



| | |
|------------------------------|---|
| Publication Year | 2011 |
| Acceptance in OA@INAF | 2023-01-20T15:32:21Z |
| Title | Insights into the high-energy γ -ray emission of Markarian 421 from multifrequency observations in the Fermi era |
| Authors | Abdo, A.A.; Ackermann, M.; Ajello, M.; Allafort, A.; Baldini, L.; et al. |
| DOI | 10.1088/0004-637X/727/2/129 |
| Handle | http://hdl.handle.net/20.500.12386/32960 |
| Journal | THE ASTROPHYSICAL JOURNAL |
| Number | 727 |

INSIGHTS INTO THE HIGH-ENERGY γ -RAY EMISSION OF MARKARIAN 501 FROM EXTENSIVE MULTIFREQUENCY OBSERVATIONS IN THE *FERMI* ERA

- A. A. ABDO^{1,2}, M. ACKERMANN³, M. AJELLO³, A. ALLAFORT³, L. BALDINI⁴, J. BALLE⁵, G. BARBIELLINI^{6,7}, M. G. BARING⁸, D. BASTIERI^{9,10}, K. BECHTOL³, R. BELLAZZINI⁴, B. BERENJI³, R. D. BLANDFORD³, E. D. BLOOM³, E. BONAMENTE^{11,12}, A. W. BORGLAND³, A. BOUVIER³, T. J. BRANDT^{13,14}, J. BREGEON⁴, A. BREZ⁴, M. BRIGIDA^{15,16}, P. BRUEL¹⁷, R. BUEHLER³, S. BUSON^{9,10}, G. A. CALIANDRO¹⁸, R. A. CAMERON³, A. CANNON^{19,20}, P. A. CARAVEO²¹, S. CARRIGAN¹⁰, J. M. CASANDJIAN⁵, E. CAVAZZUTI²², C. CECCHI^{11,12}, Ö. ÇELİK^{19,23,24}, E. CHARLES³, A. CHEKHTMAN^{1,25}, C. C. CHEUNG^{1,2}, J. CHIANG³, S. CIPRINI¹², R. CLAUS³, J. COHEN-TANUGI²⁶, J. CONRAD^{27,28,133}, S. CUTINI²², C. D. DERMER¹, F. DE PALMA^{15,16}, E. DO COUTO E SILVA³, P. S. DRELL³, R. DUBOIS³, D. DUMORA²⁹, C. FAVUZZI^{15,16}, S. J. FEGAN¹⁷, E. C. FERRARA¹⁹, W. B. FOCKE³, P. FORTIN¹⁷, M. FRAILIS^{30,31}, L. FUHRMANN³², Y. FUKAZAWA³³, S. FUNK³, P. FUSCO^{15,16}, F. GARGANO¹⁶, D. GASPARRINI²², N. GEHRELS¹⁹, S. GERMANI^{11,12}, N. GIGLIETTO^{15,16}, F. GIORDANO^{15,16}, M. GIROLETTI³⁴, T. GLANZMAN³, G. GODFREY³, I. A. GRENIER⁵, L. GUILLEMOT^{29,32}, S. GUIRIEC³⁵, M. HAYASHIDA³, E. HAYS¹⁹, D. HORAN¹⁷, R. E. HUGHES¹⁴, G. JÓHANNESSON³, A. S. JOHNSON³, W. N. JOHNSON¹, M. KADLER^{23,36,37,38}, T. KAMAE³, H. KATAGIRI³³, J. KATAOKA³⁹, J. KNÖDLSER¹³, M. KUSS⁴, J. LANDE³, L. LATRONICO⁴, S.-H. LEE³, M. LEMOINE-GOUMARD²⁹, F. LONGO^{6,7}, F. LOPARCO^{15,16}, B. LOTT²⁹, M. N. LOVELLETTE¹, P. LUBRANO^{11,12}, G. M. MADEJSKI³, A. MAKEEV^{1,25}, W. MAX-MOERBECK⁴⁰, M. N. MAZZIOTTA¹⁶, J. E. MCENERY^{19,41}, J. MEHAULT²⁶, P. F. MICHELSON³, W. MITTHUMSIRI³, T. MIZUNO³³, A. A. MOISEEV^{23,41}, C. MONTE^{15,16}, M. E. MONZANI³, A. MORSELLI⁴², I. V. MOSKALENKO³, S. MURGIA³, M. NAUMANN-GODO⁵, S. NISHINO³³, P. L. NOLAN³, J. P. NORRIS⁴³, E. NUSS²⁶, T. OHSUGI⁴⁴, A. OKUMURA⁴⁵, N. OMODEI³, E. ORLANDO⁴⁶, J. F. ORMES⁴³, D. PANEQUE^{3,47,134}, J. H. PANETTA³, D. PARENT^{1,25}, V. PAVLIDOU⁴⁰, T. J. PEARSON⁴⁰, V. PELASSA²⁶, M. PEPE^{11,12}, M. PESCE-ROLLINS⁴, F. PIRON²⁶, T. A. PORTER³, S. RAINO^{15,16}, R. RANDO^{9,10}, M. RAZZANO⁴, A. READHEAD⁴⁰, A. REIMER^{3,48}, O. REIMER^{3,48}, J. L. RICHARDS⁴⁰, J. RIPKEN^{27,28}, S. RITZ⁴⁹, M. ROTH⁵⁰, H. F.-W. SADROZINSKI⁴⁹, D. SANCHEZ¹⁷, A. SANDER¹⁴, J. D. SCARGLE⁵¹, C. SGRÒ⁴, E. J. SISKIND⁵², P. D. SMITH¹⁴, G. SPANDRE⁴, P. SPINELLI^{15,16}, Ł. STAWARZ^{45,53,134}, M. STEVENSON⁴⁰, M. S. STRICKMAN¹, K. V. SOKOLOVSKY^{32,54}, D. J. SUSON⁵⁵, H. TAKAHASHI⁴⁴, T. TAKAHASHI⁴⁵, T. TANAKA³, J. B. THAYER³, J. G. THAYER³, D. J. THOMPSON¹⁹, L. TIBALDO^{5,9,10,135}, D. F. TORRES^{18,56}, G. TOSTI^{11,12}, A. TRAMACERE^{3,57,58}, Y. UCHIYAMA³, T. L. USHER³, J. VANDENBROUCKE³, V. VASILEIOU^{23,24}, N. VILCHEZ¹³, V. VITALE^{42,59}, A. P. WAITE³, P. WANG³, A. E. WEHRLE⁶⁰, B. L. WINER¹⁴, K. S. WOOD¹, Z. YANG^{27,28}, T. YLINEN^{28,61,62}, J. A. ZENSUS³², M. ZIEGLER⁴⁹
- (THE *Fermi*-LAT COLLABORATION)
- J. ALEKSIĆ⁶³, L. A. ANTONELLI⁶⁴, P. ANTORANZ⁶⁵, M. BACKES⁶⁶, J. A. BARRIO⁶⁷, J. BECERRA GONZÁLEZ^{68,69}, W. BEDNAREK⁷⁰, A. BERDYUGIN⁷¹, K. BERGER⁶⁹, E. BERNARDINI⁷², A. BILAND⁷³, O. BLANCH⁶³, R. K. BOCK⁴⁷, A. BOLLER⁷³, G. BONNOLI⁶⁴, P. BORDAS⁷⁴, D. BORLA TRIDON⁴⁷, V. BOSCH-RAMON⁷⁴, D. BOSE⁶⁷, I. BRAUN⁷³, T. BRETZ⁷⁵, M. CAMARA⁶⁷, E. CARMONA⁴⁷, A. CAROSI⁶⁴, P. COLIN⁴⁷, E. COLOMBO⁶⁸, J. L. CONTRERAS⁶⁷, J. CORTINA⁶³, S. COVINO⁶⁴, F. DAZZI^{30,136}, A. DE ANGELIS³⁰, E. DE CEA DEL POZO¹⁸, B. DE LOTTO⁷⁶, M. DE MARIA⁷⁶, F. DE SABATA⁷⁶, C. DELGADO MENDEZ^{68,77}, A. DIAGO ORTEGA^{68,69}, M. DOERT⁶⁶, A. DOMÍNGUEZ⁷⁸, D. DOMINIS PRESTER⁷⁹, D. DORNER⁷³, M. DORO^{9,10}, D. ELSAESSER⁷⁵, D. FERENC⁷⁹, M. V. FONSECA⁶⁷, L. FONT⁸⁰, R. J. GARCÍA LÓPEZ^{68,69}, M. GARCZARCYK⁶⁸, M. GAUG⁶⁸, G. GIVITTO⁶³, N. GODINOV⁷⁹, D. HADASCH¹⁸, A. HERRERO^{68,69}, D. HILDEBRAND⁷³, D. HÖHNE-MÖNCH⁷⁵, J. HOSE⁴⁷, D. HRUPEC⁷⁹, T. JOGLER⁴⁷, S. KLEPSE⁶³, T. KRÄHENBÜHL⁷³, D. KRANICH⁷³, J. KRAUSE⁴⁷, A. LA BARBERA⁶⁴, E. LEONARDO⁶⁵, E. LINDFORS⁷¹, S. LOMBARDI^{9,10}, M. LÓPEZ^{9,10}, E. LORENZ^{73,47}, P. MAJUMDAR⁷², E. MAKARIEV⁸¹, G. MANEVA⁸¹, N. MANKUZHIYIL³⁰, K. MANNHEIM⁷⁵, L. MARASCHI⁸², M. MARIOTTI^{9,10}, M. MARTÍNEZ⁶³, D. MAZIN⁶³, M. MEUCCI⁶⁵, J. M. MIRANDA⁶⁵, R. MIRZOYAN⁴⁷, H. MIYAMOTO⁴⁷, J. MOLDÓN⁷⁴, A. MORALES⁶³, D. NIETO⁶⁷, K. NILSSON⁷¹, R. ORITO⁴⁷, I. OYA⁶⁷, R. PAOLETTI⁶⁵, J. M. PAREDES⁷⁴, S. PARTINI⁶⁵, M. PASANEN⁷¹, F. PAUSS⁷³, R. G. PEGNA⁶⁵, M. A. PEREZ-TORRES⁷⁸, M. PERSIC^{30,83}, J. PERUZZO^{9,10}, J. POCHON⁶⁸, P. G. PRADA MORONI⁶⁵, F. PRADA⁷⁸, E. PRANDINI^{9,10}, N. PUCHADES⁶³, I. PULJAK⁷⁹, T. REICHARDT⁶³, R. REINTHAL⁷¹, W. RHODE⁶⁶, M. RIBÓ⁷⁴, J. RICO^{56,63}, M. RISSI⁷³, S. RÜGAMER⁷⁵, A. SAGGION^{9,10}, K. SAITO⁴⁷, T. Y. SAITO⁴⁷, M. SALVATI⁶⁴, M. SÁNCHEZ-CONDE^{68,69}, K. SATALECKA⁷², V. SCALZOTTO^{9,10}, V. SCAPIN³⁰, C. SCHULTZ^{9,10}, T. SCHWEIZER⁴⁷, M. SHAYDUK⁴⁷, S. N. SHORE^{4,84}, A. SIERPOWSKA-BARTOSIK⁷⁰, A. SILLANPÄÄ⁷¹, J. SITAREK^{70,47}, D. SOBCHYNSKA⁷⁰, F. SPANIER⁷⁵, S. SPIRO⁶⁴, A. STAMERRA⁶⁵, B. STEINKE⁴⁷, J. STORZ⁷⁵, N. STRAH⁶⁶, J. C. STRUEBIG⁷⁵, T. SURIC⁷⁹, L. O. TAKALO⁷¹, F. TAVECCHIO⁸², P. TEMNIKOV⁸¹, T. TERZIĆ⁷⁹, D. TESCARO⁶³, M. TESHIMA⁴⁷, H. VANKOV⁸¹, R. M. WAGNER⁴⁷, Q. WEITZEL⁷³, V. ZABALZA⁷⁴, F. ZANDANEL⁷⁸, R. ZANIN⁶³
- (THE MAGIC COLLABORATION)
- V. A. ACCIARI⁸⁵, T. ARLEN⁸⁶, T. AUNE⁴⁹, W. BENBOW⁸⁵, D. BOLTUCH⁸⁷, S. M. BRADBURY⁸⁸, J. H. BUCKLEY⁸⁹, V. BUGAEV⁸⁹, A. CANNON²⁰, A. CESARINI⁹⁰, L. CIUPIK⁹¹, W. CUI⁹², R. DICKHERBER⁸⁹, M. ERRANDO⁹³, A. FALCONE⁹⁴, J. P. FINLEY⁹², G. FINNEGAN⁹⁵, L. FORTSON⁹¹, A. FURNISS⁴⁹, N. GALANTE⁸⁵, D. GALL⁹², G. H. GILLANDERS⁹⁰, S. GODAMBE⁹⁵, J. GRUBE⁹¹, R. GUENETTE⁹⁶, G. GYUK⁹¹, D. HANNA⁹⁶, J. HOLDER⁸⁷, D. HUANG⁹⁷, C. M. HUI⁹⁵, T. B. HUMENSKY⁹⁸, P. KAARET⁹⁹, N. KARLSSON⁹¹, M. KERTZMAN¹⁰⁰, D. KIEDA⁹⁵, A. KONOPELKO⁹⁷, H. KRAWCZYNSKI⁸⁹, F. KRENNRICH¹⁰¹, M. J. LANG⁹⁰, G. MAIER^{72,96}, S. McARTHUR⁸⁹, A. McCANN⁹⁶, M. McCUTCHEON⁹⁶, P. MORIARTY¹⁰², R. MUKHERJEE⁹³, R. ONG⁸⁶, A. N. OTTE⁴⁹, D. PANDEL⁹⁹, J. S. PERKINS⁸⁵, A. PICHEL¹⁰³, M. POHL^{72,104}, J. QUINN²⁰, K. RAGAN⁹⁶, L. C. REYES¹⁰⁵, P. T. REYNOLDS¹⁰⁶, E. ROACHE⁸⁵, H. J. ROSE⁸⁸, A. C. ROVERO¹⁰³, M. SCHROEDTER¹⁰¹, G. H. SEMBROSKI⁹², G. D. SENTURK¹⁰⁷, D. STEELE^{91,108}, S. P. SWORDY^{98,137}, G. TEŠIĆ⁹⁶, M. THEILING⁸⁵, S. THIBADEAU⁸⁹, A. VARLOTTA⁹², S. VINCENT⁹⁵, S. P. WAKELY⁹⁸, J. E. WARD²⁰

T. C. WEEKES⁸⁵, A. WEINSTEIN⁸⁶, T. WEISGARBER⁹⁸, D. A. WILLIAMS⁴⁹, M. WOOD⁸⁶, B. ZITZER⁹²

(THE VERITAS COLLABORATION)

M. VILLATA¹⁰⁹, C. M. RAITERI¹⁰⁹, H. D. ALLER¹¹⁰, M. F. ALLER¹¹⁰, A. A. ARKHAROV¹¹¹, D. A. BLINOV¹¹¹, P. CALCIDESE¹¹²,
W. P. CHEN¹¹³, N. V. EFIMOVA^{111,114}, G. KIMERIDZE¹¹⁵, T. S. KONSTANTINOVA¹¹⁴, E. N. KOPATSKAYA¹¹⁴, E. KOPTELOVA¹¹³,
O. M. KURTANIDZE¹¹⁵, S. O. KURTANIDZE¹¹⁵, A. LÄHTENMÄKI¹¹⁶, V. M. LARIONOV^{111,114,117}, E. G. LARIONOVA¹¹⁴,
L. V. LARIONOVA¹¹⁴, R. LIGUSTRI¹¹⁸, D. A. MOROZOVA¹¹⁴, M. G. NIKOLASHVILI¹¹⁴, L. A. SIGUA¹¹⁵, I. S. TROITSKY¹¹⁴,
E. ANGELAKIS³², M. CAPALBI²², A. CARRAMIÑANA¹¹⁹, L. CARRASCO¹¹⁹, P. CASSARO¹²⁰, E. DE LA FUENTE¹²¹, M. A. GURWELL¹²²,
Y. Y. KOVALEV^{32,54}, YU. A. KOVALEV⁵⁴, T. P. KRICHBAUM³², H. A. KRIMM^{23,38}, P. LETO¹²³, M. L. LISTER⁹², G. MACCAFERRI¹²⁴,
J. W. MOODY¹²⁵, Y. MORI¹²⁶, I. NESTORAS³², A. ORLATI¹²⁴, C. PAGANI¹²⁷, C. PACE¹²⁵, R. PEARSON III¹²⁵, M. PERRI²²,
B. G. PINER¹²⁸, A. B. PUSHKAREV^{32,111,129}, E. ROS^{32,130}, A. C. SADUN¹³¹, T. SAKAMOTO¹⁹, M. TORNIKOSKI¹¹⁶, Y. YATSU¹²⁶, AND
A. ZOOK¹³²

¹ Space Science Division, Naval Research Laboratory, Washington, DC 20375, USA² National Research Council Research Associate, National Academy of Sciences, Washington, DC 20001, USA³ W. W. Hansen Experimental Physics Laboratory, Kavli Institute for Particle Astrophysics and Cosmology, Department of Physics and SLAC National Accelerator Laboratory, Stanford University, Stanford, CA 94305, USA; dpneque@mppmu.mpg.de⁴ Istituto Nazionale di Fisica Nucleare, Sezione di Pisa, I-56127 Pisa, Italy⁵ Laboratoire AIM, CEA-IRFU/CNRS/Université Paris Diderot, Service d'Astrophysique, CEA Saclay, 91191 Gif sur Yvette, France⁶ Istituto Nazionale di Fisica Nucleare, Sezione di Trieste, I-34127 Trieste, Italy⁷ Dipartimento di Fisica, Università di Trieste, I-34127 Trieste, Italy⁸ Rice University, Department of Physics and Astronomy, MS-108, P. O. Box 1892, Houston, TX 77251, USA⁹ Istituto Nazionale di Fisica Nucleare, Sezione di Padova, I-35131 Padova, Italy¹⁰ Dipartimento di Fisica "G. Galilei," Università di Padova, I-35131 Padova, Italy¹¹ Istituto Nazionale di Fisica Nucleare, Sezione di Perugia, I-06123 Perugia, Italy¹² Dipartimento di Fisica, Università degli Studi di Perugia, I-06123 Perugia, Italy¹³ Centre d'Étude Spatiale des Rayonnements, CNRS/UPS, BP 44346, F-30128 Toulouse Cedex 4, France¹⁴ Department of Physics, Center for Cosmology and Astro-Particle Physics, The Ohio State University, Columbus, OH 43210, USA¹⁵ Dipartimento di Fisica "M. Merlin" dell'Università e del Politecnico di Bari, I-70126 Bari, Italy¹⁶ Istituto Nazionale di Fisica Nucleare, Sezione di Bari, 70126 Bari, Italy¹⁷ Laboratoire Leprince-Ringuet, École polytechnique, CNRS/IN2P3, Palaiseau, France¹⁸ Institut de Ciències de l'Espai (IEEC-CSIC), Campus UAB, 08193 Barcelona, Spain¹⁹ NASA Goddard Space Flight Center, Greenbelt, MD 20771, USA²⁰ University College Dublin, Belfield, Dublin 4, Ireland²¹ INFN-Istituto di Astrofisica Spaziale e Fisica Cosmica, I-20133 Milano, Italy²² Agenzia Spaziale Italiana (ASI) Science Data Center, I-00044 Frascati (Roma), Italy²³ Center for Research and Exploration in Space Science and Technology (CRESST) and NASA Goddard Space Flight Center, Greenbelt, MD 20771, USA²⁴ Department of Physics and Center for Space Sciences and Technology, University of Maryland Baltimore County, Baltimore, MD 21250, USA²⁵ George Mason University, Fairfax, VA 22030, USA²⁶ Laboratoire de Physique Théorique et Astroparticules, Université Montpellier 2, CNRS/IN2P3, Montpellier, France²⁷ Department of Physics, Stockholm University, AlbaNova, SE-106 91 Stockholm, Sweden²⁸ The Oskar Klein Centre for Cosmoparticle Physics, AlbaNova, SE-106 91 Stockholm, Sweden²⁹ Université Bordeaux 1, CNRS/IN2P3, Centre d'Études Nucléaires de Bordeaux Gradignan, 33175 Gradignan, France³⁰ Dipartimento di Fisica, Università di Udine and Istituto Nazionale di Fisica Nucleare, Sezione di Trieste, Gruppo Collegato di Udine, I-33100 Udine, Italy³¹ Osservatorio Astronomico di Trieste, Istituto Nazionale di Astrofisica, I-34143 Trieste, Italy³² Max-Planck-Institut für Radioastronomie, Auf dem Hügel 69, 53121 Bonn, Germany³³ Department of Physical Sciences, Hiroshima University, Higashi-Hiroshima, Hiroshima 739-8526, Japan³⁴ INFN Istituto di Radioastronomia, 40129 Bologna, Italy³⁵ Center for Space Plasma and Aeronomic Research (CSPAR), University of Alabama in Huntsville, Huntsville, AL 35899, USA³⁶ Dr. Remeis-Sternwarte Bamberg, Sternwartstrasse 7, D-96049 Bamberg, Germany³⁷ Erlangen Centre for Astroparticle Physics, D-91058 Erlangen, Germany³⁸ Universities Space Research Association (USRA), Columbia, MD 21044, USA³⁹ Research Institute for Science and Engineering, Waseda University, 3-4-1, Okubo, Shinjuku, Tokyo 169-8555, Japan⁴⁰ Cahill Center for Astronomy and Astrophysics, California Institute of Technology, Pasadena, CA 91125, USA⁴¹ Department of Physics and Department of Astronomy, University of Maryland, College Park, MD 20742, USA⁴² Istituto Nazionale di Fisica Nucleare, Sezione di Roma "Tor Vergata," I-00133 Roma, Italy⁴³ Department of Physics and Astronomy, University of Denver, Denver, CO 80208, USA⁴⁴ Hiroshima Astrophysical Science Center, Hiroshima University, Higashi-Hiroshima, Hiroshima 739-8526, Japan⁴⁵ Institute of Space and Astronautical Science, JAXA, 3-1-1 Yoshinodai, Sagami-hara, Kanagawa 229-8510, Japan; stawarz@astro.isas.jaxa.jp⁴⁶ Max-Planck Institut für extraterrestrische Physik, 85748 Garching, Germany⁴⁷ Max-Planck-Institut für Physik, D-80805 München, Germany⁴⁸ Institut für Astro- und Teilchenphysik und Institut für Theoretische Physik, Leopold-Franzens-Universität Innsbruck, A-6020 Innsbruck, Austria⁴⁹ Santa Cruz Institute for Particle Physics, Department of Physics and Department of Astronomy and Astrophysics, University of California at Santa Cruz, Santa Cruz, CA 95064, USA⁵⁰ Department of Physics, University of Washington, Seattle, WA 98195-1560, USA⁵¹ Space Sciences Division, NASA Ames Research Center, Moffett Field, CA 94035-1000, USA⁵² NYCB Real-Time Computing Inc., Lattingtown, NY 11560-1025, USA⁵³ Astronomical Observatory, Jagiellonian University, 30-244 Kraków, Poland⁵⁴ Astro Space Center of the Lebedev Physical Institute, 117810 Moscow, Russia⁵⁵ Department of Chemistry and Physics, Purdue University Calumet, Hammond, IN 46323-2094, USA⁵⁶ Institució Catalana de Recerca i Estudis Avançats (ICREA), Barcelona, Spain⁵⁷ Consorzio Interuniversitario per la Fisica Spaziale (CIFS), I-10133 Torino, Italy⁵⁸ INTEGRAL Science Data Centre, CH-1290 Versoix, Switzerland⁵⁹ Dipartimento di Fisica, Università di Roma "Tor Vergata," I-00133 Roma, Italy⁶⁰ Space Science Institute, Boulder, CO 80301, USA

- ⁶¹ Department of Physics, Royal Institute of Technology (KTH), AlbaNova, SE-106 91 Stockholm, Sweden
⁶² School of Pure and Applied Natural Sciences, University of Kalmar, SE-391 82 Kalmar, Sweden
⁶³ Institut de Física d'Altes Energies (IFAE), Edifici Cn, Universitat Autònoma de Barcelona (UAB), E-08193 Bellaterra (Barcelona), Spain
⁶⁴ INAF National Institute for Astrophysics, I-00136 Roma, Italy
⁶⁵ Università di Siena and INFN Pisa, I-53100 Siena, Italy
⁶⁶ Technische Universität Dortmund, D-44221 Dortmund, Germany
⁶⁷ Universidad Complutense, E-28040 Madrid, Spain
⁶⁸ Instituto de Astrofísica de Canarias, E38205 - La Laguna (Tenerife), Spain
⁶⁹ Departamento de Astrofísica, Universidad de La Laguna, E-38205 La Laguna, Tenerife, Spain
⁷⁰ University of Łódź, PL-90236 Łódź, Poland
⁷¹ Tuorla Observatory, University of Turku, FI-21500 Piikkiö, Finland
⁷² Deutsches Elektronen Synchrotron DESY, D-15738 Zeuthen, Germany
⁷³ ETH Zurich, CH-8093 Zurich, Switzerland
⁷⁴ Universitat de Barcelona (ICC/IEED), E-08028 Barcelona, Spain
⁷⁵ Institut für Theoretische Physik and Astrophysik, Universität Würzburg, D-97074 Würzburg, Germany
⁷⁶ Istituto Nazionale di Fisica Nucleare, Sezione di Trieste, and Università di Trieste, I-34127 Trieste, Italy
⁷⁷ Centro de Investigaciones Energéticas, Medioambientales y Tecnológicas (CIEMAT), Madrid, Spain
⁷⁸ Instituto de Astrofísica de Andalucía, CSIC, E-18080 Granada, Spain
⁷⁹ Croatian MAGIC Consortium, Institute R. Bošković, University of Rijeka and University of Split, HR-10000 Zagreb, Croatia
⁸⁰ Universitat Autònoma de Barcelona, E-08193 Bellaterra, Spain
⁸¹ Institute for Nuclear Research and Nuclear Energy, BG-1784 Sofia, Bulgaria
⁸² INAF Osservatorio Astronomico di Brera, I-23807 Merate, Italy
⁸³ INAF Osservatorio Astronomico di Trieste, I-34143 Trieste, Italy
⁸⁴ Dipartimento di Fisica "Enrico Fermi," Università di Pisa, Pisa I-56127, Italy
⁸⁵ Fred Lawrence Whipple Observatory, Harvard-Smithsonian Center for Astrophysics, Amado, AZ 85645, USA
⁸⁶ Department of Physics and Astronomy, University of California, Los Angeles, CA 90095-1547, USA
⁸⁷ Department of Physics and Astronomy and the Bartol Research Institute, University of Delaware, Newark, DE 19716, USA
⁸⁸ School of Physics and Astronomy, University of Leeds, Leeds LS2 9JT, UK
⁸⁹ Department of Physics, Washington University, St. Louis, MO 63130, USA
⁹⁰ National University of Ireland Galway, University Road, Galway, Ireland
⁹¹ Adler Planetarium and Astronomy Museum, Chicago, IL 60605, USA
⁹² Department of Physics, Purdue University, West Lafayette, IN 47907, USA
⁹³ Department of Physics and Astronomy, Barnard College, Columbia University, New York, NY 10027, USA
⁹⁴ Department of Astronomy and Astrophysics, Pennsylvania State University, University Park, PA 16802, USA
⁹⁵ Department of Physics and Astronomy, University of Utah, Salt Lake City, UT 84112, USA
⁹⁶ Department of Physics, McGill University, Montreal, PQ, Canada H3A 2T8
⁹⁷ Department of Physics, Pittsburg State University, Pittsburg, KS 66762, USA
⁹⁸ Enrico Fermi Institute, University of Chicago, Chicago, IL 60637, USA
⁹⁹ Department of Physics and Astronomy, University of Iowa, Iowa City, IA 52242, USA
¹⁰⁰ Department of Physics and Astronomy, DePauw University, Greencastle, IN 46135-0037, USA
¹⁰¹ Department of Physics and Astronomy, Iowa State University, Ames, IA 50011-3160, USA
¹⁰² Department of Life and Physical Sciences, Galway-Mayo Institute of Technology, Galway, Ireland
¹⁰³ Instituto de Astronomía y Física del Espacio, Parbellón IAFE, Cdad. Universitaria, Buenos Aires, Argentina
¹⁰⁴ Institut für Physik und Astronomie, Universität Potsdam, 14476 Potsdam, Germany
¹⁰⁵ Kavli Institute for Cosmological Physics, University of Chicago, Chicago, IL 60637, USA
¹⁰⁶ Department of Applied Physics and Instrumentation, Cork Institute of Technology, Bishopstown, Cork, Ireland
¹⁰⁷ Columbia Astrophysics Laboratory, Columbia University, New York, NY 10027, USA
¹⁰⁸ Los Alamos National Laboratory, Los Alamos, NM 87545, USA
¹⁰⁹ INAF, Osservatorio Astronomico di Torino, I-10025 Pino Torinese (TO), Italy
¹¹⁰ Department of Astronomy, University of Michigan, Ann Arbor, MI 48109-1042, USA
¹¹¹ Pulkovo Observatory, 196140 St. Petersburg, Russia
¹¹² Osservatorio Astronomico della Regione Autonoma Valle d'Aosta, Italy
¹¹³ Graduate Institute of Astronomy, National Central University, Jhongli 32054, Taiwan
¹¹⁴ Astronomical Institute, St. Petersburg State University, St. Petersburg, Russia
¹¹⁵ Abastumani Observatory, Mt. Kanobili, 0301 Abastumani, Georgia
¹¹⁶ Aalto University Metsähovi Radio Observatory, FIN-02540 Kylmälä, Finland
¹¹⁷ Isaac Newton Institute of Chile, St. Petersburg Branch, St. Petersburg, Russia
¹¹⁸ Circolo Astrofili Talmassons, I-33030 Campofornido (UD), Italy
¹¹⁹ Instituto Nacional de Astrofísica, Óptica y Electrónica, Tonantzintla, Puebla 72840, Mexico
¹²⁰ INAF Istituto di Radioastronomia, Sezione di Noto, Contrada Renna Bassa, 96017 Noto (SR), Italy
¹²¹ Instituto de Astronomía y Meteorología, Dpto. de Física, CUCEI, Universidad de Guadalajara, Jalisco, Mexico
¹²² Harvard-Smithsonian Center for Astrophysics, Cambridge, MA 02138, USA
¹²³ Osservatorio Astrofisico di Catania, 95123 Catania, Italy
¹²⁴ INAF Istituto di Radioastronomia, Stazione Radioastronomica di Medicina, I-40059 Medicina (Bologna), Italy
¹²⁵ Department of Physics and Astronomy, Brigham Young University, Provo Utah 84602, USA
¹²⁶ Department of Physics, Tokyo Institute of Technology, Meguro City, Tokyo 152-8551, Japan
¹²⁷ Department of Physics and Astronomy, University of Leicester, Leicester, LE1 7RH, UK
¹²⁸ Department of Physics and Astronomy, Whittier College, Whittier, CA, USA
¹²⁹ Crimean Astrophysical Observatory, 98409 Nauchny, Crimea, Ukraine
¹³⁰ Universitat de València, 46010 València, Spain
¹³¹ Department of Physics, University of Colorado, Denver, CO 80220, USA
¹³² Department of Physics and Astronomy, Pomona College, Claremont CA 91711-6312, USA

Received 2010 August 18; accepted 2010 November 19; published 2011 January 12

ABSTRACT

We report on the γ -ray activity of the blazar Mrk 501 during the first 480 days of *Fermi* operation. We find that the average Large Area Telescope (LAT) γ -ray spectrum of Mrk 501 can be well described by a single power-law function with a photon index of 1.78 ± 0.03 . While we observe relatively mild flux variations with the *Fermi*-LAT (within less than a factor of two), we detect remarkable spectral variability where the hardest observed spectral index within the LAT energy range is 1.52 ± 0.14 , and the softest one is 2.51 ± 0.20 . These unexpected spectral changes do not correlate with the measured flux variations above 0.3 GeV. In this paper, we also present the first results from the 4.5 month long multifrequency campaign (2009 March 15–August 1) on Mrk 501, which included the Very Long Baseline Array (VLBA), *Swift*, *RXTE*, MAGIC, and VERITAS, the F-GAMMA, GASP-WEBT, and other collaborations and instruments which provided excellent temporal and energy coverage of the source throughout the entire campaign. The extensive radio to TeV data set from this campaign provides us with the most detailed spectral energy distribution yet collected for this source during its relatively low activity. The average spectral energy distribution of Mrk 501 is well described by the standard one-zone synchrotron self-Compton (SSC) model. In the framework of this model, we find that the dominant emission region is characterized by a size $\lesssim 0.1$ pc (comparable within a factor of few to the size of the partially resolved VLBA core at 15–43 GHz), and that the total jet power ($\simeq 10^{44}$ erg s $^{-1}$) constitutes only a small fraction ($\sim 10^{-3}$) of the Eddington luminosity. The energy distribution of the freshly accelerated radiating electrons required to fit the time-averaged data has a broken power-law form in the energy range 0.3 GeV–10 TeV, with spectral indices 2.2 and 2.7 below and above the break energy of 20 GeV. We argue that such a form is consistent with a scenario in which the bulk of the energy dissipation within the dominant emission zone of Mrk 501 is due to relativistic, proton-mediated shocks. We find that the ultrarelativistic electrons and mildly relativistic protons within the blazar zone, if comparable in number, are in approximate energy equipartition, with their energy dominating the jet magnetic field energy by about two orders of magnitude.

Key words: acceleration of particles – BL Lacertae objects: general – BL Lacertae objects: individual (Mrk 501) – galaxies: active – gamma rays: general – radiation mechanisms: non-thermal

Online-only material: color figures

1. INTRODUCTION

Blazars constitute a subclass of radio-loud active galactic nuclei (AGNs), in which a jet of magnetized plasma assumed to emanate with relativistic bulk velocity from close to a central supermassive black hole points almost along the line of sight. The broadband emission spectra of these objects are dominated by non-thermal, strongly Doppler-boosted, and variable radiation produced in the innermost part of the jet. Most of the identified extragalactic γ -ray sources detected with the EGRET instrument (Hartman et al. 1999) on board the *Compton Gamma Ray Observatory* belong to this category. Blazars include flat-spectrum radio quasars (FSRQs) and BL Lacertae objects (BL Lac objects). Even though blazars have been observed for several decades at different frequencies, the existing experimental data did not permit unambiguous identification of the physical mechanisms responsible for the production of their high-energy (γ -ray) emission. Given the existing high-sensitivity detectors which allow detailed study of the low-energy (synchrotron) component of blazar sources (extending from radio up to hard X-rays), one of the reasons for the incomplete understanding of those objects was only the moderate sensitivity of previous γ -ray instruments. This often precluded detailed cross-correlation studies between the low- and high-energy emission and did not provide enough constraints on the parameters of the theoretical models. Some

of the open and fundamental questions regarding blazar sources are (1) the content of their jets, (2) the location and structure of their dominant emission zones, (3) the origin of their variability, observed on timescales from minutes to tens of years, (4) the role of external photon fields (including the extragalactic background light, EBL) in shaping their observed γ -ray spectra, and (5) the energy distribution and the dominant acceleration mechanism for the underlying radiating particles.

The Large Area Telescope (LAT) instrument (Atwood et al. 2009) on board the *Fermi Gamma-ray Space Telescope* satellite provides a large improvement in the experimental capability for performing γ -ray astronomy, and hence it is shedding new light on the blazar phenomenon. In this paper, we report on the *Fermi* observations of the TeV-emitting high-frequency-peaked—or, according to a more recent classification (Abdo et al. 2010c), high-synchrotron-peaked (HSP)—BL Lac object Markarian 501 (Mrk 501; R.A. = 16^h45^m52^s.22, decl. = 39°45′36″.6, J2000, redshift $z = 0.034$), which is one of the brightest extragalactic sources in the X-ray/TeV sky. Mrk 501 was the second extragalactic object (after Markarian 421) identified as a very high energy (hereafter VHE) γ -ray emitter (Quinn et al. 1996; Bradbury et al. 1997). After a phase of moderate emission lasting for about a year following its discovery (1996), Mrk 501 went into a state of surprisingly high activity and strong variability, becoming >10 times brighter than the Crab Nebula at energies >1 TeV, as reported by various instruments/groups (Catanese et al. 1997; Samuelson et al. 1998; Aharonian et al. 1999a, 1999b, 1999c; Djannati-Ataï et al. 1999). In 1998–1999, the mean VHE γ -ray flux dropped by an order of magnitude, and the overall VHE spectrum softened significantly (Piron 2000; Aharonian et al. 2001). In 2005, γ -ray flux variability on minute timescales was observed in the VHE band, thus establishing Mrk 501 as one of the sources with the fastest γ -ray

¹³³ Royal Swedish Academy of Sciences Research Fellow, funded by a grant from the K. A. Wallenberg Foundation.

¹³⁴ Author to whom any correspondence should be addressed.

¹³⁵ Partially supported by the International Doctorate on Astroparticle Physics (IDAPP) program.

¹³⁶ Supported by INFN Padova.

¹³⁷ Deceased.

flux changes (Albert et al. 2007a). During the 2005 VHE flux variations (when Mrk 501 was three to four times dimmer than it was in 1997), significant spectral variability was detected as well, with a clear “harder when brighter” behavior. Those spectral variations are even more pronounced when compared with the spectrum measured during the low-activity level recently reported in Anderhub et al. (2009).

The spectral energy distribution (SED) and the multifrequency correlations of Mrk 501 have been intensively studied in the past (e.g., Pian et al. 1998; Villata & Raiteri 1999; Krawczynski et al. 2000; Sambruna et al. 2000; Tavecchio et al. 2001; Katarzyński et al. 2001; Ghisellini et al. 2002; Gliozzi et al. 2006; Anderhub et al. 2009), but the nature of this object is still far from being understood. The main reasons for this lack of knowledge are the sparse multifrequency data during long periods of time, and the moderate sensitivity available in the past to study the γ -ray emission of this source. Besides, most of the previous multifrequency campaigns were triggered by an enhanced flux level in some energy band, and hence much of our information about the source is biased toward “high-activity” states, where perhaps distinct physical processes play a dominant role. In addition, until now we knew very little about the GeV emission of Mrk 501. The only detection reported at GeV energies before *Fermi* was in Kataoka et al. (1999), but the significance of this detection was too low to include Mrk 501 in the third EGRET catalog (Hartman et al. 1999). Moreover, Mrk 501 was not detected by EGRET during the large X-ray and VHE γ -ray flare which lasted for several months in 1997 (Pian et al. 1998).

The large improvement in the performance provided by the *Fermi*-LAT compared with its predecessor, EGRET, provides us with a new perspective for the study of blazars like Mrk 501. However, it is important to emphasize that blazars can vary their emitted power by one or two orders of magnitude on short timescales, and that they emit radiation over the entire observable electromagnetic spectrum (from $\sim 10^{-6}$ eV up to $\sim 10^{13}$ eV). For this reason, the information from *Fermi*-LAT alone is not enough to understand the broadband emission of Mrk 501, and hence simultaneous data in other frequency ranges are required. In particular, the frequency ranges where the low- and high-energy spectral components peak in the SED representation are of major importance. In the case of Mrk 501, those peaks are typically located around 1 keV (low-energy bump) and 100 GeV (high-energy bump), and hence simultaneous UV/X-ray and GeV/TeV observations are essential for the proper reconstruction of the overall SED of Mrk 501. At TeV energies there has been a substantial improvement in the instrumental capability as a result of the deployment of a new generation of imaging atmospheric Cherenkov telescopes (IACTs). In particular, for the study of Mrk 501, the new telescope systems MAGIC and VERITAS provide greater sensitivity, wider energy range, and improved energy resolution compared with the previous generation of instruments. Simultaneous observations with *Fermi*-LAT and IACTs like MAGIC or VERITAS (potentially covering six decades in energy, from 20 MeV to 20 TeV) can, for the first time, significantly resolve both the rising and the falling segments of the high-energy emission component of Mrk 501, with the expected location of the SED peak in the overlapping energy range between those instruments. Because of the smaller collection area, and the self-veto problem,¹³⁸ the sensitivity of

EGRET to detect γ -rays with energies larger than 10 GeV was about two orders of magnitude lower than that of *Fermi*-LAT.¹³⁹ Besides, during the period of operation of EGRET, the sensitivity of the previous generation of IACTs was only moderate, with relatively low sensitivity below 0.5 TeV. Therefore, the higher sensitivity and larger energy range of the newer γ -ray instruments have become a crucial tool for studying Mrk 501, and the blazar phenomenon in general.

In order to exploit the performance of *Fermi*-LAT and the new IACTs, as well as the capabilities of several existing instruments observing at radio-to-X-ray frequencies, a multifrequency (from radio to TeV photon energies) campaign was organized to monitor Mrk 501 during a period of 4.5 months, from 2009 mid-March to August. The scientific goal was to collect a very complete, simultaneous, multifrequency data set that would allow current theoretical models of broadband blazar emission to be tested. This, in turn, should help us to understand the origin of high-energy emission of blazar sources and the physical mechanisms responsible for the acceleration of radiating particles in relativistic jets in general. In this paper, the only reported result from the multifrequency observations is the overall SED averaged over the duration of the observing campaign. A more in-depth analysis of the multifrequency data set will be given in a forthcoming paper. The scientific results from the data collected during the two-day time interval 2009 March 23–25 (which includes extensive observations with the *Suzaku* X-ray satellite) will be reported in a separate paper (V. A. Acciari et al. 2011, in preparation). The paper is organized as follows. In Section 2, we introduce the LAT instrument and describe the LAT data analysis. In Section 3, we report on the flux/spectral variability of Mrk 501 observed during the first 16 months of *Fermi*-LAT operation, and compare it with the flux variability observed in X-rays by the all-sky instruments *Rossi X-Ray Timing Explorer* (RXTE; Bradt et al. 1993) All Sky Monitor (ASM) and the *Swift* (Gehrels et al. 2004) Burst Alert Telescope (BAT). In Section 4, we analyze the γ -ray spectrum of Mrk 501 measured by *Fermi*-LAT in the energy range 0.1–400 GeV. Section 5 reports on the overall SED obtained during the 4.5 month long multifrequency campaign organized in 2009. Section 6 is devoted to SED modeling, the results of which are further discussed in Section 7. Conclusions are presented in Section 8.

2. FERMI-LAT DATA SELECTION AND ANALYSIS

The *Fermi*-LAT is an instrument that performs γ -ray astronomy above 20 MeV. The instrument is an array of 4×4 identical towers, each one consisting of a tracker (where the photons are pair-converted) and a calorimeter (where the energies of the pair-converted photons are measured). The entire instrument is covered with an anticoincidence detector to reject charged-particle background. The LAT has a peak effective area of 0.8 m² for 1 GeV photons, an energy resolution typically better than 10%, and an FoV of about 2.4 sr, with an angular resolution (68% containment angle) better than 1° for energies above 1 GeV. Further details on the LAT can be found in Atwood et al. (2009).

The LAT data reported in this paper were collected from 2008 August 5 (MJD 54683) to 2009 November 27 (MJD 55162). During this time, the *Fermi*-LAT instrument operated

¹³⁸ The self-veto problem in EGRET is the degradation of the effective area at high energies (>5 GeV) due to backscatter of secondary particles from the calorimeter causing the anticoincidence system to veto the event. This problem

is substantially reduced in the LAT by using a segmented anticoincidence detector.

¹³⁹ This estimate includes the larger exposure from *Fermi*-LAT due to the four times larger field of view (FoV).

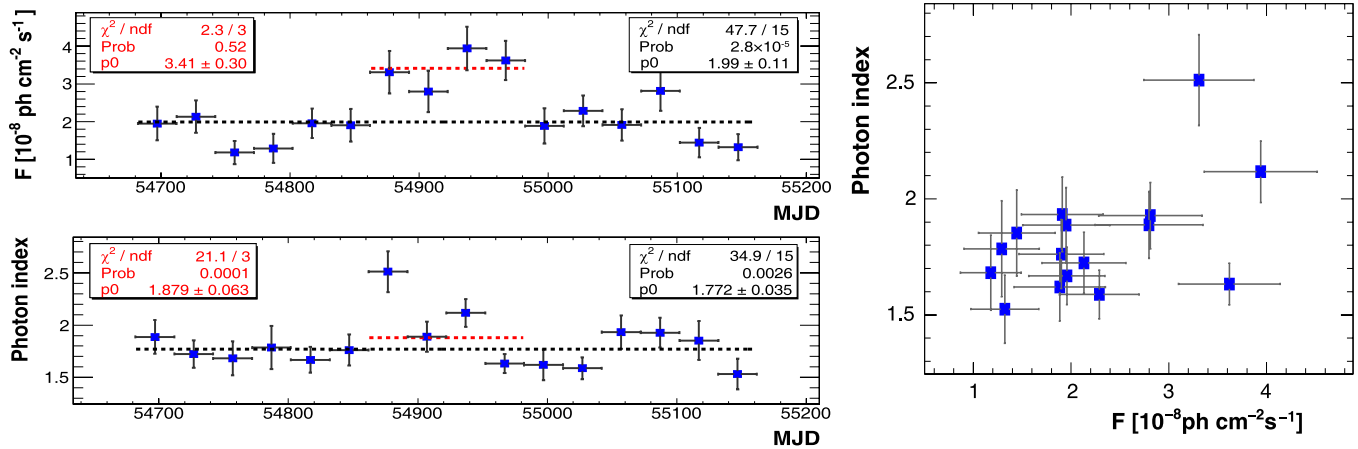


Figure 1. Left: *FERMI*-LAT γ -ray flux in the energy range 0.3–400 GeV (top panel) and spectral photon index from a power-law fit (bottom panel) for Mrk 501 for 30-day time intervals from 2008 August 5 (MJD 54683) to 2009 November 27 (MJD 55162). Vertical bars denote 1σ uncertainties and the horizontal bars denote the width of the time interval. The red dashed line and the red legend show the results from a constant fit to the time interval MJD 54862–54982, while the black dashed line and black legend show the results from a constant fit to the entire 480-day data set. Right: the scatter plot of the photon index vs. flux values.

mostly in survey mode. The analysis was performed with the Fermi Science Tools software package version v9r15p6. Only events with the highest probability of being photons—those in the “diffuse” class—were used. The LAT data were extracted from a circular region of 10° radius centered at the location of Mrk 501. The spectral fits were performed using photon energies in the energy range 0.3–400 GeV. At photon energies above 0.3 GeV, the effective area of the instrument is relatively large ($>0.5 \text{ m}^2$) and the angular resolution relatively good (68% containment angle smaller than 2°). In particular, because of the better angular resolution, the spectral fits using energies above 0.3 GeV (instead of 0.1 GeV) are less sensitive to possible contamination from unaccounted (perhaps transient), neighboring γ -ray sources and hence have smaller systematic errors, at the expense of reducing somewhat the number of photons from the source. In addition, a cut on the zenith angle ($>105^\circ$) was applied to reduce contamination from Earth-albedo γ -rays, which are produced by cosmic rays interacting with the upper atmosphere.

The background model used to extract the γ -ray signal includes a Galactic diffuse emission component and an isotropic component. The model that we adopted for the Galactic component is `gll_iem_v02.fit`.¹⁴⁰ The isotropic component, which is the sum of the extragalactic diffuse emission and the residual charged-particle background, is parameterized here with a single power-law function. To reduce systematic uncertainties in the analysis, the photon index of the isotropic component and the normalization of both components in the background model were allowed to vary freely during the spectral point fitting. Owing to the relatively small size of the region analyzed (radius 10°) and the hardness of the spectrum of Mrk 501, the high-energy structure in the standard tabulated isotropic background spectrum `isotropic_iem_v02.txt` does not dominate the total counts at high energies. In addition, we find that for this region a power-law approximation to the isotropic background results in somewhat smaller residuals for the overall model, possibly because the isotropic term, with a free spectral index, compensates for an inaccuracy in the model for the Galactic diffuse emission, which is also approximately isotropic at the high Galactic latitude of Mrk 501 ($b \sim 39^\circ$). In any case, the resulting spectral fits for Mrk 501 are not significantly different

if `isotropic_iem_v02.txt` is used for the analysis. In addition, the model also includes five nearby sources from the 1FGL catalog (Abdo et al. 2010b): 1FGL J1724.0+4002, 1FGL J1642.5+3947, 1FGL J1635.0+3808, 1FGL J1734.4+3859, and 1FGL J1709.6+4320. The spectra of those sources were also parameterized by a power-law functions, whose photon index values were fixed to the values from the 1FGL catalog, and only the normalization factors for the single sources were left as free parameters. The spectral analysis was performed with the post-launch instrument-response functions `P6_V3_DIFFUSE` using an unbinned maximum-likelihood method (Mattox et al. 1996). The systematic uncertainties on the flux were estimated as 10% at 0.1 GeV, 5% at 560 MeV and 20% at 10 GeV and above.¹⁴¹

3. FLUX AND SPECTRAL VARIABILITY

The high sensitivity and survey-mode operation of *Fermi*-LAT permit systematic, uninterrupted monitoring of Mrk 501 in γ -rays, regardless of the activity level of the source. The measured γ -ray flux above 0.3 GeV and the photon index from a power-law fit are shown in the left panel of Figure 1. The data span the time from 2008 August 5 (MJD 54683) to 2009 November 27 (MJD 55162), binned in time intervals of 30 days. The Test Statistic (TS) values¹⁴² for the 16 time intervals are all in excess of 50 (i.e., ~ 7 standard deviations, hereafter σ), with three-quarters of them greater than 100 (i.e., $\sim 10\sigma$). During this 480-day period, Mrk 501 did not show any outstanding flaring activity in the *Fermi*-LAT energy range, but there appear to be flux and spectral variations on timescales of the order of 30 days. During the 120-day period MJD 54862–54982, the photon flux above 0.3 GeV was $(3.41 \pm 0.28) \times 10^{-8} \text{ ph cm}^{-2} \text{ s}^{-1}$, which is about twice as large as the averaged flux values before and after that time period, which are $(1.65 \pm 0.16) \times 10^{-8} \text{ ph cm}^{-2} \text{ s}^{-1}$ and $(1.84 \pm 0.17) \times 10^{-8} \text{ ph cm}^{-2} \text{ s}^{-1}$, respectively. Remarkably, the photon index changed from 2.51 ± 0.20 for the first 30-day interval of this “enhanced-flux period” to 1.63 ± 0.09 for the last 30-day interval. As shown in the red legend of the bottom plot in the left panel of Figure 1, a constant fit to the photon

¹⁴¹ http://fermi.gsfc.nasa.gov/ssc/data/analysis/LAT_caveats.html

¹⁴² The TS value quantifies the probability of having a point γ -ray source at the location specified. It is roughly the square of the significance value: a TS of 25 would correspond to a signal of approximately five standard deviations (Mattox et al. 1996).

¹⁴⁰ <http://fermi.gsfc.nasa.gov/ssc/data/access/lat/BackgroundModels.html>

index values of this 120-day period gives a null probability of 10^{-4} , hence a deviation of 4σ . A constant fit to the entire 16-month period gives a null probability of 2.6×10^{-3} ; hence, spectral variability is detected for the entire data set at the level of 3σ . It is worth stressing that the spectral variability in the 480-day time interval is entirely dominated by the spectral variability occurring during the 120-day time interval of MJD 54862–54982, with no significant spectral variability before or after this “enhanced-flux period.” The right plot in Figure 1 does not show any clear correlation between the flux and the spectral variations. The discrete correlation function (DCF) computed as prescribed in Edelson & Krolik (1988) gives $\text{DCF} = 0.5 \pm 0.3$ for a time lag of zero.

Mrk 501 is known for showing spectral variability at VHE γ -ray energies. During the large X-ray/ γ -ray flare in 1997, Whipple and (especially) CAT observations showed evidence of spectral curvature and variability (Samuelson et al. 1998; Djannati-Ataï et al. 1999). The spectral changes are larger when comparing the measurements from 1997 with the low states from 1998 and 1999, as reported by CAT and HEGRA (Piron 2000; Aharonian et al. 2001). The MAGIC telescope, with lower energy threshold and higher sensitivity than the Whipple, HEGRA, and CAT telescopes, observed remarkable spectral variability in 2005, when the γ -ray activity of Mrk 501 was significantly smaller than that observed in 1997 (Albert et al. 2007a). The spectral variability is even larger when comparing the MAGIC measurements from 2005 with those from 2006 when the source was in an even lower state (Anderhub et al. 2009). However, despite the measured spectral variability at VHE γ -ray energies, the outstanding spectral steepening at GeV energies observed during the time interval MJD 54862–54892 was not envisioned in any of the previous works in the literature; the modeled spectrum of Mrk 501 at GeV energies was always assumed to be hard (photon indices ~ 1.5 – 1.8). This observational finding, further discussed in Sections 4 and 7, shows the importance of having a γ -ray instrument capable of long-term, uninterrupted, high-sensitivity monitoring of Mrk 501 and other HSP BL Lac objects, and it points to the important role *Fermi*-LAT will play in improving our understanding of the physics behind the blazar phenomenon.

The *Fermi*-LAT capability for continuous source monitoring is complemented at X-ray frequencies by *RXTE*-ASM and *Swift*-BAT, the two all-sky instruments that can probe the X-ray activity of Mrk 501 on a 30-day timescale. Figure 2 shows the fluxes measured by the ASM in the energy range 2–10 keV, by the BAT in the energy range 15–50 keV, and by the LAT in two different energy bands: 0.2–2 GeV (low-energy band) and >2 GeV (high-energy band).¹⁴³ The data from *RXTE*-ASM were obtained from the ASM Web site.¹⁴⁴ The data were filtered according to the prescription provided there, and the weighted average over all of the dwells¹⁴⁵ was determined for the 30-day time intervals defined for the *Fermi* data. The data from *Swift*-BAT were gathered from the BAT Web site.¹⁴⁶ We retrieved the daily averaged BAT values and made the weighted average over all the days from the 30-day time intervals defined for the *Fermi* data. The X-ray intensity from Mrk 501, averaged over

the 16 months, is 0.25 ± 0.01 count s^{-1} per Scanning Shadow Camera in the ASM, and $(0.52 \pm 0.05) \times 10^{-3}$ count $\text{s}^{-1} \text{cm}^{-2}$ in the BAT (close to the BAT 30-day detection limit). This X-ray activity is compatible with that recorded in recent years, but quite different from the activity of the source during 1997, when the ASM flux was above 1 count s^{-1} per Scanning Shadow Camera during most of the year, with a peak well above 2 count s^{-1} around 1997 June.

As noted previously (Section 1), Mrk 501 is not in the third EGRET catalog, although there was a marginally significant EGRET detection during the γ -ray outburst (with no clear X-ray counterpart) in 1996 (Kataoka et al. 1999). At that time, the source was detected at a level of 4.0σ at energies above 0.1 GeV and at 5.2 σ above 0.5 GeV. The flux from the EGRET 1996 flare above 0.5 GeV was $(6 \pm 2) \times 10^{-8}$ ph $\text{cm}^{-2} \text{s}^{-1}$, which is about five times higher than the average flux observed by *Fermi* from 2008 August 5 (MJD 54683) to 2009 November 27 (MJD 55162), namely $(1.39 \pm 0.07) \times 10^{-8}$ ph $\text{cm}^{-2} \text{s}^{-1}$ (also above photon energy 0.5 GeV). The *Fermi*-LAT flux measured during the 120 days with the “enhanced” γ -ray activity (MJD 54862–54982) is $(2.03 \pm 0.18) \times 10^{-8}$ ph $\text{cm}^{-2} \text{s}^{-1}$ (above photon energy 0.5 GeV), about a factor of three lower than that detected by EGRET in 1996.

In spite of the relatively low activity, the ASM and BAT fluxes show some flux variations and a positive correlation between the fluxes measured by these two instruments. The discrete correlation function for the ASM/BAT data points shown in Figure 2 is $\text{DCF} = 0.73 \pm 0.17$ for a time lag of zero. On the other hand, the X-ray ASM/BAT fluxes are not significantly correlated with the γ -ray LAT fluxes. We found, for a time lag of zero, $\text{DCF} = 0.32 \pm 0.22$ for the ASM/LAT (<2 GeV) and $\text{DCF} = 0.43 \pm 0.30$ for the ASM/LAT (>2 GeV) flux data points shown in Figure 2. It is also interesting to note that the largest flux variations occur at the highest *Fermi* energies (>2 GeV), where the γ -ray flux increased by one order of magnitude during the 120-day interval MJD 54862–54892. This trend is consistent with the photon index hardening revealed by the spectral analysis reported above (see Figure 1).

We followed the description given in Vaughan et al. (2003) to quantify the flux variability by means of the fractional variability parameter, F_{var} , as a function of photon energy. In order to account for the individual flux measurement errors ($\sigma_{\text{err},i}$), we used the “excess variance” as an estimator of the intrinsic source variance (Nandra et al. 1997; Edelson et al. 2002). This is the variance after subtracting the expected contribution from the measurement errors. For a given energy range, F_{var} is calculated as

$$F_{\text{var}} = \sqrt{\frac{S^2 - \langle \sigma_{\text{err}}^2 \rangle}{\langle F_{\gamma} \rangle^2}}, \quad (1)$$

where $\langle F_{\gamma} \rangle$ is the mean photon flux, S is the standard deviation of the N flux points, and $\langle \sigma_{\text{err}}^2 \rangle$ is the average mean square error, all determined for a given energy bin.

Figure 3 shows the F_{var} values derived for the four different energy ranges and the time window covered by the light curves shown in Figure 2. The source is variable at all energies. The uncertainty in the variability quantification for the *Swift*-BAT energies is large due to the fact that Mrk 501 is a relatively weak X-ray source, and is therefore difficult to detect above 15 keV in exposure times as short as 30 days. In contrast, the variability at the *RXTE*-ASM and, especially, *Fermi*-LAT energies, is significant ($>3\sigma$ level). The amplitude variability in the two X-ray bands is compatible within errors, and the same holds for

¹⁴³ The fluxes depicted in the *Fermi*-LAT light curves were computed fixing the photon index to 1.78 (average index during the first 480 days of *Fermi* operation) and fitting only the normalization factor of the power-law function.

¹⁴⁴ http://xte.mit.edu/ASM_lc.html

¹⁴⁵ A dwell is a scan/rotation of the ASM Scanning Shadow Camera lasting 90 s.

¹⁴⁶ <http://swift.gsfc.nasa.gov/docs/swift/results/transients/>

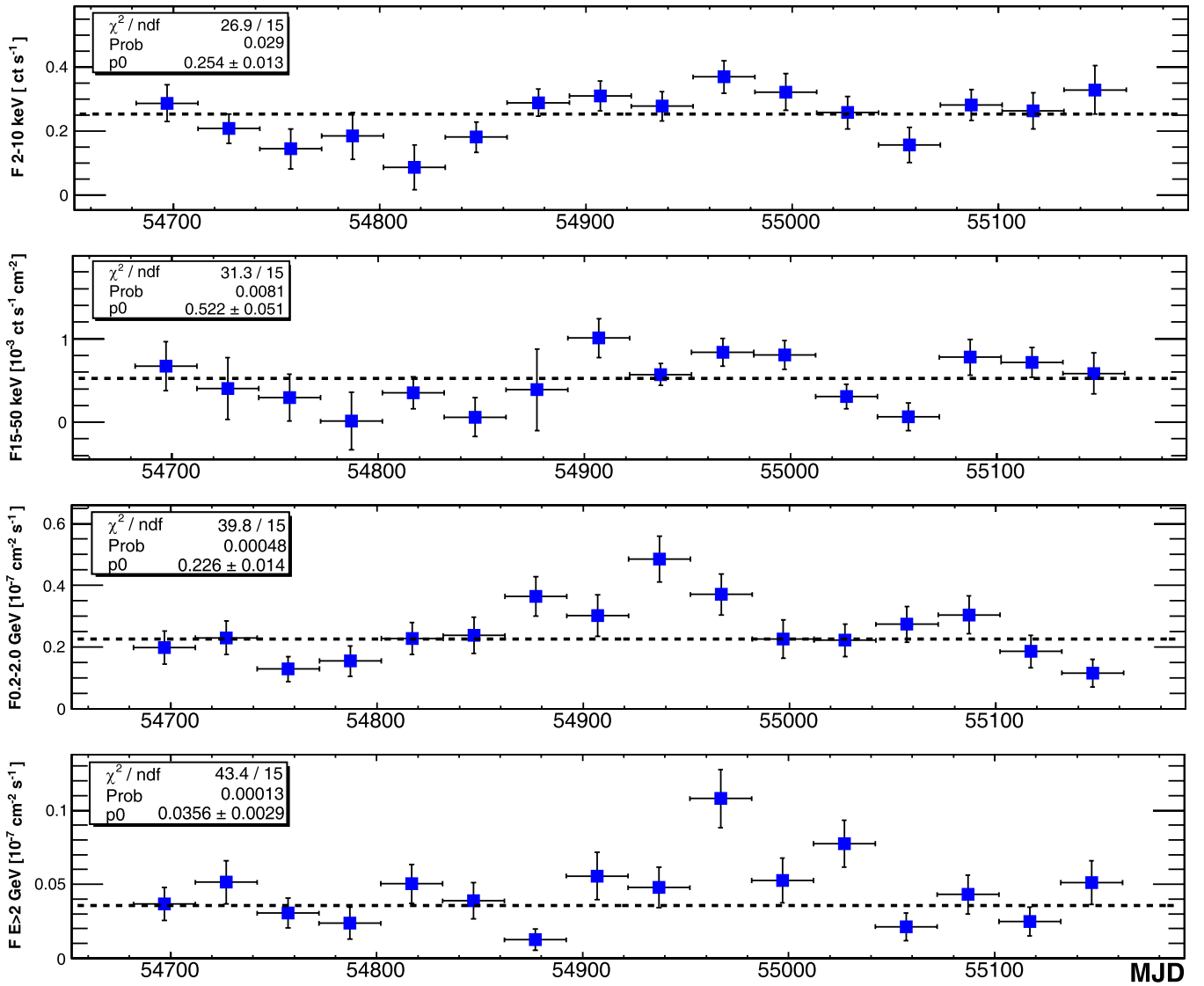


Figure 2. Multifrequency light curves of Mrk 501 with 30-day time bins obtained with three all-sky-monitoring instruments: *RXTE*-ASM (2–10 keV, first from the top); *Swift*-BAT (15–50 keV, second), and *Fermi*-LAT for two different energy ranges (0.2–2 GeV, third, and >2 GeV, fourth). The light curves cover the period from 2008 August 5 (MJD 54683) to 2009 November 27 (MJD 55162). Vertical bars denote 1σ uncertainties and horizontal bars show the width of the time interval. The horizontal dashed lines and the legends (for all the plots) show the results from a constant fit to the entire 480-day data set.

(A color version of this figure is available in the online journal.)

the variability in the two γ -ray bands. As shown in Figure 3, for the hypothesis of a constant F_{var} over the four energy bands, one obtains $\chi^2 = 3.5$ for three degrees of freedom (probability of 0.32), implying that the energy-dependent variability is not statistically significant. It is worth noticing that the limited sensitivity of the ASM and (particularly) BAT instruments to detect Mrk 501 in 30-day time intervals, as well as the relatively stable X-ray emission of Mrk 501 during the analyzed observations, precludes any detailed X-ray/ γ -ray variability and correlation analysis.

4. SPECTRAL ANALYSIS UP TO 400 GeV

The large effective area of the *Fermi*-LAT instrument permits photon energy reconstruction over many orders of magnitude. As a result, the spectrum of Mrk 501 could be resolved within the energy range 0.1–400 GeV, as shown in Figure 4. This is the first time the spectrum of Mrk 501 has been studied with high accuracy over this energy range. The fluxes were computed

using the analysis procedures described in Section 2. The black line in Figure 4 is the result of an unbinned likelihood fit with a single power-law function in the energy range 0.3–400 GeV,¹⁴⁷ and the red contour is the 68% uncertainty of the fit. The data are consistent with a pure power-law function with a photon index of 1.78 ± 0.03 . The black data points result from the analysis in differential energy ranges¹⁴⁸ ($\log \Delta E = 0.4$). The points are well within 1σ – 2σ from the fit to the overall spectrum (black line), which confirms that the entire *Fermi* spectrum is consistent with a pure power-law function. Note, however, that, due to the low photon count, the error bars for the highest energy data points are

¹⁴⁷ The unbinned likelihood fit was performed on photon energies above 0.3 GeV in order to reduce systematics. See Section 2 for further details.

¹⁴⁸ Because the analysis was carried out in small energy ranges, we decided to fix the spectral index at 1.78 (the value obtained from fitting the entire energy range) and fit only the normalization factor. We repeated the same procedure fixing the photon indices to 1.5 and 2.0 and found no significant change. Therefore, the results from the differential energy analysis are not sensitive to the photon index used in the analysis.

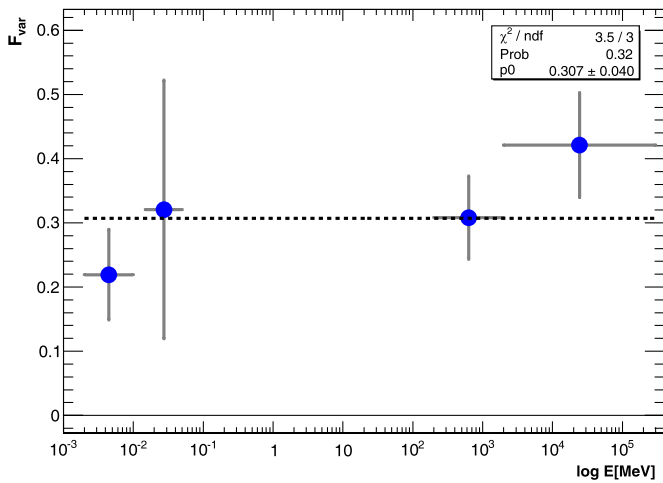


Figure 3. Fractional variability parameter for 16 months data (2008 August 5–2009 November 27) from three all-sky-monitoring instruments: *RXTE*-ASM (2–10 keV), *Swift*-BAT (15–50 keV), and *Fermi*-LAT (two energy ranges 0.2–2 GeV and 2–300 GeV). The fractional variability was computed according to Vaughan et al. (2003) using the light curves from Figure 2. Vertical bars denote 1σ uncertainties and horizontal bars indicate the width of each energy bin. The horizontal dashed line and the legend show the results from a constant fit.

(A color version of this figure is available in the online journal.)

rather large. The predicted (by the model for Mrk 501) numbers of photons detected by the LAT in the energy bins 60–160 GeV and 160–400 GeV are only 11 and 3, respectively. Therefore, even though the signal significance in the highest energy bins is very high due to the very low background (the TS values for the two highest-energy ranges is 162 and 61, respectively), the large statistical uncertainties could hide a potential turnover in the spectrum of Mrk 501 around 100 GeV photon energies. As we know from past observations, the VHE spectrum is significantly softer than the one observed by *Fermi* (e.g., Aharonian et al. 2001; Anderhub et al. 2009), and hence the spectrum of Mrk 501 must have a break around the highest *Fermi*-LAT energies.

In Section 3, we reported remarkable spectral variability during the 120-day time interval MJD 54862–54982, when Mrk 501 was characterized by a photon flux (at >0.3 GeV) twice as large as during the rest of the exposure. In order to understand better the behavior of the source during that time, we produced SED plots (analogous to that of Figure 4) for each of the 30-day time intervals from the period with the enhanced flux level. These are shown in Figure 5, together with the SED plots from the 30-day time intervals before and after this 120-day epoch, which are representative of the average source behavior during the other 360 days. The variability of the SED data points below a few GeV is rather mild (factor of two), but above a few GeV the spectra vary substantially (factor of 10). The γ -ray signal at the highest energies is suppressed during MJD 54862–54982, while it increases by a large factor during MJD 54952–54982, where the analysis model for Mrk 501 predicts 2.0 photons in the energy range 160–400 GeV. It is worth stressing that for the SED from Figure 4, which corresponds to the total exposure of 480 days, the analysis model for Mrk 501 predicts only 3.2 photons in the highest energy bin. Hence, the time interval MJD 54952–54982 holds almost all the signal detected by the LAT in the energy range 160–400 GeV during 16 months. The situation changes somewhat for the lower energy bin 60–160 GeV, for which the analysis model for Mrk 501 predicts 2.4 photons for the time interval MJD 54952–54982, while it does predict 11.3 photons for the entire 16-month time

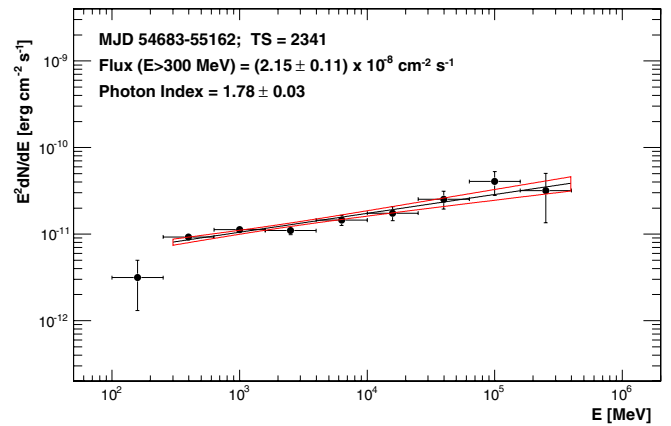


Figure 4. SED for Mrk 501 from *Fermi*-LAT during the period from 2008 August 5 (MJD 54683) to 2009 November 27 (MJD 55162). The black line depicts the result of the unbinned likelihood power-law fit, the red contour is the 68% uncertainty of the fit, and the black data points show the energy fluxes in differential energy ranges. The legend reports the results from the unbinned likelihood power-law fit in the energy range 0.3–400 GeV.

(A color version of this figure is available in the online journal.)

interval. Fortunately, the 30-day time interval characterized by hard spectrum is covered by the 4.5-month campaign that we organized, and hence simultaneous multifrequency observations (radio to TeV) are available for this particular period, as discussed further below.

5. BROADBAND SPECTRAL ENERGY DISTRIBUTION OF Mrk 501

As mentioned in Section 1, we organized a multifrequency campaign (from radio to TeV photon energies) to monitor Mrk 501 during a time period of 4.5 months. The observing campaign started on 2009 March 15 (MJD 54905) and finished on 2009 August 1 (MJD 55044). The observing goal for this campaign was to sample the broadband emission of Mrk 501 every five days, which was largely accomplished whenever the weather and/or technical limitations allowed. The underlying scientific goal has already been outlined in Section 1. A detailed analysis of the multifrequency variability and correlations, as well as the evolution of the overall SED with time, will be reported in a forthcoming paper. In this section of the manuscript, we describe the source coverage during the campaign and the data analysis for several of the participating instruments, and we report on the averaged SED resulting from the campaign. The modeling of these data and the physical implications are given in Sections 6 and 7, respectively.

5.1. Details of the Campaign: Participating Instruments and Temporal Coverage

The list of all the instruments that participated in the campaign is given in Table 1, and the scheduled observations can be found online.¹⁴⁹ In some cases, the planned observations could not be performed due to bad observing conditions, while in some other occasions the observations were performed but the data could not be properly analyzed due to technical problems or rapidly changing weather conditions. In order to quantify the actual time and energy coverage during the campaign on Mrk 501, Figure 6 shows the exposure time as a function of the

¹⁴⁹ [https://confluence.slac.stanford.edu/display/GLAMCOG/Campaign+on+Mrk501+\(March+2009+to+July+2009\)](https://confluence.slac.stanford.edu/display/GLAMCOG/Campaign+on+Mrk501+(March+2009+to+July+2009))

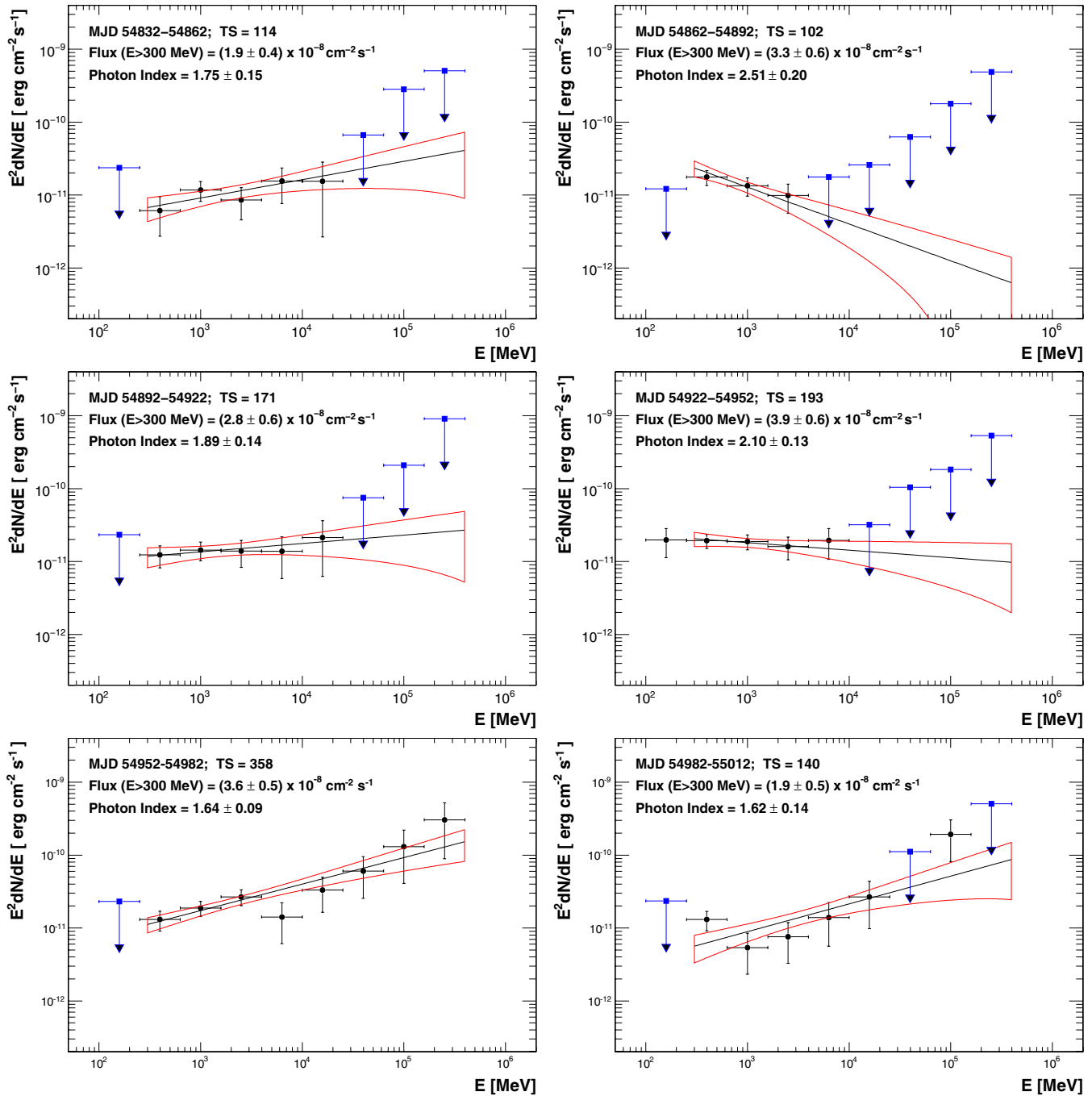


Figure 5. SED for Mrk 501 from *Fermi*-LAT for six 30-day time intervals: MJD 54832–54862 (top left), MJD 54862–54892 (top right), MJD 54892–54922 (middle left), MJD 54922–54952 (middle right), MJD 54952–54982 (bottom left), and MJD 54982–55012 (bottom right). In all the panels, the black line depicts the result of the unbinned likelihood power-law fit, the red contour denotes the 68% uncertainty of the fit, and the black data points show the energy fluxes computed for differential energy ranges. The blue arrows denote 95% upper limits, which were computed for the differential energy ranges with a signal of $TS < 4$ or less than two photons predicted by the analysis model for Mrk 501. The legend reports the results from the unbinned likelihood power-law fit in the energy range 0.3–400 GeV.

(A color version of this figure is available in the online journal.)

energy range for the instruments/observations used to produce the SED shown in Figure 8. Apart from the unprecedented energy coverage (including, for the first time, the GeV energy range from *Fermi*-LAT), the source was sampled quite uniformly with the various instruments participating in the campaign and, consequently, it is reasonable to consider the SED constructed below as the actual average (typical) SED of Mrk 501 during the time interval covered by this multifrequency campaign. The largest non-uniformity in the sampling of the source comes

from the Cherenkov Telescopes, which are the instruments most sensitive to weather conditions. Moreover, while there are many radio/optical instruments spread all over the globe, there are only three Cherenkov Telescope observatories in the northern hemisphere we could utilize (MAGIC, VERITAS, Whipple). Hence, the impact of observing conditions was more important to the coverage at the VHE γ -ray energies.

We note that Figure 6 shows the MAGIC, VERITAS, and Whipple coverage at VHE γ -ray energies, but only the MAGIC

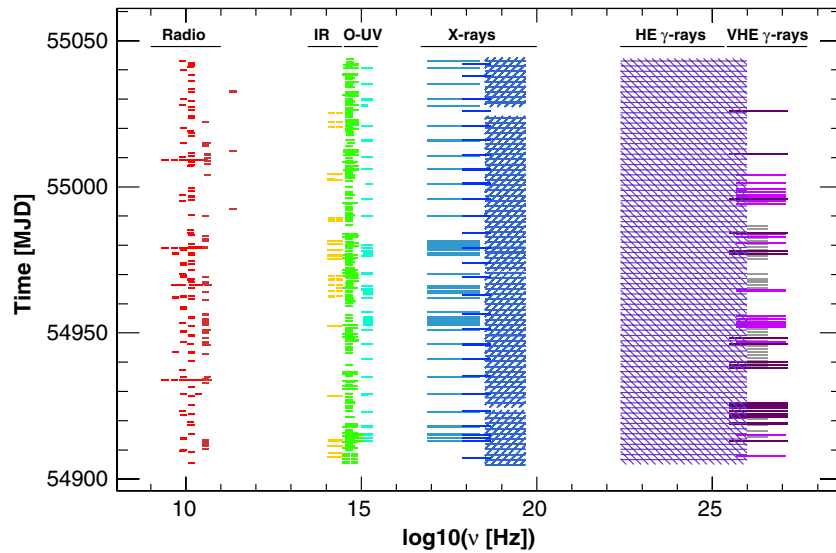


Figure 6. Time and energy coverage during the multifrequency campaign. For the sake of clarity, the minimum observing time displayed in the plot was set to half a day.

Table 1

List of Instruments Participating in the Multifrequency Campaign and Used in the Construction of the SED in Figure 8

| Instrument/Observatory | Energy Range Covered | Web Page |
|--|---|---|
| MAGIC | 0.12–5.8 TeV | http://www.magic.mppmu.mpg.de/ |
| VERITAS | 0.20–5.0 TeV | http://veritas.sao.arizona.edu/ |
| Whipple ^a | 0.4–1.5 TeV | http://veritas.sao.arizona.edu/content/blogsection/6/40/ |
| Fermi-LAT | 0.1–400 GeV | http://www-glast.stanford.edu/index.html |
| Swift-BAT | 14–195 keV | http://heasarc.gsfc.nasa.gov/docs/swift/swiftsc.html |
| RXTE-PCA | 3–28 keV | http://heasarc.gsfc.nasa.gov/docs/xte/rxte.html |
| Swift-XRT | 0.3–9.6 keV | http://heasarc.gsfc.nasa.gov/docs/swift/swiftsc.html |
| Swift-UVOT | V, B, U, UVW1, UVM2, UVW2 | http://heasarc.gsfc.nasa.gov/docs/swift/swiftsc.html |
| Abastumani (through GASP-WEBT program) | R band | http://www.oato.inaf.it/blazars/webt/ |
| Lulin (through GASP-WEBT program) | R band | http://www.oato.inaf.it/blazars/webt/ |
| Roque de los Muchachos (KVA) | R band | http://www.oato.inaf.it/blazars/webt/ |
| (through GASP-WEBT program) | | |
| St. Petersburg (through GASP-WEBT program) | R band | http://www.oato.inaf.it/blazars/webt/ |
| Talmassons (through GASP-WEBT program) | R band | http://www.oato.inaf.it/blazars/webt/ |
| Valle d'Aosta (through GASP-WEBT program) | R band | http://www.oato.inaf.it/blazars/webt/ |
| GRT | V, R, B bands | http://asd.gsfc.nasa.gov/Takanori.Sakamoto/GRT/index.html |
| MitSume | g, Rc, Ic bands | http://www.hp.phys.titech.ac.jp/mitsume/index.html |
| ROVOR | B, R, V, I bands | http://rovor.byu.edu/ |
| Campo Imperatore (through GASP-WEBT program) | H, J, K bands | http://www.oato.inaf.it/blazars/webt/ |
| OAGH | H, J, K bands | http://astro.inaoep.mx/en/observatories/oagh/ |
| WIRO | J, K bands | http://physics.uwoy.edu/~chip/wiro/wiro.html |
| SMA | 225 GHz | http://sma1.sma.hawaii.edu/ |
| VLBA | 4.8, 8.3, 15.4, 23.8, 43.2 GHz | http://www.vlba.nrao.edu/ |
| Noto | 8.4, 43 GHz | http://www.noto.ira.inaf.it/ |
| Metsähovi (through GASP-WEBT program) | 37 GHz | http://www.metsahovi.fi/ |
| VLBA (through MOJAVE program) | 15 GHz | http://www.physics.purdue.edu/MOJAVE/ |
| OVRO | 15 GHz | http://www.astro.caltech.edu/ovroblazars |
| Medicina | 8.4, 22.3 GHz | http://www.med.ira.inaf.it/index_EN.htm |
| UMRAO (through GASP-WEBT program) | 4.8, 8.0, 14.5 GHz | http://www.oato.inaf.it/blazars/webt/ |
| RATAN-600 | 2.3, 4.8, 7.7, 11.1, 22.2 GHz | http://w0.sao.ru/ratan/ |
| Effelsberg (through F-GAMMA program) | 2.6, 4.6, 7.8, 10.3, 13.6, 21.7, 31 GHz | http://www.mpifr-bonn.mpg.de/div/effelsberg/index_e.html |

Notes. The energy range shown in Column 2 is the actual energy range covered during the Mrk 501 observations, and not the instrument's nominal energy range, which might only be achievable for bright sources and excellent observing conditions.

^a The Whipple spectra were not included in Figure 8. The energy range given in the table is based on a very conservative estimate given before performing the spectral analysis of the data. See the text for further comments.

and VERITAS observations were used to produce the spectra shown in Figure 8. The more extensive (120 hr), but less sensitive, Whipple data (shown as gray boxes in Figure 6) were primarily taken to determine the light curve (Pichel et al. 2009)

and a re-optimization was required to derive the spectrum which will be reported elsewhere.

In the following paragraphs, we briefly discuss the procedures used in the data analysis of the instruments participating in the

campaign. The analysis of the *Fermi*-LAT data was described in Section 2, and the results obtained will be described in detail in Section 5.2.

5.1.1. Radio Instruments

Radio data were taken for this campaign from single-dish telescopes, 1 mm interferometer, and one VLBI array, at frequencies between 2.6 GHz and 225 GHz (see Table 1). The single-dish telescopes were the Effelsberg 100 m radio telescope, the 32 m Medicina radio telescope, the 14 m Metsähovi radio telescope, the 32 m Noto radio telescope, the Owens Valley Radio Observatory (OVRO) 40 m telescope, the 26 m University of Michigan Radio Astronomy Observatory (UMRAO) and the 600 m ring radio telescope RATAN-600. The millimeter-interferometer is the Submillimeter Array (SMA). The NRAO VLBA was used for the VLBI observations. For the single-dish instruments and SMA, Mrk 501 is point-like and unresolved at all observing frequencies. Consequently, the single-dish measurements denote the total flux density of the source integrated over the whole source extension. Details of the observing strategy and data reduction can be found in Fuhrmann et al. (2008), Angelakis et al. (2008, F-GAMMA project), Teräsranta et al. (1998, Metsähovi), Aller et al. (1985, UMRAO), Venturi et al. (2001, Medicina and Noto), Kovalev et al. (1999, RATAN-600), and Richards et al. (2011, OVRO).

In the case of the VLBA, the data were obtained at various frequencies from 5 GHz to 43 GHz through various programs (BP143, BK150, and MOJAVE). The data were reduced following standard procedures for data reduction and calibration (see, for example, Lister et al. 2009, for a description of the MOJAVE program which provided the 15 GHz data). Since the VLBA angular resolution is smaller than the radio source extension, measurements were performed for the most compact core region, as well as for the total radio structure at parsec scales. The VLBA core size was determined with two-dimensional circular or elliptical Gaussian fits to the measured visibilities. The FWHM size of the core was estimated to be in the range 0.14–0.18 mas at the highest observing frequencies, 15–43 GHz. Both the total and the core radio flux densities from the VLBA data are depicted in Figure 8.

5.1.2. Optical and Near-IR Instruments

The coverage at optical frequencies was obtained through various telescopes around the globe, and this decreased the sensitivity to weather/technical difficulties and provided good overall coverage of the source, as depicted in Figure 6. Many of the observations were performed within the GASP-WEBT program (e.g., Villata et al. 2008, 2009); that is the case for the data collected by the telescopes at Abastumani, Lulin, Roque de los Muchachos (KVA), St. Petersburg, Talmassons, and Valle d’Aosta observatories (*R* band), and also for Campo Imperatore (near-infrared frequencies, *JHK* bands). In addition, the telescopes GRT, ROVOR, and MitSume provided data with various optical filters, while OAGH and WIRO provided data at near-infrared wavelengths. See Table 1 for further details.

All the instruments used the calibration stars reported in Villata et al. (1998), and the Galactic extinction was corrected with the coefficients given in Schlegel et al. (1998). On the other hand, the flux from the host galaxy, which in the *R* band accounts for about two-thirds of the overall measured optical flux (Nilsson et al. 2007), was not subtracted. As can be seen from Figure 8, the host galaxy contribution shows up as an additional (narrow) bump in the SED with the peak located at infrared frequencies

and the flux decreasing rapidly with increasing frequency. At frequencies above 10^{15} Hz, the blazar emission again dominates the radiative output of Mrk 501.

5.1.3. Swift-UVOT

The *Swift*-Ultra-Violet/Optical Telescope (UVOT; Roming et al. 2005) data used in this analysis include all the observations performed during the time interval MJD 54905 and 55044, which amounts to 41 single pointing observations that were requested to provide UV coverage during the Mrk 501 multifrequency campaign. The UVOT telescope cycled through each of six optical and ultraviolet passbands (*V*, *B*, *U*, *UVM1*, *UVM2*, and *UVW2*). Photometry was computed using a 5 arc-sec source region around Mrk 501 using a custom UVOT pipeline that obtains similar photometric results to the public pipeline (Poole et al. 2008). The custom pipeline also allows for separate, observation-by-observation corrections for astrometric misalignments (Acciari et al. 2011). A visual inspection was also performed on each of the observations to ensure proper data quality selection and correction. The flux measurements obtained have been corrected for Galactic extinction $E_{B-V} = 0.019$ mag (Schlegel et al. 1998) in each spectral band (Fitzpatrick 1999).

5.1.4. Swift-XRT

All the *Swift*-X-ray Telescope (XRT; Burrows et al. 2005) Windowed Timing observations carried out from MJD 54905 to 55044 were used for the analysis: this amounts to a total of 41 observations performed within this dedicated multi-instrument effort to study Mrk 501. The XRT data set was first processed with the XRTDAS software package (v.2.5.0) developed at the ASI Science Data Center (ASDC) and distributed by HEASARC within the HEASoft package (v.6.7). Event files were calibrated and cleaned with standard filtering criteria with the *xrtpipeline* task using the latest calibration files available in the *Swift* CALDB. The individual XRT event files were then merged together using the XSELECT package and the average spectrum was extracted from the summed event file. Events for the spectral analysis were selected within a circle of 20-pixel (~ 47 arcsec) radius centered at the source position and enclosing about 95% of the point-spread function (PSF) of the instrument. The background was extracted from a nearby circular region of 40-pixel radius. The source spectrum was binned to ensure a minimum of 20 counts per bin to utilize the χ^2 minimization fitting technique. The ancillary response files were generated with the *xrtmkarf* task applying corrections for the PSF losses and CCD defects using the cumulative exposure map. The latest response matrices (v.011) available in the *Swift* CALDB were used.

The XRT average spectrum in the 0.3–10 keV energy band was fitted using the XSPEC package. We adopted a log-parabolic model for the photon flux spectral density (Massaro et al. 2004a, 2004b) of the form $\log[\mathcal{F}(E)] = \log K - a \log[E/\text{keV}] - b \log^2[E/\text{keV}]$, with an absorption hydrogen-equivalent column density fixed to the Galactic value in the direction of the source, namely $1.56 \times 10^{20} \text{ cm}^{-2}$ (Kalberla et al. 2005). This model provided a good description of the observed spectrum, with the exception of the 1.4–2.3 keV energy band where spectral fit residuals were present. These residuals are due to known XRT calibration uncertainties (SWIFT-XRT-CALDB-12)¹⁵⁰ and hence we decided to exclude the 1.4–2.3 keV

¹⁵⁰ http://heasarc.gsfc.nasa.gov/docs/heasarc/caldb/swift/docs/xrt/SWIFT-XRT-CALDB-09_v12.pdf

energy band from the analysis. In addition, we had to apply a small energy offset (~ 40 eV) to the observed energy spectrum. The origin of this correction is likely to be CCD charge traps generated by radiation and high-energy proton damage (SWIFT-XRT-CALDB-12), which affects mostly the lowest energies (first one or two bins) in the spectrum. The resulting spectral fit gave the following parameters: $K = (3.41 \pm 0.03) \times 10^{-2} \text{ ph cm}^{-2} \text{ s}^{-1} \text{ keV}^{-1}$, $a = 1.96 \pm 0.04$, and $b = 0.308 \pm 0.010$. The XRT SED data shown in Figure 8 were corrected for the Galactic absorption and then binned into 10 energy intervals.

5.1.5. RXTE-PCA

The *Rossi*-X-ray Timing Explorer (*RXTE*; Bradt et al. 1993) satellite performed 29 pointing observations of Mrk 501 during the time interval MJD 54905 and 55044. These observations amount to a total exposure of 52 ks, which was requested through a dedicated Cycle 13 proposal to provide X-ray coverage for our campaign. We did not find a significant signal in the *RXTE*-HEXTE data and hence we only report on the data from *RXTE*-Proportional Counter Array (*RXTE*-PCA), which is the main pointing instrument on board *RXTE*. The data analysis was performed using FT00LS v6.5 and following the procedures and filtering criteria recommended by the *RXTE* Guest Observer Facility¹⁵¹ after 2007 September. In particular, the observations were filtered following the conservative procedures for faint sources¹⁵²: Earth elevation angle greater than 10° , pointing offset less than $0^\circ 02'$, time since the peak of the last South Atlantic Anomaly (SAA) passage greater than 30 minutes, and electron contamination less than 0.1. For further details on the analysis of faint sources with *RXTE*, see the online Cook Book.¹⁵³ In the data analysis, in order to increase the quality of the signal, only the first xenon layer of PCU2 was used. We used the package *pcabackest* to model the background and the package *saextract* to produce spectra for the source and background files and the script¹⁵⁴ *pcarsp* to produce the response matrix.

The PCA average spectrum in the 3–28 keV energy band was fitted using the XSPEC package with a single power-law function $\log[\mathcal{F}(E)] = \log K - a \log[E/\text{keV}]$ with a constant neutral hydrogen column density N_{H} fixed at the Galactic value in the direction of the source, namely $1.56 \times 10^{20} \text{ cm}^{-2}$ (Kalberla et al. 2005). However, since the PCA bandpass starts at 3 keV, the value used for N_{H} does not significantly affect our results. The resulting spectral fit provided a good representation of the data for the following parameters: $K = (4.34 \pm 0.11) \times 10^{-2} \text{ ph cm}^{-2} \text{ s}^{-1} \text{ keV}^{-1}$, and $a = 2.28 \pm 0.02$. The PCA average spectrum obtained using 23 energy bins is shown in Figure 8.

5.1.6. Swift-BAT

The *Swift*-BAT (Barthelmy et al. 2005) analysis results presented in this paper were derived with all the available data during the time intervals MJD 54905 and 55044. The spectrum was extracted following the recipes presented in Ajello et al. (2008, 2009b). This spectrum is constructed by weighted

averaging of the source spectra extracted from short exposures (e.g., 300 s) and is representative of the averaged source emission over the time range spanned by the observations. These spectra are accurate to the mCrab level and the reader is referred to Ajello et al. (2009a) for more details. The *Swift*-BAT spectrum is consistent with a power-law function with the normalization parameter $K = 0.24 \pm 0.16 \text{ ph cm}^{-2} \text{ s}^{-1} \text{ keV}^{-1}$ and photon index $a = 2.8 \pm 0.4$.

5.1.7. MAGIC

MAGIC is a system of two 17 m diameter IACTs for VHE γ -ray astronomy located on the Canary Island of La Palma, at an altitude of 2200 m above sea level. At the time of the observation, MAGIC-II, the new second telescope of the current array system, was still in its commissioning phase so that Mrk 501 was observed in standalone mode by MAGIC-I, which is in scientific operation since 2004 (Albert et al. 2008). The MAGIC telescope monitored the VHE activity of Mrk 501 in the framework of the organized multifrequency campaign. The observations were performed in the so-called wobble mode (Daum 1997). In order to have a low-energy threshold, only observations at zenith angles less than 35° were used in this analysis. Bad weather and a shutdown for a scheduled hardware system upgrade during the period MJD 54948–54960 (April 27–May 13) significantly reduced the actual amount of observing time compared to what had initially been scheduled for this campaign. The data were analyzed following the prescription given in Albert et al. (2008) and Aliu et al. (2009). The data surviving the quality cuts amount to a total of 16.2 hr. The preliminary reconstructed photon fluxes for the individual observations gave an average activity of about 30% the flux of the Crab Nebula, with small (typically much less than a factor of two) flux variations. The derived spectrum was unfolded to correct for the effects of the limited energy resolution of the detector and of possible bias (Albert et al. 2007b). The resulting spectrum was fitted satisfactorily with a single power-law function of the form $\log[\mathcal{F}(E)] = \log K - a \log[E/\text{TeV}]$, giving the normalization parameter $K = (0.90 \pm 0.05) \times 10^{-11} \text{ ph cm}^{-2} \text{ s}^{-1} \text{ TeV}^{-1}$ and photon index $a = 2.51 \pm 0.05$.

5.1.8. VERITAS

VERITAS is a state-of-the-art TeV γ -ray observatory consisting of four 12 m diameter IACTs. VERITAS is located at the base camp of the F. L. Whipple Observatory in southern Arizona, USA, at an altitude of 1250 m above sea level, and the system has been fully operational since fall 2007 (Acciari et al. 2010). VERITAS observed Mrk 501 as part of the long-term monitoring campaign between 2009 March and June. The observations were performed in “wobble” mode (Daum 1997) at relatively low zenith angle ($< 40^\circ$). These data were analyzed following the prescription reported in Acciari et al. (2008). After removal of data runs with poor observing conditions, a total of 9.7 hr of good quality data were obtained between MJD 54907 and MJD 55004. Due to the long-term nature of these observations, several factors had to be taken into account when analyzing the data. The initial portion of the campaign includes data taken under standard four-telescope operating conditions. Two nights of data were taken with only two operational telescopes due to technical difficulties. For the latter portion of the campaign, data were taken over several nights with three operational telescopes because one of the telescopes was being relocated as part of an upgrade to the array (Perkins et al. 2009). The effective collection areas for the array in these three configurations

¹⁵¹ <http://www.universe.nasa.gov/xrays/programs/rxte/pca/doc/bkg/bkg-2007-saa/>

¹⁵² The average net count rate from Mrk 501 was about 7 counts $\text{s}^{-1} \text{ pcu}^{-1}$ in the energy range 3–20 keV with flux variations typically much smaller than a factor of two.

¹⁵³ http://heasarc.gsfc.nasa.gov/docs/xte/recipes/cook_book.html

¹⁵⁴ The CALDB files are located at <http://heasarc.gsfc.nasa.gov/FTP/caldb>.

were calculated using Monte Carlo simulations of extensive air showers passed through the analysis chain with detector configurations corresponding to the respective data-taking conditions.

An initial analysis of the VHE activity showed an increase in the flux by a factor of about five during MJD 54953–54956. Because of the large difference in the VHE flux, we decided to analyze this three-day data set (corresponding to a “flaring” state of Mrk 501) separately from the rest of the collected data (non-flaring). The “flaring” epoch consists of 2.4 hr of data taken during MJD 54953–54956. The “non-flaring” epoch consists of 7.3 hr of data taken during the remaining portion of the campaign. The spectra from these two data sets were each fitted with a single power-law function of the form $\log[\mathcal{F}(E)] = \log K - a \log[E/\text{TeV}]$. The resulting fit parameter values are $K = (4.17 \pm 0.24) \times 10^{-11} \text{ ph cm}^{-2} \text{ s}^{-1} \text{ TeV}^{-1}$ with $a = 2.26 \pm 0.06$ for the “flaring” state, and $K = (0.88 \pm 0.06) \times 10^{-11} \text{ ph cm}^{-2} \text{ s}^{-1} \text{ TeV}^{-1}$ with the photon index $a = 2.48 \pm 0.07$ for the “non-flaring” state.

5.2. Fermi-LAT Spectra During the Campaign

The Mrk 501 spectrum measured by *Fermi*-LAT, integrated during the time interval of the multifrequency campaign, is shown in the panel (b) of Figure 7. The spectrum can be described by a power-law function with the photon index 1.74 ± 0.05 . The flux data points resulting from the analysis in differential energy ranges are within 1σ – 2σ of the power-law fit result; this is an independent indication that a single power-law function is a good representation of the spectrum during the multifrequency campaign. On the other hand, the shape of the spectrum depicted by the differential energy flux data points suggests the possibility of a concave spectrum. As was discussed in Sections 3 and 4 (see Figures 1 and 5), Mrk 501 showed substantial spectral variability during the time period covered by the multifrequency campaign, with some 30-day time intervals characterized by relatively soft spectra (photon index ~ 2 for the 30-day intervals MJD 54892–54922 and MJD 54922–54952) and others by relatively hard spectra (photon index ~ 1.6 for the 30-day intervals MJD 54952–54982, MJD 54982–55012, and MJD 55012–55042). Panel (b) of Figure 7 presents the average spectrum over those time intervals, and hence it would not be surprising to see two slopes (instead of one) in the spectrum. In order to evaluate this possibility, a broken power-law fit was applied, yielding indices of 1.86 ± 0.08 and 1.44 ± 0.14 below and above a break energy of $10 \pm 3 \text{ GeV}$, respectively. The likelihood ratio of the broken power law and the power law is 2.2. Given that the broken power law has two additional degrees of freedom, this indicates that the broken power law is not statistically preferred over the single power-law function.

For comparison purposes we also computed the spectra for time intervals before and after the multifrequency campaign (MJD 54683–54901 and MJD 55044–55162).¹⁵⁵ These two spectra, shown in panels (a) and (c) of Figure 7, can both be described satisfactorily by single power-law functions with photon indices 1.82 ± 0.06 and 1.80 ± 0.08 . Note that the two spectra are perfectly compatible with each other, which is consistent with the relatively small flux/spectral variability shown in Figures 1 and 2 for those time periods.

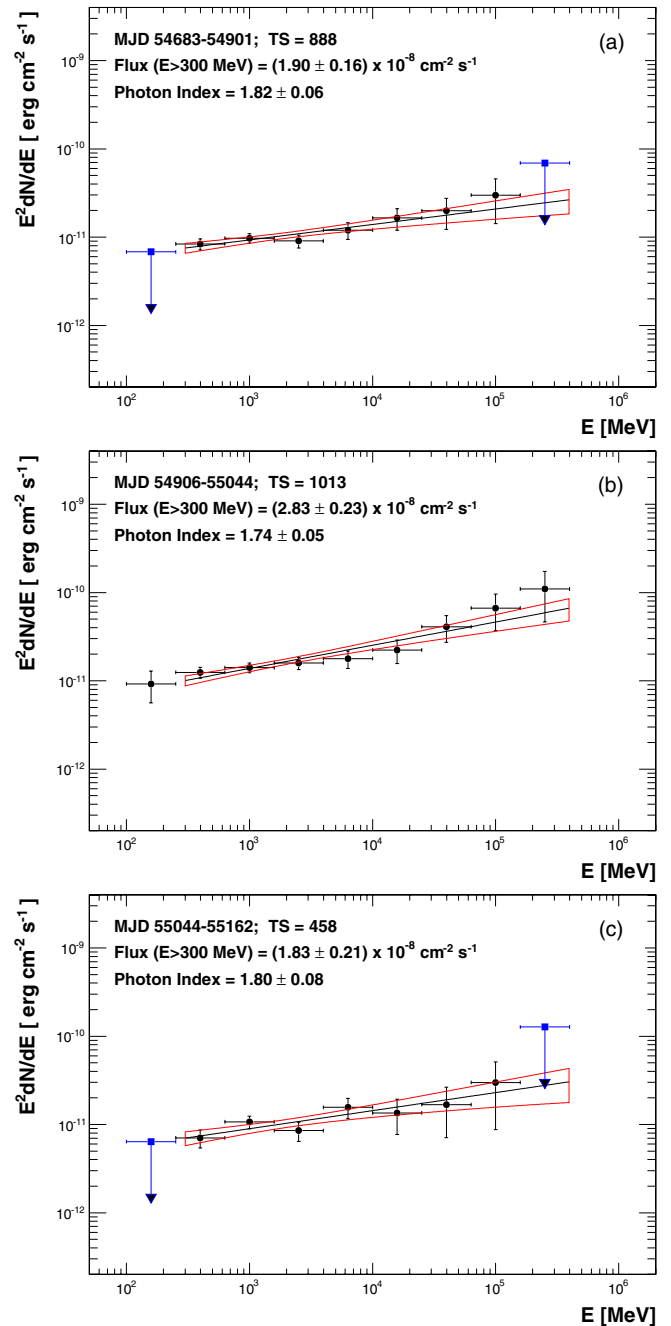


Figure 7. SED for Mrk 501 from *Fermi*-LAT for several time intervals of interest. Panel (a) shows the SED for the time period before the multifrequency campaign (MJD 54683–54901), panel (b) for the time interval corresponding to the multifrequency campaign (MJD 54905–55044), and panel (c) for the period after the campaign (MJD 55044–55162). In all panels, the black line depicts the result of the unbinned likelihood power-law fit, the red contours denote the 68% uncertainty of the power-law fit, and blue arrows denote upper limits at 95% confidence level, which were computed for the differential energy ranges with a signal of $TS < 4$ or less than two photons predicted by the analysis model for Mrk 501. The legend reports the results from the unbinned likelihood power-law fit in the energy range 0.3–400 GeV.

(A color version of this figure is available in the online journal.)

5.3. The Average Broadband SED During the Campaign

The average broadband SED of Mrk 501 resulting from our 4.5 month long multifrequency campaign is shown in Figure 8. The TeV data from MAGIC and VERITAS have been corrected for the absorption in the EBL using the particular EBL model

¹⁵⁵ Technical problems prevented the scientific operation of the *Fermi*-LAT instrument during the interval MJD 54901–54905.

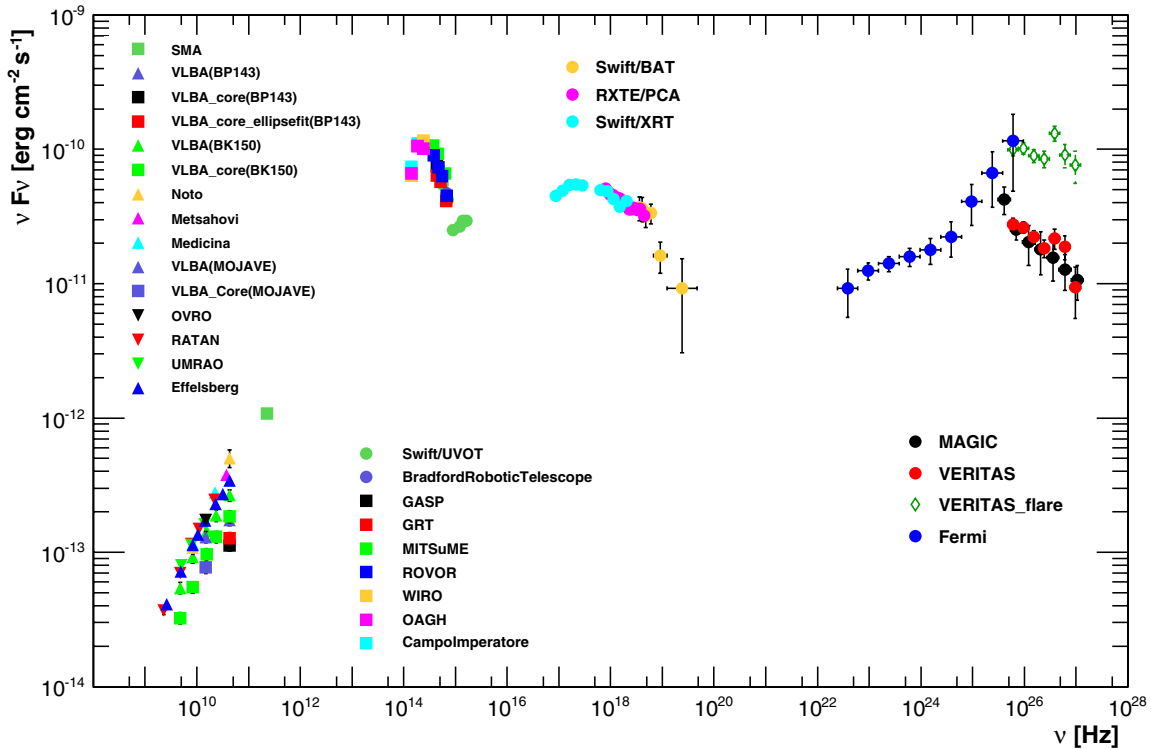


Figure 8. SED for Mrk 501 averaged over all observations taken during the multifrequency campaign performed between 2009 March 15 (MJD 54905) and 2009 August 1 (MJD 55044). The legend reports the correspondence between the instruments and the measured fluxes. Further details about the instruments are given in Section 5.1. The optical and X-ray data have been corrected for Galactic extinction, but the host galaxy (which is clearly visible at the IR/optical frequencies) has not been subtracted. The TeV data from MAGIC and VERITAS have been corrected for the absorption in the EBL using the model reported in Franceschini et al. (2008). The VERITAS data from the time interval MJD 54952.9–54955.9 were removed from the data set used to compute the average spectrum, and are depicted separately in the SED plot (in green diamonds). See the text for further details.

by Franceschini et al. (2008). The corrections given by the other low-EBL-level models (Kneiske et al. 2004; Gilmore et al. 2009; Finke et al. 2010) are very similar for the low redshift of Mrk 501 ($z = 0.034$). The attenuation factor at a photon energy of 6 TeV (the highest energy detected from Mrk 501 during this campaign) is in the range $e^{-\tau_{\gamma\gamma}} \simeq 0.4$ –0.5, and smaller at lower energies.

During the campaign, as already noted above, the source did not show large flux variations like those recorded by EGRET in 1996, or those measured by X-ray and TeV instruments in 1997. Nevertheless, significant flux and spectral variations at γ -ray energies occurred in the time interval MJD 54905–55044. The largest flux variation during the campaign was observed at TeV energies during the time interval MJD 54952.9–54955.9, when VERITAS measured a flux about five times higher than the average one during the campaign. Because of the remarkable difference with respect to the rest of the analyzed exposure, these observations were excluded from the data set used to compute the average VERITAS spectrum for the campaign; the three-day “flaring-state” spectrum (2.4 hr of observation) is presented separately in Figure 8. Such a remarkable flux enhancement was not observed in the other energy ranges and hence Figure 8 shows only the averaged spectra for the other instruments.¹⁵⁶

The top panel in Figure 9 shows a zoom of the high-energy bump depicted in Figure 8. The last two energy bins from *Fermi* (60–160 and 160–400 GeV) are systematically above (1σ – 2σ) the measured/extrapolated spectrum from MAGIC and VERITAS. Even though this mismatch is not statistically

significant, we believe that the spectral variability observed during the 4.5 month long campaign (see Sections 4 and 5.2) could be the origin of such a difference. Because *Fermi*-LAT operates in a survey mode, Mrk 501 is constantly monitored at GeV energies,¹⁵⁷ while this is not the case for the other instruments which typically sampled the source for ≤ 1 hr every five days approximately. Moreover, because of bad weather or moonlight conditions, the monitoring at the TeV energies with Cherenkov telescopes was even less regular than that at lower frequencies. Therefore, *Fermi*-LAT may have measured high activity that was missed by the other instruments. Indeed, the 2.4 hr high-flux spectrum from VERITAS depicted in Figure 8 (which was obtained during the three-day interval MJD 54952.9–54955.9) demonstrates that, during the multifrequency campaign, there were time periods with substantially (factor of five) higher TeV activity. It is possible that the highest energy LAT observations (≥ 50 GeV) include high TeV flux states which occurred while the IACTs were not observing.

If the flaring activity occurred only at the highest photon energies, then the computed *Fermi*-LAT flux (>0.3 GeV) would not change very much and the effect might only be visible in the measured power-law photon index. This seems to be the case in the presented data set. As was shown in Figure 5, the 30-day intervals MJD 54922–54952 and MJD 54952–54982 have photon fluxes above 0.3 GeV of $(3.9 \pm 0.6) \times 10^{-8}$ ph cm $^{-2}$ s $^{-1}$ and $(3.6 \pm 0.5) \times 10^{-8}$ ph cm $^{-2}$ s $^{-1}$, while their photon indices are 2.10 ± 0.13 and 1.63 ± 0.09 , respectively. Therefore, the spectral information (together with the enhanced photon flux) indicates

¹⁵⁶ The MAGIC telescope did not operate during the time interval MJD 54948–54965 due to a drive system upgrade.

¹⁵⁷ For every three hours of *Fermi* operation, Mrk 501 is in the LAT FoV for about 0.5 hr.

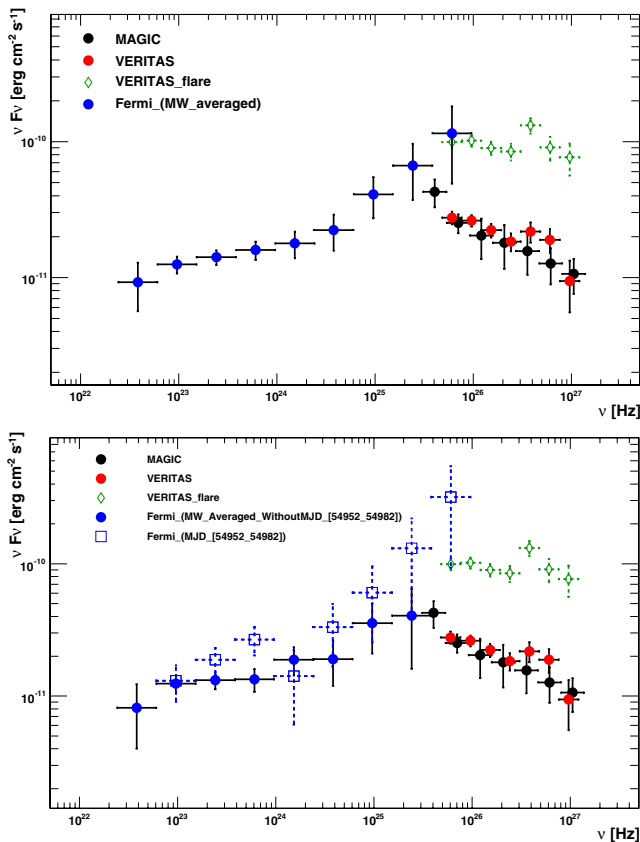


Figure 9. Top panel: enlargement of the γ -ray energy range from Figure 8. Bottom panel: same SED as in the top panel, but with the *Fermi*-LAT data from the multifrequency campaign split in two data sets: MJD 54952–54982 (open blue squares) and the rest (filled blue circles).

the presence of flaring activity at the highest γ -ray energies during the second 30-day time period. Besides the factor ~ 5 VHE flux enhancement recorded by VERITAS and Whipple at the beginning of the time interval MJD 54952–54982, MAGIC, and Whipple also recorded a factor ~ 2 VHE flux enhancement at the end of this 30-day time interval (see preliminary fluxes reported in Paneque 2009; Pichel et al. 2009). This flux enhancement was measured for the time interval MJD 54975–54977, but there were no VHE measurements during the period MJD 54970.5–54975.0. Thus, the average *Fermi*-LAT spectrum could have been affected by elevated VHE activity during the 30-day time interval MJD 54952–54982, which was only partly covered by the IACTs participating in the campaign.

For illustrative purposes, in the bottom panel of Figure 9, we show separately the *Fermi*-LAT spectra for the 30-day time interval MJD 54952–54982 (high photon flux and hard spectrum), and for the rest of the campaign. It is interesting to note that the *Fermi*-LAT spectrum without the 30-day time interval MJD 54952–54982 (blue data points in the bottom panel of Figure 9) agrees perfectly with the VHE spectrum measured by IACTs. We also want to point out that the power-law fit to the *Fermi*-LAT spectrum without the 30-day interval MJD 54952–54982 gave a photon flux above 0.3 GeV of $(2.62 \pm 0.25) \times 10^{-8}$ ph cm $^{-2}$ s $^{-1}$ with a photon index of 1.78 ± 0.07 , which is statistically compatible with the results for the power-law fit to the *Fermi*-LAT data from the entire campaign (see panel (b) in Figure 7). As discussed above, the flaring activity occurred mostly at the highest energies, where the (relatively) low photon count has little impact on the overall power-law fit performed above 0.3 GeV.

This is the most complete quasi-simultaneous SED ever collected for Mrk 501, or for any other TeV-emitting BL Lac object (see also A. A. Abdo et al. 2011, in preparation). At the highest energies, the combination of *Fermi* and MAGIC/VERITAS data allows us to measure, for the first time, the high-energy bump without any spectral gap. The low-energy spectral component is also very well characterized with *Swift*-UVOT, *Swift*-XRT, and *RXTE*-PCA data, covering the peak of the synchrotron continuum. The only (large) region of the SED with no available data corresponds to the photon energy range 200 keV–100 MeV, where the sensitivity of current instruments is not good enough to detect Mrk 501. It is worth stressing that the excellent agreement in the overlapping energies among the various instruments (which had somewhat different time coverages) indicates that the collected data are representative of the truly average SED during the multi-instrument campaign.

6. MODELING THE SPECTRAL ENERGY DISTRIBUTION OF Mrk 501

The simultaneous broadband data set resulting from the multifrequency campaign reported above offers an unprecedented opportunity of modeling the emission of an archetypal TeV blazar in a more robust way than in the past. It is widely believed that the radio-to- γ -ray emission of the BL Lac class of an AGN is produced predominantly via the synchrotron and SSC processes, and hence the homogeneous one-zone approximation of the SSC scenario is the simplest model to consider. Here, we therefore adopt the “standard” one-zone SSC model, which has had moderate success in accounting for the spectral and temporal properties of the TeV-emitting BL Lac objects analyzed so far (e.g., Finke et al. 2008; Ghisellini et al. 2009b, and references therein). We also note that one-zone SSC analyses have been widely applied before to the particular case of Mrk 501 (e.g., Bednarek & Protheroe 1999; Katarzyński et al. 2001; Tavecchio et al. 2001; Kino et al. 2002; Albert et al. 2007a). However, it is important to stress that the modeling results from the previous works related almost exclusively to the high-activity state of Mrk 501. In the more recent work by Anderhub et al. (2009), the source was studied also during its low-activity state, yet the simultaneous observations used in the modeling covered only the X-ray and TeV photon energies. In this paper, we study Mrk 501 during a relatively low-activity state, and the modeling is applied to a more complete broadband SED extending from radio to TeV energies, including the previously unavailable GeV data from *Fermi*. This constitutes a substantial difference with respect to previous works. The resulting constraints on the physical parameters of the source, together with several limitations of the applied scenario, are discussed further down in the following sections.

We want to note that modeling of the average blazar SED based on a scenario assuming a steady-state homogeneous emission zone could be an oversimplification of the problem. The blazar emission may be produced in an inhomogeneous region, involving stratification of the emitting plasma both along and across a relativistic outflow. In such a case, the observed radiative output of a blazar could be due to a complex superposition of different emission zones characterized by very different parameters and emission properties. Some first attempts to approach this problem in a more quantitative way have been already discussed in the literature (e.g., Ghisellini et al. 2005; Katarzyński et al. 2008; Graff et al. 2008; Giannios et al. 2009). The main drawback of the proposed models, however, is the increased number of free parameters (over

the simplest homogeneous one-zone scenario), which reduces considerably the predictive power of the modeling. That is particularly problematic if a “limited” (in time and energy coverage) data set is considered in the modeling. Only a truly simultaneous multifrequency data set covering a large fraction of the available electromagnetic spectrum and a wide range of timescales—like the one collected during this and future campaigns which will be further exploited in forthcoming publications—will enable us to test such more sophisticated and possibly more realistic blazar emission models in a time-dependent manner.

6.1. SSC Modeling

Let us assume that the emitting region is a homogeneous and roughly spherically symmetric moving blob, with radius R and comoving volume $V' \simeq (4\pi/3) R^3$. For this, we evaluate the comoving synchrotron and SSC emissivities, $\nu' j'_{\nu'}$, assuming isotropic distributions of ultrarelativistic electrons and synchrotron photons in the rest frame of the emitting region. Thus, we use the exact synchrotron and inverse-Compton kernels (the latter one valid in both Thomson and Klein–Nishina regimes), as given in Crusius & Schlickeiser (1986) and Blumenthal & Gould (1970), respectively. The intrinsic monochromatic synchrotron and SSC luminosities are then $\nu' L'_{\nu'} = 4\pi V' \nu' j'_{\nu'}$, while the observed monochromatic flux densities (measured in $\text{erg cm}^{-2} \text{s}^{-1}$) can be found as

$$\nu F_{\nu} = \frac{\delta^4}{4\pi d_L^2} [\nu' L'_{\nu'}]_{\nu'=\nu(1+z)/\delta} \simeq \frac{4\pi \delta^4 R^3}{3 d_L^2} [\nu' j'_{\nu'}]_{\nu'=\nu(1+z)/\delta}, \quad (2)$$

where δ is the jet Doppler factor, $z = 0.034$ is the source redshift, and $d_L = 142$ Mpc is the luminosity distance to Mrk 501. In order to evaluate the comoving emissivities $\nu' j'_{\nu'}$, the electron energy distribution $n'_e(\gamma)$ has to be specified. For this, we assume a general power-law form between the minimum and maximum electron energies, γ_{\min} and γ_{\max} , allowing for multiple spectral breaks in between, as well as for an exponential cut-off above γ_{\max} . In fact, the broadband data set for Mrk 501 requires two different electron break energies, and hence we take the electron energy distribution in the form

$$n'_e(\gamma) \propto \begin{cases} \gamma^{-s_1} & \text{for } \gamma_{\min} \leq \gamma < \gamma_{\text{br}, 1} \\ \gamma^{-s_2} & \text{for } \gamma_{\text{br}, 1} \leq \gamma < \gamma_{\text{br}, 2} \\ \gamma^{-s_3} \exp[-\gamma/\gamma_{\max}] & \text{for } \gamma_{\text{br}, 2} \leq \gamma, \end{cases} \quad (3)$$

with the normalization expressed in terms of the equipartition parameter (the ratio of the comoving electron and magnetic field energy densities), namely

$$\eta_e \equiv \frac{U'_e}{U'_B} = \frac{\int \gamma m_e c^2 n'_e(\gamma) d\gamma}{B^2/8\pi}. \quad (4)$$

The measured SED is hardly compatible with a simpler form of the electron distribution with only one break and an exponential cutoff. However, some smoothly curved spectral shape might perhaps be an alternative representation of the electron spectrum (e.g., Stawarz & Petrosian 2008; Tramacere et al. 2009).

The model adopted is thus characterized by four main free parameters (B , R , δ , and η_e), plus seven additional ones related to the electron energy distribution (γ_{\min} , $\gamma_{\text{br}, 1}$, $\gamma_{\text{br}, 2}$, γ_{\max} , s_1 , s_2 , and s_3). These seven additional parameters are determined by the spectral shape of the non-thermal emission continuum probed by the observations, predominantly by the spectral shape of the

synchrotron bump (rather than the inverse-Compton bump), and depend only slightly on the particular choice of the magnetic field B and the Doppler factor δ within the allowed range.¹⁵⁸ There is a substantial degeneracy regarding the four main free parameters: the average emission spectrum of Mrk 501 may be fitted by different combinations of B , R , δ , and η_e with little variation in the shape of the electron energy distribution. Note that, for example, $[\nu F_{\nu}]_{\text{syn}} \propto R^3 \eta_e$, but at the same time $[\nu F_{\nu}]_{\text{ssc}} \propto R^4 \eta_e^2$. We can attempt to reduce this degeneracy by assuming that the observed main variability timescale is related to the size of the emission region and its Doppler factor according to the formula

$$t_{\text{var}} \simeq \frac{(1+z) R}{c \delta}. \quad (5)$$

The multifrequency data collected during the 4.5 month campaign (see Section 5) allow us to study the variability of Mrk 501 on timescales from months to a few days. We found that, during this time period, the multifrequency activity typically varied on a timescale of 5–10 days, with the exception of a few particular epochs when the source became very active in VHE γ -rays, and flux variations with timescales of a day or shorter were found at TeV energies. Nevertheless, it is important to stress that several authors concluded in the past that the dominant emission site of Mrk 501 is characterized by variability timescales longer than one day (see Kataoka et al. 2001, for a comprehensive study of the Mrk 501 variability in X-rays), and that the power in the intraday flickering of this source is small, in agreement with the results of our campaign. Nevertheless, one should keep in mind that this object is known for showing sporadic but extreme changes in its activity that can give flux variations on timescales as short as a few minutes (Albert et al. 2007a). In this work, we aim to model the average/typical behavior of Mrk 501 (corresponding to the 4.5 month campaign) rather than specific/short periods with outstanding activity, and hence we constrained the minimum (typical) variability timescale t_{var} in the model to the range 1–5 days.

Even with t_{var} fixed as discussed above, the reconstructed SED of Mrk 501 may be fitted by different combinations of B , R , δ , and η_e . Such a degeneracy between the main model parameters is an inevitable feature of the SSC modeling of blazars (e.g., Kataoka et al. 1999), and it is therefore necessary to impose additional constraints on the physical parameters of the dominant emission zone. Here, we argue that such constraints follow from the requirement for the electron energy distribution to be in agreement with the one resulting from the simplest prescription of the energy evolution of the radiating electrons within the emission region, as discussed below.

The idea of separating the sites for the particle acceleration and emission processes is commonly invoked in modeling different astrophysical sources of high-energy radiation, and blazar jets in particular. Such a procedure is not always justified, because interactions of ultrarelativistic particles with the magnetic field (leading to particle diffusion and convection in momentum space) are generally accompanied by particle radiative losses (and vice versa). On the other hand, if the characteristic timescale for energy gains is much shorter than the timescales for radiative cooling (t'_{rad}) or escape (t'_{esc}) from the system, the particle acceleration processes may be indeed approximated as

¹⁵⁸ For example, for a given critical (break) synchrotron frequency in the observed SED, the corresponding electron break Lorentz factor scales as $\gamma_{\text{br}} \propto 1/\sqrt{B \delta}$.

Table 2
Parameters of the Blazar Emission Zone in Mrk 501

| Parameter | Main SSC Fit Considered | Alternative SSC Fit |
|--|--|--|
| Magnetic field | $B = 0.015$ G | $B = 0.03$ G |
| Emission region size | $R = 1.3 \times 10^{17}$ cm | $R = 0.2 \times 10^{17}$ cm |
| Jet Doppler and bulk Lorentz factors | $\Gamma = \delta = 12$ | $\Gamma = \delta = 22$ |
| Equipartition parameter | $\eta_e \equiv U'_e/U'_B = 56$ | $\eta_e \equiv U'_e/U'_B = 130$ |
| Minimum electron energy | $\gamma_{\min} = 600$ | $\gamma_{\min} = 300$ |
| Intrinsic electron break energy | $\gamma_{\text{br},1} = 4 \times 10^4$ | $\gamma_{\text{br},1} = 3 \times 10^4$ |
| Cooling electron break energy | $\gamma_{\text{br},2} = 9 \times 10^5$ | $\gamma_{\text{br},2} = 5 \times 10^5$ |
| Maximum electron energy | $\gamma_{\max} = 1.5 \times 10^7$ | $\gamma_{\max} = 0.3 \times 10^7$ |
| Low-energy electron index | $s_1 = 2.2$ | $s_1 = 2.2$ |
| High-energy electron index | $s_2 = 2.7$ | $s_2 = 2.7$ |
| Electron index above the cooling break | $s_3 = 3.65$ | $s_3 = 3.5$ |
| Mean electron energy | $\langle \gamma \rangle \simeq 2400$ | $\langle \gamma \rangle \simeq 1200$ |
| Main variability timescale | $t_{\text{var}} \simeq 4$ days | $t_{\text{var}} \simeq 0.35$ day |
| Comoving electron energy density | $U'_e \simeq 0.5 \times 10^{-3}$ erg cm $^{-3}$ | $U'_e \simeq 4.6 \times 10^{-3}$ erg cm $^{-3}$ |
| Comoving magnetic field energy density | $U'_B \simeq 0.9 \times 10^{-5}$ erg cm $^{-3}$ | $U'_B \simeq 3.6 \times 10^{-5}$ erg cm $^{-3}$ |
| Comoving energy density of synchrotron photons | $U'_{\text{syn}} \simeq 0.9 \times 10^{-5}$ erg cm $^{-3}$ | $U'_{\text{syn}} \simeq 3.1 \times 10^{-5}$ erg cm $^{-3}$ |
| Comoving electron number density | $N'_e \simeq 0.3$ cm $^{-3}$ | $N'_e \simeq 4.6$ cm $^{-3}$ |
| Luminosity of the host galaxy | $L_{\text{star}} \simeq 3 \times 10^{44}$ erg s $^{-1}$ | $L_{\text{star}} \simeq 3 \times 10^{44}$ erg s $^{-1}$ |
| Jet power carried by electrons | $L_e \simeq 1.1 \times 10^{44}$ erg s $^{-1}$ | $L_e \simeq 0.85 \times 10^{44}$ erg s $^{-1}$ |
| Jet power carried by magnetic field | $L_B \simeq 2 \times 10^{42}$ erg s $^{-1}$ | $L_B \simeq 0.65 \times 10^{42}$ erg s $^{-1}$ |
| Jet power carried by protons ^a | $L_p \simeq 3 \times 10^{43}$ erg s $^{-1}$ | $L_p \simeq 4.2 \times 10^{43}$ erg s $^{-1}$ |
| Total jet kinetic power | $L_j \simeq 1.4 \times 10^{44}$ erg s $^{-1}$ | $L_j \simeq 1.3 \times 10^{44}$ erg s $^{-1}$ |
| Total emitted power | $L_{\text{em}} \simeq 9.7 \times 10^{42}$ erg s $^{-1}$ | $L_{\text{em}} \simeq 2.7 \times 10^{42}$ erg s $^{-1}$ |
| Isotropic synchrotron luminosity | $L_{\text{syn}} \simeq 10^{45}$ erg s $^{-1}$ | $L_{\text{syn}} \simeq 10^{45}$ erg s $^{-1}$ |
| Isotropic SSC luminosity | $L_{\text{SSC}} \simeq 2 \times 10^{44}$ erg s $^{-1}$ | $L_{\text{SSC}} \simeq 2 \times 10^{44}$ erg s $^{-1}$ |

Note. ^a Assuming one electron–proton pair per electron–positron pair, and mean proton Lorentz factor $\langle \gamma_p \rangle \sim 1$.

being “instantaneous,” and may be modeled by a single injection term $\dot{Q}(\gamma)$ in the simplified version of the kinetic equation

$$\frac{\partial n'_e(\gamma)}{\partial t} = -\frac{\partial}{\partial \gamma} \left[\gamma n'_e(\gamma) \right] - \frac{n'_e(\gamma)}{t'_{\text{esc}}} + \dot{Q}(\gamma), \quad (6)$$

describing a very particular scenario for the energy evolution of the radiating ultrarelativistic electrons.

It is widely believed that the above equation is a good approximation for the energy evolution of particles undergoing diffusive (first-order Fermi) shock acceleration, and cooling radiatively in the downstream region of the shock. In such a case, the term $\dot{Q}(\gamma)$ specifies the energy spectrum and the injection rate of the electrons freshly accelerated at the shock front and not affected by radiative losses, while the escape term corresponds to the energy-independent dynamical timescale for the advection of the radiating particles from the downstream region of a given size R , namely $t'_{\text{esc}} \simeq t'_{\text{dyn}} \simeq R/c$. The steady-state electron energy distribution is then very roughly $n'_e(\gamma) \sim t'_{\text{dyn}} \dot{Q}(\gamma)$ below the critical energy for which $t'_{\text{rad}}(\gamma) = t'_{\text{dyn}}$, and $n'_e(\gamma) \sim t'_{\text{rad}}(\gamma) \dot{Q}(\gamma)$ above this energy. Note that in the case of a power-law injection $\dot{Q}(\gamma) \propto \gamma^{-s}$ and a homogeneous emission region with dominant radiative losses of the synchrotron type, $t'_{\text{rad}}(\gamma) \propto \gamma^{-1}$, the injected electron spectrum is expected to steepen by $\Delta s = 1$ above the critical “cooling break” energy. This provides us with the additional constraint on the free model parameters for Mrk 501: namely, we require that the position of the second break in the electron energy distribution needed to fit the reconstructed SED, $\gamma_{\text{br},2}$, should correspond to the location of the cooling break for a given chosen set of the model free parameters.

Figure 10 (black curves) shows the resulting SSC model fit (summarized in Table 2) to the averaged broadband emission

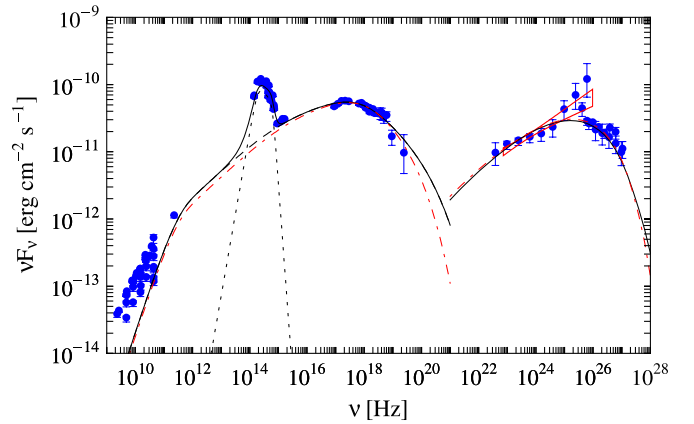


Figure 10. SSC model fits to the broadband emission spectrum of Mrk 501, averaged over all the observations made during the multifrequency campaign performed between 2009 March 15 (MJD 54905) and 2009 August 1 (MJD 55044). The red bow-tie in the figure corresponds to the 68% containment of the power-law fit to the average *Fermi*-LAT spectrum (photon index 1.74 ± 0.05). The dotted black curve denotes the fit to the starlight emission of the host galaxy assuming a template of a luminous elliptical as given in Silva et al. (1998). The details of the model are given in Section 6. The black curves correspond to the main set of the model parameters considered (variability timescale $t_{\text{var}} \simeq 4$ days), while the red dot-dashed curves to the alternative set of the model parameters with the emission region size decreased by an order of magnitude ($t_{\text{var}} \simeq 0.35$ days). See the text for further discussion.

spectrum of Mrk 501, which was obtained for the following parameters: $B = 0.015$ G, $R = 1.3 \times 10^{17}$ cm, $\delta = 12$, $\eta_e = 56$, $\gamma_{\min} = 600$, $\gamma_{\text{br},1} = 4 \times 10^4$, $\gamma_{\text{br},2} = 9 \times 10^5$, $\gamma_{\max} = 1.5 \times 10^7$, $s_1 = 2.2$, $s_2 = 2.7$, and $s_3 = 3.65$. The overall good agreement of the model with the data is further discussed in Section 6.2. Here, we note that, for these model parameters, synchrotron

self-absorption effects are important only below 1 GHz, where we do not have observations.¹⁵⁹ We also emphasize that with all the aforementioned constraints and for a given spectral shape of the synchrotron continuum (including all the data points aimed to be fitted by the model, as discussed below), and thus for a fixed spectral shape of the electron energy distribution (modulo critical electron Lorentz factors scaling as $\propto 1/\sqrt{B\delta}$), the allowed range for the free parameters of the model is relatively narrow. Namely, for the variability timescale between one and five days, the main model parameters may change within the ranges $R \simeq (0.35\text{--}1.45) \times 10^{17}$ cm, $\delta \simeq 11\text{--}14$, and $B \simeq 0.01\text{--}0.04$ G. The parameter η_e depends predominantly on the minimum Lorentz factor of the radiating electrons. Hence, it is determined uniquely as $\eta_e \simeq 50$ with the submillimeter flux included in the fitted data set. Only with a different prescription for the spectral shape of the electron energy distribution could the main free parameters of the model be significantly different from those given above.

Despite the absence of any fast variability during this multi-frequency campaign (apart from the already discussed isolated three-day-long flare), Mrk 501 is known for the extremely rapid flux changes at the highest observed photon energies (e.g., Albert et al. 2007a). Hence, it is interesting to check whether any shorter than few-day-long variability timescales can be accommodated in the framework of the simplest SSC model applied here for the collected data set. In order to do that, we decreased the minimum variability time scale by one order of magnitude (from four days to 0.4 days), and tried to model the data. A satisfactory fit could be obtained with those modified parameters, but only when we relaxed the requirement for the electron energy distribution to be in agreement with the one following from the steady-state solution to Equation (6), and in particular the resulting constraint for the second break in the electron spectrum to be equal to the cooling break. This “alternative” model fit is shown in Figure 10 (red curves) together with the “best” model fit discussed above. The resulting model parameters for the “alternative” fit are $B = 0.03$ G, $R = 2 \times 10^{16}$ cm, $\delta = 22$, $\eta_e = 130$, $\gamma_{\min} = 300$, $\gamma_{\text{br},1} = 3 \times 10^4$, $\gamma_{\text{br},2} = 5 \times 10^5$, $\gamma_{\max} = 3 \times 10^6$, $s_1 = 2.2$, $s_2 = 2.7$, and $s_3 = 3.5$. This particular parameter set—which should be considered as an illustrative one only—would be therefore consistent with a minimum variability timescale of 0.36 days, but at the price of much larger departures from the energy equipartition ($\eta_e > 100$). The other source parameters, on the other hand, would change only slightly (see Table 2). Because of the mismatch (by a factor ~ 3) between the location of the cooling break and the second break in the electron distribution, we consider this “alternative” fit less consistent with the hypothesis of the steady-state homogeneous one-zone SSC scenario, which is the framework with which we chose to model the broadband SED of Mrk 501 emerging from the campaign.

6.2. Notes on the Spectral Data Points

The low-frequency radio observations performed with single-dish instruments have a relatively large contamination from non-blazar emission due to the underlying extended jet component, and hence they only provide upper limits for the radio flux density of the blazar emission zone. On the other hand, the

flux measurements by the interferometric instruments (such as VLBA), especially the ones corresponding to the core, provide us with the radio flux density from a region that is not much larger than the blazar emission region.

The radio flux densities from interferometric observations (from the VLBA core) are expected to be close upper limits to the radio continuum of the blazar emission component. The estimated size of the partially resolved VLBA core of Mrk 501 at 15 GHz and 43 GHz is $\simeq 0.14\text{--}0.18$ mas $\simeq 2.9\text{--}3.7 \times 10^{17}$ cm (with the appropriate conversion scale 0.67 pc mas $^{-1}$). The VLBA size estimation is the FWHM of a Gaussian representing the brightness distribution of the blob, which could be approximated as 0.9 times the radius of a corresponding spherical blob (Marscher 1983). That implies that the size of the VLBA core is only a factor of two to three larger than the emission region in our SSC model fit ($R = 1.3 \times 10^{17}$ cm). Therefore, it is reasonable to assume that the radio flux density from the VLBA core is indeed dominated by the radio flux density of the blazar emission. Forthcoming multi-band correlation studies (in particular VLBA and SMA radio with the γ -rays from *Fermi*-LAT) will shed light on this particular subject. Interestingly, the magnetic field claimed for the partially resolved radio core of Mrk 501 (which has a size of $\lesssim 0.2$ mas) and its submilliarcsec jet, namely $B \simeq (10\text{--}30)$ mG (Giroletti et al. 2004, 2008), is in very good agreement with the value emerging from our model fits (15 mG), assuring self-consistency of the approach adopted.

In addition to this, in the modeling we also aimed at matching the submillimeter flux of Mrk 501, given at the observed frequency of 225 GHz, assuming that it represents the low-frequency tail of the optically thin synchrotron blazar component. One should emphasize in this context that it is not clear if the blazar emission zone is in general located deep within the millimeter photosphere or not. However, the broadband variability of luminous blazars of the FSRQ type indicates that there is a significant overlap of the blazar zone with a region where the jet becomes optically thin at millimeter wavelengths (as discussed by Sikora et al. 2008, for the particular case of the blazar 3C 454.3). We have assumed that the same holds for BL Lac objects.

The IR/optical flux measurements in the range $\sim (1\text{--}10) \times 10^{14}$ Hz represent the starlight of the host galaxy and hence they should be excluded when fitting the non-thermal emission of Mrk 501. We modeled these data points with the template spectrum of an elliptical galaxy instead (including only the dominant stellar component due to the evolved red giants, as discussed in Silva et al. 1998), obtaining a very good match (see the dotted line in Figure 10) for the bolometric starlight luminosity $L_{\text{star}} \simeq 3 \times 10^{44}$ erg s $^{-1}$. Such a luminosity is in fact expected for the elliptical host of a BL Lac object. The model spectrum of the galaxy falls off very rapidly above 5×10^{14} Hz, while the three UV data points (above 10^{15} Hz) indicate a prominent, flat power-law UV excess over the starlight emission. Therefore, it is reasonable to assume that the observed UV fluxes correspond to the synchrotron (blazar-type) emission of Mrk 501 and, consequently, we used them in our model fit. However, many elliptical galaxies do reveal in their spectra the so-called UV upturn, or UV excess, whose origin is not well known, but which is presumably related to the starlight continuum (most likely due to young stars from the residual star-forming activity within the central region of a galaxy) rather than to non-thermal (jet-related) emission processes (see, e.g., Code & Welch 1979; Atlee et al. 2009). Hence, it is possible that the UV data points

¹⁵⁹ The turnover frequency related to the synchrotron self-absorption may be evaluated using the formulae given in Ghisellini & Svensson (1991) and the parameter values from our SSC model fit as $\nu'_{\text{ssa}} \simeq 60$ MHz, which in the observer frame reads $\nu_{\text{ssa}} = \delta \nu'_{\text{ssa}} / (1 + z) \simeq 0.7$ GHz.

provided here include some additional contamination from the stellar emission, and as such might be considered as upper limits for the synchrotron radiation of the Mrk 501 jet.

The observed X-ray spectrum of Mrk 501 agrees very well with the SSC model fit, except for a small but statistically significant discrepancy between the model curve and the first two data points provided by *Swift*-XRT, which correspond to the energy range 0.3–0.6 keV. As pointed out in Section 5.1, the *Swift*-XRT data had to be corrected for a residual energy offset which affects the lowest energies. The correction for this effect could introduce some systematic differences with respect to the actual fluxes detected at those energies. These low-energy X-ray data points might also be influenced by intrinsic absorption of the X-ray photons within the gaseous environment of Mrk 501 nucleus, as suggested by the earlier studies with the ASCA satellite (see Kataoka et al. 1999). As a result, the small discrepancy between the data and the model curve within the range 0.3–0.6 keV can be ignored in the modeling.

The agreement between the applied SSC model and the γ -ray data is also very good. In particular, the model returns the γ -ray photon index 1.78 in the energy range 0.3–30 GeV, which can be compared with the one resulting from the power-law fit to the *Fermi*-LAT data above 0.3 GeV, namely 1.74 ± 0.05 . However, the last two energy bins from *Fermi* (60–160 and 160–400 GeV) are systematically above (2σ) the model curves, as well as above the averaged spectrum reported by MAGIC and VERITAS. A possible reason for mismatch between the average *Fermi*-LAT spectrum and the one from MAGIC/VERITAS was discussed in Section 5.3.

7. DISCUSSION

In this section, we discuss some of the implications of the model results presented above. After a brief analysis of the global parameters of the source resulting from the SSC fits (Section 7.1), the discussion focuses on two topics. First (Section 7.2), we show that the characteristics of the electron energy distribution emerging from our modeling can be used to constrain the physical processes responsible for the particle acceleration in Mrk 501, processes which may also be at work in other BL Lac-type objects. Second (Section 7.3), we examine the broadband variability of Mrk 501 in the framework of the model.

7.1. Main Characteristics of the Blazar Emission Zone in Mrk 501

The values for the emission region size $R = 1.3 \times 10^{17}$ cm and the jet Doppler factor $\delta = 12$ emerging from our SSC model fit give a minimum (typical) variability timescale of $t_{\text{var}} \simeq (1+z)R/c\delta \sim 4$ days, which is consistent with the variability observed during the campaign and with previous studies of the X-ray activity of Mrk 501 (Kataoka et al. 2001). At this point, it is necessary to determine whether an emission region characterized by these values of R and δ is optically thin to internal two-photon pair creation $\gamma\gamma \rightarrow e^+e^-$ for the highest TeV energies observed during the campaign. We now affirm pair transparency due to insufficient density of soft target photons.

Since Mrk 501 is a cosmologically local object, pair conversion in the EBL is not expected to prevent its multi-TeV photons from reaching the Earth, although the impact of this process is not negligible, as mentioned in Section 5. Therefore, dealing with a nearby source allows us to focus mostly on the intrinsic absorption processes, rather than on the cosmologi-

cal, EBL-related, attenuation of the γ -ray emission. Moreover, because of the absence (or weakness) of accretion-disk-related circumnuclear photon fields in BL Lac objects like Mrk 501, we only need to consider photon–photon pair production involving photon fields internal to the jet emission site. The analysis is therefore simpler than in the case of FSRQs, where the attenuation of high-energy γ -ray fluxes is dominated by interactions with photon fields external to the jet—such as those provided by the broad line regions or tori—for which the exact spatial distribution is still under debate.

Pair-creation optical depths can be estimated as follows. Using the δ -function approximation for the photon–photon annihilation cross-section (Zdziarski & Lightman 1985), $\sigma_{\gamma\gamma}(\varepsilon'_0, \varepsilon'_\gamma) \simeq 0.2 \sigma_T \varepsilon'_0 \delta[\varepsilon'_0 - (2m_e^2 c^4 / \varepsilon'_\gamma)]$, the corresponding optical depth for a γ -ray photon with observed energy $\varepsilon_\gamma = \delta \varepsilon'_\gamma / (1+z)$ interacting with a jet-originating soft photon with observed energy

$$\varepsilon_0 = \frac{\delta \varepsilon'_0}{1+z} \simeq \frac{2 \delta^2 m_e^2 c^4}{\varepsilon_\gamma (1+z)^2} \simeq 50 \left(\frac{\delta}{10} \right)^2 \left(\frac{\varepsilon_\gamma}{\text{TeV}} \right)^{-1} \text{ eV}, \quad (7)$$

may be found as

$$\begin{aligned} \tau_{\gamma\gamma} &\simeq \int^R ds \int_{m_e c^2 / \varepsilon_0} d\varepsilon_0 n'_0(\varepsilon_0) \sigma_{\gamma\gamma}(\varepsilon_0, \varepsilon_\gamma) \\ &\sim 0.2 \sigma_T R \varepsilon'_0 n'_0(\varepsilon_0), \end{aligned} \quad (8)$$

where $n'_0(\varepsilon'_0)$ is the differential comoving number density of soft photons. Noting that $\varepsilon'^2_0 n'_0(\varepsilon'_0) = L'_0 / 4\pi R^2 c$, where L'_0 is the intrinsic monochromatic luminosity at photon energy ε'_0 , we obtain

$$\begin{aligned} \tau_{\gamma\gamma} &\simeq \frac{\sigma_T d_L^2 F_0 \varepsilon_\gamma (1+z)}{10 R m_e^2 c^5 \delta^5} \simeq 0.001 \left(\frac{\varepsilon_\gamma}{\text{TeV}} \right) \\ &\times \left(\frac{F_0}{10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}} \right) \left(\frac{R}{10^{17} \text{ cm}} \right)^{-1} \left(\frac{\delta}{10} \right)^{-5}, \end{aligned} \quad (9)$$

where $F_0 = L_0 / 4\pi d_L^2$ is the observed monochromatic flux energy density as measured at the observed photon energy ε_0 . Thus, for 5 TeV γ -rays and the model parameters discussed (implying the observed $\varepsilon_0 = 15$ eV flux of Mrk 501 roughly $F_0 \simeq 3.2 \times 10^{-11} \text{ erg s}^{-1} \text{ cm}^{-2}$), one has $\tau_{\gamma\gamma}(5 \text{ TeV}) \simeq 0.005$. Therefore, the values of R and δ from our SSC model fit do not need to be adjusted to take into account the influence of spectral modifications due to pair attenuation. Note that such opacity effects, studied extensively in the context of γ -ray bursts, generally yield a broken power law for the spectral form, with the position and magnitude of the break fixed by the pair-production kinematics (e.g., Baring 2006 and references therein). The broad-band continuum of Mrk 501, and in particular its relatively flat spectrum VHE γ -ray segment, is inconsistent with such an expected break. This deduction is in agreement with the above derived transparency of the emitting region for high-energy γ -ray photons.

Next we evaluated the “monoenergetic” comoving energy density of ultrarelativistic electrons for a given electron Lorentz factor,

$$\gamma U'_e(\gamma) \equiv \gamma^2 m_e c^2 n'_e(\gamma), \quad (10)$$

and this is shown in Figure 11 (solid black line). The total electron energy density is then $U'_e = \int U'_e(\gamma) d\gamma \simeq 5 \times$

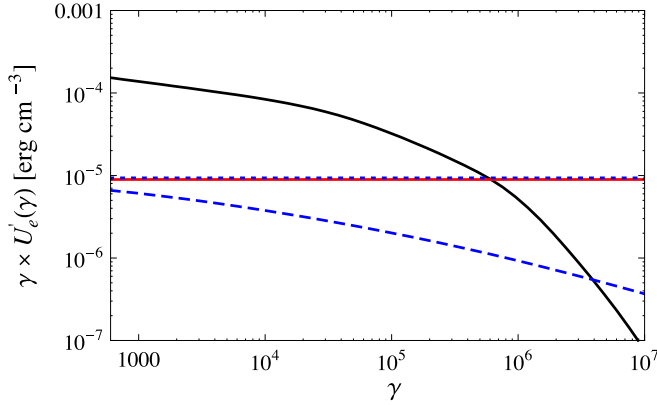


Figure 11. Jet comoving energy density of ultrarelativistic electrons per logarithmic energy bin, $\gamma U'_e(\gamma)$, as a function of the electron Lorentz factor γ (solid black curve). For comparison, the comoving energy density of the magnetic field U'_B (solid red line) and synchrotron photons U'_{syn} (dotted blue line) are shown. The dashed blue curve denotes the comoving energy density of synchrotron photons which are inverse-Compton upscattered in the Thomson regime, $U'_{syn/T}$, for a given electron Lorentz factor γ (see equation 12).

$10^{-4} \text{ erg cm}^{-3}$. As shown, most of the energy is stored in the lowest energy particles ($\gamma_{\min} \simeq 600$). For comparison, the comoving energy density of the magnetic field and that of the synchrotron photons are plotted in the figure as well (horizontal solid red line and dotted blue line, respectively). These two quantities are approximately equal, namely $U'_B = B^2/8\pi \simeq 0.9 \times 10^{-5} \text{ erg cm}^{-3}$ and

$$U'_{syn} = \frac{4\pi R}{3c} \int j'_{v',syn} dv' \simeq 0.9 \times 10^{-5} \text{ erg cm}^{-3}. \quad (11)$$

In Figure 11, we also plot the comoving energy density of synchrotron photons which are inverse-Compton upscattered in the Thomson regime for a given electron Lorentz factor γ ,

$$U'_{syn/T}(\gamma) = \frac{4\pi R}{3c} \int_{v'_{KN}(\gamma)}^{v'_{KN}(\gamma)} j'_{v',syn} dv' \quad (12)$$

(dashed blue line), where $v'_{KN}(\gamma) \equiv m_e c^2/4\gamma h$. Because of the well-known suppression of the inverse-Compton scattering rate in the Klein–Nishina regime, the scattering in the Thomson regime dominates the inverse-Compton losses.¹⁶⁰ Hence, one may conclude that even though the total energy density of the synchrotron photons in the jet rest frame is comparable to the comoving energy density of the magnetic field ($U'_{syn} \simeq U'_B$), the dominant radiative cooling for all the electrons is due to synchrotron emission, since $U'_{syn/T} < U'_B$ for any γ .

The timescale for synchrotron cooling may be evaluated as

$$t'_{syn} \simeq \frac{3m_e c}{4\sigma_T \gamma U'_B} \simeq 4 \left(\frac{\gamma}{10^7} \right)^{-1} \text{ days}. \quad (13)$$

Hence, $t'_{rad} \simeq t'_{syn}$ equals the dynamical timescale of the emitting region, $t'_{dyn} \simeq R/c$, for electron Lorentz factor $\gamma \simeq 8 \times 10^5$, i.e., close to the second electron break energy $\gamma_{br,2}$. Also, the difference between the spectral indices below and above the break energy $\gamma_{br,2}$ determined by our modeling, namely $\Delta s_{3/2} = s_3 - s_2 = 0.95$, is very close to the “classical”

¹⁶⁰ The inverse-Compton cross-section goes as $\sigma_{ic} \simeq \sigma_T$ for $v' < v'_{KN}(\gamma)$, and roughly as $\sigma_{ic} \sim \sigma_T (v'/v'_{KN})^{-1} \ln[v'/v'_{KN}]$ for $v' > v'_{KN}(\gamma)$ (e.g., Coppi & Blandford 1990).

synchrotron cooling break $\Delta s = 1$ expected for a uniform emission region, as discussed in Section 6.1. This agreement, which justifies at some level the assumed homogeneity of the emission zone, was in fact the additional constraint imposed on the model to break the degeneracy between the main free parameters. Note that in such a case the first break in the electron energy distribution around electron energy $\gamma_{br,1} = 4 \times 10^4$ is related to the nature of the underlying particle acceleration process. We come back to this issue in Section 7.2,

Another interesting result from our model fit comes from the evaluation of the mean energy of the electrons responsible for the observed non-thermal emission of Mrk 501. In particular, the mean electron Lorentz factor is

$$\langle \gamma \rangle \equiv \frac{\int \gamma n'_e(\gamma) d\gamma}{\int n'_e(\gamma) d\gamma} \simeq 2400. \quad (14)$$

This value, which is determined predominantly by the minimum electron energy $\gamma_{\min} = 600$ and by the position of the first break in the electron energy distribution, is comparable to the proton-to-electron mass ratio m_p/m_e . In other words, the mean energy of ultrarelativistic electrons within the blazar emission zone of Mrk 501 is comparable to the energy of non-relativistic/mildly-relativistic (cold) protons. This topic will be discussed further in Section 7.2 as well.

The analysis presented allows us also to access the global energetics of the Mrk 501 jet. In particular, with the given energy densities U'_e and U'_B , we evaluate the total kinetic powers of the jet stored in ultrarelativistic electrons and magnetic field as

$$L_e = \pi R^2 c \Gamma^2 U'_e \simeq 10^{44} \text{ erg s}^{-1}, \quad (15)$$

and

$$L_B = \pi R^2 c \Gamma^2 U'_B \simeq 2 \times 10^{42} \text{ erg s}^{-1}, \quad (16)$$

respectively. In the above expressions, we have assumed that the emission region analyzed occupies the whole cross-sectional area of the outflow, and that the jet propagates at sufficiently small viewing angle that the bulk Lorentz factor of the jet equals the jet Doppler factor emerging from our modeling, $\Gamma = \delta$. These assumptions are justified in the framework of the one-zone homogeneous SSC scenario. Moreover, assuming one electron–proton pair per electron–positron pair within the emission region (see Celotti & Ghisellini 2008), or equivalently $N'_p \simeq N'_e/3$, where the total comoving number density of the jet leptons is

$$N'_e = \int n'_e(\gamma) d\gamma \simeq 0.26 \text{ cm}^{-3}, \quad (17)$$

we obtain the comoving energy density of the jet protons $U'_p = \langle \gamma_p \rangle m_p c^2 N'_e/3$, and hence the proton kinetic flux $L_p = \pi R^2 c \Gamma^2 U'_p \simeq 0.3 \langle \gamma_p \rangle 10^{44} \text{ erg s}^{-1}$. This is comparable to the kinetic power carried out by the leptons for mean proton Lorentz factor $\langle \gamma_p \rangle \simeq 4$ (see Equation (15)). It means that, within the dominant emission zone of Mrk 501 (at least during non-flaring activity), ultrarelativistic electrons and mildly relativistic protons, if comparable in number, are in approximate energy equipartition, and their energy dominates that of the jet magnetic field by two orders of magnitude. It is important to compare this result with the case of powerful blazars of the FSRQ type, for which the relatively low mean energy of the radiating electrons, $\langle \gamma \rangle \ll 10^3$, assures dynamical domination of cold protons even for a smaller proton content $N'_p/N'_e \lesssim 0.1$ (see the discussion in Sikora et al. 2009 and references therein).

Assuming $\langle \gamma_p \rangle \sim 1$ for simplicity, we find that the implied total jet power $L_j = L_e + L_p + L_B \simeq 1.4 \times 10^{44} \text{ erg s}^{-1}$ constitutes only a small fraction of the Eddington luminosity $L_{\text{Edd}} = 4\pi G M_{\text{BH}} m_p c / \sigma_T \simeq (1.1\text{--}4.4) \times 10^{47} \text{ erg s}^{-1}$ for the Mrk 501 black hole mass $M_{\text{BH}} \simeq (0.9\text{--}3.5) \times 10^9 M_\odot$ (Barth et al. 2002). In particular, our model implies $L_j/L_{\text{Edd}} \sim 10^{-3}$ in Mrk 501. In this context, detailed investigation of the emission-line radiative output from the Mrk 501 nucleus by Barth et al. (2002) allowed for an estimate of the bolometric, accretion-related luminosity as $L_{\text{disk}} \simeq 2.4 \times 10^{43} \text{ erg s}^{-1}$, or $L_{\text{disk}}/L_{\text{Edd}} \sim 10^{-4}$. Such a relatively low luminosity is not surprising for BL Lac objects, which are believed to accrete at low, sub-Eddington rates (e.g., Ghisellini et al. 2009a). For a low-accretion rate AGN (i.e., those for which $L_{\text{disk}}/L_{\text{Edd}} < 10^{-2}$) the expected radiative efficiency of the accretion disk is $\eta_{\text{disk}} \equiv L_{\text{disk}}/L_{\text{acc}} < 0.1$ (Narayan & Yi 1994; Blandford & Begelman 1999). Therefore, the jet power estimated here for Mrk 501 is comfortably smaller than the available power of the accreting matter L_{acc} .

Finally, we note that the total emitted radiative power is

$$L_{\text{em}} \simeq \Gamma^2 (L'_{\text{syn}} + L'_{\text{ssc}}) = 4\pi R^2 c \Gamma^2 (U'_{\text{syn}} + U'_{\text{ssc}}) \sim 10^{43} \text{ erg s}^{-1}, \quad (18)$$

where U'_{syn} is given in Equation (11) and the comoving energy density of γ -ray photons, U'_{ssc} , was evaluated in an analogous way as $\simeq 1.7 \times 10^{-6} \text{ erg cm}^{-3}$. This implies that the jet/blazar radiative efficiency was at the level of a few percent ($L_{\text{em}}/L_j \simeq 0.07$) during the period covered by the multifrequency campaign. Such a relatively low radiative efficiency is a common characteristic of blazar jets in general, typically claimed to be at the level of 1%–10% (see Celotti & Ghisellini 2008; Sikora et al. 2009). On the other hand, the isotropic synchrotron and SSC luminosities of Mrk 501 corresponding to the observed average flux levels are $L_{\text{syn}} = \delta^4 L'_{\text{syn}} \simeq 10^{45} \text{ erg s}^{-1}$ and $L_{\text{ssc}} = \delta^4 L'_{\text{ssc}} \simeq 2 \times 10^{44} \text{ erg s}^{-1}$, respectively.

7.2. Electron Energy Distribution

The results of the SSC modeling presented in the previous sections indicate that the energy spectrum of freshly accelerated electrons within the blazar emission zone of Mrk 501 is of the form $\propto E_e^{-2.2}$ between electron energy $E_{e,\text{min}} = \gamma_{\text{min}} m_e c^2 \sim 0.3 \text{ GeV}$ and $E_{e,\text{br}} = \gamma_{\text{br},1} m_e c^2 \sim 20 \text{ GeV}$, steepening to $\propto E_e^{-2.7}$ above 20 GeV, such that the mean electron energy is $\langle E_e \rangle \equiv \langle \gamma \rangle m_e c^2 \sim 1 \text{ GeV}$. At this point, a natural question arises: is this electron distribution consistent with the particle spectrum expected for a diffusive shock acceleration process? Note in this context that the formation of a strong shock in the innermost parts of Mrk 501 might be expected around the location of the large bend (change in the position angle by 90°) observed in the outflow within the first few parsecs (projected) from the core (Edwards & Piner 2002; Piner et al. 2009). This distance scale could possibly be reconciled with the expected distance of the blazar emission zone from the center for the model parameters discussed, $r \sim R/\theta_j \sim 0.5 \text{ pc}$, for jet opening angle $\theta_j \simeq 1/\Gamma \ll 1$.

In order to address this question, let us first discuss the minimum electron energy implied by the modeling, $E_{e,\text{min}} \sim 0.3 \text{ GeV}$. In principle, electrons with lower energies may be present within the emission zone, although their energy distribution has to be very flat (possibly even inverted), in order not to overproduce the synchrotron radio photons and to account for the measured *Fermi*-LAT γ -ray continuum. Therefore, the con-

strained minimum electron energy marks the injection threshold for the main acceleration mechanism, meaning that only electrons with energies larger than $E_{e,\text{min}}$ are picked up by this process to form the higher-energy (broken) power-law tail. Interestingly, the energy dissipation mechanisms operating at the shock fronts do introduce a particular characteristic (injection) energy scale, below which the particles are not freely able to cross the shock front. This energy scale depends on the shock microphysics, in particular on the thickness of the shock front. The shock thickness, in turn, is determined by the operating inertial length, or the diffusive mean free path of the radiating particles, or both. Such a scale depends critically on the constituents of the shocked plasma. For pure pair plasmas, only the electron thermal scale enters, and this sets $E_{e,\text{min}} \sim \Gamma m_e c^2$. In contrast, if there are approximately equal numbers of electrons and protons, the shock thickness can be relatively large. Diffusive shock acceleration can then operate on electrons only above a relatively high energy, establishing $E_{e,\text{min}} \sim \epsilon \Gamma m_p c^2$. Here, ϵ represents some efficiency of the equilibration in the shock layer between shocked thermal protons and their electron counterparts, perhaps resulting from electrostatic potentials induced by charge separation of species of different masses (Baring & Summerlin 2007). Our multifrequency model fits suggest that $\epsilon \sim 0.025$, providing an important blazar shock diagnostic.

At even lower electron energies, other energization processes must play a dominant role (e.g., Hoshino et al. 1992), resulting in formation of very flat electron spectra (see the related discussion in Sikora et al. 2009). Which of these energy dissipation mechanisms are the most relevant, as well as how flat the particle spectra could be, are subjects of ongoing debates. Different models and numerical simulations presented in the literature indicate a wide possible range for the lowest-energy particle spectral indices (below $E_{e,\text{min}}$), from $s_{\text{inj}} \simeq 1.0\text{--}1.5$ down to $s_{\text{inj}} \simeq -2$, depending on the particular shock conditions and parameters (Amato & Arons 2006; Sironi & Spitkovsky 2009).

All in all, we argue that the relatively high minimum energy of the radiating electrons implied by the SSC modeling of the Mrk 501 broadband spectrum and the overall character of the electron energy evolution in this source are consistent with a proton-mediated shock scenario. In addition, the fact that the mean energy of the ultrarelativistic electrons is of the order of the proton rest energy, $\langle E_e \rangle \sim m_p c^2$, can be reconciled with such a model. Moreover, the constrained power-law slope of the electrons with energies $E_{e,\text{min}} \leq E_e \leq E_{e,\text{br}}$, namely $s_1 = 2.2$, seems to suggest a dominant role for diffusive shock acceleration above the injection energy $E_{e,\text{min}}$, as this value of the spectral index is often claimed in the literature for particles undergoing first-order Fermi acceleration at relativistic shock (Bednarz & Ostrowski 1998; Kirk et al. 2000; Achterberg et al. 2001). The caveat here is that this result for the “universal” particle spectral index holds only for particular conditions (namely, for ultrarelativistic shock velocities with highly turbulent conditions downstream of the shock: see the discussion in Ostrowski & Bednarz 2002), whereas in general a variety of particle spectra may result from the relativistic first-order Fermi mechanism, depending on the local magnetic field and turbulence parameters at the shock front, and the speed of the upstream flow (Kirk & Harens 1989; Niemiec & Ostrowski 2004; Lemoine et al. 2006; Sironi & Spitkovsky 2009; Baring 2010). Nevertheless, the evidence for ultrarelativistic electrons with spectral index $s_1 = 2.2$ in the jet of Mrk 501 may be considered as an indication that the plasma conditions within the blazar emission zone allow for efficient diffusive shock acceleration (at least in this source),

as described by the simplest asymptotic test-particle models, though only in a relatively narrow electron energy range.

If the relativistic shock acceleration plays a dominant role in the blazar Mrk 501 (as argued above), the observations and the model results impose important constraints on this mechanism, many aspects of which are still not well understood. Firstly, this process must be very efficient in the sense that all the electrons pre-accelerated/preheated to the energy $E_{e, \min} \sim 0.3$ GeV are picked up by the acceleration process so that a single electron component is formed above the injection threshold and there is no Maxwellian-like population of particles around $E_{e, \min}$ outnumbering the higher energy ones from the power-law tail.¹⁶¹ The second constraint is due to the presence of the spectral break $\Delta s_{2/1} = s_2 - s_1 \simeq 0.5$ around electron energies $E_{e, br} \sim 20$ GeV. As discussed in the previous section, this break cannot be simply a result of cooling or internal pair-attenuation effects, and hence must be accounted for by the acceleration mechanism. The discussion regarding the origin of this break—which may reflect variations in the global field orientation or turbulence levels sampled by particles of different energy—is beyond the scope of this paper. However, the presence of high-energy breaks in the electron energy distribution (intrinsic to the particle spectrum rather than forming due to cooling or absorption effects) seems to be a common property of relativistic jet sources observed by *Fermi*-LAT, such as the FSRQ objects 3C 454.3 and AO 0235+164 (Abdo et al. 2009, 2010a).

7.3. Variability

In Sections 3–4, we reported on the γ -ray flux and spectral variability of Mrk 501 as measured by the *Fermi*-LAT instrument during the first 16 months of operation. In this section, we discuss whether the observed spectral evolution can be accounted for by our simple one-zone SSC model.

Figure 12 (top panel) presents in more detail the GeV–TeV γ -ray spectrum of Mrk 501, together with the decomposition of the SSC model continuum. Here the contributions of different segments of the electron energy distribution are indicated by different colors. As shown, the low-energy electrons, $\gamma_{\min} \leq \gamma < \gamma_{br, 1}$, which emit synchrotron photons up to $\simeq 10^{15}$ Hz, dominate the production of γ -rays up to a few GeV (red line). The contribution of higher energy electrons with Lorentz factors $\gamma_{br, 1} \leq \gamma < \gamma_{br, 2}$ is pronounced within the observed synchrotron range $10^{15} - 10^{18}$ Hz, and at γ -ray energies from a few GeV up to \sim TeV (green line). Finally, the highest energy tail of the electron energy distribution, $\gamma \geq \gamma_{br, 2}$, responsible for the observed hard-X-ray synchrotron continuum (> 2 keV) in the fast cooling regime, generates the bulk of γ -rays with observed energies $> \text{TeV}$ (blue line). Interestingly, even though any sharp breaks in the underlying electron energy distribution are “smeared out” into a smoothly curved spectral continuum due to the nature of the SSC emission, the average data set does support the presence of distinct low-energy and high-energy segments in the electron spectrum.

It therefore seems reasonable to argue that the spectral variability of Mrk 501 observed by *Fermi*-LAT may be explained by postulating that the low-energy segment of the electron energy distribution ($\gamma < \gamma_{br, 1}$) is characterized by only small flux variations, while the high-energy electron tail ($\gamma > \gamma_{br, 1}$)

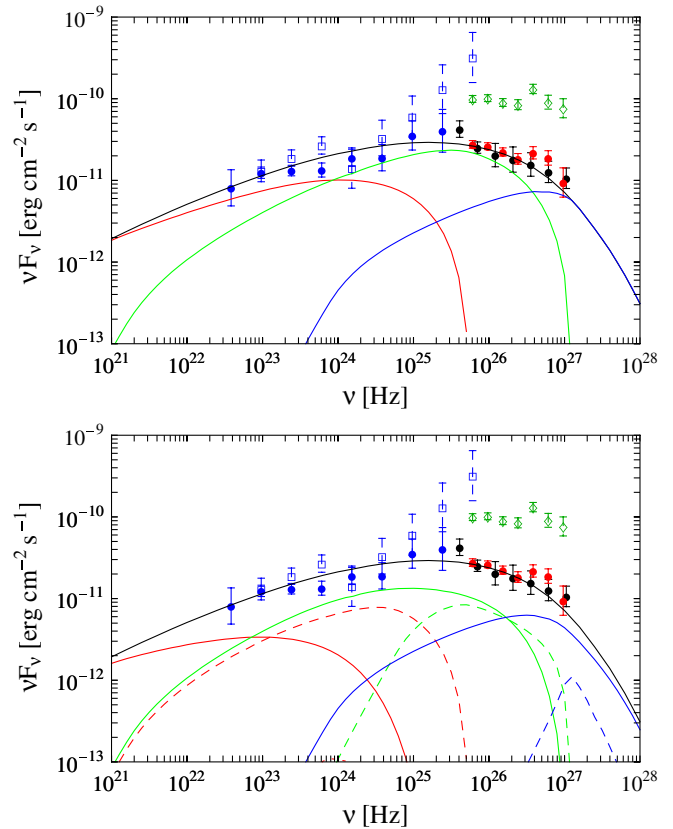


Figure 12. Decomposition of the SSC continuum for Mrk 501. The data points are the same as in the bottom panel of Figure 9. The SSC fit to the average spectrum is denoted by the solid black curve. Top: contributions of the different segments of the electron spectrum Comptonizing the whole synchrotron continuum (red curve: $\gamma_{\min} < \gamma < \gamma_{br, 1}$; green curve: $\gamma_{br, 1} < \gamma < \gamma_{br, 2}$; blue curve: $\gamma_{br, 2} < \gamma$). Bottom: contributions of the different segments of the electron spectrum (as in the top panel) Comptonizing different segments of the synchrotron continuum (solid curves: $\nu < \nu_{br, 1} \simeq 10^{15}$ Hz; dashed curves: $\nu_{br, 1} < \nu < \nu_{br, 2} \simeq 6 \times 10^{17}$ Hz; curves for $\nu > \nu_{br, 2}$ do not appear in the plot because the corresponding flux levels are all less than 10^{-13} erg cm $^{-2}$ s $^{-1}$).

varies more substantially. In such a scenario, some correlation might be expected between the fluxes in the UV-to-soft-X-ray photon energies from the synchrotron bump and the GeV–TeV fluxes from the inverse-Compton bump. This expectation is not inconsistent with the fact that we do not see any obvious relation between the ASM/BAT fluxes and the LAT (> 2 GeV) fluxes (see Figure 2), because the electrons producing X-ray synchrotron photons above 2–3 keV contribute to the SSC emission mostly at the highest photon energies in the TeV range (see Figure 12). This issue will be studied in detail (on timescales of 1 month down to 5 days) in a forthcoming publication using the data from this multifrequency campaign.

The X-ray/TeV connection has been established in the past for many BL Lac objects (and for Mrk 501 in particular). However, the exact character of the correlation between the X-ray and TeV fluxes is known to vary from object to object, and from epoch to epoch in a given source, as widely discussed in, e.g., Krawczynski et al. (2004), Błazejowski et al. (2005), Gliozzi et al. (2006), and Fossati et al. (2008). Note that the data analyzed in those papers were obtained mostly during periods of high activity. Consequently, the conclusions presented were somewhat biased towards flaring activity, and hence they might not apply to the typical (average) behavior of the source, which is the main focus of this paper. Moreover, the data set from

¹⁶¹ Note that due to low accretion rates and thus low luminosities of the accretion disks in BL Lac objects, the number of non-relativistic/mildly relativistic electrons (Lorentz factors $\gamma \sim 1$) cannot be constrained by analyzing “bulk-Compton” spectral features in the observed SED of Mrk 501, in contrast to the situation in FSRQs (see Sikora & Madejski 2000).

our campaign includes UV fluxes, soft-X-ray fluxes (down to 0.3 keV; *Swift*), and γ -ray fluxes spanning a very wide photon energy range (0.1 GeV–10 TeV; *Fermi*-LAT combined with MAGIC and VERITAS). This unique data set allows us to evaluate the multifrequency variability and correlations for Mrk 501 over an unprecedented range of photon energies.

Considering only the data set presented in this paper, we note that by steepening the high-energy electron continuum above the intrinsic break energy $\gamma_{br,1}$ (and only slightly adjusting the other model parameters), one can effectively remove photons above 10 GeV in the SSC component, leaving a relatively steep spectrum below 10 GeV, similar to the one observed by *Fermi*-LAT during the time interval MJD 54862–54892 (see Section 4). Such a change should be accompanied by a decrease in the UV-to-soft-X-ray synchrotron fluxes by a factor of a few, but the data available during that time interval are not sufficient to detect this effect.¹⁶² This statement is further justified by the bottom panel in Figure 12, where the contributions of the different segments of electrons Comptonizing the different segments of the synchrotron bump to the average γ -ray emission of Mrk 501 are shown. Note that the lowest-energy electron population ($\gamma < \gamma_{br,1}$) inverse-Compton upscatter only synchrotron photons emitted by the same population ($\nu < \nu_{br,1} \sim 10^{15}$ Hz; solid red line) may account for the bulk of the observed steep-spectrum γ -ray emission.

Another important conclusion from this figure is that Comptonization of the highest-energy synchrotron photons ($\nu > \nu_{br,2} \sim 6 \times 10^{17}$ Hz) by electrons with arbitrary energies produces only a negligible contribution to the average γ -ray flux of Mrk 501 due to the Klein–Nishina suppression. Thus, the model presented here explains in a natural way the fact that the X-ray and TeV fluxes of TeV-emitting BL Lac objects are rarely correlated according to the simple scaling $F_{TeV} \propto F_{keV}^2$ which would be expected from the class of SSC models in which the highest-energy electrons upscatter (in the Thomson regime) their own synchrotron photons to the TeV band (see, e.g., Gliazzi et al. 2006; Fossati et al. 2008). In addition, it opens a possibility for accommodating short-timescale variability ($t_{var} < 4$ days) at the highest synchrotron and inverse-Compton frequencies (hard X-rays and TeV photon energies, respectively). The reason for this is that, in the model considered here, these high-energy tails of the two spectral components are produced by the highest-energy electrons which are deep in the strong cooling regime (i.e., for which $t'_{rad} \ll R/c$), and thus the corresponding flux changes may occur on timescales shorter than $R/c \delta$ (see in this context, e.g., Chiaberge & Ghisellini 1999; Kataoka et al. 2000).

A more in-depth analysis of the multifrequency data set (including correlation studies of the variability in different frequency ranges) will be presented in a forthcoming paper. The epoch of enhanced γ -ray activity of Mrk 501 (MJD 54952–54982; see Section 4) may be more difficult to explain in the framework of the one-zone SSC model, because a relatively flat *Fermi*-LAT spectrum above 10 GeV, together with an increased TeV flux as measured by the VERITAS and Whipple 10 m telescopes around this time, may not be easy to reproduce with a set of model parameters similar to that discussed in previous sections. This is mostly due to Klein–Nishina effects, which tend to steepen the high-energy tail of the SSC component, thus precluding the formation of

a flat power law extending beyond the observed TeV energies. Hence, detailed modeling and data analysis will be needed to determine whether the enhanced VHE γ -ray activity period can be accommodated within a one-zone SSC model, or whether it will require a multi-zone approach.

8. CONCLUSIONS

We have presented a study of the γ -ray activity of Mrk 501 as measured by the LAT instrument on board the *Fermi* satellite during its first 16 months of operation, from 2008 August 5 (MJD 54683) to 2009 November 27 (MJD 55162). Because of the large leap in capabilities of LAT with respect to its predecessor, EGRET, this is the most extensive study to date of the γ -ray activity of this object at GeV–TeV photon energies. The *Fermi*-LAT spectrum (fitted with a single power-law function) was evaluated for 30-day time intervals. The average photon flux above 0.3 GeV was found to be $(2.15 \pm 0.11) \times 10^{-8}$ ph cm⁻² s⁻¹, and the average photon index 1.78 ± 0.03 . We observed only relatively mild (factor less than 2) γ -ray flux variations, but we detected remarkable spectral variability. In particular, during the four consecutive 30-day intervals of the “enhanced γ -ray flux” (MJD 54862–54982), the photon index changed from 2.51 ± 0.20 (for the first interval) down to 1.63 ± 0.09 (for the fourth one). During the whole period of 16 months, the hardest spectral index within the LAT energy range was 1.52 ± 0.14 , and the softest one was 2.51 ± 0.20 . Interestingly, this outstanding (and quite unexpected) variation in the slope of the GeV continuum did not correlate with the observed flux variations at energies above 0.3 GeV.

We compared the γ -ray activity measured by LAT in two different energy ranges (0.2–2 GeV and > 2 GeV) with the X-ray activity recorded by the all-sky instruments *RXTE*-ASM (2–10 keV) and *Swift*-BAT (15–50 keV). We found no significant difference in the amplitude of the variability between X-rays and γ -rays, and no clear relation between the X-ray and γ -ray flux changes. We note, however, that the limited sensitivity of the ASM and (particularly) the BAT instruments to detect Mrk 501 in a 30-day time interval, together with the relatively stable X-ray emission of Mrk 501 during the observations, precludes any detailed X-ray/ γ -ray variability or correlation analysis.

In this paper we also presented the first results from a 4.5 month multifrequency campaign on Mrk 501, which lasted from 2009 March 15 (MJD 54905) to 2009 August 1 (MJD 55044). During this period, the source was systematically observed with different instruments covering an extremely broad segment of the electromagnetic spectrum, from radio frequencies up to TeV photon energies. In this manuscript, we have focused on the average SED emerging from the campaign. Further studies on the multifrequency variability and correlations will be covered in a forthcoming publication.

We have modeled the average broadband spectrum of Mrk 501 (from radio to TeV) in the framework of the standard one-zone SSC model, obtaining a satisfactory fit to the experimental data. We found that the dominant emission region in this source can be characterized by a size of $R \simeq 10^3 r_g$, where $r_g \sim 1.5 \times 10^{14}$ cm is the gravitational radius of the black hole ($M_{BH} \simeq 10^9 M_\odot$) hosted by Mrk 501. The intrinsic (i.e., not affected by cooling or absorption effects) energy distribution of the radiating electrons required to fit the data was found to be of a broken power-law form in the energy range 0.3 GeV–10 TeV, with spectral indices 2.2 and 2.7 below and above the break energy of $E_{e,br} \sim 20$ GeV, respectively. In addition, the model parameters

¹⁶² The 4.5 month multifrequency campaign started 13 days after the end of the 30-day time interval MJD 54862–54892. Therefore, for this epoch, the only additional multifrequency data are from *RXTE*-ASM and *Swift*-BAT, which have only moderate ability to detect Mrk 501 on short timescales.

imply that all the electrons cool predominantly via synchrotron emission, forming a cooling break at 0.5 TeV. We argue that the particular form of the electron energy distribution emerging from our modeling is consistent with the scenario in which the bulk of the energy dissipation within the dominant emission zone of Mrk 501 is related to relativistic, proton-mediated shock waves. The low-energy segment of the electron energy distribution ($E_e < E_{e,br}$) formed thereby, which dominates the production of γ -rays observed below a few GeV, seems to be characterized by low and relatively slow variability. On the other hand, the high-energy electron tail ($E_e > E_{e,br}$), responsible for the bulk of the γ -rays detected above a few GeV, may be characterized by more significant variability.

Finally, we found that ultrarelativistic electrons and mildly-relativistic protons within the blazar zone of Mrk 501, if comparable in number, are in approximate energy equipartition, with their energy dominating the energy in the jet magnetic field by about two orders of magnitude. The model fit implies also that the total jet power, $L_j \simeq 10^{44}$ erg s $^{-1}$, constitutes only a small fraction of the Eddington luminosity, $L_j/L_{\text{Edd}} \sim 10^{-3}$, but is an order of magnitude larger than the bolometric, accretion-related luminosity of the central engine, $L_j/L_{\text{disk}} \sim 10$. Finally, we estimated the radiative efficiency of the Mrk 501 jet to be at the level of a few percent, $L_{\text{em}}/L_j \lesssim 0.1$, where L_{em} is the total emitted power of the blazar. The results from this study could perhaps be extended to all HSP BL Lac objects.

The *Fermi*-LAT Collaboration acknowledges the generous support of a number of agencies and institutes that have supported the *Fermi*-LAT Collaboration. These include the National Aeronautics and Space Administration and the Department of Energy in the United States, the Commissariat à l’Energie Atomique and the Centre National de la Recherche Scientifique / Institut National de Physique Nucléaire et de Physique des Particules in France, the Agenzia Spaziale Italiana and the Istituto Nazionale di Fisica Nucleare in Italy, the Ministry of Education, Culture, Sports, Science and Technology (MEXT), High Energy Accelerator Research Organization (KEK) and Japan Aerospace Exploration Agency (JAXA) in Japan, and the K. A. Wallenberg Foundation, the Swedish Research Council and the Swedish National Space Board in Sweden. Additional support for science analysis during the operations phase is gratefully acknowledged from the Istituto Nazionale di Astrofisica in Italy and the Centre National d’Études Spatiales in France.

The MAGIC collaboration would like to thank the Instituto de Astrofísica de Canarias for the excellent working conditions at the Observatorio del Roque de los Muchachos in La Palma. The support of the German BMBF and MPG, the Italian INFN, the Swiss National Fund SNF, and the Spanish MICINN is gratefully acknowledged. This work was also supported by the Marie Curie program, by the CPAN CSD2007-00042 and MultiDark CSD2009-00064 projects of the Spanish Consolider-Ingénio 2010 programme, by grant DO02-353 of the Bulgarian NSF, by grant 127740 of the Academy of Finland, by the YIP of the Helmholtz Gemeinschaft, by the DFG Cluster of Excellence “Origin and Structure of the Universe,” and by the Polish MNiSzW Grant N N203 390834.

VERITAS is supported by grants from the US Department of Energy, the US National Science Foundation, and the Smithsonian Institution, by NSERC in Canada, by Science Foundation Ireland, and by STFC in the UK. The VERITAS Collaboration also acknowledges the support of the *Fermi*/LAT team through the Guest Investigator Program Grant NNX09AT86G.

We acknowledge the use of public data from the *Swift* and *RXTE* data archive. The Metsähovi team acknowledges the support from the Academy of Finland to the observing projects (numbers 212656, 210338, among others). This research has made use of data obtained from the National Radio Astronomy Observatory’s Very Long Baseline Array (VLBA), projects BK150, BP143, and MOJAVE. The National Radio Astronomy Observatory is a facility of the National Science Foundation operated under cooperative agreement by Associated Universities, Inc. St. Petersburg University team acknowledges support from Russian RFBR foundation via grant 09-02-00092. AZT-24 observations are made within an agreement between Pulkovo, Rome and Teramo observatories. This research is partly based on observations with the 100-m telescope of the MPIfR (Max-Planck-Institut für Radioastronomie) at Effelsberg, as well as with the Medicina and Noto telescopes operated by INAF–Istituto di Radioastronomia. The Submillimeter Array is a joint project between the Smithsonian Astrophysical Observatory and the Academia Sinica Institute of Astronomy and Astrophysics and is funded by the Smithsonian Institution and the Academia Sinica. M. Villata organized the optical-to-radio observations by GASP-WEBT as the president of the collaboration. Abastumani Observatory team acknowledges financial support by the Georgian National Science Foundation through grant GNSF/ST07/4-180. The OVRO 40 m program was funded in part by NASA (NNX08AW31G) and the NSF (AST-0808050).

REFERENCES

- Abdo, A. A., et al. 2009, *ApJ*, **699**, 817
 Acciari, V. A. 2011, *ApJ*, in press (arXiv:1012.2200)
 Abdo, A. A., et al. 2010a, *ApJ*, **710**, 1271
 Abdo, A. A., et al. 2010b, *ApJS*, **188**, 405
 Abdo, A. A., et al. 2010c, *ApJ*, **716**, 30
 Acciari, V. A., et al. 2008, *ApJ*, **679**, 1427
 Acciari, V. A., et al. 2010, *ApJ*, **714**, 163A
 Achterberg, A., Gallant, Y. A., Kirk, J. G., & Guthmann, A. W. 2001, *MNRAS*, **328**, 393
 Aharonian, F. A., et al. 1999a, *A&A*, **342**, 69
 Aharonian, F. A., et al. 1999b, *A&A*, **349**, 11
 Aharonian, F. A., et al. 1999c, *A&A*, **349**, 29
 Aharonian, F. A., et al. 2001, *ApJ*, **546**, 898
 Ajello, M., et al. 2008, *ApJ*, **673**, 96
 Ajello, M., et al. 2009a, *ApJ*, **690**, 367
 Ajello, M., et al. 2009b, *ApJ*, **699**, 603
 Albert, J., et al. 2007a, *ApJ*, **669**, 862
 Albert, J., et al. 2007b, *Nucl. Instrum. Methods A*, **583**, 494
 Albert, J., et al. 2008, *ApJ*, **674**, 1037
 Aliu, E., et al. 2009, *Astropart. Phys.*, **30**, 293
 Aller, H. D., Aller, M. F., Latimer, G. E., & Hodge, P. E. 1985, *ApJS*, **59**, 513
 Amato, E., & Arons, J. 2006, *ApJ*, **653**, 325
 Anderhub, H., et al. 2009, *ApJ*, **705**, 1624
 Angelakis, E., Fuhrmann, L., Marchili, N., Krichbaum, T. P., & Zensus, J. A. 2008, *Mem. Soc. Astron. Ital.*, **79**, 1042
 Atlee, D. W., Assef, R. J., & Kochanek, C. S. 2009, *ApJ*, **694**, 1539
 Atwood, W. B., et al. 2009, *ApJ*, **697**, 1071
 Baring, M. G. 2006, *ApJ*, **650**, 1004
 Baring, M. G. 2010, *Adv. Space Res.*, in press (arXiv:1002.3848)
 Baring, M. G., & Summerlin, E. J. 2007, *Ap&SS*, **307**, 165
 Barth, A. J., Ho, L. C., & Sargent, W. L. W. 2002, *ApJ*, **566**, L13
 Barthelmy, S. D., et al. 2005, *Space Sci. Rev.*, **120**, 143
 Bednarek, W., & Protheroe, R. J. 1999, *MNRAS*, **310**, 577
 Bednarz, J., & Ostrowski, M. 1998, *Phys. Rev. Lett.*, **80**, 3911
 Blandford, R. D., & Begelman, M. C. 1999, *MNRAS*, **303**, L1
 Blazejowski, M., et al. 2005, *ApJ*, **630**, 130
 Blumenthal, G. R., & Gould, R. J. 1970, *Rev. Mod. Phys.*, **42**, 237
 Bradbury, S. M., et al. 1997, *A&A*, **320**, L5
 Bradt, H. V., Rothschild, R. E., & Swank, J. H. 1993, *A&AS*, **97**, 355
 Burrows, D. N., et al. 2005, *Space Sci. Rev.*, **120**, 165
 Catanese, M., et al. 1997, *ApJ*, **487**, L143

- Celotti, A., & Ghisellini, G. 2008, *MNRAS*, **385**, 283
- Chiaberge, M., & Ghisellini, G. 1999, *MNRAS*, **306**, 551
- Code, A. D., & Welch, G. A. 1979, *ApJ*, **228**, 95
- Coppi, P. S., & Blandford, R. D. 1990, *MNRAS*, **245**, 453
- Crusius, A., & Schlickeiser, R. 1986, *A&A*, **164**, L16
- Daum, A., et al. 1997, *Astrophys. J.*, **8**, 1
- Djannati-Ataï, A., et al. 1999, *A&A*, **350**, 17
- Edelson, R. A., & Krolik, J. H. 1988, *ApJ*, **333**, 646
- Edelson, R., et al. 2002, *ApJ*, **568**, 610
- Edwards, P. G., & Piner, B. G. 2002, *ApJ*, **579**, L67
- Finke, J. D., Dermer, C. D., & Böttcher, M. 2008, *ApJ*, **686**, 181
- Finke, J. D., Razzaque, S., & Dermer, C. D. 2010, *ApJ*, **712**, 238
- Fitzpatrick, E. L. 1999, *PASP*, **111**, 63
- Fossati, G., et al. 2008, *ApJ*, **677**, 906
- Franceschini, A., et al. 2008, *A&A*, **487**, 837
- Fuhrmann, L., et al. 2008, *A&A*, **490**, 1019
- Gehrels, N., et al. 2004, *ApJ*, **611**, 1005
- Ghisellini, G., Maraschi, L., & Tavecchio, F. 2009a, *MNRAS*, **396**, L105
- Ghisellini, G., & Svensson, R. 1991, *MNRAS*, **252**, 313
- Ghisellini, G., Tavecchio, F., Bodo, G., & Celotti, A. 2009b, *MNRAS*, **393**, L16
- Ghisellini, G., Tavecchio, F., & Chiaberge, M. 2005, *A&A*, **432**, 401
- Ghisellini, G., et al. 2002, *A&A*, **386**, 833
- Giannios, D., Uzdensky, D. A., & Begelman, M. C. 2009, *MNRAS*, **395**, L29
- Gilmore, R. C., Madau, P., Primack, J. R., Somerville, R. S., & Haardt, F. 2009, *MNRAS*, **399**, 1694
- Giroletti, M., et al. 2004, *ApJ*, **600**, 127
- Giroletti, M., et al. 2008, *A&A*, **488**, 905
- Giozzi, M., et al. 2006, *ApJ*, **646**, 61
- Graff, P. B., Georganopoulos, M., Perlman, E. S., & Kazanas, D. 2008, *ApJ*, **689**, 68
- Hartman, R. C., et al. 1999, *ApJS*, **123**, 79
- Hoshino, M., Arons, J., Gallant, Y. A., & Langdon, A. B. 1992, *ApJ*, **390**, 454
- Kalberla, P. M. W., et al. 2005, *A&A*, **440**, 775
- Kataoka, J., et al. 1999, *ApJ*, **514**, 138
- Kataoka, J., et al. 2000, *ApJ*, **528**, 243
- Kataoka, J., et al. 2001, *ApJ*, **560**, 659
- Katarzyński, K., Lenain, J.-P., Zech, A., Boisson, C., & Sol, H. 2008, *MNRAS*, **390**, 371
- Katarzyński, K., Sol, H., & Kus, A. 2001, *A&A*, **367**, 809
- Kino, M., Takahara, F., & Kusunose, M. 2002, *ApJ*, **564**, 97
- Kirk, J. G., Guthmann, A. W., Gallant, Y. A., & Achterberg, A. 2000, *ApJ*, **542**, 235
- Kirk, J. G., & Heavens, A. F. 1989, *MNRAS*, **239**, 995
- Kneiske, T. M., Bretz, T., Mannheim, K., & Hartmann, D. H. 2004, *A&A*, **413**, 807
- Kovalev, Y. Y., Nizhelsky, N. A., Kovalev, Y. A., Berlin, A. B., Zhekanis, G. V., Mingaliev, M. G., & Bogdantsov, A. V. 1999, *A&AS*, **139**, 545
- Krawczynski, H., et al. 2000, *A&A*, **353**, 97
- Krawczynski, H., et al. 2004, *ApJ*, **601**, 151
- Lemoine, M., Pelletier, G., & Revenu, B. 2006, *ApJ*, **645**, L129
- Lister, M. L., et al. 2009, *AJ*, **137**, 3718
- Marscher, A. P. 1983, *ApJ*, **264**, 296
- Massaro, E., Perri, M., Giommi, P., & Nesci, R. 2004, *A&A*, **413**, 489
- Massaro, E., Perri, M., Giommi, P., Nesci, R., & Verrecchia, F. 2004, *A&A*, **422**, 103
- Mattox, J. R., et al. 1996, *ApJ*, **461**, 396
- Nandra, K., et al. 1997, *ApJ*, **476**, 70
- Narayan, R., & Yi, I. 1994, *ApJ*, **428**, L13
- Niemiec, J., & Ostrowski, M. 2004, *ApJ*, **610**, 851
- Nilsson, K., Pasanen, M., Takalo, L. O., Lindfors, E., Berdyugin, A., Ciprini, S., & Pforr, J. 2007, *A&A*, **475**, 199
- Ostrowski, M., & Bednarz, J. 2002, *A&A*, **394**, 1141
- Panque, D. 2009, in ASP Conf. Ser. 427, *Accretion and Ejection in AGNs: A Global View*, ed. L. Maraschi et al. (San Francisco, CA: ASP), 277
- Perkins, J. S., et al. 2009, 2009 Fermi Symposium, eConf Proceedings C091122, arXiv:0912.3841v1
- Pian, E., et al. 1998, *ApJ*, **492**, L17
- Pichel, A. (for the VERITAS Collaboration) 2009, Proc. 31st ICRC (Łódź), arXiv:0908.0010
- Piner, B. G., Pant, N., Edwards, P. G., & Wiik, K. 2009, *ApJ*, **690**, L31
- Piron, F. 2000, in AIP Conf. Proc. 515, *Gev–TeV Gamma Ray Astrophysics Workshop: Towards a Major Atmospheric Cherenkov Detector VI*, ed. B. L. Dingus, M. H. Salamon, & D. B. Kieda (Melville, NY: AIP), 113
- Poole, T. S., et al. 2008, *MNRAS*, **383**, 627
- Quinn, J., et al. 1996, *ApJ*, **456**, L83
- Richards, J. L. 2011, *ApJS*, submitted (arXiv:1011.3111)
- Roming, P. W. A., et al. 2005, *Space Sci. Rev.*, **120**, 95
- Sambruna, R., et al. 2000, *ApJ*, **538**, 127
- Samuelson, F. W., et al. 1998, *ApJ*, **501**, L17
- Schlegel, D. J., Finkbeiner, D. P., & Davis, M. 1998, *ApJ*, **500**, 525
- Sikora, M., & Madejski, G. 2000, *ApJ*, **534**, 109
- Sikora, M., Moderski, R., & Madejski, G. M. 2008, *ApJ*, **675**, 71
- Sikora, M., Stawarz, Ł., Moderski, R., Nalewajko, K., & Madejski, G. M. 2009, *ApJ*, **704**, 38
- Silva, L., Granato, G. L., Bressan, A., & Danese, L. 1998, *ApJ*, **509**, 103
- Sironi, L., & Spitkovsky, A. 2009, *ApJ*, **698**, 1523
- Stawarz, Ł., & Petrosian, V. 2008, *ApJ*, **681**, 1725
- Tavecchio, F., et al. 2001, *ApJ*, **554**, 725
- Teräsanta, H., et al. 1998, *A&AS*, **132**, 305
- Tramacere, A., Giommi, P., Perri, M., Verrecchia, F., & Tosti, G. 2009, *A&A*, **501**, 879
- Vaughan, S., Edelson, R., Warwick, R. S., & Uttley, P. 2003, *MNRAS*, **345**, 1271
- Venturi, T., et al. 2001, *A&A*, **379**, 755
- Villata, M., & Raiteri, C. M. 1999, *A&A*, **347**, 30
- Villata, M., Raiteri, C. M., Lanteri, L., Sobrito, G., & Cavallone, M. 1998, *A&AS*, **130**, 305
- Villata, M., et al. 2008, *A&A*, **481**, L79
- Villata, M., et al. 2009, *A&A*, **504**, L9
- Zdziarski, A. A., & Lightman, A. P. 1985, *ApJ*, **294**, L79