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Kiloparsec-scale gaseous clumps and star formation at $z=5-7$

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ABSTRACT

We investigate the morphology of the [CII] emission in a sample of “normal” star-forming galaxies at $5 < z < 7.2$ in relation to their UV (rest-frame) counterpart. We use new ALMA observations of galaxies at $z \sim 6-7$, as well as a careful re-analysis of archival ALMA data. In total 29 galaxies were analysed, 21 of which are detected in [CII]. For several of the latter the [CII] emission breaks into multiple components. Only a fraction of these [CII] components, if any, is associated with the primary UV systems, while the bulk of the [CII] emission is associated either with fainter UV components, or not associated with any UV counterpart at the current limits. By taking into account the presence of all these components, we find that the $L_{[\text{CII}]}$ -SFR relation at early epochs is fully consistent with the local relation, but it has a dispersion of 0.48 ± 0.07 dex, which is about two times larger than observed locally. We also find that the deviation from the local $L_{[\text{CII}]}$ -SFR relation has a weak anti-correlation with the $\text{EW}(\text{Ly}\alpha)$. The morphological analysis also reveals that [CII] emission is generally much more extended than the UV emission. As a consequence, these primordial galaxies are characterised by a [CII] surface brightness generally much lower than expected from the local $\Sigma_{[\text{CII}]} - \Sigma_{\text{SFR}}$ relation. These properties are likely a consequence of a combination of different effects, namely: gas metallicity, [CII] emission from obscured star-forming regions, strong variations of the ionisation parameter, and circumgalactic gas in accretion or ejected by these primeval galaxies.

Key words: galaxies: evolution - galaxies: high-redshift - galaxies: ISM

1 INTRODUCTION

The morphological investigation of galaxies in the early Universe can provide important information on their formation and evolutionary processes. For instance, a disturbed and multi-clump morphology, especially in the cold phase, may suggest the presence of disc instabilities, and can give indication of feedback processes, as well as minor and/or major merger events during the galaxy assembly (e.g., Tamburello et al. 2015; Fiacconi et al. 2016; Ceverino et al. 2017; Pallottini et al. 2017b,a). In this context, the advent of facilities delivering high angular resolution observations has enabled us to probe the internal structure of galaxies in the distant Universe, revealing that clumpy morphologies are more common at

higher redshift than at $z = 0$ (e.g. Förster Schreiber et al. 2006; Genzel et al. 2008). For instance, the fraction of galaxies at $0.5 < z < 3$ exhibiting kpc-scale clumps with sustained star-formation activity is higher than 30% (Ravindranath et al. 2006; Elmegreen et al. 2009; Guo et al. 2015). At higher redshifts the fraction of galaxies showing disturbed morphology or multi clumps is even higher. On a sample of 51 Ly α emitters (LAEs) and 16 Lyman break galaxies (LBGs) at $z > 5.7$ Jiang et al. (2013) found that roughly half of the brightest galaxies ($M_{1500} < -20.5$ mag) are made of multiple components that may be merging. Near-infrared (NIR) high-resolution imaging have also revealed that irregular shapes with multi-clump morphology are prevalent in LAEs and LBGs at $z \sim 7$

within the epoch of reionization (Ouchi et al. 2010; Sobral et al. 2015; Matthee et al. 2017; Bowler et al. 2017).

Additional identification of clumpy systems in the early epoch and a detailed characterisation of high- z clumps or satellites is fundamental to constrain galaxy assembly. The extended Atacama Large Millimetre/submillimetre Array (ALMA) configurations enable us to reach high-angular resolution and exploit far-infrared (FIR) fine structure emission lines, such as [CII] at $158\mu\text{m}$, as powerful diagnostics to assess the morphology of primeval galaxies. [CII] is emitted primarily in the (mostly neutral) atomic and molecular gas associated with Photon Dominated Regions (excited by the soft UV photons), but also in partly ionised regions and it is one of the primary coolants of the ISM. Indeed, it is generally the strongest emission lines observed in the spectra of galaxies. Since its first detection at high redshift (Maiolino et al. 2005) this transition has then been detected in large samples of distant galaxies. However, until recently, the [CII] emission was only detected in extreme environments, such as quasar host galaxies and SMGs, characterised by SFRs of several hundred solar masses per year, not really representative of the bulk of the galaxy population at these epochs (e.g. Maiolino et al. 2009, 2012; De Breuck et al. 2011; Wagg et al. 2012; Gallerani et al. 2012; Carilli & Walter 2013; Carniani et al. 2013; Williams et al. 2014; Riechers et al. 2014; Yun et al. 2015; Schreiber et al. 2017; Trakhtenbrot et al. 2017; Decarli et al. 2017). Detecting [CII] in “normal” galaxies has required the sensitivity delivered by ALMA. To date, the [CII] line has been detected in several galaxies at $z > 5$ and it is spatially resolved in most of these targets (Capak et al. 2015; Willott et al. 2015b; Maiolino et al. 2015; Knudsen et al. 2016; Pentericci et al. 2016; Bradač et al. 2016; Smit et al. 2017; Carniani et al. 2017a; Matthee et al. 2017). Smit et al. (2017) recently presented [CII] observations of two galaxies at $z \sim 7$ characterised by a gradient of velocity consistent with an undisturbed rotating gas disk. However, most of the $z > 5$ galaxies show extended and clumpy [CII] emission with velocities consistent with the systematic redshift of the galaxy ($|\Delta v_{\text{Ly}\alpha}| < 500 \text{ km/s}$) but spatially offset relative to the rest-frame UV counterpart (Maiolino et al. 2015; Willott et al. 2015b; Capak et al. 2015; Carniani et al. 2017a,b; Jones et al. 2017b). In many cases these offsets have been ignored or ascribed to astrometric uncertainties. However, based on detailed astrometric analysis, it has been shown that most of these offsets are physical (a revised analysis will be given in this paper), hence they should be taken as an important signature of the evolutionary processes in the early phases of galaxy formation. Various scenarios have been proposed to explain the positional offsets between [CII] and star-forming regions such as stellar feedback clearing part of the ISM, gas accretion, wet mergers, dust obscuration and variations of the ionisation parameter (e.g. Vallini et al. 2015; Katz et al. 2017). Barišić et al. (2017) and Faisst et al. (2017) have recently found (rest-frame) UV faint companions whose locations is consistent with the displaced [CII] emission, suggesting that the carbon line traces star-forming regions where the UV light is absorbed by dust.

As mentioned, most previous studies have attempted to assess the nature of [CII] emission in primeval galaxies neglecting the positional offsets between the FIR line and rest-frame UV emission. The goal of this paper is to assess the connection between [CII] and SFR in the early Universe by taking into account the multi-clump morphology of galaxies at $z > 5$ and by associating the components with their proper optical-UV counterparts (if detected). This is achieved by re-analysing ALMA [CII] observations of $z > 5$ star-forming galaxies, and by performing a detailed kinematical analysis of the [CII] line, in order to deblend the different

components of the multi-clump systems. In addition to previous ALMA observations, partly discussed in literature, we also make use of new ALMA data targeting five $z \sim 6$ star-forming galaxies with $\text{SFR} < 20 M_{\odot} \text{ yr}^{-1}$. In Section 2 we detail the sample and the analysis of ALMA observations. The morphological analysis is presented in Section 3, while the relation between the [CII] and SFR is discussed in Section 4. In Section 5 we investigate the connection between [CII] luminosity and Ly α strength in our sample. Section 6 focuses on the spatial extension of the [CII] and UV emission and the correlation between [CII] surface brightness and SFR surface density. We discuss the findings in Section 7, while the conclusions of this work are reported in Section 8

Throughout this paper we assume the following cosmological parameters: $H_0 = 67.8 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.308$, $\Omega_{\Lambda} = 0.685$ (Planck Collaboration et al. 2016)

2 SAMPLE, OBSERVATIONS AND ANALYSIS

2.1 Archival data

The sample is mainly drawn from the archive and literature by selecting only spectroscopically confirmed star-forming galaxies at $z > 5$ observed with ALMA in the [CII] line. We limit our sample to those systems with $\text{SFR} \lesssim 100 M_{\odot} \text{ yr}^{-1}$ since they are representative of the bulk of the galaxy population in the early Universe (e.g. Robertson et al. 2015; Carniani et al. 2015). The list of selected sources is given in Table 1. The sample does not include lensed systems (Knudsen et al. 2016; González-López et al. 2014; Bradač et al. 2016; Schaerer et al. 2015) since magnification factor uncertainties may lead to large errors on SFR and [CII] luminosity estimates, as well as on the morphology analysis.

For the purpose of our investigation, which focuses on the nature and implications of the positional offsets between [CII] and UV emission, we have retrieved and re-analysed ALMA data revealing a [CII] detection at the systemic velocity of the galaxy (see Table 1). For these objects, ALMA observations have been calibrated following the prescriptions presented in previous works.

In addition to the rest-frame FIR images, we have also used Hubble Space Telescope (*HST*) and Visible and Infrared Survey Telescope for Astronomy (*VISTA*) NIR observations (rest-frame UV at $z > 5$). ALMA, *HST* and *VISTA* data have been aligned based on the location of serendipitous sources detected in both ALMA continuum and NIR images by assuming that the millimetre emission of these sources is cospatial to the near-infrared map. This is also supported by the fact that all foreground sources used for registering millimetre and NIR images do not exhibit any multi-clump or merger-like morphologies indicating that astrometric offsets between the ALMA and NIR images are likely associated to astrometric calibrations. For those observations revealing the presence of two (or more) serendipitous sources we have verified that the astrometric shift for each source is consistent with that estimated from the other source(s) in the same map. In all cases we have checked that the estimated astrometric offset is consistent with those obtained by aligning NIR foreground sources and ALMA phase calibrators to their astrometric position from the GAIA Data Release 1 catalogue (Gaia Collaboration et al. 2016). For those sources whose ALMA continuum map showing no serendipitous sources, we have matched the NIR foreground sources and ALMA phase calibrators to either GAIA Data Release 1 catalogue (Gaia Collaboration et al. 2016) or AllWISE catalogue (Cutri & et al. 2013). We note that in all ALMA datasets the locations of the various phase calibrators are in agreement within the error with the

Table 1. Overview of the $z > 5$ star-forming galaxies observed with ALMA used in this paper, ordered by name.

Target ^(a)	Ra ^(b)	Dec ^(c)	Ref. ^(d)	[CII] ^(e)	Clumpy ^(f)
<i>Literature Sample</i>					
BDF3299	337.0511	-35.1665	1,2	✓	✓
BDF512	336.9444	-35.1188	1		
CLM1	37.0124	-4.2717	3	✓	
COSMOS13679	150.0990	2.3436	4	✓	
COSMOS24108	150.1972	2.4786	4	✓	✓
COS-2987030247	150.1245	2.2173	5	✓	
COS-3018555981	150.1245	2.2666	5	✓	
CR7	150.2417	1.8042	6	✓	✓
Himiko	34.4898	-5.1458	7,12	✓	✓
HZ8	150.0168	2.6266	8	✓	✓
HZ7	149.8769	2.1341	8	✓	
HZ6	150.0896	2.5864	8	✓	✓
HZ4	149.6188	2.0518	8	✓	
HZ3	150.0392	2.3371	8	✓	
HZ9	149.9654	2.3783	8	✓	
HZ10	150.2470	1.5554	8	✓	✓
HZ2	150.5170	1.9289	8	✓	✓
HZ1	149.9718	2.1181	8	✓	
IOK-1	200.9492	27.4155	9		
NTTDF6345	181.4039	-7.7561	4	✓	
SDF46975	200.9292	27.3414	1		
SXDF-NB1006-2	34.7357	-5.3330	10		
UDS16291	34.3561	-5.1856	4	✓	
WMH5	36.6126	-4.8773	3,11	✓	✓
<i>Additional new data</i>					
BDF2203	336.958	-35.1472		✓	
GOODS3203	53.0928	-27.8826			
COSMOS20521	150.1396	2.4269			
NTTDF2313	181.3804	-7.6935			
UDS4812	34.4768	-5.2472			

Notes. (a) Name of the source. (b, c) J2000 coordinates. (d) References in which ALMA observations are presented ([1] Maiolino et al. 2015; [2] Carniani et al. 2017a; [3] Willott et al. 2015b; [4] Pentericci et al. 2016; [5] Smit et al. 2017; [6] Matthee et al. 2017; [7] Ouchi et al. 2013; [8] Capak et al. 2015; [9] Ota et al. 2014; [10] Ouchi et al. 2013; [11] Jones et al. 2017b [12] Carniani et al. 2017b) (e) Check mark, ✓, indicates that [CII] emission has been detected at the redshift of the galaxy. (f) Check mark, ✓, indicates that the galaxy has a clumpy morphology.

GAIA and AllWISE catalogues, implying that, when a systemic (not physical) offset is seen, this is generally due to some small astrometric uncertainties in the optical-NIR data. These astrometric issues have been also discussed by Dunlop et al. (2016) who analysed ALMA images targeting the Hubble Ultra Deep Field. We therefore applied the astrometric shifts, which span a range between $0.1''$ and $0.25''$, to the NIR images.

We note that additional systems with low SFRs have also been tentatively detected in 14 [CII] line emitting candidates at $6 < z < 8$ (Aravena et al. 2016), which are not included in this analysis as they are not spectroscopically confirmed yet and $\sim 60\%$ of these objects are expected to be spurious.

2.2 Additional new ALMA data

In addition to the archival/literature sample, we have also included new [CII] observations of five star-forming galaxies at $z \sim 6$ with a

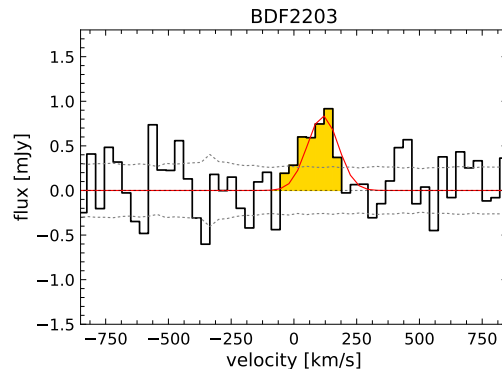


Figure 1. ALMA spectrum of BDF2203 showing a new [CII] detection. The velocity reference is set to the redshift defined by the Ly α . The dotted grey lines shows the 1σ and -1σ per channel. The red line indicates the best fit 1D Gaussian line profile.

SFR $\sim 10 M_{\odot} \text{ yr}^{-1}$ observed with ALMA in Cycles 3 and 4 (P.I. Pentericci). The five new sources, listed in Table 1, have been selected from a sample of > 120 LBGs at $z \sim 6 - 7$. These have been spectroscopically confirmed thanks to recent ultra-deep spectroscopic observations from the ESO Large Program CANDELSz7 (P.I. L. Pentericci; De Barros et al. 2017). The proposed ALMA programs aimed at observing [CII] emission in ten star-forming galaxies at $z > 6$ with UV luminosities lower than -21 mag (UV SFR $< 20 M_{\odot} \text{ yr}^{-1}$) and spectroscopic redshift uncertainties < 0.03 , but only five galaxies have been observed. The observations and data calibrations are presented in the Appendix A. We have registered NIR images to ALMA observations by matching the location of the foreground, serendipitous continuum sources and ALMA calibrators to the position given by the GAIA Data Release 1 catalogue (Gaia Collaboration et al. 2016).

While the continuum emission is not detected at the location of any of the five galaxies, we detect two and one serendipitous sources in the COSMOS20521 and BDF2313 continuum maps, respectively. The positions of the serendipitous continuum sources is in agreement with the location of NIR foreground galaxies, thus confirming the astrometric shifts estimated from the catalogue.

The [CII] emission is undetected in all but one of these sources, that is, BDF2203. Figure 1 shows the spectrum of the detected [CII] emission, with a spectral rebinning of 40 km s^{-1} , while the spectra of the non-detections are shown in the Appendix (Figure A1).

For the four non-detections, we assume a full-width-at-half-maximum of $\text{FWHM} = 100 \text{ km s}^{-1}$, which is consistent with [CII] line widths observed in other $z > 6$ galaxies (e.g. Pentericci et al. 2016; Carniani et al. 2017a,b), and we infer 3σ upper limits on the [CII] luminosity (Table A1).

The redshift of the [CII] emission detected in BDF2203 is $z_{[\text{CII}]} = 6.1224 \pm 0.0005$, which is in agreement with that inferred from Ly α ($z_{\text{Ly}\alpha} = 6.12 \pm 0.03$). By fitting a 1D Gaussian profile to the [CII] line we estimate a $\text{FWHM} = 150 \pm 50 \text{ km/s}$ that is similar to those estimated in high- z [CII]-emitting galaxies ($\text{FWHM} = 50-250 \text{ km/s}$; Willott et al. 2015b; Pentericci et al. 2016; Matthee et al. 2017; Carniani et al. 2017b). The flux map of the [CII] line, extracted with a spectral width of 200 km/s , is shown in Figure 2. The emission is detected in the map with a $\text{S/N} = 5$ and has an integrated flux density of $140 \pm 35 \text{ mJy km/s}$, which corresponds to $L_{[\text{CII}]} = (12.5 \pm 2.5) \times 10^7 L_{\odot}$ at $z = 6.12$. By fitting a 2D gaussian profile to the flux map, we measure a size of $(2.3 \pm 0.5)'' \times (1.2 \pm 0.2)''$ with $\text{PA} = 88 \text{ deg} \pm 8 \text{ deg}$

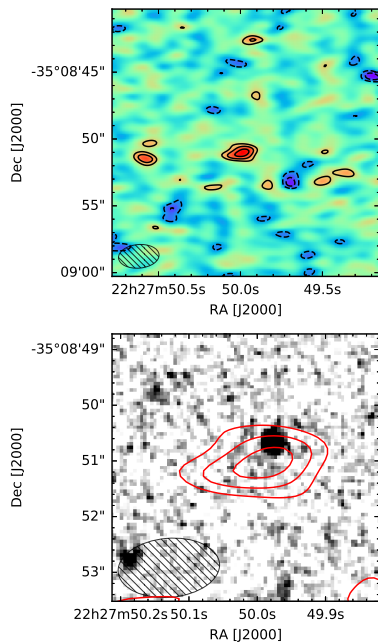


Figure 2. Top: [CII] flux map of BDF2203 obtained integrating the ALMA cube over the velocity range between -60 km/s and 132 km/s. Red solid contours are at levels of 2σ , 3σ and 4σ , where $\sigma = 23$ mJy km/s. Dashed black contours indicate negative values at -3 and -2 times the noise level in the same map. Bottom: zoom of the central $2.5''$ region around BDF2203. [CII] flux map contours are over imposed on the NIR image. Red contours are at the same levels as in the top panel. The ALMA synthesised beam in the lower left corner in both panels.

and a beam-deconvolved size of $(1.4 \pm 0.8)'' \times (0.5 \pm 0.4)''$ with a PA = $80 \text{ deg} \pm 20 \text{ deg}$. The [CII] emission is thus marginally resolved and has a diameter size of ~ 5 kpc. The [CII] flux map overlaps with the rest-frame UV emission that has an extension of about 1 kpc. The centroids are separated by $0.25''$. However, given the ALMA beam size ($1.90'' \times 1.11''$), such positional offset is consistent with the positional uncertainties $\Delta\theta \approx \frac{\langle \text{beam size} \rangle}{S/N} \sim 0.2'' - 0.4''$. The low angular resolution and sensitivity of current observations are not sufficient to assess the morphology of the FIR line emission.

2.3 SFR estimates and morphology analysis

The continuum emission at rest-frame wavelengths around $158 \mu\text{m}$ is associated to thermal emission from dust heated by the UV emission of young stellar populations in galaxies. The continuum emission is detected only in four sources within the selected sample. When compared with the typical UV-IR SED of galaxies, the weak rest-frame far-IR continuum indicates that these galaxies are on average characterised by low dust masses.

The spectral energy distribution (SED) of the thermal dust emission can be modelled with a greybody with dust temperature T_d and spectral index emissivity β . However, one single photometric measurement is not sufficient to perform the SED fitting and constrain the two free parameters. We thus assume $T_d = 30\text{K}$ and $\beta = 1.5$, which are consistent with those observed in local dwarf galaxies (Ota et al. 2014), to estimate the far-infrared (FIR) luminosity from the ALMA continuum observations. For those galaxies that are not detected in ALMA continuum images we infer a 3σ upper limit on L_{FIR} . We note that the FIR emission strongly de-

pends on the assumed dust temperature, yielding to a luminosity uncertainty of $\Delta \log(L_{\text{FIR}}) = 0.6$ (Faisst et al. 2017).

For each galaxy we can infer the star-formation rate from both the UV (SFR_{UV}) and FIR (SFR_{FIR}) emission by adopting the calibrations presented in Kennicutt & Evans (2012). Excluding HZ9 and HZ10, all sources in the sample have $\text{SFR}_{\text{FIR}}/\text{SFR}_{\text{UV}} \lesssim 1$ with an average value of $\text{SFR}_{\text{FIR}}/\text{SFR}_{\text{UV}} \approx 0.6$. Given that most of the FIR luminosity estimates are 3σ upper limits, the average $\text{SFR}_{\text{FIR}}/\text{SFR}_{\text{UV}}$ is actually much lower than 0.6 .

Given the small contribution of SFR_{FIR} (either in the IR-detected sources or those with upper limits), we assume total $\text{SFR} \approx \text{SFR}_{\text{UV}}$ with no dust correction for those galaxies without continuum detection. While we infer total $\text{SFR} = \text{SFR}_{\text{UV}} + \text{SFR}_{\text{FIR}}$ for those galaxies revealing continuum emission at $158 \mu\text{m}$.

In order to assess the multi-clump morphology of the [CII] emission in our sources, we perform a kinematical analysis on the retrieved ALMA cube, extracting channel maps at different velocities relative to the redshift of the galaxies. This analysis enables us to disentangle the emission of complex systems having multiple components at different velocities (e.g. Carniani et al. 2013; Riechers et al. 2014). We also estimate the size of the various UV and [CII] emission by fitting a 2D elliptical Gaussian profile to emission maps and by taking into account the angular resolution of the observations.

The properties of the [CII], FIR and UV emission, such as redshift, flux density, luminosity, SFR, and radius, are reported in Table 2

3 MULTI-COMPONENTS SYSTEMS

Carniani et al. (2017a) discuss the origin of the positional and spectral offsets between optical emission and FIR lines observed in high- z systems, suggesting that the observational properties can arise from distinct regions (or components) of galaxies. In order to understand the nature of these offsets, we have re-analysed ALMA [CII] observations for those galaxies showing [CII] detections and have performed a morphology analysis as discussed in Section 2.3. We have found a clear multi-clump morphology in nine out of 21 $z > 5$ galaxies having [CII] detections (Table 2). The UV rest-frame images and [CII] maps are shown in Figure 3, sorted by redshift. The stamps are $3''$ across, or ~ 17.5 kpc at the average redshift of this sample ($\langle z \rangle = 6$), and show the contours of the [CII] maps, obtained from our analysis, superimposed on the rest-frame UV emission in grayscale.

New *HST* near-IR images of HZ6, HZ8, and HZ10 have already been presented in Barišić et al. (2017) and Faisst et al. (2017) revealing multi-component structures. The location of the individual rest-frame UV clumps is consistent with the peak positions of the [CII] emission extracted at different velocities relative to the redshift of the brightest component (which is labelled “a” in all stamps). The [CII] emission detected in all individual clumps has a level of significance¹ higher than 5σ and it is spatially resolved (see Table 2). The channel map analysis confirms that HZ8b and HZ10b (Hz8W and Hz10W in previous works) are at the same redshift of HZ8a and HZ10a, respectively. We note that the kinematic properties of HZ10 are consistent with the analysis reported by Jones et al. (2017a), who claim that the velocity gradient observed in this

¹ σ is the rms of the channel map in which we detect the [CII] emission

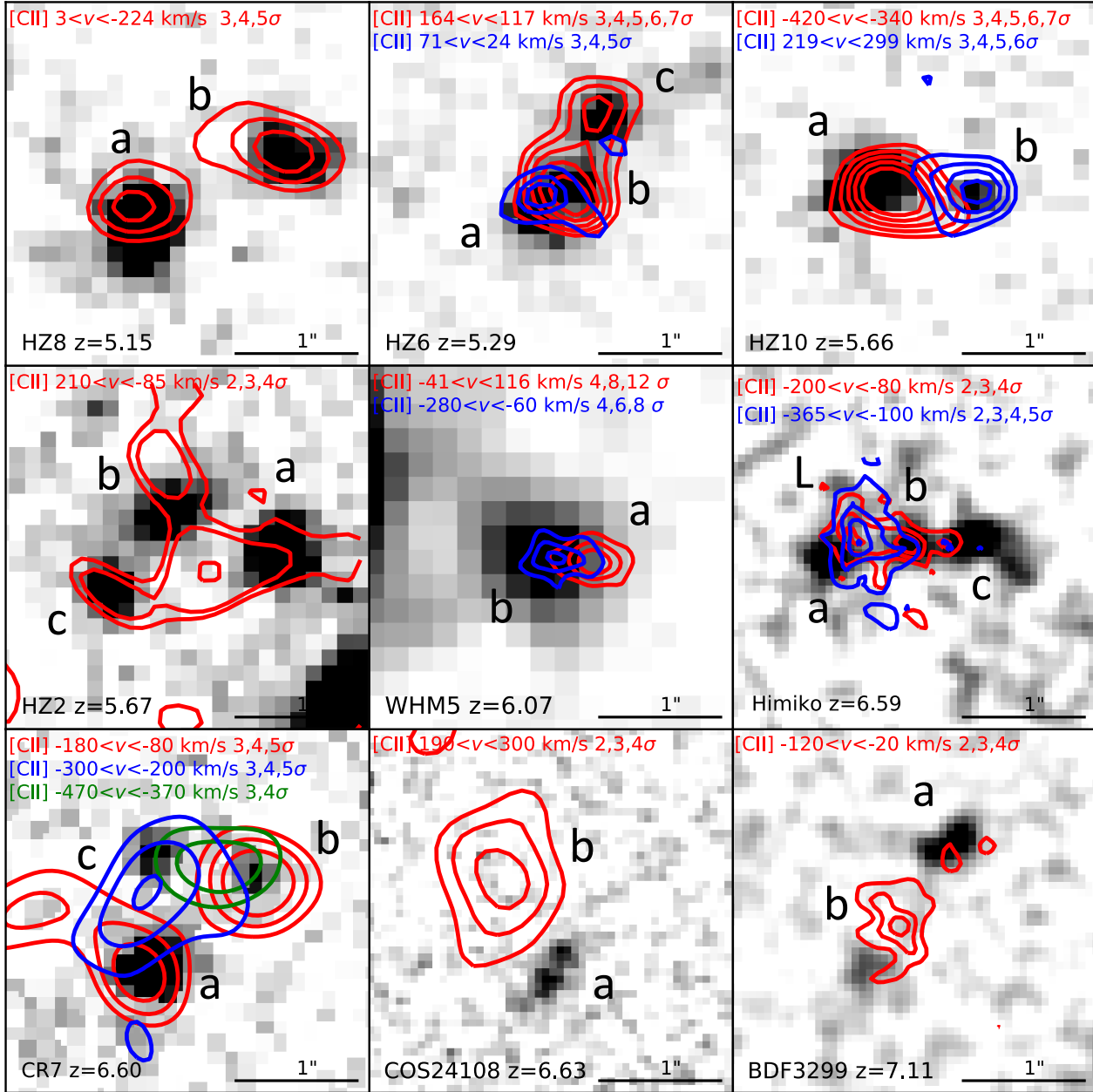


Figure 3. Rest-frame UV images of the nine $z > 5$ star-forming galaxies showing multi-clump morphology in [CII] and/or rest-frame UV emission. The galaxies are ordered by redshift and each stamp is $3''$ on the side, with North to the top and East to the left. The red, blue, and green contours show the [CII] channel maps at different velocity intervals. The levels of the contours and channel maps velocities are indicated in the legends. The properties of each galaxy are reported in Table 2.

galaxy matches a merger scenario rather than a rotating gas disk model.

The deeper *HST* observations reveal also two faint companions close to the HZ2 galaxy within a projected distance of ~ 6 kpc. These sources are not discussed in previous studies (Capak et al. 2015; Barišić et al. 2017), since their redshifts were not fully spectroscopically confirmed by the [CII] line. However, a detailed kinematic investigation of the carbon line shows an extended emission with a morphology consistent with the rest-frame UV emission. Although the two faint companions are detected with a low level of significance ($\sim 3.5\sigma$), the match between the three UV peaks and the [CII] emission supports the reliability of the ALMA detections.

The three sources, dubbed in this work as HZ2a, HZ2b and HZ2c, are at the same redshift and form a multi-component system similar to that observed in HZ6. The angular resolution and the low sensitivity of current ALMA observations are not sufficient to spatially resolve the [CII] emission in HZ2b and HZ2c. HZ2a is instead spatially resolved, with a diameter of 2.5 kpc.

An additional star-forming galaxy with a complex morphology is WHM5 at $z = 6.0695$, which appears to consist of multiple components in [CII]: a compact source, seen also in dust emission, and an extended component at the location of the rest-frame UV emission (Willott et al. 2015a; Jones et al. 2017b). The two components are separated by a projected distance of ~ 3 kpc and a

velocity of ~ 200 km/s. Both [CII] emission components are spatially resolved (~ 1.3 and ~ 2.8 kpc) in the ALMA observations with an angular resolution of $0.3''$ (Jones et al. 2017b).

At higher redshift, the presence of multiple [CII] components has recently been reported in CR7 at $z=6.6$ by Sobral et al. (2015, 2017) and Matthee et al. (2015, 2017). Three out of four detected [CII] clumps coincide with the location of UV clumps. In two cases the [CII] emission is spatially resolved with a radius of $\sim 3 - 3.8$ kpc (see Matthee et al. 2017 for details).

An additional multi-clump galaxy observed with ALMA is Himiko at $z \sim 6.695$ (Ouchi et al. 2013). The rest-frame UV image (see Figure 3 of Ouchi et al. 2013) reveals that the galaxy comprises three sub-components with SFR spanning in range between 5 and $8 M_{\odot} \text{ yr}^{-1}$. The projected distance between the sources is of about $3-7$ kpc. Out of the three sub-components have a Ly α EW=68 Å while the other two have EW less than 8 Å. Although early observations had reported non-detections of [CII] in this source, recently Carniani et al. (2017b) have reported a clear detection with extended/multi-clump morphology. The primary [CII] emission is coincident with the Ly α peak, while the UV clumps are much weaker in [CII] or even undetected. More generally the extended [CII] emission does not resembles the UV clumpy distribution.

For the remaining systems (COS24108 and BDF3299) showing a clear positional offsets between [CII] and UV emission in Figure 3, we cannot speculate much more than what has been already done in previous works due to the lack of deeper ALMA and/or *HST* observations. As discussed by Carniani et al. 2017a, these offsets are certainly associated with physically distinct sub-components, and the [CII] clumps with no UV counterpart may either be tracing star forming regions that are heavily obscured at UV wavelengths (Katz et al. 2017), or associated with accreting/ejected gas.

3.1 On the nature of the kpc-scale sub-components

The deeper [CII] and rest-frame UV observations have unveiled the real multi-component nature of nine star-forming galaxies at $z > 5$ further highlighting and (partly) explaining positional offsets between UV and [CII] emission in previous studies (Capak et al. 2015; Maiolino et al. 2015; Willott et al. 2015b; Pentericci et al. 2016; Carniani et al. 2017a). These results suggest that future evidence of displaced FIR line and UV emission should be not ignored since it generally reveals the presence of sub-components with different physical properties.

Such sub-components can be ascribed to either satellites in the process of accreting (Pallottini et al. 2017a) or clumps ejected by past galactic outflows (Gallerani et al. 2016). However, in many cases both the SFR_{UV} and the size of the various sub-components are comparable to those estimated for the central galaxies, hence suggesting a major merger scenario in many cases. Future ALMA observations with higher resolution and sensitivity are necessary to detect further sub-components in these and other systems and to perform a detailed dynamical analysis, which allow us to estimate the dynamic mass and assess the nature of sub-components.

4 THE L_[CII] - SFR RELATION AT Z=5-7

A tight relation between the [CII] luminosity and the global SFR is seen in local galaxy observations, at least when excluding extreme (ULIRG-like) cases (De Looze et al. 2014; Kapala et al. 2015; Herrera-Camus et al. 2015). This finding makes the [CII]

line a promising tool to investigate the properties of early galaxies and to trace their star formation. However, the behaviour of the [CII] line emission at $z > 5$ seems to be more complex than observed in the local Universe. Previous studies have shown that only a fraction of [CII] detections of early galaxies agree with the local relation, while most high- z galaxies are broadly scattered, with claims that most of them are [CII]-deficient relative to the local relation. However, most of previous high- z studies classified multi-component systems as single objects in the L_[CII]-SFR diagram. If we associate each clump and/or galaxy with its proper UV counterparts (or lack thereof), then the resulting location on the L_[CII]-SFR diagram changes significantly for these objects

Figure 4 shows L_[CII] as a function of SFR. The green line illustrates the local relation obtained by De Looze et al. (2014) and its dispersion is given by the shaded area. Results for $z > 5$ galaxies, as listed in Table 2, are shown with various symbols in Figure 4. The SFR estimation for the $z > 5$ galaxies (and their sub-components) is discussed in Section 2.3. The multiple-component objects (HZ2, HZ6, HZ8, HZ10, WHM5, Himiko, CR7, COS24108, and BDF3299) are split into several individual components with their own SFRs and L_[CII]. The nine complex systems discussed in Section 3 are broken into 20 sub-components distributed on different regions of the L_[CII]-SFR plane. The location of the individual subcomponents is indicated with yellow stars, while the location of these systems by integrating the whole [CII] and UV emission (i.e. ignoring that these are actually composed of different subsystems) is indicated with blue stars. The four new [CII] non-detections presented in Section 2.2 fall below the local relation, while the L_[CII] for BDF2203 places this galaxy along the De Looze et al. (2014) relation.

Once the association between [CII] emission and optical-UV counterparts is properly done, we find that the resulting distribution occupies a large area of the L_[CII]-SFR plot with a large scatter both above and below the local relation. About 19 objects of the total sample are in agreement within 1 sigma with the local relation, but the remaining 24 systems have deviations, either above or below the relation, up to 3 sigma.

In order to quantify the L_[CII]-SFR offset of the high- z sample from the relation found in the local population, we investigate the distribution of offsets relative to the local relation. More specifically, for each galaxy we calculate the offset from the relation as $\Delta\text{Log}([\text{CII}]) = \text{Log}([\text{CII}]) - \text{Log}([\text{CII}]_{\text{expect-local}})$, where $\text{Log}([\text{CII}]_{\text{expect-local}})$ is the [CII] luminosity expected from the local relation according to the SFR measured in the galaxy or sub-component. The result of this distribution is shown in the left panel of Figure 5, while the right panel shows the distribution of the offsets. In contrast with some previous claims based on fewer targets, the $\Delta\text{Log}([\text{CII}])$ distribution, which includes both detections and upper limits, does not exhibit any clear shift relative to the local relation (and whose distribution is shown with the dotted green histogram. The number of objects below (23) and above (20) the L_[CII]-SFR relation are comparable. However, the dispersion of the is 0.48 ± 0.07 , which is about two times larger than the uncertainty reported by De Looze et al. (2014) for the local relation. Such larger dispersion may be associated to the presence of kpc-scale sub-components that are not common in the local Universe. However, we will discuss the possible origin of this dispersion in the next sections.

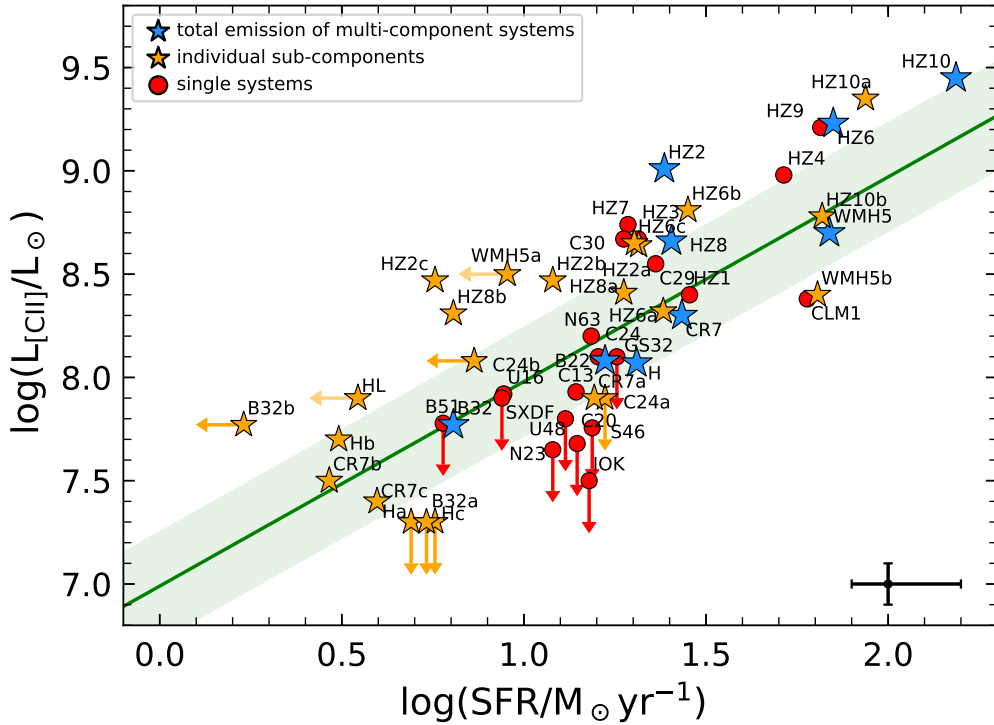


Figure 4. $L_{\text{[CII]}}$ as a function of SFR at $z = 5 - 7$. Blue stars indicate the global SFR and [CII] luminosity for system having multiple sub-components, as discussed in Section 3; in these cases the location of the individual sub-components are shown with yellow stars. Red circles show the location of the remaining high- z galaxies (i.e. single systems) listed in Table 2, that do not exhibit any positional offsets and/or disturbed morphology. In the bottom-right of the figure we show an error bar that is representative for the whole sample. The green line is the local relation for local star-forming galaxies (De Looze et al. 2014).

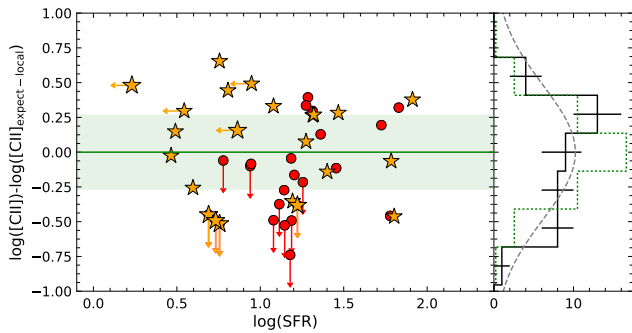


Figure 5. Deviations from the local $L_{\text{[CII]}}$ -SFR relation as a function of SFR. The 1σ dispersion of the local relation is indicated with the shaded green region. The distribution of the offsets is shown in the right panel. The black curve represents the $\text{Log}([\text{CII}]) - \text{Log}([\text{CII}]_{\text{expect-local}})$ distribution from our sample, while the dotted green histogram is the distribution of the local relation. The dashed grey line shows a Gaussian fit to the high- z distribution.

5 THE RELATION BETWEEN $L_{\text{Ly}\alpha}$ EW AND [CII] EMISSION

It is well known that $\text{Ly}\alpha$ emission depends on the level of ionizing photons produced by star formation (or AGN activity) and radiative transfer effects in the ISM. Models and observations suggest that the $\text{Ly}\alpha$ EW increases with decreasing metallicity and dust content (Raiter et al. 2010; Song et al. 2014). Since the [CII] emission is sensitive on the ISM properties as well, and in particular the ISM heating through photoelectric ejection from dust grains, we

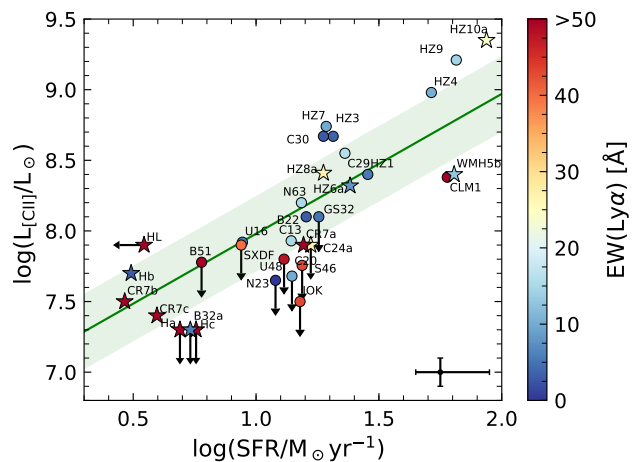


Figure 6. $L_{\text{[CII]}}$ as a function of SFR at $z = 5 - 7$. Stars and circles represent sub-components and individual galaxies, respectively. Symbols are colour-coded according to their $\text{Ly}\alpha$ EWs (not corrected for the intergalactic medium absorption), as indicated on the colour bar on the right. In the bottom-right corner we show an error bar that is representative for the whole sample. The green line are the local relation for local star-forming galaxies (De Looze et al. 2014) while its dispersion is indicated by the shaded green region.

expect a relation between the [CII] luminosity and the $\text{Ly}\alpha$ strength (Harikane et al. 2017).

In Figure 6 we show $L_{\text{[CII]}}$ as function of SFR by colour-coding

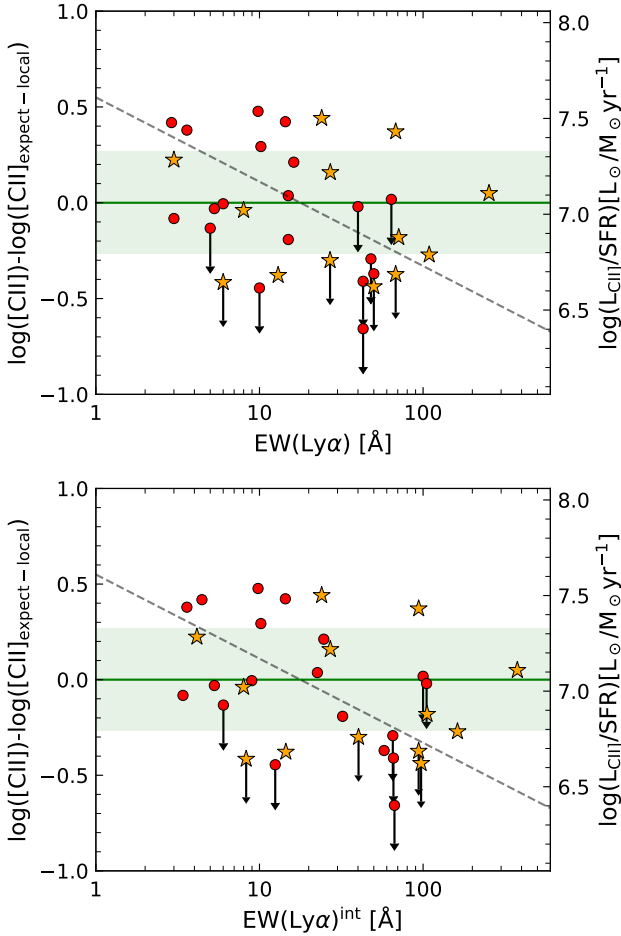


Figure 7. Top panel: Offset from the local $L_{[\text{CII}]}$ -SFR relation as a function of $\text{EW}(\text{Ly}\alpha)$. The right hand axis shows the $[\text{CII}]/\text{SFR}$ ratio corresponding to the deviations from the local $L_{[\text{CII}]}$ -SFR relation. The 1σ dispersion of the local relation is indicated by the shaded green region. Star and circles indicate sub-components and individual galaxies, respectively. The grey dashed relationship shows our linear fit (see text). Bottom panel: Same as the top panel but the $\text{EW}(\text{Ly}\alpha)^{\text{int}}$ has been corrected for the IGM absorption by following Harikane et al. 2017.

the different symbols according to their $\text{Ly}\alpha$ EWs (not corrected for the inter-galactic medium absorption). We include only those sub-components and galaxies having a $\text{Ly}\alpha$ EW measurement. There is a weak tendency for galaxies with high $\text{EW}(\text{Ly}\alpha)$ to lie below the local $L_{[\text{CII}]}$ -SFR relation, and vice-versa. This is shown better in Figure 7 where the offset from the local $L_{[\text{CII}]}$ -SFR relation is plotted as a function of $\text{EW}(\text{Ly}\alpha)$. Note that, since the slope of the local relation is one, plotting the deviation from the local relation ($\Delta \log([\text{CII}]) = \log([\text{CII}]) - \log([\text{CII}]_{\text{expect} - \text{local}})$) is equivalent to plotting the ratio between $[\text{CII}]$ luminosity and SFR, i.e. $L_{[\text{CII}]/\text{SFR}}$, which is indeed given on the right hand axis of Figure 6. Although there is a large dispersion, there is a tentative indication that the offset from the local relation (hence $L_{[\text{CII}]/\text{SFR}}$) anti-correlates with $\text{Ly}\alpha$ EW. A linear fit gives

$$\Delta \log([\text{CII}]) = (0.55 \pm 0.20) - (0.44 \pm 0.15) \log(\text{EW}(\text{Ly}\alpha))$$

and the dispersion around this best-fit is 0.25 dex.

The result does not change significantly if we attempt to correct the $\text{EW}(\text{Ly}\alpha)$ for IGM absorption, by following the prescription given by Harikane et al. (2017). In this case the relation is

shown in the bottom panel of Figure 6 and the resulting best-fit linear relation is

$$\Delta \log([\text{CII}]) = (0.56 \pm 0.20) - (0.41 \pm 0.14) \log(\text{EW}(\text{Ly}\alpha)^{\text{int}}) \quad (0.30 \text{ dex})$$

and the dispersion around this best-fit is even larger, 0.3 dex.

Harikane et al. (2017) find a relation between $L_{[\text{CII}]/\text{SFR}}$ and $\text{EW}(\text{Ly}\alpha)$ steeper than ours, though consistent within errors. The steeper relation found by Harikane et al. (2017) is probably associated with the fact that in their work they combine the global properties of galaxies and do not extract the sub-components.

Some anti-correlation between deviation from the local $L_{[\text{CII}]}$ -SFR relation and $\text{EW}(\text{Ly}\alpha)$ (or, equivalently, between $L_{[\text{CII}]}$ -SFR and $\text{EW}(\text{Ly}\alpha)^{\text{int}}$) is expected from the dependence of these quantities from the metallicity, either directly or through the associated dust content, as already predicted by some models (e.g. Vallini et al. 2015; Pallottini et al. 2017b; Matthee et al. 2017). Indeed, lower metallicity implies lower amount of carbon available for cooling, but also less dust content. Indeed, since the heating of PDRs occurs primarily through photoelectric effect on dust grains, the lower is the dust content the lower is the heating efficiency of the gas in the PDR, hence the lower is the emissions of the $[\text{CII}]$ cooling line. On the other hand, the lower the dust content the lower is the absorption of the $\text{Ly}\alpha$ resonant line, hence the higher is the $\text{EW}(\text{Ly}\alpha)$.

6 SPATIALLY RESOLVED $L_{[\text{CII}]}$ -SFR RELATION

In the previous sections, we show that the $L_{[\text{CII}]}$ -SFR relation at $z > 5$ has a intrinsic dispersion larger than observed in the local Universe. Such a large scatter suggests that the $[\text{CII}]$ luminosity may not be good tracer of the SFR at least at early epochs.

Recent spatially resolved studies have claimed that the $[\text{CII}]$ -SFR relation is better behaved in terms of SFR surface density and $[\text{CII}]$ surface brightness than in global proprieties ($L_{[\text{CII}]}$ and SFR), since the surface brightness calibration is more closely related to the local UV field (Herrera-Camus et al. 2015; Smith et al. 2017). It is thus worth to analyse the relation $\Sigma_{[\text{CII}]} - \Sigma_{\text{SFR}}$ at $z > 5$.

In this section we compare the spatial extension of the $[\text{CII}]$ and UV emission and, then, we investigate the correlation between $[\text{CII}]$ surface brightness and SFR surface density.

6.1 Spatial extension of the $[\text{CII}]$ emission

Figure 8 shows the extension of the $[\text{CII}]$ emission compared with the extension of star formation traced by the UV counterpart. $[\text{CII}]$ emission is generally much more extended than the UV emission tracing unobscured star formation. This discrepancy may be partially associated with observational effects. Indeed, while the high angular resolution of *HST* enables to resolve small clumps, it may have low sensitivity to diffuse, extended emission. However, even by smoothing the UV images in the deepest observations available to us, we still do not recover the extension observed in $[\text{CII}]$. On the other hand, the resolution of the $[\text{CII}]$ observations may, in some case, smear out clumps and result in an overall extended distribution. However, in many cases the ALMA observations achieve a resolution comparable, or even higher, than *HST* at UV rest-frame wavelengths and, despite this, we measure clearly larger $[\text{CII}]$ sizes. Moreover, when high angular resolution observation are used, these do reveal that a significant fraction of the $[\text{CII}]$ flux is resolved out on large scales (see e.g. discussion in Carniani et al. 2017a). In conclusion, we believe that the different sizes between $[\text{CII}]$ and

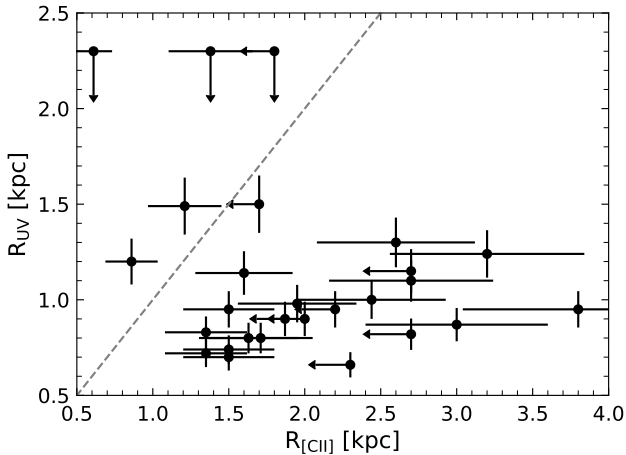


Figure 8. Half-light radii of star formation regions, as measured from the (rest-frame) UV light, compared with the half-light radii of the associated [CII] emission. The dashed line indicates the 1:1 relation.

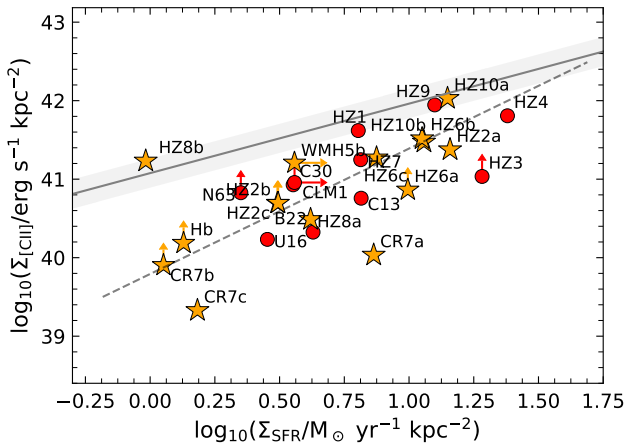


Figure 9. $\Sigma_{[CII]}$ vs Σ_{SFR} for $z > 5$ star-forming galaxies detected both in [CII] and rest-frame UV. Symbols are as in Figure 4. The grey line indicates the local relation by Herrera-Camus et al. (2015). The dashed line indicates the best linear fit on the $z > 5$ sample.

UV emission are tracing truly different distribution of the [CII] and UV emission on different scales.

There could be various explanations for these differences. If the star formation associated with the [CII] emission on large scale is heavily obscured, the UV light does not trace this component (Katz et al. 2017). This is certainly a possibility, although one might expect that the bulk of the obscuration should affect the central region more heavily than the outer parts, hence one would expect the opposite trend. In alternative, the extended component of [CII] may not be directly associated with star forming regions, but with circumgalactic gas, either in accretion or ejected by the galaxy, and which is illuminated by the strong radiation field produced by the galaxy.

6.2 Surface brightness

Once we have measured the extension of the [CII] emission and of the SFR regions, we can estimate the [CII] surface brightness and SFR surface density of each sub-component and of the individual sources detected in [CII] and UV. Figure 9 shows the $\Sigma_{[CII]}$ - Σ_{SFR} relation, where we have included only those galaxies detected in both [CII] and UV emission. Systems that are not spatially resolved in [CII] (or UV) emission are indicated with lower limits. We also show the local relation by Herrera-Camus et al. (2015) and its dispersion. In contrast to the $L_{[CII]}$ -SFR diagram, there are no galaxies located significantly above the local relation, only a few galaxies are located on the local relation, and most galaxies spread largely below the local relation. This is primarily due to the large extension of the [CII] emission in these high redshift systems, as discussed in the previous section. By fitting the $\Sigma_{[CII]}$ - Σ_{SFR} measurements for our sample we obtain the following relation:

$$\log(\Sigma_{SFR}) = (0.63 \pm 0.11) \times (\log(\Sigma_{[CII]}) - (25 \pm 6))$$

Part of the $\Sigma_{[CII]}$ deficit may be ascribed to the metallicity of the gas. Indeed a similar deviations have been observed in local low-metallicity galaxies, in which the $\Sigma_{[CII]}$ - Σ_{SFR} calibration overpredicts the $\Sigma_{[CII]}$ by up to a factor of six. We note that Faisst et al. (2017) estimated a metallicity of $12+\log(O/H) > 8.5$ for HZ10 and HZ9, and a metallicity of $12+\log(O/H) < 8.5$ for HZ1, HZ2, and HZ4. The former agree with the local $\Sigma_{[CII]}$ - Σ_{SFR} relation while HZ1, HZ2, and HZ4 have low $\Sigma_{[CII]}$ - Σ_{SFR} ratio. This interpretation is also supported by the fact that all CR7 clumps, which have metallicities $12+\log(O/H) < 8.2$ (see discussion in Sobral et al. 2017), are located well below the local relation. The $\Sigma_{[CII]}$ deficit is also akin to the simulations by Pallottini et al. (2017a) who investigated the physical properties of a simulated galaxy at $z = 6$, with metallicity $12+\log(O/H) = 8.35$. Such galaxy has a $\log(\Sigma_{[CII]}/\text{erg s}^{-1} \text{ kpc}^{-2}) = 40.797$ and $\log(\Sigma_{SFR}/M_{\odot} \text{ yr}^{-1}) = 1.027$, which places the mock galaxy below the $\Sigma_{[CII]}$ - Σ_{SFR} relation found for local star-formation galaxies.

However, the offset observed in high- z galaxies is significantly larger (one order of magnitude or more) than the one observed in local low metallicity galaxies, so other effects are likely in place, which will be discussed in the next section. Finally, still within the context of the surface brightness and SFR surface density, we mention that, for local galaxies, Smith et al. (2017) found a dependence between the $L_{[CII]}/L_{FIR}$ ratio and the SFR surface density, suggesting that the [CII] deficit increases strongly with increasing Σ_{SFR} . Therefore the Σ_{SFR} dependence has a strong impact also on the $L_{[CII]}$ -SFR relation. In our sample, we cannot verify this dependence as most galaxies have only an upper limit on the L_{FIR} . However, we can investigate the relation $L_{[CII]}/L_{UV}$ and Σ_{SFR} since the contribution of SFR_{FIR} to the total SFR is negligible (see discussion in Section 2.3). Figure 10 shows the observed $L_{[CII]}/L_{UV}$ ratio spanning a range between 0.002% to 0.4% (over two orders of magnitude), while the SFR surface density spans the range between 1 and $30 M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$. There is only a very weak correlation between $L_{[CII]}/L_{UV}$ and Σ_{SFR} , with very large dispersion (much larger than what observed locally for the $L_{[CII]}/L_{FIR}$ and Σ_{SFR} relation). This indicates that in these high- z systems the $L_{[CII]}/L_{UV}$ line ratio (and $L_{[CII]}/SFR$) does not strongly depend on the areal density with which galaxies form stars at $z > 5$, and that other effects (metallicity, and other phenomena discussed in the next section) may contribute to the very large dispersion.

- We have studied the relation between [CII] and SFR on the high- z sample by taking into account the presence of these sub-components and the proper associations between [CII] and UV components. The distribution of $z > 5$ galaxies on the $L_{[CII]}$ -SFR diagram follows the local relation, but the dispersion is 1.8 times larger than that observed in nearby galaxies.

- The deviation from the local $L_{[CII]}$ -SFR relation shows a weak anti-correlation with $EW(Ly\alpha)$ though shallower and with larger dispersion than what found in other studies that did not account for the multi-component nature of these systems.

- Most of the objects in the high- z sample are spatially resolved in [CII] and UV emission. The extension of the [CII] emission is generally much larger than the extension of star forming regions traced by the UV emission.

- In the $\Sigma_{[CII]}$ - Σ_{SFR} diagram $z > 5$ galaxies are characterised by a large scatter with respect to local galaxies, and are mostly distributed below the local relation (i.e. fainter $\Sigma_{[CII]}$ at a given Σ_{SFR}).

We have suggested that a combination of different effects may be responsible for the different properties of high- z galaxies in terms of [CII]-SFR properties relative to local galaxies. More specifically: 1) the low metallicity of high- z galaxies may be responsible (also indirectly through the lower dust photoelectric heating) for part of the scatter towards lower [CII] emission relative to the local relations; 2) the presence of circumnuclear gas in accretion and/or expelled from the galaxy may be responsible for the larger size in [CII] relative to the SFR distribution and may also be responsible for the scatter of the $L_{[CII]}$ -SFR distribution above the local relation as a consequence of lower ionisation parameter; 3) in compact young star forming regions the increased ionisation parameter and higher gas density may be responsible for the suppression of [CII] for galaxies which are below the local relation; 4) dust obscuration may be responsible for both the different morphology between [CII] and UV emission and also for the scatter of sources above the local $L_{[CII]}$ -SFR relation.

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four sources. The flux calibrators were J2056-4714, Ganymede, J1229+0203, J1229+0623, and J0334-4008 while bandpass calibrations were carried out through the observations of J2258-2758, J1058+0133, J1229+0203, J0238+1636, and J0522-3627.

ALMA observations were calibrated by using CASA software version v4.5.2 (McMullin et al. 2007). Continuum and data cube images were obtained by using the CASA task `CLEAN` and natural weighting. The final angular resolution and sensitivity reached in each set of data are listed in Table A1.

We registered NIR images to ALMA observations by matching the location of the foreground sources and ALMA calibrators to the position given by the GAIA Data Release 1 catalogue (Gaia Collaboration et al. 2016).

As discussed in Section 2.2, [CII] emission was detected only in one target, BDF3203. Figures 1 and A1 show the ALMA [CII] spectra for the five star-forming galaxies extracted from a region as large as the ALMA beam and centred at the location of the rest-frame UV regions.

APPENDIX A: ADDITIONAL ALMA DATA - OBSERVATION AND DATA REDUCTION

The five $z \sim 6$ star-forming galaxies listed in Table A1 were observed with ALMA in band 6 during Cycle 3 and Cycle 4 (program ID #2015.1.01105.S and #2015.1.01240.S). ALMA observations were carried out with a semi-compact array configuration with angular resolutions ranging from 0.3'' and 0.8''. The sources were observed for a total on source integration time of 0.1-0.8 h with a precipitable water vapour of 0.4-1.4 mm, depending on the specific observation. For each target we used four spectral windows (SPWs) set up in frequency division mode with a spectral resolution of ~ 30 MHz (~ 35 km/s) and bandwidth of 1.875 GHz. One of the four SPWs was tuned to the expected frequency of the [CII] line. The phase of each observation was centred at the NIR position of the respectively source.

J2248-3235, J0948-002, J1147-0724, J0239-0234, and J0552-3627 were observed as phase calibrator, respectively for the

Table 2. Star-forming galaxies at $z > 5$

NAME	ID	Redshift	$\log(L_{UV})$ [L_{\odot}]	SFR_{UV} [$M_{\odot} \text{ yr}^{-1}$]	$\log(L_{[CII]})$ [L_{\odot}]	r_{UV} [kpc]	$r_{[CII]}$ [kpc]	S_{mm} [μJy]	$\log(L_{FIR})$ [L_{\odot}]	(k)
(a)	(b)	(c)	(d)	(e)	(f)	(g)	(h)	(i)	(j)	(k)
HZ8*	HZ8	5.1533	11.17	25.4	8.7	–	–	< 90	< 10.7	[1,2,3]
HZ8a	HZ8a	5.1533	11.04	18.8	8.4	1.2	3.2	< 90	< 10.7	[1,2,3]
HZ8b	HZ8b	5.1533	10.57	6.4	8.3	1.5	1.2	< 90	< 10.7	[1,2,3]
HZ7	HZ7	5.2532	11.05	19.3	8.7	1.0	1.9	< 108	< 10.8	[1,2,3]
HZ6*	HZ6	5.2928	11.47	50.7	9.2	–	–	220	11.1	[1,2,3]
HZ6a	HZ6a	5.2928	11.11	22.1	8.3	0.9	< 1.9	30	10.3	[1,2,3]
HZ6b	HZ6b	5.2928	11.0	17.2	8.8	0.8	1.6	120	10.9	[1,2,3]
HZ6c	HZ6c	5.2928	10.81	11.1	8.7	0.8	1.7	100	10.8	[1,2,3]
HZ9	HZ9	5.541	10.95	15.3	9.2	0.9	1.5	516	11.5	[1,2,3]
HZ3	HZ3	5.5416	11.08	20.6	8.7	0.7	< 2.3	< 153	< 11.0	[1,2,3]
HZ4	HZ4	5.544	11.28	32.7	9.0	0.7	1.4	202	11.1	[1,2,3]
HZ10*	HZ10	5.6566	11.19	26.6	9.4	–	–	1261	11.9	[1,2,3]
HZ10a	HZ10a	5.6566	11.14	23.7	9.3	1.1	1.6	630	11.6	[1,2,3]
HZ10b	HZ10b	5.6566	10.23	2.9	8.8	0.7	1.5	630	11.6	[1,2,3]
HZ2*	HZ2	5.6597	11.15	24.3	9.0	–	–	< 87	< 10.8	[1,2,3]
HZ2a	HZ2a	5.6597	11.08	20.6	8.6	0.7	1.5	< 87	< 10.8	This Work
HZ2b	HZ2b	5.6597	10.85	12.0	8.5	1.1	< 2.7	< 87	< 10.8	This Work
HZ2c	HZ2c	5.6597	10.52	5.7	8.5	0.8	< 2.7	< 87	< 10.8	This Work
HZ1	HZ1	5.6885	11.21	28.5	8.4	1.2	0.9	< 90	< 10.8	[1,2,3]
WMH5*	WMH5	6.0695	11.36	59.0	8.7	–	–	91	10.8	[6,7]
WMH5a	WMH5a	6.0695	< 10.5	< 5.0	8.5	< 2.3	0.6	42	10.5	[6,7]
WMH5b	WMH5b	6.0695	11.36	59.0	8.4	< 2.3	1.4	49	10.6	[6,7]
NTTDF2313	N23	6.07	10.85	12.0	< 7.7	< 1.4	–	< 54	< 10.6	This Work
BDF2203	B22	6.12	10.97	16.0	8.1	1.1	2.7	< 69	< 10.7	This Work
CLM1	CLM1	6.1657	11.37	60.0	8.4	< 2.3	< 1.8	< 78	< 10.8	[6]
GOODS3203	GS32	6.27	11.02	18.0	< 8.1	1.1	–	< 123	< 11.0	This Work
COSMOS20521	C20	6.36	10.89	14.0	< 7.7	0.7	–	< 60	< 10.7	This Work
UDS4812	U48	6.561	10.87	13.0	< 7.8	0.7	–	< 72	< 10.8	This Work
Himiko*	H	6.595	11.07	20.4	8.1	–	–	< 27	< 10.4	[13,14]
Himiko-a	Ha	6.595	10.45	4.9	< 7.3	0.7	–	< 27	< 10.4	[13,14]
Himiko-b	Hb	6.595	10.26	3.1	7.7	0.9	< 2.0	< 27	< 10.4	[13,14]
Himiko-c	Hc	6.595	10.49	5.4	< 7.3	0.7	–	< 27	< 10.4	[13,14]
Himiko-Lya	HL	6.595	< 10.31	< 3.5	7.9	–	3.4	< 27	< 10.4	[13,14]
CR7*	CR7	6.604	11.2	27.1	8.3	–	–	< 21	< 10.2	[10]
CR7a	CR7a	6.604	10.9	15.6	7.9	0.9	3.0	< 21	< 10.2	[10]
CR7b	CR7b	6.604	10.2	2.9	7.5	0.9	< 2.2	< 21	< 10.2	[10]
CR7c	CR7c	6.604	10.3	4.0	7.4	0.9	3.8	< 21	< 10.2	[10]
COSMOS24108*	C24	6.6294	10.99	16.7	8.1	–	–	< 54	< 10.7	[4]
COSMOS24108a	C24a	6.6294	10.99	16.7	< 7.9	1.1	–	< 54	< 10.7	[4]
COSMOS24108b	C24b	6.6294	< 10.37	< 7.3	8.1	–	2.1	< 54	< 10.7	[4]
UDS16291	U16	6.6381	10.71	8.8	7.9	1.0	2.4	< 60	< 10.7	[4]
NTTDF6345	N63	6.701	10.95	15.3	8.2	1.5	< 1.7	< 48	< 10.6	[4]
COS-2987030247	C29	6.8076	11.11	23.0	8.6	–	3.1	< 75	< 10.8	[5]
SDF46975	S46	6.844	10.94	15.4	< 7.8	< 2.0	–	< 58	< 10.7	[8]
COS-3018555981	C30	6.854	11.04	18.8	8.7	1.3	2.6	< 87	< 10.9	[5]
IOK-1	IOK	6.96	10.94	15.1	< 7.5	0.6	–	< 63	< 10.8	[11]
BDF512	B51	7.008	10.54	6.0	< 7.8	0.5	–	< 52	< 10.7	[8]
BDF3299*	B32	7.107	0.0	6.4	7.8	–	–	< 23	< 10.3	[8,9]
BDF3299a	B32a	7.107	10.52	5.7	< 7.3	0.9	–	< 23	< 10.3	[8,9]
BDF3299b	B32b	7.107	< 10.0	< 1.7	7.8	–	1.0	< 23	< 10.3	[8,9]
COSMOS13679	C13	7.1453	10.91	13.9	7.9	0.8	1.4	< 42	< 10.6	[4]
SXDF-NB1006-2	SXDF	7.212	11.06	8.7	< 7.9	–	–	< 42	< 10.6	[12]

Notes: (a) Name of the source; the asterisk mark (*) indicates that the source has a multi-clump morphology. (b) ID used to indicate the source in the figures of this paper. (c) Redshift of the galaxy (or system) inferred from Ly α . (d) Rest-frame UV luminosity at 1600Å. (e) SFR based on the UV emission: $\log(SFR/M_{\odot} \text{ yr}^{-1}) = \log(L_{UV}/\text{erg s}^{-1}) - 43.35$ (Murphy et al. 2011; Hao et al. 2011; Kennicutt & Evans 2012) (f) [CII] luminosity. (g, h) Half-light radius in kpc for UV and [CII] emission. (i) Continuum emission (or 3σ upper limit) at rest-frame 158 μm . (j) FIR luminosity (or 3σ upper limit) estimated from ALMA observations. (k) References: [1] Capak et al. (2015), [2] Barišić et al. (2017), [3] Faisst et al. (2017), [4] Willott et al. (2015b), [5] Jones et al. (2017b), [6] Ouchi et al. (2013), [7] Matthee et al. (2017), [8] Pentericci et al. (2016), [9] Smit et al. (2017), [10] Maiolino et al. (2015), [11] Ota et al. (2014), [12] Carniani et al. (2017a), [13] Inoue et al. (2016), [14] Carniani et al. (2017b).

Table A1. UV and far-IR properties of the new five $z \sim 6 - 7$ sources observed with ALMA.

Name	RA(J2000) [deg]	DEC(J2000) [deg]	$z_{\text{Ly}\alpha}^a$	SFR_{UV}^b $M_{\odot} \text{ yr}^{-1}$	$\nu_{\text{obs}}(\text{[CII]})^c$ [GHz]	beam ^d [min'' \times maj'']	t_{exp}^e [hours]	σ_{cont}^f [μJy]	σ_{line}^g [mJy]	$L_{\text{[CII]}}^h$ [$10^7 L_{\odot}$]
NTTDF2313	181.3804	-7.6935	6.07	12	268.817	0.98×0.71	0.7	18	0.15	<4.5
BDF2203	336.958	-35.1472	6.12	16	266.93	1.90×1.11	0.4	23	0.2	12.5 ± 2.5
GOODS3203	53.0928	-27.8826	6.27	18	250.236	1.26×1.03	0.1	41	0.4	<12.0
COSMOS20521	150.1396	2.4269	6.36	14	258.225	1.46×1.20	0.8	20	0.15	<4.8
UDS4821	34.4768	-5.24728	6.561	13	251.361	0.24×0.22	0.3	24	0.2	<6.7

Notes: ^aRedshift from either Ly α line or spectroscopic Lyman break. The uncertainty is < 0.04 . ^bSFR inferred from the rest-frame UV continuum adopting the calibration discussed in Kennicutt & Evans (2012). ^cExpected [CII] frequency according to $z_{\text{Ly}\alpha}$. ^dALMA synthesised beam. ^eOn-source integration time. ^fSensitivity in ALMA continuum map. ^gSensitivity in spectral channels of 100 km/s. ^h[CII] luminosity. The upper limits on the $L_{\text{[CII]}}$ are at 3σ , and are calculated on a width of 100 km/s.

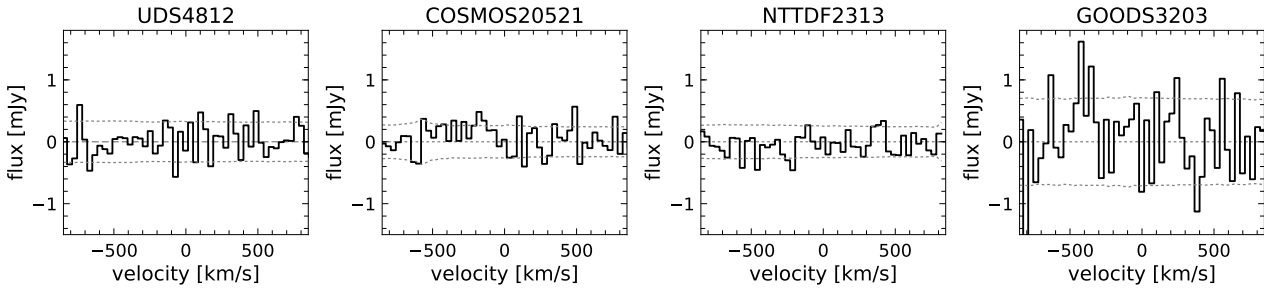


Figure A1. ALMA spectra of UDS4812, COSMOS20521, NTTDF2313, and GOODS3203. The spectra have been extracted at the location of the UV emission within a ALMA beam region. The velocity reference is set to the redshift inferred from spectroscopic rest-frame optical observations. The dotted grey lines shows the 1σ and -1σ