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Physical properties of the ambient medium and of dense cores in the Perseus star-forming region derived from *Herschel* Gould Belt Survey observations^{★,★★}

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ABSTRACT

The complex of star-forming regions in Perseus is one of the most studied due to its proximity (about 300 pc). In addition, its regions show variation in star-formation activity and age, with formation of low-mass and intermediate-mass stars. In this paper, we present analyses of images taken with the *Herschel* ESA satellite from 70 μm to 500 μm . From these images, we first constructed column density and dust temperature maps. We then identified compact cores in the maps at each wavelength, and characterised the cores using modified blackbody fits to their spectral energy distributions (SEDs): we identified 684 starless cores, of which 199 are bound and potential prestellar cores, and 132 protostars. We also matched the *Herschel*-identified young stars with *Gaia* sources to model distance variations across the Perseus cloud. We measure a linear gradient function with right ascension and declination for the entire cloud. This function is the first quantitative attempt to derive the gradient in distance across Perseus, from east to west, in an analytical form. We derived mass and temperature of cores from the SED fits. The core mass function can be modelled with a log-normal distribution that peaks at 0.82 M_{\odot} suggesting a star formation efficiency of 0.30 for a peak in the system initial mass function of stars at 0.25 M_{\odot} . The high-mass tail can be modelled with a power law of slope ~ -2.32 , which is close to the Salpeter's value. We also identify the filamentary structure of Perseus and discuss the relation between filaments and star formation, confirming that stars form preferentially in filaments. We find that the majority of filaments with ongoing star formation are transcritical against their own internal gravity because their linear masses are below the critical limit of 16 $M_{\odot} \text{pc}^{-1}$ above which we expect filaments to collapse. We find a possible explanation for this result, showing that a filament with a linear mass as low as 8 $M_{\odot} \text{pc}^{-1}$ can already be unstable. We confirm a linear relationship between star formation efficiency and the slope of dust probability density function, and we find a similar relationship with the core formation efficiency. We derive a lifetime for the prestellar core phase of 1.69 ± 0.52 Myr for the whole Perseus complex but different regions have a wide range in prestellar core fractions, suggesting that star formation began only recently in some clumps. We also derive a free-fall time for prestellar cores of 0.16 Myr.

Key words. stars: formation – circumstellar matter – stars: protostars – dust, extinction – submillimeter: ISM

1. Introduction

The aim of this paper is to derive a catalogue of cold compact cores of condensed dust and gas of $\lesssim 0.1$ pc in size in the Perseus molecular complex and to present and discuss their physical properties. This work uses *Herschel* observations and is part of the *Herschel* Gould Belt Survey (HGBS, André et al. 2010), which aims to probe the origin of the stellar initial mass function (IMF) by finding the physical mechanisms influencing star formation at its earliest stages. The HGBS team has already published core catalogues for many clouds: Aquila (Könyves et al. 2015, the reference paper for HGBS); L1495 in Taurus (Marsh et al. 2016); Lupus (Benedettini et al. 2018); Corona Australis (Bresnahan et al. 2018); Orion B (Könyves et al. 2020), Ophiuchus (Ladjelate et al. 2020);

Cepheus (Di Francesco et al. 2020); and Serpens (Fiorellino et al. 2021).

The analysis of these data has led to the idea that gravitational fragmentation of the cloud filaments is the dominant physical mechanism generating prestellar cores within interstellar filaments (see André et al. 2014, and references therein). It has been suggested in the past that the filamentary structure of the interstellar medium is fragmenting to form cores (see, e.g. Román-Zúñiga et al. 2009), but only with *Herschel* did it become clear that filament fragmentation is the dominant mode of (solar-type) star formation, because the physical properties and the spatial distribution of both the diffuse dust and the compact cold cores in the clouds can be derived simultaneously, with sufficiently high spatial resolution, namely of 36". Moreover, the large areas observed, combined with the high sensitivity of *Herschel* instruments, allow the construction of catalogues containing between two and nine times more cores than in previous ground-based surveys (e.g. Johnstone et al. 2000; Stanke et al. 2006; Enoch et al. 2006; Nutter & Ward-Thompson 2007; Alves et al. 2007).

* Full Tables A.1–A.3 are only available at the CDS via anonymous ftp to cdsarc.u-strasbg.fr (130.79.128.5) or via <http://cdsarc.u-strasbg.fr/viz-bin/cat/J/A+A/645/A55>

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The first prestellar core mass functions (CMFs) in star-forming regions, (e.g. [Motte et al. 1998](#); [Testi & Sargent 1998](#)) revealed a similarity between the shape of the CMF and the initial mass function (IMF) of the stars. The system IMF shows ([Chabrier 2005](#)) a log-normal distribution with a peak at $0.25 M_{\odot}$, and a high-mass tail for $M \geq 1 M_{\odot}$ modelled with a power-law function $dN/d\log(M) \propto M^{-1.35}$, also known as the Salpeter slope ([Salpeter 1955](#)).

A quantitative assessment of the similarity between CMF and IMF was carried out for the Pipe nebula by [Alves et al. \(2007\)](#) who found that the peak of the CMF was about three times higher than the peak of the IMF. Assuming a one-to-one correspondence between cores and stars, the shift in the respective peaks of the CMF and the IMF was interpreted by [Alves et al. \(2007\)](#) as an indication of a star formation efficiency (SFE) per single core of about 0.3, meaning that 30% of the core mass eventually becomes star mass. Later studies confirmed that values around 0.3–0.4 are typical in many other star-forming regions even if different values are also reported (see Sect. 5.1).

The similarity between CMF and IMF led to the hypothesis that the IMF is a consequence of the physical mechanisms that produce the prestellar core population. Testing this hypothesis requires derivation of the CMF in many star-forming regions in the most uniform way possible with respect to how observations are executed as well as to how cores are detected and their intensities are measured. Only in this way can the different CMFs be compared consistently (see, e.g. [Swift & Beaumont 2010](#)).

Another issue on a proper CMF derivation is that it is very important to assign each core its temperature rather than using a constant temperature for all the cores in a cloud, an assumption that leads to inaccurate core masses. Unless we determine both mass and temperature at the same time, the stability of cores cannot be investigated on a purely photometric basis, and the derived CMF can contain collapsing cores and transient structures that are not forming stars, as pointed out 20 yr ago by [Johnstone et al. \(2000\)](#). Deriving the CMF for the most important nearby star-forming regions in a robust and uniform way is the main aim of the HGBS.

The Gould Belt is a complex of neutral gas and star-forming molecular clouds covering an area of $\sim 2.6 \times 10^5 \text{ pc}^2$ (e.g. [Bobylev 2016](#)); the molecular clouds are located within distances of $d \lesssim 500 \text{ pc}$ of the Sun, which are close enough to observe individual cores. Among them, the system of clouds in Perseus (Fig. 1) are at distances in the range of $\sim 230\text{--}320 \text{ pc}$ ([Hirota et al. 2008, 2011](#); [Strom et al. 1974](#)) with many sites of active low-mass and intermediate-mass star formation. The review by [Bally et al. \(2008\)](#) gives a complete presentation of the region, including its surroundings, with a long list of references; here we limit the literature references to papers that are pertinent to our work or when their results are used in this study.

Perseus was observed as part of the *Spitzer* Cores to Disks (c2d) program ([Evans et al. 2009](#)) with the InfraRed Array Camera (IRAC, [Jørgensen et al. 2006](#)) and the Multiband Imaging Photometer (MIPS, [Rebull et al. 2007](#)). The analysis of both datasets was later presented by [Young et al. \(2015\)](#). Perseus was also mapped at sub-millimetre (sub-mm) wavelengths with SCUBA ([Hatchell et al. 2005](#)) and later with SCUBA-2 ([Chen et al. 2016](#)), and at 1.1 mm with BOLOCAM ([Enoch et al. 2006](#)).

Examples of previous publications related to HGBS data are the study of B1-E ([Sadavoy et al. 2012](#)) and of the relation between *Spitzer* Class 0 and the distribution of dense gas ([Sadavoy et al. 2014](#), see also Sect. 3.2.4 of this work); [Pezzuto et al. \(2012\)](#) discussed the properties of two sources, B1-bN and

B1-bS, that were proposed to be first hydrostatic core candidates (see Sect. 4.3).

The present paper is organised as follows. In Sect. 2, we give an overview of the observations of Perseus and present our data reduction and analysis. In Sect. 3, we derive a first quantitative estimate of the distance gradient across the cloud using *Gaia* Data Release 2 (DR2) results. We then derive and discuss the column density and temperature maps, and the filamentary structure of the region. In Sect. 4, we present the source extraction and the physical properties of the cores derived via SED fitting. We also discuss the stability of cores against their internal gravity. In Sect. 5, we present the CMF and link the cores to the filaments. We also analyse the core and star formation efficiencies and estimate the lifetimes of the different phases of star formation. In Sect. 6, we summarise our conclusions. We include several appendices with the full source catalogue, completeness testing, region definitions, and the catalogue of sources found in our maps that are not related to star formation, such as galaxies for example.

It is worth concluding this section by introducing some of the terms that are used in the following sections. We define a core as a compact over-density of gas and dust that is round in shape and exceeds the density of its local diffuse background emission appearing as a local maximum in an intensity image or in a column density map. Cores are most easily identified using mid-infrared to mm wavelengths ([di Francesco et al. 2007](#); [Ward-Thompson et al. 2007](#); [André et al. 2014](#)).

A core is defined as starless if there is no internal source of energy (e.g. a protostar). Starless cores are instead warmed by the interstellar radiation field and their spectral energy distribution (SED) can be modelled as a single modified blackbody $I_{\nu} \propto \nu^{\beta} B_{\nu}(T)$ at temperature T . We use our source extraction technique and the modified blackbody fits to derive the radius R , mass M , and temperature T of the cores, and we use these physical properties to determine the dynamical state of the cores. If the core self-gravity exceeds its pressure support, we consider the source to be bound. We define bound, starless cores as prestellar cores. These objects can collapse, or are collapsing, and will likely form one or more stars. If the self-gravity of a core is insufficient to balance its internal pressure, the core is classified as unbound and may be a transient structure that will dissipate in the future unless it is confined by another mechanism (e.g. pressure confined).

Once a star forms, it warms up the surrounding envelope, whose emission in the first phases still resembles a modified blackbody for $\lambda \gtrsim 100 \mu\text{m}$ while at shorter wavelengths the SED is no longer compatible with such a model. For this reason, a core with compact $70 \mu\text{m}$ emission is considered protostellar because a central source must be present to warm the dust. However, we note that in principle a starless core can be detected at $70 \mu\text{m}$ if it is sufficiently warm ($T \gtrsim 20 \text{ K}$), and so the shape of the SED at short wavelengths determines if the object is already a protostar. On the other hand, an object undetected at $70 \mu\text{m}$ is always considered a starless core even if the lack of detection in the PACS band(s) may be just a matter of sensitivity.

The focus of this paper is on starless cores. Protostars are mentioned here when necessary, for instance to derive star formation efficiency in Sect. 5.3, but a full discussion on the protostars in Perseus is postponed to a forthcoming paper.

2. Observations and data reduction

Perseus was observed with *Herschel* ([Pilbratt et al. 2010](#)) as part of the HGBS ([André et al. 2010](#)) in two overlapping mosaics:

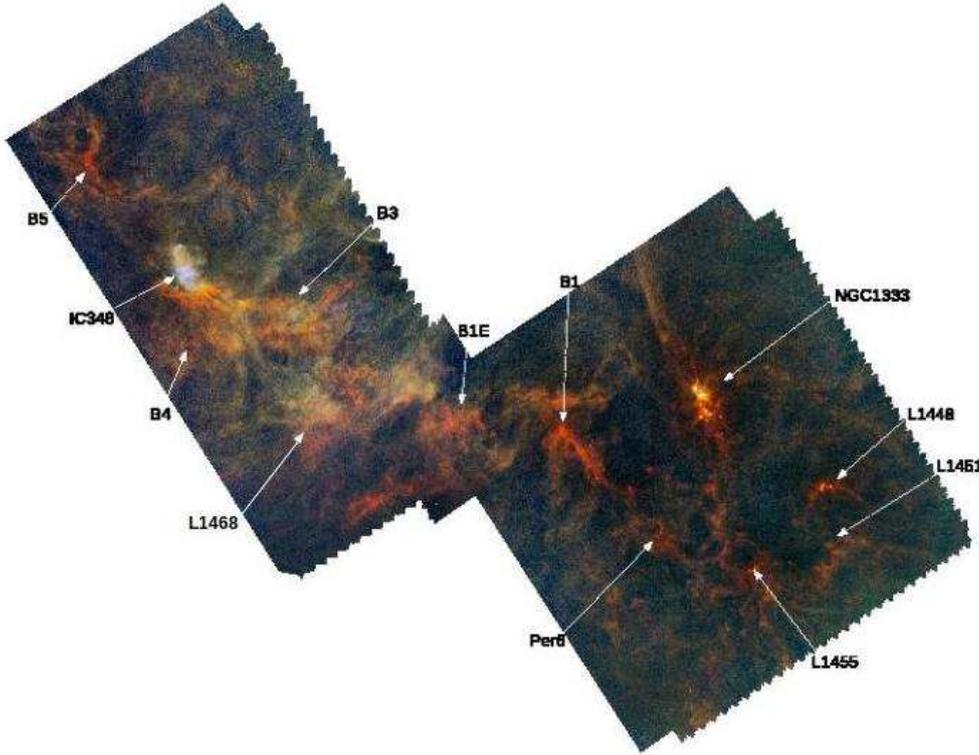


Fig. 1. Composite image of the star-forming region in Perseus. Blue is PACS 70 μm ; green is PACS 160 μm ; red is SPIRE 250 μm ; north is up, east is to the left. The latter two maps have been processed to simulate having the highest resolution of the 70 μm . The resulting image is therefore not meant to be used for scientific analysis. Labels on the map identify the approximate position of main subregions. Monochromatic *Herschel* intensity maps are reproduced in Appendix D with coordinate grids.

Table 1. Log of the observations.

	West Perseus	East Perseus
OBSID	1342190326 1342190327	1342214504 1342214505
Date	09/02/2010	19/02/2011
Centre	3 ^h 29 ^m 34 ^s +30 ^d 54 ['] 07 ["]	3 ^h 42 ^m 51 ^s +32 ^d 00 ['] 49 ["]
Size	135' \times 150'	90' \times 184'

Notes. OBSID: identifier in the *Herschel* Science Archive. Date: start of the observation; Centre: the J2000 central coordinates of each field; Size: requested size.

the western field (mainly NGC 1333, B1, L1448, L1455) and the eastern field (L1468, IC348, B5). Results from these observations were initially presented in Sadavoy et al. (2012, 2014), Pezzuto et al. (2012), and Zari et al. (2016). Both fields were observed with PACS (Poglitsch et al. 2010) at 70 μm (blue) and 160 μm (red), and with SPIRE (Griffin et al. 2010), at 250 μm (PSW), 350 μm (PMW), and 500 μm (PLW) in parallel mode with the telescope scanning at a speed of 60''s⁻¹. The total area common to both instruments is about 13 square degrees. Table 1 gives the observation log.

A composite image using PACS bands and the SPIRE PSW band is shown in Fig. 1. All maps are reported to the same spatial resolution as the 70 μm map by using the method presented in Li Causi et al. (2016).

Each field was observed twice along two almost orthogonal directions to remove better instrumental 1/*f* noise. The *Herschel* raw data were reduced with HIPE (Ott 2010) version 10. For PACS, we reduced the data to Level 1 with HIPE using version 45 of the set of calibration files. Level 1 data consist of the calibrated timelines of the detectors to which the celestial coordinates were added. Afterwards, the data were exported outside HIPE and maps were generated with *Unimap* (Piazzo et al. 2015)

version 6.5.3. The adopted spatial grid is 3'' and 4'' for the blue and red bands, respectively.

For SPIRE, we used the calibration files version 10.1 and maps were generated with the *destriper* module in HIPE. We adopted the default values for the spatial grids: 6'', 10'', and 14'' for PSW, PMW, and PLW bands, respectively. The SPIRE maps are given in Jy beam⁻¹, which we converted to MJy sr⁻¹ using the nominal beamsizes of 426 arcsec², 771 arcsec², and 1626 arcsec², respectively. Additional details on the PACS and SPIRE data processing can be found in Könyves et al. (2015).

The alignment of each image was checked by comparing the *Herschel* position at 70 μm of 22 bright and isolated sources, with their corresponding positions reported in the *Spitzer* catalogue at 24 μm (Evans et al. 2009), that is, 11 from both the western and eastern parts. For the western field, we found a shift in RA of $\Delta\alpha = 4''.82 \pm 0''.56$ which was added to our coordinates. We also applied a similar shift to the red map. SPIRE maps required no correction. The final residuals in the five maps run from -1/4 of a pixel in the PSW map to +1/5 of a pixel in the red map for right ascension, and from 1/10 of a pixel in the PSW map to 1/3 of a pixel in the blue map for declination.

The PACS and SPIRE maps were made using a common coordinates system, at each band, for both Perseus east and west. A zero-level offset was applied to each map, obtained by correlating the *Herschel* data with *Planck* and IRAS data (Bernard et al. 2010). All five maps shown in Appendix D are available on the HGBS archive¹.

3. Data analysis

3.1. Distance to Perseus

Given that physical parameters such as mass and radius depend on the distance to the sources, we start this section with a discussion on the distance to Perseus. For IC 348, in the eastern

¹ <http://gouldbelt-herschel.cea.fr/archives>

Table 2. Sources from our protostar catalogue with a *Gaia* counterpart within 10".

ID	RA	Dec	Distance	ID	RA	Dec	Distance	ID	RA	Dec	Distance
1)	51.27800	31.11428	299 ± 10	10)	52.23893	31.23733	291 ± 10	19)	55.41336	31.60310	300 ± 11
2)	51.28956	30.77220	285 ± 65	h2)	52.26552	31.26772	240 ± 36	20)	55.49376	31.81556	312 ± 25
h1)	51.41199	30.73479	237 ± 26	11)	52.37169	31.30924	315 ± 24	21)	55.72903	31.72932	330 ± 13
3)	51.45769	31.17290	292 ± 11	12)	52.47598	31.34729	296 ± 17	22)	55.73202	31.97893	313 ± 10
4)	51.61836	31.20151	292 ± 14	13)	52.71848	30.90509	327 ± 66	23)	56.11088	32.07526	330 ± 19
5)	51.94876	30.20121	250 ± 29	14)	52.82618	30.82677	302 ± 12	24)	56.14806	32.15777	333 ± 89
6)	52.15382	31.29370	374 ± 60	15)	52.87012	30.51453	295.1 ± 8.6	25)	56.31799	32.10562	308 ± 23
7)	52.18035	31.29248	205 ± 36	16)	53.14148	31.01549	307 ± 19	26)	56.33484	32.10952	310 ± 20
8)	52.21291	31.30397	296 ± 31	17)	55.19518	32.53161	308 ± 13	27)	56.45099	32.40305	311 ± 13
9)	52.23529	31.30954	345 ± 40	18)	55.28775	31.74389	326 ± 28	28)	56.94595	33.06746	343 ± 76

Notes. We listed also h1 (L1448 C) and h2 (SVS13 in NGC1333) which host masers from which an astrometric distance was obtained (Hirota et al. 2008, 2011). For all the sources, the distance in parsec was derived from parallax according to Luri et al. (2018).

part of the complex, Herbig (1998) discussed a variety of distance measurements in the literature and concluded that the value 316 ± 22 pc derived by Strom et al. (1974) is the most reliable. This distance was later confirmed indirectly by Ripepi et al. (2002) who detected δ -Scuti oscillations in one star of IC 348, and concluded that the derived star properties were not compatible with distances ≤ 260 pc, but were compatible with 316 pc. On the other hand, the western half of the Perseus region is known to be closer, which led to the conclusion that the molecular clouds in Perseus form a chain with increasing distance from west to east (see Černis 1993). This conjecture was definitively confirmed with the measure of maser parallaxes for a couple of sources in NGC 1333 and L1448 (Hirota et al. 2008, 2011) giving a distance of ~ 240 pc.

To gain some insight into this gradient in distance, we exploited the *Gaia* archive second data release (DR2) of astrometric data (Gaia Collaboration 2016). We looked for sources in our protostar catalogue (see Appendix A) that are spatially close to a source in the *Gaia* archive with an angular separation of less than 10". Objects with negative parallax were excluded. All sources had spatial correspondences of $< 8''$, with only two having separations $> 6''$ (6.5" and 7.4"). Conversely, 18 are spatially separated by $< 2''$ and 7 are separated by $< 1''$. Parallaxes were converted into distances following Luri et al. (2018). In particular, we used the `DistanceEstimatorApplication.py` tool² with the median of the *Exponentially Decreasing Space Density Prior* distribution option. The tool also returns the 5% and 95% quantiles that are not symmetrical around the median (this reflects the obvious prior that the distance must be positive). As the difference between the median minus 5% quantile and the 95% quantile minus the median is negligible, we adopted the latter for the distance uncertainty which is slightly larger than the former. No attempt was made to exclude projection effects. Therefore, some associations between our sources and *Gaia* stars may not be physical. On the other hand, it is unlikely that a proximity of a few arcseconds between a *Gaia* source at the distance expected for Perseus and a protostar is only due to projection.

In Table 2 we show distances and uncertainties for the 28 sources possibly associated with our protostars. We also report the distances for the maser sources in L1448 (Hirota et al. 2011, h1 in the table) and NGC 1333 (Hirota et al. 2008, h2) derived

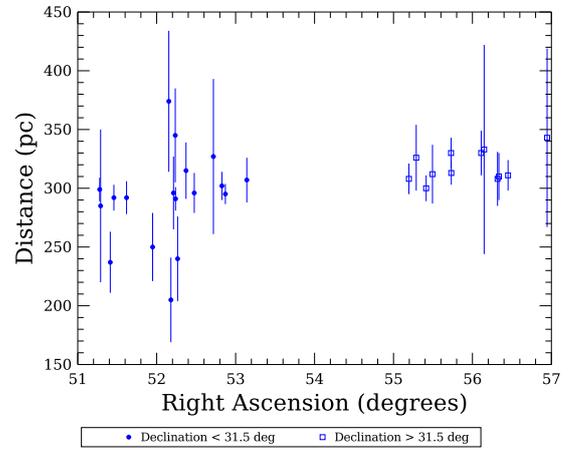


Fig. 2. Relation between distance and right ascension for the 30 sources in Table 2.

using the python tool reported above to convert the parallax to distance.

In Fig. 2 the 30 distances are plotted as a function of their right ascension. Starting from the observational finding that a distance gradient exists going from west to east, and noting that a gradient also exists going from south to north (see the different symbols of the points in the figure), we made a multivariate linear fit to the data and found the following relation:

$$d(\text{pc}) = 281.8 \pm 3.4 + (5.13 \pm 0.65) \times (\alpha - \alpha_0) + (5.5 \pm 1.9) \times (\delta - \delta_0), \quad (1)$$

where $\alpha_0 = 51:27800$ and $\delta_0 = 30:20121$ are the smallest right ascension and declination of sources 1 and 5 in Table 2, respectively. Such a relation, namely $d \equiv d(\alpha, \delta)$, is the first quantitative estimate of the distance gradient in Perseus even if the large scatter in distances could be due to false associations between our protostars and *Gaia* sources. In particular, if all the associations in NGC 1333 are physical, an exceptional depth of more than 100 pc is implied, or an overlap along the line of sight of at least two different clouds (see also Appendix C). Further, there are no data points (no associations) in the range $53:14 \leq \alpha \leq 55:20$, and so the middle part of the fit is not constrained. A refinement of Eq. (1) will be attempted in a forthcoming paper on protostars, where Class II pre-main sequence stars – which are more easily detected by *Gaia* – will be distinguished from younger sources,

² <https://github.com/agabrown/astrometry-inference-tutorials/tree/master/single-source/tutorial>

Table 3. Comparison between distances, in pc, derived by different author.

Name	RA	Dec	P	Z	O
L1451	51.0	30.5	282.3 ± 3.5	279 ± 14	
L1448	51.2	30.9	285.5 ± 3.7	288 ± 14	
NGC1333	52.2	31.2	292.3 ± 4.0	299 ± 14	293^{+24}_{-21}
B1	53.4	31.1	297.9 ± 4.1	301 ± 15	
IC348	55.8	31.8	314.0 ± 5.5	295 ± 15	320^{+28}_{-24}
B5	56.9	32.9	325.7 ± 7.2	302 ± 16	

Notes. RA and Dec are in degrees as quoted by Zucker et al. (2018); P distance based on our Eq. (1); Z distance from Zucker et al. (2018); O distance from Ortiz-León et al. (2018).

Table 4. Comparison between the distance, in parsecs, derived in this paper (column P) and that derived by Zucker et al. (2019) (column Z) for the four lines of sight reported in their Table 3 for the Perseus cloud.

RA	Dec	P	Z
51.5455	30.6747	285.7 ± 3.5	289 ± 15
52.0097	30.1235	285.1 ± 3.4	279 ± 14
52.0315	30.6110	287.9 ± 3.5	289^{+16}_{-15}
52.4212	31.3616	294.0 ± 4.1	281^{+15}_{-14}

that is, Class 0 objects. The latter, being heavily obscured in the visible light, are more difficult to detect with *Gaia* and thus an association between a Class 0 source and a *Gaia* object may be due to a projection effect.

Distance measurements to the main Perseus clouds were recently reported by Zucker et al. (2018) and Ortiz-León et al. (2018). Table 3 compares these distances: column “ P ” gives the distance based on our Eq. (1) using right ascension and declination from Zucker et al. (2018), column “ Z ” gives the Zucker et al. (2018) distances, and column “ O ” gives the Ortiz-León et al. (2018) values. The distances from Zucker et al. (2018) are derived by combining data from PAN-STARRS1, 2MASS, *Gaia* and the ^{12}CO data from the COMPLETE survey (Ridge et al. 2006a) (we use the ^{13}CO data in Sect. 3.2.2 and Appendix C). Ortiz-León et al. (2018) measure distances using VLBA and *Gaia* data. Uncertainties in column P , derived through error propagation from Eq. (1), can be misleading. These are small because they give the (estimated) distance of each point in the sky assuming a 2D geometry for Perseus. These uncertainties do not take into account the dispersion in distance caused by the depth of each cloud, and they do not take into account the fact that there are no protostars in the region with right ascension between $\sim 53^\circ$ and $\sim 55^\circ$, and therefore we do not have an indication of the distance. Uncertainties in column Z are the sum in quadrature of the random and systematic errors given by Zucker et al. (2018).

We find excellent agreement (<3% differences) between our measurements and those of Zucker et al. (2018) and Ortiz-León et al. (2018). Only IC348 and B5 show more significant differences but still <10% and, in any case, within the uncertainties. Zucker et al. (2019) derived a distance map of Perseus and other molecular clouds through star-specific distance-extinction estimates improved with *Gaia* parallaxes. The result for a small set of lines of sight is reported in Table 4 where we compare their distances with ours: the agreement is very good, better than 5% in all cases.

The full set of distances derived by Zucker et al. (2019) is shown in Fig. 3 where we give a graphical representation of Eq. (1): our distance was computed for each point of the column density map (see Sect. 3.2) and colour-coded according to the bar on the right of the figure. Zucker et al. (2019) give the distance averaged over boxes of 0.84 deg^2 , and these are shown in green at the position of the centre of each box in Fig. 3.

While Zucker et al. (2019) computed accurate distances to Perseus that are more precise than those given by our Eq. (1), it is nevertheless not easy to extract a distance for a given point in space from their map; for instance, B5 is seen to be at 302 pc in Zucker et al. (2018), a distance that can hardly be derived from the grid of distances shown in Fig. 3. Conversely, our Eq. (1) has been derived in a less rigorous way, but it immediately gives a reference distance at any point (α, δ) in space, a distance that captures the distance gradient in the complex, and closely matches the distances derived by other authors.

For the rest of the paper, we adopted the mean of the 30 sources, 300 pc, as a *representative* value of the distance for the whole Perseus complex (294 pc in Zucker et al. 2019). However, below we show how the physical description of the whole region changes when distances are varied according to our formula.

3.2. Dust temperature and column density maps

Dust emission in the far-infrared can be modelled as a modified blackbody $I_\nu = \kappa_\nu \Sigma B_\nu(T_d)$, where κ_ν is the dust opacity per unit mass expressed as a power law normalised to the value κ_0 at a reference frequency ν_0 , such that $\kappa_\nu = \kappa_0(\nu/\nu_0)^\beta \text{ cm}^2 \text{ g}^{-1}$. Here, Σ is the gas surface-density distribution $\mu_{\text{H}_2} m_{\text{H}} N(\text{H}_2)$ with μ_{H_2} being the molecular weight, m_{H} the mass of the hydrogen atom, and $N(\text{H}_2)$ the column density assuming the gas is fully molecular hydrogen. Finally, $B_\nu(T_d)$ is the Planck function at the dust temperature T_d . Following the conventions of the HGBS (Könyves et al. 2015), we adopted an opacity of $0.1 \text{ cm}^2 \text{ g}^{-1}$ at $300 \mu\text{m}$ (gas-to-dust ratio of 100), a dust opacity index $\beta = 2$, and $\mu_{\text{H}_2} = 2.8$ representing the mean weight per hydrogen molecule (Kauffmann et al. 2008).

We used the four intensity maps from $160 \mu\text{m}$ to $500 \mu\text{m}$ to derive $N(\text{H}_2)$ and T_d maps. The $70 \mu\text{m}$ map was not included because emission at this wavelength may include contributions from very small grains (VSGs), which cannot be modelled with a simple single-temperature modified blackbody. Indeed, Schnee et al. (2008) found that in Perseus contribution from VSGs may elevate the $70 \mu\text{m}$ emission by 70% at 17 K and by 90% below 14 K. Contribution from VSGs is not expected for $\lambda > 100 \mu\text{m}$ (Li & Draine 2001).

The three intensity maps shortwards of $500 \mu\text{m}$ were degraded to the spatial resolution of the PLW band ($36''$), and then the four maps were regridded onto a common coordinate system using $14''$ pixels at all wavelengths.

The SED fitting procedure was executed pixel by pixel with a code that takes the four images in FITS format as input. The code creates a grid of models, as in Pezzuto et al. (2012), by varying only the temperature in the range $5 \leq T_d(\text{K}) \leq 50$ in steps of 0.01 K. For each temperature T_j , the code computes the intensity at far-infrared wavelengths for a fixed column density of 1 cm^{-2} .

As I_ν is linear with $N(\text{H}_2)$, we can compute the column density at each pixel using a straightforward application of the least-squares technique,

$$N(\text{H}_2)_j = \frac{\sum_i f_i q_i(T_j) / \sigma_i^2}{\sum_i q_i^2(T_j) / \sigma_i^2}, \quad (2)$$

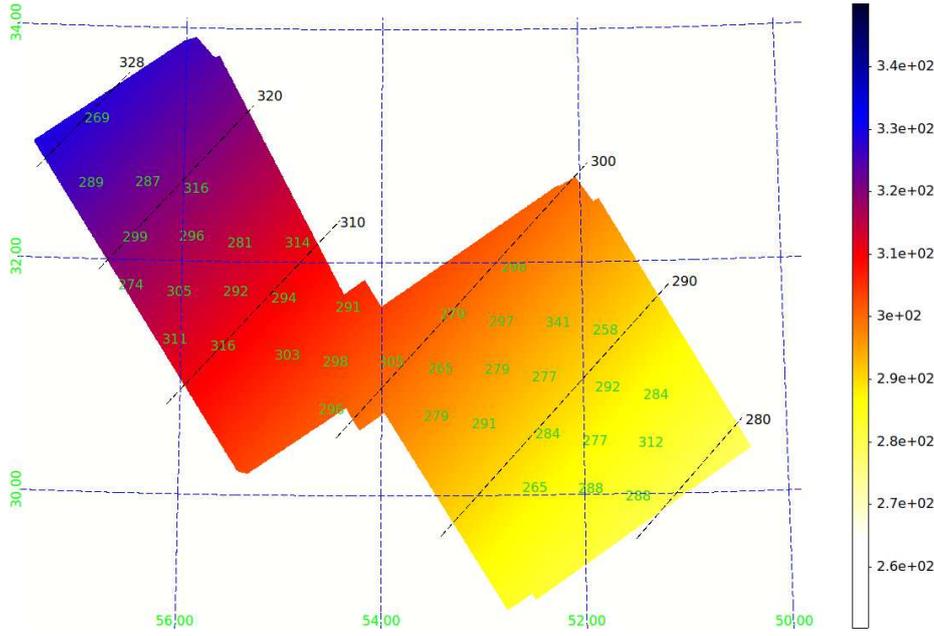


Fig. 3. Color-coded 2D visualisation of Eq. (1). Color bar is in parsec, coordinates grid in degrees. Black dashed lines connect points at the same distance, shown in the labels. The coordinates grid, in blue, is in steps of 2° . Green labels are the distances derived by Zucker et al. (2019) in boxes of $0.84''$, the labels are at the centre of the boxes.

where f_i is the observed SED I_ν at each frequency i , and $q_i(T_j)$ is the synthetic SED model at frequency i and temperature T_j . The uncertainty σ_i is 10 and 20% for SPIRE and PACS respectively (Könyves et al. 2015). Index i runs from 1 to 4, index j runs from 1 to 4501, which is the number of models in the grid.

At each pixel of the intensities maps, the model j with the smallest residuals is kept as the best-fit model. To be consistent with the other HGBS papers, we ran the code without applying colour corrections but in Sect. 3.2.2 we show how much the results depend on this assumption and on the choice of $\beta = 2$. The code, available on request, outputs the column density map, the temperature map, and the map of the uncertainty of $N(\text{H}_2)$.

Another column density map at higher spatial resolution (i.e. $18''.2$ corresponding to the SPIRE $250 \mu\text{m}$ band) was also obtained using the method described in Palmeirim et al. (2013). The procedure is based on a multi-scale decomposition technique. The high-resolution column density map used to optimise the source extraction described in Sect. 4.1 is available for download from the HGBS website. This map, shown in Appendix D, has a higher spatial resolution than those previously obtained with the same *Herschel* data at $36''.1$ (Sadavoy et al. 2014; Zari et al. 2016).

The column density and temperature maps of Perseus at the PLW spatial resolution are shown in the two panels of Fig. 4. The top panel shows the column density map and the bottom panel shows the temperature map. Contour levels in the column density map are $3 \times 10^{21} \text{ cm}^{-2}$ and 10^{22} cm^{-2} . The first level follows the border of the known main regions relatively well, even if B3 and B4 form a single complex with IC348, while the other marks the densest part of the clouds.

The histogram of T_d is shown in Fig. 5. There are two peaks seen at 16.4 K and 17.1 K, and a third, very broad one, indeed a plateau, at 19.2 K. The first two peaks correspond to the temperatures of the diffuse medium in west and east Perseus, respectively. In other words, the dust temperature is slightly lower in the western half of Perseus than in the eastern half. The peak at 19.2 K reflects the inner parts of NGC 1333 and IC 348, and a few other regions that we discuss below.

To check our results, we compared the *Herschel* $N(\text{H}_2)$ map with the all-sky *Planck* map of the optical depth τ_p at

$850 \mu\text{m}$ (Planck Collaboration XI 2014)³. First we convolved our map to the *Planck* resolution of $5'$ and projected the result onto the *Planck* grid. We then computed the ratio $r = \tau_H / \tau_p$ where $\tau_H = 0.1(3/8.5)^2 \mu\text{m}_H N(\text{H}_2)$. Because of the convolution, the *Herschel* column density and, as a consequence, the optical depth, have very low values close to map borders, with τ_H as low as 6×10^{-21} . To exclude these points, we made the comparison in the region where $\tau_H \geq 1.38 \times 10^{-5} = \min(\tau_p)$.

In Fig. 6 we show the ratio r versus τ_p ; we also highlight the small values of r corresponding to $\tau_H < 1.38 \times 10^{-5}$ (green points). The blue histogram shows the mean ratios, averaged in bins of 10^{-4} , excluding the green points. The weighed mean of the histogram values is $\bar{r} = 1.174 \pm 0.040$. The blue points are compatible with a constant ratio for r . On the other hand, an increasing trend of r with τ_p seems present in the figure, with $r \lesssim 1.05$ for $\tau_p < 3 \times 10^{-4}$, and increasing up to 1.30 when $\tau_p \sim 8 \times 10^{-4}$. In particular, 98% of the points of the optical depth map have $1.38 \times 10^{-5} \leq \tau_p \leq 3 \times 10^{-4}$, and for them the ratio r is 1.05 ± 0.15 . A change in r might be evidence of a change in opacity at high column density.

3.2.1. Possible contribution of very small grains

As written, the $70 \mu\text{m}$ map was not used to derive the column density map, because of possible VSG emission at this wavelength. However, we can estimate this contribution by computing the intensity map at $70 \mu\text{m}$ from the column density and temperature maps obtained for $\lambda \geq 160 \mu\text{m}$. Figure 7 compares the extrapolated $70 \mu\text{m}$ map with the difference map between the extrapolated and observed data. All maps have been convolved to $36''.1$ resolution. The most striking feature of Fig. 7 is that the dense subregions in Perseus are expected to be dark against the brighter diffuse background emission if observed with an instrument of high-enough sensitivity. Such $70 \mu\text{m}$ dark clouds are typically classified as infrared dark clouds (IRDCs, Carey et al. 1998). However, the observed data (see Appendix D) do not show the subregions as silhouettes.

³ HFI_CompMap_ThermalDustModel_2048_R1.20.fits available at <http://pla.esac.esa.int/pla>

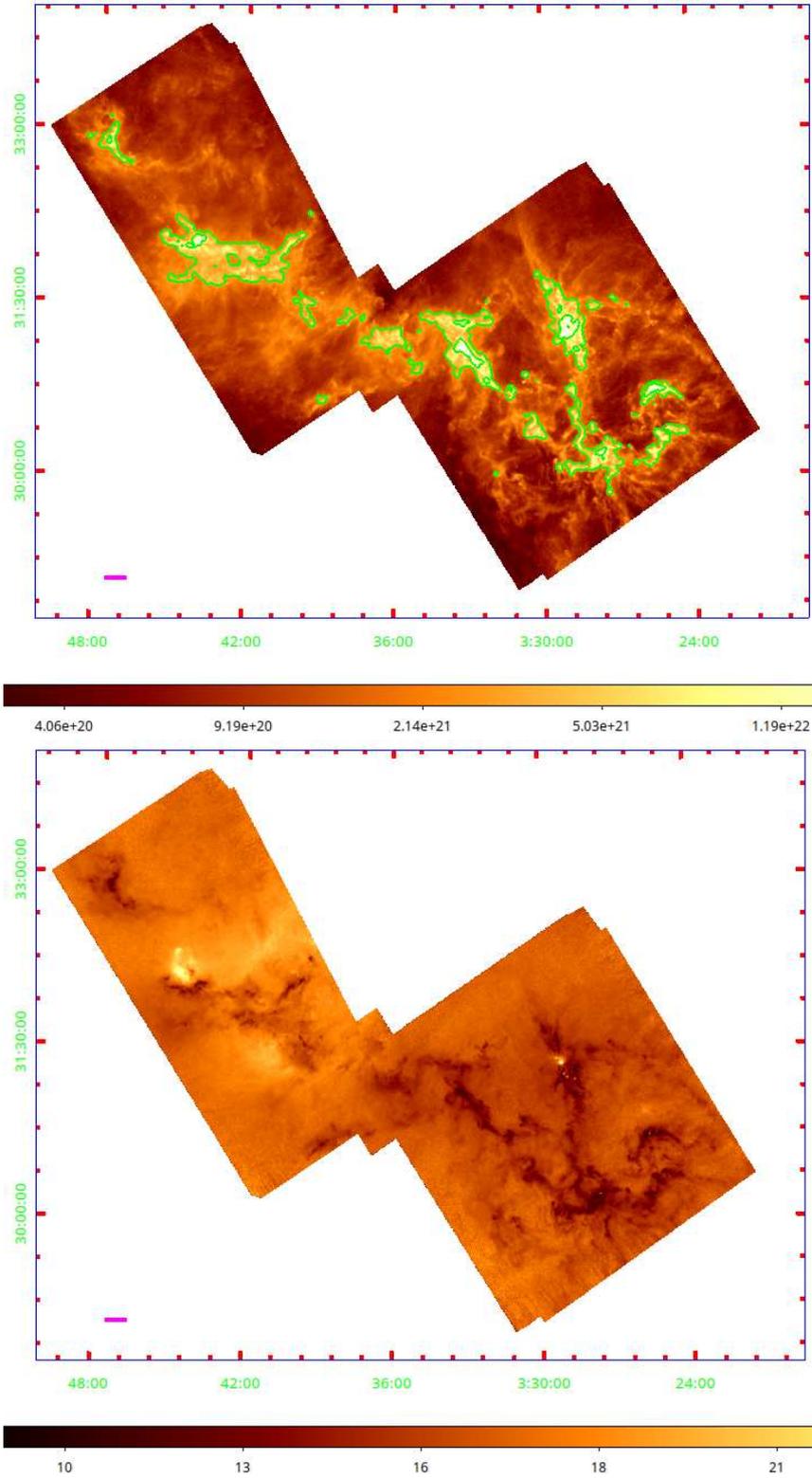


Fig. 4. *Top panel:* column density map with contours at $3 \times 10^{21} \text{ cm}^{-2}$ and 10^{22} cm^{-2} . *Bottom panel:* dust temperature map. In both panels, the magenta line in the bottom-left corner shows the angular scale corresponding to 1 pc at 300 pc; J2000.0 coordinates grid is shown. Both maps have a spatial resolution of $36''$. The anticorrelation between $N(\text{H}_2)$ and T_d is evident: regions at high/low column density are cold/warm, with few exceptions in IC348 and NGC 1333.

The right panel of Fig. 7 shows the difference between the observed $70 \mu\text{m}$ map, degraded in spatial resolution and projected onto the $500 \mu\text{m}$ spatial grid, and the computed $70 \mu\text{m}$ map. If the far-infrared dust emission can be extrapolated by a modified blackbody distribution, then the difference map should approach zero everywhere. This is indeed true for most of Perseus with the exception of NGC 1333 and the complex IC348/B3/B4/L1468, with differences as high as 10^3 MJy sr^{-1} in NGC 1333.

As a first hypothesis we suppose that there is not a population of VSGs and that the differences in observed and computed $70 \mu\text{m}$ flux are due to poor estimation of T_d . In fact, for a modified blackbody, T , β , and the wavelength where the SED peaks λ_{peak} are related through the relation (Elia & Pezzuto 2016)

$$\lambda_{\text{peak}} = \frac{1.493}{T(K)(3 + \beta)} \text{ cm.} \quad (3)$$

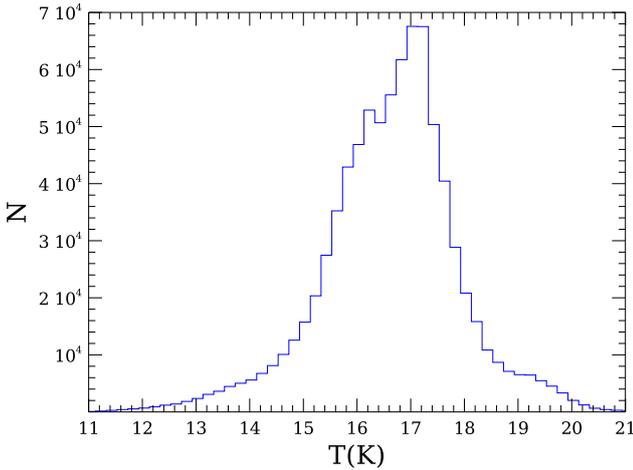


Fig. 5. Histogram of dust temperature in the range 11–21 K (minimum is 10.1 K, maximum is 28.3 K).

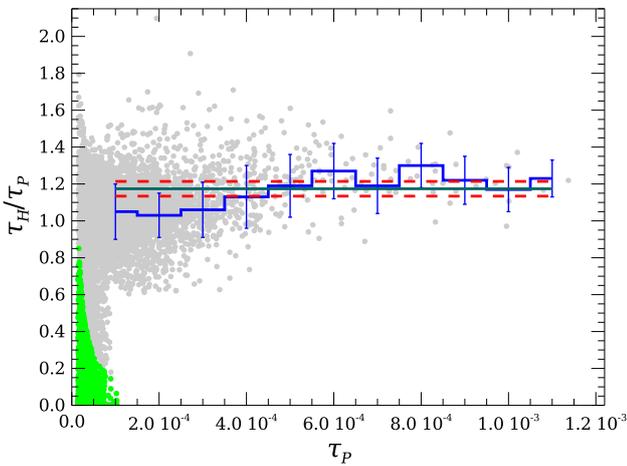


Fig. 6. Ratio of τ_H , the optical depth derived from our *Herschel* observations, to τ_P , the optical depth derived from *Planck* observations (Planck Collaboration XI 2014) (grey points), with both τ values computed at $850 \mu\text{m}$. The ratio is plotted vs. τ_P . Green points correspond to low values of τ_H due to the convolution applied to the *Herschel* column density map (see text). Mean values \pm one standard deviation of the grey points, excluding the green ones, are shown with the blue histogram. The dark green and the two red lines show the weighted mean of the ratio: 1.174 ± 0.040 .

If $T = 18.7 \text{ K}$, the SED peaks at $\lambda_{\text{peak}} = 160 \mu\text{m}$ if $\beta = 2$. At higher temperatures, λ_{peak} moves shortwards of $160 \mu\text{m}$, meaning that the peak wavelength, and therefore T , is poorly determined from our dataset built for $\lambda \geq 160 \mu\text{m}$. To quantify this effect, we derived the column density and temperature maps using only PACS data. To this aim we used an additional PACS intensity map at $100 \mu\text{m}$, also observed as part of the HGBS and that will be the subject of a future paper focused on Class 0 objects.

The $70 \mu\text{m}$ and $100 \mu\text{m}$ maps were degraded to the red band spatial resolution, projected onto the spatial grid of the latter image, and then fitted following the same pixel-by-pixel technique described previously. As the $100 \mu\text{m}$ observation covered a smaller area than the PACS/SPIRE parallel-mode observations, and because the PACS intensities have larger uncertainties in the zero-level of the diffuse emission, the resulting $N(\text{H}_2)$ and T maps offer less coverage and higher noise when compared to the “nominal” maps obtained for $\lambda \geq 160 \mu\text{m}$.

The new temperature map, $T_{\text{PACSoonly}}$, was then degraded to the $500 \mu\text{m}$ spatial resolution and projected onto the spatial grid of the nominal T map, $T_{160+\text{SPIRE}}$. Figure 8 shows the mean of the ratio $T_{\text{PACSoonly}}/T_{160+\text{SPIRE}}$ as a function of $T_{\text{PACSoonly}}$ in bins of 1° K . At low temperatures, $T_{\text{PACSoonly}}$ is a poor estimate of T , as expected. However, the trend for $T_{\text{PACSoonly}} > 20 \text{ K}$ is more interesting. In this regime, corresponding to $\lambda_{\text{peak}} \lesssim 150 \mu\text{m}$, the peak of the SED falls outside the range of wavelengths used to derive $T_{160+\text{SPIRE}}$. On the other hand, the PACS bands span the SED peak, making $T_{\text{PACSoonly}}$ a more reliable measurement of the dust temperature, at least as long as $T \lesssim 43 \text{ K}$ (above this temperature the SED peak moves at $\lambda < 70 \mu\text{m}$). From the figure, we see that $T_{160+\text{SPIRE}}$ starts being colder than $T_{\text{PACSoonly}}$ for $T > 20 \text{ K}$ (for $T_{\text{PACSoonly}} > 28 \text{ K}$, the value of the mean is likely affected by the small number of points).

The blue contours in the right panel of Fig. 7 show the regions where $T_{\text{PACSoonly}} \geq 20 \text{ K}$. These regions are confined to the inner parts of IC348 and NGC 1333. Clearly, underestimation of T_d cannot explain the large difference where the observed emission at $70 \mu\text{m}$ is higher than that computed from the $160 + \text{SPIRE}$ column density and temperature maps. The VSG population in Perseus may be the cause of this excess. However, if so, it is clear that VSGs are detected primarily in the eastern field. We also note that the bubble CPS5 (Arce et al. 2011), for which Ridge et al. (2006b) derived $T \sim 29 \text{ K}$ with IRAS 60 and $100 \mu\text{m}$ data, is not visible in the right panel of Fig. 7, suggesting that its $70 \mu\text{m}$ emission seen by *Herschel* is, rather, compatible with a modified blackbody at lower temperatures, that is, those less than $\sim 20 \text{ K}$. HD 278942, thought to be the driving source of the bubble, has a distance of $\sim 520 \text{ pc}$ (Gaia Collaboration 2016), excluding that this star can indeed be related to the shell.

The fact that T can be underestimated when $T > 20 \text{ K}$ could in principle have an impact on the core temperature derived from SED fits when a source is not visible in the PACS blue band. However, for the starless cores, almost all of their temperatures are $< 20 \text{ K}$ with only a few exceptions. Sources with temperatures $> 20 \text{ K}$ are detected at $70 \mu\text{m}$, meaning that the observed SED covers the region of the peak intensity.

3.2.2. Masses and temperatures of the Perseus subclouds

From the intensity maps at the different wavelengths or the column density map, it is not clear how to define the borders of individual subregions of Perseus. The contour at $3 \times 10^{21} \text{ cm}^{-2}$ shown in Fig. 4 provides a good first-order estimate of the dense material, but it does not separate the individual subregions well. Further, our choice of $3 \times 10^{21} \text{ cm}^{-2}$ does not have a physical meaning and it may be possible to adopt another level of $N(\text{H}_2)$ and measure a different mass.

To find a solution for the border definition, we started drawing a polygon enclosing each $3 \times 10^{21} \text{ cm}^{-2}$ contour. Indeed, a polygon is much easier to handle in a computer code given the fact that it is defined with much fewer vertices. We therefore used such polygons to obtain a first guess for the borders. Then we visualised the ^{13}CO 1–0 (Ridge et al. 2006a) spectrum inside each area, the perimeters of which were varied manually to find locations with one velocity component. However, this approach was only possible in a few cases because most of regions have multiple CO components. In the end, we therefore defined the different subclouds in an attempt to obtain one bright line with few contaminants. The areas thus defined are shown in Fig. 9 while in Figs. C.1 and C.2 they are overplotted onto the ^{13}CO intensity maps at different velocity components to make it clear why a certain region was defined with that border. We introduced

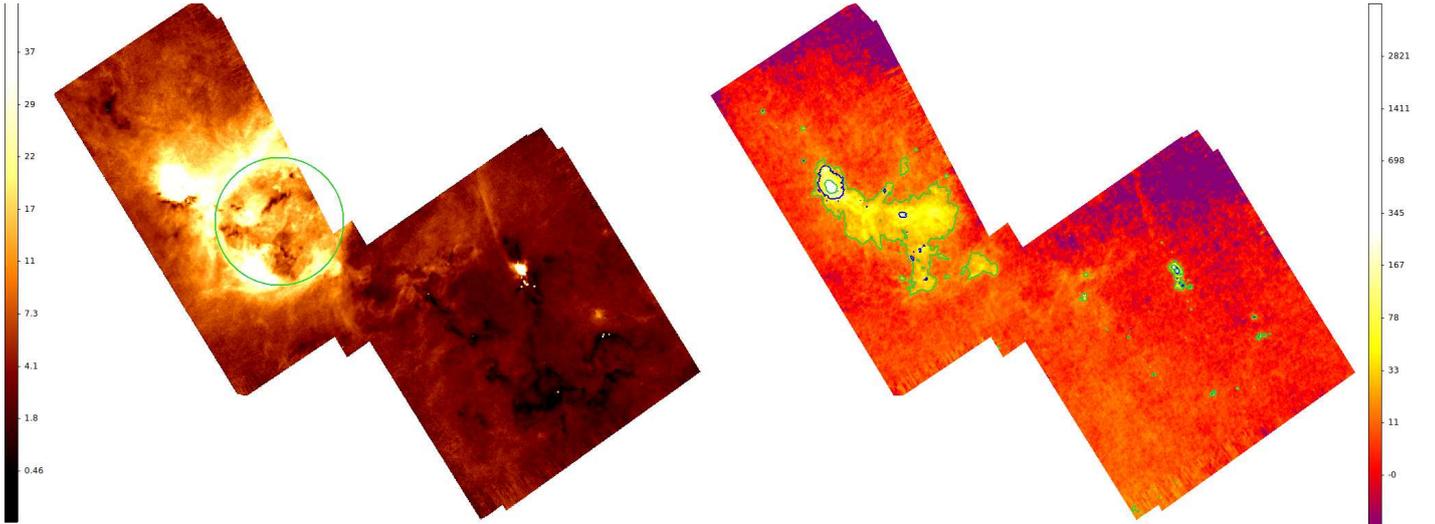


Fig. 7. *Left panel:* inferred 70 μm intensity map extrapolated from the column density and temperature maps obtained for $\lambda \geq 160 \mu\text{m}$. The green circle shows the bubble found by [Ridge et al. \(2006b\)](#) (CPS5 in the list by [Arce et al. 2011](#)). *Right panel:* difference between the observed 70 μm map and the extrapolated map. Green contours are 20 and 200 MJy sr^{-1} ; blue contours show the region where $T_{\text{PACSonly}} \geq 20 \text{ K}$ (see text). Colour bars are in MJy sr^{-1} .

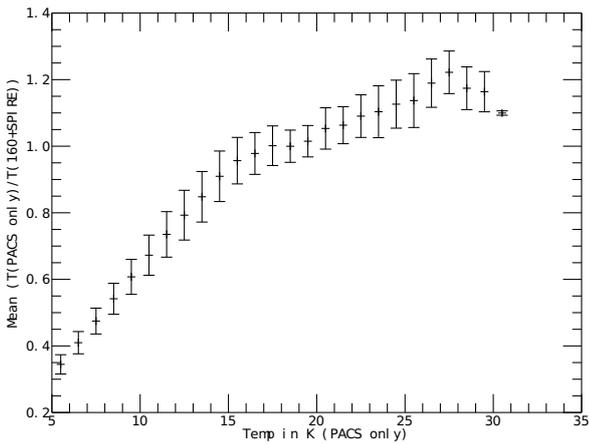


Fig. 8. Change in dust temperature when only PACS data are used (70, 100 and 160 μm) instead of PACS 160 μm plus SPIRE bands. Error bars are one standard deviation. The point at $T \sim 30 \text{ K}$ is not reliable because only a few values are available to compute the mean.

additional zones not associated with already known subregions, naming them HPZ#, that is, *Herschel* Perseus Zone number #. In Appendix C we also give the coordinates of the corners for all regions in DS9 format.

In Table 5 we report the main physical properties of each identified subregion. Columns labelled RA and Dec give, in degrees, the position of the peak in column density reported in the fourth column (for HPZ1, the peak is located on the far-east side of the subregion, and so we computed the geometrical centroid of the zone). In the next columns we give mass, area, and median temperature for the whole region and within the denser part where $N(\text{H}_2) > 3 \times 10^{21} \text{ cm}^{-2}$. The last two columns show the effect of assigning a different distance to each point of the column density map, following Eq. (1). Specifically, $M_{d(\alpha,\delta)}$ shows how the total mass of the cloud is affected, and column “ \bar{d} ” gives the distance at which $M_{d(\alpha,\delta)} = M_{\bar{d}}$, where $\bar{d} = \sqrt{\sum N_i(\text{H}_2) d^2(\alpha_i, \delta_i) / \sum N_i(\text{H}_2)}$, with i running over all the

pixels of each subregion. The total mass, area, median temperature, and the mean distance for the entire cloud within specific column density ranges are given in Table 6. From the values of \bar{d} in this table we conclude that our choice of 300 pc as representative distance is reasonable.

The mass for $N(\text{H}_2) > 1 \times 10^{22} \text{ cm}^{-2}$ is $954 M_{\odot}$ at 300 pc. With the same set of data, [Sadavoy et al. \(2014\)](#) found $1171 M_{\odot}$ when adopting 235 pc, which translates into $1908 M_{\odot}$ at our distance of 300 pc, a factor of two higher than our value. The main difference between the two analyses is that we adopted a newer version of the calibration files with slightly different SPIRE beamsizes. As the mass enclosed within a certain contour scales with the area defined by that contour, small variations in column density can cause large variations in mass. In particular, scaling the column density derived by [Sadavoy et al. \(2014\)](#) by 30%, the area decreases from 600 arcmin^2 to 303 arcmin^2 and the mass reduces to $944 M_{\odot}$, in agreement with the values we derive.

We also compare the *Herschel*-derived column density maps with the near-infrared extinction maps based on star counts (see, e.g. [Cambr esy 1999](#); [Schneider et al. 2011](#)). However, such a comparison is not immediate. [Benedettini et al. \(2015\)](#) found large discrepancies between the masses derived in Lupus with these two methods, which they attributed to uncertainties on the opacity law used to compute the *Herschel* column density map. The situation is more difficult for Perseus because of the known variability of $R \equiv A_V/E(B-V)$. [Foster et al. \(2013\)](#) found a strong correlation between R and A_V in Perseus, with the former increasing from ~ 3 to ~ 5 when the latter goes from 2 to 10 mag. Without a priori knowledge of the relation $R = R(A_V)$, it is not possible to translate $N(\text{H}_2)$ to A_V . Moreover, A_V maps are also subject to a zero-level uncertainty, in the sense that stars must be counted relative to a fiducial zone where $A_V = 0 \text{ mag}$. As Perseus is a large cloud, it is difficult to find such a clean region.

To investigate further, we derived the mass enclosed within a certain contour of A_V using both our column density map and a 2MASS-based extinction map ([Cambr esy 2015, priv. comm.](#)), taking into account the different spatial resolution and pixel size. We use the nominal conversion

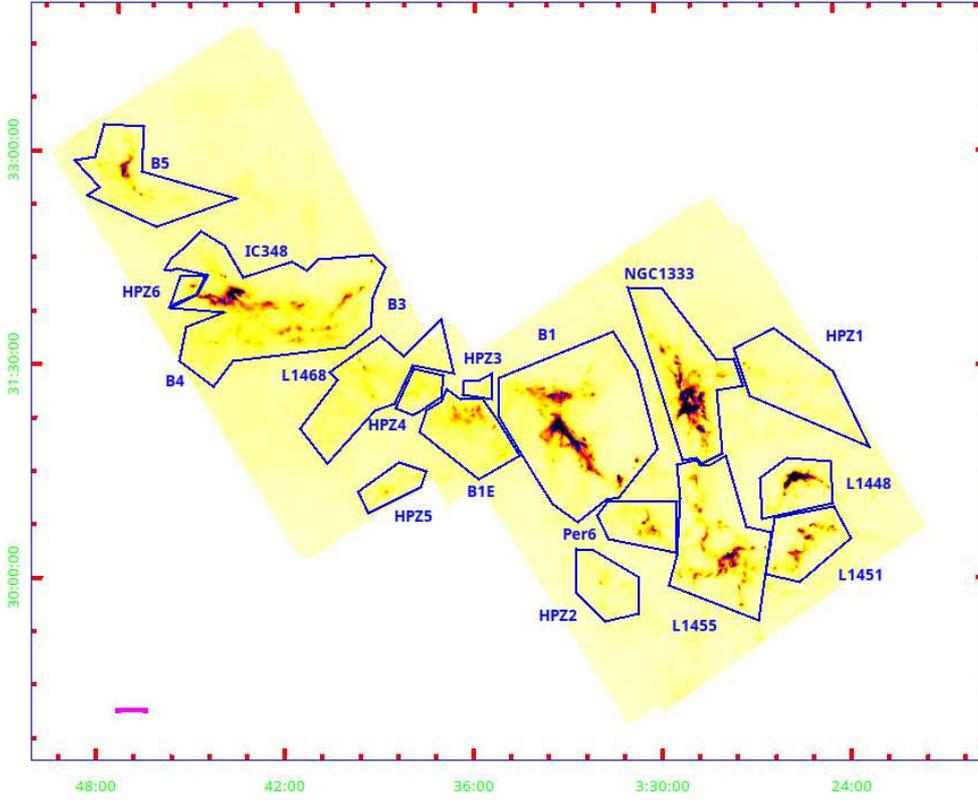


Fig. 9. Regions identified in Perseus overplotted on the column density map. The magenta line in the bottom left corner shows the angular scale corresponding to 1 pc at 300 pc. The J2000.0 coordinates grid is also shown. Coordinates of the regions are given in Appendix C.

Table 5. Properties of the sub-regions of the Perseus molecular cloud.

Name	RA J2000	Dec J2000	Peak (10^{21} cm^{-2})	M (M_{\odot})	Area (arcmin 2)	Med(T) (K)	$M_{d(\alpha,\delta)}$ (M_{\odot})	\bar{d} (pc)	Pro	Pre	UB			
L1451	51.32	30.32	16.9	276	125	746	140	15.1	13.3	246	283	0	19	8
L1448	51.40	30.76	106	313	194	610	122	15.2	13.2	285	285	4	18	7
L1455	51.92	30.20	37.0	740	375	1897	420	15.0	13.3	677	286	8	36	17
NGC1333	52.29	31.23	138	1060	727	1661	538	15.2	13.8	1021	294	44	23	23
Perseus6	52.64	30.44	19.1	203	74	579	93	15.5	13.6	191	290	5	4	2
B1	53.34	31.13	112	1443	780	3444	657	15.8	14.2	1423	297	15	31	38
B1E	53.98	31.24	9.65	389	157	887	215	15.6	14.7	395	301	0	8	9
L1468	54.91	31.53	7.87	455	49	1391	79	17.4	15.7	479	307	0	4	18
IC348 ^(a)	55.98	32.01	57.6	1568	881	3258	1082	17.2	16.5	1730	314	26	41	51
B5	56.62	32.49	23.9	410	106	1225	122	16.4	14.6	484	325	3	6	6
HPZ1 ^(b)	51.37	31.40	7.98	275	8	1282	13	15.9	14.3	257	289	0	1	9
HPZ2	52.98	30.01	6.73	116	3	575	5	16.1	13.9	108	289	0	1	5
HPZ3	53.89	31.40	3.88	24	1	92	2	16.3	14.4	24	302	0	0	1
HPZ4	54.54	31.34	4.70	93	26	243	44	17.5	17.1	96	305	1	0	1
HPZ5	54.72	30.66	6.33	90	12	315	18	15.9	14.3	92	302	0	1	0
HPZ6	56.32	32.08	11.4	40	17	101	23	16.8	15.9	46	318	2	2	2

Notes. Coordinates give the position of the peak in column density reported in the “Peak” column. The next six columns give the mass, area, and median temperature, with two values for each parameter. The first value refers to the entire subregion within the boundaries shown in Fig. 9, and the second value applies only for the data with $N(\text{H}_2) > 3 \times 10^{21} \text{ cm}^{-2}$. The column labelled $M_{d(\alpha,\delta)}$ gives the mass of each subregion using a modified distance according to Eq. (1), with \bar{d} being the mean distance for the entire region. The last three columns are the number of cores detected: protostars, prestellar cores, and unbound cores. ^(a)Includes regions B3 and B4; ^(b)coordinates of the geometrical centre.

$N(\text{H}_2) = 9.4 \times 10^{20} A_V \text{ mag}^{-1} \text{ cm}^{-2}$ from Bohlin et al. (1978), which assumes fully ionised hydrogen and $R = 3.1$. The mass found for $A_V > 3$ mag is $3620 M_{\odot}$ in our map and $9980 M_{\odot}$ in the extinction map, a factor 2.76 higher. However, for $A_V > 5$ mag, the ratio decreases to 2.35 ($1830 M_{\odot}$ and $4290 M_{\odot}$). The comparison improves if we derive the mass enclosed within the same area instead of the same A_V . The area is found projecting the

A_V contour derived on the extinction map onto our column density map. In this way, we measure in our map $5523 M_{\odot}$ and $2548 M_{\odot}$ for $A_V > 3$ mag and 5 mag, respectively, with a ratio mass(extinction map)/mass(column density map) of 1.81 and 1.68 for the two contours. A qualitatively similar trend is found in Orion B (Könyves et al. 2020). Overall, the masses derived from our *Herschel*-based column density map are accurate to

Table 6. Properties of the Perseus molecular cloud as a whole in three different regimes of column density.

$N(\text{H}_2)$ (10^{21} cm^{-2})	M (M_\odot)	Area (degrees ²)	Med(T) (K)	$M_{d(\alpha,\delta)}$ (M_\odot)	\bar{d} (pc)
≤ 3 ^(a)	8116	11.5	16.8	8300	303
> 3	3546	1.0	14.5	3560	300
> 10	954	303 ^(b)	12.7	937	296

Notes. The table lists: mass M , area, and median temperature $\text{Med}(T)$; the column labelled $M_{d(\alpha,\delta)}$ gives the mass as in Table 5 using Eq. (1) for the distance; \bar{d} is the mean distance. ^(a)Minimum value in our map is $2.51 \times 10^{20} \text{ cm}^{-2}$; ^(b)area is in arcmin².

within a factor of ~ 1.5 – 2 for $A_V > 3$ mag. At low column densities, $A_V < 3$ mag, we may overestimate the dust opacity and consequently underestimate the masses by a factor of ~ 2 – 3 (Benedettini et al. 2015).

To test the hypothesis that R can play a role in causing column density discrepancies, we varied R until good agreement was found between the extinction map and the thermal dust map. We chose a fiducial extinction level of $A_V = 7$ mag to make this comparison, so that any uncertainties in both maps from the zero-point corrections are negligible relative to the measured column of dust. We then searched iteratively for the value of R that equalises the masses found from the two maps. We find a best match for $R = 3.95$ which implies $N(\text{H}_2) = 7.34 \times 10^{20} A_V \text{ mag}^{-1} \text{ cm}^{-2}$. With this conversion, the cloud mass for $A_V > 7$ mag is roughly $1600 M_\odot$ from both maps. Assuming the nominal value $R = 3.1$, the thermal dust map gives a mass of $1213 M_\odot$ and the extinction map gives $2046 M_\odot$.

Another study of the extinction in the Perseus molecular cloud has been done by Zari et al. (2016) using the same set of *Herschel* images. These latter authors give the mass enclosed within a set of A_K contours based on the relation $A_K = 1.67 \times 10^{22} [2N(\text{H}_2) + N(\text{H})] \text{ mag cm}^{-2}$, which becomes $A_K = 8.35 \times 10^{21} N(\text{H}_2) \text{ mag cm}^{-2}$ assuming fully molecular hydrogen. To take into account the different mean molecular weight (1.37×2 instead of our value 2.8) and the different adopted distance (240 pc instead of 300 pc), we increased their masses by 1.53.

Finally, we had to consider the different dust opacity, which was more difficult. By combining Eq. (9) from Zari et al. (2016), namely the relation between A_K and $N(\text{H}_2)$, with Eq. (4) from the same paper, which relates A_K with the optical depth at $850 \mu\text{m}$, we derived

$$\kappa_{850} = \frac{A_K - \delta}{8.35 \times 10^{21} \mu_{\text{H}_2} m_{\text{H}} A_K}. \quad (4)$$

The fact that in the above equation the zero-point δ is not zero makes κ_{850} a function of A_K ; on the other hand, if $A_K \gg \|\delta\| = 0.05$ mag, then $\kappa_{850} \rightarrow 0.0065 \text{ cm}^2 \text{ g}^{-1}$. As we have assumed a dust opacity index $\beta = 2$, $\kappa_{300} = 0.052 \text{ cm}^2 \text{ g}^{-1}$. Also, because $I_\nu \propto \kappa_\nu \Sigma$, lowering κ_ν by a factor $0.052/.1 = 0.52$ implies increasing our values of $N(\text{H}_2)$ by the factor $1/0.52 = 1.92$.

We downloaded the τ_{850} map of Zari et al. (2016) from the CDS and converted it into A_K following their prescription. Our $N(\text{H}_2)$ was projected onto that of Zari et al. (2016) in order to take into account the different pixel size, and then we computed the mass enclosed within the A_K contours used by Zari et al. (2016). For $A_K > 0.2$ mag, the ratio between the mass M_Z derived by Zari et al. (2016) and the mass estimated from our column density map, M_P , is 1.72 instead of the expected value of 1.92. This

Table 7. Mass cloud derived with two dust emissivity indices β , with and without colour corrections applied, and with a different spatial sampling of the column density map.

	$\beta = 2$		$\beta = 1.7$		
	No CC	CC	No CC	CC	
	(Masses in $10^3 M_\odot$)				
(1)	(2)	(3)	(4)	(5)	(6)
$N(\text{H}_2) \leq 3 \times 10^{21}$	8.12	8.38	7.27	6.41	5.66
$N(\text{H}_2) > 3 \times 10^{21}$	3.55	3.28	2.59	1.73	1.29
$N(\text{H}_2) > 10^{22}$	0.95	0.59	0.66	0.72	0.37
$A_V > 7$ mag	1.54	1.14	1.14	0.74	0.56

Notes. The value $\beta = 2$ is the HGBS standard, $\beta = 1.7$ is the peak of the distribution of β in the Perseus region (Planck Collaboration XI 2014). Masses in Cols. (2), (3), and (5) were derived without applying colour corrections, while in Cols. (4) and (6) the reported masses were computed with colour corrections. In all cases but Col. (3), the column density map was created on a grid of size $14''$ and spatial resolution of $36''$, default values for HGBS. In Col. (3) the column density map was projected onto the grid used for *Planck* data: $90''$ pixels and $5'$ resolution. The 7 mag contour has been drawn assuming $R = 3.1$.

small difference of 10% in mass could be due to the different way in which the zero-level of the intensity maps is derived, combined with the fact that for 0.6 mag, the zero-point δ in Eq. (4) is not yet negligible. The ratio of the masses increases with increasing A_K until, for $A_K > 6.4$ mag, it reaches 1.93, as expected.

We finish this section by showing how the mass derived from the intensity maps depends on the spatial resolution of the column density map, on the dust opacity index β fixed to a value of 2, and on the fact that colour corrections are neglected. In Table 7, we report the masses enclosed within different $N(\text{H}_2)$ levels and for $A_V > 7$ mag converted into column density values using $R = 3.1$. The first set of columns show the masses from our data for $\beta = 2$ and the second set show our data for $\beta = 1.7$, which is the peak of the distribution of β in the Perseus region (Planck Collaboration XI 2014). For each β case, we measure masses before (No CC) and after applying a colour correction (see below for details). Finally, in the middle column for the $\beta = 2$ set, we also give the mass when the column densities are convolved to the Planck resolution ($5'$) with a $90''$ pixel size.

In general, smoothing the map decreases the column density in the densest parts and increases the column density in the more diffuse regions surrounding the denser material. For example, at $N(\text{H}_2) > 10^{22} \text{ cm}^{-2}$, we recover only 62% of the mass measured in the smoothed *Herschel* map at $5'$ resolution compared to the original map at $36''$ resolution, whereas for $N(\text{H}_2) < 3 \times 10^{21} \text{ cm}^{-2}$, we find a slightly higher mass in the $5'$ resolution map. With $\beta = 1.7$, the column density and cloud mass decreases substantially. We find values that are roughly 75–80% of what was obtained with $\beta = 2$.

Table 7 shows that the colour corrections (CC) decrease the estimated masses by 70–90%. Such corrections are in general necessary because the flux calibration of PACS and SPIRE was performed assuming that the SED of a source displays a flat νF_ν spectrum⁴. For all other kinds of SEDs, the derived fluxes must be colour-corrected according to the intrinsic source spectrum. To compute the column density with colour corrections applied, our fitting code integrates the synthetic SEDs over the PACS and

⁴ This is a common assumption when calibrating a photometer.

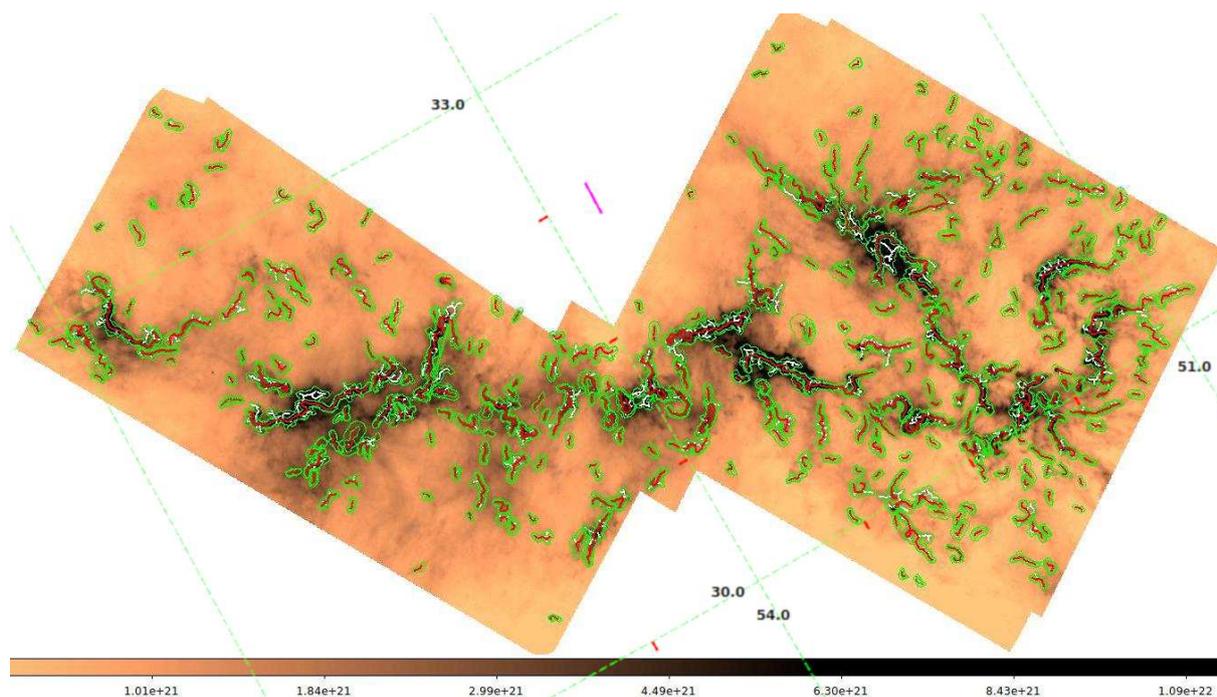


Fig. 10. Network of filaments overlaid on the column density map. Red lines show the spine of the filaments, white lines show the branches, and the green contours show the width of the filaments and branches. The map has been rotated by 28° . The magenta line in the centre shows the angular scale corresponding to 1 pc at 300 pc. The J2000.0 coordinates grid is shown.

SPIRE response filters during the generation of the grid. In this way, each model has its own CC built in. In the remainder of this paper, we use the maps at the spatial resolution of $500 \mu\text{m}$ derived as in the other HGBS works: no CC applied and β fixed to a value of 2.

3.2.3. Filamentary structure of Perseus

One of the main results of *Herschel* in the field of star formation is the discovery of the deep link between the filamentary structure of molecular clouds and the sites where stars form. In particular, star-forming cores are found preferentially in denser filaments (André et al. 2010; Molinari et al. 2010; Rayner et al. 2017).

Based on *Herschel* data, Polychroni et al. (2013) derived two distinct core mass functions in L1641, part of the Orion A complex, for sources inside and outside of filaments. The mass distribution of the sources on the filaments was found to peak at $4 M_\odot$ with a CMF at higher masses modelled with a power law $dN/d\log M \propto M^{-1.4}$. The mass distribution of the off-filament sources has a peak at $0.8 M_\odot$ and a flat CMF at masses lower than $\sim 4 M_\odot$.

Roy et al. (2015) found a possible link between the 1D power spectrum of filaments and the origin of the high-mass tail in the core mass function (very close to the Salpeter law $dN/dM \propto M^{-2.5}$).

A complete study of the filaments in Perseus is not within the scope of this paper; nevertheless, given the aforementioned results, it is important to give here at least a short summary of the main properties of the filaments, which are discussed in Sect. 5 where the relation with the core population is addressed.

We used the filament-detection algorithm⁵ of Schisano et al. (2020, 2014) to identify filamentary structures and their

properties across our column density map. Here we summarise how the algorithm works, adopting the nomenclature described in those papers.

By thresholding the minimum eigenvalue of the Hessian of the column density map, the code is able to find and encompass regions where there are maximum variations of the contrast, that is, the bright features on the map. Filamentary structures are picked among these features through selection criteria on the elongation and coverage of the regions of interest.

Once the regions of interest are identified, the *IDL*⁶ morphological operator “THIN” is applied to them. The result of the operator is the “skeleton” region. Because of the nested morphology of filamentary features, the skeleton in each region is composed by one or more “branches”. The spine of the filament is defined as the group of consecutive branches, which are connected to each other and trace the longest possible path over the filamentary region.

Schisano et al. (2014) found that the area of the filament is underestimated by a mere thresholding of the eigenvalue, which only traces where the emission is concave downwards. These latter authors also found that by enlarging the border by three pixels in both directions perpendicular to the spine gives a better estimate of the position where filaments merge with the background. We verified that such an approach is valid also for our case and we applied it to our filament sample.

The code was run on the $36''$ -resolution column density map. Among the identified features, we selected as filaments the elongated regions having a spine longer than 12 pixels, corresponding to 0.24 at 300 pc, or about four times the spatial resolution at $500 \mu\text{m}$.

The result of the extraction is shown in Fig. 10 where the filaments are overlapped on the column density map. Green lines show the filament borders, and red and white lines are the spines and the branches, respectively.

⁵ The code is freely available at the following URL <http://vialactea.iaps.inaf.it/vialactea/eng/tools.php>

⁶ *IDL* is registered trademark of Harris Corporation.

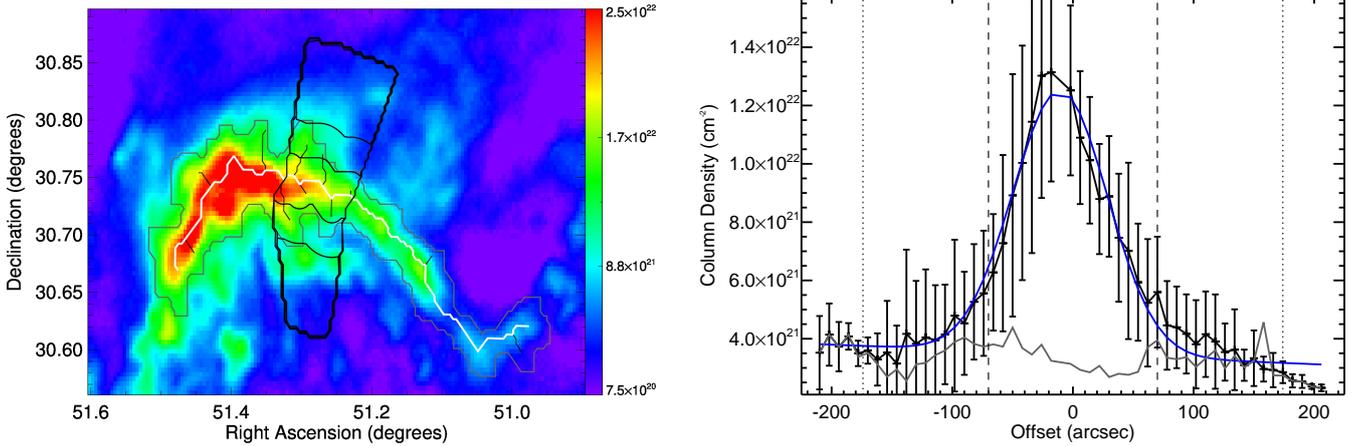


Fig. 11. *Left panel:* zoom-in of the filament in the region of L1448. The white line is the spine, while the grey line is the border of the filament. The short black lines are branches. To explain how the radial profile of the filament is derived, we defined, with the thick black line, an arbitrary region R . *Right panel:* radial profile of the filament averaged longitudinally within the black region shown in the left panel. The error bars are the standard deviation of the column density averaged within ten pixels. The blue line is the Gaussian fit, and the grey line the estimated background. The dashed and dotted vertical lines are explained in the text.

In the left panel of Fig. 11 we show the details of all the features previously defined: the filament border (thin grey line), the main spine (white line), and the system of branches (short black lines), for the case of the L1448 cloud. The identified filament runs along almost the entire cloud. The thick black line defines an arbitrary region R used in the right panel of Fig. 11 to show the radial profile of the filament, averaged along the filament itself, and to assess, at least in one case, the validity of the three-pixel border expansion.

The filament-detection code expands the filament border radially by three pixels to identify the position where the filament merges with the surroundings. The average radial profile inside R is shown in the right panel of Fig. 11. The error bars are the standard deviation of the column density for each bin of distance from the spine. Because of the irregular shape of expansion, which reflects the irregular filament profile, the border does not have a constant offset from the spine. The minimum and the maximum radial distance from the spine reached by the filament border over R , r_1 and r_2 , respectively, are shown with thin black lines, running parallel to the spine, in the left panel of Fig. 11. For distances $|r| \leq r_1$, within the two vertical dashed black lines in the right panel of Fig. 11, all the pixels of the map belong to the filament area, while for distances $|r| \geq r_2$, marked with the two dotted black lines, all the pixels are external to the filament which contains only background pixels. At intermediate distances $r_1 \leq |r| \leq r_2$, to be part of the filament or not becomes a local property, depending on the position along the spine.

The background is estimated from linear interpolation of the pixels that are outside the filament area, where $|r| \geq r_2$, or locally where $r_1 \leq |r| \leq r_2$. This estimated background is shown with a thick grey line in the right panel of Fig. 11. The blue line in the same figure is the Gaussian fit to the radial profile. The filament width is estimated as the FWHM of the fit. The Gaussian fit gives another estimate of the background, but we prefer to use the linear interpolation because the Gaussian fit is less suited to catching possible asymmetries, as is visible in our example. These asymmetries are stronger in the wings of the profile than in the inner part, meaning that it is reasonable to use the Gaussian FWHM to estimate the width.

In Fig. 12 we show the filament widths averaged over the spine and deconvolved by the FWHM at 500 μm , 36". The

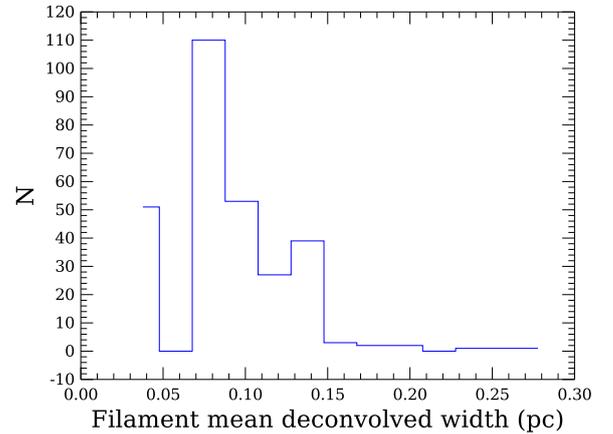


Fig. 12. Histogram of the filament widths averaged over the spine (bin size 1 pixel or 0.02 pc at $d = 300$ pc).

width is given in parsecs assuming a distance of 300 pc. The peak of the distribution is ~ 0.08 pc, consistent with the finding of [Arzoumanian et al. \(2019\)](#) of a characteristic width of 0.10 ± 0.03 pc for the filaments in eight HGBS regions.

An important physical parameter for filaments is the mass per unit length. Theoretical models of isothermal infinite cylindrical filaments confined by the external pressure of the ambient medium predict the existence of a maximum equilibrium value to the mass per unit length ([Ostriker 1964](#)):

$$M_{\text{line,max}} = \frac{2kT}{\mu_{\text{H}}G} = 16.4 \left(\frac{T}{10 \text{ K}} \right) M_{\odot} \text{ pc}^{-1}. \quad (5)$$

Above this value the filament is unstable and can fragment to form cores.

In our case, $M_{\text{line}} = M/L$ where the mass M of the filament is $\mu_{\text{H}} \sum_i n_i(\text{H}_2)A$ with A the area of one pixel in cm^2 , and the sum is over all the pixels in the filament area, while L is the spine length. We note that A depends on the distance as d^2 while L depends on d , and so M/L increases with d , and, for a constant $d = 300$ pc, is likely overestimated in the western half of Perseus, and underestimated in the eastern half.

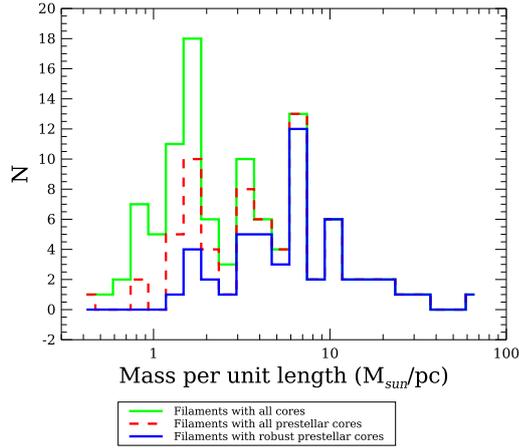


Fig. 13. Histograms of mass per unit length for filaments with cores (green line) and prestellar cores (blue and red lines, see text).

In Fig. 13 we show the distribution of the mass per unit length for all the filaments we find in Perseus that contain at least one core of any type (for core definition and extraction, see Sect. 4.1 below). The mass per unit length distribution peaks at $\sim 1.7 M_{\odot} \text{ pc}^{-1}$, well below the typical value of $16 M_{\odot} \text{ pc}^{-1}$ of Eq. (5). If we consider that 10 K is not representative of the dust temperature and that we should adopt $T \gtrsim 12$ K (see Fig. 5; see also Arzoumanian et al. 2019, where a range between 11 and 30 K is found for the dust temperature of filaments in many HGBS regions) the peak in the distribution is more than a factor of ten smaller than the maximum linear mass.

An important issue here is the background subtraction because, if the background is overestimated, the linear mass of the filament is underestimated. To address this problem we made use of the results found by Hacar et al. (2017). They derive 14 kinematically coherent structures in NGC 1333 through observations of the N_2H^+ (1–0) line. These structures, named fibres, are similar to what we call branches. However, the comparison between fibres and branches is not easy because of the different tracers used (molecular gas vs. dust emission), and because of the different algorithm used to derive the structures. Only in one case did we find a fibre and a branch that have a reasonable spatial overlap with a similar length. This happens for fibre number 9: Hacar et al. (2017) derived a length of 0.47 pc and $M_{\text{lin}} = 58.5 M_{\odot} \text{ pc}^{-1}$, while we find a length of 0.50 pc and $M_{\text{lin}} = 62.5 M_{\odot} \text{ pc}^{-1}$, and so, at least for this case, the background estimate appears to be reliable.

As only prestellar cores are likely star-forming in nature (see Sect. 4.3), it is perhaps more reasonable to consider the filaments that contain prestellar cores instead of the filaments that contain only unbound cores. In this case (blue histogram in Fig. 13), the peak of the distribution increases to $\sim 7 M_{\odot} \text{ pc}^{-1}$. This value is still less than the value $M_{\text{lin}}/M_{\text{lin,max}}$. Adding the candidate prestellar cores (see again Sect. 4.3) to enlarge the samples of bound cores changes the shape of the histogram (dashed red line), but only in the region of smaller masses. In any case, the majority of filaments are below $M_{\text{lin,max}}$. However, we note that Fischera & Martin (2012) found that fragmentation and core formation can occur in filaments when $M_{\text{lin}}/M_{\text{lin,max}} > 0.5$, that is, close to our value.

Similarly, Benedettini et al. (2018) found that the majority of filaments with bound cores in Lupus have $M_{\text{lin}} < M_{\text{lin,max}}$. These latter authors suggest that the mass per unit length of a filament should be considered locally instead of giving one value

for a whole filament. Indeed, if we compute M_{lin} for the single branches instead of considering the entire filament, we derive a much broader distribution toward high values, $\gtrsim 130 M_{\odot} \text{ pc}^{-1}$, of M_{lin} . Nonetheless, the median of the distribution is $\sim 1.5 M_{\odot} \text{ pc}^{-1}$ with a peak at less than $1 M_{\odot} \text{ pc}^{-1}$.

Nevertheless, we note that the formula given in Eq. (5) gives the maximum line mass for isothermal equilibrium when the non-dimensional radius ξ of a filament goes to infinity. However, in the more general case of finite ξ , Eq. (5) reads (Ostriker 1964)

$$M(\xi) = \frac{2kT}{\mu m_{\text{H}} G} \frac{1}{1 + 8/\xi^2} \leq 16.4 \left(\frac{T}{10 \text{ K}} \right) M_{\odot} \text{ pc}^{-1}, \quad (6)$$

where the equality holds only for $\xi \rightarrow \infty$. Therefore, in the more general case of finite ξ , the mass $M(\xi)$ of a self-gravitating cylinder in thermal equilibrium is smaller than the asymptotic $M_{\text{lin,max}}$ for the same temperature.

The non-dimensional radius ξ can be transformed back to a physical radius once we know T and ρ_0 , the temperature and central density of the filament, respectively. Alternatively, one can estimate ρ_0 from the observed values. For example, if we assume $T = 16$ K and $M_{\text{lin,max}} = 8 M_{\odot} \text{ pc}^{-1}$, Eq. (6) solved for ξ gives $\xi \sim 2.15$. Subsequently, if we assume 0.08 pc to be the typical filament radius, solving Eq. (45) in Ostriker (1964) for the central density yields $\rho_0 \sim 4.3 \times 10^{-20} \text{ g cm}^{-3}$ or $n_0 \sim 9 \times 10^3 \text{ cm}^{-3}$ for $\mu = 2.8$. Strictly speaking, 0.08 pc is the width of the filament, defined as the FWHM of the Gaussian fit. If instead we define the radius as $1.29 \times \text{FWHM}$ (Pezzuto et al. 2012), then the typical filament radius becomes ~ 0.10 pc and $n_0 \sim 5.5 \times 10^3 \text{ cm}^{-3}$. In any case, we conclude that as long as $n \gtrsim n_0$, a filament with radius 0.08 pc and $T = 16$ K has $M_{\text{lin,max}} = 8 M_{\odot} \text{ pc}^{-1}$, similar to the peak value we find in the histogram of Fig. 13.

3.2.4. The probability density function of the column density map

The probability density function (PDF) derived from the column density map has previously been used to probe the physics governing the diffuse medium (see Schneider et al. 2015, for a review on PDF). Briefly, the PDF shows a log-normal behaviour at low densities due to the turbulence in the cloud, while a power-law tail develops at higher densities as a consequence of self-gravity and star-formation activity.

In Fig. 14, we present the PDF for the whole molecular cloud. We model the observed distribution with log-normal and power-law functions, fitting each curve independently. The brown line shows the best log-normal fit to the low-column density part of the distribution. The peak is at $N(\text{H}_2) = (9.738 \pm 0.061) \times 10^{20} \text{ cm}^{-2}$. However, the significance of this fit is limited by the fact that the smallest value for column density with a closed contour in our map is $N(\text{H}_2) \sim 1.6 \times 10^{21} \text{ cm}^{-2}$. As a consequence, the shape of the PDF in the region where the log-normal fit is derived can be distorted by incompleteness in the data (Alves et al. 2017 argued that the overall log-normal shape is created rather than distorted by incompleteness).

Moving now to the high-column-density tail in the PDF, Federrath & Klessen (2013) showed that the following relation exists between s , the power-law slope in the PDF, and q , the power-law slope $\rho(r) \propto r^{-q}$ of the volume density for cores:

$$q = 1 - \frac{2}{s}, \quad (7)$$

meaning that if s is derived, information can be obtained on how the star-formation process in a cloud is proceeding.

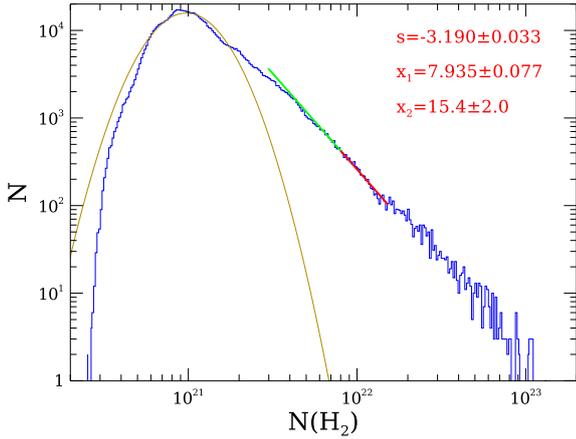


Fig. 14. Column density PDF for the whole Perseus molecular cloud: s is the slope of the power-law fit, shown in red, over the interval x_1 – x_2 , both in 10^{21} cm^{-2} . As shown with the green line, the fit can be extended down to $\sim 4 \times 10^{21} \text{ cm}^{-2}$. The brown line is a log-normal fit to the low-density PDF whose parameters are given in the text. Note that $N(\text{H}_2) \sim 1.6 \times 10^{21} \text{ cm}^{-2}$ is the smallest column density having a closed contour in our map.

To find the exponent of a power law, the easiest and most often used solution is to linearise the problem in log-log space, obtaining an equation of the kind $\log y = \log c - \gamma \log x$. The unknowns c and γ are immediately found by fitting a straight line to the dataset. Nevertheless, this procedure should be avoided for many reasons (Bauke 2007; Clauset et al. 2009); for instance, it violates the assumption that uncertainties on the dependent variable follow a Gaussian distribution. This assumption is the basis for a least-square fit.

Our strategy instead was to directly fit a function $y = cx^\gamma$ using a non-linear fitting routine. We used the linearisation scheme described above to obtain only a first estimate of the parameters. As we do not know the extent of the interval over which the fit should be applied, we treated the interval extremes, $N(\text{H}_2)_{\min}$ and $N(\text{H}_2)_{\max}$, as free parameters.

In this way, it is not possible to directly compare the χ^2 corresponding to each model, because the number of data points is not constant. Furthermore, the use of the reduced χ^2 with non-linear models is questionable because the degrees of freedom in this case are generally not known⁷. To derive the best model we adopted the cross-validation method (Andrae et al. 2010). Namely, for a given $N(\text{H}_2)_{\min}$ and $N(\text{H}_2)_{\max}$, we computed the best c, γ excluding one point of the dataset, and from the derived best-fit model we computed the difference between the expected value and the model-derived value. This procedure is repeated for all the points in the dataset and the product of all the differences gives an estimate of the likelihood for that model. We then looked for the maximum likelihood value among all the models ($N(\text{H}_2)_{\min}, N(\text{H}_2)_{\max}, c, \gamma$).

The validity of this strategy is limited by the fact that the PDF often does not show a clear power-law trend. As a consequence, the fit procedure tends to minimise the interval $N(\text{H}_2)_{\min} - N(\text{H}_2)_{\max}$. To impose physical constraints on the fit, we varied $N(\text{H}_2)_{\min}$ around the value reported in column “ $N(\text{H}_2)$ ” of Table 10, which is discussed below. This value in

column density is the smallest background column density⁸ found for prestellar cores in each sub-cloud. As the power-law tail should trace the region where gravity is strong enough for cores to form, it seems reasonable to impose that $N(\text{H}_2)_{\min}$ is not much different from the minimum background column density.

To derive an uncertainty for $N(\text{H}_2)_{\min}$ and $N(\text{H}_2)_{\max}$, dubbed x_1 and x_2 in Fig. 14, we created the PDF histogram with five different bin sizes from 0.8 to 1.2, in steps of 0.1, in units of 10^{20} cm^{-2} (the one shown in the figure corresponds to the choice $1 \times 10^{20} \text{ cm}^{-2}$). The power-law slopes and limits were found for the five histograms and similar solutions were sought. The red line in Fig. 14 shows the power-law whose slope is the weighted average of the five slopes and is measured over an interval of 7.935×10^{21} and $1.54 \times 10^{22} \text{ cm}^{-2}$. These intervals are the mean of the five starting and end points over which the fit extends.

For the whole Perseus cloud, the minimum background column density for prestellar cores is $9.4 \times 10^{20} \text{ cm}^{-2}$ (see Table 10). We were not able to find any good power-law fit starting from such small values. Indeed, the best fit we found (slope -3.190) is limited to $7.9 \times 10^{21} \text{ cm}^{-2}$ but it is nonetheless a good solution down to $\sim 4 \times 10^{21} \text{ cm}^{-2}$, as can be seen from the green line that extends the best fit to smaller column densities.

Our slope -3.190 ± 0.033 is in good agreement with the value -3 found by Zari et al. (2016), and translates into $q = 1.6269 \pm 0.0065$, similar to the 1.5 slope expected for spherically symmetric free-fall collapse. According to Schneider et al. (2015), this slope should be relatively insensitive to the spatial resolution and to the histogram bin size.

Sadavoy et al. (2014) found a correlation between the slope of the PDFs and the star formation efficiency. We discuss this topic below. Here, we show the results for the individual subregions defined in Fig. 9. Figure 15 shows the PDFs for each of the subregions that hosts at least one protostar. In many cases, the low-density portion cannot be fit with a log-normal function, while the high-density tail shows a wide range of values (labels s in Fig. 15). The slopes s for the single clouds are different from the one found for the whole Perseus region, as predicted by Chen et al. (2018). Through Eq. (7), the interval in slopes translates into an interval for q , limited by 1.5 in IC348, and by 2.3 in HPZ6 (however, for this region the result of the fit is not meaningful).

It is not possible to directly compare our slopes with those found by Sadavoy et al. (2014), given the different strategy adopted. To make the comparison, we derived the slope of the power law following the method of these latter authors: first, we fixed $N(\text{H}_2)_{\min}$ to $7 \times 10^{21} \text{ cm}^{-2}$, which is 30% (see Sect. 3.2.2) of $1 \times 10^{22} \text{ cm}^{-2}$, the value adopted by Sadavoy et al. Second, we made a linear fit in the log-log plane. The slopes are compared in Table 8.

In this table we report both the slopes found with our data according to the procedure written above (column P) and the slopes found with our data but following the same procedure as in Sadavoy et al. (2014) (column P_S). Naming the values reported in columns S and P_S , $s \pm \sigma_s$ and $p \pm \sigma_p$, respectively, we see that $s > p$ always. In three cases $s - \sigma_s < p < s$ and in another three case the intervals $s - \sigma_s$ and $p + \sigma_p$ overlap. Only for L1448 do we find a difference that is slightly larger than $1\sigma_s$, namely $p = s - 1.4\sigma_s$. We conclude that the slopes reported in Fig. 15 differ from those reported in Sadavoy et al. (2014) because of the

⁷ The usual assumption of $N - n$ degrees of freedom if N is the number of data and n is the number of parameters may be valid, but in general is not.

⁸ The background column density is computed by *getsources* at the position of each core in the high-resolution column density map, and reported in the catalogue, see Table A.1.

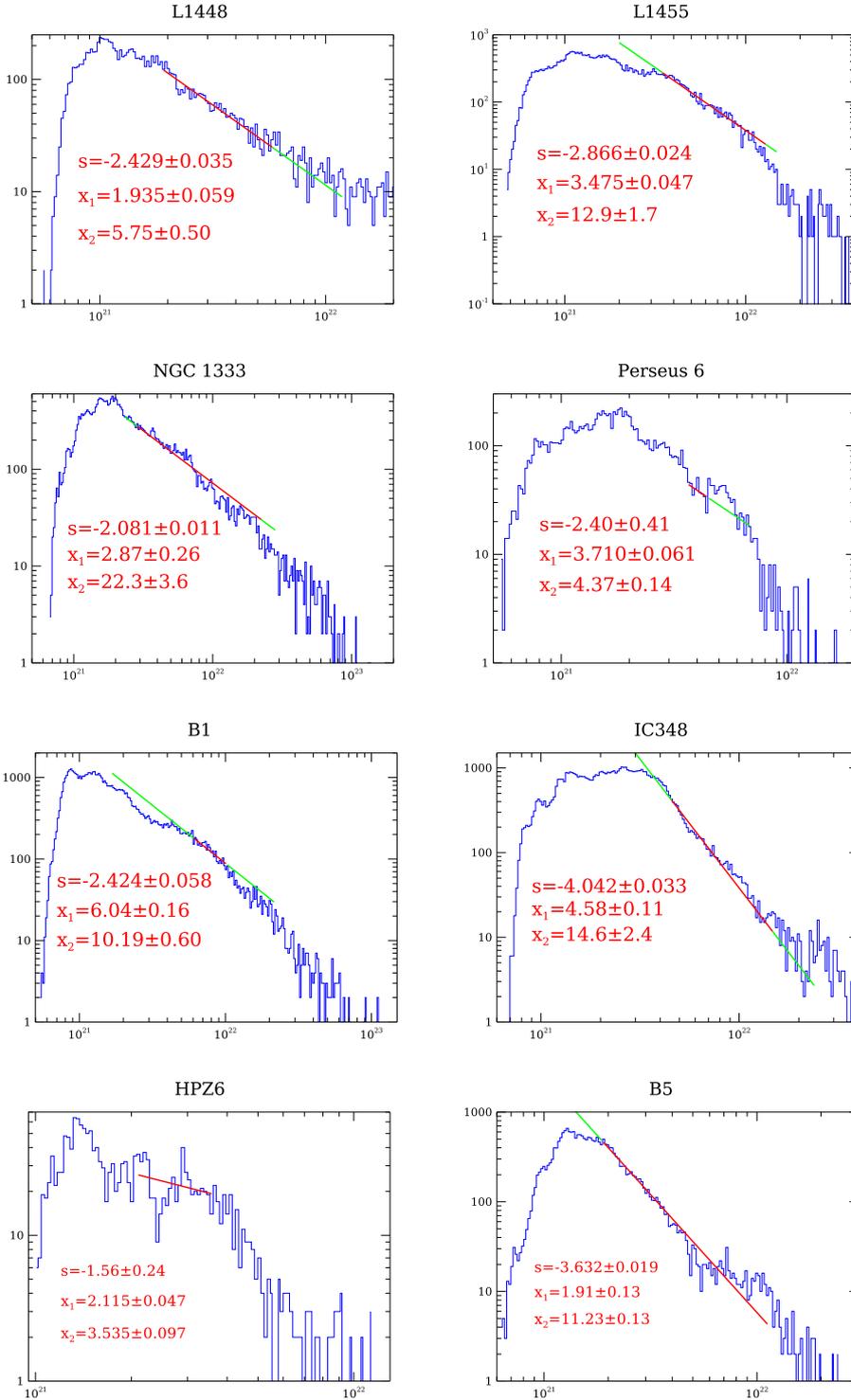


Fig. 15. Probability density functions for the subregions identified in Perseus where protostars were detected. In each panel, s gives the slope of the power-law fit, shown in red, in the interval delimited by x_1 and x_2 , both in 10^{21} cm^{-2} ; the fitting procedure is detailed in the text. The x -axis is $N(\text{H}_2)$ in cm^{-2} . Green lines extend the best fit over the interval of column densities where protostars were found (see Col. $N(\text{H}_2)$ in Table 9).

different strategy adopted to fit the power-law tail in the PDF. When used in the rest of this paper, the slopes of the PDFs are those reported in Fig. 15.

4. Physical properties of the cores

4.1. Source extraction

Sources were extracted from the intensity images using release 1.140127 of *getsources* (Men'shchikov et al. 2012;

Men'shchikov 2013), a multi-scale, multi-wavelength source-extraction algorithm. Its application to the fields of the HGBS is described in Könyves et al. (2015). We followed their strategy to detect starless cores and protostars by running *getsources* twice with two separate sets of parameters to optimise the extraction. In the following we summarise the extraction procedure.

For starless cores, we extracted sources on the SPIRE intensity maps and the high-spatial-resolution column density map (Palmeirim et al. 2013, see also Sect. 3.2). We also use the $160 \mu\text{m}$ data via a temperature-corrected map, which is obtained

Table 8. Comparison of the PDF power-law slopes found by Sadavoy et al. (2014) (column S) and in this paper (column P).

Region	S	P_S	P
L1448	-1.4 ± 0.3	-0.982 ± 0.053	-2.429 ± 0.035
L1455	-2.9 ± 0.2	-2.67 ± 0.11	-2.866 ± 0.024
NGC1333	-1.8 ± 0.2	-1.643 ± 0.041	-2.081 ± 0.011
B1	-2.2 ± 0.2	-1.983 ± 0.052	-2.424 ± 0.058
B1E	-9.6 ± 1.0	-7.8 ± 1.2	-4.26 ± 0.30
IC348	-2.5 ± 0.2	-2.308 ± 0.082	-4.042 ± 0.033
B5	-2.8 ± 0.7	-2.34 ± 0.14	-3.632 ± 0.019

Notes. Column P_S reports the slope found with our data following the Sadavoy et al. (2014) procedure.

by convolving the PACS red image to the resolution of the SPIRE 250 μm band and then combining them to derive a two-colour column density map (Könyves et al. 2015). The purpose of this approach is to avoid strong gradients of intensity that can be found in proximity to photon-dominated regions or hot sources. For protostars, only the 70 μm band was used for detection.

After source detection is completed, flux intensity, or simply flux, is measured for each source by *getsources* at all *Herschel* wavelengths. In particular, at 160 μm we use the intensity map, and not the colour-corrected map which is only used for detection. To both catalogues, we applied the selection criteria used for Aquila (Könyves et al. 2015). Starless cores were selected using the following criteria:

- column density detection significance greater than 5;
- global detection significance over all wavelengths greater than 10;
- global “goodness” ≥ 1 ;
- column density measurement with S/N greater than 1 in the high-resolution column density map;
- monochromatic detection significance greater than 5 in at least two bands between 160 μm and 500 μm ;
- flux measurement with $S/N > 1$ in at least one band between 160 μm and 500 μm for which the monochromatic detection significance is simultaneously greater than 5.

Protostars were selected using the following criteria:

- monochromatic detection significance greater than 5 in the 70 μm band;
- positive peak and integrated flux densities at 70 μm ;
- global “goodness” ≥ 1 ;
- flux measurement with $S/N > 1.5$ in the 70 μm band;
- FWHM source size at 70 μm smaller than 1.5 times the 70 μm beam size (i.e., $< 12''.6$);
- estimated source elongation < 1.30 at 70 μm .

The two catalogues were cross-matched to look for associations within $6''$. We found 60 sources present in both catalogues. These sources were removed from the core analysis, as they are considered as candidate protostars. In Appendix A we explain in more detail how the SEDs were built for the objects detected in both catalogues. Sources with no counterpart at 70 μm were classified as tentative starless cores. We then visually inspected the intensity maps and column density map toward each candidate to be sure that the source is not an artefact. This step removes non-existent sources due to local fluctuations in the maps and also finds candidate protostars present only in the starless cores catalogue (e.g. protostars that are not well detected at 70 μm). The reason for this depends on the slightly different criteria used

to build the catalogues (see Appendix A for details). To remove known galaxies we queried the NED⁹ and SIMBAD¹⁰ archives.

Starless cores not visible at 70 μm but with a SIMBAD counterpart at shorter wavelengths were removed from the catalogue because starless cores are not expected to be detected $\lesssim 160 \mu\text{m}$. We use a distance criterion of $6''$. As the association with a SIMBAD source may be only spatial and not physical when a starless core with no detection at 70 μm also has a submm counterpart within $6''$ beside a counterpart at short wavelengths, the source was left in our catalogue.

We also scanned the WISE¹¹ (Wright et al. 2010) archive within a radius of $6''$ from our objects, requiring a S/N greater than 0 in bands 3 and 4. This approach is relatively conservative because the threshold of S/N is very low. Moreover, there remains the possibility that two different sources appear close to each other in projection. In any case, sources with a WISE counterpart were removed. All the starless cores removed from our catalogue are reported in Appendix E.

In the end, we were left with 816 candidates: 684 starless cores and 132 protostars. A search for counterparts within $6''$ in the *c2d* archive (Evans et al. 2009)¹² shows no associations for the starless cores, while about 70% of our protostars have a possible counterpart. The catalogue, presented in Appendix A and available at the CDS, reports data for all 816 sources, both starless cores and protostars. In the following we discuss the physical properties of the starless cores while the protostars will be the subject of a forthcoming paper.

As an additional check on the reliability of the detected sources, we made a second extraction with a completely different source-identification algorithm, namely CuTEX (Molinari et al. 2011). This method identifies sources via second-order differentiation of the signal image in four different directions, looking for minima of the curvature. Since CuTEX makes detection and measurements on single band images, it is not straightforward to make comparisons with extractions by *getsources*, which builds its catalogue by combining the results of monochromatic detections. Given that the peak of the SED of cold cores is expected around 250 μm , we looked for agreement between *getsources* and PSW CuTEX catalogues. Indeed, we find that 467 (68%) of *getsources* cores have a CuTEX-detected object located within the 250 μm elliptical extent measured by *getsources*.

We find excellent agreement (98%) between the protostar catalogues derived with *getsources* and CuTEX. As both algorithms use only the 70 μm data to extract the candidate protostars, where diffuse emission is less intense, it is not surprising that the two methods find similar objects.

4.2. SED fitting

Once the source catalogue had been produced, we fitted the measured SEDs using a grid of theoretical models. This approach, already used in Benedettini et al. (2018) and equivalent to that used in Sect. 3.2, consists of generating a number of modified blackbodies

$$F_\nu = \frac{\kappa_\nu M B_\nu(T)}{d^2}; \quad \kappa_\nu = \kappa_0 \left(\frac{\lambda_0}{\lambda} \right)^\beta, \quad (8)$$

⁹ <http://ned.ipac.caltech.edu/>

¹⁰ <http://simbad.u-strasbg.fr/simbad/>

¹¹ <https://irsa.ipac.caltech.edu/cgi-bin/Gator/nph-scan?mission=irsa&submit=Select&projshort=WISE>

¹² <http://vizier.u-strasbg.fr/viz-bin/VizieR?-source=J/JApJS/181/321>

where M is the mass of the isothermal emitting dusty envelope at temperature T and distance d . For the dust opacity, we used the same parametrisation as for the column density map: $\kappa_0 = 0.1 \text{ cm}^2 \text{ g}^{-1}$ at $\lambda_0 = 300 \mu\text{m}$ with $\beta = 2$.

The grid of models was constructed fixing $M = 1 M_\odot$ and $d = 1 \text{ pc}$, while T was varied in the range 5–50 K in steps of 0.01 K (4501 models). The distance is scaled to the required value, namely of 300 pc for Perseus, during the fitting procedure.

Using a grid instead of a non-linear fitting procedure confers two advantages. First, it is not necessary to give initial guesses of the parameters, which are more or less arbitrary. Second, as long as the grid is large and dense, we can be sure to find the global minimum and not a local one.

As Eq. (8) is linear in M , the mass can be found with straightforward application of the least squares method when comparing a given SED to the grid of models, as shown in Eq. (2). Among the 4501 models, the best fit corresponds to the pair (T, M) with the minimum χ^2 . Uncertainty on T is derived considering all the models that have $\chi^2 \leq \chi_{\text{min}}^2 + 1$ (Andrae 2010). Then, if T_1 and T_2 are minimum and maximum temperatures of these models we give $T_{-(T-T_1)}^{+(T_2-T)}$.

Uncertainty in M was derived both according to the least squares method, and from the variation in mass corresponding to T_1 and T_2 . Clearly, if $T_1 < T_2$ then $M_2 < M_1$. The larger of the two uncertainties was then adopted.

The fitting procedure was not applied for SEDs with less than three significant flux measurements (see Appendix A). We also did not look for the best fit in cases where $F_\nu(350 \mu\text{m}) < F_\nu(500 \mu\text{m})$. For all these cases, we adopted a fiducial temperature equal to $10.4 \pm 1.0 \text{ K}$, the median temperature of all the reliable SEDs, excluding the protostars. The uncertainty on this median is arbitrarily large (the median $\Delta T/T$ for reliable sources is $\sim 3\%$). For these sources, the mass was computed from the measured integrated flux density at the longest significant wavelength, assuming $T = 10.4 \text{ K}$. The number of sources with reliable SEDs is 519 cores and 85 protostars. However, for the latter Eq. (8) is not an appropriate model, even if it can provide an estimate of the envelope properties. As noted earlier, discussion of the protostars catalogue is beyond the scope of this paper.

4.3. Core classification and stability

As explained in Sect. 4.1, we classify a core as protostellar if it is in emission at $70 \mu\text{m}$. This approach is physically plausible because in general cold cores are not visible at this wavelength, often not even at $160 \mu\text{m}$. On the contrary, once the source is visible at $70 \mu\text{m}$, a star has already formed, warming up its surrounding envelope. However, this rule does not always apply. In principle, all cores emit in the PACS bands, and their detection at $70 \mu\text{m}$ depends on its physical properties combined with the sensitivity of the instrument. Indeed, we find seven objects (ID #48, 692, 745, 754, 756, 761, 769 in the catalogue) with $70 \mu\text{m}$ emission that is either uncertain or at a level compatible with or below the best-fit SED found with Eq. (8). These sources were classified post facto as starless cores.

First hydrostatic cores (FHSC) are expected to have excess $70 \mu\text{m}$ with respect to a modified blackbody, even if they cannot be considered genuine Class 0 objects. They have masses comparable to those of giant planets and radii of only a few astronomical units (Larson 1969). Source #500 is indeed such a candidate (B1-bS, Pezzuto et al. 2012). However, in our catalogue we classified this source as a protostar because its properties are not compatible with those of the starless cores.

Finally, source #777 was first classified as starless because its $70 \mu\text{m}$ emission was an upper limit. From visual inspection this source is visible at $70 \mu\text{m}$, but as the uncertainty on the flux is higher than the flux itself, the code used to make the SED fitting classified it as starless, putting an upper limit on the blue band. We re-define this source as a protostar.

Among the starless sources, we want to identify cores that are not stable against their own gravity. As a result, they are likely to be collapsing or on the verge of collapse, forming one or more stars in the near future. As in other HGBS studies (e.g. Könyves et al. 2015), we adopted the Bonnor-Ebert (BE) critical mass as an indicator of the maximum mass a core can have whilst remaining stable against gravitational collapse. For an isothermal sphere consisting of particles of mass m , the BE mass is (Bonnor 1956):

$$M_{\text{BE,crit}} = \frac{2.4R_{\text{BE}}kT}{\mu m_{\text{H}}G}, \quad (9)$$

where $\mu = 2.33$ is the mean molecular weight per free particle (Kauffmann et al. 2008), R_{BE} is the BE radius estimated as the deconvolved geometrical mean radius of each source measured in the high-resolution column density map, k is the Boltzmann constant, and G is the gravitational constant. We imposed a deconvolved radius of $6''.1$, which is one-third of the FWHM at $250 \mu\text{m}$, to sources that are unresolved. We note that in the above equation, non-thermal effects like turbulence or magnetic fields are not taken into account. If the source mass is M , we define an object as bound (or prestellar) if it has $M_{\text{BE,crit}}/M \equiv \alpha_{\text{BE}} \leq 2$ (see Könyves et al. 2015). We also note that in the catalogue tables presented in Appendix A, we used the short name prestellar for bound cores and starless for unbound cores, even if both are starless, for consistency with the other HGBS catalogues.

For unresolved cores, the deconvolved radius tends to zero, causing $M_{\text{BE,crit}} \rightarrow 0$, and therefore Könyves et al. (2015) used a set of simulated bound cores to find that bound, unresolved cores can be identified if $\alpha_{\text{BE}} \leq 5(\text{HPBW}/\text{FWHM})^{0.4}$, where FWHM is the non-deconvolved source radius estimated in the high-resolution column density map and HPBW is the linear map resolution, $18''.2$. Sources that fulfilled this criterion were named candidate bound cores while cores with $\alpha_{\text{BE}} \leq 2$ were named robust. In this paper we adopt the same nomenclature.

We simulated bound cores as in Könyves et al. (2015) to derive the completeness limit in mass of our survey. In Appendix B we present our simulations. We find that, for sources with masses greater than $\sim 0.32 M_\odot$, our survey is complete at $\geq 90\%$ in general, but this limit depends on the contrast between source intensity and background level. We also used the simulations to verify that the above formula for candidate bound cores also works in our case.

4.4. Mass, temperature, and radius

Figure 16 shows the relationship between mass and radius for the whole population of 519 cores. At any radius, bound cores are more massive than unbound cores: this is not a consequence of using Eq. (9) as the stability criterion. A massive starless core can be unbound if sufficiently warm and, conversely, a low-mass starless core can be bound if sufficiently cold. The fact that bound cores are more massive implies some relation between M and T , an anticorrelation that is in fact present in our data (see below Fig. 18).

For the population of unbound cores, we find $M \propto R^{1.28}$, shown with the red line in Fig. 16, while for candidate bound

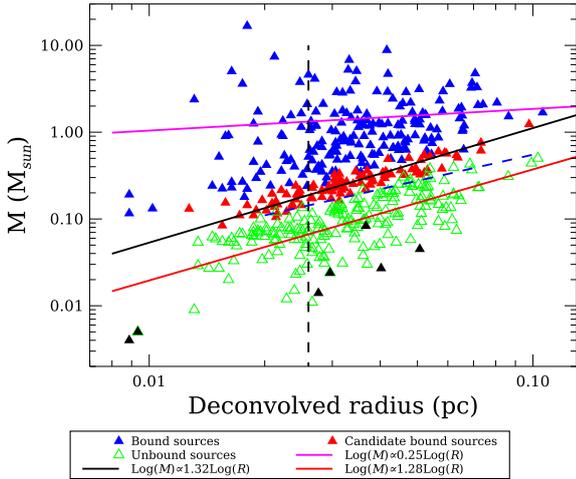


Fig. 16. Relation between mass and radius for all the 519 detected cores: vertical dashed line corresponds to 0.026 pc, the linear distance corresponding to the FWHM ($18''.2$) of the high-resolution column density map at 300 pc. The three power laws are fit to our data. Black triangles are hot starless cores discussed in the text; also the blue dashed line is discussed in the text.

cores, we find $M \propto R^{1.32}$, the solid black line in the same figure. Bound cores do not show a clear relation between mass and radius: indeed, the best-fit slope is $M \propto R^{0.25}$. The lack of correlation is not surprising given that: (i) Eq. (9) is a relation between mass, temperature, and radius. When, as in our case, T is not assumed constant, as it generally was before *Herschel*, there is not an obvious relation between M and R ; (ii) If a bound core is collapsing, it is not even expected to be described by Eq. (9). In fact, our stability criterion is not that M is equal to M_{BE} , but rather that M is greater than M_{BE} by some factor. Coming back to the original paper by Bonnor, in particular the plot $p - V$ in Fig. 1 of Bonnor (1956), bound cores should populate the curve on the left of point A, where instability is dominating. Equation (9) of the present article, on the contrary, describes the condition of a core that is found exactly on point A.

In Fig. 16 the dashed blue line shows the $M-R$ relation for cores at $T = 6.7$ K with $\alpha_{BE} = 2$. This temperature is the lowest we derived for prestellar cores. We consider this line to be the lower-limit for bound cores in our sample. However, we note that for $T = 6.7$ K and $\beta = 2$, the modified black body distribution peaks at $\lambda > 430 \mu\text{m}$, which means that we may lack colder, bound cores in our data because we imposed that $F_\nu(350 \mu\text{m}) > F_\nu(500 \mu\text{m})$ in Sect. 4.2.

The smallest radius in our sample, 0.0089 pc or $6''.1$ at 300 pc, simply reflects our choice to impose a deconvolved radius of one third of the FWHM at $250 \mu\text{m}$ to sources that are unresolved. The largest observed radius is 0.11 pc for bound cores. During the detection and measurement of the sources with *getsources*, we imposed a maximum radius of $220''$ (0.32 pc), which is three to four times larger than the largest core size that we detect. Therefore, this size constraint does not affect our results. Instead, we find that the cores in Perseus have a maximum size of ~ 0.1 pc. This limiting size could be physical or may reflect another, unaccounted for bias. For instance, larger cores may have smaller density contrasts between their peaks and the background, which makes them more difficult to detect.

In Fig. 17 we show the relation between the temperature T of the cores and their mean column density $N(\text{H}_2)$ (defined as $M/\pi R^2$, where T , M , and R are derived from source SED fitting).

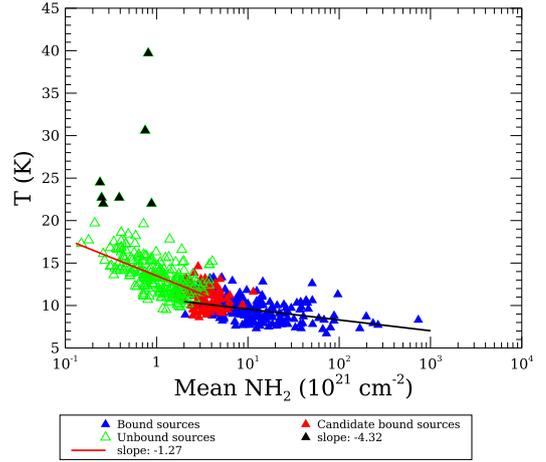


Fig. 17. Relation between T and mean column density of the sources, where symbols are defined as in the previous figure. The two lines are linear fits of the form $T = s \log N(\text{H}_2) + b$. The slopes s are given in the legend and in the text where we also report the uncertainties. Black triangles are hot starless cores discussed in the text.

We modelled the trend seen with the equation $T = s \log N(\text{H}_2) + b$. Two different slopes are found for prestellar and unbound cores: -1.27 ± 0.21 and -4.32 ± 0.41 for bound and unbound, respectively. This result is in contrast with what was found by Marsh et al. (2016) in L1495. These latter authors found that the trend from the unbound populations continues smoothly when entering the bound region, with a single slope being enough to describe the overall trend. Their slope itself, of -3.6 , is different from the value of -4.32 we find for unbound cores. Such a difference could be due to the different environments in which cores are embedded.

In deriving the $M - T$ slope for the unbound cores, we excluded seven objects that are clearly warmer than other cores, namely those with $T > 20$ K. Including these points gives $s = -5.43 \pm 0.67$. These sources are the seven starless cores that show emission at $70 \mu\text{m}$ (see Sect. 4.3). As noted earlier, detection in the PACS blue band does not necessarily mean that a source has evolved into a protostar. In fact, a starless core with $T > 20$ K has the peak of the SED approaching the PACS blue band ($\lambda < 144 \mu\text{m}$ for $\beta = 2$, Elia & Pezzuto 2016) while at low T the peak moves at longer wavelengths so that the emission at $70 \mu\text{m}$ becomes too faint to be detected. Moreover, the SED increases at higher T as $L \propto T^{4+\beta}$ (see Elia & Pezzuto 2016). In principle, one could have bright enough starless core emission at $70 \mu\text{m}$ by also increasing the mass, but the intensity depends only linearly on M , meaning that huge masses are required to make an unbound and cold core visible at short λ . The seven sources described here are warm and have small masses (see black triangles in Fig. 17).

In principle, there is no reason why warm unbound cores should not exist. However, their position in Fig. 17 seems to define a class of sources that are independent from the other unbound cores. It is possible that these sources have poorly determined intensities, which means that T was not estimated correctly. For example, they could be objects falling very close to strong sources, meaning that their envelope is unusually warm. Alternatively, these sources could have some intrinsic properties that make them different from the others. As our goal here is a statistical description of the cold core population in Perseus, we do not further investigate this small sample.

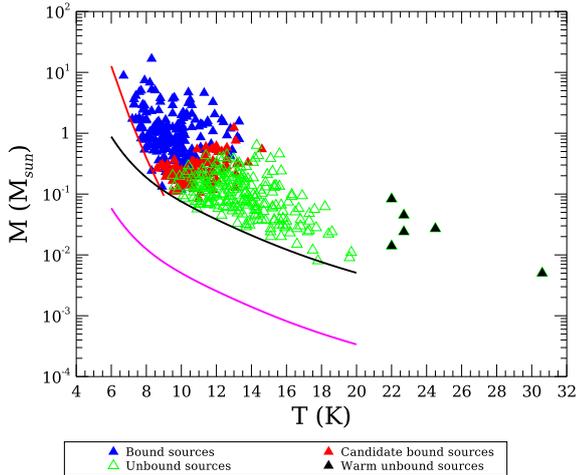


Fig. 18. Relation between T and M of the sources. The warmest unbound source at ~ 40 K is not shown to limit the extension of the x -axis. Bound sources are shown in blue, unbound sources in green, and candidate bound sources in red.

Finally, in Fig. 18 we show the relationship between M and T , which shows similar distinctions between the populations to those seen in Fig. 17. The anticorrelation between M and T is evident (see the discussion about Fig. 16).

For any given temperature, there is a mass below which cores are not detected. To investigate this phenomenon from a quantitative point of view, we studied whether or not, and to what extent, the sensitivity of the *Herschel* instruments plays any role. First we constructed a SED with the SPIRE-PACS Parallel Mode $1\text{-}\sigma$ instrument sensitivity for a point source¹³ divided by $\sqrt{2}$ because of the two directions. The sensitivities are then 14.8 and 33.2 mJy for PACS, and 8.91, 7.42, and 10.6 mJy for SPIRE. We then made a fit to this SED with the same routine and grid of models used for the detected cores, and the same distance of 300 pc. The best fit model has $T = 23.9$ K and $M = 1.7 \times 10^{-4} M_{\odot}$. For each T , we then derived the mass that scales the model to the SED of sensitivities. For instance, at 20 K the mass is $M = 3.4 \times 10^{-4} M_{\odot}$, at 10 K $M = 5.3 \times 10^{-3} M_{\odot}$, and so on. Clearly, the match between models and sensitivity-based SED is increasingly poor as T decreases, but to first approximation, and without considering measurement uncertainties, these values give an estimate of the smallest detectable mass for a given T compatible with *Herschel* sensitivities.

All these values give the magenta line in Fig. 18, which is accurately described by $M \propto T^{-4}$. This line, derived from *Herschel* sensitivities for a point source, can be considered as the locus of the smallest detectable mass of a compact source as a function of temperature. The line is well below the detected masses: indeed, in star forming regions it is well known that the *Herschel* detection of cores is not driven by the sensitivity but by the contrast between core and diffuse medium emission. Interestingly, a simple scaling by the arbitrary factor of 15, which was found by eye, gives the black line that accurately reproduces the smallest detected mass of starless cores for $T \gtrsim 8$ K.

We can see that for $T \lesssim 9$ K, unbound cores are not detected; indeed, the minimum temperature in our catalogue for an unbound core is 9.5 K. This could be an observational bias,

¹³ <http://herschel.esac.esa.int/Docs/PMODE/html/ch02s03.html>. We adopted the sensitivities when scanning in the nominal direction, which were slightly poorer than those in the orthogonal direction.

because at low temperatures, M_{BE} decreases while the smallest detectable mass of a source increases. Therefore, an unbound cold core must have high mass to be detected and large radius so as not to be unbound, but low T and high R equate to a faint and diffuse object that could remain undetected.

The magenta line does not work with bound cores for $T \lesssim 8$ K. The smallest detected mass scales as $T^{-12.2}$, shown in Fig. 18 with a red line. This slope was found by choosing two arbitrary bound cores by eye, with $T = 8.9$ K and $M = 0.13 M_{\odot}$, and with $T = 7.2$ K, $M = 1.73 M_{\odot}$. A possible explanation for this behaviour is based on the finding that more massive prestellar cores are found in regions of high column density (see Fig. 21 in Sect. 5.2). As a consequence, a bound core is detected if massive enough to show up against the bright diffuse medium, and the minimum mass no longer reflects the law $M \propto T^{-4}$ which is more related to the instrument sensitivity. This conclusion is consistent with the result shown in Appendix B.2 where we discuss the completeness in mass.

5. Discussion

5.1. Core mass function

The histograms in Fig. 19 show the CMFs for robust prestellar cores (blue line), and for robust and candidate prestellar cores (red line) in Perseus in the form $dN/d\log M$. The histograms use logarithmic bin sizes of 0.1. The most massive candidate prestellar core has $M = 1.24 M_{\odot}$, meaning that blue and red lines overlap at high masses.

The results found by Swift & Beaumont (2010) show that with 199 sources it is difficult to discern between a power-law tail and a log-normal to model the CMF shape. Therefore, we separately fit both a power law and a log-normal function to the CMF shown in Fig. 19.

The high-mass tail was fit with a power law for $M > 1 M_{\odot}$. Given the limited number of sources, the size of the bin of the histogram may play a role. We therefore followed the prescription of Maíz Apellániz & Úbeda (2005) and used a variable bin size with a fixed number of values in each bin. In this procedure, N values, here representing the core masses, are arranged in m bins whose size is chosen such that each bin contains $n = N/m$ values. This means that each bin will have its own specific size. Maíz Apellániz & Úbeda (2005) showed that this procedure is robust even when $n = 1$, that is, when each bin contains one single value, or the number m of bins is as low as 3.

There are 96 prestellar cores with $M > 1 M_{\odot}$. We constructed different histograms, with 4, 6, 8, 12, 16 or 24 sources per bin. For the fit we adopted the functional form $dN/dM \propto M^{\alpha}$ for the CMF, shown as $dN/d\log M \propto M^{\alpha+1}$ in Fig. 19, and we find mean of the slope for the six cases to be $\alpha = -2.321 \pm 0.035$, in agreement at the 1σ level with the Salpeter slope of -2.35 .

Another survey of cold cores in Perseus was performed by Enoch et al. (2006) at 1.1 mm at CSO using BOLOCAM, with a spatial resolution of $31''$. These latter authors fitted the high-mass part of their CMF with two power laws merging at $2.5 M_{\odot}$. They also tried a fit with only one slope for $M > 0.8 M_{\odot}$. Although they find a single power-law slope for dN/dM of -2.1 ± 0.1 , which is close to the Salpeter value, they note that the CMF is better fit with two power laws. In this respect, it is worth mentioning that the three most massive cores in the survey of Enoch et al. (2006), namely Bolo 48, Bolo 8, and Bolo 43 with masses of $25.6 M_{\odot}$, $11.3 M_{\odot}$, and $10.8 M_{\odot}$ (for an assumed dust temperature of 10 K, distance of 250 pc, and a different opacity law), respectively, were detected by PACS in its blue band

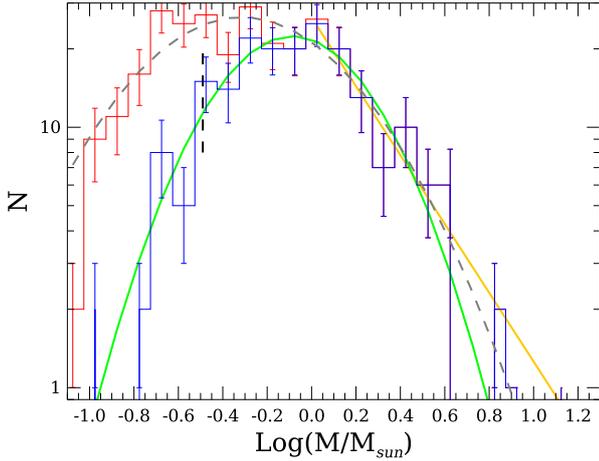


Fig. 19. Core mass function ($dN/d\log M$) for the 199 robust prestellar cores (blue histogram), and for the 299 robust and candidate prestellar cores (red histogram) in Perseus. The two distributions overlap above $M \sim 1.2 M_{\odot}$. Histograms use a bin size of $\log M = 0.1$. Green and dashed grey lines show a log normal fit, and the orange line shows a power-law fit (see text for details). The dashed black line shows the completeness limit at $\sim 0.32 M_{\odot}$.

(ID 330, 68 and 303 in our catalogue; see Appendix A). For this reason, we classified them as protostars, and did not include them in our CMF. It is possible that removing these sources would have made the BOLOCAM-derived CMF better described with a single slope.

Sadavoy et al. (2010) also produced CMFs of Perseus using SCUBA observations at $850 \mu\text{m}$ from the SCUBA Legacy Survey. These latter authors found a $dN/d\log M$ slope of -0.95 ± 0.20 , which is within $2\text{-}\sigma$ of Salpeter. Like Enoch et al. (2006), Sadavoy et al. (2010) also assumed a fixed dust temperature (11 K based on NH_3 observations of cores in Perseus).

Nevertheless, the SCUBA Legacy catalogue was incomplete. The sample did not include B5 and also missed several other high-density regions. To account for the missing areas, Sadavoy et al. (2010) produced predicted CMFs. The predicted CMF was still consistent with Salpeter, and Sadavoy et al. (2010) estimated the number of predicted starless cores to be 232, which is not far from the number of prestellar cores that we find. Even if the predicted number of 232 were to include some unbound cores, SCUBA $850 \mu\text{m}$ data are most sensitive to dense cores that are likely to be bound rather than fluffy unbound objects.

Our CMF in Fig. 19 for robust prestellar cores (blue histogram) can be fit with a log-normal function (green line) with peak at -0.086 ± 0.028 (or $M_{\text{peak}} = 0.820 \pm 0.053 M_{\odot}$ and width 0.347 ± 0.028). The system IMF for field stars in the Galaxy disc was parametrised by Chabrier (2005) in terms of a log-normal function with a peak at $0.25 M_{\odot}$ and $\sigma = 0.55$. If we assume that the IMF of Perseus resembles that of field stars, we derive a star formation efficiency per single core, ϵ_{core} , of 0.30. Adding the 199 candidate prestellar cores gives the CMF shown as the red histogram in Fig. 19. As these cores have $M < 1.2 M_{\odot}$, the effect is to enlarge the distribution and shift the peak to a smaller value. The log-normal function (dashed grey line) has a peak at -0.314 ± 0.033 (or $M_{\text{peak}} = 0.485 \pm 0.037 M_{\odot}$) and a width of 0.471 ± 0.036 ; ϵ_{core} increases to 0.52. In case not all 199 candidate prestellar cores are indeed bound, especially at low masses, we can reasonably estimate ϵ_{core} in the range 0.3–0.4. A similar value, 0.4, was found in Aquila (Könyves et al. 2015), but higher

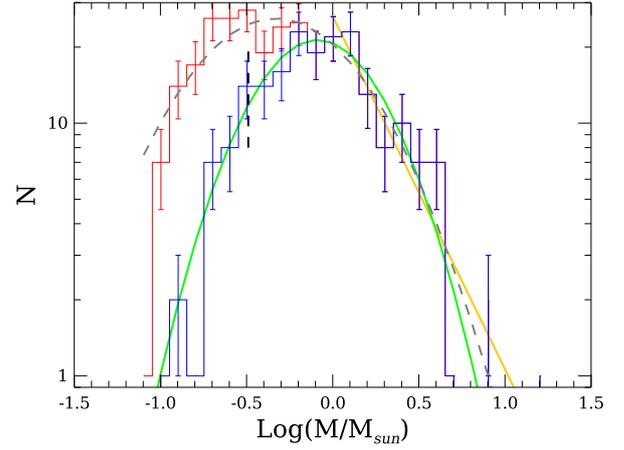


Fig. 20. Prestellar core mass function ($dN/d\log M$) for sources in Perseus when each core is assigned a different distance based on Eq. (1). Lines as in Fig. 19.

values are also reported in HGBS regions (e.g. Marsh et al. 2016; Benedettini et al. 2018, with $\epsilon_{\text{core}} \sim 1$).

Another estimate of ϵ_{core} in Perseus was derived by Mercimek et al. (2017) using data from the literature. These latter authors found $\epsilon_{\text{core}} \sim 0.16$, about half of what we find, but using a different method: they estimated ϵ_{core} as the ratio between the total mass of protostars (Class 0 and Class I objects) and the total mass of starless cores. To compare these two values for ϵ_{core} is not easy because the two methods depend on different assumptions. Here we simply note that, on one side, Mercimek et al. (2017) considered their value a lower limit; on the other side, our value can be lowered if the peak of the IMF is chosen differently.

In fact, instead of using the peak of the IMF for field stars, we can compare the peak of our CMF with the IMF derived in two regions in Perseus: Alves de Oliveira et al. (2013) fit the IMF of IC348 with a log-normal function finding $M_{\text{peak}} = 0.21 \pm 0.02 M_{\odot}$ and $\sigma = 0.52 \pm 0.03$; Scholz et al. (2013) derived instead the median mass of the IMF for IC348 and NG C1333, finding $\bar{M} = 0.27 M_{\odot}$ for both clouds. Since the IMF has a high-mass tail, clearly $\bar{M} > M_{\text{peak}}$, but even assuming $M_{\text{peak}} = \bar{M}$ we conclude that $0.26 \lesssim \epsilon_{\text{core}} \lesssim 0.56$.

Enoch et al. (2006) also fit their CMF with a log-normal distribution for $M > 0.8 M_{\odot}$. They gave two values for M_{peak} , namely $0.9 M_{\odot}$ and $0.3 M_{\odot}$, corresponding to $T_{\text{d}} = 10 \text{ K}$ and 20 K , respectively, because they could not derive the dust temperature of each core, as we did with our data. If we want to compare their M_{peak} with ours, we have to take into account both the different opacities ($0.0074 \text{ cm}^2 \text{ g}^{-1}$ following our opacity law instead of $0.0114 \text{ cm}^2 \text{ g}^{-1}$) and the different distance (300 pc here vs. 250 pc). Considering both differences, the two values for M_{peak} become $\sim 2 M_{\odot}$ and $\sim 0.7 M_{\odot}$ for $T_{\text{d}} = 10$ and 20 K , respectively. Our M_{peak} is closer to the value corresponding to 20 K even if our range in T_{d} (see Fig. 18) suggests that $T_{\text{d}} = 10 \text{ K}$ should be more appropriate.

In Sect. 3.1, we derived a relation between distance and source coordinates. In Fig. 20, we show the resulting CMF for prestellar cores when each source is assigned a different distance. We note that physical radius increases with d while mass increases as d^2 , and therefore, in general, $\alpha_{\text{BE}} \propto d^{-1}$. As a consequence, objects that are closer will generally have lower α_{BE} and objects that are further will have higher α_{BE} . In any case, we find that changing the distance affects our prestellar core counts only minimally as the number of robust prestellar sources remains at

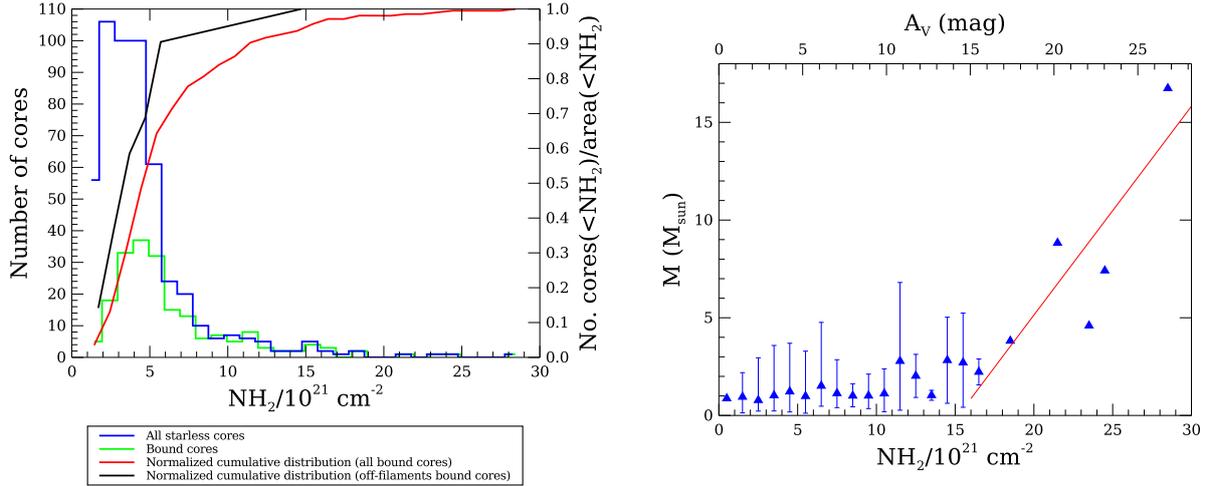


Fig. 21. *Left panel:* comparison of robust bound cores and starless cores. Histograms with the y -axis on the left show the distributions of all starless cores (blue) and all bound cores (green). The red and black lines (right y -axis) show cumulative surface density distributions for all bound cores and the off-filament bound cores, respectively. These distributions were derived by dividing the number of cores found at a certain $N(\text{H}_2)$ level by the total area in the image where the density is below that value. *Right panel:* prestellar core mass vs. background $N(\text{H}_2)$. Triangles are the mean mass averaged over bins of 10^{21} cm^{-2} . The linear fit is shown as a red line.

199 while the number of candidate bound cores increases by 1, from 100 to 101.

Including a varying distance, the log-normal fits do not change significantly: the peak for the robust prestellar cores is $M_{\text{peak}} = 0.816 \pm 0.038 M_{\odot}$ while that for all the bound cores, that is both robust and candidate bound cores, is $M_{\text{peak}} = 0.464 \pm 0.034 M_{\odot}$. The two sets of values (with fixed and varying distance) are equivalent within the uncertainties. On the contrary, the slope of the high-mass power law changes: for $M > 1.03 M_{\odot}$ there are 92 sources and the average of the slopes found with 4 and 23 bins becomes -2.398 ± 0.026 , different from the Salpeter value at 1.8σ . Even if Eq. (1) is only an approximation, it nevertheless highlights the importance of assuming a non-uniform distance for a complex like Perseus.

5.2. Link between filamentary structure and star formation

One of the key results from *Herschel* observations is the deep link between filamentary structures and the star formation in clouds. Numerous studies have shown that star formation preferentially occurs within filaments, with dense prestellar cores primarily in denser filaments (e.g. André et al. 2014, and references therein). This correlation has led to a possible paradigm, where filaments fragment into prestellar cores by gravitational instability (André et al. 2019).

We derived the number of robust prestellar cores inside a filament in two ways: in the most conservative approach, a core is considered internal to the filament if its distance from the spine is less than the width of the filament, or, more precisely, $1.29 \times FWHM$. In the relaxed approach, we simply require that the core be inside the filament border (refer to Sect. 3.2.3 and Fig. 11). The first, conservative sample consists of 167 bound cores (84% of the total) with a median mass of $1.0 M_{\odot}$. The second, relaxed sample contains an additional 23 (12%) bound cores, with a median mass of $0.96 M_{\odot}$. The remaining 9 (4%) prestellar cores have a median mass of $0.47 M_{\odot}$. Bound cores are therefore preferentially found within filaments, where the more massive cores are found.

For the unbound cores, the percentages for the three samples (conservative, relaxed, and out of filaments) are 47%, 14%, and

39%, respectively, with median masses of $0.10 M_{\odot}$, $0.15 M_{\odot}$, and $0.073 M_{\odot}$.

Figure 21 compares the background column density measured by getsources for the 519 starless cores (blue line) and 199 robust prestellar cores (green line; adding the 100 candidate prestellar cores shifts the distribution upwards but does not change the conclusions). The low number of cores at high density reflects to some extent the fact that regions with high $N(\text{H}_2)$ have an intrinsically smaller area than regions at low density. To check for this geometrical bias, we also plot in the same figure the cumulative distribution of bound cores normalised to area. The red curve shows the distribution for all 199 robust prestellar cores and the black curve shows the distribution for the 9 off-filament prestellar cores.

This cumulative distribution shows a few interesting features. First, about 90% of prestellar cores are found above $\sim 2 \times 10^{21} \text{ cm}^{-2}$, in line with the value of $\sim 2 \text{ mag}$ found in Lupus (Benedettini et al. 2018). Higher values are instead found in Aquila (7 mag, Könyves et al. 2015) and in L1495 (6 mag, Marsh et al. 2016). Second, we find two separate regimes in the distribution: an almost linear regime up to $\sim 5 \times 10^{21} \text{ cm}^{-2}$ below which $\sim 60\%$ of bound sources are found, followed by a flatter distribution.

The distribution of the off-filament prestellar cores shows a different trend (clearly, the on-filament distribution overlaps almost perfectly with the cumulative distribution for all bound cores). For the off-filament cores, all but one of the sources are found at $N(\text{H}_2) < 5.5 \times 10^{21} \text{ cm}^{-2}$.

One might speculate that off-filament sources are less massive than on-filament sources because the former are found at lower background column density than the latter. However, this is not the case because if we consider the median of the mass for sources that have $N(\text{H}_2) < 5.5 \times 10^{21} \text{ cm}^{-2}$, we find $\sim 0.9 M_{\odot}$ for on-filament sources. For off-filament sources, the median of the mass is only $0.24 M_{\odot}$.

The right panel of Fig. 21 shows the mass of prestellar cores as a function of the background column density. We find a relatively flat mass distribution up to a threshold of $15 \times 10^{21} \text{ cm}^{-2}$. Above this threshold, sources are found only on-filament (more precisely 8 out of 9 off-filament bound cores are

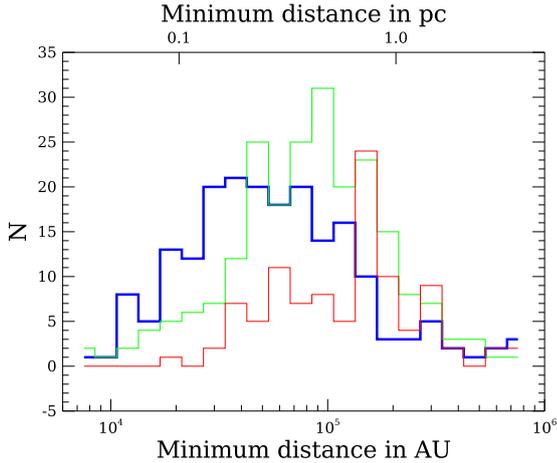


Fig. 22. Minimum distance between bound cores (blue), unbound cores (green), and tentative bound cores (red).

below $5.5 \times 10^{21} \text{ cm}^{-2}$) and their masses grow with the background column density with a slope $1.06 \pm 0.23 M_{\odot} 10^{-21} \text{ cm}^{-2}$, if a linear trend is assumed (red line in figure).

The distributions shown in Fig. 21 raise the question of why we do not detect prestellar cores below $0.9 \times 10^{21} \text{ cm}^{-2}$. For example, the prestellar core with the smallest level of column density has a mass of $0.87 M_{\odot}$, which is 2.7 times above the completeness limit. The area of the map at column densities lower than this limit is 40% of the total map. If cores were distributed randomly, we would expect 90 prestellar cores of any mass to be found at low column density values¹⁴. The fact that we do not find bound cores at such low column densities implies that column densities $N(\text{H}_2) \gtrsim 10^{21} \text{ cm}^{-2}$ or $A_V \gtrsim 1$ are necessary for prestellar cores to form.

The possible existence of an A_V threshold for star formation in Perseus was already investigated in the past with contrasting results. Hatchell et al. (2005) found no column density cutoff for dense cores based on C^{18}O (1–0) observations, whereas Kirk et al. (2006) found that submillimetre clumps were detected only at $A_V > 5\text{--}7$ mag. Sadavoy et al. (2010) found that 94% of SCUBA-identified sources were located at $N(\text{H}_2) > 3 \times 10^{21} \text{ cm}^{-2}$. Above the same level of column density, we find 63% of our sources.

In Fig. 22 we show the histogram of the minimum distance between cores. For robust bound cores, the distribution has a median value of 0.29 pc with a broad peak in the range 0.15–0.37 pc. However, for unbound cores the distribution has a median of 0.47 pc with a broad peak shifted to slightly larger separations than in the previous case, namely between 0.23 and 0.73 pc. For the tentative bound cores the median is 0.74 pc.

If cores were distributed uniformly across the observed region, one would expect a mean distance among bound cores of 0.69 pc, a factor 2.4 larger than the observed value of 0.29 pc. For starless cores, the mean distance is 0.65 pc, a factor 1.4 larger than 0.47 pc. This difference implies that prestellar cores cluster to some extent. This clustering is evident in Fig. 23, which shows the position of all the prestellar cores (red crosses) and of the unbound cores (green crosses). We find that sources do not

¹⁴ Strictly speaking, background column density and column density are not the same thing in the sense that the former is estimated at the position where a source is found and the latter is measured in regions where there are no sources.

Table 9. Slope of the PDF and SFE for the subregions with known protostars.

Region	N	$N(\text{H}_2)$	M	s	SFE
L1448	4	17.1	87	-2.429 ± 0.035	2.2
L1455	8	2.14	477	-2.866 ± 0.024	0.8
NGC1333	44	4.26	613	-2.081 ± 0.011	3.5
Per6	5	1.68	146	-2.40 ± 0.41	1.4
B1	15	0.99	1336	-2.424 ± 0.058	0.6
IC348	26	2.06	1241	-4.042 ± 0.033	1.0
B5	3	2.96	108	-3.632 ± 0.019	1.4
HPZ6	2	2.71	20	-1.56 ± 0.24	4.8
Perseus	126	0.64	11229	-3.190 ± 0.033	0.6

Notes. N : number of protostars (126 instead of 132 for whole Perseus because for six protostars the column density could not be computed, see text); $N(\text{H}_2)$: minimum background column density in 10^{21} cm^{-2} ; M : mass in M_{\odot} for $N(\text{H}_2) > \text{min}$; s slope of PDF; SFE in %.

appear uniformly distributed, but with few exceptions they are confined inside the regions defined in Fig. 9.

5.3. Star-formation rate and efficiency, and core formation efficiency

The star formation rate (SFR) and star formation efficiency (SFE) are two fundamental properties derived from observations that any star-formation theory has to address, explain, and predict. The former is usually estimated as the total mass of young stellar objects (YSOs) divided by the mean duration of star formation (e.g. 2 Myr, Evans et al. 2009).

The total mass of YSOs is usually computed as the number N of YSOs multiplied by the representative value of $0.5 M_{\odot}$. It is clearly necessary to determine N as precisely as possible. Indeed, *Herschel* observations aimed to achieve better completeness for starless cores rather than YSOs, and so from our data alone we cannot derive a reliable SFR for Perseus.

The SFE is defined as the total mass of YSOs divided by the total mass M_{cloud} of the cloud, that is, the mass of protostars plus gas. For each subcloud, the SFE depends on how the border is defined. As we have already discussed in Sect. 3.2.2, defining such borders is not obvious. Sadavoy et al. (2014) computed dense cloud masses for the Perseus subclouds using a column density threshold of $N(\text{H}_2) = 10^{22} \text{ cm}^{-2}$, chosen to select material primarily found in the high-column-density power-law tail. Here we follow another approach based on the local properties of each subcloud. Namely, we find for each subcloud the minimum background column density at the positions of the protostars in each region and compute M_{cloud} for $N(\text{H}_2) > \text{min}$. In Table 9, we report for each zone the number of protostars, the minimum background column density where a protostar is found, the mass of the subcloud within this minimum value, the exponent s of the power-law fit to the associated PDF (see Sect. 3.2.4), and the respective $\text{SFE}(\%) = \sum M_{\text{YSO}} / (\sum M_{\text{YSO}} + M_{\text{cloud}})$. The number of protostars associated to the whole Perseus cloud is 126 instead of 132 because five sources fall in the region observed with PACS only, and one source is in a region masked out for $250 \mu\text{m}$ extraction because of its close proximity to the border of the SPIRE map. In both cases the column density was not computed. L1451, B1E, and L1468 were not included because no protostars were found.

For the whole Perseus cloud, the minimum column density associated with a protostar is $0.64 \times 10^{21} \text{ cm}^{-2}$, below the limit

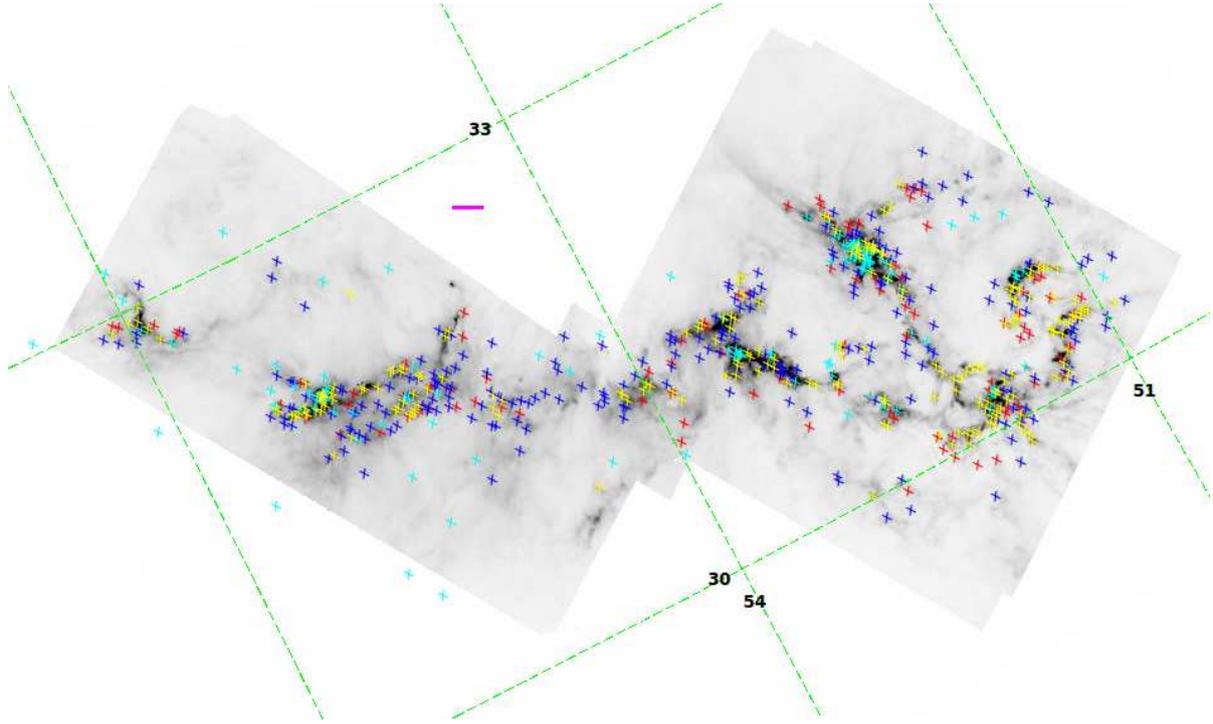


Fig. 23. Column density map overlaid with the positions of the 299 prestellar cores (yellow crosses and red crosses are robust and tentative cores, respectively), 220 unbound cores (blue hashes), and 132 protostars (cyan asterisks; we note that these sources were detected on the $70\ \mu\text{m}$ map, which does not overlap completely with the SPIRE field of view, and so five cores are found outside the column density map). The map is rotated by 28° , with the top and bottom corners cut out. The magenta line in the centre shows the angular scale corresponding to 1 pc at 300 pc; J2000.0 coordinates grid is shown.

of $1.6 \times 10^{21}\ \text{cm}^{-2}$ that defines the last closed contour in the density column map. For this reason, the total mass of the cloud for $N(\text{H}_2) > \text{min}$ is likely underestimated. If we consider only protostars above the last closed column density contour, we find 113 protostars whose minimum $N(\text{H}_2)$ is $1.65 \times 10^{21}\ \text{cm}^{-2}$. The total mass is now $6198\ M_\odot$ and the SFE becomes 0.9%.

Sadavoy et al. (2014) found a linear correlation between SFE and slope s of the PDF, $\text{SFE} \propto 0.02s$, and they discussed the implications of such correlation. These latter authors identified candidate Class 0 sources through their detectability at $70\ \mu\text{m}$, excluding objects that could be more evolved, namely Class I and Class II sources. As our column density map and slopes s are different from those derived by Sadavoy et al. (2014), as discussed in Sects. 3.2.2 and 3.2.4, one may wonder whether or not their result still holds with our data. This question cannot be properly addressed in this paper because in our work we used the criterion of $70\ \mu\text{m}$ visibility to pick up potential protostars in our catalogue, without separating sources into Class 0, I, or II (see Sect. 4.1). Nonetheless, to give at least an indication of whether or not the Sadavoy et al. (2014) linear correlation is also present in our data, we made a linear fit using fifth and sixth columns of Table 9, finding $\text{SFE} \propto 0.013s$; see Fig. 24. Therefore, even with the caveat on the used sample, we confirm the correlation between SFE and s .

As the *Herschel*-based catalogue is more complete in starless cores than in YSOs, it is more reliable to derive from our sample the core formation efficiency (CFE) than the SFE. Similarly to the latter, the former is computed as $\text{CFE}(\%) = \sum M_{\text{pre}} / (\sum M_{\text{pre}} + M_{\text{cloud}})$ where we use only prestellar cores because unbound cores are less likely to form stars. For the CFE, we used the core masses computed during SED fitting instead of assuming a constant $0.5\ M_\odot$ value as done for

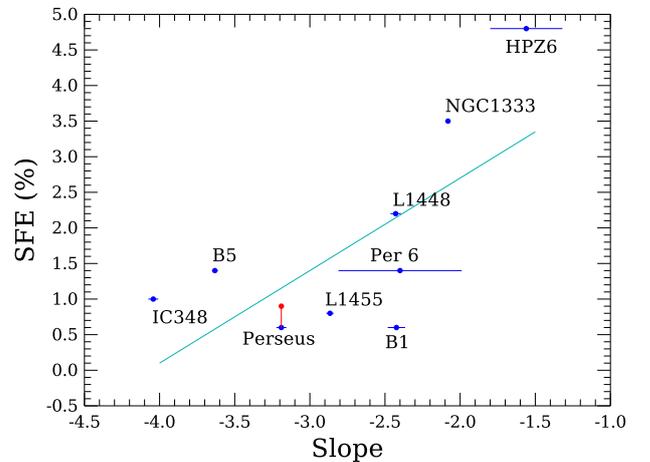


Fig. 24. Star formation efficiency as a function of the PDF slope for the different subclouds in Perseus, and for the whole cloud. The red point is an alternative SFE for Perseus (see text). The cyan line is a linear fit without taking into account uncertainties on the x -axis.

SFE. The results are reported in Table 10. Column N gives the number of prestellar cores per subcloud. In a few cases there is also an uncertainty for N , which we explain below. N_{II} is the number of Class II objects given by Young et al. (2015). These latter authors also counted N_{II} for the rest of the cloud (RC) defined as the whole of Perseus without IC348 and NGC 1333. In Table 10 we define RC as the whole of Perseus but without the four subregions for which the number of Class II objects was found. Column $N(\text{H}_2)$ has the same meaning as in Table 9: the minimum background column density at the position of the

Table 10. CFE for all the subregions in Perseus with at least one bound starless core; where protostars are observed, the lifetime of the prestellar phase is also shown.

Region	N	N_{II}	$N(\text{H}_2)$ (10^{21} cm^{-2})	M (M_{\odot})	CFE (%)	Area (pc^2)	$\Sigma(\text{gas})$ ($M_{\odot} \text{ pc}^{-2}$)	t_{pre} (Myr)
L1451	19		2.15	165	13	1.77	93	
L1448	18		2.37	213	10	1.24	172	
L1455	36^{+8}_{-4}	4	1.55	575	7.9	7.44	77	18.0 ± 4.0
NGC1333	23^{+10}_{-1}	59	3.18	708	7.4	3.82	186	0.77 ± 0.34
Per6	4		3.59	56	6.9	0.47	120	
B1	31^{+6}_{-2}	8	1.67	1053	3.9	10.6	99	7.7 ± 1.5
B1E	8		3.11	146	4.4	1.47	99	
L1468	4		2.64	82	4.3	1.11	73	
IC348	41^{+6}_{-2}	104	3.62	660	8.5	5.25	126	0.78 ± 0.12
B5	6		2.93	109	5.8	0.98	111	
HPZ1	1		3.61	4.5	11	0.04	106	
HPZ2	1		2.08	13	10	0.21	61	
HPZ5	1		2.55	17	2.0	0.23	76	
HPZ6	2		3.99	9.5	21	0.08	126	
Perseus	196^{+60}_{-17}	231	1.67	5960	4.5	76	77	1.69 ± 0.52
RC	68^{+30}_{-9}	56						2.4 ± 1.1

Notes. N : number of prestellar cores; N_{II} : number of Class II objects from Young et al. (2015); $N(\text{H}_2)$: minimum background column density; M : mass for $N(\text{H}_2)$ larger than minimum; CFE: core formation efficiency; Area: area for $N(\text{H}_2)$ larger than minimum; $\Sigma(\text{gas})$: surface density; t_{pre} : lifetime of prestellar phase. RC (rest of cloud) refers to Perseus without the four regions with independent lifetime estimates (we note that Young et al. 2015 define RC as the whole of Perseus excluding only IC348 and NGC1333).

prestellar cores in each region. Mass M , CFE, and area are then computed for $N(\text{H}_2) > \text{min}$. The ratio M/area gives the surface density reported in the $\Sigma(\text{gas})$ column. Finally, t_{pre} is the estimated lifetime of the prestellar core phase, derived by scaling the number of bound cores to the number of Class II objects, and assuming a constant SFR: $t_{\text{pre}} = N_{\text{pre}}/N_{\text{II}} \times t_{\text{II}}$. As before, we fixed t_{II} to the usually adopted value of 2 Myr for the duration of the Class II.

The uncertainty in t_{pre} was derived based on α_{BE} , the criterion used to define bound a core (see Sect. 4.3). We first derived the uncertainty $\Delta\alpha_{\text{BE}}$ from the uncertainties on mass and temperature for each source. Among the bound cores we then counted the number of sources with $\alpha_{\text{BE}} + \Delta\alpha_{\text{BE}} > 2$: within the uncertainty, these sources could be unbound cores. Let us call this number N_- . Similarly, unbound or candidate bound sources with a reliable SED and with $\alpha_{\text{BE}} - \Delta\alpha_{\text{BE}} \leq 2$ could be potential prestellar cores. We define N_+ as the number of these sources. Table 10 provides the number of prestellar cores as $N_{N_-}^{N_+}$ for the subclouds for which N_{II} has been estimated. Here, t_{pre} was not computed for the other subclouds for which an estimate of N_{II} is not available. The uncertainty on N is assumed as $\max(N_+, N_-)$. This uncertainty is then propagated to t_{pre} . The resulting lifetime for the whole Perseus cloud is $t_{\text{pre}} = 1.69 \pm 0.52$ Myr, which overlaps, within 1σ , with the value 1.2 ± 0.3 Myr found in Aquila by Könyves et al. (2015).

We see that for L1455 and B1, t_{pre} is much higher than for the other regions. This could be explained if both clouds were very young, such that star formation started very recently. In this case, the number of Class II objects is less than expected assuming a constant SFR (for a similar case reported in Lupus I and Lupus IV, see Rygl et al. 2013). On the contrary, in NGC 1333 and IC348 we observe a deficit of prestellar cores compared to Class II objects. This can be interpreted as a decelerating SFR in these two regions, with IC348 having the shortest lifetime, in line with the known fact that this cloud is the oldest of the Perseus

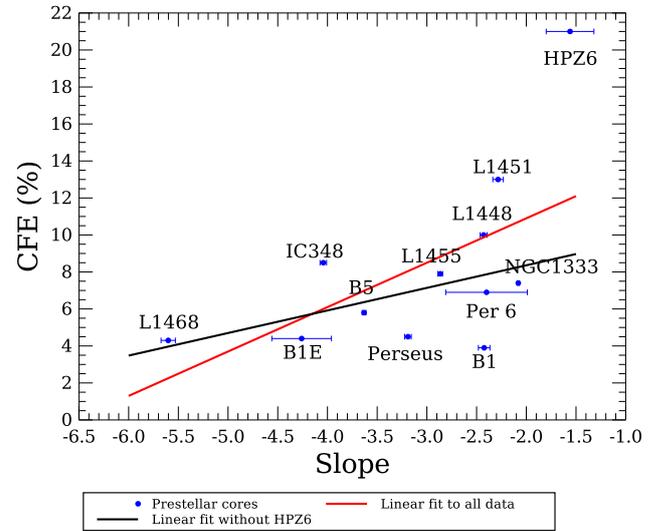


Fig. 25. Core formation efficiency as a function of the power-law slope of the PDFs for the different subclouds in Perseus, and for the whole cloud.

complex. We do not have the number of Class II stars for the other regions, but the relatively high t_{pre} for RC might suggest that star formation also started recently in the other subclouds.

In Fig. 25 we show the CFE as a function of the PDF power-law exponent. Similarly to the case of SFE, for CFE we can also consider the threshold in column density of $1.6 \times 10^{21} \text{ cm}^{-2}$. In doing so, N decreases to 196 with a CFE of 4.5%, which is the value used in Fig. 25. The red line is a linear fit to all the data with a slope of 2.4. Here, HPZ6 appears to be an outlier; if its value is not used in the fit then the result is the black line, with a slope 1.22.

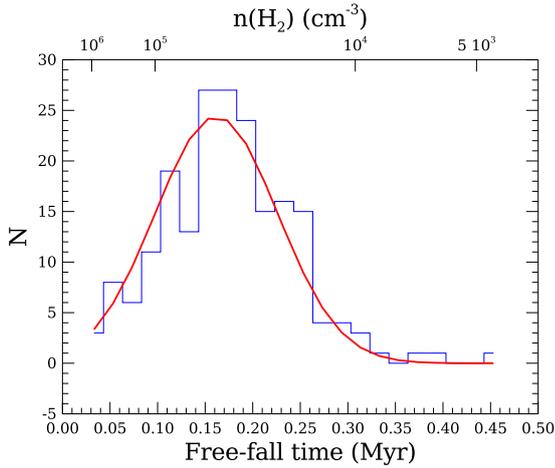


Fig. 26. Distribution of free-fall time for robust prestellar cores. Top-axis is the uniform density for spherical cores ($3M/4\pi R^3$). The red line is a Gaussian fit.

The mean value of all the subclouds excluding Perseus as a whole is $8.2 \pm 4.8\%$ or $7.2 \pm 2.8\%$ if we exclude the HPZ6 zone: the CFE of the whole cloud is within 1σ of these values.

Another important timescale to compare with our estimate of the prestellar phase duration is the free-fall time, that is, the time required for a uniform sphere of density ρ and pressureless gas, i.e. gas at zero temperature, to collapse to infinite density (e.g. Krumholz 2015):

$$t_{\text{ff}} = \sqrt{\frac{3\pi}{32G\rho}}, \quad (10)$$

where G is the gravitational constant and ρ is the uniform density.

In Fig. 26 we show the distribution of t_{ff} for the 199 robust bound cores. The Gaussian fit to the distribution, plotted in red, shows a peak at 0.16 ± 0.01 Myr and $\sigma = 0.06 \pm 0.01$ Myr, where we assigned an uncertainty of 0.01 Myr (half the size of the histogram bin) to both parameters. For a star of $0.25 M_{\odot}$, corresponding to the peak of the IMF, this implies a mean mass accretion rate of $\dot{M} \sim 1.6 \times 10^{-6} M_{\odot} \text{ yr}^{-1}$. Our peak of 0.16 Myr is close to the value of 0.10 Myr derived by Mercimek et al. (2017) using the sample of 69 cores defined by Sadavoy et al. (2010).

6. Conclusions

In this paper, we analyse photometric observations of the Perseus complex executed with PACS at 70 and 160 μm , and SPIRE at 250 μm , 350 μm , and 500 μm . We can summarise our main conclusions as follows.

1. Using distances to protostars from *Gaia* DR2, we find a linear correlation between coordinates (α, δ) in the cloud and distance across the whole Perseus cloud. Although qualitatively already suspected, this correlation is the first estimate of such a trend. Distances estimated from our relation do not take into account the depth of the region at any point and are meant to give representative values; but they are not limited to a few lines of sight and are in very good agreement with other recent and more statistically robust estimates reported in the literature. We confirm that 300 pc is a good mean distance value for Perseus.

2. We produce dust-temperature and column density maps. We find dust temperature to be in the range 10–28 K. Two main peaks in T_{d} distribution are present at 16.4 and 17.1 K, corresponding to the average temperature of the diffuse medium in western and eastern halves of Perseus, respectively. A small bump is present at 19.2 K, which is associated with the inner part of NGC 1333 and IC348. The column density range is $2.5 \times 10^{20} - 1.4 \times 10^{23} \text{ cm}^{-2}$ with the smallest closed contour at $1.6 \times 10^{21} \text{ cm}^{-2}$. T_{d} and $N(\text{H}_2)$ are found to be anticorrelated throughout the complex, with the exception of the inner parts of NGC 1333 and IC348, which are both warm and dense. We also show that a lack of data at $\lambda < 160 \mu\text{m}$ causes T_{d} to be underestimated by more than 10% for $T_{\text{d}} > 22$ K.
3. Extrapolating the column density and temperature maps to 70 μm shows that most of the cloud would appear as infrared dark clouds if observed with enough sensitivity. We also find excess emission at 70 μm above the extrapolated fluxes, which we attribute to very small dust grains.
4. We find that the optical depth derived from our column density map, and extrapolated at 850 μm , agrees with that measured with *Planck* for $\tau_{\text{P}} < 3 \times 10^{-4}$, where τ_{P} is the optical depth measured by *Planck*. At higher τ_{P} , the optical depth found with *Herschel* is higher than that found with *Planck* data. This could be due to a change in the dust opacity. On the contrary, our column density map disagrees with the column density maps derived from near-infrared extinction maps if masses within the same A_{V} contour are compared; better agreement is found when masses are derived from the same areas in the two maps. In any case, the comparison is not easy because of the varying dust opacity, or because in star-forming regions $R \equiv A_{\text{V}}/E(B - V) > 3.1$ and in Perseus it is also variable throughout the complex.
5. We find a typical filament width of 0.08 pc for an assumed distance of 300 pc, confirming the general result of previous HGBS studies that a characteristic filament width of 0.1 pc exists in star-forming regions.
6. On average, star-forming filaments in Perseus have mean line masses significantly below the critical value of $16 M_{\odot} \text{ pc}^{-1}$, above which an isothermal filament at 10 K and infinite radius becomes unstable to radial collapse. However, we suggest that with a proper internal volume density, a filament with a line mass as low as $8 M_{\odot} \text{ pc}^{-1}$, which is half the critical line mass, can fragment.

We extracted compact cores from the intensity maps and used SED fitting to determine their masses and temperatures. Our findings are as follows.

7. We find 199 starless cores have greater masses than their corresponding Bonnor-Ebert masses, above which an isothermal sphere of gas cannot provide sufficient support against gravity. These cores are therefore thought to be collapsing, or to be close to collapse, and to be forming stars. We define these sources as robust bound or prestellar cores. Another 100 cores may also be collapsing and we define these as candidate prestellar cores. On the other hand, 220 sources have masses below their corresponding Bonnor-Ebert threshold and so they are not thought to be dominated by gravity. We name these objects unbound cores. We also declare 132 cores as protostars, because they show, through emission at 70 μm , that an internal source of energy should already be present. Finally, the remaining 165 sources cannot be satisfactorily fit with a modified blackbody because their SEDs are too uncertain.

8. On the $M - R$ plane, starless cores do not trace a clear path but show a large scatter with a faint hint of a correlation between M and R . Ignoring the scatter and fitting a power law anyway, we find $M \propto R^{1.28}$ for unbound sources. A similar relation is found for candidate prestellar cores: $M \propto R^{1.32}$. However, for bound sources, we find $M \propto R^{0.25}$, a relation that is shallower than expected for Bonnor-Ebert spheres ($M \propto R$). We suggest that our slope of 0.25 likely reflects the lack of a correlation for the prestellar cores as a whole, which is probably a hint that these sources are indeed collapsing. From M and R , we derive a mean column density for sources that appears anticorrelated with temperature, with different exponents for bound and unbound sources.
9. We find that M and T are also anticorrelated. We show that, for a given T , the smallest detectable mass of an unbound core is given by a combination of the *Herschel* instruments sensitivity, which gives a relation $M \propto T^{-4}$, scaled by an empirical factor that is likely due to the presence of the bright diffuse medium in which cores are embedded. However, for bound cores, the relation is much steeper $M \propto T^{-12}$, which is probably a consequence of the fact that more massive and colder bound cores are found in brighter regions of the clouds.
10. The prestellar CMF can be modelled with a log-normal curve with a peak at $M \sim 0.82 M_{\odot}$, which is above the completeness limit in mass estimated at $0.32 M_{\odot}$. Assuming a one-to-one mapping of the CMF over the IMF, this peak implies a core mass efficiency of 0.30. The high-mass tail of the CMF can be modelled with a power law with a slope of -2.321 ± 0.035 , which is in agreement, at 1σ , with the Salpeter IMF value of -2.35 . The two distributions, log-normal and power law, were modelled separately because the number of bound cores is not large enough to make a combined fit; for the same reason, we cannot decipher whether or not one distribution alone can fit the whole dataset.
11. More than 84% of robust bound cores are found within filaments versus 47% for unbound cores. On-filament bound sources are on average more massive than the off-filament ones. No bound cores are found below $0.9 \times 10^{20} \text{ cm}^{-2}$. Should sources be distributed uniformly on the map, 90 prestellar cores of any mass should have been found at column densities lower than that value. We propose that a limit of $\sim 1 \times 10^{21} \text{ cm}^{-2}$ in column density is required to form bound cores.
12. We find a correlation between SFE and the slope of the column density PDF for different Perseus subclouds. The correlation has a slope of 1.3%. We also find a correlation between SFE and CFE, with a slope of between 1.2 and 2.4%, depending on whether we use all the subclouds or exclude one extreme value from the sample.
13. Assuming a constant SFR we derive a lifetime for the prestellar phase t_{pre} of 1.69 Myr for Perseus but with significant scatter for the individual subclouds.
14. We find a typical free fall time of 0.16 Myr, which implies a mean mass accretion rate of $\dot{M} \sim 1.6 \times 10^{-6} M_{\odot} \text{ yr}^{-1}$ to form a star of $0.25 M_{\odot}$, which is at the peak of the IMF.

The catalogue of all 816 detected sources is presented in the Appendix A where we give an example of a protostar, a prestellar core, and an unbound core. The full catalogue can be found at the CDS.

We also list in Appendix E sources that were excluded from our analysis, such as galaxies for example. This catalogue is not part of the HGBS official products, and is presented here in its entirety.

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References

- Alves, J., Lombardi, M., & Lada, C. J. 2007, *A&A*, 462, L17
- Alves de Oliveira, C., Moraux, E., Bouvier, J., et al. 2013, *A&A*, 549, A123
- Alves, J., Lombardi, M., & Lada, C. J. 2017, *A&A*, 606, L2
- Andrae, R. 2010, ArXiv e-prints [arXiv:1009.2755]
- Andrae, R., Schulze-Hartung, T., & Melchior, P. 2010, ArXiv e-prints [arXiv:1012.3754]
- André, P., Men’shchikov, A., Bontemps, S., et al. 2010, *A&A*, 518, L102
- André, P., Di Francesco, J., Ward-Thompson, D., et al. 2014, *Protostars and Planets VI*, eds. H. Beuther, R. S. Klessen, C. P. Dullemond, & T. Henning (Tucson, AZ: University of Arizona Press), 27
- André, P., Arzoumanian, D., Könyves, V., Shimajiri, Y., & Palmeirim, P. 2019, *A&A*, 629, L4
- Arce, H. G., Borkin, M. A., Goodman, A. A., Pineda, J. E., & Beaumont, C. N. 2011, *ApJ*, 742, 105
- Arzoumanian, D., André, P., Könyves, V., et al. 2019, *A&A*, 621, A42
- Bally, J., Walawender, J., Johnstone, D., Kirk, H., & Goodman, A. 2008, *Handbook of Star Forming Regions*, ed. B. Reipurth (San Francisco: ASP), 4, 308
- Bauke, H. 2007, *Eur. Phys. J. B*, 58, 167
- Benedettini, M., Schisano, E., Pezzuto, S., et al. 2015, *MNRAS*, 453, 2036
- Benedettini, M., Pezzuto, S., Schisano, E., et al. 2018, *A&A*, 619, A52
- Bernard, J.-P., Paradis, D., Marshall, D. J., et al. 2010, *A&A*, 518, L88
- Bobylev, V. V. 2016, *Astron. Lett.*, 42, 544
- Bohlin, R. C., Savage, B. D., & Drake, J. F. 1978, *ApJ*, 224, 132
- Bonnor, W. B. 1956, *MNRAS*, 116, 351
- Bresnahan, D., Ward-Thompson, D., Kirk, J. M., et al. 2018, *A&A*, 615, A125
- Cambrésy, L. 1999, *A&A*, 345, 965
- Carey, S. J., Clark, F. O., Egan, M. P., et al. 1998, *ApJ*, 508, 721
- Černis, K. 1993, *Balt. Astron.*, 2, 214
- Chabrier, G. 2005, *Astrophys. Space Sci. Lib.*, 327, 41
- Chen, M. C.-Y., Di Francesco, J., Johnstone, D., et al. 2016, *ApJ*, 826, 95
- Chen, H. H.-H., Burkhardt, B., Goodman, A., & Collins, D. C. 2018, *ApJ*, 859, 162
- Clauset, A., Rohilla Shalizi, C., & Newman, M. E. J. 2009, *SIAM Rev.*, 51, 661
- COMPLETE team 2012, *FCRAO Perseus 13CO cubes and map*

- di Francesco, J., Evans, N. J., I., Caselli, P., et al. 2007, in *Protostars and Planets V*, eds. B. Reipurth, D. Jewitt, & K. Keil (Tucson, AZ: University of Arizona Press), 17
- Di Francesco, J., Keown, J., Fallscheer, C., et al. 2020, *ApJ*, 904, 172
- Elia, D., & Pezzuto, S. 2016, *MNRAS*, 461, 1328
- Enoch, M. L., Young, K. E., Glenn, J., et al. 2006, *ApJ*, 638, 293
- Evans, II, N. J., Dunham, M. M., Jørgensen, J. K., et al. 2009, *ApJS*, 181, 321
- Federrath, C., & Klessen, R. S. 2013, *ApJ*, 763, 51
- Fiorellino, E., Elia, D., André, P., et al. 2021, *MNRAS*, 500, 4257
- Fischera, J., & Martin, P. G. 2012, *A&A*, 542, A77
- Foster, J. B., Mandel, K. S., Pineda, J. E., et al. 2013, *MNRAS*, 428, 1606
- Gaia Collaboration (Prusti, T., et al.) 2016, *A&A*, 595, A1
- Griffin, M. J., Abergel, A., Abreu, A., et al. 2010, *A&A*, 518, L3
- Hacar, A., Tafalla, M., & Alves, J. 2017, *A&A*, 606, A123
- Hatchell, J., Richer, J. S., Fuller, G. A., et al. 2005, *A&A*, 440, 151
- Herbig, G. H. 1998, *ApJ*, 497, 736
- Hirota, T., Bushimata, T., Choi, Y. K., et al. 2008, *PASJ*, 60, 37
- Hirota, T., Honma, M., Imai, H., et al. 2011, *PASJ*, 63, 1
- Johnstone, D., Wilson, C. D., Moriarty-Schieven, G., et al. 2000, *ApJ*, 545, 327
- Jørgensen, J. K., Harvey, P. M., Evans, II, N. J., et al. 2006, *ApJ*, 645, 1246
- Joye, W. A., & Mandel, E. 2003, *ASP Conf. Ser.*, 295, 489
- Kauffmann, J., Bertoldi, F., Bourke, T. L., Evans, N. J., I., & Lee, C. W. 2008, *A&A*, 487, 993
- Kirk, H., Johnstone, D., & Di Francesco, J. 2006, *ApJ*, 646, 1009
- Könyves, V., André, P., Men'shchikov, A., et al. 2015, *A&A*, 584, A91
- Könyves, V., André, P., Arzoumanian, D., et al. 2020, *A&A*, 635, A34
- Krumholz, M. R. 2015, ArXiv e-prints [arXiv:1511.03457]
- Ladjelate, B., André, P., Könyves, V., et al. 2020, *A&A*, 638, A74
- Larson, R. B. 1969, *MNRAS*, 145, 271
- Li, A., & Draine, B. T. 2001, *ApJ*, 554, 778
- Li Causi, G., Schisano, E., Liu, S. J., Molinari, S., & Di Giorgio, A. 2016, *Proc. SPIE*, 9904, 99045V
- Luri, X., Brown, A. G. A., Sarro, L. M., et al. 2018, *A&A*, 616, A9
- Maíz Apellániz, J., & Úbeda, L. 2005, *ApJ*, 629, 873
- Marsh, K. A., Kirk, J. M., André, P., et al. 2016, *MNRAS*, 459, 342
- Men'shchikov, A. 2013, *A&A*, 560, A63
- Men'shchikov, A., André, P., Didelon, P., et al. 2012, *A&A*, 542, A81
- Mercimek, S., Myers, P. C., Lee, K. I., & Sadavoy, S. I. 2017, *AJ*, 153, 214
- Molinari, S., Swinyard, B., Bally, J., et al. 2010, *A&A*, 518, L100
- Molinari, S., Schisano, E., Faustini, F., et al. 2011, *A&A*, 530, A133
- Motte, F., Andre, P., & Neri, R. 1998, *A&A*, 336, 150
- Nutter, D., & Ward-Thompson, D. 2007, *MNRAS*, 374, 1413
- Ortiz-León, G. N., Loinard, L., Dzib, S. A., et al. 2018, *ApJ*, 865, 73
- Ostriker, J. 1964, *ApJ*, 140, 1056
- Ott, S. 2010, *ASP Conf. Ser.*, 434, 139
- Palmeirim, P., André, P., Kirk, J., et al. 2013, *A&A*, 550, A38
- Pezzuto, S., Elia, D., Schisano, E., et al. 2012, *A&A*, 547, A54
- Piazzo, L., Calzoletti, L., Faustini, F., et al. 2015, *MNRAS*, 447, 1471
- Pilbratt, G. L., Riedinger, J. R., Passvogel, T., et al. 2010, *A&A*, 518, L1
- Planck Collaboration XI. 2014, *A&A*, 571, A11
- Poglitsch, A., Waelkens, C., Geis, N., et al. 2010, *A&A*, 518, L2
- Polychroni, D., Schisano, E., Elia, D., et al. 2013, *ApJ*, 777, L33
- Rayner, T. S. M., Griffin, M. J., Schneider, N., et al. 2017, *A&A*, 607, A22
- Rebull, L. M., Stapelfeldt, K. R., Evans, II, N. J., et al. 2007, *ApJS*, 171, 447
- Ridge, N. A., Di Francesco, J., Kirk, H., et al. 2006a, *AJ*, 131, 2921
- Ridge, N. A., Schnee, S. L., Goodman, A. A., & Foster, J. B. 2006b, *ApJ*, 643, 932
- Ripepi, V., Palla, F., Marconi, M., et al. 2002, *A&A*, 391, 587
- Román-Zúñiga, C. G., Lada, C. J., & Alves, J. F. 2009, *ApJ*, 704, 183
- Roy, A., André, P., Arzoumanian, D., et al. 2015, *A&A*, 584, A111
- Rygl, K. L. J., Benedettini, M., Schisano, E., et al. 2013, *A&A*, 549, L1
- Sadavoy, S. I., Di Francesco, J., Bontemps, S., et al. 2010, *ApJ*, 710, 1247
- Sadavoy, S. I., di Francesco, J., André, P., et al. 2012, *A&A*, 540, A10
- Sadavoy, S. I., Di Francesco, J., André, P., et al. 2014, *ApJ*, 787, L18
- Salpeter, E. E. 1955, *ApJ*, 121, 161
- Schisano, E., Rygl, K. L. J., Molinari, S., et al. 2014, *ApJ*, 791, 27
- Schisano, E., Molinari, S., Elia, D., et al. 2020, *MNRAS*, 492, 5420
- Schnee, S., Li, J., Goodman, A. A., & Sargent, A. I. 2008, *ApJ*, 684, 1228
- Schneider, N., Bontemps, S., Simon, R., et al. 2011, *A&A*, 529, A1
- Schneider, N., Ossenkopf, V., Csengeri, T., et al. 2015, *A&A*, 575, A79
- Scholz, A., Geers, V., Clark, P., Jayawardhana, R., & Muzic, K. 2013, *ApJ*, 775, 138
- Stanke, T., Smith, M. D., Gredel, R., & Khanzadyan, T. 2006, *A&A*, 447, 609
- Strom, S. E., Strom, K. A., & Carrasco, L. 1974, *PASP*, 86, 798
- Swift, J. J., & Beaumont, C. N. 2010, *PASP*, 122, 224
- Testi, L., & Sargent, A. I. 1998, *ApJ*, 508, L91
- Ward-Thompson, D., André, P., Crutcher, R., et al. 2007, *Protostars and Planets V*, eds. B. Reipurth, D. Jewitt, & K. Keil (Tucson, AZ: University of Arizona Press), 33
- Wright, E. L., Eisenhardt, P. R. M., Mainzer, A. K., et al. 2010, *AJ*, 140, 1868
- Young, K. E., Young, C. H., Lai, S.-P., Dunham, M. M., & Evans, II, N. J. 2015, *AJ*, 150, 40
- Zari, E., Lombardi, M., Alves, J., Lada, C. J., & Bouy, H. 2016, *A&A*, 587, A106
- Zucker, C., Schlafly, E. F., Speagle, J. S., et al. 2018, *ApJ*, 869, 83
- Zucker, C., Speagle, J. S., Schlafly, E. F., et al. 2019, *ApJ*, 879, 125

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Appendix A: The catalogue

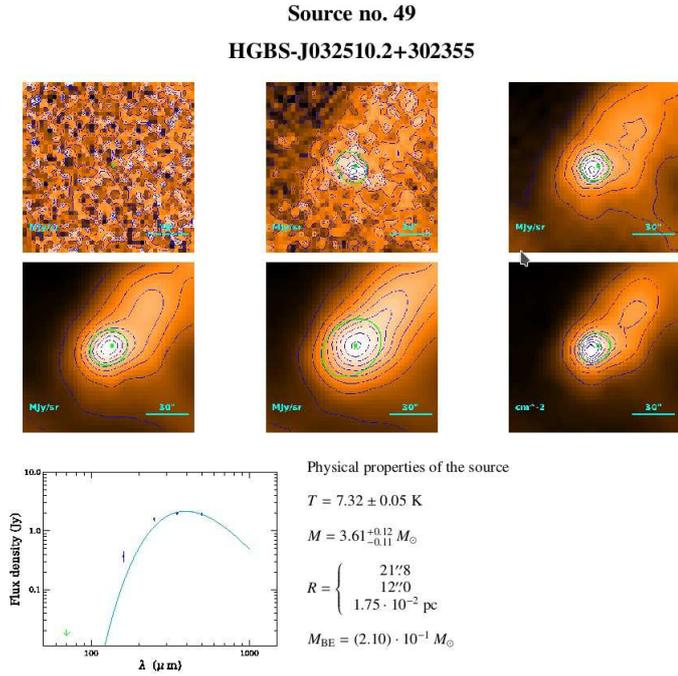


Fig. A.1. An example of the atlas of the Perseus cores showing the prestellar core #49. The original size of the page is A4, here the page has been shrunk to one third.

The CDS material provides data for all 816 candidate cores in the form of an atlas, one page per object. Figure A.1 shows an example for a prestellar core (#49) and Fig. A.2 shows an example for a protostellar core (#55). For each core we provide the following:

- A figure showing a $2' \times 2'$ cutout around each source at $70 \mu\text{m}$ (top left), $160 \mu\text{m}$ (top middle), $250 \mu\text{m}$ (top right), $350 \mu\text{m}$ (bottom left), $500 \mu\text{m}$ (bottom middle), and from the high-resolution column density map (bottom right). If the source is detected by *getsources*, a green ellipse shows the source size from the extraction. We note that some sources were not observed at all wavelengths, because PACS and SPIRE have slightly different fields of view.
- Below the six cutouts, the figure on the left shows the observed SED (blue points) and its best-fit modified black body function (cyan curve). On the right we give the best-fit value for dust temperature (T) and mass (M) from the SED fit. Then, the source radius (R), that is the geometrical mean of major and minor FWHMs both measured from the high-resolution column density map, is given in arcseconds, before and after the deconvolution with the instrumental FWHM ($18''$); the deconvolved radius is given also in parsecs in the third row. If a source is unresolved, we use a source size of $< 6''.1$ (or one-third of the beam). In this case the physical radius in parsecs is computed assuming $R = 6''.1$. Protostars were detected using the $70 \mu\text{m}$ intensity map only.

As a residual misalignment may remain among the maps (see Sect. 2), the $70 \mu\text{m}$ coordinates of the centre may appear offset in the SPIRE intensity maps. The misalignment is smaller than the instrument spatial resolution, but before running *getsources* all the maps are reprojected onto a grid with pixel size $3''$, and so the displacement, if present, appears enhanced. Adopting the coordinates derived during the starless core detection the protostar would appear well centred at the SPIRE wavelengths but offset in the PACS bands. To show this effect, we use Fig. A.3

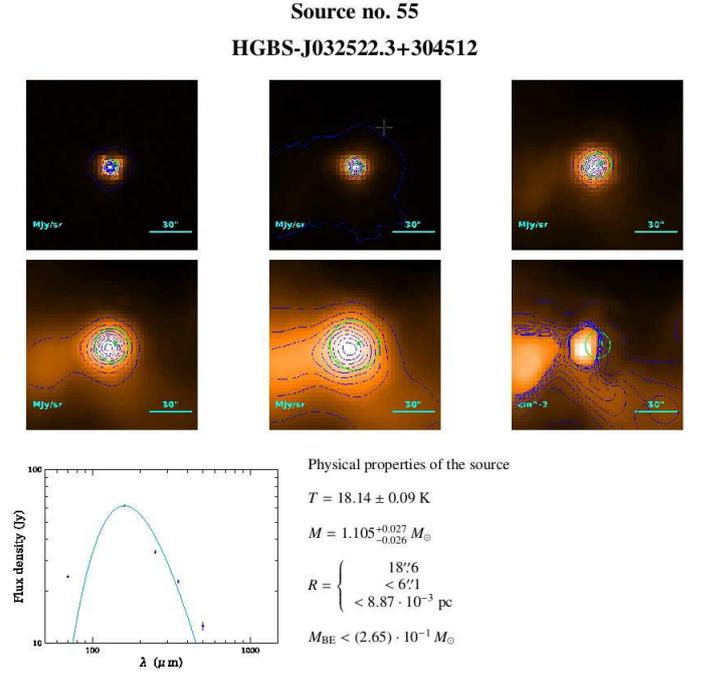


Fig. A.2. An example of the atlas of the Perseus cores showing the protostar core #55. Also in this case the page has been shrunk.

where the background image in both panels is the $250 \mu\text{m}$ cutout. In the upper panel, we show the ellipse and centre from the protostar catalogue (obtained from the PACS extraction), whereas in the bottom panel, we show the ellipse and centre from the starless core catalogue (obtained from the SPIRE extraction). Clearly, at $250 \mu\text{m}$ the source is better centred in the SPIRE extraction. Therefore, for protostars, data in the PACS bands were taken from the protostars catalogue, while data in the SPIRE bands were taken from the starless catalogue.

A similar problem occurs at $160 \mu\text{m}$ for starless cores. In this case, the detection is done in the SPIRE maps, in the high-resolution column density map and in the $160 \mu\text{m}$ temperature-corrected map, but not in the $160 \mu\text{m}$ intensity map that is used for photometry extraction at the position found in the other maps. This strategy could be what causes many of the sources to have only an upper limit at $160 \mu\text{m}$ and in some cases the upper limit poses a problem when fitting the SED. In Fig. A.4 we show such an example. The top-left panel shows the cutout in the PACS red band where *getsources* could not detect any source. However, even by eye it is possible to see that some emission is present at the position of the source that was not recognised as a real object. The bottom-left panel shows the corresponding SED with the 3σ upper limit estimated with *getsources*.

We test our strategy by performing an alternative extraction using all five *Herschel* intensity maps simultaneously. The benefit of this approach is that the direct detection at $160 \mu\text{m}$ may help to extract photometry at this wavelength for weak sources. This alternative catalogue was used only to see if we could assign a measurement to sources with a stringent upper limit at $160 \mu\text{m}$, like the one in Fig. A.4, with the conventional strategy. We found measurements for 12 sources using the alternative extraction. The top-right panel of Fig. A.4 shows that now the starless core is detected also at $160 \mu\text{m}$. In the bottom-right panel we show how the SED fitting benefits from this measurement.

The new $160 \mu\text{m}$ measurement has a large impact on the source physical parameters, with the temperature increasing

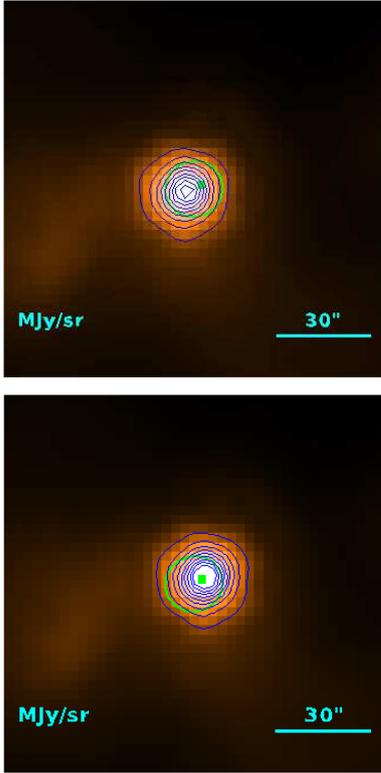


Fig. A.3. Coordinates of source #55 at $250 \mu\text{m}$ as derived in the protostars (top) and starless cores (bottom) catalogues, respectively.

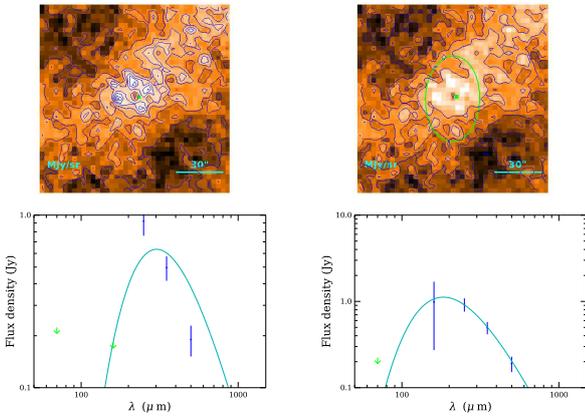


Fig. A.4. Effects on the SED when using an alternative extraction strategy (see text). Top panels: undetected and detected source at $160 \mu\text{m}$. Bottom panels: SEDs with an upper limit and a measure at $160 \mu\text{m}$.

from 10.42 K to 16.14 K and the mass decreasing from $0.627 M_{\odot}$ to $0.073 M_{\odot}$. As there are 22 sources in the final catalogue with masses $0.6 \leq M/M_{\odot} < 0.7$, two of which have no reliable SED fit, it is clear that removing even one single source may have some impact on the CMF.

Another approach with which we tried to solve the problem of the upper limits at $160 \mu\text{m}$ was to exclude this value during the fitting procedure. When *getsources* cannot detect a source found in other images at a certain wavelength, it sets the so-called monochromatic significance to the special value $9.999\text{E}-31$. When we found it evident that the upper limit at $160 \mu\text{m}$ was influencing the SED fit much more than the SPIRE measures, we changed this value to $9.999\text{E}-30$. In this case the

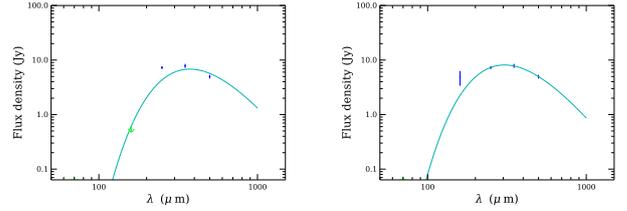


Fig. A.5. Effects on the SED when the measurement at $160 \mu\text{m}$ is used instead of the upper limit for undetected sources (see text). Left panel: best-fit with upper limit; right panel: best-fit with a measure.

fitting code uses the “extracted” $160 \mu\text{m}$ intensity as if it is reliable, but assigns it a very low weight such that this band is not used to constrain the fit. This strategy was already used in a study of Lupus (Benedettini et al. 2018). We show the result of this procedure when applied to source #58 in Fig. A.5.

In the left panel we show the SED and best-fit modified blackbody function of this core as derived by *getsources*: it is clear that the upper limit in the red band forces the best-fit to have a low temperature. The best-fit has $T = 7.74 \pm 0.02 \text{ K}$ and $M = 8.59 \pm 0.45 M_{\odot}$, a large mass indeed. On the other hand, using the total flux density at $160 \mu\text{m}$ that *getsources* measured at the position of the core in the other bands gives the SED shown in the right panel of Fig. A.5. The $160 \mu\text{m}$ value has a very low weight and so the best-fit is now determined by the SPIRE intensities only and the parameters become $T = 11.39^{+0.04}_{-0.05} \text{ K}$ and $M = 0.330 \pm 0.014 M_{\odot}$, much more reasonable values for Perseus.

We used this method to “correct” the SED of 139 unbound cores with reliable SED fitting, 51 prestellar, and 5 likely bound sources. A much more simple approach would be to completely ignore the upper limit, but we find that in some cases the best-fit model agrees well with the $160 \mu\text{m}$ measurement even if it is not relevant during the fitting procedure. This means that in this case *getsources* is able to produce a reasonable measurement even without directly detecting the source. The reader will recognise the ‘corrected’ sources because in these cases the ellipse showing the size of the core (see Figs. A.1 and A.2) is drawn in red instead of green.

In the following section on the completeness limit, we provide a precise assessment of the validity of these two approaches to the problem of the $160 \mu\text{m}$ upper limit, because in that case we deal with simulated sources for which we know the true values of the physical parameters.

The master catalogue listing the observed properties of all the cores is available at the CDS. A template is provided in Table A.1 to illustrate its form and content which follows Könyves et al. (2015). There are cases in which two *Spitzer* sources can be associated with one *Herschel* protostar. In this case we chose the *Spitzer* source detected at $70 \mu\text{m}$, exploiting the fact that in no case both sources had detections at $70 \mu\text{m}$. If none was revealed in this band, we associated the source with the higher $24 \mu\text{m}$ detection.

The derived properties (physical radius, mass, SED dust temperature, peak column density at the resolution of the $500 \mu\text{m}$ data, average column density, peak volume density, and average density) are given in another table also available at the CDS. Table A.2 gives an example. We note that for protostars the derived properties refer to the external envelope. A more detailed SED fitting will be the object of a forthcoming paper.

In a third table we provide additional data for each source; see the sample in Table A.3. Columns (2) and (3) give the source luminosity in two ways. First, for a modified blackbody

Table A.1. Catalogue of dense cores identified in the HGBS maps of Perseus molecular complex.

No.	Name	RA ₂₀₀₀ (h m s)	Dec ₂₀₀₀ (° ′ ″)	Sig ₀₇₀	$S_{070}^{\text{peak}}/S_{\text{bg}}$	$S_{070}^{\text{cont},500}$ (Jy beam ⁻¹)	$S_{070}^{\text{peak}}/S_{\text{bg}}$	$S_{070}^{\text{cont},500}$ (Jy beam ⁻¹)	S_{070}^{int} (Jy)	r_{070}^{int} (″)	r_{070}^{peak} (″)	$r_{070}^{\text{cont},500}$ (″)	PA ₀₇₀ (°)
(1)	(2)	(3)	(4)	(5)	(8)	(6) ± (7)	(8)	(9)	(10) ± (11)	(12)	(13)	(14)	(14)
49	032510.2+302355	03:25:10.28	+30:23:55.4	0.000	0.88	2.93e-03	1.9e-03	3.19e-03	2.14e-01	1.4e-01	8	8	90
55	032522.3+304512	03:25:22.31	+30:45:12.0	735.700	227.14	1.48e+01	2.2e-02	1.61e+01	2.42e+01	3.7e-02	8	8	49
60	032529.2+303108	03:25:29.27	+30:31:08.4	0.000	1.04	5.70e-03	3.1e-03	2.13e-02	-7.21e-02	3.9e-02	25	12	38

Sig ₁₆₀	S_{160}^{peak} (Jy beam ⁻¹)	$S_{160}^{\text{cont},500}$ (Jy beam ⁻¹)	$S_{160}^{\text{peak}}/S_{\text{bg}}$	$S_{160}^{\text{cont},500}$ (Jy beam ⁻¹)	Sig ₁₆₀	PA ₁₆₀ (°)	r_{160}^{peak} (″)	$r_{160}^{\text{cont},500}$ (″)	S_{160}^{int} (Jy)	r_{160}^{int} (″)	$S_{160}^{\text{peak}}/S_{\text{bg}}$	$S_{160}^{\text{cont},500}$ (Jy beam ⁻¹)	S_{160}^{int} (Jy)	r_{160}^{int} (″)	PA ₁₆₀ (°)	
(15)	(16) ± (17)	(19)	(18)	(20) ± (21)	(22)	(23)	(22)	(23)	(20) ± (21)	(22)	(25)	(26) ± (27)	(28)	(30) ± (31)	(34)	
6.897	2.02e-01	4.80e-01	0.52	3.70e-01	7.7e-02	25	20	53	23.880	8.09e-01	4.8e-02	1.01	8.93e-01	1.59e+00	9.4e-02	21
1673	5.44e+01	1.5e-01	43.22	6.20e+01	1.7e-01	14	14	0	1426	3.20e+01	6.9e-01	7.22	3.16e+01	3.36e+01	7.2e-01	19
0.000	3.82e-02	3.2e-02	0.11	8.45e-02	7.1e-02	35	30	-89	6.138	1.64e-01	3.5e-02	0.31	2.17e-01	2.55e-01	5.5e-02	30

Sig ₃₅₀	S_{350}^{peak} (Jy beam ⁻¹)	$S_{350}^{\text{cont},500}$ (Jy beam ⁻¹)	$S_{350}^{\text{peak}}/S_{\text{bg}}$	$S_{350}^{\text{cont},500}$ (Jy beam ⁻¹)	Sig ₃₅₀	PA ₃₅₀ (°)	r_{350}^{peak} (″)	$r_{350}^{\text{cont},500}$ (″)	S_{350}^{int} (Jy)	r_{350}^{int} (″)	$S_{350}^{\text{peak}}/S_{\text{bg}}$	$S_{350}^{\text{cont},500}$ (Jy beam ⁻¹)	S_{350}^{int} (Jy)	r_{350}^{int} (″)	PA ₃₅₀ (°)
(35)	(36) ± (37)	(39)	(38)	(40) ± (41)	(42)	(43)	(42)	(43)	(40) ± (41)	(42)	(44)	(46) ± (47)	(48)	(49) ± (50)	(53)
54.300	1.49e+00	7.1e-02	1.50	1.56e+00	9.6e-02	28	25	-32	66.400	1.67e+00	7.9e-02	1.66	1.95e+00	9.1e-02	47
849.200	2.02e+01	4.8e-01	4.87	2.00e+01	5.4e-01	25	25	-10	407.900	1.17e+01	6.4e-01	4.82	1.26e+01	6.9e-01	36
6.427	1.05e-01	3.4e-02	0.19	1.29e-01	4.8e-02	43	25	78	0.000	6.68e-02	2.5e-02	0.11	4.95e-02	1.8e-02	36

Sig ₆₇₅	$N_{\text{H}_2}^{\text{peak}}$ (10 ²¹ cm ⁻²)	$N_{\text{H}_2}^{\text{cont},500}$ (10 ²¹ cm ⁻²)	$N_{\text{H}_2}^{\text{peak}}/N_{\text{bg}}$	$N_{\text{H}_2}^{\text{cont},500}$ (10 ²¹ cm ⁻²)	$r_{\text{H}_2}^{\text{peak}}$ (″)	PA _{H₂} (°)	$r_{\text{H}_2}^{\text{cont},500}$ (″)	$r_{\text{H}_2}^{\text{int}}$ (″)	$N_{\text{H}_2}^{\text{int}}$ (10 ²¹ cm ⁻²)	Core type	CUTEx	N_{SED}	SIMBAD	Spitzer	Comments	
(54)	(55)	(57)	(56)	(58)	(59)	(60)	(60)	(59)	(58)	(61)	(62)	(63)	(64)	(65)	(66)	(67)
166.200	14.190	4.591	3.32	4.272	24	20	-41	4	4	1	prestellar	[KJT2007] SMM1032516+30238				
883.600	64.580	16.770	2.80	23.080	19	18	-89	5	5	1	protostellar			J032522.32+304513.9		
6.689	0.683	0.256	0.53	1.284	35	18	76	2	2	1	starless				no SED fit	

Notes. Columns: (1) core running number; (2) name = HGBS_J followed by J2000 source coordinates in sexagesimal format; (3) and (4) J2000 RA and Dec of source centre; (5), (15), (25), (35) and (45): detection significance from monochromatic single scales in the *Herschel* bands (0.0 when the core is not visible in clean single scales); (6) and (7), (16) and (17), (26) and (27), (36) and (37), (46) and (47): peak flux density and its uncertainty estimated by *getsources*; (8), (18), (28), (38) and (48): contrast over the local background, defined as the ratio of the background-subtracted peak intensity to the local background intensity; (9), (19), (29) and (39): peak flux density smoothed to a 36′′ beam; (10) and (11), (20) and (21), (30) and (31), (40) and (41), (49) and (50): integrated flux density and its uncertainty estimated by *getsources*; (12)-(13), (22)-(23), (32)-(33), (42)-(43) and (51)-(52): major and minor FWHM of the core, respectively, estimated by *getsources*; (14), (24), (34), (44) and (53): position angle, measured east of north, of the core major axis; (54): detection significance in the high-resolution column density map; (55) peak H₂ column density as estimated by *getsources* in the high-resolution column density map; (56) column density contrast over the local background estimated by *getsources* in the high-resolution column density map; (57) peak column density smoothed to a 36′′ beam; (58) local background H₂ column density estimated by *getsources* in the high-resolution column density map; (59)-(60)-(61): major & minor FWHM and position angle of the core, respectively, estimated by *getsources* in the high-resolution column density map; (62): number of *Herschel* bands in which the core is significant (Sig_i > 5) and has a positive flux density, excluding the column density plane; (63): CUTEx (Molinari et al. 2011) counterpart I = counterpart found with distance between the two peaks less than 6′′; 2 = counterpart found within the ellipse defined by the FWHMs estimated on the 250 μm map, with no constraint on the distance between the two peaks; 0 = no counterpart found by CUTEx; (64): core classification, starless means unbound core; (65) closest counterpart found in SIMBAD, if any, up to 6′′ from the *Herschel* peak position (the identifier is copied as is in the SIMBAD archive); (66) closest *Spitzer*-identified YSO from the c2d survey (Evans et al. 2009) within 6′′ from the *Herschel* peak position, if any. When present, the leading part of the name (SSTc2d) has been removed; (67) comments. Full catalog is at CDS.

Table A.2. Catalogue of dense cores identified in the HGBS maps of the Perseus molecular complex.

No.	Name	RA ₂₀₀₀	Dec ₂₀₀₀	R_{core}	M_{core}	T_{dust}	$N_{\text{H}_2}^{\text{peak}}$	$N_{\text{H}_2}^{\text{ave}}$	$n_{\text{H}_2}^{\text{peak}}$	$n_{\text{H}_2}^{\text{ave}}$	α_{BE}	Core type	Comments					
(1)	(2)	(h m s)	(° ' ")	(pc)	(M_{\odot})	(K)	(10^{21} cm^{-2})	(10^{21} cm^{-2})	(10^4 cm^{-3})	(10^4 cm^{-3})	(17)	(18)	(19)					
49	032510.2+302355	03:25:10.28	+30:23:55.4	1.75e-02	3.17e-02	3.611	0.120	7.3	0.1	36.333	168.27	51.12	20.949	234.31	39.24	0.058	prestellar	
55	032522.3+304512	03:25:22.31	+30:45:12.0	8.87e-03	2.70e-02	1.105	0.028	18.1	0.1	33.650	199.34	21.45	19.402	546.05	19.27	0.239	protostellar	
60	032529.2+303108	03:25:29.27	+30:31:08.4	2.56e-02	3.68e-02	0.047	0.025	10.4	1.0	0.262	1.01	0.49	0.151	0.96	0.32	9.421	starless	no SED fit

Notes. Columns: (1) core running number; (2) name = HGBS_J followed by J2000 source coordinates in sexagesimal format; (3) and (4) J2000 RA and Dec of source centre; (5) and (6) geometrical average of the two FWHMs measured in the high-resolution column density map, after and before deconvolution with HPBW of 18''2, respectively (NB: both values provide estimates of the object's outer radius when the core can be approximately described by a Gaussian distribution, as is the case for the critical Bonnor-Ebert spheroid); (7) and (8) estimated core mass; (9) and (10) SED dust temperature; (11) peak H_2 column density, at the resolution of 500 μm intensity map, derived from a modified blackbody SED fit to the core peak flux densities measured in a common 36'3 beam at all wavelengths; (12) average column density, calculated as $M_{\text{core}}/(\mu m_{\text{H}} \pi R_{\text{core}}^2)$, where M_{core} is the estimated core mass (Col. 7), R_{core} the estimated core radius after deconvolution (Col. 5), $\mu = 2.8$; (13) average column density, calculated is the same way as for Col. (12), but using the observed radius (Col. 6); (14) beam-averaged peak volume density at the resolution of 500 μm intensity map, derived from the peak column density (Col. 11) assuming a Gaussian spherical distribution: $n_{\text{H}_2}^{\text{peak}} = \sqrt{\frac{4 \ln 2}{\pi}} \frac{N_{\text{H}_2}^{\text{peak}}}{\text{FWHM}_{500}}$; (15) average volume density, calculated as $n_{\text{H}_2}^{\text{ave}} = \frac{M_{\text{core}}}{4/3 \pi R_{\text{core}}^3 \mu m_{\text{H}}}$ using the estimated core radius after deconvolution; (16) average volume density calculated as previous column but using the observed core radius; (17) Bonnor-Ebert mass ratio $\alpha_{\text{BE}} = M_{\text{BE,crit}}/M_{\text{obs}}$; (18) core type: starless, prestellar or protostellar; (19) comments. Full catalog is at CDS.

Table A.3. Catalogue of dense cores identified in the HGBS maps of the Perseus molecular complex.

No.	$L_{\text{bol,ana}}$	L_{SED}	λ_0	λ_*	Core	SED	Region
(1)	(L_{\odot})	(L_{\odot})	(μm)	(μm)	type	fit	(8)
49	0.052	0.049	82.307	131	1	1	1
55	3.722	4.152	89.583	53	2	1	2
60	0.000	0.002	0.000	0	0	0	2

Notes. Columns: (1) core running number; (2) bolometric luminosity found from Eq. (A.2); (3) bolometric luminosity found by numerical integration of the SED; (4) λ_0 as given by Eq. (A.1); (5) λ_* see text; (6) core type: 0=starless; 1=prestellar; 2=protostellar; 3=candidate bound; (7) SED fit: 0=no reliable SED fit; 1=reliable SED fit; (8) region: 1=L1451; 2=L1448; 3=L1455; 4=NGC1333; 5=Perseus6; 6=B1; 7=B1E; 8=L1468; 9=IC348; 10=B5; 11=HPZ1; 12=HPZ2; 13=HPZ3; 14=HPZ4; 15=HPZ5; 16=HPZ6; 0=outside all the regions. Full catalog is at CDS.

the bolometric luminosity L_{bol} , can be computed analytically as shown by Elia & Pezzuto (2016). For $\beta = 2$,

$$L_{\text{bol}} = \frac{30}{63} \pi R^2 \left(\frac{4\pi k}{h\nu_0} \right)^2 \sigma T^6,$$

where σ is the Stefan-Boltzmann constant, k and h are the Boltzmann and Planck constants, T is the dust temperature, and $\nu_0 = c/\lambda_0$ where λ_0 is the wavelength above which the model becomes optically thin. For a given dust opacity parametrised with a power law of parameters ($\lambda_{\text{ref}}, \kappa_{\text{ref}}, \beta$), λ_0 is (Pezzuto et al. 2012)

$$\lambda_0 = \lambda_{\text{ref}} \left(\kappa_{\text{ref}} \frac{M}{\pi R^2} \right)^{1/\beta} = 0.756 \left(\frac{R}{1 \text{ pc}} \right)^{-1} \sqrt{\frac{M}{1 M_{\odot}}} \mu\text{m}, \quad (\text{A.1})$$

with R and M being the radius and mass of the (spherical) source. The bolometric luminosity then becomes

$$L_{\text{bol}} = \frac{30}{63} \kappa_{\text{ref}} M \left(\frac{4\pi k \lambda_{\text{ref}}}{ch} \right)^2 \sigma T^6 = 9.45 \times 10^{-8} \left(\frac{M}{1 M_{\odot}} \right) T^6 L_{\odot}. \quad (\text{A.2})$$

Equation (A.2) implies an integration of the whole SED, even at $\lambda < \lambda_0$ where the modified blackbody becomes optically thick. However, Elia & Pezzuto (2016) showed that the SED does not necessarily have to be optically thin at all wavelengths for Eq. (A.2) to be valid, and that it is enough that optical depth is less than 1 for $\lambda = \lambda_*$ where λ_* is defined by the condition

$$\int_0^{\infty} \frac{x^{3+\beta}}{e^x - 1} dx \approx \int_0^{x_*} \frac{x^{3+\beta}}{e^x - 1} dx, \quad (\text{A.3})$$

with $x_* = hc/k\lambda_*T$. For $\beta = 2$ the integral becomes negligible (< 0.01) at $x_* = 15$. This implies that the dusty envelope should already be optically thin at $\lambda_* = 960/T \mu\text{m}$.

Columns (4) and (5) in Table A.3 report λ_0 and λ_* for each model, while Col. (3) gives the luminosity obtained by numerical integration of the SED at *Herschel* wavelengths. For the first source in the table, the model becomes optically thin at 82 μm , while λ_* is 131 μm . In other words, the analytical formula for L_{bol} is efficient as long as the model is optically thin at 131 μm which is correct given that $\lambda_0 = 82 \mu\text{m}$.

The second source is a protostar. In this case L_{bol} cannot be computed with Eq. (A.2), meaning that Col. (2) is not correct and one can use Col. (3) as a measure of the FIR luminosity. The third source does not have a reliable SED fit, meaning that M and T were not derived from the SED. In this case, Cols. (2), (4), and (5) report a zero while Col. (3) gives a more precise FIR luminosity.

Column (6) reports the kind of core: "0" is unbound, "1" is prestellar, "2" is protostellar, and "3" is candidate. Column (7) gives the reliability of the SED fitting: "1" means that the SED could be fit with a modified blackbody, "0" otherwise. Column (8) gives the region to which each source is associated, the numerical correspondence is given in the caption. These last three columns allow us to make selections on the sample of cores easily: for instance, to select prestellar cores in IC348 one can look for sources with (1,1,9) in these columns.

Appendix B: Simulation of prestellar cores

A total of 432 simulated prestellar cores with Bonnor-Ebert density profile and mass between 0.05 M_{\odot} and 0.96 M_{\odot} were

injected into the source-free intensity maps to verify the correctness of our extraction procedure and to derive the mass completeness limit. The simulations were performed as in previous HGBS studies (Könyves et al. 2015; Marsh et al. 2016; Benedettini et al. 2018). The extraction of the simulated sources was performed nominally and the post-selection checks for cores described in Sect. 4.1 were executed. This catalogue was then cross-checked against the truth table of the synthetic cores and 315 sources out of the 432 were recovered.

B.1. Restoring the $160\mu\text{m}$ intensity

As discussed in the previous section, we found some problems when dealing with the $160\mu\text{m}$ upper limit often found in the standard extraction. To prevent the upper limit from forcing the SED fitting to overly low temperatures, we adopted two different strategies previously described and here we discuss their validity.

First, we carried out an alternative extraction using all five of the *Herschel* bands simultaneously. Compared to the nominal extraction, two sources with only upper limits at $160\mu\text{m}$ ended up with intensity measurements in the alternative extraction. For one source, the mass derived from the SED fitting using the upper limit at $160\mu\text{m}$ is 1.46 times the true mass. Using the measurement from the alternative catalogue, the derived mass is 1.36 times the true mass, which is slightly better. For the second source, using the upper limit causes the derived mass to be overestimated by more than a factor of three. Using the measured flux density in the alternative catalogue, the ratio is 0.75, which means that the mass is only underestimated by 25%, which is a better result in any case.

The second approach was to pretend that the upper limit in the nominal extraction is indeed a real measurement but with a negligible weight during the SED fitting so that it does not force the fitting. For 96 sources out of 147 simulated cores with an upper limit, nothing changed because the best-fit model fulfilled the upper limit, or because the SED could not be fit anyway. For the remaining 51 sources, the mean ratio between the SED-fitting derived mass and true mass is 1.5 ± 1.1 when using the upper limit, and 0.92 ± 0.78 when the $160\mu\text{m}$ intensity is weighted down in the SED fit. A comparison between the two means show that the upper limit causes, on average, an overestimation of the mass by 50% while weighing down the $160\mu\text{m}$ intensity ensures an agreement with the input mass within 8% on average. We conclude that the SED fitting improves in both approaches, even if in the second approach the large standard deviation of the means shows that the results are valid for the whole sample and not applicable to single sources.

B.2. Completeness limit in mass

In Fig. B.1 we show the fraction of detected sources with respect to the number of simulated sources as a function of the true core mass. To improve the statistics we summed the number of detected sources in three adjacent bins with similar masses, while the mass of each bin was the average of the three masses. The last two points are the sum/average of four bins.

The horizontal dashed line shows the 90% level of completeness while the vertical line corresponds to a mass of $0.323 M_{\odot}$, found by linear interpolation of the two bins at $0.290 M_{\odot}$ and $0.420 M_{\odot}$.

The trend at high-mass in Fig. B.1 is also interesting, where the number of detected sources decreases. As shown by Könyves et al. (2015), the likelihood of detecting a sources does not only depend on its intensity, but also on the contrast, i.e. how well

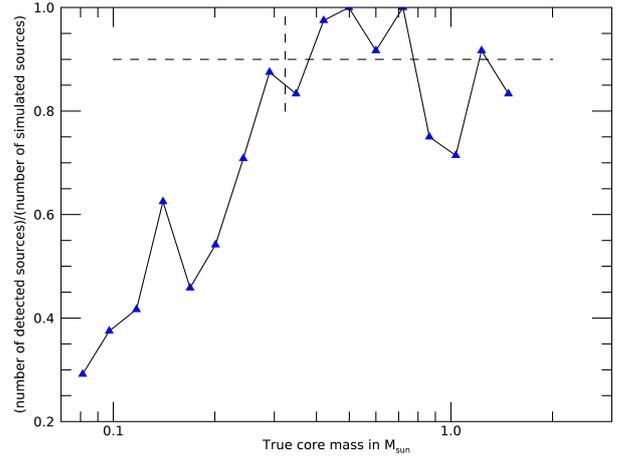


Fig. B.1. Number of detected sources over number of simulated sources vs. true core mass. The completeness limit is defined as the mass at which completeness is 90%: in our case this happens at $0.323 M_{\odot}$, the limit shown as intersection between the two dashed lines. The decrease in the number of detected sources at high mass is discussed in the text (see also next figure).

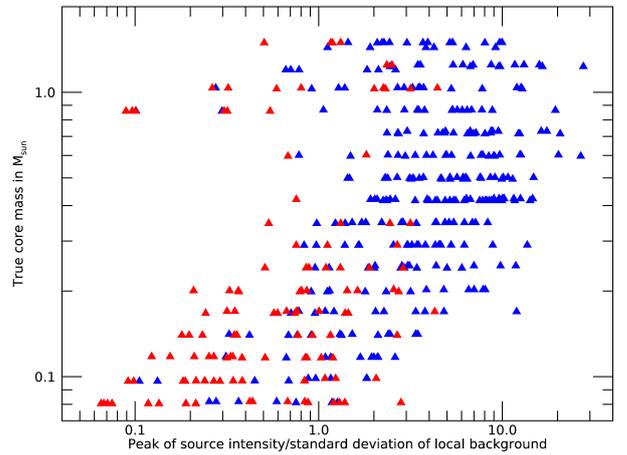


Fig. B.2. Contrast level (see text) for detected sources (blue triangles) and undetected sources (red triangles).

the source stands out above the emission of the diffuse medium. In general, high-mass cores are found in high-column-density regions which are very bright. To estimate the contrast we used the signal-to-noise ratio (S/N) expressed as the peak intensity of the simulated sources over the standard deviation of the local background, both in MJy sr^{-1} . This ratio, for the fiducial wavelength of $250\mu\text{m}$, is shown in Fig. B.2: blue and red triangles refer to detected/undetected sources, respectively. We can see that for $S/N \geq 3$ we detect $\sim 98\%$ of the sources independent on their mass. The percentage is as low as 31% for $S/N \leq 1$. At high mass we explored low levels of S/N that resulted in a small fraction of detected sources. This figure shows that high-mass cores are not necessarily found easier than low-mass cores. The interested reader is invited to read the in-depth discussion on this topic in Könyves et al. (2015).

Appendix C: Definition of Perseus subregion borders through ^{13}CO observations

The Perseus molecular complex is known to host different regions with distances varying between ≥ 200 pc and ≤ 350 pc.

From our intensity maps, because the dust emission is optically thin, it is not possible to estimate which regions are closer or further. Also, their borders are arbitrarily defined.

To obtain some insight into these problems, especially for tracing the borders, we combined our column density map (Sect. 3.2) with the ^{13}CO map of Perseus (COMPLETE team 2012). In particular, we started from the $N(\text{H}_2)$ contour at $3 \times 10^{21} \text{ cm}^{-2}$ and modified it in order to maximise one single velocity component in each region wherever possible. The result of this exercise is visualised in Figs. C.1–C.2, where we show how the borders have been defined to accommodate the observed velocity components, and in Fig. C.3 which shows the ^{13}CO spectra within these borders. The regions are shown over the column density map in Fig. 9.

From these figures it is clear that only in a few cases is a region associated to one single ^{13}CO component; in general, two or even three components are present, often in the form of broad bump. The most complex case is NGC 1333 which appears bright from 6 to 10 km s^{-1} . It is not possible to assess whether these different velocity components are associated to different distances or simply reflect the presence of complex motions of gas at the same distance.

Border coordinates are reported here in DS9 format.

```
# Region file format: DS9 version 4.1
global color=green dashlist=8 3 width=1 font=
"helvetica 10 normal roman" select=1 highlite=1 dash=0
fixed=0 edit=1 move=1 delete=1 include=1 source=1
image
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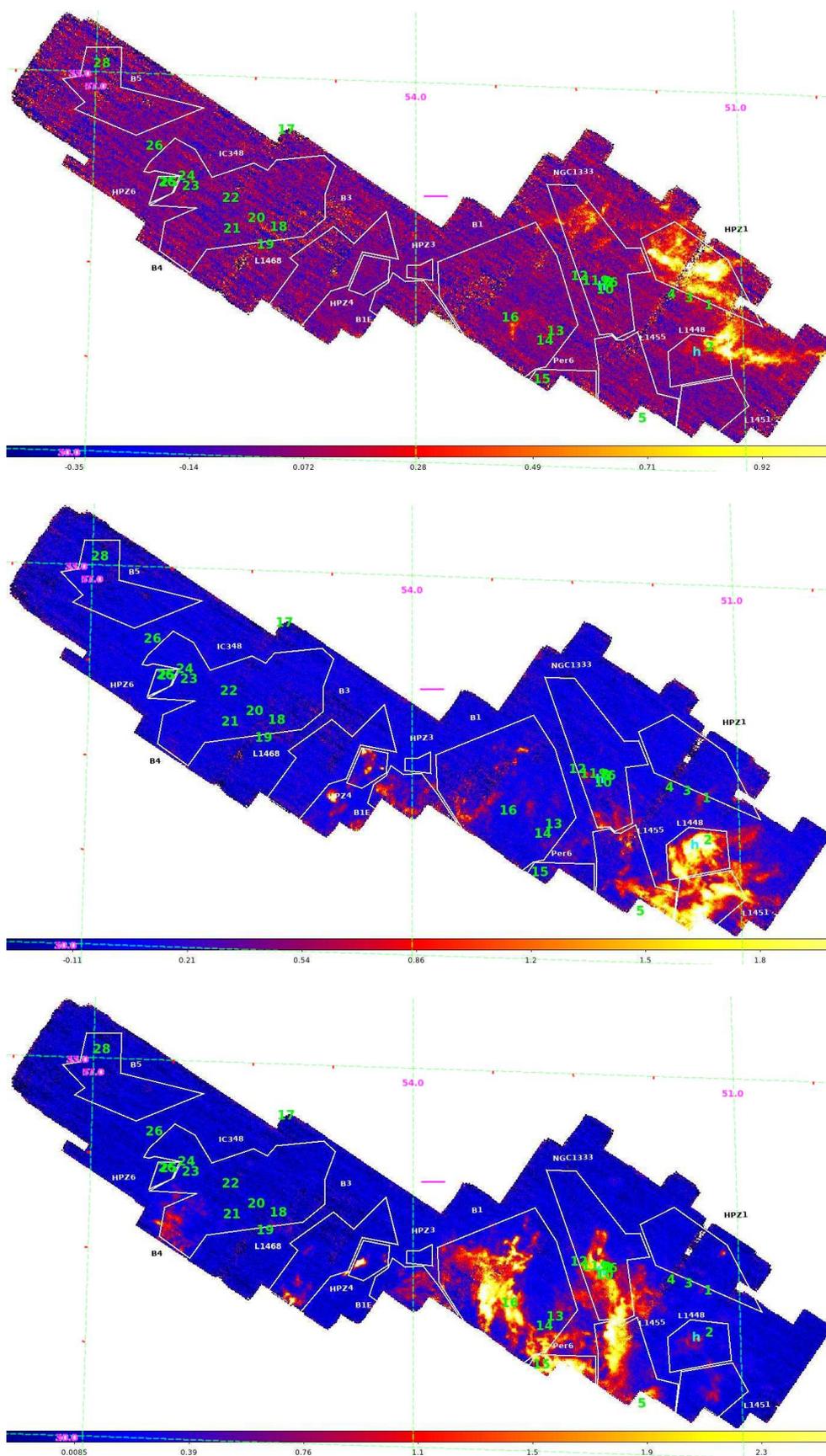


Fig. C.1. Maps of ^{13}CO in three velocity channels: 1.1 Km s^{-1} (top), 4.35 Km s^{-1} (centre), 6.0 Km s^{-1} (bottom); colour bars are uncorrected antenna temperature in K. Green labels mark the position of *Gaia* sources (see Table 2); *h* labels show the position of Hirota et al. (2008) and Hirota et al. (2011) sources (SVS13 is in the cloud of sources 6-10). The magenta lines at the centre of each map show the angular scale corresponding to 1 pc at 300 pc; J2000.0 coordinates grid is shown.

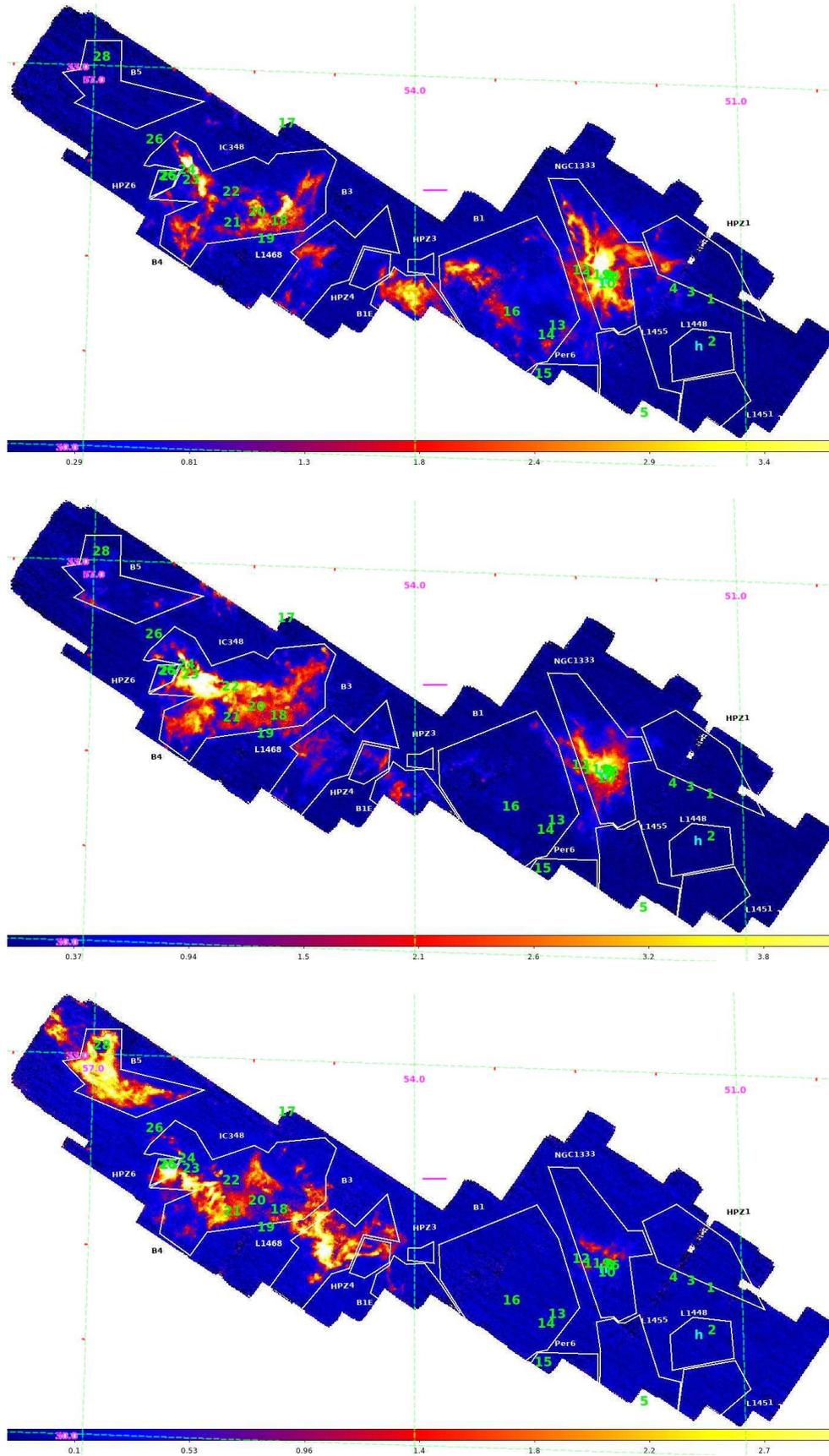


Fig. C.2. Maps of ^{13}CO in three velocity channels: 7.9 km s^{-1} (top), 8.8 km s^{-1} (centre), 10.0 km s^{-1} (bottom). In the latter panel, north of B5 and slightly covered by label 28, the bubble CPS12 (Arce et al. 2011) is evident. The magenta lines at the centre of each map show the angular scale corresponding to 1 pc at 300 pc; J2000.0 coordinates grid is shown. Labels as in Fig. C.1.

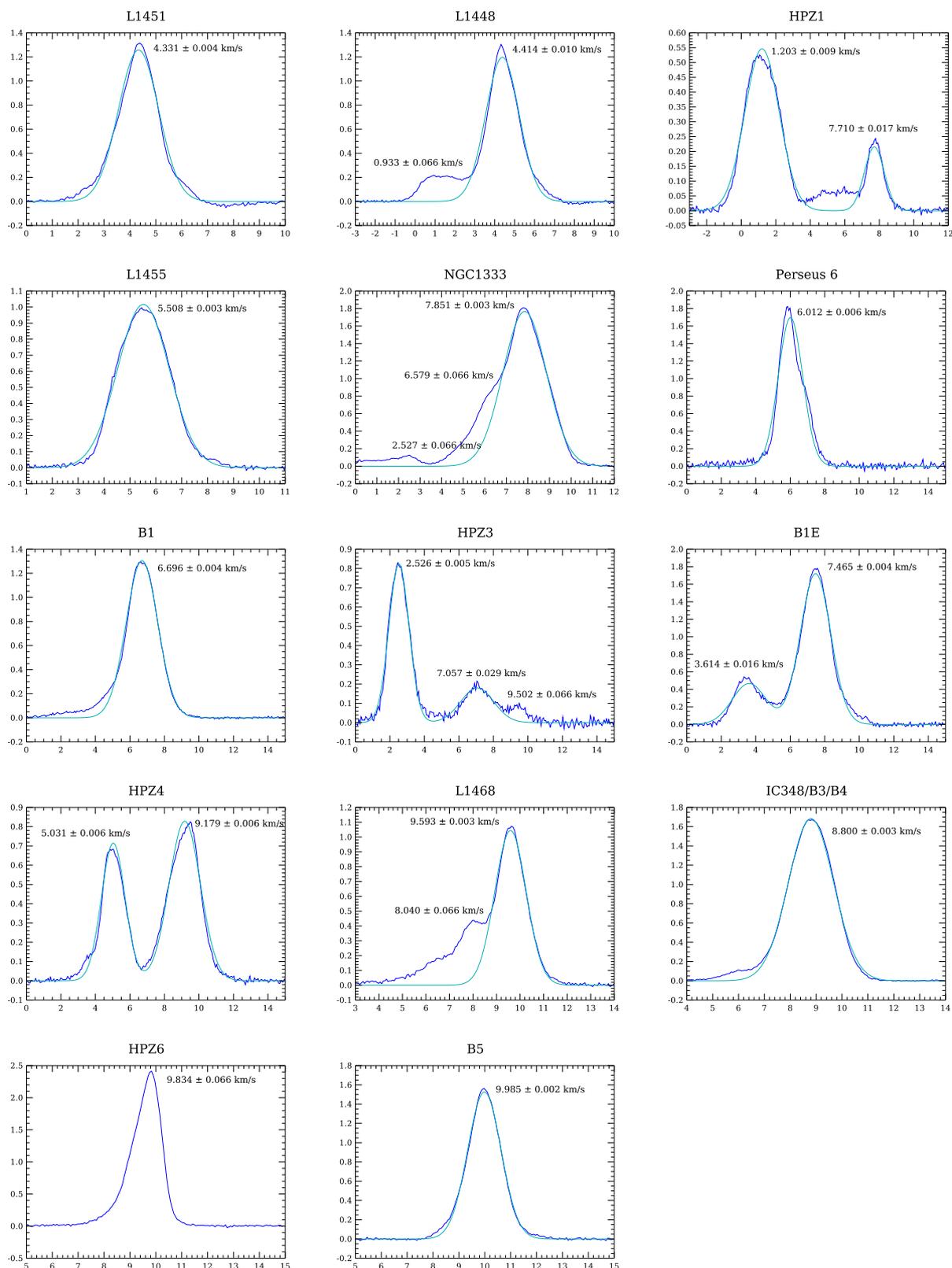


Fig. C.3. ^{13}CO spectra for all the subregions we defined in Perseus. Whenever possible we made a Gaussian fit to one or two profiles; for the broad bumps we simply derived the position of the maximum and assigned it an uncertainty of 0.066 km s^{-1} , which is equal to the step in the velocity grid.

Appendix D: *Herschel* intensity maps

In this Appendix we show the PACS intensity maps at $70\ \mu\text{m}$ and $160\ \mu\text{m}$ in Fig. D.1; SPIRE intensity map at $250\ \mu\text{m}$ and the high-resolution column density map with a resolution of $18''.2$ in Fig. D.2 (the column density map shown in Fig. 4 is at $36''.1$); and SPIRE intensity maps at $350\ \mu\text{m}$ and $500\ \mu\text{m}$ in Fig. D.3. To help to localise the region of Perseus in other surveys, we provide a grid of coordinates in different systems: FK5 (equatorial) for $70\ \mu\text{m}$ and $250\ \mu\text{m}$, Galactic for $160\ \mu\text{m}$, $350\ \mu\text{m}$ and for the high-resolution column density map, ecliptic for $500\ \mu\text{m}$.

Maps centres, as given in FITS header and converted to the specific coordinate system are: ($3^{\text{h}}36^{\text{m}}43^{\text{s}}, +31^{\circ}35'27''$), or ($54.1775, +31.5908$) for blue band; ($159.5082, -19.3010$) for red band; ($3^{\text{h}}35^{\text{m}}06^{\text{s}}, +31^{\circ}31'20''$), or ($53.7732, +31.5221$) for PSW band; ($159.3639, -19.4821$) for the column density map; ($159.4321, -19.5946$) for PMW band; ($59.3449, +11.7751$) for PLW band.

From Figs. D.1–D.3 it is evident that the diffuse emission intensity increases at longer wavelengths, being marginally visible at $\lambda = 70\ \mu\text{m}$.

Table D.1. Zero-levels of the diffuse emission for each map.

λ (μm)	East (MJy sr $^{-1}$)	West (MJy sr $^{-1}$)
70	3.320	-3.492
160	54.39	68.77
250	66.32	45.32
350	31.69	23.06
500	13.35	10.14

The zero-level intensities for the diffuse emission were found through comparison with the predicted Planck+IRAS emission (Bernard et al. 2010) for the same region. For each map and wavelength, the zero-level intensity was found separately for west and east Perseus and the offset added to the map. Finally, the two maps were combined with a simple pixel-by-pixel mean (the two maps were created on the same spatial grid). The offsets are reported in Table D.1.

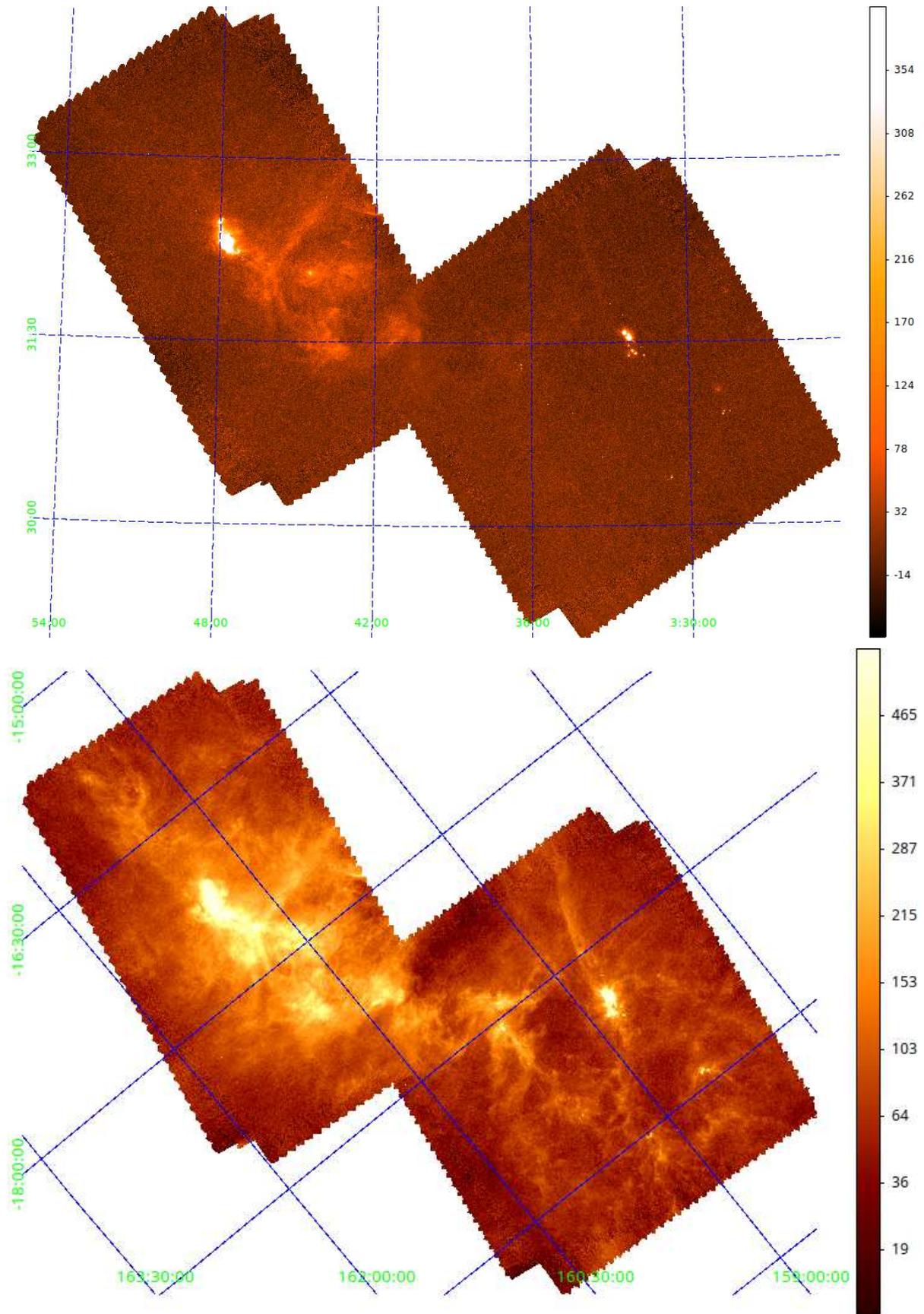


Fig. D.1. Combined maps of Perseus at $70\ \mu\text{m}$ (*top*, grid of equatorial J2000.0 coordinates, HPBW $8''.4$) and $160\ \mu\text{m}$ (*bottom*, grid of Galactic coordinates, HPBW $13''.5$). Colour bars in MJy sr^{-1} .

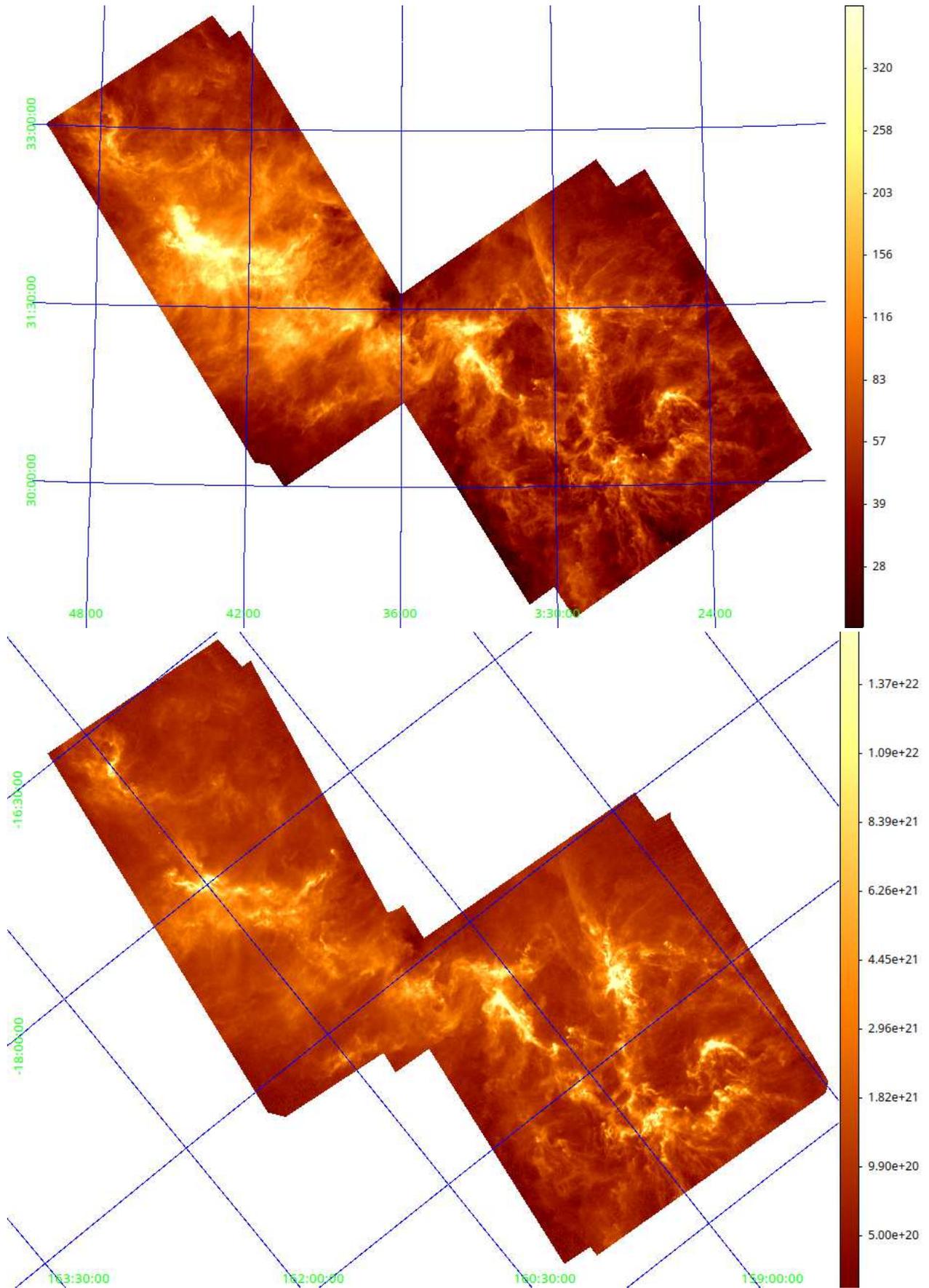


Fig. D.2. SPIRE intensity map of Perseus at $250 \mu\text{m}$ (*top*, grid of equatorial J2000.0 coordinates) and the high-resolution column density map (*bottom*, grid of Galactic coordinates). For both maps the HPBW is $18''.2$. Colour bars are in MJy sr^{-1} (*top*) and cm^{-2} (*bottom*).

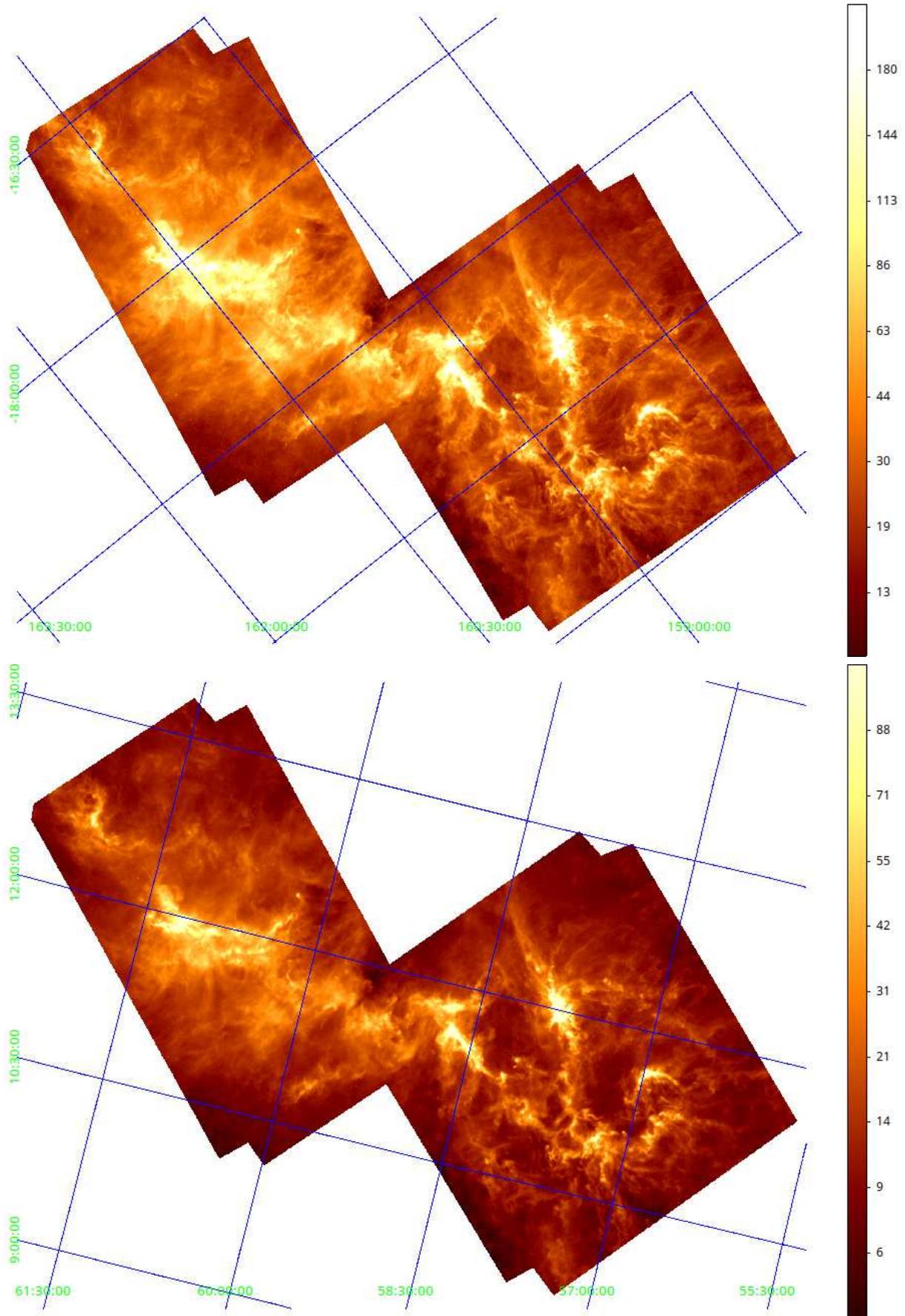


Fig. D.3. SPIRE intensity maps of Perseus at 350 μm (*top*, grid of Galactic coordinates, HPBW 24'9) and 500 μm (*bottom*, grid of ecliptic coordinates, HPBW 36'1). Colourbars in MJy sr^{-1} .

Appendix E: Additional catalogue of sources

Here we report basic data for sources eliminated from the main catalogue; as these data are not part of the official HGBS delivered products they are reported here. The columns in the following tables are a subset of the columns in Table A.1, and the numbering is kept the same as for that table: e.g. the contents of Col. (12) in Table E.1 is the same as Col. (12) in Table A.1. The only difference is units in Col. (55): 10^{21} cm^{-2} in Table A.1 and 10^{20} cm^{-2} here.

The other difference is Col. (2) in this table: here an identifier is given that explains why a source was removed

(acronyms as in Simbad database): G – galaxy; W – counterpart in WISE catalogue; 2 – counterpart in 2MASS; HH – Herbig-Haro object; M – star from catalogue MBO; HL – candidate YSO from [HL2013]; DS – source in [DS95]; LRL – source in “CI* IC 348 LRL”; PSZ – source in [PSZ2003]; NTC – source in “CI* IC 348 NTC”; U – W UMa star; δ – δ Scuti star.

In particular, “W” and “2” means that an infrared sources was found in WISE or 2MASS catalogues within $6''$ of an *Herschel* source not detected at $70 \mu\text{m}$ nor in the millimetre band. No attempt was made to distinguish between physical and projected associations.

Table E.1. Catalogue of additional sources excluded from the main catalogue.

No.	Type	RA ₂₀₀₀ (degree)	Dec ₂₀₀₀ (degree)	Sig ₀₇₀	S_{070}^{peak} (Jy beam ⁻¹)		S_{070}^{tot} (Jy)		r_{070}^a (")	r_{070}^b (")	PA ₀₇₀ (°)	Sig ₁₆₀	S_{160}^{peak} (Jy beam ⁻¹)		S_{160}^{tot} (Jy)		r_{160}^a (")	r_{160}^b (")	PA ₁₆₀ (°)
(1)	(2)	(3)	(4)	(5)	(6) ± (7)	(8) ± (9)	(10) ± (11)	(12)	(13)	(14)	(15)	(16) ± (17)	(18) ± (19)	(20) ± (21)	(22)	(23)	(24)		
1	W	51.03003	30.27615	0.000	2.02e-2	1.3e-2	-1.58e-2	5.6e-2	60.8	47.2	133.5	0.000	1.81e-1	4.1e-2	1.30e+0	1.1e-1	47.9	37.0	149.7
2	W	51.18427	30.27827	0.000	2.75e-2	1.5e-2	4.30e-1	6.6e-2	54.0	47.5	69.3	0.000	6.13e-2	3.2e-2	8.90e-1	8.6e-2	57.0	51.0	52.4
3	W	51.30787	30.76503	0.000	1.53e-2	1.3e-2	-1.42e-1	4.7e-2	46.1	40.1	58.2	6.779	9.46e-2	8.3e-2	5.66e-1	1.2e-1	51.9	18.6	108.4
4	HH	51.30889	30.78195	0.000	1.81e-2	1.4e-2	-7.24e-2	3.4e-2	35.3	23.6	67.6	0.000	5.55e-2	4.4e-2	2.87e-1	7.4e-2	32.3	28.1	57.1
5	W	51.41323	31.32864	0.000	6.48e-3	8.1e-3	-5.62e-2	6.0e-2	91.8	84.2	80.8	7.148	2.09e-1	4.7e-2	3.29e+0	2.0e-1	65.8	45.4	50.7
6	W	51.41401	31.50636	0.000	2.11e-2	1.2e-2	-2.59e-1	4.3e-2	43.4	39.3	123.9	0.000	6.73e-2	3.1e-2	2.81e-1	6.7e-2	40.6	36.1	101.0
7	W	51.42242	30.55880	3.062	-3.14e-2	1.0e-2	1.65e-3	2.0e-2	43.0	11.7	63.1	0.000	2.02e-2	2.9e-2	2.67e-2	6.4e-2	48.0	37.2	175.1
8	W	51.42375	30.66788	4.781	-4.76e-2	9.9e-3	-7.01e-1	3.4e-2	58.7	25.4	150.8	0.000	5.29e-2	3.4e-2	1.13e+0	9.8e-2	57.5	49.6	102.9
9	2	51.42510	30.26197	4.096	-4.36e-2	1.0e-2	-3.01e-2	1.8e-2	35.1	9.2	92.7	0.000	5.43e-2	3.7e-2	1.89e-1	6.5e-2	21.9	20.7	44.0
10	W	51.43138	31.31290	0.000	2.35e-2	1.3e-2	-8.08e-2	4.0e-2	38.6	32.9	90.3	0.000	5.98e-2	3.6e-2	3.63e-1	7.0e-2	39.7	33.2	85.2
11	W	51.44525	30.73913	0.000	2.79e-2	1.9e-2	-3.47e-1	8.6e-2	56.7	51.3	74.4	6.322	5.49e-2	7.4e-2	1.92e-1	1.2e-1	57.8	14.2	147.7
12	W	51.51405	30.63757	0.000	1.58e-2	1.2e-2	-2.86e-1	4.4e-2	45.7	41.5	89.3	0.000	1.27e-1	4.7e-2	5.51e-1	8.8e-2	42.3	28.2	0.5
13	W	51.55092	29.93898									0.000	6.21e-2	3.5e-2	1.54e-1	1.4e-1	71.3	67.0	71.2
14	2	51.55744	30.48739	3.231	-2.90e-2	8.9e-3	2.38e-2	1.7e-2	43.8	10.3	13.5	0.000	9.52e-2	3.4e-2	2.52e-1	6.5e-2	44.2	30.6	178.1
15	W	51.55879	31.04921	8.126	-9.46e-3	1.1e-3	2.00e-2	2.5e-3	42.9	8.4	79.6	0.000	1.31e-1	3.3e-2	1.13e+0	8.5e-2	54.3	40.1	175.8
16	U	51.58218	31.10986	7.527	1.76e-1	1.9e-2	1.30e-2	7.0e-2	28.4	15.2	145.0	0.000	1.67e-1	3.4e-2	8.40e-1	9.8e-2	24.6	22.0	120.2
17	G	51.59041	29.97570	8.853	7.83e-2	1.9e-2	1.19e-2	7.6e-2	57.2	33.3	103.7	11.61	3.65e-1	4.1e-2	5.29e-1	1.1e-1	15.9	13.5	84.2
18	W	51.61117	30.58375	0.000	1.63e-2	1.1e-2	-4.89e-1	1.2e-1	132.3	123.3	72.6	0.000	9.06e-2	3.7e-2	5.76e+0	2.3e-1	133.3	108.7	178.6
19	W	51.64251	31.47244	4.392	-1.43e-2	3.2e-3	6.45e-2	1.6e-2	78.0	37.1	131.1	0.000	1.21e-1	4.4e-2	2.66e+0	1.8e-1	68.8	62.9	18.4
20	G	51.68151	31.39189	6.808	1.02e-1	1.8e-2	2.74e-1	3.0e-2	16.8	16.3	5.6	15.54	4.48e-1	3.8e-2	6.01e-1	4.3e-2	14.2	13.5	157.1
21	G	51.72476	31.82794	6.785	9.09e-2	2.2e-2	1.71e-1	3.0e-2	23.2	8.4	57.2	14.70	5.52e-1	3.0e-2	4.99e-1	3.5e-2	13.5	13.5	74.3
22	W	51.82867	30.28879	0.000	4.30e-2	1.9e-2	1.96e-1	3.7e-2	27.1	22.0	132.2	0.840	-2.21e-2	2.6e-2	1.11e-1	3.0e-2	30.1	13.5	110.5
23	G	51.83795	30.20311	0.000	3.56e-2	1.8e-2	4.82e-2	4.5e-2	33.9	27.4	175.1	6.947	1.47e-1	4.7e-2	4.18e-1	6.1e-2	29.6	15.8	77.5
24	W	51.84337	30.09334	0.000	1.90e-3	7.8e-3	2.80e-1	1.5e-2	32.8	11.9	29.8	0.000	9.78e-2	3.4e-2	4.72e-1	6.6e-2	32.3	30.3	126.9
25	W	51.88653	30.15626	0.000	5.44e-3	7.6e-3	1.46e-1	2.0e-2	43.4	30.1	107.5	0.000	6.93e-2	2.6e-2	1.23e-1	5.0e-2	39.2	33.8	29.7
26	W	51.91415	30.09052	0.000	4.92e-2	2.1e-2	-4.58e-1	8.7e-2	50.8	46.6	124.6	0.000	8.01e-2	3.3e-2	1.19e+0	7.1e-2	57.8	38.1	172.3
27	W	51.92840	31.49610	2.335	-3.31e-2	1.4e-2	6.68e-2	2.5e-2	44.6	11.3	97.8	0.000	7.98e-2	3.7e-2	4.67e-1	7.1e-2	36.4	33.0	118.6
28	W	51.96375	31.46276	4.345	-9.43e-3	2.1e-3	-4.26e-3	9.7e-3	48.4	36.0	103.8	0.000	1.07e-1	4.5e-2	1.27e+0	1.3e-1	61.5	45.6	148.6
29	W	51.99873	29.97346	8.369	-1.64e-2	1.9e-3	2.42e-1	8.6e-3	68.1	39.2	21.1	0.000	1.30e-1	4.1e-2	1.63e+0	1.1e-1	71.8	47.5	8.2
30	W	52.04854	30.65368	5.072	-1.68e-2	3.3e-3	2.66e-1	2.5e-2	122.5	60.0	6.3	0.000	1.24e-1	4.3e-2	5.65e+0	2.6e-1	114.3	92.6	165.9
31	G	52.06091	30.49836	17.50	3.51e-1	2.3e-2	6.30e-1	3.2e-2	12.1	8.4	35.8	39.13	1.24e+0	3.4e-2	1.41e+0	3.5e-2	14.9	13.5	51.7
32	W	52.07263	31.08168	2.872	-1.80e-2	6.2e-3	8.21e-3	1.7e-2	51.1	20.1	112.1	0.000	6.90e-2	3.7e-2	1.28e+0	9.8e-2	56.7	46.9	177.8
33	W	52.08719	29.66298									0.000	8.68e-2	3.4e-2	9.63e-1	8.5e-2	58.5	42.6	45.9
34	W	52.13651	31.01901	4.340	-8.16e-2	1.8e-2	-1.58e-1	3.1e-02	37.8	10.3	21.4	0.000	2.86e-1	4.3e-2	1.10e+0	6.6e-2	25.1	20.0	177.9
35	W	52.13744	31.18561	6.046	4.55e-2	1.7e-2	4.52e-1	4.3e-02	33.8	23.5	101.0	29.41	6.40e-1	5.5e-2	2.41e+0	9.1e-2	26.0	19.9	83.5
36	W	52.14268	31.26396	0.000	1.39e-2	8.0e-3	6.37e-2	1.4e-02	30.4	20.0	30.9	8.477	1.68e-1	9.7e-2	3.51e-1	1.1e-1	28.8	14.1	110.4
37	2	52.14460	31.11770	6.208	1.01e-1	2.2e-2	1.40e-1	3.4e-02	22.2	9.0	65.7	26.54	7.97e-1	6.3e-2	1.18e+0	9.2e-2	13.6	13.5	26.5
38	2	52.15573	31.29497	0.000	4.19e-2	2.2e-2	3.34e-1	4.0e-02	24.3	20.4	106.7	16.93	4.58e-1	9.1e-2	1.26e+0	1.0e-1	24.1	13.5	102.7
39	2	52.15975	31.44214	2.277	-1.68e-2	7.3e-3	-1.73e-1	2.3e-02	48.7	21.5	83.1	0.000	9.03e-2	5.3e-2	1.31e+0	1.7e-1	62.7	53.2	74.2
40	W	52.16198	31.09982	0.000	8.76e-2	2.1e-2	-2.31e-1	9.5e-02	68.7	47.4	142.0	12.06	4.28e-1	6.9e-2	1.11e+0	1.8e-1	23.4	14.9	112.9
41	W	52.16796	31.29740	4.120	-1.32e-2	3.2e-3	-9.47e-2	8.3e-03	65.0	10.3	170.1	21.53	3.21e-1	1.1e-1	3.80e+0	2.7e-1	60.3	27.8	84.3
42	W	52.17626	31.06439	0.000	1.56e-2	7.6e-3	1.84e-1	2.0e-02	35.7	28.0	141.3	0.000	9.08e-2	5.7e-2	9.35e-2	9.7e-2	36.9	31.0	175.9
43	W	52.20424	31.26801	0.000	2.22e-2	1.6e-2	1.69e-1	4.8e-02	45.4	34.1	172.4	9.987	2.89e-1	1.1e-1	1.31e+0	2.0e-1	32.8	23.9	50.8
44	W	52.20630	31.49733	4.718	-1.73e-2	3.6e-3	-1.84e-1	1.2e-02	50.5	24.0	109.1	6.075	8.84e-2	5.1e-2	9.91e-1	1.2e-1	50.6	33.8	57.7
45	W	52.22218	31.56480	19.38	-1.44e-2	7.4e-4	6.70e-4	1.4e-03	32.1	10.8	129.6	0.000	1.06e-1	3.5e-2	1.43e-1	6.6e-2	33.4	31.9	66.2
46	W	52.22956	31.39457	0.000	3.61e-2	2.1e-2	2.93e-2	4.6e-02	25.4	24.5	131.7	9.366	3.60e-1	8.0e-2	6.60e-1	9.4e-2	25.1	14.5	135.8
47	W	52.23603	31.32068	0.000	1.52e-2	2.0e-2	8.66e-2	8.5e-02	53.8	44.9	163.3	34.99	9.83e-1	1.5e-1	4.04e+0	2.6e-1	40.4	15.4	20.0
48	2	52.23690	31.34821	0.000	4.07e-2	2.4e-2	-2.07e-1	3.6e-02	34.3	14.8	9.6	0.000	4.95e-3	1.4e-1	-4.32e-1	1.9e-1	13.5	13.5	138.5
49	2	52.23980	31.36661	2.957	-5.17e-2	1.7e-2	-2.26e-1	3.1e-02	24.3	12.4	141.8	17.42	7.40e-1	2.5e-1	2.14e+0	3.1e-1	22.6	15.9	21.3
50	M	52.24869	31.40059	0.470	-9.72e-3	2.0e-2	3.28e-1	7.5e-02	49.0	37.2	40.5	15.93	5.55e-1	1.9e-1	2.26e+0	3.7e-1	33.6	14.8	0.2
51	W	52.27934	31.25992	0.000	1.00e-2	2.5e-2	1.14e+0	4.2e-02	53.0	13.6	11.6	19.52	5.98e-1	3.7e-1	2.34e+0	4.8e-1	30.5	14.9	128.7
52	2	52.28681	31.38513	197.3	1.95e+0	2.9e-1	1.37e+1	7.3e-01	25.3	19.5	9.2	144.4	4.21e+0	1.0e+0	9.10e+0	1.3e+0	22.8	13.7	46.6
53	W	52.29690	30.33289	0.000	4.67e-2	2.0e-2	-2.27e-1	6.4e-02	42.0	32.7	98.0	0.000	8.93e-2	4.0e-2	4.58e-1	8.2e-2	25.8	22.5	101.2
54	W	52.29768	31.31080	195.7															

Table E.1. continued.

No.	Type	RA ₂₀₀₀ (degree)	Dec ₂₀₀₀ (degree)	Sig ₀₇₀	S_{070}^{peak} (Jy beam ⁻¹)	S_{070}^{tot} (Jy)	r_{070}^a ($''$)	r_{070}^b ($''$)	PA ₀₇₀ ($^\circ$)	Sig ₁₆₀	S_{160}^{peak} (Jy beam ⁻¹)	S_{160}^{tot} (Jy)	r_{160}^a ($''$)	r_{160}^b ($''$)	PA ₁₆₀ ($^\circ$)				
(1)	(2)	(3)	(4)	(5)	(6) \pm (7)	(10) \pm (11)	(12)	(13)	(14)	(15)	(16) \pm (17)	(20) \pm (21)	(22)	(23)	(24)				
57	W	52.30389	31.61624	3.026	-4.58e-2	1.5e-2	-6.65e-3	4.0e-2	44.2	19.3	41.0	0.000	7.78e-2	3.3e-2	3.21e-1	8.4e-2	48.7	43.1	2.7
58	2	52.30625	31.39899	23.45	7.17e-2	9.2e-2	2.64e-1	1.8e-1	42.6	9.7	98.8	96.97	2.24e+0	6.6e-1	8.01e+0	1.0e+0	40.6	16.5	78.2
59	W	52.31406	31.44606	3.076	-5.08e-2	1.6e-2	-4.53e-1	4.4e-2	34.7	24.9	81.1	8.201	2.16e-1	9.8e-2	6.01e-1	1.3e-1	35.4	15.3	161.9
60	M	52.31499	31.38011	7.363	1.08e-1	1.1e-1	1.70e-1	1.6e-1	24.6	8.4	156.7	32.09	1.08e+0	5.9e-1	1.75e+0	6.7e-1	22.3	17.5	134.5
61	W	52.32361	31.67231	0.000	2.53e-2	6.7e-3	9.30e-1	4.8e-2	88.8	80.0	52.0	7.074	2.36e-1	4.4e-2	4.65e+0	1.9e-1	65.7	51.6	118.4
62	2	52.33308	31.40231	81.15	1.01e+0	4.4e-1	4.18e+0	5.9e-1	22.0	12.2	2.4	192.9	3.96e+0	1.9e+0	7.36e+0	2.0e+0	21.3	14.2	20.3
63	W	52.35596	31.25907	0.000	1.47e-2	1.1e-2	1.73e-1	4.2e-2	46.2	40.1	159.4	0.000	9.56e-2	4.0e-2	6.68e-1	8.7e-2	48.5	39.4	30.7
64	G	52.36137	32.48733	16.59	3.45e-1	2.3e-2	3.70e-1	3.1e-2	8.4	8.4	41.5	23.76	8.92e-1	3.3e-2	7.04e-1	3.3e-2	13.5	13.5	57.3
65	W	52.38830	31.26762	0.000	5.24e-4	3.4e-3	9.46e-2	1.5e-2	8.4	8.4	0.0	0.000	1.17e-1	4.5e-2	9.36e-1	1.1e-1	46.5	40.9	42.5
66	2	52.40749	31.46313	0.000	5.14e-3	2.6e-3	2.13e-1	5.9e-3	41.4	15.7	36.9	0.000	1.20e-1	5.7e-2	5.77e-1	1.2e-1	35.7	27.6	136.8
67	2	52.45188	30.67810	0.000	1.14e-2	1.1e-2	5.15e-1	4.9e-2	58.3	49.4	58.7	0.000	1.54e-1	4.3e-2	9.88e-1	1.0e-1	50.7	35.4	24.9
68	2	52.47598	31.34729	8.573	1.10e-1	2.2e-2	2.87e-1	2.8e-2	20.0	9.2	115.7	7.763	2.68e-1	3.8e-2	2.80e-1	3.8e-2	18.1	13.5	133.3
69	G	52.53849	32.04833	50.18	1.25e+0	2.4e-2	1.20e+0	3.0e-2	8.4	8.4	9.7	45.16	1.55e+0	4.0e-2	1.34e+0	4.0e-2	13.5	13.5	89.0
70	W	52.55178	31.78570	3.597	-2.02e-2	5.6e-3	2.07e-1	1.3e-2	30.0	24.5	15.0	0.000	1.00e-7	4.3e-2	1.76e-2	6.6e-2	33.9	21.8	55.3
71	W	52.55590	31.73387	0.000	4.41e-2	1.7e-2	6.15e-1	6.8e-2	51.3	43.0	18.4	0.000	2.21e-4	4.5e-2	1.56e-3	1.1e-1	29.7	27.1	111.2
72	W	52.55624	30.38622	0.000	1.18e-2	1.0e-2	3.22e-1	3.8e-2	49.5	42.4	102.8	7.979	2.08e-1	4.9e-2	5.83e-1	7.6e-2	43.3	16.3	81.6
73	W	52.55964	30.82271	3.510	-3.72e-2	1.0e-2	-4.67e-1	6.3e-2	88.5	51.7	44.1	0.000	8.81e-2	3.5e-2	3.83e+0	1.5e-1	93.1	76.0	149.6
74	W	52.56779	31.73597	0.000	3.23e-2	1.7e-2	6.23e-1	5.3e-2	48.2	32.3	5.8	0.000	1.23e-1	4.5e-2	6.19e-1	1.0e-1	22.7	22.1	82.8
75	G	52.58504	31.84034	27.54	6.95e-1	2.2e-2	8.63e-1	2.8e-2	8.4	8.4	57.1	27.36	9.61e-1	3.3e-2	8.51e-1	3.3e-2	13.5	13.5	66.6
76	W	52.58644	30.37985	0.000	2.17e-2	1.6e-2	9.03e-2	7.8e-2	56.3	53.2	16.9	0.000	8.22e-2	3.4e-2	4.15e-1	1.0e-1	51.3	44.5	163.2
77	W	52.59140	30.72418	0.000	5.01e-3	8.9e-3	-6.39e-2	3.1e-2	50.3	41.3	82.3	2.138	-4.50e-2	2.1e-2	1.15e-1	3.2e-2	42.0	14.4	134.2
78	W	52.59727	30.52785	3.320	-4.06e-2	1.2e-2	-1.16e-1	4.1e-2	52.2	24.2	17.8	0.000	1.96e-1	3.6e-2	1.52e+0	8.4e-2	45.2	26.9	75.6
79	W	52.60111	30.50754	22.53	-3.57e-2	1.5e-3	-1.73e-1	2.8e-3	37.6	8.4	121.0	6.096	7.93e-2	4.0e-2	4.03e-1	5.6e-2	37.8	18.4	27.7
80	2	52.66060	30.53677	7.358	1.11e-1	2.2e-2	3.34e-1	3.0e-2	19.6	10.9	110.2	9.062	2.66e-1	3.8e-2	3.63e-1	4.0e-2	18.0	13.5	84.6
81	W	52.66645	30.75864	7.023	-2.54e-2	3.6e-3	-3.88e-3	1.1e-2	54.1	21.7	89.0	0.000	5.58e-2	2.8e-2	6.21e-1	7.1e-2	58.7	42.5	24.8
82	W	52.68080	30.84786	4.359	-1.90e-2	4.3e-3	-1.27e-3	9.6e-3	38.7	14.7	59.1	0.003	-8.06e-5	2.2e-2	3.46e-1	5.5e-2	13.5	13.5	56.6
83	2	52.71907	30.90486	6.484	1.26e-1	2.3e-2	4.25e-1	8.6e-2	45.7	25.7	12.3	8.846	2.95e-1	4.1e-2	1.06e+0	9.8e-2	27.8	18.4	154.1
84	W	52.71975	30.45809	0.000	6.91e-2	2.3e-2	2.12e-1	1.5e-1	76.1	74.4	3.9	0.000	6.88e-2	3.7e-2	2.12e+0	1.3e-1	72.5	63.3	149.7
85	c2d	52.72979	30.52994	1.936	-8.51e-3	4.3e-3	8.88e-2	1.5e-2	64.1	25.4	46.5	0.000	6.60e-2	4.2e-2	5.54e-1	1.2e-1	59.3	54.5	18.6
86	W	52.73148	30.93104	2.736	-1.26e-2	4.6e-3	-5.09e-2	8.7e-3	51.5	10.7	71.0	0.000	4.77e-2	2.7e-2	2.91e-1	5.2e-2	43.5	32.6	74.2
87	W	52.79658	30.01275	1.763	-1.13e-2	6.4e-3	-2.45e-1	3.8e-2	94.6	49.1	52.7	0.000	1.02e-1	4.5e-2	2.04e+0	2.1e-1	108.5	80.9	1.7
88	c2d	52.82660	30.82787	10.36	1.70e-1	2.2e-2	1.56e-1	4.5e-2	11.7	8.4	103.0	6.517	2.38e-1	3.5e-2	2.45e-1	4.3e-2	18.6	13.5	71.4
89	W	52.84505	30.60484	0.000	4.55e-3	1.1e-2	-1.53e-1	3.2e-2	50.1	30.4	43.0	0.000	9.23e-2	3.6e-2	4.43e-1	7.2e-2	44.7	35.3	170.5
90	2	52.87120	30.51474	7.605	1.80e-1	2.1e-2	2.95e-1	3.7e-2	14.6	9.9	79.2	16.52	4.84e-1	3.8e-2	5.43e-1	4.5e-2	17.9	13.5	83.3
91	HL	52.88158	30.72896	0.000	2.63e-2	1.5e-2	1.79e-1	3.5e-2	33.0	23.4	79.2	0.000	2.61e-3	3.7e-2	6.61e-2	4.8e-2	45.1	13.5	131.6
92	G	52.90477	32.33035	8.630	1.86e-1	2.0e-2	2.32e-1	2.8e-2	8.4	8.4	111.7	15.16	5.27e-1	3.4e-2	5.08e-1	4.4e-2	13.5	13.5	113.1
93	W	52.93117	30.74336	0.000	1.61e-2	5.0e-3	1.29e-1	1.3e-2	14.9	12.8	113.6	0.000	4.37e-2	2.8e-2	9.13e-2	4.4e-2	31.6	26.2	32.1
94	W	52.94554	30.03632	0.000	2.99e-2	1.6e-2	-1.27e-1	5.3e-2	36.1	35.7	18.7	7.009	1.36e-1	4.1e-2	7.49e-1	7.2e-2	33.2	27.2	135.6
95	W	52.94828	30.01694	0.000	2.10e-2	1.5e-2	-8.09e-2	4.2e-2	30.6	26.3	148.9	0.000	6.78e-2	3.3e-2	2.47e-1	5.5e-2	34.3	29.6	34.4
96	G	53.00470	31.31602	10.65	1.65e-1	2.0e-2	1.98e-1	2.6e-2	10.5	8.4	39.4	11.44	3.69e-1	4.3e-2	2.78e-1	4.3e-2	13.5	13.5	114.8
97	W	53.06107	30.98750	0.000	3.48e-2	1.7e-2	-4.13e-1	8.8e-2	62.4	56.9	91.8	6.004	8.09e-2	3.2e-2	1.28e+0	9.7e-2	60.0	50.0	155.5
98	W	53.06163	30.91518	0.000	3.64e-2	1.9e-2	3.60e-1	1.4e-1	93.3	84.7	134.0	8.069	1.89e-1	3.7e-2	3.90e+0	1.6e-1	76.0	59.7	40.5
99	W	53.08332	30.88540	0.000	4.24e-3	1.2e-2	-3.48e-2	2.9e-2	50.2	25.2	28.5	0.000	4.34e-2	3.0e-2	1.78e-1	5.8e-2	40.1	33.2	43.0
100	W	53.11192	30.80632	0.000	2.30e-2	1.5e-2	2.56e-1	4.2e-2	37.1	31.8	114.4	13.71	3.12e-1	4.2e-2	1.09e+0	7.4e-2	26.2	22.1	134.1
101	W	53.11801	31.42308	1.811	-1.47e-2	8.1e-3	2.83e-4	2.3e-2	59.5	20.1	79.5	6.019	7.08e-2	4.3e-2	8.16e-1	9.4e-2	59.9	31.0	87.3
102	W	53.26341	30.91314	0.000	6.75e-3	9.7e-3	2.14e-1	3.8e-2	48.8	40.3	27.2	0.000	8.81e-2	3.3e-2	7.23e-1	8.8e-2	47.1	45.5	45.0
103	W	53.29845	30.99134	0.000	1.10e-2	1.3e-2	-5.06e-1	7.1e-2	81.6	63.0	47.2	6.021	1.08e-1	4.4e-2	1.50e+0	1.1e-1	78.0	27.5	133.3
104	W	53.30688	31.33361	0.000	2.96e-3	1.2e-2	-2.20e-1	4.8e-2	60.1	43.1	31.8	11.93	3.88e-1	6.6e-2	5.98e-1	1.4e-1	19.4	13.5	4.2
105	W	53.30898	31.19856	0.000	1.87e-3	1.0e-2	1.09e-1	3.0e-2	69.5	40.8	77.9	0.000	1.09e-1	4.0e-2	5.37e-1	9.3e-2	40.9	37.7	82.5
106	W	53.35184	31.43971	0.013	-1.15e-4	8.2e-3	-5.02e-2	3.5e-2	8.4	8.4	0.0	0.000	4.42e-2	2.7e-2	2.59e-1	6.7e-2	45.2	37.5	148.2
107	W	53.36533	31.37536	0.000	5.93e-2	1.8e-2	6.52e-1	8.1e-2	50.9	47.9	44.6	7.039	1.73e-1	4.1e-2	7.67e-1	1.0e-1	38.9	31.3	75.2
108	2	53.42230	31.22838	9.578	1.37e-1	1.9e-2	1.94e-1	3.2e-2	29.6	8.4	68.1	6.864	2.02e-1	3.5e-2	2.36e-1	5.5e-2	21.6	13.5	76.7
109	W	53.44435	30.46080	0.000	2.54e-3	7.6e-3	1.71e+0	5.4e-2	100.7	72.4	163.3	7.073	1.57e-1	6.5e-2	3.09e+0	2.6e-1	83.9	48.7	126.0
110	2	53.46041	31.43651	0.000	1.30e-2	1.0e-2	3.76e-1	4.4e-2	58.3	48.2	35.8	0.000	1.01e-1	3.6e-2	1.32e+0	9.9e-2	55.2	50.0	117.7
111	W	53.46134	31.20830	6.387	-5.84e-2	9.1e-3	6.48e-3	2.8e-2	70.4	21.0	158.1	0.000	1.42e-1	3.3e-2	1.90e+0	1.0e-1	62.1	46.1	63.9
112	W	53.46958	31.23553	0.000	3.78e-2	1.7e-2	7.42e-1	1.3e-1	95.2	87.5	168.8	0.000	4.74						

Table E.1. continued.

No.	Type	RA ₂₀₀₀ (degree)	Dec ₂₀₀₀ (degree)	Sig ₀₇₀	S_{070}^{peak} (Jy beam ⁻¹)	S_{070}^{tot} (Jy)	r_{070}^a (")	r_{070}^b (")	PA ₀₇₀ (°)	Sig ₁₆₀	S_{160}^{peak} (Jy beam ⁻¹)	S_{160}^{tot} (Jy)	r_{160}^a (")	r_{160}^b (")	PA ₁₆₀ (°)				
(1)	(2)	(3)	(4)	(5)	(6) ± (7)	(10) ± (11)	(12)	(13)	(14)	(15)	(16) ± (17)	(20) ± (21)	(22)	(23)	(24)				
130	W	54.81504	31.49644	1.609	-1.65e-2	1.0e-2	-2.45e-3	3.7e-2	56.7	29.0	47.1	0.000	7.39e-2	3.4e-2	1.20e+0	9.2e-2	57.1	50.3	146.9
131	W	54.87118	31.53178	1.159	-1.15e-2	9.9e-3	-1.59e-4	1.7e-2	22.9	8.4	113.8	0.000	7.68e-2	5.1e-2	1.50e-1	7.1e-2	35.2	22.9	127.1
132	W	54.89657	31.88566	0.000	5.56e-2	2.1e-2	1.69e-1	1.4e-1	79.9	79.2	91.2	8.829	1.86e-1	3.8e-2	2.56e+0	1.3e-1	60.1	41.6	4.9
133	W	54.96321	31.89303	0.000	2.20e-2	1.5e-2	1.77e+0	1.6e-1	126.0	19.6	73.5	7.135	1.11e-1	3.6e-2	3.49e+0	1.6e-1	121.7	53.4	152.3
134	W	55.04875	32.00877	3.659	-2.17e-2	5.9e-3	5.23e-3	1.2e-2	31.1	15.1	67.6	0.000	1.08e-1	3.8e-2	2.18e-1	7.0e-2	36.0	26.9	132.1
135	G	55.11486	31.72087	17.89	1.56e-1	2.2e-2	3.38e-1	5.1e-2	16.5	13.8	140.8	28.53	9.23e-1	4.3e-2	1.11e+0	6.9e-2	13.5	13.5	117.1
136	W	55.31937	31.70944	1.929	-1.59e-2	8.2e-3	1.23e-2	3.2e-2	53.5	32.9	52.0	0.000	9.29e-2	3.4e-2	1.05e+0	8.9e-2	59.5	45.0	158.8
137	W	55.42577	31.84510	6.058	5.73e-2	2.1e-2	8.26e-1	9.7e-2	57.5	41.2	39.0	11.88	2.06e-1	6.3e-2	2.30e+0	1.7e-1	54.7	39.0	88.7
138	W	55.48222	31.70643	0.000	2.40e-3	1.2e-2	-5.98e-2	2.6e-2	39.6	24.3	139.2	6.257	2.19e-1	7.0e-2	4.44e-1	1.0e-1	24.2	19.0	4.8
139	G	55.52977	32.76461	18.95	3.01e-1	2.2e-2	2.80e-1	2.7e-2	9.0	8.4	130.5	23.84	7.11e-1	4.2e-2	5.98e-1	4.2e-2	13.5	13.5	84.0
140	W	55.53155	31.96546	0.000	9.51e-4	1.4e-2	-1.57e-1	2.9e-2	73.2	20.4	100.1	0.000	8.06e-2	3.5e-2	5.71e-1	7.1e-2	48.5	34.4	148.6
141	W	55.68771	32.65634	0.000	2.68e-3	1.3e-2	5.16e-1	5.4e-2	67.9	53.4	129.5	12.36	2.62e-1	3.8e-2	9.92e-1	7.5e-2	35.7	16.9	31.0
142	G	55.71746	31.18435	20.41	3.49e-1	2.1e-2	5.26e-1	2.5e-2	11.2	8.4	33.4	16.84	5.55e-1	6.7e-2	5.67e-1	6.7e-2	13.5	13.5	25.1
143	W	55.76333	32.73297	0.000	6.70e-2	2.1e-2	-5.44e-2	1.0e-1	61.9	56.1	18.6	6.174	5.46e-2	3.0e-2	9.58e-1	7.0e-2	67.5	30.1	47.4
144	W	55.81280	31.67497	0.275	-3.22e-3	1.1e-2	1.09e-1	5.5e-2	108.1	24.5	64.9	0.000	1.51e-1	5.6e-2	1.56e+0	2.0e-1	82.6	56.3	11.6
145	G	55.87451	31.67162	0.000	1.15e-3	9.5e-3	2.55e-1	2.6e-2	8.4	8.4	0.0	7.384	1.64e-1	4.8e-2	3.57e-1	6.3e-2	27.6	15.5	54.5
146	W	55.89279	31.69999	0.000	1.14e-8	1.1e-2	-4.09e-2	2.7e-2	51.8	24.0	1.7	6.356	9.64e-2	4.0e-2	4.34e-1	7.1e-2	34.5	22.8	32.9
147	2	55.95272	31.84960	8.141	-4.03e-2	4.9e-3	2.49e-2	1.4e-2	72.9	16.7	40.6	8.233	1.13e-1	4.1e-2	1.45e+0	1.0e-1	65.6	33.9	130.2
148	W	55.98131	32.79943	0.000	1.01e-2	1.5e-2	-3.44e-2	4.3e-2	37.3	29.3	160.8	6.247	1.39e-1	3.4e-2	4.35e-1	5.2e-2	33.0	19.6	147.3
149	LRL	56.02004	32.08123	24.28	6.02e-2	3.5e-2	2.79e+0	1.5e-1	76.9	35.3	130.8	63.53	1.50e+0	2.0e-1	2.15e+1	6.4e-1	55.1	36.4	103.6
150	PSZ	56.06644	32.12152	1.676	-1.85e-2	1.1e-2	1.62e-2	4.5e-2	56.0	35.9	19.5	9.312	1.94e-1	7.9e-2	7.23e-1	1.3e-1	55.1	16.4	152.9
151	W	56.10446	31.85356	0.000	4.32e-2	2.0e-2	2.94e-1	9.2e-2	56.0	51.5	114.8	6.915	1.20e-1	4.3e-2	1.16e+0	9.8e-2	54.4	31.4	165.0
152	δ	56.13131	32.10462	26.73	3.11e-1	4.0e-2	1.47e+0	9.5e-2	15.1	12.5	8.6	10.88	2.90e-1	1.3e-1	7.01e-1	1.7e-1	24.9	17.6	110.0
153	LRL	56.14112	32.21060	0.000	1.05e-2	1.8e-2	1.12e+0	6.7e-2	61.8	42.0	145.9	15.79	6.56e-1	1.5e-1	4.70e+0	3.6e-1	38.6	30.2	157.8
154	NTC	56.14466	32.14748	179.4	2.54e-1	2.0e-1	7.37e+0	7.7e-1	70.3	20.4	76.7	80.58	2.17e+0	3.6e-1	3.22e+1	1.1e+0	62.9	34.9	76.3
155	W	56.17117	32.00816	4.733	-6.67e-2	1.4e-2	-7.53e-2	2.9e-2	37.2	13.4	167.4	0.000	4.28e-2	6.3e-2	1.95e-1	1.2e-1	39.0	35.4	132.6
156	W	56.18055	32.28282	0.000	2.41e-2	2.1e-2	4.07e-1	9.3e-2	63.5	49.4	169.5	11.36	3.49e-1	5.1e-2	1.59e+0	1.2e-1	33.0	21.0	17.6
157	c2d	56.18371	32.02561	345.6	5.22e+0	2.6e-2	1.38e+1	4.6e-2	18.1	8.4	105.9	935.7	2.81e+0	2.0e-1	4.66e+1	4.1e-1	14.9	13.5	107.1
158	W	56.18745	31.68569	6.127	-5.17e-2	8.4e-3	-7.65e-4	2.7e-2	52.1	28.1	121.6	0.000	7.32e-2	3.3e-2	1.20e+0	8.4e-2	49.8	44.1	31.6
159	PSZ	56.21042	32.06547	1.819	-3.71e-2	2.0e-2	-2.16e-2	5.4e-2	58.2	18.0	78.9	36.09	6.27e-1	1.4e-1	4.62e+0	2.7e-1	52.0	22.7	1.5
160	LRL	56.28257	32.00969	13.66	-5.20e-2	3.8e-3	-8.86e-2	7.8e-3	57.4	10.8	64.8	13.42	2.63e-1	4.2e-2	1.53e+0	9.8e-2	36.2	27.3	95.3
161	G	56.28887	33.13569	10.76	2.04e-1	2.1e-2	3.34e-1	4.8e-2	27.9	8.4	129.9	14.11	3.78e-1	3.4e-2	4.82e-1	9.3e-2	13.5	13.5	56.7
162	PSZ	56.33264	32.04403	0.000	5.29e-2	2.3e-2	1.55e-1	7.5e-2	41.7	36.5	172.3	9.917	1.65e-1	5.3e-2	8.38e-1	8.7e-2	38.9	20.6	165.6
163	W	56.33779	31.72232	0.000	4.83e-2	1.9e-2	-4.78e-2	7.1e-2	46.2	40.0	92.5	0.000	6.34e-2	3.3e-2	3.37e-1	7.4e-2	45.8	37.6	88.7
164	W	56.48066	32.68635	0.000	2.50e-6	1.2e-2	7.17e-5	5.9e-2	62.5	54.1	106.3	0.000	2.93e-2	2.7e-2	3.55e-1	7.0e-2	54.8	37.0	112.2
165	G	56.54580	33.14685	11.31	1.37e-1	2.1e-2	4.01e-1	3.2e-2	16.4	13.7	43.4	27.39	8.10e-1	3.2e-2	1.06e+0	3.8e-2	14.3	13.5	51.5
166	G	56.72649	33.34245	8.096	1.27e-1	2.0e-2	3.69e-1	7.5e-2	35.1	21.5	89.3	0.000	2.15e-1	3.1e-2	4.66e-1	7.3e-2	42.9	26.2	6.6
167	W	56.80190	32.66893	0.000	3.32e-2	1.7e-2	5.06e-2	5.8e-2	41.1	38.0	101.2	0.000	1.65e-1	3.7e-2	3.45e-1	6.9e-2	28.8	21.5	84.2
168	W	56.87434	32.94919	0.000	2.69e-2	1.6e-2	-3.13e-1	8.9e-2	67.4	60.1	29.3	0.000	1.30e-1	3.6e-2	1.13e+0	1.1e-1	59.9	45.4	37.3
169	W	56.89496	32.93415	0.000	3.07e-2	1.6e-2	1.54e-1	8.0e-2	62.1	57.0	138.8	0.000	4.46e-3	2.2e-2	1.17e-1	6.0e-2	77.7	55.3	46.9
170	W	57.11232	32.72454	0.000	2.51e-2	1.5e-2	1.85e-1	3.9e-2	39.9	27.7	139.9	0.000	1.07e-1	3.5e-2	3.87e-1	6.2e-2	34.4	30.0	169.0

Table E.2. Same as Table E.1 for 250 μm and 350 μm.

No.	Sig ₂₅₀	S_{250}^{peak} (Jy beam ⁻¹)	S_{250}^{tot} (Jy)	r_{250}^a (")	r_{250}^b (")	PA ₂₅₀ (°)	Sig ₃₅₀	S_{350}^{peak} (Jy beam ⁻¹)	S_{350}^{tot} (Jy)	r_{350}^a (")	r_{350}^b (")	PA ₃₅₀ (°)				
(1)	(25)	(26) ± (27)	(30) ± (31)	(32)	(33)	(34)	(35)	(36) ± (37)	(40) ± (41)	(42)	(43)	(44)				
1	9.220	2.64e-1	3.9e-2	1.05e+0	7.1e-2	35.8	30.5	67.6	10.40	2.58e-1	3.9e-2	6.09e-1	5.5e-2	36.2	32.1	91.2
2	7.151	2.02e-1	2.8e-2	6.30e-1	4.3e-2	37.4	24.2	75.1	8.910	2.21e-1	1.8e-2	4.18e-1	2.2e-2	35.2	26.9	70.2
3	18.46	5.30e-1	1.9e-1	1.41e+0	2.3e-1	42.2	18.2	102.8	29.90	9.23e-1	1.7e-1	1.76e+0	1.8e-1	42.8	24.9	106.7
4	12.45	2.62e-1	1.0e-1	4.85e-1	1.0e-1	28.9	19.2	32.4	15.93	3.08e-1	1.2e-1	3.96e-1	1.0e-1	31.4	24.9	42.0
5	12.08	4.41e-1	6.4e-2	4.32e+0	1.6e-1	64.6	43.5	53.1	16.96	5.37e-1	6.6e-2	2.70e+0	1.2e-1	64.1	40.9	58.5
6	6.539	1.49e-1	5.0e-2	3.15e-1	5.7e-2	33.3	18.3	87.5	6.998	1.51e-1	5.0e-2	2.16e-1	5.0e-2	33.1	24.9	74.3
7	6.285	1.75e-1	3.8e-2	5.39e-1	5.1e-2	39.5	27.1	123.8	6.814	1.78e-1	2.8e-2	3.01e-1	2.9e-2	36.8	27.2	154.3
8	7.144	2.52e-1	8.5e-2	9.39e-1	1.5e-1	36.3	28.9	28.9	11.79	4.04e-1	9.0e-2	9.09e-1	1.2e-1	38.2	30.2	34.7
9	6.719	1.55e-1	3.7e-2	2.44e-1	4.0e-2	22.8	19.6	175.0	9.234	2.36e-1	4.8e-2	2.79e-1	4.4e-2	24.9	24.9	15.4
10	6.246	1.27e-1	7.0e-2	2.57e-1	8.1e-2	30.3	21.8	148.8	6.023	1.56e-1	8.0e-2	2.21e-1	7.8e-2	31.4	24.9	121.5
11	20.12	7.28e-1	2.3e-1	2.32e+0	4.1e-1	32.1	26.4	60.0	39.01	1.21e+0	4.9e-1	2.31e+0	6.3e-1	34.9	26.4	52.1
12	14.80	3.29e-1	6.4e-2	7.10e-1	7.8e-2	31.2	21.6	8.2	22.62	4.98e-1	5.4e-2	6.86e-1	5.7e-2	30.0	24.9	4.1
13	0.000	9.55e-2	3.0e-2	6.78e-1	6.7e-2	74.1	51.9	59.4	7.361	1.99e-1	3.5e-2	6.62e-1	6.1e-2	45.7	33.6	92.9
14	6.915	1.78e-1	5.5e-2	2.92e-1	6.3e-2	29.8	18.2	152.6	7.412	1.75e-1	5.7e-2	1.99e-1	5.8e-2	30.3	24.9	139.7
15	7.615	2.05e-1	3.8e-2	6.18e-1	5.9e-2	37.2	24.0									

Table E.2. continued.

No.	Sig ₂₅₀	S_{250}^{peak} (Jy beam ⁻¹)		S_{250}^{tot} (Jy)		r_{250}^a ($''$)	r_{250}^b ($''$)	PA ₂₅₀ ($^\circ$)	Sig ₃₅₀	S_{350}^{peak} (Jy beam ⁻¹)		S_{350}^{tot} (Jy)	r_{350}^a ($''$)	r_{350}^b ($''$)	PA ₃₅₀ ($^\circ$)	
(1)	(25)	(26) ±	(27)	(30) ±	(31)	(32)	(33)	(34)	(35)	(36) ±	(37)	(40) ±	(41)	(42)	(43)	(44)
26	14.46	3.75e-1	3.0e-2	1.00e+0	3.8e-2	37.6	20.7	142.8	20.00	5.49e-1	5.5e-2	1.10e+0	6.0e-2	43.0	25.9	148.4
27	7.800	1.72e-1	5.6e-2	2.92e-1	6.0e-2	32.8	18.9	99.1	7.131	1.92e-1	9.3e-2	3.00e-1	8.9e-2	34.6	26.9	34.2
28	7.986	4.51e-1	1.2e-1	2.61e+0	2.3e-1	51.9	34.2	147.2	17.81	7.22e-1	1.6e-1	2.32e+0	2.3e-1	50.5	34.6	158.4
29	7.291	2.55e-1	5.6e-2	9.93e-1	9.1e-2	46.8	22.8	9.6	9.402	2.57e-1	5.7e-2	6.66e-1	7.2e-2	51.3	26.1	0.1
30	0.000	2.17e-1	4.6e-2	5.15e+0	1.9e-1	85.2	76.9	3.5	6.449	2.28e-1	4.0e-2	2.58e+0	1.3e-1	76.7	71.6	31.2
31	32.41	6.63e-1	2.6e-2	6.00e-1	2.3e-2	18.2	18.2	27.1	14.28	2.84e-1	1.9e-2	2.22e-1	1.7e-2	24.9	24.9	36.2
32	9.334	2.18e-1	4.3e-2	7.67e-1	5.5e-2	48.4	23.4	13.8	13.32	3.48e-1	7.2e-2	6.96e-1	8.2e-2	43.7	25.8	4.1
33	7.302	1.45e-1	4.0e-2	3.86e-1	5.1e-2	48.7	23.2	87.1	7.571	2.13e-1	4.2e-2	3.96e-1	4.9e-2	38.1	24.9	34.9
34	12.68	5.34e-1	1.0e-1	9.55e-1	1.0e-1	26.4	20.4	165.9	20.41	6.66e-1	9.5e-2	9.51e-1	8.7e-2	34.0	24.9	4.0
35	80.78	2.13e+0	1.1e-1	3.57e+0	1.3e-1	22.7	19.2	81.7	80.77	2.19e+0	2.1e-1	2.82e+0	2.0e-1	27.9	24.9	66.7
36	27.49	5.97e-1	1.0e-1	7.22e-1	9.8e-2	21.9	18.2	29.9	26.90	7.51e-1	1.6e-1	8.28e-1	1.4e-1	24.9	24.9	67.5
37	91.96	2.00e+0	1.1e-1	1.99e+0	1.1e-1	18.2	18.2	166.3	82.59	1.71e+0	1.0e-1	1.53e+0	9.8e-2	24.9	24.9	171.2
38	39.96	1.50e+0	2.0e-1	1.98e+0	1.8e-1	21.9	18.2	79.2	30.46	2.09e+0	1.7e-1	3.41e+0	1.6e-1	32.3	27.8	95.2
39	10.36	2.72e-1	6.3e-2	1.36e+0	1.0e-1	58.9	30.6	17.3	15.61	4.08e-1	6.9e-2	1.01e+0	8.6e-2	61.6	28.2	16.1
40	34.73	6.15e+0	2.7e-1	4.25e+0	4.6e-1	38.1	22.6	101.4	52.24	1.97e+0	3.3e-1	4.36e+0	4.2e-1	41.6	24.9	90.3
41	67.50	2.74e+0	1.7e-1	1.14e+1	3.9e-1	36.3	31.8	78.6	138.1	4.32e+0	2.5e-1	6.07e+0	4.3e-1	26.2	24.9	85.2
42	9.194	2.40e-1	9.4e-2	3.41e-1	1.0e-1	24.1	18.2	5.8	11.25	2.13e-1	1.5e-1	2.21e-1	1.4e-1	28.7	24.9	32.9
43	36.02	7.40e-1	4.4e-1	1.20e+0	5.3e-1	24.9	18.6	81.0	64.90	1.31e+0	5.7e-1	1.57e+0	5.6e-1	28.6	24.9	53.5
44	15.92	6.15e-1	8.7e-2	2.07e+0	1.3e-1	42.5	27.1	133.3	21.61	7.34e-1	1.1e-1	1.39e+0	1.2e-1	39.1	26.8	142.4
45	6.942	1.84e-1	4.2e-2	3.10e-1	4.6e-2	23.7	20.7	162.6	6.484	1.41e-1	2.2e-2	1.69e-1	2.0e-2	30.3	24.9	1.7
46	22.41	5.48e-1	1.2e-1	5.51e-1	1.1e-1	23.4	18.2	145.0	12.97	3.38e-1	1.4e-1	3.40e-1	1.2e-1	30.5	24.9	134.8
47	88.95	2.57e+0	5.0e-1	6.01e+0	6.7e-1	36.8	19.2	21.4	110.7	2.98e+0	8.1e-1	5.92e+0	9.2e-1	41.9	25.8	17.5
48	25.92	4.92e-1	3.3e-1	1.91e-1	2.9e-1	26.7	18.2	59.3	16.25	4.20e-1	6.7e-1	1.32e-1	6.1e-1	28.0	24.9	162.7
49	117.6	3.82e+0	6.4e-1	5.08e+0	5.9e-1	26.3	18.2	7.9	77.60	4.17e+0	6.7e-1	5.02e+0	6.1e-1	28.0	24.9	162.5
50	23.52	6.89e-1	1.9e-1	1.63e+0	2.4e-1	37.5	18.2	9.2	22.44	6.13e-1	2.1e-1	9.72e-1	2.3e-1	38.1	24.9	8.8
51	25.04	6.67e+0	1.1e+0	9.64e+0	1.1e+0	22.5	18.2	141.7	334.0	8.95e+0	2.0e+0	1.07e+1	1.8e+0	30.0	24.9	70.5
52	11.27	2.88e+0	7.9e-1	4.17e+0	7.9e-1	33.1	18.2	14.0	51.02	8.63e-1	5.8e-1	7.62e-1	5.2e-1	35.2	24.9	19.1
53	6.067	2.00e-1	4.2e-2	6.59e-1	4.7e-2	45.1	25.7	105.9	8.250	2.52e-1	3.7e-2	5.29e-1	3.7e-2	44.3	31.1	106.9
54	954.2	1.82e+1	5.9e-1	1.84e+1	5.2e-1	18.3	18.2	83.4	750.6	1.03e+1	9.0e-1	1.03e+1	8.2e-1	24.9	24.9	73.7
55	1269	2.83e+1	5.3e-1	4.26e+1	5.9e-1	25.1	18.2	52.6	954.0	2.01e+1	9.4e-1	2.30e+1	9.2e-1	25.3	24.9	48.5
56	32.90	6.74e-1	2.4e-1	8.53e-1	2.4e-1	27.6	18.2	139.7	28.85	5.22e-1	3.1e-1	4.15e-1	2.9e-1	25.0	24.9	133.7
57	6.838	1.88e-1	4.1e-2	6.43e-1	6.6e-2	33.8	27.3	28.1	8.195	1.98e-1	4.5e-2	3.52e-1	5.6e-2	31.1	26.6	29.2
58	138.1	3.61e+0	8.3e-1	7.82e+0	1.0e+0	30.1	20.4	96.5	97.34	2.21e+0	5.8e-1	2.97e+0	6.2e-1	30.4	24.9	110.4
59	34.82	1.09e+0	3.0e-1	1.80e+0	3.2e-1	29.5	18.2	154.7	31.67	6.19e-1	3.5e-1	5.36e-1	3.2e-1	35.8	24.9	119.3
60	54.22	1.25e+0	5.5e-1	1.20e+0	4.9e-1	20.0	18.2	88.2	7.798	4.50e-1	1.8e-1	3.10e-1	1.7e-1	28.5	24.9	102.8
61	11.56	5.74e-1	4.7e-2	5.33e+0	1.3e-1	50.4	44.6	136.0	17.42	6.53e-1	4.1e-2	3.25e+0	8.9e-2	50.1	44.6	131.8
62	265.3	5.68e+0	1.8e+0	7.05e+0	1.6e+0	19.1	18.2	137.1	199.8	4.21e+0	1.0e+0	4.91e+0	9.6e-1	24.9	24.9	20.4
63	8.690	2.51e-1	9.4e-2	5.78e-1	1.1e-1	34.6	20.3	3.2	9.368	2.61e-1	6.6e-2	3.73e-1	6.8e-2	36.4	24.9	176.9
64	23.15	4.44e-1	1.4e-2	3.74e-1	1.2e-2	18.2	18.2	159.1	9.039	1.90e-1	1.2e-2	1.61e-1	1.1e-2	24.9	24.9	122.9
65	7.830	2.01e-1	5.1e-2	4.76e-1	6.4e-2	46.8	18.2	52.6	11.06	3.04e-1	7.7e-2	6.06e-1	8.5e-2	42.3	25.3	47.2
66	8.786	2.21e-1	4.6e-2	5.18e-1	5.2e-2	43.0	18.2	143.7	8.878	1.86e-1	3.5e-2	2.28e-1	3.4e-2	39.3	24.9	143.4
67	7.192	2.20e-1	7.0e-2	8.00e-1	1.0e-1	41.7	26.2	15.2	8.916	2.39e-1	5.8e-2	5.02e-1	7.0e-2	38.8	27.1	15.8
68	15.35	3.09e-1	1.9e-2	1.41e-1	1.7e-2	18.2	18.2	141.1	10.26	1.76e-1	4.9e-2	9.84e-2	4.5e-2	24.9	24.9	125.6
69	29.68	6.26e-1	2.9e-2	4.90e-1	2.5e-2	18.2	18.2	66.4	11.39	2.47e-1	3.1e-2	1.95e-1	2.8e-2	24.9	24.9	106.3
70	6.873	1.49e-1	7.3e-2	2.85e-1	7.7e-2	27.6	20.6	87.1	6.534	1.47e-1	7.5e-2	1.77e-1	7.0e-2	25.8	24.9	66.2
71	6.162	3.52e-1	2.7e-2	1.71e+0	3.9e-2	48.0	33.2	96.6	6.244	3.76e-1	3.1e-2	9.56e-1	3.7e-2	40.4	33.6	136.6
72	34.16	9.77e-1	1.2e-1	2.36e+0	1.4e-1	40.2	19.9	75.6	47.29	1.47e+0	1.5e-1	2.56e+0	1.6e-1	41.2	24.9	76.8
73	11.42	4.43e-1	5.3e-2	2.05e+0	1.2e-1	46.2	24.2	94.8	18.49	5.17e-1	5.7e-2	1.45e+0	9.9e-2	48.8	24.9	93.6
74	6.795	2.65e-1	2.6e-2	6.32e-1	3.8e-2	26.1	23.5	75.5	6.317	3.53e-1	3.9e-2	5.12e-1	4.4e-2	26.1	24.9	22.6
75	13.58	2.55e-1	2.6e-2	2.01e-1	2.3e-2	18.2	18.2	137.3	0.000	5.82e-2	2.7e-2	1.90e-2	2.5e-2	24.9	24.9	162.4
76	7.896	1.72e-1	3.9e-2	5.97e-1	5.8e-2	49.3	26.9	152.4	13.07	2.92e-1	5.7e-2	5.04e-1	6.8e-2	42.6	24.9	133.2
77	9.077	2.29e-1	3.0e-2	5.59e-1	4.2e-2	28.4	22.4	53.9	12.21	2.74e-1	2.5e-2	3.64e-1	2.9e-2	27.0	24.9	57.1
78	0.000	2.03e-1	4.2e-2	1.19e+0	7.5e-2	49.6	40.3	102.5	8.692	2.95e-1	7.6e-2	8.17e-1	9.8e-2	47.8	37.8	179.6
79	18.75	5.40e-1	8.3e-2	9.91e-1	9.0e-2	28.9	18.2	16.4	23.37	7.50e-1	1.1e-1	1.13e+0	1.0e-1	34.4	24.9	8.4
80	15.29	2.98e-1	8.0e-2	2.28e-1	7.1e-2	18.2	18.2	166.2	9.651	2.12e-1	7.2e-2	1.78e-1	6.6e-2	26.8	24.9	11.3
81	6.405	1.66e-1	3.9e-2	6.48e-1	6.9e-2	37.0	31.1	0.5	8.212	2.21e-1	4.0e-2	4.94e-1	5.5e-2	35.6	31.4	144.0
82	0.000	1.70e-1	6.5e-2	5.91e-1	7.8e-2	45.5	25.1	174.3	11.14	3.19e-1	1.1e-1	5.57e-1	1.1e-1	45.7	24.9	175.4
83	13.41	5.06e-1	7.5e-2	1.33e+0	1.1e-1	30.3	18.2	155.1	17.97	4.71e-1	7.8e-2	8.86e-1	9.0e-2	38.0	24.9	160.1
84	6.478	2.17e-1	4.1e-2	1.54e+0	9.1e-2	53.0	38.0	34.4	9.606	2.43e-1	7.3e-2	9.58e-1	1.1e-1	60.3	37.5	33.8
85	0.000	1.38e-1	4.0e-2	5.48e-1	5.9e-2	57.8	28.4	141.3	6.600	1.94e-1	6.2e-2	4.42e-1	7.9e-2	44.9	28.7	144.1
86	7.162	2.11e-1	1.9e-2	4.18e-1	2.1e-2	40.3	18.2	20.0	8.944	2.12e-1	2.0e-2	3.09e-1	2.0e-2	37.7	24.9	22.1
87	6.218	3.12e-1	7.9e-2	3.15e+0	2.5e-1	56.7	44.5	110.1	11.94	3.26e-1	1.0e-1	1.40e+0	1.9e-1	73.7	33.1	179.3
88	12.48	2.50e-1	1.6e-2	2.11e-1	1.5e-2	18.2	18.2	155.7	13.23	2.72e-1	1.6e-2	2.09e-1	1.4e-2	24.9	24.9	130.7
89	7.045	1.89e-1	4.3e-2	4.28e-1	4.9e-2	41.2	18.2	174.5	7.230	1.62e-1	3.3e-2	2.46e-1	3.4e-2	39.8	24.9	175.2
90	28.73	5.93e-1	1.9e-2	4.26e-1	1.7e-2	18.2	18.2	170.3	21.46	4.43e-1	3.5e-2	3.44e-1	3.2e-2	24.9	24.9	171.2
91	6.648	1.67e-1	7.1e-2	2.34e-1	7.0e-2	31.4	18.7	63.1	10.41	2.06e-1	8.9e-2	2.47e-1	8.1e-2	33.4</		

Table E.2. continued.

No.	Sig ₂₅₀	S_{250}^{peak} (Jy beam ⁻¹)		S_{250}^{tot} (Jy)		r_{250}^a ($''$)	i_{250}^b ($''$)	PA ₂₅₀ ($^\circ$)	Sig ₃₅₀	S_{350}^{peak} (Jy beam ⁻¹)		S_{350}^{tot} (Jy)	r_{350}^a ($''$)	i_{350}^b ($''$)	PA ₃₅₀ ($^\circ$)	
(1)	(25)	(26) ±	(27)	(30) ±	(31)	(32)	(33)	(34)	(35)	(36) ±	(37)	(40) ±	(41)	(42)	(43)	(44)
98	20.29	6.51e-1	7.3e-2	2.76e+0	1.5e-1	52.5	23.1	3.6	35.72	1.04e+0	5.8e-2	2.56e+0	9.8e-2	52.3	24.9	10.0
99	11.14	3.17e-1	5.2e-2	6.51e-1	7.2e-2	23.8	22.3	38.6	16.64	4.20e-1	7.6e-2	5.47e-1	7.6e-2	28.3	24.9	23.3
100	30.69	7.77e-1	6.7e-2	1.13e+0	7.6e-2	21.8	18.2	156.3	33.09	7.14e-1	7.3e-2	7.51e-1	6.8e-2	24.9	24.9	156.6
101	11.07	4.01e-1	2.4e-2	1.68e+0	4.2e-2	39.5	33.5	99.8	13.47	5.35e-1	5.1e-2	1.39e+0	6.4e-2	45.8	34.6	24.4
102	6.219	1.66e-1	2.3e-2	4.82e-1	3.0e-2	41.2	22.2	32.8	7.663	1.94e-1	3.7e-2	3.43e-1	4.1e-2	39.5	24.9	32.8
103	13.08	4.06e-1	8.5e-2	1.96e+0	1.4e-1	60.9	26.5	141.2	23.46	6.71e-1	1.1e-1	2.28e+0	1.7e-1	56.4	32.8	143.2
104	44.56	1.38e+0	1.5e-1	3.62e+0	2.3e-1	29.6	19.7	167.6	89.14	2.37e+0	2.5e-1	4.53e+0	3.0e-1	36.0	26.3	165.3
105	11.12	2.64e-1	6.0e-2	5.70e-1	8.6e-2	26.9	22.0	114.2	12.86	2.61e-1	9.6e-2	3.44e-1	1.0e-1	30.7	24.9	87.4
106	6.242	1.64e-1	2.9e-2	5.55e-1	3.6e-2	54.6	20.1	139.7	9.088	2.63e-1	3.0e-2	5.20e-1	3.3e-2	49.8	24.9	143.1
107	8.721	2.41e-1	1.0e-1	5.91e-1	1.6e-1	28.0	20.9	100.5	7.949	1.65e-1	1.4e-1	2.49e-1	1.6e-1	43.2	24.9	65.2
108	10.62	2.01e-1	2.9e-2	1.78e-1	2.9e-2	32.1	18.2	124.5	11.02	2.37e-1	3.9e-2	2.33e-1	3.6e-2	30.8	24.9	107.0
109	7.029	3.40e-1	1.0e-1	4.19e+0	3.3e-1	61.6	52.4	7.1	8.239	3.26e-1	1.0e-1	2.19e+0	2.4e-1	60.0	51.2	14.8
110	9.722	2.98e-1	5.6e-2	1.01e+0	9.0e-2	35.8	24.9	110.8	10.89	2.64e-1	5.3e-2	5.01e-1	6.1e-2	38.9	24.9	117.3
111	11.99	3.31e-1	4.3e-2	1.49e+0	6.8e-2	55.3	25.5	51.2	23.38	5.95e-1	5.9e-2	1.58e+0	7.5e-2	54.9	26.9	52.6
112	0.000	2.52e-1	4.1e-2	2.69e+0	1.1e-1	67.0	48.2	108.2	11.82	3.19e-1	5.2e-2	1.37e+0	7.9e-2	82.6	29.9	96.0
113	10.85	3.26e-1	7.2e-2	8.85e-1	1.0e-1	30.3	23.5	139.0	12.25	3.13e-1	8.2e-2	5.36e-1	9.5e-2	31.5	25.6	109.9
114	6.197	1.38e-1	8.3e-2	2.47e-1	1.0e-1	43.4	18.2	108.7	10.39	2.72e-1	9.1e-2	5.35e-1	1.0e-1	38.9	24.9	5.1
115	6.282	1.84e-1	8.9e-2	7.77e-1	1.2e-1	47.1	32.7	93.1	11.10	2.07e-1	4.9e-2	3.69e-1	5.4e-2	41.9	24.9	126.3
116	23.96	4.93e-1	1.1e-2	4.13e-1	1.0e-2	18.2	18.2	177.6	8.756	1.70e-1	1.5e-2	1.34e-1	1.3e-2	24.9	24.9	4.5
117	6.041	1.77e-1	6.7e-2	1.48e+0	1.4e-1	60.2	43.2	91.2	6.299	1.43e-1	6.2e-2	3.97e-1	8.4e-2	60.1	29.2	105.7
118	18.88	5.23e-1	1.1e-1	1.24e+0	1.5e-1	42.9	19.9	146.1	25.52	6.95e-1	1.6e-1	1.27e+0	1.8e-1	40.5	25.4	141.5
119	8.119	2.34e-1	3.8e-2	1.69e+0	9.0e-2	54.8	38.4	163.4	8.714	2.42e-1	3.7e-2	9.76e-1	7.0e-2	52.6	39.1	158.0
120	9.627	3.34e-1	6.2e-2	1.22e+0	1.0e-1	54.6	18.2	135.2	16.25	4.43e-1	1.1e-1	1.34e+0	1.6e-1	62.0	24.9	134.5
121	9.711	2.09e-1	4.7e-2	2.59e-1	4.5e-2	22.7	18.2	9.5	7.487	1.41e-1	8.4e-2	1.17e-1	7.7e-2	24.9	24.9	168.6
122	18.28	5.93e-1	6.5e-2	3.13e+0	1.4e-1	39.4	32.2	121.8	26.69	7.37e-1	1.2e-1	2.28e+0	2.0e-1	41.1	33.8	132.7
123	6.611	1.44e-1	5.5e-2	2.95e-1	6.1e-2	34.1	20.4	99.4	6.325	1.36e-1	4.1e-2	1.43e-1	4.1e-2	31.9	24.9	95.5
124	7.168	2.04e-1	6.6e-2	7.18e-1	1.0e-1	39.0	25.1	33.9	8.817	2.27e-1	8.8e-2	4.71e-1	1.0e-1	40.8	26.9	31.9
125	15.86	3.73e-1	3.8e-2	9.69e-1	5.0e-2	45.5	18.2	147.3	11.58	3.33e-1	8.7e-2	6.36e-1	1.0e-1	45.6	24.9	149.6
126	59.48	1.22e+0	1.5e-2	9.84e-1	1.5e-2	18.2	18.2	40.9	23.85	4.94e-1	1.4e-2	4.34e-1	1.3e-2	24.9	24.9	114.3
127	8.700	3.33e-1	9.5e-2	3.44e+0	2.4e-1	83.3	37.0	48.1	10.62	3.51e-1	1.1e-1	2.07e+0	2.4e-1	70.7	39.7	35.9
128	13.82	6.18e-1	1.3e-1	1.05e+1	5.3e-1	85.3	54.2	138.7	14.58	5.62e-1	1.3e-1	5.98e+0	3.8e-1	90.3	56.5	139.2
129	12.60	3.52e-1	3.4e-2	1.03e+0	4.3e-2	42.1	21.0	136.4	12.19	3.11e-1	2.6e-2	6.34e-1	3.0e-2	46.3	24.9	136.9
130	6.523	2.55e-1	4.5e-2	1.17e+0	7.9e-2	40.5	34.0	154.9	9.752	2.84e-1	8.3e-2	7.82e-1	1.1e-1	41.6	36.1	87.2
131	7.346	1.56e-1	6.7e-2	1.93e-1	6.7e-2	23.4	18.2	121.8	8.679	1.67e-1	9.4e-2	1.55e-1	8.5e-2	31.1	24.9	108.0
132	6.470	2.45e-1	4.5e-2	1.51e+0	1.0e-1	55.7	34.1	162.1	8.688	2.11e-1	3.4e-2	7.60e-1	5.3e-2	62.0	31.3	150.3
133	6.799	2.99e-1	2.4e-2	4.46e+0	8.9e-2	75.1	56.0	158.4	8.165	3.16e-1	3.4e-2	2.71e+0	9.6e-2	75.1	56.5	162.3
134	7.686	1.89e-1	8.4e-2	3.31e-1	9.2e-2	36.7	18.2	134.6	8.177	1.29e-1	1.4e-1	1.30e-1	1.3e-1	42.2	24.9	120.0
135	28.83	5.83e-1	3.2e-2	5.22e-1	3.3e-2	18.2	18.2	162.4	11.82	2.11e-1	2.9e-2	1.66e-1	2.6e-2	24.9	24.9	150.3
136	7.476	1.96e-1	4.5e-2	5.89e-1	6.5e-2	40.2	22.1	177.0	8.544	1.84e-1	4.2e-2	4.02e-1	4.9e-2	47.4	25.7	173.8
137	11.37	2.19e-1	3.5e-2	6.61e-1	4.8e-2	51.7	20.2	76.4	11.66	2.59e-1	4.4e-2	6.16e-1	5.4e-2	48.5	28.4	74.6
138	18.71	4.58e-1	8.3e-2	6.30e-1	8.9e-2	19.1	18.2	66.1	20.98	4.99e-1	1.2e-1	5.24e-1	1.1e-1	30.0	24.9	62.3
139	14.27	2.99e-1	5.2e-2	2.65e-1	4.6e-2	18.2	18.2	91.6	0.000	1.37e-2	4.5e-2	6.89e-3	4.1e-2	24.9	24.9	19.3
140	8.379	2.22e-1	4.4e-2	5.87e-1	5.1e-2	43.2	19.2	159.4	11.49	2.94e-1	5.9e-2	4.81e-1	6.1e-2	42.8	24.9	153.7
141	10.23	3.00e-1	6.8e-2	7.67e-1	8.8e-2	40.0	18.4	32.2	7.819	2.59e-1	6.4e-2	5.17e-1	7.3e-2	42.3	24.9	33.0
142	12.96	3.10e-1	6.4e-2	2.84e-1	5.6e-2	18.2	18.2	135.0	6.493	1.44e-1	6.0e-2	1.62e-1	5.4e-2	27.1	24.9	129.3
143	7.395	2.08e-1	6.8e-2	1.18e+0	1.2e-1	50.9	33.8	74.5	8.008	2.10e-1	5.0e-2	5.95e-1	6.8e-2	47.1	32.3	85.8
144	6.079	2.55e-1	7.8e-2	1.82e+0	1.6e-1	57.6	34.0	10.9	7.791	2.18e-1	8.5e-2	9.83e-1	1.4e-1	64.2	37.9	4.2
145	10.14	2.51e-1	7.5e-2	4.06e-1	7.5e-2	27.6	18.2	156.4	7.260	1.54e-1	6.0e-2	1.64e-1	5.5e-2	26.6	24.9	168.4
146	8.140	2.29e-1	8.2e-2	4.85e-1	9.3e-2	31.4	19.3	81.5	7.103	1.69e-1	6.1e-2	2.18e-1	6.3e-2	29.6	24.9	68.2
147	7.573	2.46e-1	2.3e-2	1.20e+0	4.2e-2	54.1	30.6	133.3	8.001	2.23e-1	2.1e-2	6.38e-1	3.0e-2	49.5	30.5	138.5
148	8.602	2.17e-1	4.0e-2	4.00e-1	4.2e-2	33.3	18.2	112.3	8.324	2.04e-1	5.6e-2	2.46e-1	5.3e-2	29.3	24.9	129.8
149	44.99	2.03e+0	7.8e-1	1.49e+1	1.7e+0	50.0	39.2	92.2	43.92	1.55e+0	6.4e-1	6.19e+0	1.1e+0	48.0	39.2	85.5
150	23.35	6.36e-1	6.1e-2	1.82e+0	1.1e-1	28.6	22.5	153.3	24.93	6.33e-1	6.2e-2	1.31e+0	8.8e-2	32.7	28.1	160.1
151	6.800	1.94e-1	4.7e-2	8.25e-1	7.1e-2	48.3	29.7	178.5	7.337	1.76e-1	4.6e-2	3.57e-1	5.3e-2	48.6	24.9	145.2
152	0.000	6.23e-2	9.1e-2	9.14e-2	8.7e-2	28.1	22.9	136.7	0.000	6.06e-2	6.7e-2	6.62e-2	6.1e-2	28.1	24.9	57.2
153	17.49	8.13e-1	1.3e-1	2.79e+0	2.2e-1	34.1	29.0	61.1	18.88	6.81e-1	1.0e-1	1.38e+0	1.2e-1	37.0	28.7	39.7
154	53.16	1.57e+0	2.5e-1	8.79e+0	4.6e-1	60.5	26.1	83.7	47.25	1.42e+0	1.4e-1	5.08e+0	2.4e-1	53.8	32.0	81.4
155	14.36	3.26e-1	2.6e-1	6.12e-1	2.9e-1	36.4	20.8	113.1	19.96	4.73e-1	3.4e-1	5.56e-1	3.5e-1	25.9	24.9	109.1
156	14.28	3.94e-1	6.1e-2	1.11e+0	1.0e-1	28.4	21.5	8.8	14.24	3.10e-1	6.7e-2	5.94e-1	8.8e-2	34.9	24.9	21.6
157	966.9	2.23e+1	5.6e-1	2.42e+1	5.9e-1	18.2	18.2	129.6	535.7	1.21e+1	5.0e-1	1.26e+1	4.5e-1	24.9	24.9	138.2
158	11.29	3.49e-1	6.6e-2	9.72e-1	9.6e-2	32.4	22.4	136.8	16.16	4.93e-1	7.4e-2	8.44e-1	8.5e-2	33.1	24.9	134.9
159	22.61	8.46e-1	1.5e-1	2.89e+0	1.9e-1	48.4	22.4	175.4	19.46	6.09e-1	1.4e-1	1.47e+0	1.7e-1	54.1	24.9	176.9
160	30.13	8.47e-1	7.2e-2	2.64e+0	9.8e-2	37.6	23.8	106.8	39.73	1.09e+0	5.9e-2	2.00e+0	6.5e-2	39.1	24.9	105.6
161	8.824	1.65e-1	2.7e-2	9.83e-2	4.9e-2	18.2	18.2	116.6	0.000	8.74e-2	2.6e-2	5.51e-2	3.6e-2	24.9	24.9	49.5
162	16.03	4.15e-1	8.1e-2	7.40e-1	9.1e-2	29.3	18.2	161.8	15.59	3.45e-1	8.8e-2	4.04e-1	8.7e-2	31.3	24.9	167.3
163	6.268	1.84e-1	6.6e-2	5.27e-1	9.6e-2</											

Table E.3. Same as Table E.1 for 500 μm and high-resolution column density map.

No.	Sig ₅₀₀	S_{500}^{peak} (Jy beam ⁻¹)		S_{500}^{tot} (Jy)		r_{500}^a ($''$)	r_{500}^b ($''$)	PA ₅₀₀ ($^\circ$)	Sig _{NH₂}	$N_{\text{H}_2}^{\text{peak}}$ (10 ²⁰ cm ⁻²)	$r_{\text{H}_2}^a$ ($''$)	$r_{\text{H}_2}^b$ ($''$)	PA _{H₂} ($^\circ$)
(1)	(45)	(46) \pm (47)		(49) \pm (50)		(51)	(52)	(53)	(54)	(55)	(59)	(60)	(61)
1	8.059	1.77e-1	4.8e-2	2.49e-1	4.7e-2	44.3	37.6	131.8	12.36	8.45e-2	37.2	31.0	88.2
2	7.044	1.38e-1	2.2e-2	1.63e-1	2.0e-2	49.4	36.3	78.7	10.68	6.56e-2	41.0	32.8	87.7
3	24.31	8.16e-1	2.0e-1	8.29e-1	1.8e-1	38.6	36.3	106.4	61.64	4.17e-1	34.5	18.2	113.3
4	13.34	4.69e-1	1.2e-1	5.12e-1	1.1e-1	38.2	36.3	3.0	24.97	1.34e-1	22.6	18.2	49.7
5	14.81	4.36e-1	7.0e-2	1.10e+0	9.5e-2	63.0	42.3	54.9	22.97	1.69e-1	61.1	42.3	52.9
6	0.000	9.99e-2	4.1e-2	1.12e-1	3.7e-2	40.6	36.3	101.5	8.158	4.87e-2	31.4	24.9	82.3
7	7.058	1.47e-1	4.1e-2	1.53e-1	3.8e-2	51.2	36.3	148.9	11.29	6.09e-2	40.9	21.1	142.8
8	14.44	3.52e-1	1.7e-1	4.69e-1	1.6e-1	41.1	36.3	59.5	21.53	1.67e-1	39.1	21.3	14.8
9	9.602	2.25e-1	5.2e-2	2.23e-1	4.8e-2	36.3	36.3	93.1	16.70	1.11e-1	29.2	21.7	74.3
10	0.000	1.42e-1	6.9e-2	1.06e-1	6.3e-2	36.3	36.3	108.9	8.772	5.51e-2	26.7	18.2	117.8
11	48.13	1.22e+0	6.5e-1	1.49e+0	6.6e-1	37.3	36.3	7.2	95.37	7.37e-1	37.2	29.5	59.3
12	16.08	3.69e-1	1.4e-1	3.39e-1	1.3e-1	36.3	36.3	163.2	46.37	2.90e-1	29.8	19.4	179.5
13	6.232	2.07e-1	5.2e-2	3.66e-1	6.2e-2	50.3	36.3	91.0	10.62	1.21e-1	68.2	22.0	131.0
14	0.000	1.21e-1	4.4e-2	1.29e-1	4.0e-2	41.9	36.3	169.9	10.22	6.08e-2	37.8	19.8	177.5
15	6.844	1.72e-1	5.6e-2	2.27e-1	5.4e-2	44.5	36.3	177.7	8.745	5.29e-2	42.2	26.8	164.4
16	0.000	5.20e-2	4.4e-2	6.57e-2	4.5e-2	36.5	36.3	168.5	0.000	2.71e-2	38.0	26.8	179.5
17	0.903	-1.75e-2	1.9e-2	-4.07e-2	1.8e-2	70.4	37.3	13.9	0.000	2.36e-3	61.4	40.5	115.6
18	6.874	2.05e-1	3.4e-2	9.38e-1	7.3e-2	80.2	62.0	162.6	11.26	9.32e-2	98.0	62.6	173.6
19	15.60	5.29e-1	9.6e-2	1.30e+0	1.3e-1	52.4	45.2	157.9	23.30	1.93e-1	44.2	36.3	177.3
20	0.000	7.19e-3	1.7e-2	1.00e-2	1.5e-2	41.7	36.3	69.8	0.000	2.94e-2	20.3	18.2	137.3
21	0.000	7.54e-3	9.1e-3	6.93e-3	8.3e-3	36.3	36.3	41.4	0.000	1.39e-2	29.8	18.2	172.7
22	8.784	2.30e-1	6.5e-2	2.70e-1	6.0e-2	43.7	36.3	145.7	15.43	8.89e-2	19.5	18.2	39.7
23	0.000	7.89e-2	7.4e-2	3.38e-2	6.8e-2	36.3	36.3	29.8	10.77	6.60e-2	23.6	18.2	25.3
24	0.000	1.76e-1	7.9e-2	1.05e-1	7.3e-2	36.3	36.3	162.4	12.03	7.16e-2	32.8	18.2	159.9
25	0.000	8.55e-2	8.7e-2	4.35e-2	8.0e-2	36.3	36.3	75.6	15.58	1.15e-1	29.7	18.2	173.8
26	16.71	4.12e-1	8.9e-2	4.87e-1	8.2e-2	49.0	36.3	155.4	30.22	2.09e-1	34.7	21.1	165.2
27	0.000	1.08e-1	5.7e-2	8.55e-2	5.2e-2	36.3	36.3	18.4	10.22	7.54e-2	31.9	22.8	36.9
28	20.64	7.07e-1	1.4e-1	1.26e+0	1.4e-1	54.7	36.8	148.2	35.64	3.22e-1	49.6	32.7	162.3
29	8.206	2.64e-1	5.5e-2	4.17e-1	5.8e-2	52.9	36.3	4.8	9.995	7.54e-2	43.4	20.2	178.2
30	6.578	2.10e-1	3.1e-2	1.08e+0	7.0e-2	73.1	67.8	52.3	7.433	5.57e-2	90.1	60.2	48.4
31	0.000	5.31e-2	1.6e-2	1.26e-2	1.5e-2	36.3	36.3	26.2	7.972	3.98e-2	28.1	18.2	31.4
32	14.20	3.46e-1	7.2e-2	4.33e-1	6.8e-2	46.0	36.3	168.7	23.42	1.48e-1	42.1	24.1	1.8
33	6.797	1.57e-1	5.1e-2	1.96e-1	4.9e-2	41.9	36.3	26.9	8.250	5.32e-2	42.3	26.0	24.7
34	19.11	6.48e-1	1.8e-1	6.85e-1	1.7e-1	36.3	36.3	179.8	46.28	3.06e-1	19.4	18.2	18.4
35	62.35	1.55e+0	1.1e-1	1.79e+0	1.0e-1	41.9	36.3	61.7	201.9	1.35e+0	22.2	18.2	7.9
36	9.273	4.32e-1	1.3e-1	3.84e-1	1.2e-1	36.3	36.3	3.5	46.62	3.46e-1	25.1	18.2	24.2
37	49.79	1.02e+0	1.1e-1	8.14e-1	1.0e-1	36.3	36.3	177.5	216.5	1.31e+0	18.2	18.2	156.2
38	16.50	1.92e+0	2.5e-1	2.19e+0	2.3e-1	36.3	36.3	130.7	46.25	8.62e-1	25.1	22.1	108.3
39	13.50	4.01e-1	9.3e-2	6.33e-1	9.7e-2	61.9	36.3	12.8	23.49	1.31e-1	59.8	21.9	19.6
40	63.93	2.06e+0	2.6e-1	3.35e+0	2.7e-1	48.6	36.3	90.7	144.1	1.32e+0	34.2	25.4	96.3
41	176.8	4.26e+0	3.5e-1	4.67e+0	4.3e-1	36.3	36.3	62.2	365.7	3.39e+0	27.4	21.8	27.0
42	0.000	7.94e-2	9.2e-2	3.69e-2	8.4e-2	37.4	36.3	165.2	19.11	1.33e-1	27.1	18.2	17.9
43	48.25	1.18e+0	4.3e-1	8.94e-1	3.9e-1	36.5	36.3	47.3	128.7	9.49e-1	26.3	19.3	52.7
44	24.28	6.43e-1	1.1e-1	8.15e-1	1.0e-1	47.0	36.3	148.0	35.30	3.06e-1	37.6	22.2	135.6
45	0.000	7.21e-2	2.9e-2	4.99e-2	2.7e-2	36.3	36.3	0.8	9.488	6.55e-2	31.5	18.2	6.4
46	0.000	5.60e-2	1.1e-1	1.42e-2	1.0e-1	36.3	36.3	106.9	25.94	1.75e-1	25.5	18.2	154.9
47	86.77	3.04e+0	6.3e-1	4.50e+0	5.9e-1	46.5	39.0	43.5	176.4	1.18e+0	37.7	24.4	0.5
48	0.000	4.86e-1	3.7e-1	1.68e-1	3.4e-1	36.3	36.3	152.1	99.82	5.74e-1	24.6	18.2	84.8
49	100.9	2.48e+0	6.0e-1	3.43e+0	5.5e-1	45.7	36.3	163.5	161.7	1.43e+0	25.2	18.2	156.5
50	9.532	2.26e-1	2.2e-1	2.16e-1	2.1e-1	56.0	36.3	15.7	18.92	1.27e-1	44.6	18.2	9.2
51	134.8	9.07e+0	1.0e+0	8.86e+0	1.0e+0	36.3	36.3	13.8	1049	7.86e+0	20.7	18.2	175.5
52	15.52	1.17e+0	5.5e-1	1.83e+0	5.0e-1	43.0	37.3	78.1	8.558	7.75e-2	35.0	18.2	177.6
53	7.803	2.93e-1	4.3e-2	4.65e-1	3.9e-2	52.4	37.5	111.5	10.71	7.90e-2	38.3	30.4	104.0
54	469.2	5.27e+0	7.9e-1	5.25e+0	7.2e-1	36.3	36.3	33.4	703.1	3.16e+0	18.2	18.2	130.6
55	573.4	1.12e+1	8.0e-1	1.18e+1	7.3e-1	36.3	36.3	19.1	1138	7.07e+0	25.2	18.2	13.3
56	11.39	3.81e-1	1.2e-1	3.31e-1	1.1e-1	42.8	36.3	173.6	34.16	2.36e-1	35.6	18.2	135.9
57	7.179	1.38e-1	3.4e-2	1.44e-1	3.2e-2	37.0	36.3	43.9	9.986	6.91e-2	28.4	22.5	30.4
58	51.54	1.34e+0	4.7e-1	1.56e+0	4.3e-1	38.6	36.3	77.5	57.46	3.78e-1	34.9	18.2	160.0
59	12.69	7.14e-1	2.2e-1	8.74e-1	2.0e-1	44.7	37.8	117.9	49.84	4.53e-1	28.3	18.2	150.0
60	0.000	1.40e-1	2.7e-1	-2.54e-1	2.5e-1	52.2	36.3	23.8	25.59	2.11e-1	18.8	18.2	118.9
61	17.24	5.64e-1	4.5e-2	1.43e+0	6.5e-2	55.5	46.0	151.7	22.63	2.31e-1	45.7	36.5	167.9
62	124.1	2.59e+0	4.9e-1	2.97e+0	4.5e-1	40.9	36.3	34.3	79.09	2.52e-1	21.0	18.2	76.7
63	0.000	6.54e-2	8.4e-2	6.50e-3	7.7e-2	36.4	36.3	94.8	14.35	9.45e-2	35.5	18.2	174.4
64	0.000	4.96e-2	1.5e-2	1.98e-2	1.3e-2	36.3	36.3	112.3	9.256	5.38e-2	19.0	18.2	2.3
65	10.44	2.71e-1	2.2e-2	3.00e-1	2.0e-2	41.2	36.3	43.5	14.91	1.01e-1	34.9	18.2	35.4
66	0.000	1.49e-1	4.5e-2	1.01e-1	4.1e-2	36.3	36.3	153.8	11.12	7.76e-2	45.0	18.2	143.5
67	7.217	1.85e-1	4.7e-2	2.37e-1	4.5e-2	52.1	36.3	17.3	10.29	6.33e-2	40.1	21.6	169.5
68	0.000	6.05e-2	7.0e-2	1.52e-2	6.5e-2	36.3	36.3	93.5	13.51	9.51e-2	18.2	18.2	38.8
69	0.000	4.42e-3	2.7e-2	1.53e-3	2.5e-2	36.3	36.3	129.3	0.000	3.01e-2	18.3	18.2	128.8
70	0.000	6.12e-2	6.1e-2	3.65e-2	5.6e-2	36.3	36.3	175.3	8.139	5.05e-2	23.3	20.6	164.5
71	7.228	3.73e-1	6.4e-2	6.18e-1	5.9e-2	47.5	40.4	99.4	6.048	1.01e-1	43.0	31.7	109.9
72	44.90	1.46e+0	1.5e-1	1.79e+0	1.4e-1	45.2	36.3	81.3	122.3	8.54e-1	37.0	19.2	74.1

Table E.3. continued.

No.	Sig ₅₀₀	S_{500}^{peak} (Jy beam ⁻¹)		S_{500}^{tot} (Jy)		r_{500}^a (")	r_{500}^b (")	PA ₅₀₀ (°)	Sig _{N_{H2}}	$N_{\text{H}_2}^{\text{peak}}$ (Jy beam ⁻¹)		$r_{\text{H}_2}^a$ (")	$r_{\text{H}_2}^b$ (")	PA _{H₂} (°)
(38)	(39)	(40) ±	(41)	(42) ±	(43)	(44)	(45)	(46)	(47)	(48) ±	(49)	(50)	(51)	(52)
73	19.81	5.04e-1	5.4e-2	1.05e+0	7.0e-2	60.1	36.3	100.0	35.28	2.43e-1	46.3	25.2	90.9	
74	6.735	2.61e-1	5.6e-2	3.64e-1	5.1e-2	43.7	38.1	118.7	7.408	1.08e-1	23.9	20.7	43.1	
75	0.000	5.35e-3	2.3e-2	1.41e-3	2.1e-2	36.3	36.3	106.8	3.413	-9.23e-3	33.1	18.7	70.4	
76	13.44	3.74e-1	1.3e-1	5.99e-1	1.3e-1	54.1	37.4	137.3	28.05	1.41e-1	3.4e-2	49.6	26.8	146.9
77	9.046	1.66e-1	2.5e-2	1.51e-1	2.3e-2	36.3	36.3	71.3	15.30	9.70e-2	1.1e-2	25.7	24.6	64.6
78	10.13	3.74e-1	8.2e-2	6.34e-1	8.0e-2	51.5	41.1	75.4	12.41	9.02e-2	2.0e-2	46.0	25.9	85.0
79	18.57	6.69e-1	1.1e-1	8.91e-1	1.0e-1	43.3	36.3	174.1	43.75	3.14e-1	2.7e-2	28.2	18.5	7.9
80	0.000	8.59e-2	5.0e-2	2.91e-2	4.6e-2	36.3	36.3	20.9	16.30	1.00e-1	3.3e-2	18.2	18.2	14.0
81	8.246	2.13e-1	3.5e-2	2.57e-1	3.5e-2	37.3	36.3	170.2	13.42	9.48e-2	1.7e-2	35.5	24.9	3.7
82	8.907	1.91e-1	8.1e-2	1.95e-1	7.4e-2	53.0	36.3	164.2	20.89	1.62e-1	5.7e-2	44.6	18.2	168.5
83	18.16	4.60e-1	5.4e-2	5.59e-1	5.2e-2	41.5	36.3	160.6	28.84	2.36e-1	2.7e-2	30.7	18.2	164.7
84	9.397	2.36e-1	6.8e-2	4.64e-1	7.6e-2	62.0	37.9	32.1	13.37	9.20e-2	1.9e-2	52.6	45.3	22.3
85	6.658	1.58e-1	5.0e-2	1.98e-1	5.0e-2	52.6	36.3	135.1	9.321	6.20e-2	1.6e-2	46.0	28.3	146.3
86	7.215	1.53e-1	2.2e-2	1.57e-1	2.1e-2	45.9	36.3	20.4	12.03	8.09e-2	7.3e-3	41.0	18.2	17.2
87	11.89	4.26e-1	8.3e-2	1.29e+0	1.3e-1	66.1	48.7	169.8	20.37	7.07e-2	3.8e-2	90.3	26.4	2.6
88	8.953	2.02e-1	2.0e-2	1.65e-1	1.8e-2	36.3	36.3	124.2	19.58	1.29e-1	7.0e-3	18.5	18.2	156.2
89	0.000	7.67e-2	2.1e-2	9.30e-2	1.9e-2	49.3	36.3	164.9	7.733	4.69e-2	9.3e-3	39.4	18.2	3.7
90	11.77	2.56e-1	5.2e-2	1.89e-1	4.7e-2	36.3	36.3	5.3	30.45	1.94e-1	8.8e-3	18.4	18.2	171.8
91	12.83	3.27e-1	8.9e-2	3.66e-1	8.2e-2	46.6	36.3	94.6	20.68	1.28e-1	5.1e-2	29.1	20.9	63.9
92														
93	0.000	1.42e-1	3.9e-2	7.62e-2	3.5e-2	36.3	36.3	60.8	17.08	1.03e-1	1.6e-2	27.2	18.2	67.6
94	0.000	2.38e-1	7.9e-2	1.81e-1	7.3e-2	36.3	36.3	107.0	10.62	8.12e-2	1.7e-2	28.7	18.2	104.4
95	0.000	1.34e-1	5.6e-2	8.71e-2	5.1e-2	36.3	36.3	79.0	13.52	7.73e-2	1.4e-2	26.8	19.0	55.1
96	0.000	6.13e-2	4.8e-2	3.85e-2	4.4e-2	45.5	36.3	38.0	6.256	4.96e-2	1.9e-2	37.7	20.7	33.4
97	21.61	5.16e-1	5.7e-2	7.41e-1	5.7e-2	48.7	36.3	47.1	33.61	2.39e-1	1.4e-2	43.9	26.6	46.5
98	42.13	9.73e-1	5.8e-2	1.59e+0	7.5e-2	63.5	36.3	17.4	71.61	4.89e-1	4.0e-2	53.8	25.2	4.7
99	12.87	3.36e-1	1.3e-1	3.20e-1	1.2e-1	36.3	36.3	163.6	35.80	2.55e-1	3.4e-2	25.2	23.0	168.9
100	21.33	4.81e-1	9.3e-2	4.08e-1	8.5e-2	36.3	36.3	158.2	50.61	3.36e-1	1.5e-2	21.5	18.2	159.6
101	12.49	4.51e-1	5.8e-2	6.04e-1	5.6e-2	41.0	36.3	43.4	17.85	1.61e-1	2.3e-2	38.5	24.0	34.9
102	6.898	1.46e-1	6.2e-2	1.46e-1	5.7e-2	39.5	36.3	35.4	9.206	6.34e-2	9.8e-3	38.2	22.1	32.4
103	27.23	7.39e-1	1.4e-1	1.45e+0	1.5e-1	56.8	38.6	146.9	42.69	3.69e-1	5.5e-2	53.0	39.3	150.4
104	99.20	2.63e+0	3.0e-1	3.56e+0	2.8e-1	45.8	36.3	179.5	226.3	1.58e+0	1.0e-1	35.0	25.0	173.9
105	8.023	1.91e-1	1.3e-1	2.05e-1	1.2e-1	47.1	36.3	166.1	18.83	9.04e-2	3.7e-2	37.8	19.6	21.0
106	10.21	2.34e-1	5.9e-2	2.94e-1	5.4e-2	52.2	36.3	149.4	14.43	8.62e-2	1.3e-2	37.5	19.2	143.8
107	0.000	1.11e-1	1.0e-1	1.38e-1	1.0e-1	51.8	45.0	12.4	6.828	4.71e-2	3.3e-2	40.6	22.4	63.4
108	9.185	1.89e-1	4.6e-2	1.65e-1	4.2e-2	51.6	36.3	89.1	16.70	1.01e-1	1.4e-2	22.7	18.2	154.3
109	7.702	1.86e-1	9.3e-2	6.71e-1	1.4e-1	76.9	53.5	37.9	6.027	4.39e-2	1.1e-2	73.2	38.9	44.1
110	8.644	2.06e-1	7.1e-2	2.38e-1	6.8e-2	41.6	36.3	114.8	11.90	8.52e-2	9.1e-3	33.4	23.0	106.7
111	26.11	7.15e-1	3.8e-2	1.15e+0	4.0e-2	57.7	36.3	48.2	41.43	2.71e-1	3.4e-2	52.2	23.3	49.1
112	15.22	4.70e-1	5.4e-2	1.47e+0	9.1e-2	63.3	57.9	108.9	25.79	1.32e-1	3.0e-2	83.9	27.1	97.4
113	9.524	2.24e-1	8.6e-2	2.40e-1	7.9e-2	38.3	36.3	104.1	17.93	1.29e-1	2.1e-2	32.1	24.1	102.9
114	12.86	3.18e-1	8.4e-2	5.49e-1	7.9e-2	56.9	42.0	84.4	16.66	1.19e-1	3.4e-2	35.4	22.0	24.2
115	10.79	2.98e-1	7.8e-2	4.32e-1	7.2e-2	51.2	37.2	128.5	15.52	9.47e-2	1.7e-2	33.4	24.1	153.1
116	0.000	9.31e-3	2.5e-2	2.41e-3	2.3e-2	36.3	36.3	90.3	0.000	1.32e-3	4.6e-3	18.2	18.2	0.0
117	7.231	1.75e-1	7.0e-2	3.85e-1	8.1e-2	60.7	44.7	103.7	7.703	5.27e-2	1.7e-2	62.9	57.1	155.4
118	26.48	6.79e-1	1.5e-1	9.47e-1	1.4e-1	45.8	36.3	162.8	44.11	2.42e-1	5.7e-2	50.2	22.6	150.8
119	6.884	1.83e-1	3.2e-2	3.80e-1	4.4e-2	52.9	43.5	171.2	6.132	3.40e-2	6.0e-3	69.1	34.8	175.4
120	18.91	5.52e-1	1.3e-1	1.32e+0	1.6e-1	65.2	40.6	141.1	22.58	1.89e-1	2.5e-2	56.0	18.2	137.2
121	0.000	6.21e-2	7.4e-2	3.58e-2	6.8e-2	36.3	36.3	58.2	7.617	4.84e-2	2.1e-2	21.9	18.2	167.0
122	26.72	6.67e-1	1.0e-1	1.08e+0	1.2e-1	45.2	36.3	152.0	37.36	2.92e-1	3.1e-2	43.2	36.5	151.0
123	0.000	7.10e-2	4.6e-2	5.52e-2	4.2e-2	36.3	36.3	102.5	8.009	5.09e-2	1.3e-2	33.4	21.7	93.9
124	7.701	1.76e-1	1.0e-1	2.27e-1	1.0e-1	49.0	36.3	46.8	7.949	6.69e-2	2.9e-2	44.2	25.6	40.6
125	9.803	2.76e-1	9.8e-2	3.61e-1	9.4e-2	54.4	36.3	157.6	13.94	8.93e-2	2.2e-2	50.4	18.2	147.3
126	8.046	1.86e-1	1.9e-2	1.86e-1	1.8e-2	36.3	36.3	94.3	10.44	6.27e-2	5.5e-3	20.1	18.2	173.3
127	8.752	3.29e-1	1.3e-1	1.16e+0	2.1e-1	73.6	49.7	32.2	6.356	4.96e-2	1.2e-2	68.4	35.5	33.2
128	16.25	6.17e-1	1.6e-1	4.01e+0	3.7e-1	94.8	73.9	141.1	9.796	8.49e-2	2.1e-2	113.8	59.7	146.1
129	9.889	2.54e-1	6.4e-2	3.07e-1	6.0e-2	52.9	36.3	138.3	6.717	4.81e-2	8.4e-3	42.2	19.9	145.6
130	11.09	2.70e-1	1.2e-1	3.80e-1	1.1e-1	43.9	36.3	77.9	10.46	9.09e-2	1.5e-2	40.0	33.1	94.0
131	0.000	7.04e-2	1.1e-1	4.03e-2	1.0e-1	37.9	36.3	102.3	10.97	6.32e-2	3.0e-2	23.4	18.6	83.2
132	8.377	2.18e-1	3.3e-2	4.09e-1	3.8e-2	64.1	36.3	158.3	6.509	3.25e-2	7.3e-3	75.9	26.1	146.2
133	9.067	2.97e-1	3.5e-2	1.40e+0	6.6e-2	87.0	59.4	171.1	7.485	7.28e-2	1.1e-2	89.8	56.8	171.8
134	0.000	1.88e-3	1.0e-1	-2.34e-2	9.3e-2	36.3	36.3	7.4	7.662	6.23e-2	4.2e-2	41.0	18.2	129.4
135	0.000	4.73e-2	2.2e-2	2.07e-2	2.0e-2	36.3	36.3	164.1	0.000	1.96e-2	8.8e-3	22.3	18.2	133.7
136	0.000	1.14e-1	2.8e-2	1.40e-1	2.7e-2	55.4	36.3	173.1	6.970	5.13e-2	1.2e-2	56.7	26.2	3.8
137	8.232	1.72e-1	5.0e-2	2.29e-1	5.1e-2	45.5	36.3	77.1	7.073	3.09e-2	4.9e-3	57.8	21.1	73.5
138	7.136	4.68e-1	8.6e-2	5.39e-1	7.9e-2	39.0	36.3	83.6	21.71	1.44e-1	2.8e-2	22.6	18.2	73.3
139	0.000	6.26e-3	3.9e-2	1.49e-3	3.6e-2	36.3	36.3	175.0	0.000	1.56e-3	7.1e-3	20.5	18.2	6.4
140	8.812	2.23e-1	7.9e-2	2.31e-1	7.2e-2	51.0	36.3	146.1	9.515	6.09e-2	1.2e-2	46.8	20.0	150.6
141	6.779	1.84e-1	6.1e-2	2.35e-1	5.7e-2	49.4	36.3	32.1	7.095	5.46e-2	1.0e-2	32.9	25.1	27.9
142	0.000	1.28e-2	4.6e-2	3.67e-3	4.2e-2	36.3	36.3	93.4	0.000	7.68e-3	9.6e-3	18.2	18.2	3.6
143	6.948	1.78e-1	5.1e-2	2.45e-1	5.2e-2	52.1	36.3	78.7	6.444	4.06e-2	1.3e-2	48.3	27.1	87.6
144	7.915	2.14e-1	9.1e-2	4.57e-1	1.0e-1	62.4	38.6	8.0	6.374	5.32e-2	1.6e-2	56.6	40.8	1.8

Table E.3. continued

No.	Sig ₅₀₀	S_{500}^{peak} (Jy beam ⁻¹)		S_{500}^{tot} (Jy)		r_{500}^a ($''$)	r_{500}^b ($''$)	PA ₅₀₀ ($^\circ$)	Sig _{N₂H₂}	$N_{\text{H}_2}^{\text{peak}}$ (Jy beam ⁻¹)	$r_{\text{H}_2}^a$ ($''$)	$r_{\text{H}_2}^b$ ($''$)	PA _{H₂} ($^\circ$)	
(38)	(39)	(40)	± (41)	(42)	± (43)	(44)	(45)	(46)	(47)	(48) ± (49)	(50)	(51)	(52)	
145	0.000	5.39e-2	4.8e-2	3.18e-2	4.4e-2	36.3	36.3	69.3	6.051	3.48e-2	1.4e-2	37.0	18.2	157.4
146	0.000	1.45e-1	6.8e-2	1.05e-1	6.2e-2	36.3	36.3	36.0	6.769	4.29e-2	1.7e-2	25.7	18.2	71.5
147	7.163	1.83e-1	2.3e-2	3.08e-1	2.5e-2	51.9	36.3	149.0	6.210	4.19e-2	5.0e-3	49.2	21.3	130.7
148	0.000	8.19e-2	5.3e-2	5.81e-2	4.8e-2	36.9	36.3	113.4	6.691	4.03e-2	6.9e-3	27.0	18.2	120.1
149	30.61	1.05e+0	4.4e-1	2.01e+0	5.7e-1	46.8	43.2	59.4	6.124	1.11e-1	3.7e-2	45.7	36.1	64.3
150	20.05	4.55e-1	1.4e-1	5.44e-1	1.4e-1	39.8	36.3	10.5	23.23	1.75e-1	9.2e-3	34.1	27.8	3.5
151	6.347	1.42e-1	3.3e-2	2.06e-1	3.1e-2	54.0	36.3	162.9	6.085	2.81e-2	9.6e-3	55.1	24.1	144.5
152	0.000	5.76e-2	4.9e-2	2.74e-2	4.5e-2	36.3	36.3	80.6	1.284	-5.48e-3	4.2e-3	32.4	21.2	120.2
153	15.88	4.49e-1	1.4e-1	5.23e-1	1.3e-1	40.2	36.3	34.5	9.661	5.70e-2	1.7e-2	49.2	20.4	27.9
154	42.25	1.11e+0	2.1e-1	2.26e+0	2.5e-1	56.5	37.9	81.8	21.64	1.64e-1	1.1e-2	46.5	33.7	89.8
155	12.69	4.53e-1	2.5e-1	5.41e-1	2.3e-1	40.6	36.3	115.3	28.74	1.53e-1	1.0e-1	30.9	18.2	90.1
156	8.965	1.85e-1	9.9e-2	2.34e-1	9.8e-2	46.5	36.3	14.8	6.594	4.99e-2	6.3e-3	30.5	25.0	29.6
157	254.4	5.71e+0	2.3e-1	5.61e+0	2.1e-1	36.3	36.3	151.7	311.4	2.14e+0	1.2e-1	24.0	18.2	144.7
158	17.24	4.57e-1	7.6e-2	5.71e-1	7.0e-2	37.9	36.3	175.7	24.95	1.66e-1	1.2e-2	32.9	20.4	126.1
159	15.92	4.73e-1	1.2e-1	6.52e-1	1.2e-1	53.1	36.3	177.9	8.980	7.10e-2	8.3e-3	43.1	20.5	172.8
160	27.85	9.44e-1	5.2e-2	1.08e+0	4.8e-2	37.8	36.3	97.2	46.41	3.12e-1	2.9e-2	36.7	23.1	101.9
161	0.000	3.11e-2	2.3e-2	2.89e-2	2.1e-2	50.4	43.2	137.6	0.000	2.09e-3	3.5e-3	61.2	46.2	91.4
162	11.35	2.03e-1	7.8e-2	1.55e-1	7.2e-2	36.3	36.3	7.9	18.50	1.12e-1	2.2e-2	33.3	18.2	158.1
163	5.910	1.62e-1	9.0e-2	1.99e-1	8.9e-2	42.5	36.3	102.9	7.426	5.49e-2	1.9e-2	36.9	20.8	98.4
164	0.000	1.06e-1	4.9e-2	1.52e-1	4.5e-2	46.6	39.3	133.4	7.131	3.89e-2	1.2e-2	33.0	21.4	56.0
165	5.001	1.18e-1	1.7e-2	7.38e-2	1.5e-2	36.3	36.3	138.4	0.000	3.08e-2	4.0e-3	21.9	18.2	174.1
166	0.000	8.77e-5	7.1e-3	8.60e-3	6.9e-3	85.5	36.3	56.6	0.000	1.20e-3	2.4e-3	69.5	52.1	79.9
167	0.000	1.02e-1	8.2e-2	7.79e-2	7.6e-2	36.3	36.3	46.2	6.998	4.18e-2	1.1e-2	27.2	18.2	53.3
168	33.86	9.05e-1	1.9e-1	1.64e+0	2.0e-1	54.8	38.0	127.4	49.94	3.68e-1	3.9e-2	51.5	26.4	107.4
169	30.62	8.76e-1	3.3e-1	1.37e+0	3.3e-1	46.2	37.9	126.4	59.11	3.93e-1	8.4e-2	48.7	31.6	104.6
170	0.000	1.03e-1	4.6e-2	7.03e-2	4.2e-2	38.2	36.3	112.0	8.611	5.87e-2	1.5e-2	21.1	18.2	77.7