



Publication Year	2018
Acceptance in OA	2020-12-21T12:15:17Z
Title	Ghost of a shell: magnetic fields of Galactic supershell GSH 006 - 15 + 7
Authors	Thomson, Alec J. M., McClure-Griffiths, N. M., Federrath, Christoph, Dickey, John M., CARRETTI, ETTORE, Gaensler, B. M., Haverkorn, Marijke, Kesteven, M. J., Staveley-Smith, Lister
Publisher's version (DOI)	10.1093/mnras/sty1865
Handle	http://hdl.handle.net/20.500.12386/29054
Journal	MONTHLY NOTICES OF THE ROYAL ASTRONOMICAL SOCIETY
Volume	479

Ghost of a shell: magnetic fields of Galactic supershell GSH 006 – 15 + 7

Alec J. M. Thomson,¹★ N. M. McClure-Griffiths,¹ Christoph Federrath,¹
John M. Dickey,² Ettore Carretti,³ B. M. Gaensler,⁴ Marijke Haverkorn,⁵
M. J. Kesteven⁶ and Lister Staveley-Smith⁷

¹Research School of Astronomy and Astrophysics, Australian National University, Canberra, ACT 2611, Australia

²School of Mathematics and Physics, University of Tasmania, Hobart TAS 7001, Australia

³INAF, Osservatorio Astronomico di Cagliari, Via della Scienza 5, I-09047 Selargius (CA), Italy

⁴Dunlap Institute for Astronomy and Astrophysics, University of Toronto, 50 St. George Street, Toronto ON, M5S 3G4, Canada

⁵Department of Astrophysics/IMAPP, Radboud University, PO Box 9010, NL-6500 GL Nijmegen, the Netherlands

⁶CSIRO Astronomy and Space Science, PO Box 76, Epping, NSW 1710, Australia

⁷International Centre for Radio Astronomy Research (ICRAR), The University of Western Australia, 35 Stirling Hwy, Crawley, WA 6009, Australia

Accepted 2018 July 9. Received 2018 July 9; in original form 2018 February 5

ABSTRACT

We identify a counterpart to a Galactic supershell in diffuse radio polarization, and use this to determine the magnetic fields associated with this object. GSH 006 – 15 + 7 has perturbed the polarized emission at 2.3 GHz, as observed in the S-band Polarization All Sky Survey (S-PASS), acting as a Faraday screen. We model the Faraday rotation over the shell, and produce a map of Faraday depth over the area across it. Such models require information about the polarized emission behind the screen, which we obtain from the Wilkinson Microwave Anisotropy Probe (WMAP), scaled from 23 to 2.3 GHz, to estimate the synchrotron background behind GSH 006 – 15 + 7. Using the modelled Faraday thickness, we determine the magnitude and the plane-of-the-sky structure of the line-of-sight magnetic field in the shell. We find a peak line-of-sight field strength of $|B_{\parallel}|_{\text{peak}} = 2.0_{-0.7}^{+0.01} \mu\text{G}$. Our measurement probes weak magnetic fields in a low-density regime (number densities of $\sim 0.4 \text{ cm}^{-3}$) of the ISM, thus providing crucial information about the magnetic fields in the partially-ionized phase.

Key words: polarization – ISM: bubbles – ISM: magnetic fields.

1 INTRODUCTION

H I shells, bubbles, supershells, and superbubbles are large structures in the interstellar medium (ISM) blown out by hot OB star clusters and supernovae. Both a supernova and the winds from a massive star in its main sequence lifetime will each inject around $10^{51} - 10^{53}$ ergs into the ISM. These winds and shocks ionize what will become the cavity of the shell, and sweep out the neutral material (McClure-Griffiths et al. 2002). It is now understood that these objects are strongly influenced by magnetic fields in their formation (Tomisaka 1990, 1998; Ferriere, Mac Low & Zweibel 1991; Slavin & Cox 1992; Stil et al. 2009; Gao et al. 2015; Ntormousi et al. 2017). Magnetic fields both oppose the expansion of the shell from the exterior, and prevent the collapse of the swept up shell walls (Ferrière 2001). H I shells have been discovered throughout our Galaxy (Heiles 1979; Hu 1981; Koo, Heiles & Reach 1992; Maciejewski et al. 1996; Uyaniker et al. 1999; McClure-Griffiths et al. 2000, 2002, 2006; Pidopryhora, Lockman & Shields 2007),

as well as external galaxies. These objects play a large role in determining the dynamics, evolution, and overall structure of the ISM (McClure-Griffiths et al. 2002).

Supershells and superbubbles are the largest classification of H I shells, with radii between 10^2 and 10^3 pc (Heiles 1979). Such objects occupy an intermediate size-scale within the ISM; a scale at which the role of magnetic fields in the magneto-ionic medium (MIM) is not well understood. Supershells are most commonly found in H I surveys, and appear as cavities in the neutral hydrogen (e.g. McClure-Griffiths et al. 2002); but they can also have multi-wavelength properties. Supershells can have associated emission in H α , soft X-rays, far ultra-violet, polarized radio continuum, and $100 \mu\text{m}$ (Heiles 1979, 1984; Reynolds et al. 1998; Heiles, Haffner & Reynolds 1999; Boumis et al. 2001; McClure-Griffiths et al. 2001, 2002; Ehlerová & Palouš 2005, 2013; Jo et al. 2011; Moss et al. 2012; Suad et al. 2014; Jo, Min & Seon 2015). It is thought that once bubbles expand far enough to break out of the gas of the Galactic plane, the shell breaks open into a Galactic chimney, allowing the flow of hot gas into the halo (Norman & Ikeuchi 1989). This transition from bubble to chimney is slowed by the presence

* E-mail: alec.thomson@anu.edu.au

of magnetic fields, which tend to confine the expanding shell in the disc (Tomisaka 1998).

A powerful method for probing the magnetic fields of the ISM is the study of Faraday rotation. This phenomenon describes how the polarization angle (PA) of a linearly polarized wave will rotate as the wave propagates through an MIM. Faraday rotation is measured from the linear Stokes parameters, Q and U . Following Burn (1966) and Brentjens & de Bruyn (2005), these can be parametrised as the complex polarization, \mathcal{P} :

$$\mathcal{P} = Q + iU = \text{PI}e^{2i\text{PA}}. \quad (1)$$

From this, the polarized intensity (PI) and the PA can therefore be defined as:

$$\begin{aligned} \text{PI} &= \sqrt{Q^2 + U^2} \\ \text{PA} &= \frac{1}{2} \arctan\left(\frac{U}{Q}\right). \end{aligned} \quad (2)$$

The amount of Faraday rotation at a wavelength λ is described by the Faraday depth (ϕ) times λ^2 :

$$\phi(L) \equiv 0.812 \int_0^L n_e B_{\parallel} dr \text{ rad m}^{-2}, \quad (3)$$

where B_{\parallel} is the line-of-sight magnetic field in μG , n_e is the electron density in cm^{-3} , and r is the distance along the line of sight integrated through the Faraday rotating medium with length L in pc. The sign of B_{\parallel} is taken to be positive when the field is aligned towards the observer and vice versa. In the simplest Faraday rotation case, where all the emission is in the background and all the Faraday rotation occurs in the foreground, the Faraday depth is given by the rotation measure (RM):

$$\text{RM} = \frac{\Delta\text{PA}}{\Delta(\lambda^2)}. \quad (4)$$

This also assumes no depolarization occurs along the line of sight. A method of obtaining Faraday depth in more complex scenarios is described in Brentjens & de Bruyn (2005).

Despite the important role magnetic fields play in the Galactic ISM, several mysteries remain unresolved. First, as summarized by Han (2017), the magnetic fields of a number of extended diffuse objects have been studied. However, obtaining the complete scale and structure of the magnetic fields associated with these objects is difficult. This arises as a result of both the methods used to measure these fields, and line-of-sight confusion from other magneto-ionic objects. Studies of RMs from both extragalactic point sources and pulsars can suffer from line-of-sight field reversals and confusion, and intrinsic Faraday rotation. These line-of-sight effects are a particular concern towards the Galactic plane. In this sense, diffuse polarization studies have a particular advantage in revealing large-scale, extended structures.

GSH 006 – 15 + 7 is a recently discovered Galactic supershell located near the Galactic plane (Moss et al. 2012). This supershell was discovered in H I observations with a central velocity around $v_{\text{LSR}} \approx 7 \text{ km s}^{-1}$, and subtends about 25 deg on the sky. Moss et al. (2012) constrain the age and distance to this object at $15 \pm 5 \text{ Myr}$ and $1.5 \pm 0.5 \text{ kpc}$, respectively; this gives the shell an approximate diameter of $670 \pm 220 \text{ pc}$, making it one of the largest discovered H I shells near the Sun. From their analysis, Moss et al. (2012) find that GSH 006 – 15 + 7 is likely in a break-out phase between a supershell and a chimney structure.

In this paper, we present the counterpart to GSH 006 – 15 + 7 in diffuse polarized radio emission data at 2.3 GHz. The shell appears as a ‘shadow’ in polarized emission, and shows evidence of Faraday

rotation of background synchrotron radiation. The data we use in this analysis is described in Section 2. In Section 3, we discuss this morphological association found in diffuse polarization at 2.3 GHz. We continue to use these polarization data to model the Faraday rotation through GSH 006 – 15 + 7 as a Faraday screen, following Sun et al. (2007) and Gao et al. (2015). From this, in Section 4 we constrain both the magnitude and the plane-of-sky structure of the line-of-sight magnetic fields associated with GSH 006 – 15 + 7. We use this information to estimate the thermal and magnetic pressures within the Galactic supershell. This dynamical information, as well as the magnetic field strengths themselves, will aid in the future modelling of supershells and the ISM, as well as in our overall understanding of ISM magnetohydrodynamics. Our conclusions are given in Section 5.

2 DATA

2.1 Diffuse H I emission

2.1.1 Galactic All-Sky Survey

We make use of H I data from the third release of the Parkes Galactic All-Sky Survey (GASS; Kalberla, McClure-Griffiths & Kerp 2009; McClure-Griffiths et al. 2009), now incorporated in HI4PI (Ben Bekhti et al. 2016). GASS is a fully sampled survey of H I emission over the entire sky south of declination zero with an angular resolution of 14.4 arcmin. The survey was conducted with the Parkes radio telescope using the 13-beam multibeam receiver. The data cover the velocity range $-468 \leq v_{\text{LSR}} \leq 468 \text{ km s}^{-1}$ with a velocity resolution of 1 km s^{-1} and a typical rms noise of 57 mK . In the third data release of GASS used here, the calibration and stray radiation correction were refined by Kalberla & Haud (2015). We use the GASS data to extract velocity separated H I emission of GSH 006 – 15 + 7.

2.2 Radio continuum polarization

2.2.1 S-band Polarization All Sky Survey

The S-band Polarization All Sky Survey (S-PASS; Carretti et al. 2013, Carretti et al. in preparation) was completed in 2010, and provides a highly sensitive polarization (Stokes Q and U) map of the Southern sky at 2.3 GHz. The survey was conducted using the Parkes 64 m Telescope with its ‘Galileo’ receiver and covers the Southern sky at declinations $\delta < -1 \text{ deg}$. This receiver operates in S band (13 cm) and is sensitive to circularly polarized radiation; allowing for the linear Stokes parameters, Q and U , to be measured. Table 1 lists the observational parameters for S-PASS. Initial morphological analysis was conducted by Carretti et al. (2013), and here we provide additional morphological descriptions; specifically we find a morphological correlation between the structure of GSH 006 – 15 + 7 in H I and the structure found in Stokes Q and U from S-PASS.

S-PASS supplies a number of significant improvements over previous polarization surveys. Observations of polarized emission at 2.3 GHz are inherently less prone to depolarization effects (Burn 1966; Sokoloff et al. 1998) with respect to lower frequencies. This higher observing frequency allows for greater angular resolution, with S-PASS presenting a gridded angular resolution (full width at half maximum FWHM) of 10.75 arcmin. The original data resolution was 8.9 arcmin, these were then smoothed with a Gaussian window of FWHM = 6 arcmin producing a map with a final resolution

Table 1. Observational parameters of S-PASS (e.g. Carretti et al. 2013) and WMAP band (Bennett et al. 2013). ^aMap noise is per observation.

Property	Symbol	S-PASS	WMAP
Reference frequency	ν	2307 MHz	22.69 GHz
Bandwidth	$\delta\nu$	184 MHz	4 GHz
Telescope beamwidth	FWHM_{tel}	8.9'	0.88 deg
Map beamwidth	FWHM_{map}	10.75'	–
Map RMS noise (Stokes Q/U)	σ	$\lesssim 1 \text{ mJy beam}^{-1}$	$1.435 \text{ mK} * (Q/U) - 6.00 \text{ mK}^a (I)$
Gain (Jy/K) at ν	A	$1 \text{ mJy} = 0.58 \text{ mK}$	–
System temperature	T_{sys}	$\approx 20 \text{ K}$	29 K

of 10.75 arcmin. Additionally, a much lower system temperature was achieved with respect to previous surveys in the same band. For example, S-PASS achieves a factor of two improvement over the Parkes 2.4 GHz polarization survey by Duncan et al. (1997), and they observed only a belt across the Galactic Plane ($|b| < 5^\circ$) not covering the area subject of this work. We use the S-PASS Stokes Q and U maps for computation of the RM associated with GSH 006 – 15 + 7.

2.2.2 Wilkinson Microwave Anisotropy Probe

The Wilkinson Microwave Anisotropy Probe (WMAP) survey released its final, 9-year data in 2013 (Bennett et al. 2013). The observational properties of WMAP are also summarized in Table 1. The project was focused on measuring cosmological parameters and the cosmic microwave background (CMB), however, the foreground data provides invaluable Galactic information. In particular, WMAP K band, centred on 23 GHz, gives a calibrated, whole-sky map of polarized synchrotron emission at high frequency. Discussion of the structure present in the WMAP K -band images was also conducted by Carretti et al. (2013), with a particular focus on comparison with the structure present in the S-PASS data. Specifically, the Northern and Southern Fermi bubbles feature prominently in the polarized emission of WMAP K band. We make use of the high-frequency polarization information provided by WMAP and compare these results with S-PASS data in the region of GSH 006 – 15 + 7.

2.3 Radio continuum

2.3.1 Continuum H I Parkes All-Sky Survey

The Continuum H I Parkes All-Sky Survey (CHIPASS) is a map of the radio continuum at 1.4 GHz across the whole sky below declination of $\delta = 25 \text{ deg}$ (Calabretta, Staveley-Smith & Barnes 2014). CHIPASS is a combination and reprocessing of the H I Parkes All-Sky Survey (HIPASS) and the H I Zone of Avoidance (HIZOA) survey, the result of which is a highly sensitive (sensitivity = 40 mK), all-sky, total intensity survey at a resolution of 14.4 arcmin, with very well-treated artefacts. We use these data to derive information on the synchrotron spectrum in the region of GSH 006 – 15 + 7.

3 RESULTS AND ANALYSIS

3.1 Polarization morphology

The most prominent feature in the diffuse polarization in the region of GSH 006 – 15 + 7 is the Southern lobe of the Fermi bubbles (Su, Slatyer & Finkbeiner 2010; Carretti et al. 2013; Ackermann et al. 2014), as seen in Fig. 2. There are number of structures, however, that show significant morphological correlation with the

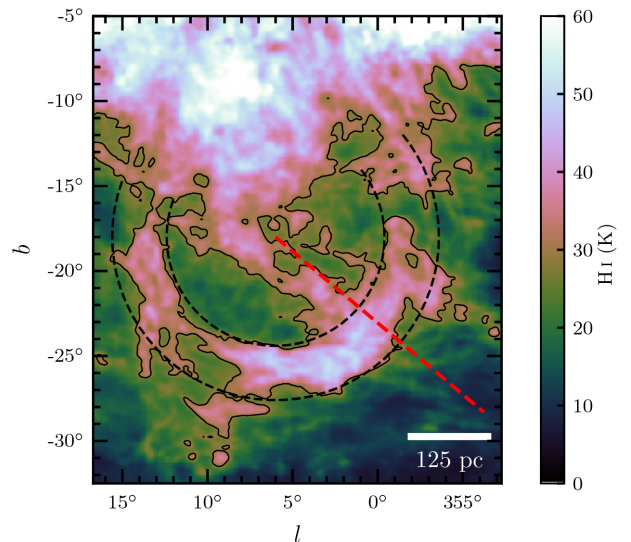


Figure 1. GASS – Brightness temperature of the H I 21 cm line in the region of the GSH 006 – 15 + 7 shell at $v_{\text{LSR}} = 6.6 \text{ km s}^{-1}$. All maps are given in Galactic longitude and latitude (l, b). Contour given at 30 K, which well outlines the H I emission associated with the shell. The black dashed lines give the approximate inner and outer bounds of the shell; approximately matching the 30 K contour. The scalebar gives the approximate size scale assuming a distance of 1.5 kpc. The red dashed line corresponds to the profile given in Fig. 4. This profile was chosen to match the one given in Fig. 2a.

supershell GSH 006 – 15 + 7. We claim that these features are perturbations caused by the MIM of GSH 006 – 15 + 7 to the polarized emission from behind it. This supershell has a distance estimate of $\sim 1.5 \pm 0.5 \text{ kpc}$ (Moss et al. 2012), and therefore is situated in the foreground relative to the Fermi bubbles (distance of the front surface from Sun $> 2.5 \text{ kpc}$; Carretti et al. 2013).

The spatial correlation between S-PASS polarization and GSH 006 – 15 + 7, as it appears in H I (see Fig. 1), is most apparent in Stokes parameters Q and U . The Stokes U image (see Fig. 2c) shows the strongest morphological correlation with GSH 006 – 15 + 7, although the structure seen in Stokes U is also similar to the structure seen in PA (see Fig. 2b). Along the upper and right-hand outer edge of the shell there is a significant shift in the values of Stokes U , from $U \approx +0.05 \text{ K}$ to $U \approx -0.05 \text{ K}$. A circular edge can be seen following the right-hand boundary of the shell, where predominately positive U switches to negative inside the boundary. On the left side of the image, where U appears negative, the magnitude of U increases on the shell. There is a bright, polarized region in the centre of the image, which corresponds to the tip of the Fermi bubble.

The Stokes Q image at 2.3 GHz (Fig. 2d) is dominated by emission from the Fermi bubble. Similar to Stokes U , there is a change

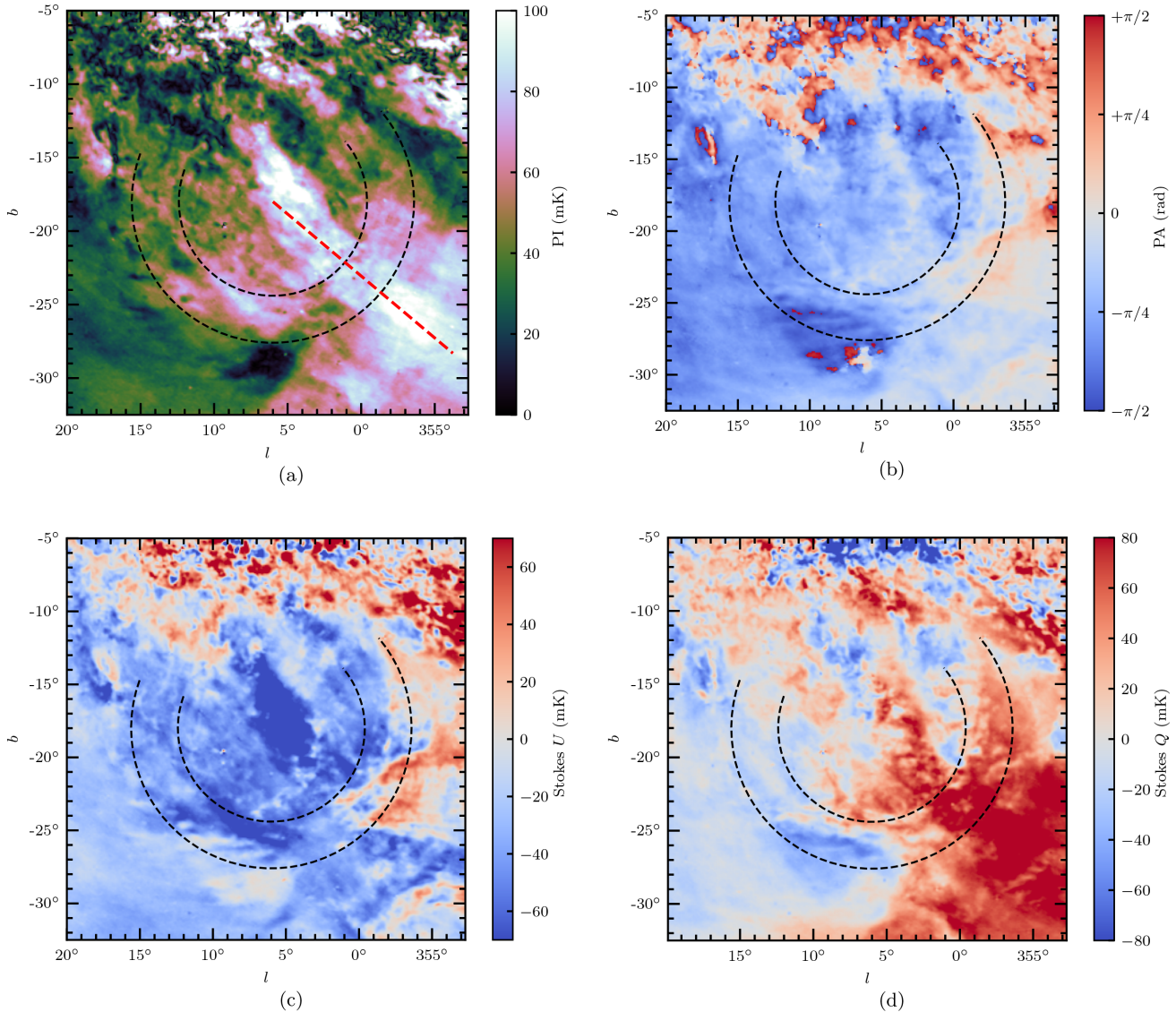


Figure 2. Maps of S-PASS polarization in the region of GSH 006 – 15 + 7. (a) S-PASS - PI. The red dashed line corresponds to the profile given in Fig. 4. This profile runs along the ridge of the polarized emission of the Fermi bubble from the centre of the shell region. (b) S-PASS - PA in radians. (c) S-PASS - Stokes U . (d) S-PASS - Stokes Q . In each panel, the black dashed lines give the approximate inner and outer bounds of the lower part of the shell.

of sign in Q along the right hand, inner boundary of the shell as defined by Hi. This region in Stokes Q , however, is not as clearly defined as Stokes U . Additionally, the change of sign occurs along the inner boundary of the shell in Q , rather than the outer boundary. A similar feature can be found along the bottom left, inner boundary of the shell. Here, Q again appears to change sign across the inner boundary of the shell. Along the left inner and outer boundaries of the shell, we find a weak change in sign of Stokes Q .

Inspecting the total linear polarization intensity (PI) at 2.3 GHz (Fig. 2a) the brightest polarized feature is the Southern ridge of the Fermi bubble (Carretti et al. 2013). This feature runs from the centre of the image to the bottom-right corner. Where the Fermi bubble appears to intersect the shell, however, the PI is reduced 10–15 per cent relative to the rest of the lobe. This is indicative that near the boundary of the shell, polarized emission is being perturbed and depolarized. Carretti et al. (2013) presented the first

morphological description of S-PASS, focusing particularly on the polarized emission from the Fermi bubbles. They also noted the perturbation feature and suggested that this feature could be due to a line-of-sight reversal of the magnetic field, associated with the Fermi bubbles. We argue however, that due to the high degree of spatial correlation, this feature is explained by the presence of GSH 006 – 15 + 7 in the foreground. This spatial correlation is exemplified in the profiles shown in Fig. 4. Here, we can see the anticorrelation of S-PASS PI with Hi brightness temperature, especially across the thickness of the shell from ~ 5 deg to ~ 10 deg along the profile. This demonstrates that the shell is weakly depolarizing the background polarized emission from the Fermi bubble. In addition, as depolarization is a Faraday rotation effect, we also note a change in the PA across the same boundary in Fig. 2b. Depolarization towards the Galactic plane is also apparent in the PI image, as noted by Carretti et al. (2013), as well as in Stokes Q

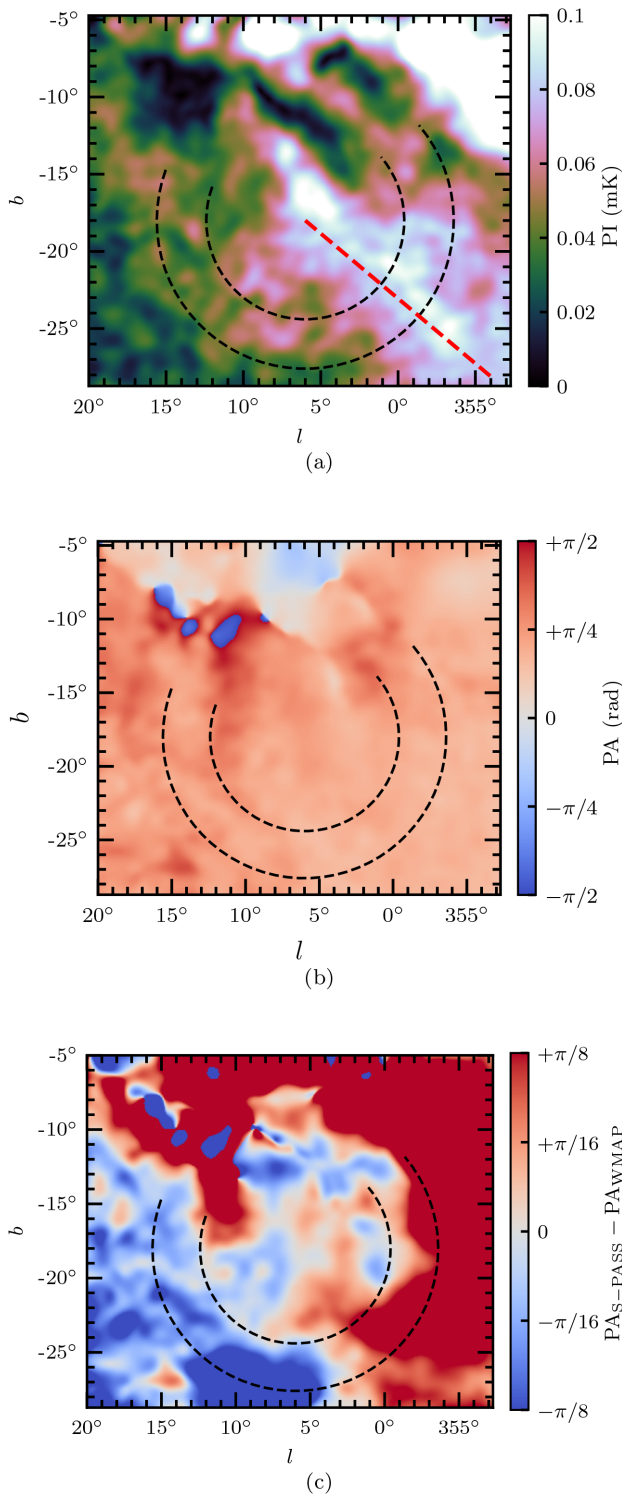


Figure 3. Panels (a) and (b): Maps of polarization in the region of GSH 006 – 15 + 7 at 23 GHz from WMAP K-band observations. Here, we have applied a Gaussian smoothing to the data with a FWHM of 1 deg. We applied this filter to increase the signal-to-noise ratio (SNR) in the WMAP data, which originally had a very high noise level. (a) WMAP – PI. Again, the red dashed line corresponds to the profile given in Fig. 4. (b) WMAP – PA in radians. (c) PA difference between S-PASS and WMAP in radians. Here, the S-PASS data were smoothed to common spatial resolution of the smoothed WMAP data. As in Fig. 2, the dashed lines give the approximate bounds of GSH 006 – 15 + 7.

and U . Significant depolarization is visible in the region near GSH 006 – 15 + 7 down to a latitude of $b \approx -5$ deg. This region is associated with large $H\text{II}$ regions in the Galactic plane, as seen in $H\alpha$ emission. Additional depolarization can be seen at latitudes as low as $b \approx -10$ deg. The modulation from this effect makes structure difficult to interpret in the polarization images near the Galactic plane.

Unlike the S-PASS polarization, we see no correlation with GSH 006 – 15 + 7 in any of the WMAP polarization maps. As seen in Fig. 3 the polarized emission at 23 GHz from the Southern Fermi bubble is the dominant source of polarized emission in the region of GSH 006 – 15 + 7. Carretti et al. (2013) also provide a detailed description of the polarized emission from WMAP in this region, again in regards to emission from the Fermi bubbles. We note that there is no noticeable depolarization across the thickness of the shell, unlike the S-PASS observations. This is particularly clear in Fig. 4 where, unlike the S-PASS profile, there is no appreciable drop in the WMAP PI profile as it intersects the shell. The lack of depolarization is to be expected, as high frequency observations are far less affected by Faraday rotation. We find that the PA is relatively uniform across the region of GSH 006 – 15 + 7 (see Fig. 3b), with a median value of ~ 37 deg and a standard deviation of ~ 9 deg. We also note the significant lack of polarized emission in both WMAP and S-PASS in a few regions above the shell along $b \approx -10$ deg. The signal to noise in these regions is therefore very low, particularly in WMAP, and also corresponds to large changes in the PA which are probably spurious. These regions will likely propagate large errors through this analysis.

The strongest morphological indication appears when the polarization information from S-PASS and WMAP are combined. Inspecting the map of the PA difference ($PA_{\text{S-PASS}} - PA_{\text{WMAP}}$) in Fig. 3(c) reveals the circular structure we expect to be associated with GSH 006 – 15 + 7. Again, here the right-hand portion of the shell is most prominent but the entire region is visible in Fig. 3(c), where the angle difference is of relatively low magnitude. This change in PA is indicative that the shell is Faraday rotating the background polarized emission. The motivation for using this PA difference is expanded upon in Section 3.3. In short, as the WMAP data was at high frequency, the polarization as measured should have encountered very little Faraday rotation, in contrast to the S-PASS polarization. Thus, the difference in PA contains information on how much Faraday rotation has occurred along the line of sight. As this difference bears a great deal of morphological similarity to GSH 006 – 15 + 7, we therefore assume that, within this region, it is the supershell alone that is causing the observed rotation effect.

3.2 Total intensity morphology

We also carefully inspect the total intensity data from S-PASS, CHIPASS, and WMAP for any morphological indication of GSH 006 – 15 + 7. The presence of such emission, especially in the S-PASS data, would mean that the interpretation that the shell was acting as a pure Faraday screen would be false. Faraday screens, by definition, do not produce any emission of their own. We find, however, no such emission in any of these data. The only structure of note is emission from the Corona Australis molecular cloud, visible in WMAP K-band total intensity. This comet-shaped reflection nebula is a common feature in infra-red observations (Schlegel, Finkbeiner & Davis 1998; Neuhäuser & Forbrich 2008), and appears to lie along the bottom edge of GSH 006 – 15 + 7. However, its overall morphology is very distinct from GSH 006 – 15 + 7, and

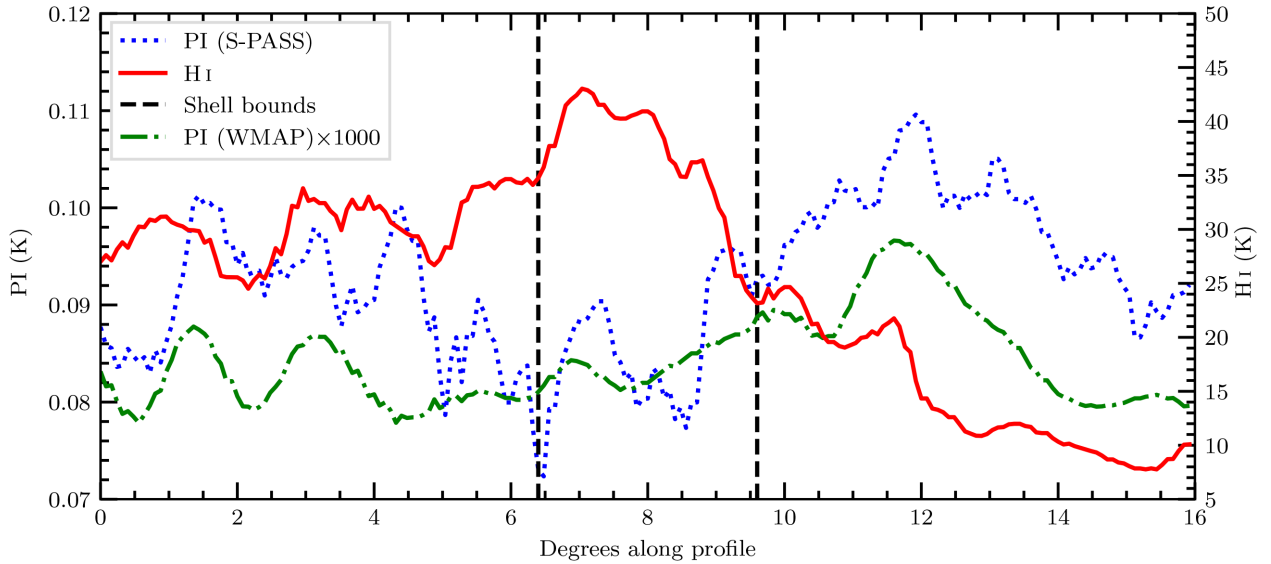


Figure 4. Brightness temperature profiles of PI at 2.3 GHz (blue, dotted), 23 GHz (green, dash-dotted), and GASS HI at $v_{\text{LSR}} = 6.6 \text{ km s}^{-1}$ (red, solid). The region from which the profiles are taken is shown in Figs 1 and 2 as the red dashed line. The profile runs from the centre of the shell region to well past the outer boundary, along the bright polarized emission from the Fermi bubble. The black dashed lines correspond to the inner and outer shell boundaries and also shown in Figs 1 and 2.

sits far closer to us at a distance of $\sim 130 \text{ pc}$ (Neuhäuser & Forbrich 2008).

3.3 Faraday screen model

3.3.1 General concept of the Faraday screen model

As discussed above, GSH 006 – 15 + 7 does present a clear polarization signature at 2.3 GHz. This signature however is not directly evident as polarized emission, nor is there a clear correlation in H α emission (Moss et al. 2012). This is indicative of a relatively low free electron density (compared to an H II region). This type of interaction is dubbed a ‘Faraday screen’ effect (Wolleben & Reich 2004; Sun et al. 2007). Faraday screens are simple Faraday rotating regions that affect the synchrotron emission that is produced from behind them (along the line of sight). Faraday screens do not produce polarized emission themselves, instead they rotate the background PA, and/or reduce the PI through depolarization effects. Sun et al. (2007) provide a model for determining the Faraday rotation that occurs due to the presence of a Faraday screen. This Faraday screen model is visualized in Fig. 5. In this scenario, an observer can measure polarization (\mathcal{P}) either ‘on’ or ‘off’ the screen. The terms ‘on’ and ‘off’ refer to whether the polarized emission has been affected by the Faraday screen or not, respectively. In either case the model assumes that the observer will measure the superposition of the polarized emission from the ‘background’ and ‘foreground’, relative to the screen. When observing ‘on’ the screen the background polarized emission will be perturbed. Specifically, the PI will be multiplied by the depolarization factor ‘ f ’ (where $f < 1$), and the PA will be rotated by an amount ‘ ψ ’. As such, the Stokes parameters ‘on’ the screen are given by (from Sun et al. 2007):

$$\begin{aligned} U_{\text{on}} &= \text{PI}_{\text{fg}} \sin(2\psi_0) + f \text{PI}_{\text{bg}} \sin[2(\psi_0 + \psi)] \\ Q_{\text{on}} &= \text{PI}_{\text{fg}} \cos(2\psi_0) + f \text{PI}_{\text{bg}} \cos[2(\psi_0 + \psi)] \end{aligned} \quad (5)$$

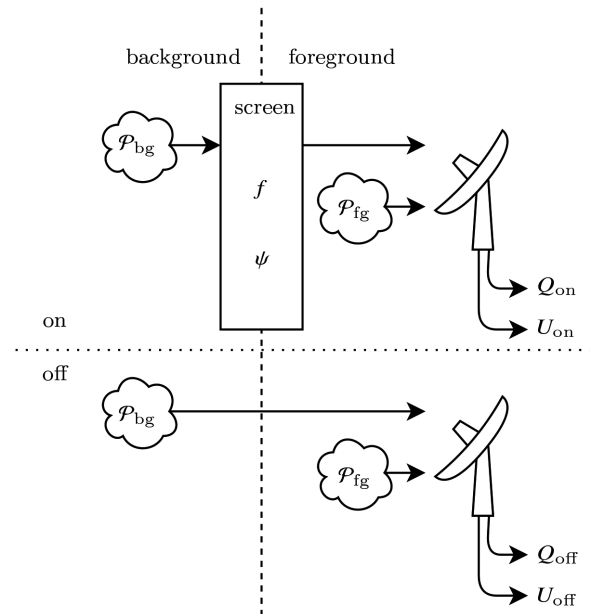


Figure 5. Cartoon of the Faraday screen model. The ‘foreground’/‘background’ divide refers to the line-of-sight region where the polarized emission is produced, relative to the screen. The ‘on’/‘off’ divide refers to whether the line-of-sight observes through the screen or not. In both the ‘on’ and ‘off’ case the measured polarization (\mathcal{P}) is the superposition of the ‘background’ and ‘foreground’ emission. In the ‘on’ case, however, the background emission is perturbed by the screen.

These equations assume that the intrinsic PAs in the background and foreground are related by $\text{PA}_{\text{bg}} = \text{PA}_{\text{fg}} = \psi_0$. This assumption is in contrast to Sun et al. (2007) who assumed that $\psi_0 \approx 0 \text{ deg}$. We are unable to make this same assumption, as it implies that $\text{PA}_{\text{off}} \approx 0 \text{ deg}$. If we take the WMAP PA information as an estimate

of PA_{off} , it is clear from Fig. 3(b) that this is not the case. Our assumption is reasonable in the case of either: (1) a random field or (2) a dominant coherent field in a single direction. From this, we derive the same model of a Faraday screen as Sun et al. (2007), using observations of PI and PA, ‘on’ and ‘off’ the screen:

$$\frac{PI_{\text{on}}}{PI_{\text{off}}} = \sqrt{f^2(1-c)^2 + c^2 + 2fc(1-c)\cos 2\psi}$$

$$PA_{\text{on}} - PA_{\text{off}} = \frac{1}{2} \arctan \left(\frac{f(1-c)\sin 2\psi}{c + f(1-c)\cos 2\psi} \right). \quad (6)$$

This model describes the change in PI and PA between observations ‘on’ and ‘off’ the screen. Here, the four observables are PI_{on} , PI_{off} , PA_{on} , and PA_{off} . The model parameters are f , c , and ψ . The parameters f and ψ are as described above, where f is the factor of depolarization that has occurred at the observed frequency, with $f \in [0, 1]$, while ψ is the amount of Faraday rotation through the screen, with $\psi \in [-\pi/2, +\pi/2]$. The parameter c is the fraction of foreground PI given by:

$$c = \frac{PI_{\text{fg}}}{(PI_{\text{fg}} + PI_{\text{bg}})}. \quad (7)$$

The range of this parameter is therefore $c \in [0, 1]$. The Faraday thickness (ϕ) of the screen can be estimated from this model by:

$$\phi = \frac{\psi}{\lambda^2}. \quad (8)$$

Here, ‘Faraday thickness’ refers to the Faraday depth of the shell alone. This equation applies in the case of a rotating-only Faraday screen, as the RM of the screen is equal to the Faraday depth. This model has an advantage over the model presented in Wolleben & Reich (2004), as it only requires observations at a single frequency.

In previous uses of this model the ‘on’ and ‘off’ components were taken from a common set of radio polarization observations at the same frequency, but were spatially offset to sample ‘on’ and ‘off’ the proposed location of the Faraday screen (Sun et al. 2007, 2011; Gao et al. 2010, 2015). There are two primary assumptions made by this technique: (1) that the polarized emission located beside (‘off’) the screen region accurately represents the superposition of synchrotron emission from the foreground and background, and (2) that all Faraday rotation along the line of sight occurs at the Faraday screen. In our analysis, the Faraday screen region of GSH 006 – 15 + 7 appears in the direction of the Galactic plane. As the line of sight approaches the plane, the amount of Faraday rotating structure increases; further, the angular extent of GSH 006 – 15 + 7 is enormous. As such, assumption (1) is not easily satisfied here.

3.3.2 Using WMAP to obtain ‘off’ polarization

In contrast to previous uses of the Sun et al. (2007) Faraday screen model, we use S -band-scaled, high frequency (23 GHz), polarization data as the ‘off’ measurements to the S-PASS ‘on’ measurements. By utilizing polarization data from a high frequency survey we bypass the need for assumption (1) to be satisfied. That is, rather than using a spatially offset pointing at the same frequency for the ‘off’ measurement, we use pointings towards the screen, but at a different frequency that is negligibly affected by the screen, thus serving as the ‘off’ measurement. Additionally, this allows for the computation of a Faraday depth for each line of sight through the Faraday screen.

With this consideration, we obtain high frequency polarization data from the WMAP survey. As Faraday rotation increases as λ^2 , only strong Faraday depth sources ($\phi > 100 \text{ rad m}^{-2}$) can produce

non-negligible Faraday rotation at 23 GHz. Since GSH 006 – 15 + 7 is an object in the diffuse ISM, likely occupying some combination of the warm neutral and warm ionized medium (Heiles & Haverkorn 2012), we do not expect significant Faraday thickness to be produced by this object.

A limitation of a Faraday screen model is that this technique is only sensitive to Faraday rotation values of $-90 \text{ deg} < \psi < +90 \text{ deg}$, due to assumption (2). Any rotations greater than this through the screen would create ‘ $n\pi$ ’ ambiguities in the measured PA. As such, we are only able to obtain Faraday depths between $\phi = \pm(\pi/2\lambda^2) \approx \pm 92 \text{ rad m}^{-2}$. We can exclude the possibility of ‘ $n\pi$ ’ ambiguities by inspecting the morphology of the PA image from S-PASS (see Fig. 2b). Regions that strongly Faraday rotate appear to have ‘onion skin’ structure in PA. That is, across these regions the PA changes sign multiple times towards the centre of the object. This corresponds to multiple revolutions of the polarization vector as the physical and Faraday depth of the object increases towards the centre. The GSH 006 – 15 + 7 region does not exhibit this morphology, rather it has an approximately uniform PA structure within the boundaries of the shell. This is therefore consistent with the assumption that the amount of rotation through the shell is $< |90 \text{ deg}|$.

WMAP K band does suffer low SNR as it measures synchrotron radiation at the far end of the spectrum. The low signal to noise of the WMAP data must be addressed in order to be used in analysis of the Galactic synchrotron emission. We do this by applying Gaussian smoothing to these images. We convolve a Gaussian smoothing function with a FWHM = 1 deg with both the Stokes Q and U maps.

3.3.3 Frequency scaling of WMAP to S band

In order to estimate the synchrotron background at S band, we need to scale WMAP data to 2.3GHz. To do this we require the synchrotron spectral index (β) between these frequencies. Between any two frequencies (ν) with total intensity (T), the spectral index is given by:

$$\beta = \frac{\log(T_1/T_2)}{\log(\nu_1/\nu_2)}. \quad (9)$$

As the zero-offset calibration total intensity map of S-PASS is not yet finalized, we instead obtain the synchrotron spectral index between 1.4 and 23 GHz. This assumes a spectrally constant β between these frequencies. We then use this value of β to scale the WMAP K -band data to S band. To find the spectral index, we use the total intensity data at 1.4 GHz from the CHIPASS survey (Calabretta et al. 2014). The synchrotron spectral index can usually be obtained from the slope (m) of the T–T plot of the total intensity (Stokes I) between two frequencies using (Turtle et al. 1962):

$$\beta \approx \frac{\log m}{\log(\nu_w/\nu_c)}, \quad (10)$$

where ‘ W ’ stands for WMAP, ‘ C ’ for CHIPASS. We find, however, that the slope of the T–T plot is not constant in the region of GSH 006 – 15 + 7, indicating thermal emission processes (see Section A1).

Rather than using this result, we compute the spectral index for each point in the region of GSH 006 – 15 + 7 using equation (9). We then analyse the indices as a function of latitude, as shown in Fig. 6. The spectral index remains relatively constant for $b < -10 \text{ deg}$, after which β starts to increase to values greater than -3 , as also indicated in the T–T plot. Following this, we bin the obtained

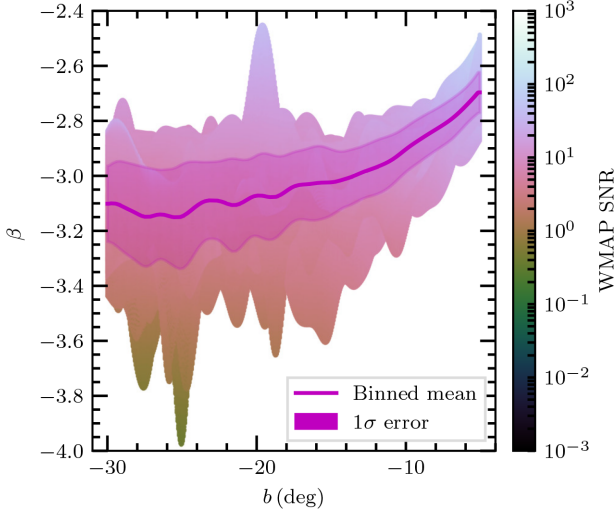


Figure 6. Spectral index for each point in the region of GSH 006 – 15 + 7 as a function of Galactic latitude. The longitude range used is $45 \text{ deg} < l < 315 \text{ deg}$. The points are coloured by the SNR of WMAP Stokes I . Despite the smoothing applied to the data, the effects of point sources and regions of low intensity are still present. These effects cause a large variance in the derived spectral index. However, after we apply a box-car smooth to the data, we recover a consistent spectral index of $\beta \sim -3.1$ in the region of the shell. Additionally, we also find a clear trend of the spectral index flattening towards the Galactic plane.

indices, and compute the error from this binning (shown as the solid line with the 1σ uncertainty band in Fig. 6). Based on Fig. 6, we use the binned spectral index as a function of Galactic latitude and assume that it is constant in Galactic longitude. We use this to scale WMAP Q and U to S Band, and carry the associated errors forward. The presence of thermal contamination in the spectral index with $b \gtrsim -10 \text{ deg}$ is irrelevant for the following analyses, because the region of GSH 006 – 15 + 7 that is depolarized in S band is also at $b \gtrsim -10 \text{ deg}$ and is thus excluded. We note that using this technique requires that the CMB emission must be subtracted from CHIPASS, as it is not present in the WMAP K -band data. A T–T plot technique usually side-steps this requirement, but since we are unable to use such a technique in this case we subtract 2.7 K uniformly from the CHIPASS total intensity data.

3.3.4 Summary of our Faraday screen observables

Finally, we obtain the four observables, as required for the Faraday screen model, as follows:

$$\begin{aligned}
 \text{PI}_{\text{on}} &= \sqrt{Q_S^2 + U_S^2} \\
 \text{PI}_{\text{off}} &= \left(\frac{\nu_S}{\nu_W}\right)^\beta \sqrt{(Q_W)^2 + (U_W)^2} \\
 \text{PA}_{\text{on}} &= \frac{1}{2} \arctan\left(\frac{U_S}{Q_S}\right) \\
 \text{PA}_{\text{off}} &= \frac{1}{2} \arctan\left(\frac{U_W}{Q_W}\right)
 \end{aligned} \tag{11}$$

here Q and U refer to the Stokes Q and U maps from S-PASS (‘S’) and WMAP (‘W’), β is the spectral index binned as a function of latitude (see Fig. 6 above), and ν refers to the frequency of each survey. We note there are some regions that have $\text{PI}_{\text{on}} > \text{PI}_{\text{off}}$,

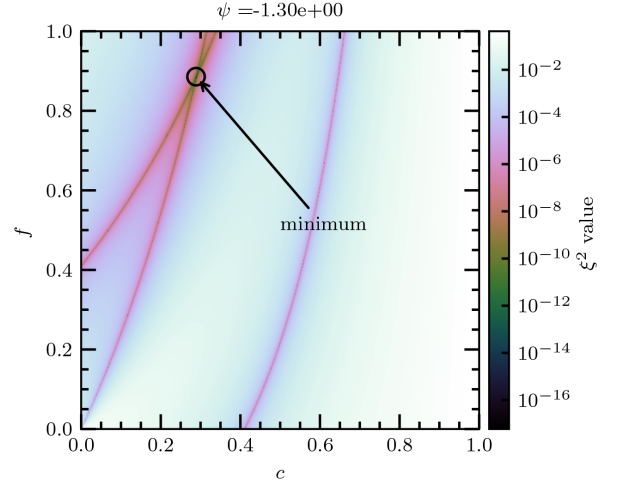


Figure 7. Slice through ξ^2 cube for a random line of sight ($l, b = [0.0 \text{ deg}, -24.96 \text{ deg}]$) at mean input value. The cube shown is a 500^3 grid, whose $[x, y, z]$ coordinates correspond to the parameters $[f, c, \psi]$. The cube is sliced through the ψ coordinate at the minimum value of ξ^2 . Note that in our final analysis, we use a 50^3 grid, which still provides robust results.

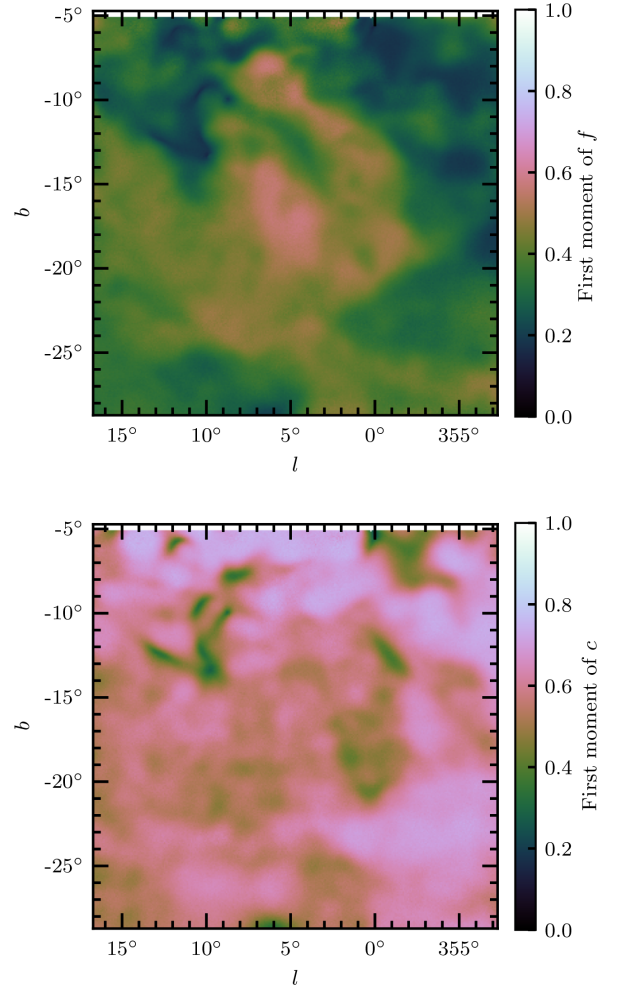


Figure 8. First moment map of model parameters in the region of GSH 006 – 15 + 7: the depolarization factor, f (upper panel), and the fraction of foreground polarization, c (lower panel).

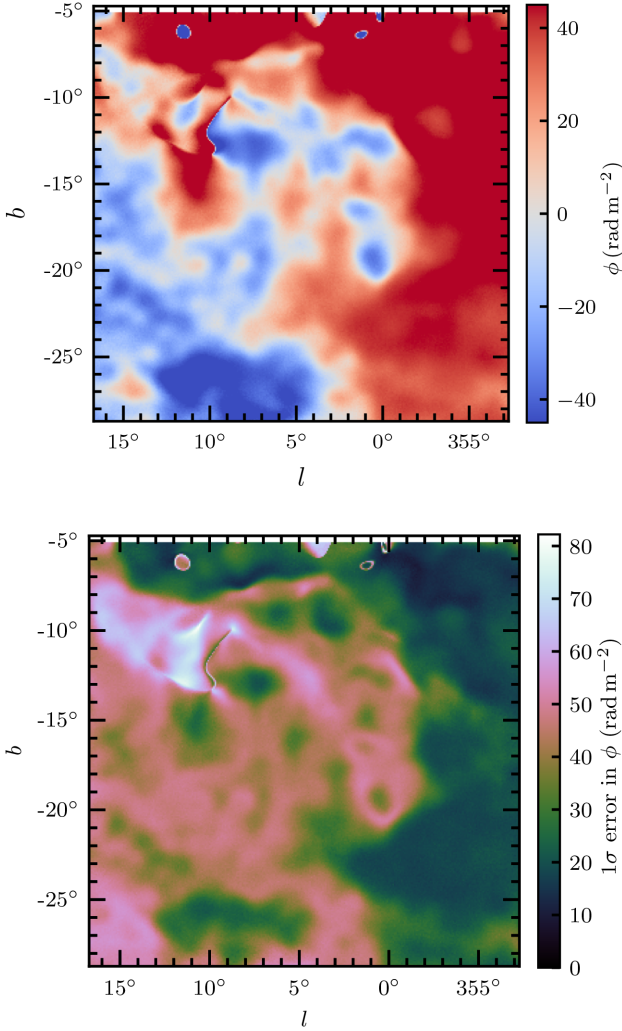


Figure 9. First moment map of Faraday depth in the region of GSH 006 – 15 + 7 (upper panel) and 1σ error associated with this Faraday depth from the model fit (lower panel).

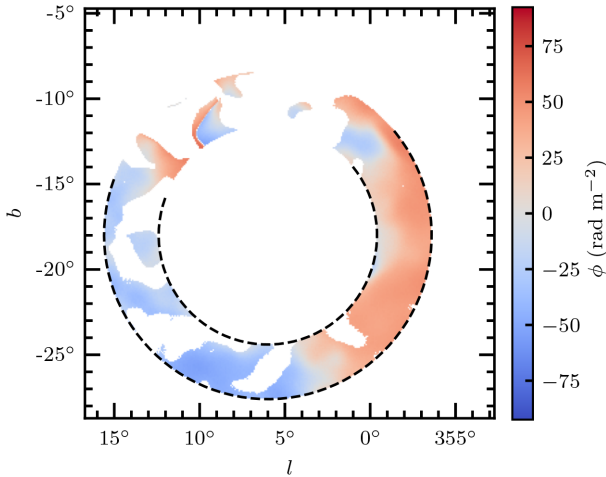


Figure 10. First moment map of Faraday thickness of GSH 006 – 15 + 7. A mask was applied to remove regions where the Faraday screen model does not apply.

which cannot be accommodated by the model. These regions will be excluded from later analysis.

3.4 Best-fit procedure

From the Faraday screen model (equation 6), we have four observables and three model parameters. We obtain the observables from the data using equation (11). To obtain the model parameters, we need to fit the Faraday screen model to the data. In order to determine the ‘best fit’ of the Faraday screen model, we have developed a ‘brute-force’ method that determines which values of the parameters f , c , and ψ produce the closest fit to the input data. To this end, we implement a grid search technique of the model hyperspace. Specifically, we sample the model for a large number of possible parameter values and find the combination of the parameter values that provide the closest fit to the data. For the purposes of our analysis, each parameter is allowed to take values in the ranges $f \in [0, 1]$, $c \in [0, 1]$, and $\psi \in [-\frac{\pi}{2}, +\frac{\pi}{2}]$. We refer to the number of samples of parameters f , c , and ψ as N_f , N_c , and N_ψ , respectively. For simplicity, we use the same number of samples for each parameter; namely $N_f = N_c = N_\psi = N_{\text{samp}}$. This is possible because the final results are insensitive to the number of samples per parameter, provided a large enough value is used. We then evaluate the Faraday screen model (equation 6) for all of these parameter values. This results in two hypersurface cubes of size N_{samp}^3 ; one containing all the possible model (mod) PI ratio ($\text{PIR} = \text{PI}_{\text{on}}/\text{PI}_{\text{off}}$) values and the other the model PA difference ($\text{PAD} = \text{PA}_{\text{on}} - \text{PA}_{\text{off}}$) values. For a given set of observed (obs) PIR and PAD, we find the following ‘ ξ^2 ’ quantity:

$$\xi^2 = [\text{PIR}_{\text{obs}} - \text{PIR}_{\text{mod}}]^2 \times [\text{PAD}_{\text{obs}} - \text{PAD}_{\text{mod}}]^2. \quad (12)$$

This quantity is constructed similarly to χ^2 , but with a few key differences. Each component in ξ^2 is not weighted by the variance from an underlying distribution. This is because the observed values are themselves sampled from a probability distribution function (PDF) which we produce. In principle, we could compute the variance from this distribution and then find and minimize the χ^2 . In doing so, however, we would lose the ability to propagate a PDF through the modelling process. As the components of ξ^2 are not weighted by the variance, each factor still has the dimensions of the observable. That is, we have the dimensionless PIR quantity and the PAD quantity in radians. Thus, unlike χ^2 , we take the product of the two factors. Like χ^2 the smallest value within our ξ^2 cube corresponds to the values of f , c , and ψ which provide the best fit to the data. This minimum value should also correspond to $\xi^2 = 0$, provided a solution exists, within numerical precision. We check for multiple solutions of (f, c, ψ) that could produce $\xi^2 = 0$, but we only ever find a unique solution for (f, c, ψ) for which $\xi^2 = 0$. Additionally, we check for the existence of an exact solution for each input pixel, and exclude pixels for which no solution exists from further analysis. We provide an example of where we find a minimum in the ξ^2 cube for a randomly sampled line-of-sight input in Fig. 7. This example is typical of how the hypersurface appears; with the location of the minimum in the cube occurring at the intersection of two ‘troughs’ of local minima, and being numerically very close to 0.

Our method does not rely on choosing reasonable initial conditions, nor does it require iteration for convergence. Instead, once the grid is built, the best fit is immediately obtained in the (f, c, ψ) parameter grid as the location of the global minimum within the ξ^2 cube.

To obtain the error in the fit of this model, we apply a Monte Carlo approach. First, we obtain the PDF of each input parameter, assuming that the starting input data (S-PASS Stokes Q and U , WMAP Stokes Q and U , and β) have Gaussian distributed errors. For each pixel in each of the input data maps, we produce a Gaussian PDF. To compute the PDFs of the observables, as required for the Faraday screen model, we apply the arithmetic described in equation (11) to each value in each PDF. These PDFs become the priors, which we input into the best-fit algorithm; where we find a best fit for the entire prior PDF for every pixel in the region around GSH 006 – 15 + 7. We sample the PDFs as a histogram with N_{PDF} bins.

By propagating the entire prior PDF through the best-fit process, we are able to obtain the full posterior PDF of the model parameters f , c , and ψ . Thus, we can gain an estimation of the uncertainty in each parameter. Examples of these posterior PDFs are given in Fig. 11. We obtain the number of samples, N_{PDF} and N_{samp} , from a convergence test on the first moment values of random pixels in the region of GSH 006 – 15 + 7. We tested values of N_{PDF} between 100 and 10^6 , and N_{samp} between 1 and 200. We found $N_{\text{PDF}} = 3000$ and $N_{\text{samp}} = 50$ to be the minimum values that still provide robust posterior results. For the entirety of our analysis, we use these two sample values across the entire region.

This process is repeated for each pixel available in the S-PASS map in the vicinity of GSH 006 – 15 + 7; the result of which is a PDF for each model parameter (f , c , and ψ) for each pixel on the map. Additionally, we also produce the best-fit values from a χ^2 minimization for comparison with the ξ^2 results. We find that the model values from this χ^2 minimization are consistent with centroid values from the distributions of the model parameters. We present here a map of the first moment of the model parameters f and c (Fig. 8). We also present a map of the first moment of Faraday thickness, as given by equation (8), and its 1σ error map (Fig. 9).

We exclude regions in the output data where the Faraday screen model does not apply, i.e. where $\text{PIR}_{\text{obs}} > 1$, and regions outside the outer boundary of GSH 006 – 15 + 7 or inside of the inner boundary, as defined by its H I emission. The resulting map is given in Fig. 10. We also apply the same mask to the maps of f and c . Within this masked region both f and c have a roughly uniform value, with spatial standard deviations of 0.08 and 0.06, respectively. We therefore compute the mean values of f and c and their associated errors. From this, we find the average depolarization factor $f_{\text{av}} = 0.4 \pm 0.3$ and fraction of foreground polarization $c_{\text{av}} = 0.6 \pm 0.3$. The errors given here are derived from the uncertainties in the mean, and not from the spatial variation. These values imply that on average roughly half of the background emission is depolarized by GSH 006 – 15 + 7, and that the background and foreground polarized emission are approximately on parity, with the foreground being slightly dominant.

3.5 Error propagation of best-fit parameters

We use a Monte Carlo error propagation method to carry the errors forward through our best-fit procedure. As such, we are able to constrain regions within the output data where the fit is poor. Fig. 11 shows the PDFs of f , c , and ψ for two pixels sampled at $l, b = [4.96 \text{ deg}, -18.0 \text{ deg}]$ (blue area) and $l, b = [0.0 \text{ deg}, -24.96 \text{ deg}]$ (orange area). Recall f is the depolarization factor, where a value of 0 refers to total depolarization, and a value of 1 to no depolarization. The parameter c is the fraction of foreground polarization, where values of 0 implies that all the emission is from the background, and 1 implies that all the emission originated in the fore-

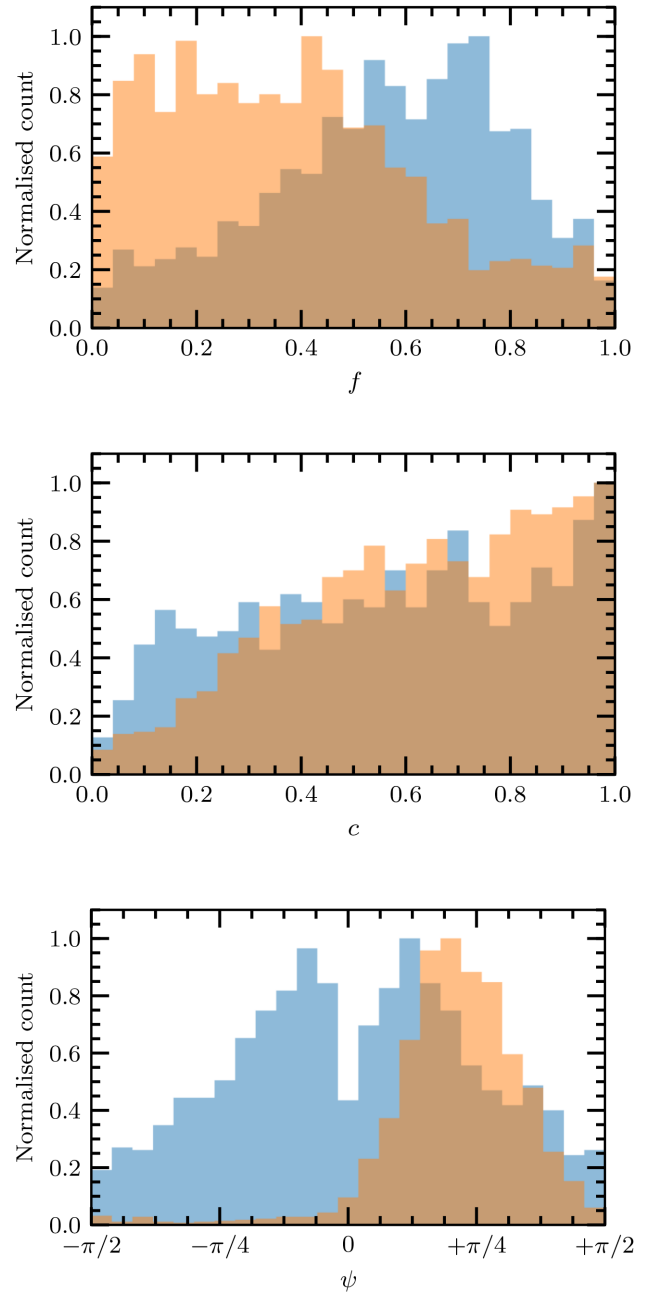


Figure 11. Posterior PDFs for two lines of sight in the f , c , and ψ data. Blue: $l, b = [4.96 \text{ deg}, -18.0 \text{ deg}]$, Orange: $l, b = [0.0 \text{ deg}, -24.96 \text{ deg}]$. Recall that f is the depolarization factor, c is the fraction of foreground polarization, and ψ is the amount of Faraday rotation through the screen.

ground. These two pixels provide typical examples of the results we see across the region of the shell. Here, we see the pixel sampled at $l, b = [0.0 \text{ deg}, -24.96 \text{ deg}]$ has a well-constrained value of ψ , whereas the pixel at $l, b = [4.96 \text{ deg}, -18.0 \text{ deg}]$ does not. The pixel at $l, b = [4.96 \text{ deg}, -18.0 \text{ deg}]$ has an input PA difference very close to 0 deg. In general, we find that the value of ψ is poorly constrained in regions where the input PA difference ($\text{PA}_{\text{on}} - \text{PA}_{\text{off}}$) is very close to 0 deg, and in regions away from the bright background polarized emission of the Fermi bubbles. The former case is a resultant behaviour of the Faraday screen model, which is better able to constrain values of ψ which are significantly greater or smaller than 0 deg. Additionally, in the lower panel of Fig. 11, the

blue distribution of the pixel at $l, b = [0.0 \text{ deg}, -24.96 \text{ deg}]$ has a dip around $\psi \approx 0$. We find similar distributions for all pixels where the input PA difference is 0 deg. This can be seen in the lower panel of Fig. 9 as high values of the 1σ error. In this figure, we note that the outline of the shell region is visible due to the increased error inside the inner boundary of the shell, where $\text{PA}_{\text{on}} - \text{PA}_{\text{off}} \approx 0$ deg. The uncertainties in f and c do follow similar trends to ψ , but overall the errors are much more uniform. Of note, the region around $l, b \approx [11 \text{ deg}, -11 \text{ deg}]$ corresponds to the highest error in ψ and unusual values in all the output parameters. The region is also associated with very low PI from the WMAP and S-PASS observations, resulting in high uncertainty in that region.

4 DISCUSSION

4.1 Line-of-sight magnetic field

The magnetic fields associated with GSH 006 – 15 + 7 can be determined from its Faraday thickness. To compute this magnetic field strength we evaluate equation (3) along the line of sight:

$$B_{\parallel} = \frac{\phi}{0.812 \langle n_e \rangle L}, \quad (13)$$

where B_{\parallel} and $\langle n_e \rangle$ are the line-of-sight averages of the magnetic field and the electron density, respectively.

The problem that equation (13) presents is an ambiguity between ϕ , $\langle n_e \rangle$, and L when trying to determine B_{\parallel} . We have already determined ϕ using the Faraday screen model, which also sets the sign of the magnetic field along the line of sight; as $\langle n_e \rangle$, and L are positive definite.

To constrain the path-length through the shell (L) we assume a simple spherical model. This model takes the path-length through the shell to be the chord between two concentric spheres. These spheres are co-centred and have an inner and outer radius such that they align with the centre, and the inner and outer bound of GSH 006 – 15 + 7, as observed in H I. This path-length is also a function of the distance to the centre of the shell. We provide a derivation for this path-length in Appendix A2.

To obtain the electron density information various different methods can be employed, each with its own advantages and disadvantages. We use the emission measure (EM) as determined by H α emission, following the procedure set out by Harvey-Smith, Madsen & Gaensler (2011) and Gao et al. (2015). Emission from H α can be observed in a diffuse manner over the entire sky; providing finely sampled spatial information on electron density. GSH 006 – 15 + 7 is not clear in H α however, which means associating emission with the shell is difficult. Additionally, as a thermal process, the H α line is very broad. This makes constraining the emission to a particular velocity challenging; especially to velocities near the local standard of rest. If instead we use velocity integrated H α observations, we would of course be including emission from both in front of and behind our region of interest. In either case, making use of H α emission will at least provide an upper limit on the electron density. Computing an emission measure also requires information in addition to H α intensity. As such, the uncertainty associated with the electron density that is derived from an emission measure will be high.

As an alternative, we did consider the electron density measured from the dispersion measure (DM) of pulsars behind GSH 006 – 15 + 7. To obtain the required DM data, we consult the ATNF Pulsar Catalogue (Manchester et al. 2005), and find 13 pulsars within a 10 deg radius of the centre of the spherical region of

GSH 006 – 15 + 7.¹ Only three of these pulsars have distance estimates independent of the dispersion measure. Of these three, one sits on the other side of the Galactic centre, meaning that most of its dispersion measure likely comes from the Galactic central region. This makes any simplifying assumptions, such as a roughly constant n_e unusable. The other two pulsars sit in front of the shell region. This means no usable electron density information was available from pulsar dispersion measures in the region of the shell. We also considered electron density information from other tracers such as S II, S III, and O II. However, we find from Shull et al. (2009) that intermediate velocity observations of ionized tracers, from both the *Hubble Space Telescope* and *Far Ultraviolet Spectroscopic Explorer*, do not probe near to the Galactic plane, and thus exclude the region around GSH 006 – 15 + 7. As each of these alternatives cannot provide us with a value of n_e in the region of GSH 006 – 15 + 7, we continue to use the EM to estimate this value.

4.1.1 Emission measure

The emission measure is defined as the path integral of n_e^2 along the line of sight:

$$\text{EM} \equiv \int_0^L n_e^2 dr. \quad (14)$$

By constraining the limits of integration to within the path-length of the shell (L), we limit the EM obtained to be from the shell only. EM can be determined from observations of H α emission as follows (Haffner, Reynolds & Tufté 1998; Valls-Gabaud 1998; Finkbeiner 2003):

$$\text{EM} = 2.75 T_4^{0.9} I_{\text{H}\alpha} \exp[2.44E(B - V)] \text{ pc cm}^{-6}. \quad (15)$$

Again, if we constrain the measured H α emission to within the shell, the resulting EM will be similarly constrained. Where T_4 is the thermal electron temperature in 10^4 K, $I_{\text{H}\alpha}$ is the H α intensity in Rayleighs, and $E(B - V)$ is the colour excess (Finkbeiner 2003). Note that to use this method, we require both an estimate of the electron temperature in the warm ionized medium (WIM) and measurements of the extinction from dust reddening. Here, we assume a typical WIM temperature of $T_4 \approx 0.8 \times 10^4$ K, following Gao et al. (2015) and Haffner et al. (1998). We obtain the colour excess from infrared dust measurements by Schlegel et al. (1998) and find a mean value within the region of the shell of ~ 0.17 .

To find the electron density associated with the supershell, we adopt the same formalism as Harvey-Smith et al. (2011). We allow the thermal electrons to be ‘clumped’ along the line of sight. Outside of a clump, we take $n_e = 0$ and within a clump $n_e = n_{e,c}$. Using this information, equation (14) can be solved for the electron density inside a clump:

$$n_{e,c} = \sqrt{\frac{\text{EM}}{f_e L}}. \quad (16)$$

Note the addition of the thermal electron filling factor f_e . This term quantifies the line-of-sight distribution of thermal electrons. If the electrons are uniformly distributed then $f_e = 1$; however, if the electrons are ‘clumped’, or if the shell has an ionized layer, then $f_e < 1$. That is, the electron clumps inhabit a column of $f_e L$ along the line of sight. From this, the average line-of-sight electron density is

¹Catalogue version: 1.57, Accessed 2017 August 23.

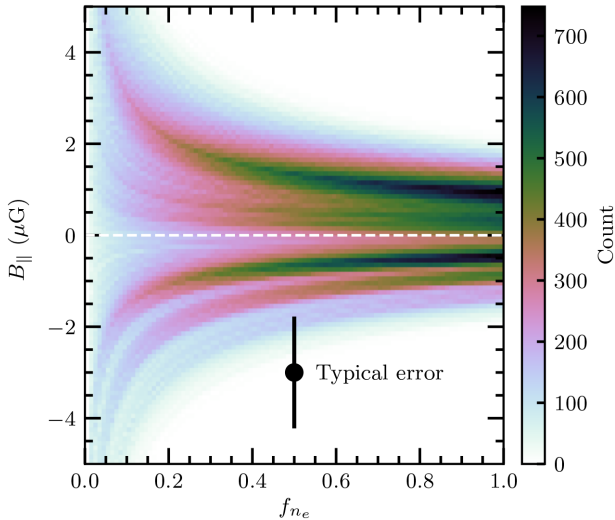


Figure 12. Distribution of the line-of-sight magnetic field ($B_{||}$) as determined by the velocity separated H α , using equation (18). We find distribution from the histogram of $B_{||}$ between $\pm 5 \mu\text{G}$, with a bin width of $0.1 \mu\text{G}$, across the region of the shell, for each value of f_{n_e} . Note that the spread of values here does not necessarily correspond to error in $B_{||}$, rather it shows the range of values found in the region in GSH 006 – 15 + 7. The ‘typical error’ shown is the mean error in $B_{||}$ when $f_{n_e} = 0.5$.

given by:

$$\langle n_e \rangle = f_e n_{e,c}. \quad (17)$$

Combining equations (13), (16), and (17) yields the line-of-sight magnetic field as:

$$B_{||} = \frac{\phi}{0.812 \sqrt{\text{EM} f_e L}}. \quad (18)$$

We use the EM as obtained from velocity separated H α data, as it provides a more accurate estimate of the electron density within GSH 006 – 15 + 7. We obtain the kinematic data from the Wisconsin H-Alpha Mapper (WHAM) survey (Haffner et al. 2003, 2010). The H α data from the WHAM kinematic survey will produce an upper limit on the electron density due to its broad line-width. We provide a map of EM in the Appendix A3. In the shell, we find a mean EM of 12.6 pc cm^{-6} .

We now compute the line-of-sight magnetic field magnitude ($B_{||}$) using equation (18), taking a distance to the shell of 1.5 kpc to evaluate the line-of-sight distance L (see Section A2). We note here that the $B_{||}$ we obtain is a lower limit, as the electron density derived from EM is an upper limit. In addition f_e is not constrained, we therefore present $B_{||}$ as a function of this factor. The spatial distribution of $B_{||}$ over the shell for each value of f_e is given in Fig. 12. We also indicate in this Figure that the mean 1σ error in $B_{||}$ is $\sim 1.2 \mu\text{G}$, when $f_e = 0.5$. As expected from equation (18), the magnitude of $B_{||}$ remains relatively constant as a function of f_e , until the filling factor becomes very small.

The value of f_e is not well constrained, and as such a value is often assumed in the literature. A value of $f_e = 1$, implying a uniform distribution of electrons, is unlikely; as are small values of f_e , since they imply large magnetic fields for a given Faraday depth measurement. We summarize some recent values of f_e from the literature in Table 2. Purcell et al. (2015) determined f_e from an MCMC model fit to their data; as such, they constrain lower limit of $f_e = 0.24$ and a mean value of $f_e = 0.3$. They note, however, that a value of around 0.5 provided a better match to dispersion measure

Table 2. A summary of recent literature values of f_e . Recall f_e is the filling factor of thermal electrons. (HVC - High velocity cloud).

f_e	Phenomena	Work
0.04 ± 0.01	Mid-plane WIM	Gaensler et al. (2008)
~ 0.3	Off-plane WIM	Gaensler et al. (2008)
0.5	HVC	McClure-Griffiths et al. (2010)
0.1	H II regions	Harvey-Smith et al. (2011)
0.4–1	W4 Superbubble	Gao et al. (2015)
≥ 0.24	Gum nebula	Purcell et al. (2015)
0.5	Magellanic Bridge	Kaczmarek et al. (2017)

data from pulsars. Kaczmarek et al. (2017) adopted their value of 0.5 following McClure-Griffiths et al. (2010). We note that some of these values are not directly derived, but rather chosen based on previous studies. Considering these values, and the range over which our derived field strength remains relatively constant, we will now adopt a value of $f_e = 0.5$ for further analysis.

4.2 Line-of-sight magnetic field structure

We provide a map of the line-of-sight magnetic field structure in Fig. 13(a) taking $f_e = 0.5$ and a distance of 1.5 kpc. The uncertainty of the values in this map is provided in Fig. 13(b). The structure presented in the map of the line-of-sight B-field appears to follow the general trend of away from the observer (negative) in the bottom left of the shell; transitioning towards the observer in the top right of the shell. This structure is, upon simple consideration, consistent with a field that is azimuthally wrapped around the surface of the shell.

Stil et al. (2009) analysed the role of magnetic fields in expanding superbubbles through MHD simulations. Their work also provided simulations of RM signatures from superbubbles expanding from the Galactic plane with a magnetic field parallel to the plane. This provides two examples to compare with our Faraday depth map (Fig. 10): an observation looking perpendicular to the Galactic magnetic field, and an observation looking parallel. In the former case, the strongest RM values are found in lines of sight through the cavity of the simulated bubble. Most notably, the sign of the RM in the shell of the bubble reverses along the line that goes through the centre of the bubble, parallel to the external field direction. In the case of observing parallel to the Galactic magnetic field the RM values are greatly amplified overall, and the wall of the bubble exhibits the strongest RM signature. Between these two scenarios, our Faraday thickness map is in better agreement with a shell which has expanded into a field perpendicular to the line of sight as we see a very clear sign change in the shell from lower left to upper right. It is likely, however, that we observe this shell at some angle between the two cases described by Stil et al. (2009). This is an area that could be further probed by simulation work. If the geometry and orientation of the shell can be well determined, and therefore the total magnetic field structure, this information could provide insight into the Galactic field into which the shell expanded.

4.3 Implications

4.3.1 Dynamical role of B-fields in GSH 006 – 15 + 7

The dynamical importance of magnetic fields in an ionized medium can be quantified by the plasma beta (β_{th}):

$$\beta_{\text{th}} = \frac{P_{\text{th}}}{P_{\text{mag}}}, \quad (19)$$

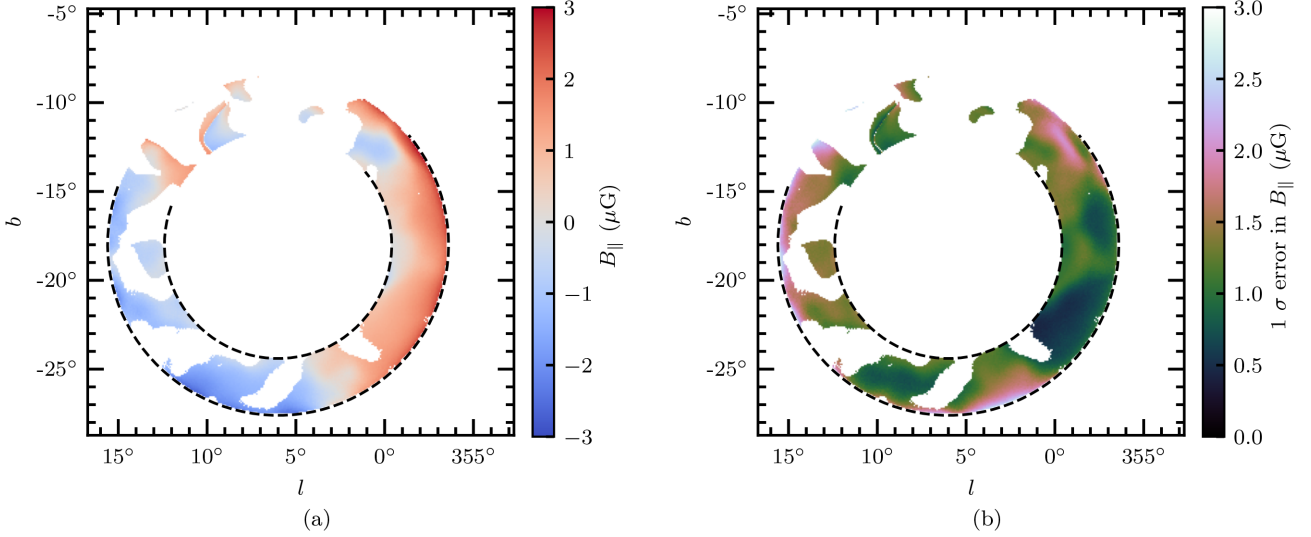


Figure 13. (a) The line-of-sight magnetic field map through GSH 006 – 15 + 7. Here, B_{\parallel} is determined from equation (18) using EM from velocity separated H α , taking $f_e = 0.5$ and a distance to the shell of 1.5 kpc. We provide a map of EM in Fig. A3. Note that some extreme values of the magnetic field occur towards the outer boundary of the shell, and arise due to the path-length L becoming very small at the limb. (b) The map of uncertainty in the line-of-sight magnetic field.

which is the ratio of the thermal pressure (P_{th}) to the magnetic pressure (P_{mag}). In this analysis, the plasma beta quantifies the dynamical role of the magnetic fields in the shell itself.

We will assume that the shell contains a mixture of a warm neutral medium (WNM) and WIM as described by Heiles & Haverkorn (2012) and that the value of B_{\parallel} remains the same in both phases. This is motivated by the observations of H I and H α in this region. We are therefore also assuming that the measured H I emission, H α emission, and Faraday depth, all arise from the same location. The thermal pressure is therefore the sum of the partial pressures of the ionized ($P_{\text{th},i}$) and neutral media ($P_{\text{th},n}$):

$$P_{\text{th}} = \langle P_{\text{th},i} \rangle + \langle P_{\text{th},n} \rangle, \quad (20)$$

where the terms in angular parentheses refer to the line-of-sight averages of those values. The partial pressures are given by:

$$\begin{aligned} \langle P_{\text{th},i} \rangle &= 2\langle n_e \rangle k T_i = 2f_e n_e k T_i \\ \langle P_{\text{th},n} \rangle &= \langle n_H \rangle k T_n \end{aligned} \quad (21)$$

where n_e is the electron number density derived from EM, n_H is the neutral hydrogen number density, k is Boltzmann's constant, and T_i and T_n are the temperatures of the ionized and neutral phase, respectively. We take values of $T_i = 8000$ K for the ionized medium and $T_n = 6000$ K for the neutral medium. We obtain the number density of neutral hydrogen from the column density derived by Moss et al. (2012) from GASS H I. They find a mean column density of $N_{H,av} \sim 2 \times 10^{20} \text{ cm}^{-2}$ in the shell of GSH 006 – 15 + 7. From this, we find the line-of-sight averaged number density of H I in the shell from:

$$\langle n_H \rangle = \frac{N_{H,av}}{L}, \quad (22)$$

where L is the path-length through the shell (see Appendix A2). Note that L varies across the projected area of the shell, therefore we also obtain a spatially varying value of n_H . In this region, we find mean values of $\langle n_e \rangle$ and $\langle n_H \rangle$ of 0.15 and 0.28 cm^{-3} , respectively. Recall, however, that our value of $\langle n_e \rangle$, as derived from EM, is an upper limit.

The magnetic pressure in the shell is given by:

$$P_{\text{mag}} = \frac{B_{\text{tot}}^2}{8\pi}, \quad (23)$$

where B_{tot} is the total local magnetic field. Our observations have provided us with the line-of-sight field, however. If we consider the case of an azimuthally wrapped magnetic field within the shell, we expect the line-of-sight field to have a distribution across the shell. That is, maximum when the total field is aligned with the line of sight, and null when the field is perpendicular. Additionally, as our model for the path-length through the shell has a hard boundary, the values of this length become very small towards the edge of the shell and thus resulting in large $|B_{\parallel}|$. Overall what we expect from the distribution of $|B_{\parallel}|$ over the shell is a smooth peak near small values of $|B_{\parallel}|$, a peak at the value which corresponds to the field being aligned with the line of sight, and a tail of more extreme values of $|B_{\parallel}|$. We find a similar distribution to this across GSH 006 – 15 + 7, as shown in the left-hand panel of Fig. 14. This distribution is bi-modal; we interpret the first peak to correspond to regions where the total field is close to perpendicular to the line of sight and the second peak to where the total field is close to parallel. To find the locations of these peaks, we fit a double Gaussian (i.e. the sum of two Gaussians) to the distribution. To obtain the uncertainty of the peak value, we perform the same analysis of the distribution of $|B_{\parallel} + \sigma_{B_{\parallel}}|$ and $|B_{\parallel} - \sigma_{B_{\parallel}}|$, as shown in the middle and right-hand panels of Fig. 14. From this, we obtain a value of the second peak of $|B_{\parallel}|_{\text{peak}} = 2.0^{+0.01}_{-0.7} \mu\text{G}$. The error range given here includes uncertainties arising from our best-fit ϕ value, $\langle n_e \rangle$ estimate, and our model L . We now assume that $B_{\text{tot}} = |B_{\parallel}|_{\text{peak}}$, since this value is likely associated with the total field being aligned with the line of sight in the case of a coherent total field. Recall that as B_{\parallel} depends on the value we found for EM, which is an upper limit, the value of B_{\parallel} and our approximation of B_{tot} are therefore lower limits. We note it would be possible for a stronger magnetic field to be obtained from our modelled Faraday depth if the electron density or f_e were demonstrated to be smaller than our current estimates. Additional data would be required, however, to motivate a different estimation.

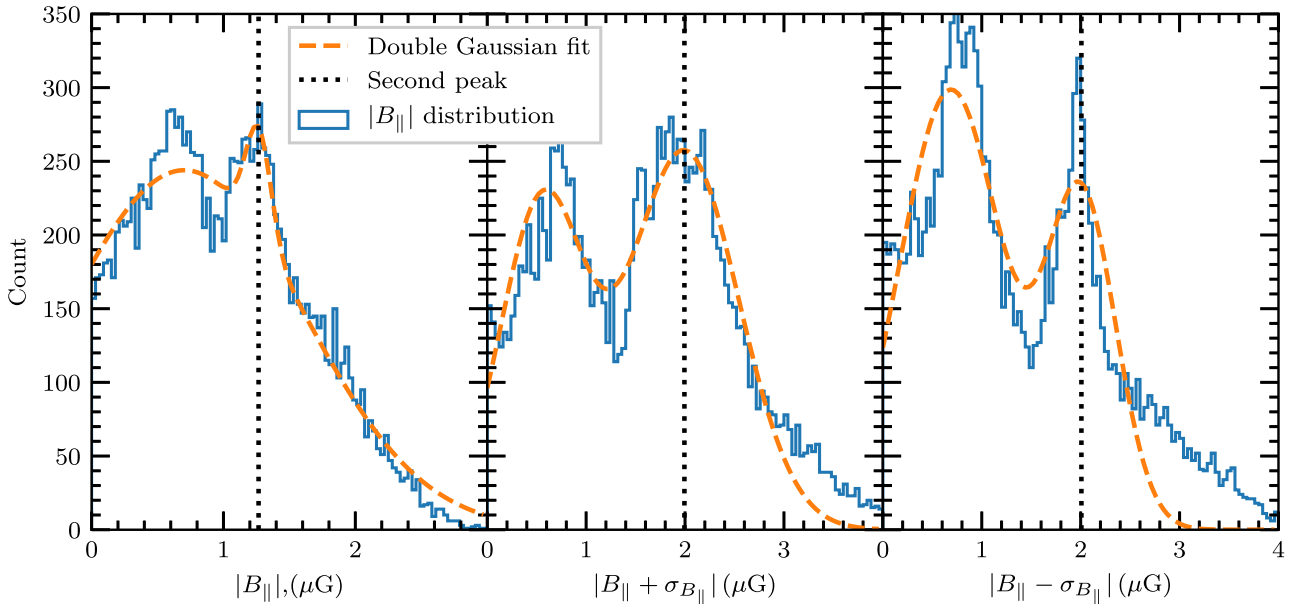


Figure 14. The distribution of the absolute value line-of-sight magnetic field ($|B_{||}$) across GSH 006 – 15 + 7. The left-hand panel shows the distribution of $|B_{||}$ taking $f_e = 0.5$ and a distance to the shell of 1.5 kpc (i.e. the distribution of values from Fig. 13a). The middle and right-hand panels show the distribution of $|B_{||}$ plus and minus the error in $B_{||}$ (as shown in Fig. 13b), respectively. Each distribution has been fitted with a double Gaussian (i.e. the sum of two Gaussians), shown in orange. The centre of the second fitted peak is shown by the black dashed line. The locations of the second peak for the left-hand, middle, and right-hand panels are 1.27, 1.99, and 2.00 μG , respectively.

We compute the mean plasma beta across GSH 006 – 15 + 7 using both the ionized and neutral partial pressures as described above. From this, we obtain a plasma beta of $\beta_{\text{th}} = 4_{-2}^{+11}$. We note that this value has a high variance, which is due to the sensitivity of β_{th} to small values of B_{tot} . Additionally, the value of β_{th} is an upper limit only, as B_{tot} was a lower limit and $\langle n_e \rangle$ was an upper limit. The errors given here include uncertainties in $\langle n_e \rangle$, $\langle n_{\text{H}} \rangle$, and B_{tot} . This value of β_{th} implies that magnetic field pressures in GSH 006 – 15 + 7 are dominated by thermal pressures in the region of the shell. Due to the large uncertainties involved it is hard to draw further conclusions regarding the dynamical role of magnetic fields in GSH 006 – 15 + 7. Such analysis would also require additional information about thermal pressures in the cavity of the shell (Ferriere et al. 1991), which is beyond the scope of this work. What is of note, however, is that despite how relatively weak the magnetic fields in GSH 006 – 15 + 7 are, this technique has allowed their detection.

4.3.2 Comparison to other results

Two structures of similar origin to GSH 006 – 15 + 7 have recently had measurements of their associated B-fields. Gao et al. (2015) analyse the magnetic fields associated with the W4 superbubble, and Purcell et al. (2015) study the Gum Nebula. Gao et al. (2015) find strong magnetic fields in association with the W4 superbubble ($B_{||} = -5.0 \mu\text{G}/\sqrt{f_e} \pm 10$ per cent, $B_{\text{tot}} > 12 \mu\text{G}$), which generally dominate the thermal pressures in the Western wall of the shell. They also find that towards the high-latitude region of W4 the magnetic fields weaken; making the magnetic and thermal pressures comparable. In the Gum Nebula, Purcell et al. (2015) find a total magnetic field strength of $B_{\text{tot}} = 3.9_{-2.2}^{+4.9} \mu\text{G}$. From this, they compute a plasma beta $\beta_{\text{th}} = 4.8$, which is relatively high, meaning that

the thermal pressures dominate the region, similar to our findings for GSH 006 – 15 + 7.

To place these values in a broader context, we compare these results to magnetic fields found in molecular and H I clouds by Crutcher et al. (2010) and in H II regions by Harvey-Smith et al. (2011). In Fig. 15, we add this work on GSH 006 – 15 + 7, as well as the W4 and Gum Nebula results, to the comparison of density against magnetic field strength. We find the H-nuclei number density ($n_{\text{H-nuclei}}$) in GSH 006 – 15 + 7 using $n_{\text{H-nuclei}} = \langle n_{\text{H}} \rangle + \langle n_{\text{p}} \rangle = \langle n_{\text{H}} \rangle + f_e n_e$, where n_{p} is the number density of protons in the ionized phase. We note that the largest objects appear in a cluster together in the lowest density region of this Figure. The H II region that appears along with GSH 006 – 15 + 7, W4, and the Gum Nebula, is Sivan 3; which is the largest H II region analysed by Harvey-Smith et al. (2011). We find that GSH 006 – 15 + 7 has comparable magnetic fields amongst these objects, but is slightly weaker and occupies a lower density regime. The resulting plasma beta for this object is therefore indicative that thermal pressures dominate the shell.

Finally, we compare the Faraday thickness derived from the Faraday screen model with RMs from extragalactic sources. If GSH 006 – 15 + 7 is the dominant Faraday component along the line of sight, we expect to see evidence of it in these data. We obtain the all-sky map of extragalactic RMs from Oppermann et al. (2015). We find a correlation between these extragalactic RMs and the Faraday thicknesses from the screen model, as shown in Fig. 16. There is significant scatter present in this correlation; which is as expected as extragalactic RMs probe the entire line of sight through the Galaxy, and thus multiple Faraday components. We fit a linear model to these data, and find that $\text{RM} \sim 0.6\phi - 15.5$, with a coefficient of determination $R^2 = 0.4$. The physical reason for the slope correlation is not obvious, as a factor of 1/2 is usually expected for regions of mixed emission and rotation (e.g. Sokoloff et al. 1998).

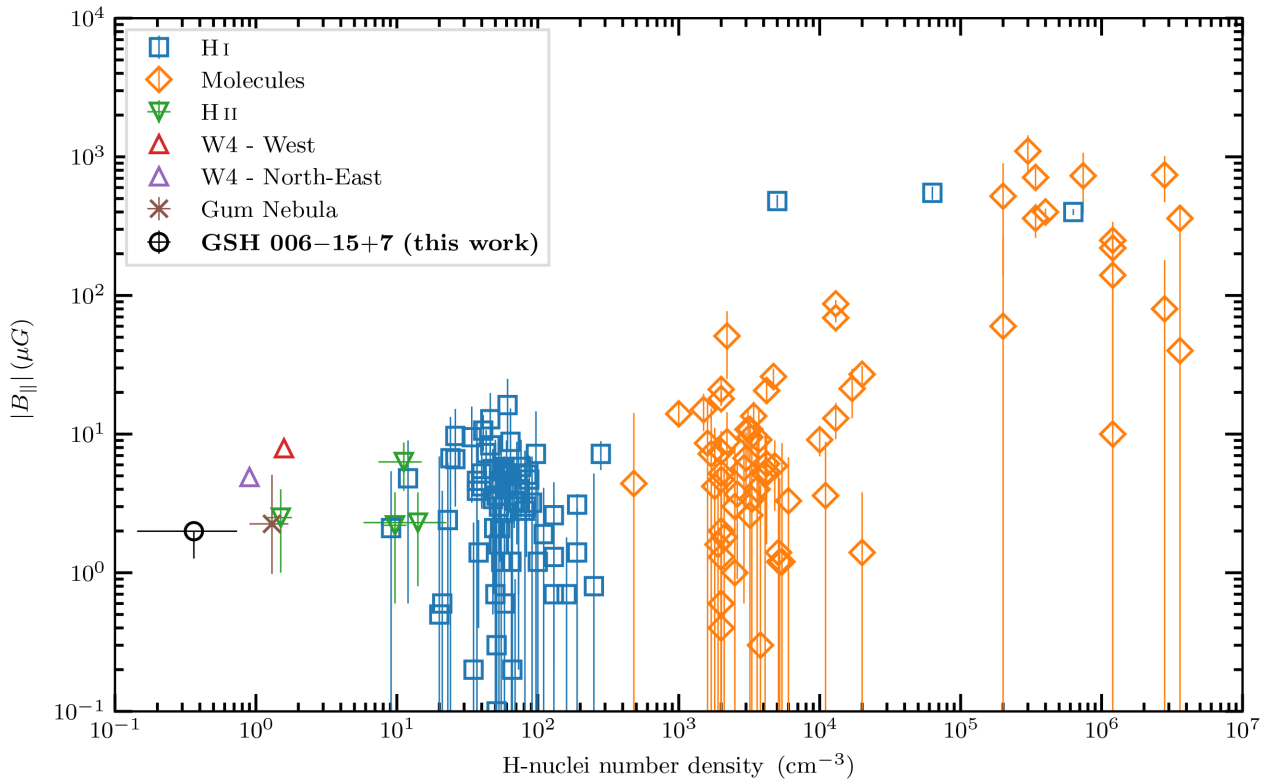


Figure 15. Line-of-sight magnetic field amplitude against number density in the associated ISM. The H I and molecular data were originally compared by Crutcher et al. (2010), and the H II data were added by Harvey-Smith et al. (2011). We make the addition of GSH 006 – 15 + 7 (this work) using the RMS value of the field, as well as the W4 superbubble (Gao et al. 2015) and the Gum Nebula (Purcell et al. 2015). Note that we have computed the line-of-sight B-field strength in the Gum Nebula from the total field strength by $B_{\parallel} \sim B_{\text{tot}}/\sqrt{3}$.

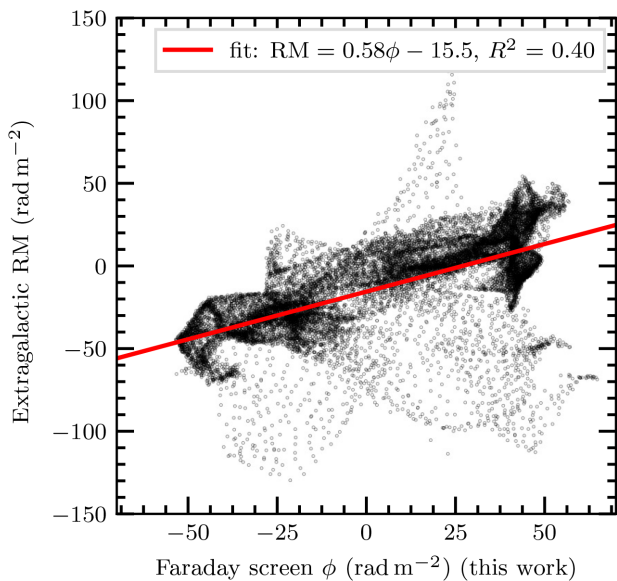


Figure 16. The comparison of extragalactic RMs from Oppermann et al. (2012) with the Faraday depth found from the Faraday screen model in the region of GSH 006 – 15 + 7. The solid red line corresponds to a least-squares linear fit to the data.

5 SUMMARY AND CONCLUSIONS

We found the polarized signature at 2.3 GHz of Galactic supershell GSH 006 – 15 + 7 in S-PASS. The morphological correlation indicates a direct detection of the MIM in association with this Galactic supershell. The ‘shadow’ of GSH 006 – 15 + 7 is most obvious in the Stokes U images. While there are signatures of the shell in PI and in Stokes Q , they are not as obvious. This highlights the importance of investigating multiple polarization modes when searching for polarized features.

We have provided a method of obtaining a Faraday depth map from two single-frequency observations; modifying the approach for modelling a Faraday screen. To estimate the background synchrotron emission of this object we used high-frequency polarization observations from WMAP K-band scaled to S band. This has allowed us to obtain Faraday rotation information with fewer assumptions than just using single-frequency observations. This method can be applied wherever a bright polarized background source illuminates purely Faraday-rotating foreground object. The source itself can be extended and complex in structure; so long as the Faraday rotation in the screen remains between $-90 \text{ deg} < \psi < +90 \text{ deg}$, the Faraday depth can be successfully recovered. This condition can be verified through inspection of a single-frequency polarization-angle map. In Section 3.4, we describe our best-fit procedure, which we find to be robust and parallelisable on multicore computers.

From this Faraday depth map, we determine the line-of-sight magnetic field and structure in association with the Galactic supershell GSH 006 – 15 + 7. We derive the line-of-sight field in

Section 4.1, and discuss the structure in Section 4.2. We find a peak line-of-sight field strength of $|B_{\parallel}|_{\text{peak}} = 2.0^{+0.01}_{-0.7} \mu\text{G}$. From these results, we have gained insight into the dynamical role of the magnetic fields associated with GSH 006 – 15 + 7. In the region of the shell, we find that the magnetic pressures are likely dominated by thermal pressures. The line-of-sight magnetic field structure indicates that the Galactic magnetic field has a significant component perpendicular to the line of sight in the region of GSH 006 – 15 + 7.

We find that the line-of-sight field strength is comparable to similarly sized objects with similar densities. This indicates that by using diffuse polarization observations, we are able to probe the magnetic fields in low-density regimes of the ISM. The method we have developed has a number of advantages in comparison to observations of point-source RMs. Most relevant to this work is that our method allows the Faraday thickness of a single, extended object to be constrained, even though it lies close to the Galactic plane.

As this object was illuminated by an extended, bright polarized source in the background, we are still able to detect Faraday rotation and assume a Faraday screen interaction. Point-source RMs probe all Faraday depths along the line of sight, which becomes very complex near the Galactic plane. Secondly, the magnetic fields detected here are weak relative to most of the previous measurements from other ISM sources reported in the literature. As such, we have shown this method to be useful in the detection of weak magnetic fields in large and diffuse areas.

ACKNOWLEDGEMENTS

The authors would like to acknowledge the significant contributions of the reviewer, Katia Ferriere, whose comments have made a substantial improvement to this work.

AT acknowledges the support of the Australian Government Research Training Program (RTP) Scholarship. NM M-G acknowledges the support of the Australian Research Council through grant FT150100024. CF gratefully acknowledges funding provided by the Australian Research Council's Discovery Projects (grants DP150104329 and DP170100603) and the Australian National University (ANU) Futures Scheme. The Dunlap Institute is funded through an endowment established by the David Dunlap family and the University of Toronto. BMG acknowledges the support of the Natural Sciences and Engineering Research Council of Canada (NSERC) through grant RGPIN-2015-05948, and of the Canada Research Chairs program. The parallel data analysis presented in this work used high-performance computing resources provided by the Australian National Computational Infrastructure (grant ek9), and the Pawsey Supercomputing Centre with funding from the Australian Government and the Government of Western Australia, in the framework of the National Computational Merit Allocation Scheme and the ANU Allocation Scheme. The Parkes Radio Telescope is part of the Australia Telescope National Facility, which is funded by the Commonwealth of Australia for operation as a National Facility managed by CSIRO. This work has been carried out in the framework of the S-PASS collaboration.

The authors acknowledge the use of the following archival data:

- The Legacy Archive for Microwave Background Data Analysis (LAMBDA), part of the High Energy Astrophysics Science Archive Centre (HEASARC). HEASARC/LAMBDA is a service of the Astrophysics Science Division at the NASA Goddard Space Flight Centre.
- The Southern H-Alpha Sky Survey Atlas (SHASSA), which is supported by the National Science Foundation.

- The Wisconsin H α Mapper and its H α Sky Survey, funded primarily by the National Science Foundation. The facility was designed and built with the help of the University of Wisconsin Graduate School, Physical Sciences Lab, and Space Astronomy Lab. NOAO staff at Kitt Peak and Cerro Tololo provided on-site support for its remote operation.

REFERENCES

- Ackermann M. et al., 2014, *ApJ*, 793, 64
 Ben Bekhti N. et al., 2016, *A&A*, 594, A116
 Bennett C. L. et al., 2013, *ApJS*, 208, 20
 Boumis P., Dickinson C., Meaburn J., Goudis C. D., Christopoulou P. E., Lopez J. A., Bryce M., Redman M. P., 2001, *MNRAS*, 320, 61
 Brentjens M. A., de Bruyn A. G., 2005, *A&A*, 441, 1217
 Burn B. J., 1966, *MNRAS*, 133, 67
 Calabretta M. R., Staveley-Smith L., Barnes D. G., 2014, *PASA*, 31
 Carretti E. et al., 2013, *Nature*, 493
 Crutcher R. M., Wandelt B., Heiles C., Falgarone E., Troland T. H., 2010, *ApJ*, 725, 466
 Duncan A. R., Haynes R. F., Jones K. L., Stewart R. T., 1997, *MNRAS*, 291, 279
 Ehlerová S., Palouš J., 2005, *A&A*, 437, 101
 Ehlerová S., Palouš J., 2013, *A&A*, 550, A23
 Ferrière K. M., 2001, *Rev. Mod. Phys.*, 73, 1031
 Ferriere K. M., Mac Low M.-M., Zweibel E. G., 1991, *ApJ*, 375, 239
 Finkbeiner D. P., 2003, *ApJS*, 146, 407
 Gaensler B. M., Madsen G. J., Chatterjee S., Mao S. A., 2008, *PASA*, 25, 184
 Gao X. Y. et al., 2010, *A&A*, 515, A64
 Gao X. Y., Reich W., Reich P., Han J. L., Kothes R., 2015, *A&A*, 578, A24
 Haffner L. M., Reynolds R. J., Tufté S. L., 1998, *ApJ*, 501, L83
 Haffner L. M., Reynolds R. J., Tufté S. L., Madsen G. J., Jaehnig K. P., Percival J. W., 2003, *ApJS*, 149, 405
 Haffner L. M. et al., 2010, in Kothes R., Landecker T. L., Willis A. G., eds, ASP Conf. Ser. Vol. 438. The Dynamic Interstellar Medium: A Celebration of the Canadian Galactic Plane Survey. Astron. Soc. Pac., San Francisco, p. 388
 Han J., 2017, *ARA&A*, 55, 111
 Harvey-Smith L., Madsen G. J., Gaensler B. M., 2011, *ApJ*, 736, 83
 Heiles C., 1979, *ApJ*, 229, 533
 Heiles C., 1984, *ApJs*, 55, 585
 Heiles C., Haverkorn M., 2012, *Space Sci. Rev.*, 166, 293
 Heiles C., Haffner L. M., Reynolds R. J., 1999, in Taylor A. R., Landecker T. L., Joncas G., eds, ASP Conf. Ser. Vol. 168, *New Perspectives on the Interstellar Medium*. Astron. Soc. Pac., San Francisco, p. 211
 Hu E. M., 1981, *ApJ*, 248, 119
 Jo Y.-S., Min K.-W., Seon K.-I., Edelstein J., Han W., 2011, *ApJ*, 738, 91
 Jo Y.-S., Min K.-W., Seon K.-I., 2015, *ApJ*, 807, 68
 Kaczmarek J. F., Purcell C. R., Gaensler B. M., McClure-Griffiths N. M., Stevens J., 2017, *MNRAS*, 467, 1776
 Kalberla P. M. W., Haud U., 2015, *A&A*, 578, A78
 Kalberla P. M. W., McClure-Griffiths N. M., Kerp J., 2009, in Heald G., Serra P., eds, *Panoramic Radio Astronomy: Wide-field 1-2 GHz Research on Galaxy Evolution*, Proceedings of Science, Groningen, the Netherlands, p. 26
 Koo B.-C., Heiles C., Reach W., 1992, *ApJ*, 390, 108
 Maciejewski W., Murphy E., Lockman F., Savage B., 1996, *ApJ*, 469, 238
 Manchester R. N., Hobbs G. B., Teoh A., Hobbs M., 2005, *AJ*, 129, 1993
 McClure-Griffiths N. M., Dickey J. M., Gaensler B. M., Green A. J., Haynes R. F., Wieringa M. H., 2000, *AJ*, 119, 21
 McClure-Griffiths N. M., Dickey J. M., Gaensler B. M., Green A. J., 2001, *ApJ*, 562, 424
 McClure-Griffiths N. M., Dickey J. M., Gaensler B. M., Green A. J., 2002, *ApJ*, 578, 176

- McClure-Griffiths N. M., Ford A., Pisano D. J., Gibson B. K., Staveley-Smith L., Calabretta M. R., Dedes L., Kalberla P. M. W., 2006, *ApJ*, 638, 196
- McClure-Griffiths N. M. et al., 2009, *ApJS*, 181, 398
- McClure-Griffiths N. M., Madsen G. J., Gaensler B. M., McConnell D., Schnitzler D. H. F. M., 2010, *ApJ*, 725, 275
- Moss V. A., McClure-Griffiths N. M., Braun R., Hill A. S., Madsen G. J., 2012, *MNRAS*, 421, 3159
- Neuhäuser R., Forbrich J., 2008, in Reipurth B., ed., *ASP Monograph Publications*, Vol. 5, *Handbook of Star Forming Regions, Volume II: The Southern Sky*. ASP Monograph Publ., San Francisco, CA, p. 735
- Norman C. A., Ikeuchi S., 1989, *ApJ*, 345, 372
- Ntormousi E., Dawson J. R., Hennebelle P., Fierlinger K., 2017, *A&A*, 599, A94
- Oppermann N. et al., 2012, *A&A*, 542, A93
- Oppermann N. et al., 2015, *A&A*, 575, A118
- Pidopryhora Y., Lockman F. J., Shields J. C., 2007, *ApJ*, 656, 928
- Purcell C. R. et al., 2015, *ApJ*, 804, 22
- Reynolds R. J., Tuftes S. L., Haffner L. M., Jaehnig K., Percival J. W., 1998, *PASA*, 15, 14
- Schlegel D. J., Finkbeiner D. P., Davis M., 1998, *ApJ*, 500, 525
- Shull J., Jones J. R., Danforth C. W., Collins J. A., 2009, *ApJ*, 699, 754
- Slavin J. D., Cox D. P., 1992, *ApJ*, 392, 131
- Sokoloff D. D., Bykov a. a., Shukurov A., Berkhuijsen E. M., Beck R., Poezd a. D., 1998, *MNRAS*, 299, 189
- Stil J., Wityk N., Ouyed R., Taylor A. R., 2009, *ApJ*, 701, 330
- Su M., Slatyer T. R., Finkbeiner D. P., 2010, *ApJ*, 724, 1044
- Suad L. A., Caiafa C. F., Arnal E. M., Cichowolski S., 2014, *A&A*, 564, A116
- Sun X. H., Han J. L., Reich W., Reich P., Shi W. B., Wielebinski R., Fürst E., 2007, *A&A*, 463, 993
- Sun X. H., Reich W., Han J. L., Reich P., Wielebinski R., Wang C., Müller P., 2011, *A&A*, 527, A74
- Tomisaka K., 1990, *ApJ*, 361, L5
- Tomisaka K., 1998, *MNRAS*, 298, 797
- Turtle A. J., Pugh J. F., Kenderdine S., Pauliny-Toth I. I. K., 1962, *MNRAS*, 124, 297
- Uyaniker B., Fürst E., Reich W., Reich P., Wielebinski R., 1999, *A&AS*, 138, 31
- Valls-Gabaud D., 1998, *PASA*, 15, 111
- Wolleben M., Reich W., 2004, *A&A*, 427, 537

APPENDIX A: APPENDIX

A1 T–T Plot

There is an increase in the slope of the T–T plot above $b \approx -10$ deg, as shown in Fig. A1, which occurs above $T_{\text{CHIPASS}} \approx 5$ K. Additionally, there are multiple ‘bifurcations’ present in the scatter. As such, the slope, as fitted to this T–T plot, does not correspond to the synchrotron emission alone. Additionally, there is significant noise present below $b \approx -10$ deg in the WMAP data; resulting in a poor fit with a coefficient of determination $R^2 = 0.5$.

A2 Derivation of path-length through shell

To obtain the path length L through GSH 006 – 15 + 7, we consider a spherical shell with an inner radius of R_i , a thickness of R_s , and therefore an outer radius of $R_o = R_i + R_s$. We consider both the near and far side of the shell. We also consider the centre of the shell to be at a distance of d_c from the Sun and at a Galactic latitude of $b = b_c$ and longitude $l = l_c$. Initially, we calculate L through the line through the centre of shell at $l = l_c$, and then use axisymmetry to find L over the entire shell. Taking $l = l_c$, we obtain the following

equation:

$$R^2 = (r \cos b - d_c \cos b_c)^2 + (r \sin b - d_c \sin b_c)^2, \quad (\text{A1})$$

where $R = R_o = R_i + R_s$ for the outer boundary, and $R = R_i$ for the inner boundary. This simplifies to:

$$R^2 = r^2 - 2d_c r [\cos(b - b_c)] + d_c^2. \quad (\text{A2})$$

Now, solving for r

$$r = d_c \cos(b - b_c) \pm \sqrt{(R^2 - d_c^2 \sin^2(b - b_c))}. \quad (\text{A3})$$

The two exact solutions to this equation correspond to the near (r_-) and far (r_+) intersections with the line of sight. Such solutions only exist within the considered boundary (outer boundary if $R = R_o$ and inner boundary if $R = R_i$). So, the path-length through the shell (L)

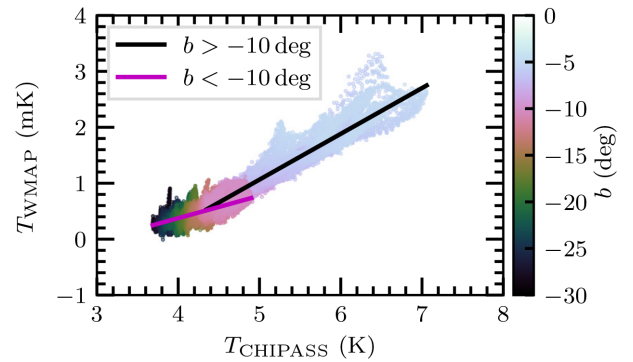


Figure A1. T–T plot: 1.4 GHz (CHIPASS) against 23 GHz (WMAP) total intensity (Stokes I) in the region of GSH 006 – 15 + 7. Linear fit to $b > -10$ deg (black): $m = 8.2 \times 10^{-4}$, $\beta = -2.6$, $R^2 = 0.9$. Linear fit to $b < -10$ deg (magenta): $m = 4.0 \times 10^{-4}$, $\beta = -2.8$, $R^2 = 0.5$. Note: point source at $l \approx [9 \text{ deg}, -19 \text{ deg}]$ is excluded. Here, R^2 is the statistical coefficient of determination. Here, we find that the slope of the T–T plot steepens towards the Galactic plane; which is due to non-thermal emission from warm gas in the disc. This emission, as well as emission from point sources, also causes ‘filaments’ in the T–T plot and increases the overall scatter.

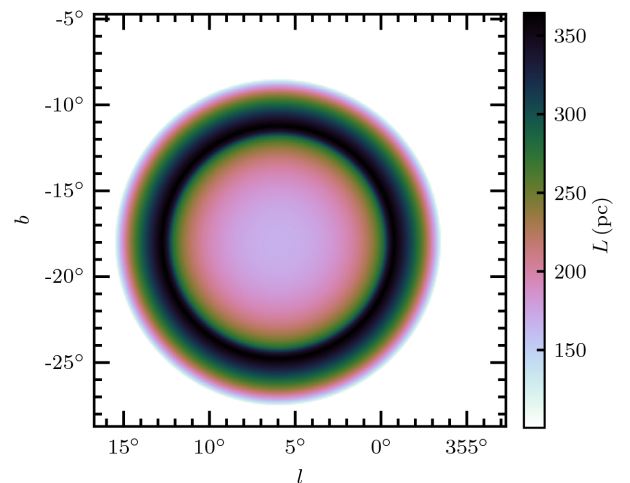


Figure A2. A map of the path-length (L) through GSH 006 – 15 + 7, as modelled by equation (A4) and (A5). Here, we adopt a distance to the shell of $d_c = 1.5$ kpc. The inner and outer radii (R_i , R_o) were chosen to match the inner outer bounds of the shell as projected on the sky. These have radii of 6.4 deg and 9.6 deg, respectively.

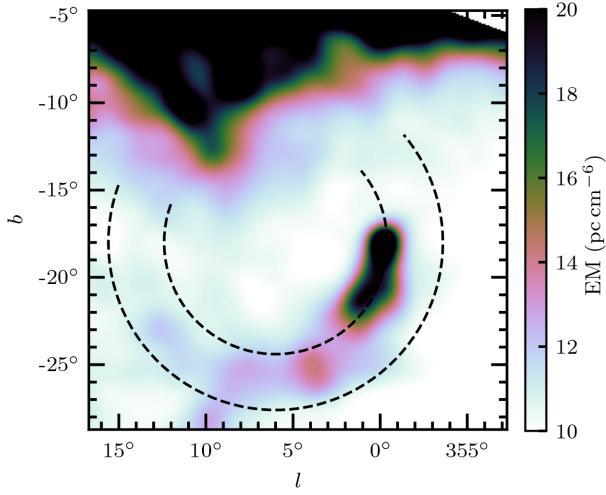


Figure A3. Map of emission measure (EM) in the region of GSH 006 – 15 + 7. The black dashed lines give the inner and outer bounds of the shell.

is the chord between these two boundaries and is given by:

$$L(l_c, b) = (r_+(R_o) - r_+(R_i)) + (r_-(R_i) - r_-(R_o)), \quad (\text{A4})$$

when the line of sight intersects both boundaries. When the line of sight intersects only the outer boundary the path length is:

$$L(l_c, b) = r_+(R_o) - r_-(R_o). \quad (\text{A5})$$

To obtain the path-length as a function of l and b ($L(l, b)$), we assume spherical symmetry. Meaning we simply apply our solution for L on the line $l = l_c$ axisymmetrically across the entire region.

A3 Emission measure map

Fig. A3 shows the map of EM as obtained using equation (15). This map has been smoothed to the spatial resolution of WMAP and then further smoothed with a 1 deg Gaussian to match the other data used.

This paper has been typeset from a $\text{\TeX}/\text{\LaTeX}$ file prepared by the author.