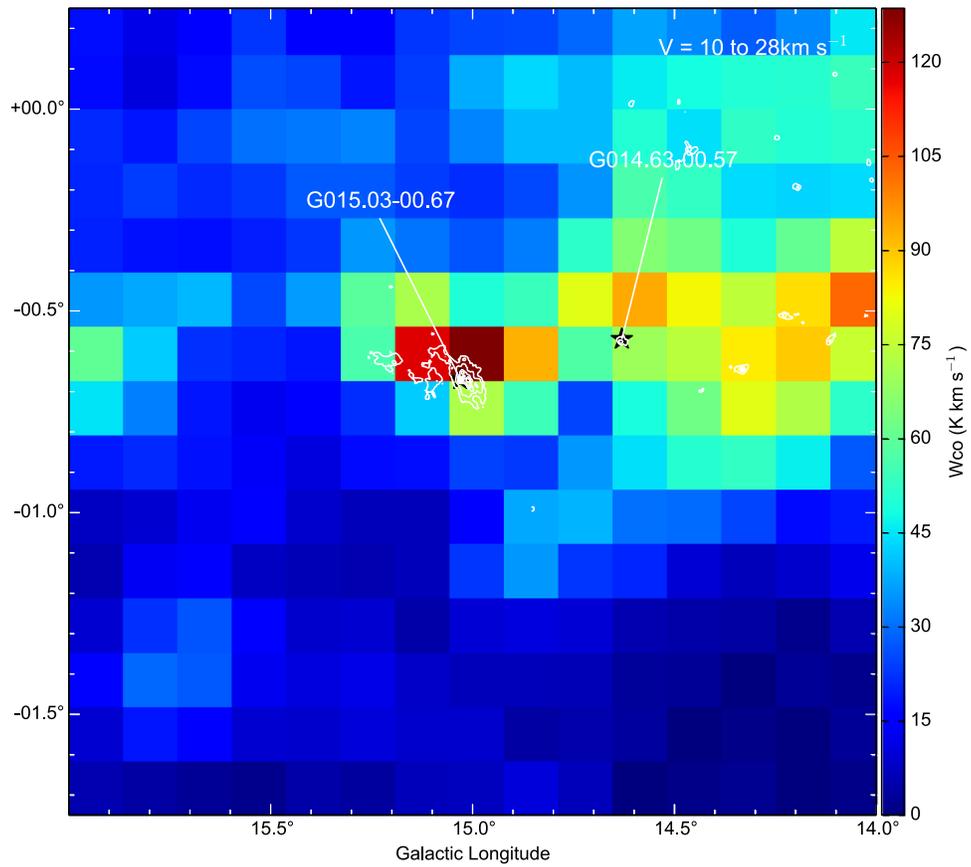
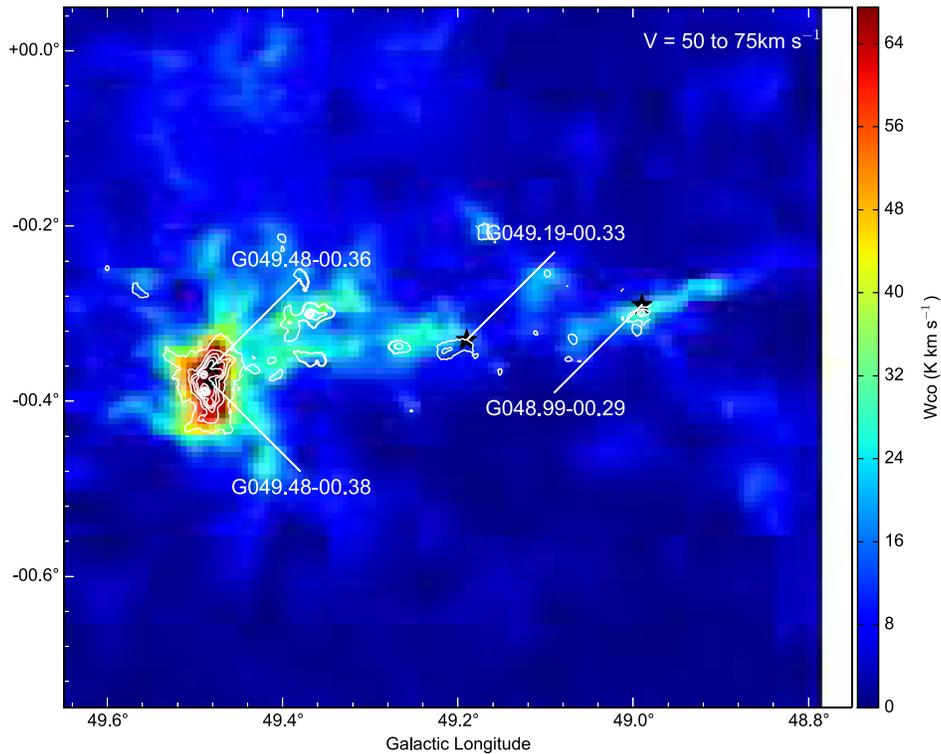


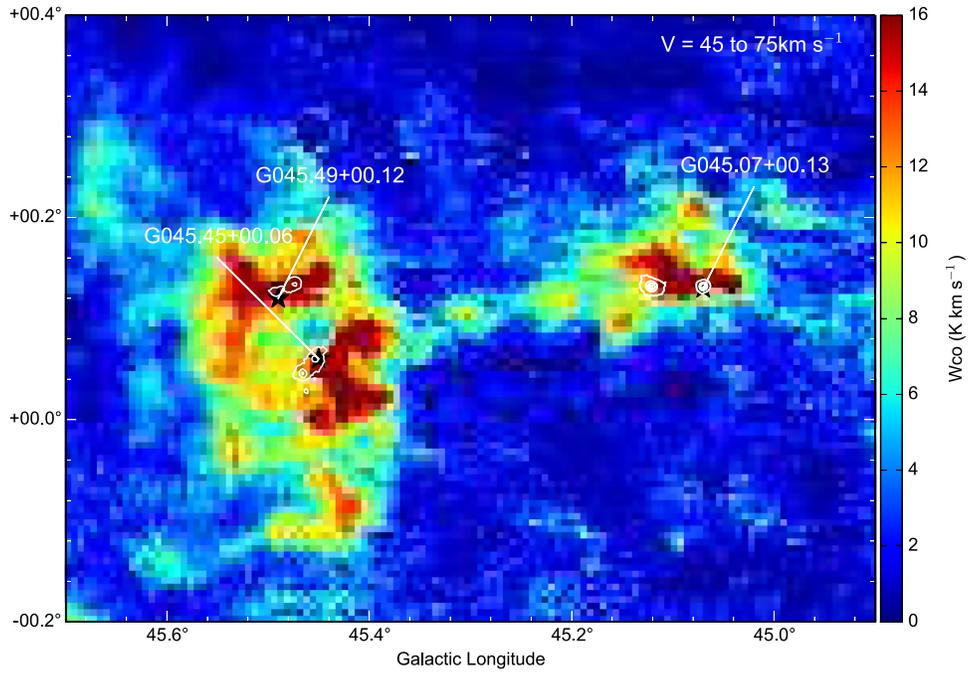
Figure 1. (Continued.)



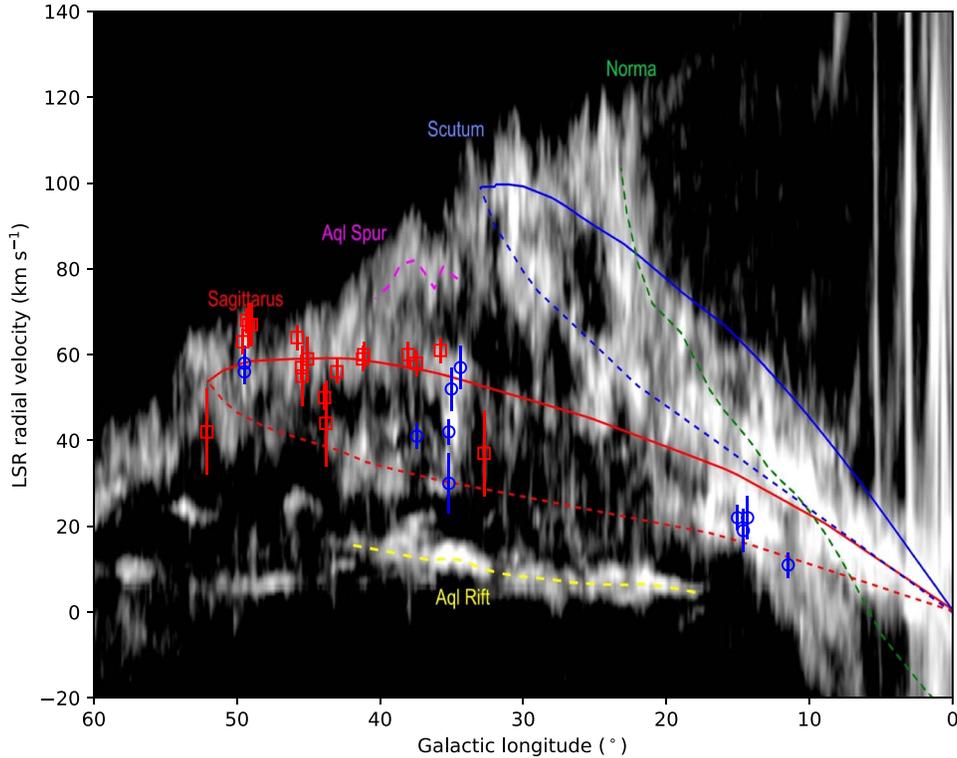
**Figure 2.** False-color image of the CO ( $J = 1-0$ ) emission integrated over 10–28  $\text{km s}^{-1}$  from the CFA 1.2 m Survey (Dame et al. 2001). Overplotted are contours of 870  $\mu\text{m}$  dust emission from the ATLASGAL survey (Schuller et al. 2009), with contour levels of 1, 3, 7, 11, 15, and 19  $\text{Jy beam}^{-1}$ . The positions of the masers are indicated with black stars.



**Figure 3.** False-color image of  $^{13}\text{CO}$  ( $J = 1-0$ ) emission from the GRS  $^{13}\text{CO}$  Survey Archive (Jackson et al. 2006) integrated over  $50 \text{ km s}^{-1} < V_{\text{LSR}} < 75 \text{ km s}^{-1}$ . Overplotted are contours of 870  $\mu\text{m}$  dust emission from the ATLASGAL survey (Schuller et al. 2009), with contour levels of 0.5, 1, 3, 5, 7, 10, 20, ..., 70  $\text{Jy beam}^{-1}$ . The position of masers are indicated with black stars.



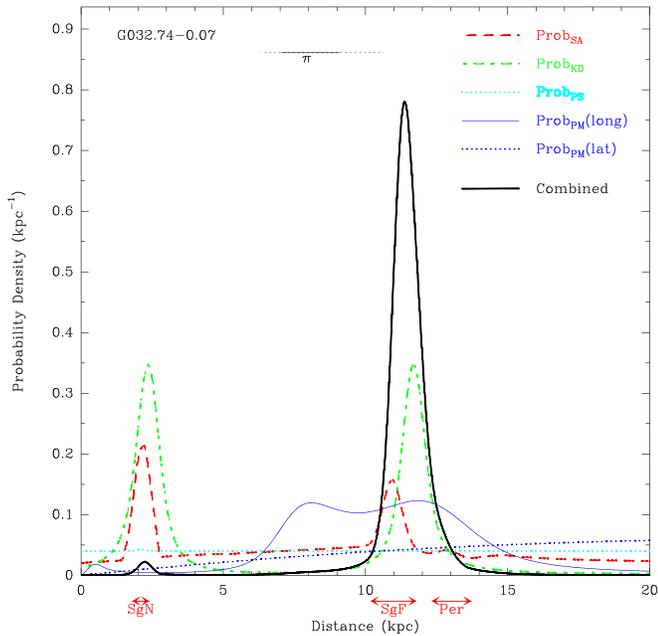
**Figure 4.** False-color image of  $^{13}\text{CO}$  ( $J = 1-0$ ) emission from the GRS  $^{13}\text{CO}$  Survey Archive (Jackson et al. 2006) integrated over  $45 \text{ km s}^{-1} < V_{\text{LSR}} < 75 \text{ km s}^{-1}$ . Overplotted are contours of  $870 \mu\text{m}$  dust emissions from the ATLASGAL survey (Schuller et al. 2009), with contour levels of 0.2, 0.5, 1, 3, 5, and  $7 \text{ Jy beam}^{-1}$ . The position of masers are indicated with black stars.



**Figure 5.** Overlay of the Sgr Far (red squares) sources and Sgr Near (blue circles) on a longitude–velocity plot of CO emission from the CfA 1.2 m survey (Dame et al. 2001). This figure is adapted from Figure 7 of Reid et al. (2016), where the Sagittarius, Scutum, and Norma arms in the first Galactic quadrant, as well as the Aquila spur and Aquila rift substructures, are traced by colored lines. The near and far segments of the arms, with respect to the arm tangent point, are indicated by dashed and solid lines, respectively. Our assignment of sources to the near or far segments used this and other information (see Section 4 for discussion).

similar for the Sgr Near and the Sgr Far segments. However, the Galactic longitude motion,  $\text{Prob}_{\text{PM}}(\text{long})$ , favors the Sgr Far segment, as does the measured parallax, denoted by the

horizontal line at the top of Figure 6 (with solid and dotted lines indicating its  $1\sigma$  and  $2\sigma$  uncertainty ranges, respectively), albeit with about a  $2\sigma$  deviation. Putting all information together, we



**Figure 6.** Distance pdf for G032.74–0.07, a 22 GHz water maser at  $V_{\text{LSR}} = 47 \text{ km s}^{-1}$ . The  $(l, b, v)$  values slightly favor the Sgr Near (SgN) arm segment. However, the combination (thick black line) of arm assignment probability (red dashed line) with the radial velocity kinematic distance (green dotted-dashed line) and the proper motion in the Galactic longitude and latitude directions (blue solid and dotted lines) favors location in the Sgr Far (SgF) arm segment over the SgN segment and Perseus arm (Per). Our measured parallax for this source is shown at the top of the figure by the horizontal solid and dotted lines, indicating the  $1\sigma$  and  $2\sigma$  uncertainty ranges. The parallax also favors the Sgr Far arm segment, albeit with  $2\sigma$  tension.

assigned G032.74–00.07 to the Sgr Far arm segment. Based on similar evidence, we also assigned G043.79–0.12 to Sgr Far.

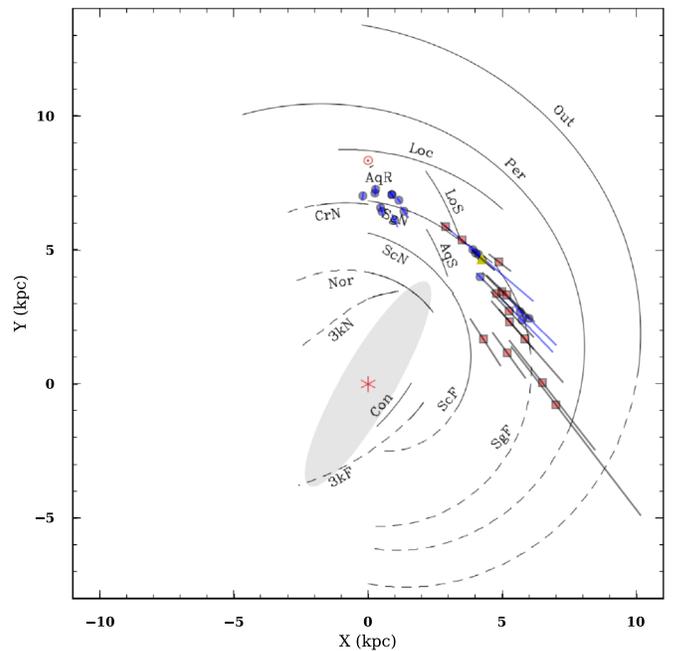
Other apparent anomalies in the arm segment assignments are seen for the following sources near  $l = 35^\circ$  in Figure 5: G034.39+00.22, G035.02+00.34, and G035.20–01.73. This is a very interesting region, with sources projected close to the W44 supernova remnant. It is possible that the precursor O star’s wind may have accelerated gas that has now formed the stars we observe (Reid & Ho 1985). If so, this might explain anomalous velocities that can confuse  $(l, v)$  locations. For sources in this region, we rely mostly on parallaxes to assign to an arm segment.

## 5. Geometry of the Sagittarius Arm

### 5.1. Galactic Locations of HMSFRs

In Figure 7, we show locations of all HMSFRs associated with the Sgr arm with measured parallaxes superposed on a plan view of the Galaxy. There now are 18 HMSFRs located beyond the tangent point of the Sgr arm: 13 sources reported in this paper, 4 reported by Wu et al. (2014), and 1 reported by Nagayama et al. (2015). The locations of all sources are consistent with the Reid et al. (2016) model for the Sagittarius arm.

We note that 16 of 18 sources in the far portion of the Sagittarius arm are below the IAU-defined Galactic plane. In order to investigate this asymmetry, we looked at the Galactic distribution of all the 6.7 GHz masers thought to be in this section of the arm based on their location in the  $l$ - $v$  plane. We assigned sources to the far arm section if their longitudes were



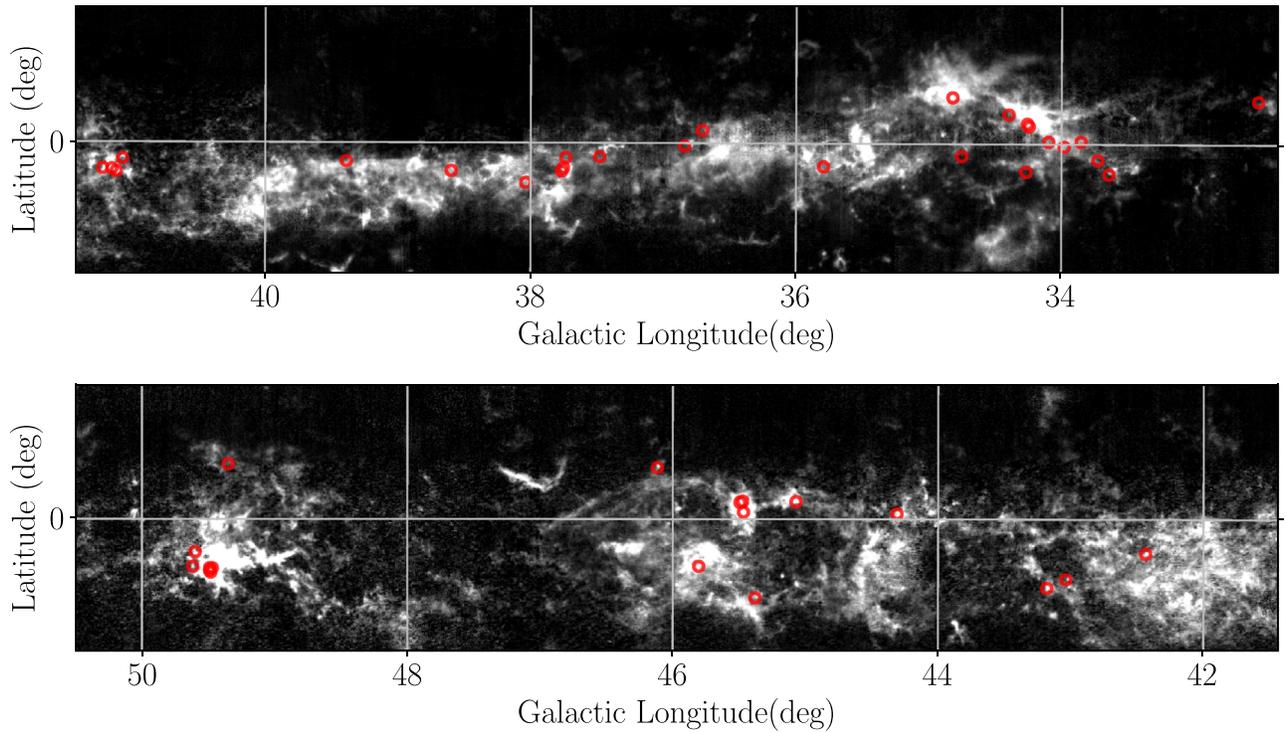
**Figure 7.** Locations of HMSFRs in the Sagittarius arm on a schematic plan view of the Milky Way from Reid et al. (2016). Red squares are 13 sources reported in this paper; blue circles are 18 HMSFRs from Wu et al. (2014). The yellow triangle is G048.99–00.29 from Nagayama et al. (2015). Sources are located by the parallax measurement, assuming that distance is the inverse of the parallax, and the lines indicate distance uncertainty by adding and subtracting the parallax uncertainty from the parallax.

within the range  $32^\circ < l < 50^\circ$  and their velocities were within the range  $50 \text{ km s}^{-1} < V_{\text{LSR}} < 70 \text{ km s}^{-1}$ . Searching the Methanol Multi-Beam Survey maser catalog of Breen et al. (2015), we identified 42 6.7 GHz masers in the far portion of the arm; they are shown as red circles in Figure 8.

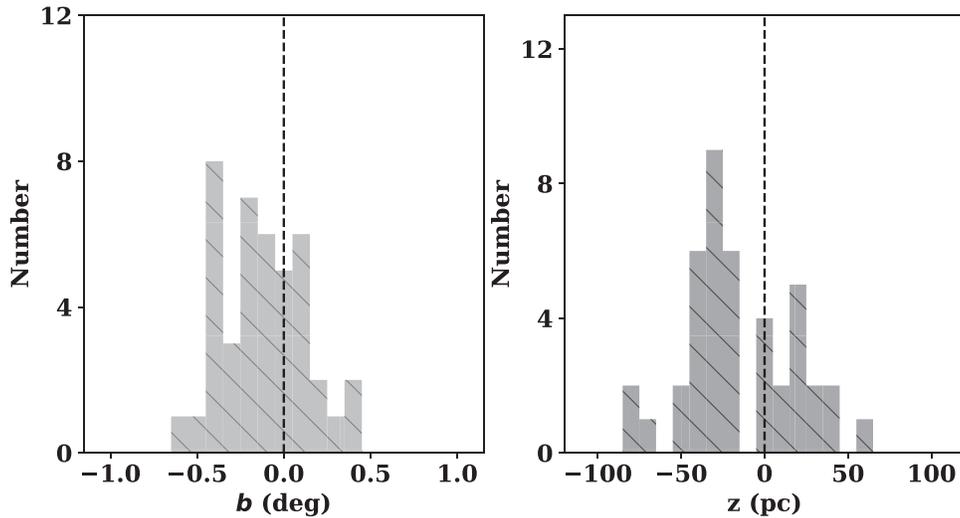
In the two panels of Figure 9 we show the histograms of Galactic latitude,  $b$ , and the vertical distances,  $z$ , of these masers from the IAU Galactic plane. For sources without maser parallaxes, we used distances calculated with the Bayesian distance estimator of Reid et al. (2016). Figure 9 suggests a systematic  $z$ -offset for the Sgr Far arm section. The number of sources below and above the Galactic midplane are 28 and 14, respectively; the mean vertical distance with respect to the midplane,  $\langle z \rangle$ , of these 6.7 GHz methanol masers is  $-15 \text{ pc}$ .

### 5.2. Pitch Angle of the Sagittarius Arm

A spiral arm pitch angle,  $\psi$ , can be defined as the angle between an arm segment and a tangent to a galactocentric circle. For a log-periodic spiral,  $\ln(R) = \ln(R_{\text{ref}}) - (\beta - \beta_{\text{ref}})\tan\psi$ , where  $R$  is the galactocentric radius in kpc at azimuth  $\beta$  in radians,  $\beta_{\text{ref}}$  is a reference azimuth, and  $R_{\text{ref}}$  is the radius of the arm at that azimuth. With parallaxes for 30 HMSFRs, we estimate a global pitch angle of  $7.2^\circ \pm 1.9^\circ$ , which is consistent with our previous estimate of  $7.1^\circ \pm 1.5^\circ$  (Wu et al. 2014). The increase in uncertainty between the previous and current estimates suggests that the pitch angle may not be constant over the larger range of azimuth that we now have sampled. Such effects are seen in other spiral galaxies (Honig & Reid 2015). If we divide the arm into two sections, we estimate pitch angles of  $5.4^\circ \pm 3.3^\circ$  and  $9.4^\circ \pm 3.9^\circ$  for the far and near portions of the Sgr arm,



**Figure 8.** Intensity map of the  $^{13}\text{CO}$  ( $J = 1-0$ ) emission integrated over  $50 \text{ km s}^{-1} < V_{\text{LSR}} < 70 \text{ km s}^{-1}$  from the Galactic Ring Survey (Jackson et al. 2006). Superposed with red circles are methanol masers from the Parkes Methanol Multibeam Survey (Breen et al. 2015) with  $50 \text{ km s}^{-1} < V_{\text{LSR}} < 70 \text{ km s}^{-1}$ .



**Figure 9.** Histograms of the Galactic latitudes and  $z$ -heights of 6.7 GHz methanol masers from the methanol multibeam survey (Breen et al. 2015) with  $32^\circ < l < 50^\circ$  and  $50 \text{ km s}^{-1} < V_{\text{LSR}} < 70 \text{ km s}^{-1}$ .

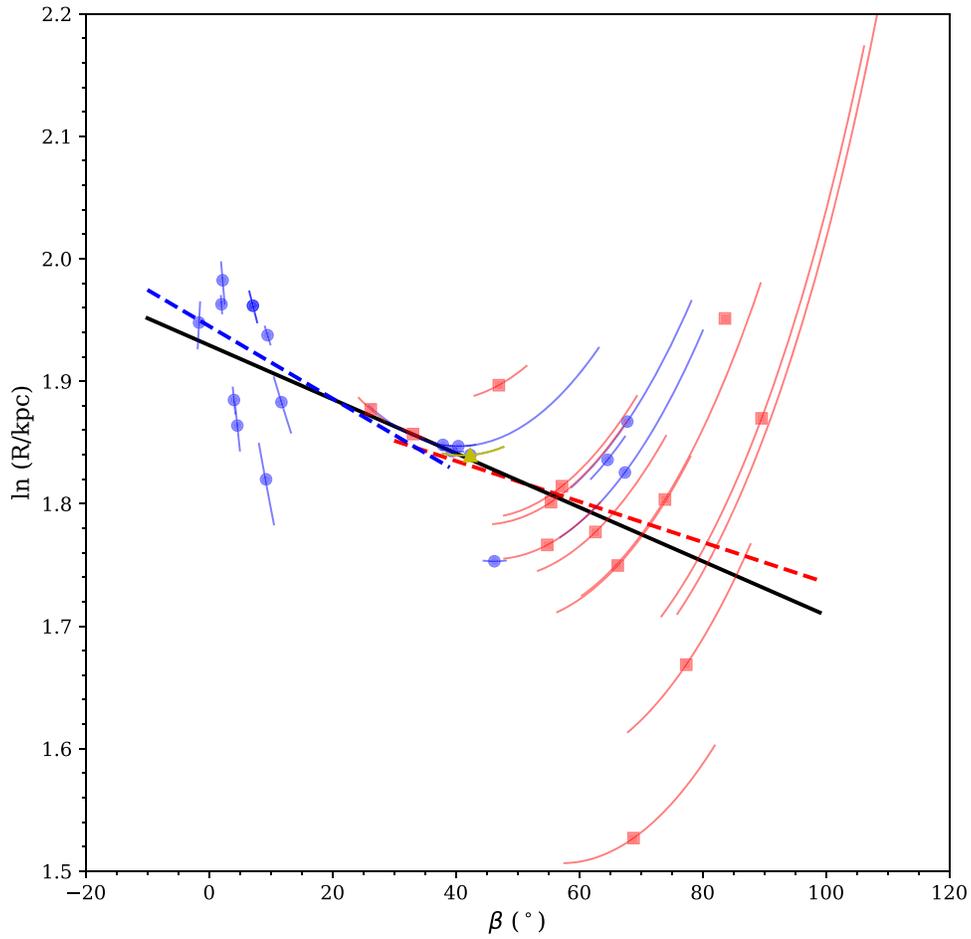
respectively, as shown in Figure 10. For comparison, pitch angles of spiral arms in the Milky Way, determined by maser parallaxes, range from  $7^\circ$  to  $20^\circ$ , with the Sagittarius arm having the smallest pitch angle (Reid et al. 2014).

## 6. Kinematics of the Sagittarius Arm

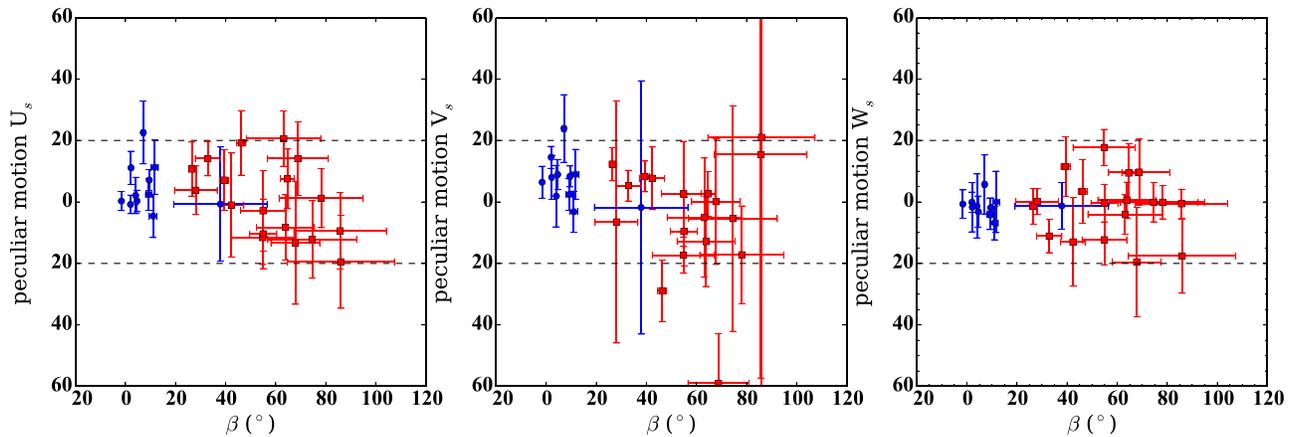
We now investigate the 3D motions of HMSFRs in the Sgr arm. We follow the methodology given in the Appendix of Reid et al. (2009b) to calculate peculiar motion:  $(U_s, V_s, W_s)$ ,

where  $U_s$  is toward the Galactic center,  $V_s$  is in the direction of Galactic rotation, and  $W_s$  is toward the north Galactic pole. We adopt the “universal” form for the rotation curve of Persic et al. (1996), with a circular rotation speed at the Sun of  $\Theta_0 = 241 \text{ km s}^{-1}$ ,  $R_0 = 8.31 \text{ kpc}$ , and solar motion components of  $(U_\odot, V_\odot, W_\odot) = (10.5, 14.4, 8.9) \text{ km s}^{-1}$  (see the “Univ” model in Table 5 of Reid et al. 2014).

In Table 4, we list peculiar motion components for the 30 (18 Sgr Far and 12 Sgr Near) HMSFRs in the Sgr arm with measured parallaxes and proper motions. Plots of the peculiar



**Figure 10.** Logarithm of Galactic radius (in kpc units),  $R$ , vs. galactocentric azimuth,  $\beta$ . Symbols are the same as in Figure 7. The black solid line denotes a pitch angle of  $7.2^\circ$ , fitted with all points; the red and blue dashed lines denote pitch angles of  $5.4^\circ$  and  $9.4^\circ$  fitted to the Sgr Far and Near arm segments, respectively.



**Figure 11.** Peculiar (noncircular) motion components in  $\text{km s}^{-1}$  of the Sgr arm HMSFRs as a function of the galactocentric azimuth,  $\beta$ . Blue circles and red squares denote sources on the near and far side of the tangent point, respectively. Except for one source, G032.74–00.07, the peculiar motions of all HMSFRs are within about  $\pm 20 \text{ km s}^{-1}$  (denoted by dashed lines).

motion components versus galactocentric azimuth  $\beta$  are shown in Figure 11. All values are consistent with  $0 \pm 20 \text{ km s}^{-1}$ , except for the  $V$  component of G032.74–00.07, which is  $-59 \pm 16 \text{ km s}^{-1}$ . The variance-weighted average peculiar

motion components for the far, near, and all Sgr arm sources are  $(4.3 \pm 2.4, -0.6 \pm 2.2, -0.2 \pm 1.6) \text{ km s}^{-1}$ ,  $(2.0 \pm 1.4, 7.1 \pm 1.5, -1.1 \pm 1.5) \text{ km s}^{-1}$ , and  $(2.6 \pm 1.1, 4.5 \pm 1.3, -0.7 \pm 1.1) \text{ km s}^{-1}$ , respectively.

**Table 4**  
Peculiar Motions

Source Name	Arm Section	$U_s$ ( $\text{km s}^{-1}$ )	$V_s$ ( $\text{km s}^{-1}$ )	$W_s$ ( $\text{km s}^{-1}$ )
G351.44+00.65	Near	$+0.3 \pm 3.1$	$+6.4 \pm 5.2$	$-0.7 \pm 4.6$
G011.49-01.48	Near	$-0.8 \pm 3.0$	$+14.5 \pm 3.6$	$-0.1 \pm 3.3$
G014.33-00.64	Near	$+11.1 \pm 5.4$	$+8.0 \pm 7.2$	$-1.7 \pm 8.0$
G014.63-00.57	Near	$+2.1 \pm 5.9$	$+1.9 \pm 10.0$	$-1.3 \pm 10.4$
G015.03-00.67	Near	$+0.3 \pm 3.5$	$+8.9 \pm 4.9$	$-3.2 \pm 5.1$
G034.39+00.22	Near	$+22.6 \pm 10.2$	$+23.9 \pm 11.0$	$+5.7 \pm 9.6$
G035.02+00.34	Near	$+11.3 \pm 8.8$	$+9.0 \pm 8.1$	$0.0 \pm 10.0$
G035.19-00.74	Near	$-4.6 \pm 7.0$	$-3.2 \pm 6.7$	$-6.8 \pm 5.5$
G035.20-01.73	Near	$+2.4 \pm 5.2$	$+2.4 \pm 5.1$	$-3.9 \pm 5.1$
G037.42+01.51	Near	$+7.1 \pm 3.4$	$+8.3 \pm 3.4$	$-1.9 \pm 3.2$
G049.48-00.36	Near	$-0.7 \pm 18.6$	$-1.8 \pm 41.2$	$-1.3 \pm 7.6$
G049.48-00.38	Near	$-10.5 \pm 6.1$	$0.0 \pm 4.0$	$+6.3 \pm 5.1$
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G032.74-00.07	Far	$+14.2 \pm 12.0$	$-59.0 \pm 16.2$	$+9.7 \pm 10.7$
G035.79-00.17	Far	$+1.2 \pm 9.6$	$-17.2 \pm 15.9$	$-0.1 \pm 5.4$
G037.47-00.10	Far	$-19.5 \pm 15.1$	$+21.0 \pm 113.9$	$-17.6 \pm 12.1$
G038.03-00.30	Far	$-9.4 \pm 12.4$	$+15.4 \pm 72.7$	$-0.6 \pm 4.7$
G041.15-00.20	Far	$-8.3 \pm 10.6$	$-12.9 \pm 14.6$	$+0.6 \pm 5.8$
G041.22-00.19	Far	$-12.2 \pm 12.6$	$-5.4 \pm 36.7$	$-0.1 \pm 6.4$
G043.03-00.45	Far	$+20.6 \pm 9.0$	$-5.1 \pm 19.4$	$-4.1 \pm 6.4$
G043.79-00.12	Far	$+19.2 \pm 10.5$	$-28.9 \pm 10.0$	$+3.4 \pm 10.4$
G043.89-00.78	Far	$-11.6 \pm 8.7$	$-17.3 \pm 5.9$	$+17.7 \pm 5.9$
G045.07+00.13	Far	$+7.5 \pm 9.7$	$+2.6 \pm 7.3$	$+9.4 \pm 9.6$
G045.45+00.06	Far	$-13.3 \pm 20.0$	$+0.1 \pm 20.4$	$-19.6 \pm 17.8$
G045.49+00.12	Far	$-10.4 \pm 11.3$	$-9.6 \pm 11.2$	$-0.4 \pm 6.1$
G045.80-00.35	Far	$-2.9 \pm 13.1$	$+2.6 \pm 17.1$	$-12.3 \pm 8.2$
G048.99-00.29	Far	$-1.0 \pm 16.9$	$+7.7 \pm 10.2$	$-13.0 \pm 14.4$
G049.19-00.33	Far	$+7.1 \pm 9.9$	$+8.3 \pm 5.0$	$+11.5 \pm 9.7$
G049.34+00.41	Far	$+10.7 \pm 8.8$	$+12.3 \pm 5.4$	$-1.4 \pm 5.8$
G049.59-00.24	Far	$+14.1 \pm 5.6$	$+5.3 \pm 5.0$	$-11.1 \pm 5.4$
G052.10+1.04	Far	$+3.8 \pm 7.9$	$-6.5 \pm 39.4$	$+0.1 \pm 4.2$
<hr/>				
Average	Near	$2.0 \pm 1.4$	$7.1 \pm 1.5$	$-1.1 \pm 1.5$
Average	Far	$4.3 \pm 2.4$	$-0.6 \pm 2.2$	$-0.2 \pm 1.6$
Average	All	$2.6 \pm 1.1$	$4.5 \pm 1.3$	$-0.7 \pm 1.1$

**Note.**  $U_s$ ,  $V_s$ , and  $W_s$  give the estimated peculiar (noncircular) velocity components of the central star that excites the masers, assuming  $R_0 = 8.31$ ,  $\Theta_0 = 241 \text{ km s}^{-1}$ , and solar motion components  $(U_\odot, V_\odot, W_\odot) = (10.5, 14.4, 8.9) \text{ km s}^{-1}$ .

## 7. Conclusions

We have measured trigonometric parallaxes and proper motions for 13 HMSFRs, which are located on the far side (past the tangent point) of the Sagittarius arm. Together with five published maser parallaxes, we studied the location, shape, and kinematics of the far portion of the Sgr arm. We find that the 6.7 GHz masers in this arm section are on average 15 pc below the IAU-defined Galactic midplane. The average pitch angle of the Sgr arm, now measured over  $\approx 10$  kpc, is  $7.2 \pm 1.9$ , but there is evidence that the pitch angle may not be constant along the arm. Peculiar motions for all but one source are less than  $20 \text{ km s}^{-1}$ , but there is a small average motion toward the Galactic center and in the direction of Galactic rotation.

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Facility: VLBA.

## ORCID iDs

Y. W. Wu (吴元伟)  <https://orcid.org/0000-0003-3188-5983>  
M. J. Reid  <https://orcid.org/0000-0001-7223-754X>  
T. M. Dame  <https://orcid.org/0000-0003-0109-2392>  
K. M. Menten  <https://orcid.org/0000-0001-6459-0669>  
A. Brunthaler  <https://orcid.org/0000-0003-4468-761X>  
B. Ho (胡波)  <https://orcid.org/0000-0002-3286-5469>  
B. Zhang (张波)  <https://orcid.org/0000-0003-1353-9040>

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