



<b>Publication Year</b>	2015
<b>Acceptance in OA</b>	2020-03-18T18:02:26Z
<b>Title</b>	The ASTRI SST-2M prototype for the Cherenkov Telescope Array: opto-mechanical test results
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<b>Publisher's version (DOI)</b>	10.1117/12.2189767
<b>Handle</b>	<a href="http://hdl.handle.net/20.500.12386/23378">http://hdl.handle.net/20.500.12386/23378</a>
<b>Serie</b>	PROCEEDINGS OF SPIE
<b>Volume</b>	9603

# The ASTRI SST-2M prototype for the Cherenkov Telescope Array: opto-mechanical test results

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## ABSTRACT

The Cherenkov Telescope Array (CTA) observatory, with a combination of large-, medium-, and small-size telescopes (LST, MST and SST, respectively), will represent the next generation of imaging atmospheric Cherenkov telescopes. It will explore the very high-energy domain from a few tens of GeV up to few hundreds of TeV with unprecedented sensitivity, angular resolution and imaging resolution.

In this framework, the Italian ASTRI program, led by the Italian National Institute of Astrophysics (INAF), is currently developing a scientific and technological SST dual-mirror end-to-end prototype named ASTRI SST-2M. It is a 4-meter class telescope; it adopts an aplanatic, wide-field, double-reflection optical layout in a Schwarzschild-Couder configuration. The ASTRI SST-2M telescope structure and mirrors have been already installed at the INAF observing station at Serra La Nave, on Mt. Etna (Sicily, Italy).

In this contribution we report about the on-site deployment and the latest results on the opto-mechanical performance test conducted soon after the telescope installation.

**Keywords:** imaging atmospheric Cherenkov telescope, very high-energy gamma-rays, wide-field aplanatic telescope, CTA, ASTRI, segmented optics

## 1. INTRODUCTION

Recent advances in Gamma-ray astronomy have shown that the 10 GeV – 100 TeV energy band is crucial to investigate scientific topics such as pulsars and binary systems, blazars, radio-galaxies, the galactic center, pulsar wind nebulae and star-forming galaxies, as well as the physics in extreme conditions. A cost-effective way to perform observations in this band is with ground-based experiments through imaging atmospheric Cherenkov telescopes. Very-high energy photons (and cosmic rays) interact with the Earth atmosphere and generate an  $e^+e^-$  couple. The pair travels faster than light in the medium and polarizes the molecules of the air. A faint flash of beamed photons in the 300-550 nm wavelength band is emitted at the de-excitation. This phenomenon is known as the Cherenkov effect, which is effective until the energy of

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the gamma-rays goes below the threshold for pair production in air. Air Cherenkov telescopes collect the flash of photons using suitable multi-pixel sensors and make morphological analysis of the images.

At present, H.E.S.S. [1], VERITAS [2] and MAGIC [3] are the state of the art of such ground-based experiments. The Cherenkov Telescope Array (CTA) will represent a breakthrough in this field [4].

CTA has the goal of increasing the sensitivity by a factor of 10, to expand the detection bandwidth both in the low- and in the high-energy limits, and improve the angular resolution with respect to the present installations. The CTA observatory is a project involving a worldwide consortium of research Institutes and Universities that will make use of cutting-edge technologies. The observatory will be deployed over two sites, one for each hemisphere, in order to ensure full sky coverage observations. CTA will be composed of telescopes with different sizes. At the CTA southern site, 70 Small Size Telescopes (4 m primary mirror diameter), 25 Medium Size Telescopes (12 m) and 4 Large Size Telescopes (23 m) are envisaged. Additional 24 Schwarzschild-Couder Telescopes (9 m) will be possibly installed to improve the sensitivity at about 1 TeV and to exploit high-precision angular resolution studies. In the northern site, 15 Medium Size Telescopes (12 m) and 4 Large Size Telescopes (23 m) are planned. The Medium Size and the Schwarzschild-Couder telescopes are discussed in [5] and [6].

The Italian National Institute for Astrophysics (INAF) is leading the design and development of an end-to-end telescope prototype for the SST class. This activity is carried out within the ASTRI flagship project funded by the Italian Ministry of Education, University and Research. The follow-up of the ASTRI project consists of the deployment of the first pre-production telescopes on the CTA southern site. This follow up is partially funded also by the Fundação de Amparo à Pesquisa do Estado de São Paulo and by the Department of Science and Technology of the Republic of South Africa. The mini-array will be the seed from which the entire CTA southern site will be deployed.

INAF has committed the electro-mechanical design of the telescope prototype, named ASTRI SST-2M, to BCV progetti ([www.bcv.it](http://www.bcv.it)) and Tomelleri ([www.tomelleri.com](http://www.tomelleri.com)) companies, while INAF has kept in house the design and the development of optics, mirrors, detector and software. Details on the design are reported in [7], [8] and [9]. The manufacturing of the electro-mechanical structure of the telescope has been awarded to Galbiati group ([www.galbiatigroup.it](http://www.galbiatigroup.it)) and EIE group ([www.eie.it](http://www.eie.it)) companies forming the GEC consortium. Details on the manufacturing are in [10] and [11].

The site of the ASTRI SST-2M opto-mechanical structure is the “M. G. Fracastoro” observing station in Serra La Nave (Sicily, Italy), at about 1735 m a.s.l. [12]. It is in operation since September 2014.

In the following sections we give a short overview of the opto-mechanical solutions implemented (see section 2) and then we report the results from the initial on-site tests (see section 3).

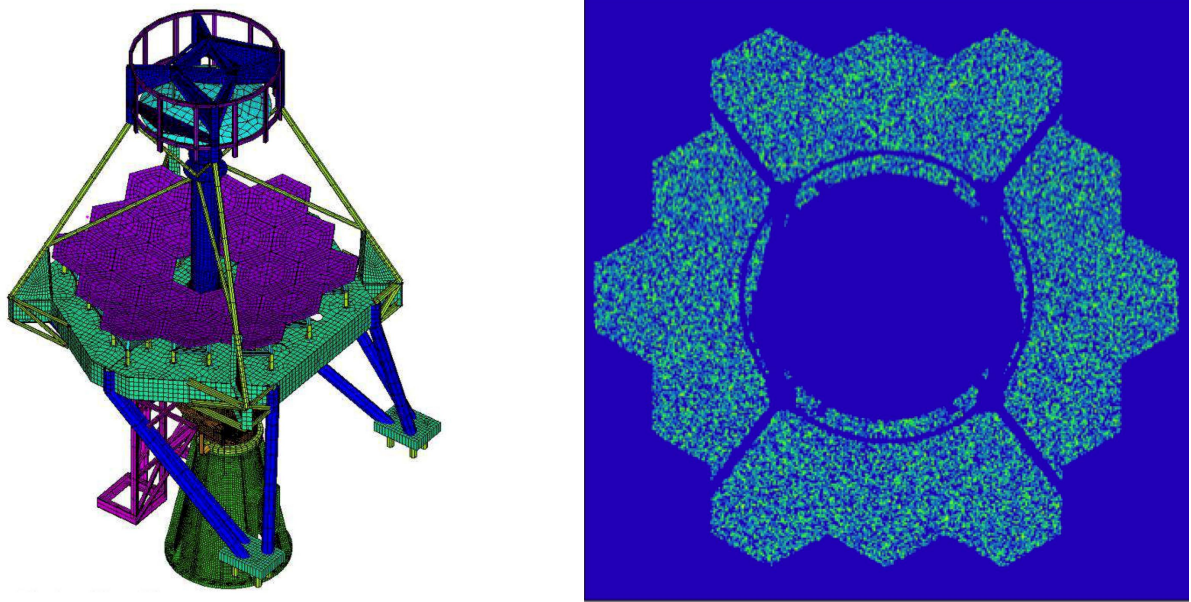
## 2. THE TELESCOPE DESIGN

The mechanical design of the ASTRI SST-2M telescope prototype is the outcome of trade-off activities among different configurations. The chosen solution has been further refined and optimized until the results of the analyses met the design specifications. Finite Elements numerical approach (FE Model and FE Analysis) is the design verification tool adopted for the static evaluation of all the telescope structural components. The integrated FE model of the telescope has about 700.000 degrees of freedom. It takes into account also the characteristic of the mechanical components, the mirrors and other devices on board of the telescope (shields, baffle, scientific detector, etc.). An example of the FE model developed within the design phase of the project is shown in the left panel of Figure 1. Simple dynamical analyses have also been performed in order to investigate the non-static behavior of the telescope.

The dimensioning of the structural components was driven by the CTA project requirements such as the design loads, the optical tolerances, reliability-availability-maintainability-safety criteria and others. Indeed, the structural checks (e.g. winds, earthquakes, etc.) have been performed in accordance with the most updated national and international norms and standards (i.e. Eurocodes, EC machinery directive). The evaluation of the telescope in terms of optical-related performance (vignetting, mis-alignments/mis-orientations of the mirrors) has been computed by means of ray-tracing analysis.

The Eigen-frequencies have a nearly constant behavior with respect to the azimuth axis; the first modes have values in any case higher than 4 Hz. They are a twist around the azimuth axis and tilts around elevation and cross-elevation axes ascribed to the stiffness of the mechanical components. Dynamic behavior on tracking shows stick and slip phenomenon due to the difference between static and dynamic friction on the azimuth bearing. However, the global tracking errors of the order of few arcsec rms also in presence of wind disturbances are well below the specification.

Concerning optical degradations due to structural performance, we observe that the point spread function remains within the requirement in any observing condition while the obstruction is at most of about 14%. The right panel of Figure 1 shows the ray-tracing of the ZEMAX configuration implemented for this purpose.



**Figure 1:** (left) Global finite element model of the ASTRI SST-2M telescope prototype. (right) Evaluation of the obstruction due to the optical elements and the mechanical structures of the telescope.

## 2.1 The electro-mechanical configuration

The ASTRI SST-2M telescope structure is primarily made of steel grade S355 material. It has an alt-azimuthal mounting type. The base is a truncated cone with reinforcing ribs to distribute loads. The base interfaces the concrete foundation by means of 24 anchor bolts, earthquake devices and electrical/communication terminals. At the opposite end, the azimuth bearing and main gear are hosted. Inside the base, the azimuth strip encoder with 4 scanning heads, the references for the motion's limit switches and the cables routing are disposed. A door is available to give access to these devices.

The azimuth fork, one of the main structural elements, is mounted on top of the telescope base together with the azimuth drive system (motors and gear boxes). The drive system has two units that work in master-slave configuration during both slewing and tracking. This operative mode neglects the backlash and guarantees adequate precision during the motion. The fork interfaces the primary mirror cell and the elevation actuator. The support frame for the electrical cabinets, the cable trays and the elevation cable wraps are also hosted on the fork. The fork can rotate around the azimuth axis for  $540^\circ$ .

The cell of the primary mirror is a thick ribbed flat dish of about 4.3 meter in side. It is moved around the elevation axis by a long and stiff linear actuator within the range  $-5^\circ$  and  $+95^\circ$ . The actuator, manufactured from Servomech company [13], has a single motor with preloaded gearbox. Backlash is completely removed because the telescope is slightly unbalanced about the elevation axis. A rotary absolute encoder is directly coupled to the elevation axis. The dish hosts 18 reflective segments, that compose the primary mirror; the quadrupode, that supports the secondary mirror; and the central tube, that supports the focal plane instrument. The quadrupode has an eccentric symmetry with radial bracings toward the central tube; this solution improves the bending stiffness against gravity deformations introduced by the large secondary mirror.

A triangular prism and a proper arrangement of tapered I-beams compose the cell of the secondary mirror. The triangular prism supports the mirror subsystem itself while the tapered beams provide interfaces toward the quadrupode legs. The cell also has supports the snow shields, windshields and optical baffles.

Electromechanical stow pins are designed for locking the telescope in parking position. They are designed for locking the telescope in position up to the worst loading conditions. Elevation bumpers are also envisaged.

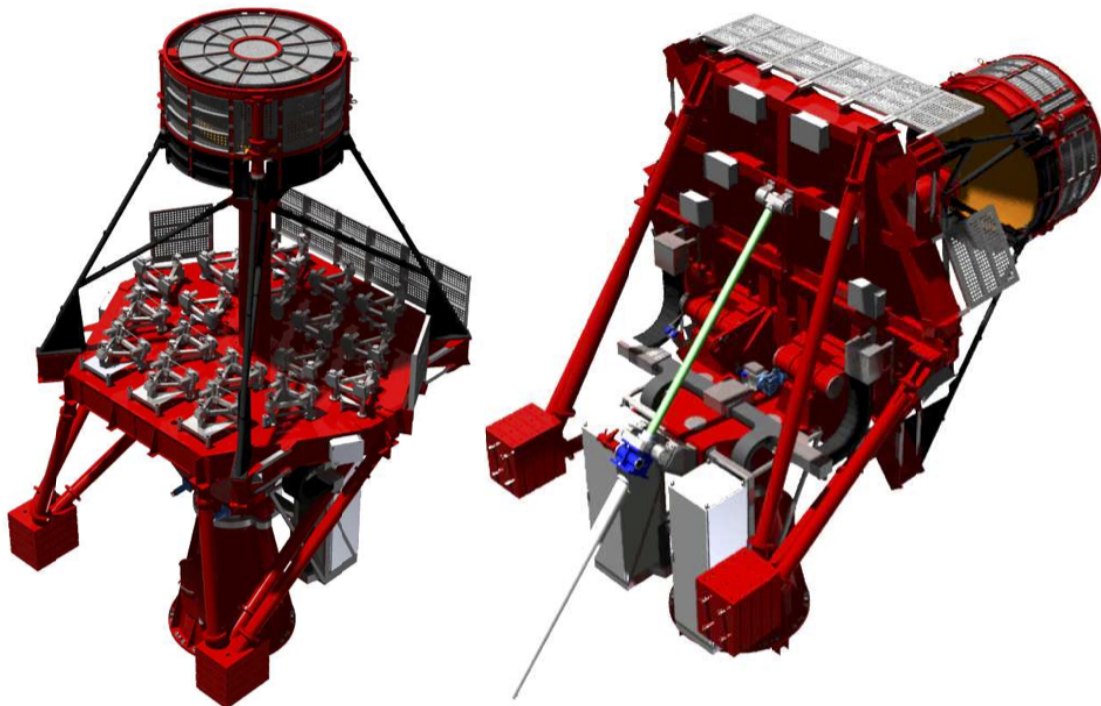
Concerning the drive system of the telescope a solution based on SEW-Eurodrives components has been adopted for both azimuth and elevation. In azimuth, the ratio-gear is SEW R97R57 with the motor SEW CM71S; while in elevation the motor SEW CM90S is used. A brake (braking torque equal to 5 Nm) with automatic and manual release is also available on both axes; moreover a secondary motor shaft allows the driving of the telescope with the aid of a battery drill for emergency situations (power cuts, drive or software failures etc.).

The motion control, safety and health monitoring systems are based on PC/PLC solution by Beckhoff Automation company [14]. The fieldbus is based on EtherCAT standard.

The telescope has two electrical cabinets: the high power cabinet and the low power cabinet. They are designed accordingly to IEC 61439 and EN 60204-1. Both cabinets are designed for outdoor applications with a protection degree IP65. The electrical cabinets are thermally controlled. They guarantee normal operation of the installed equipment between 0°C to 40°C for any environmental conditions. Heating resistors and fans have been opportunely chosen and integrated inside the cabinets; moreover, an anti-condensation resistor is present opportunely commanded by a hygostat. The high power cabinet provides the main power distribution to the entire telescope. At its entrance a switch fuse is installed. This device guarantees the protection of the downstream circuits against over-currents together with the disconnection function, which ensures that the electrical circuit is completely de-energized for service or maintenance. The high power cabinet also hosts drives and inverter electronics of the motion system.

The low power cabinet is a complex subsystem that hosts a number of devices such as: all the electronic hardware in charge to control and monitor the telescope; the hardware in charge to power some specific and auxiliary equipment and the servo minor drives (primary and secondary mirrors alignment system); the hardware in charge of the data/communication. At the entrance of the cabinet a load disconnecting switch is installed. It guarantees the possibility to open the circuit also when energized while the over-current protection is provided by the circuit breaker installed in the other cabinet.

The rendering of the ASTRI SST-2M telescope structure and electro-mechanical subsystems is shown in Figure 2.



**Figure 2** Rendering of the ASTRI SST-2M telescope structure and electro-mechanical subsystems

## 2.2 The optical configuration

The ASTRI SST-2M prototype implements an aplanatic, wide field, double reflection optical layout known as Schwarzschild-Couder configuration. Vassiliev et al. firstly proposed the use of this solution to Cherenkov telescopes [15]. The actual layout implemented in the ASTRI SST-2M telescope follows the requirements for the SST array of the CTA Observatory. In particular, the pixel size is  $0.165^\circ$ , the large field of view is  $9.6^\circ$  and the effective area is larger than  $4.5 \text{ m}^2$  (enough to trigger at  $1 - 3 \text{ TeV}$ ). The telescope has an equivalent focal length of  $2150 \text{ mm}$  with  $F\# = 0.5$ ; the plate scale is about  $37.5 \text{ mm}/^\circ$ . The trend of the angular resolution along the field of view is shown in the left panel of Figure 3.

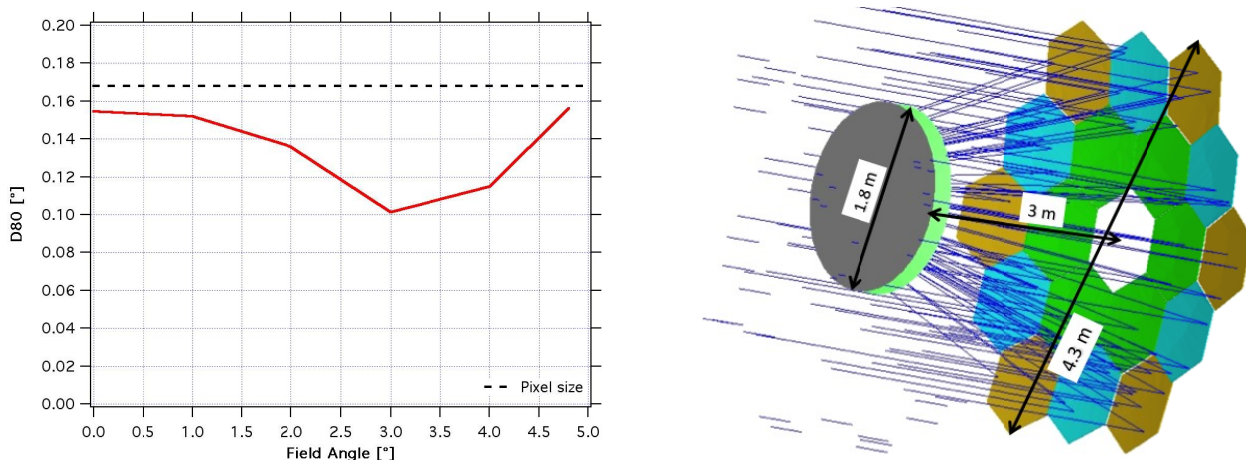
The primary mirror is tasseled in 18 panels distributed on three coronae; they compose the full aperture of the telescope of about  $4.3 \text{ m}$ . On the contrary, the secondary mirror is a large monolithic element of  $1.8 \text{ m}$  in diameter. Both optical surfaces have a concave profile with very pronounced deviations from the generating sphere. In fact, the asphericity reaches deviations in terms of peak-to-valley of  $PV_{M1} = 27 \text{ mm}$  and  $PV_{M2} = 1.3 \text{ mm}$ , for primary and secondary mirrors respectively. The focal surface has a spherical convex profile with a radius of curvature of about  $1060 \text{ mm}$ . The layout is shown in the right panel of Figure 3.

Each mirror segment is mounted on a triangular spacer that also introduces a tilt to bring the mirror segment into its rough position. The supports are equipped with two active actuators and one fixed point. The actuators have stepper motors to perform the fine alignment of the mirror. Concerning the secondary mirror, a total of 9 points distributed over the mirror area on a whiffletree configuration act as support. Three active actuators with stepper motors and absolute encoders provide tip-tilt and piston motion to the subsystem. All the active optics devices have drives communicating through CANbus.

The mirrors are manufactured adopting in both cases a somewhat similar technology. Both are made in glass applying processes that take advantage by the concept of replica: the profile of a master shape is copied to produce several identical mirrors.

Concerning the primary mirror, the technology adopted is the glass cold shaping process recently developed for making the mirrors of the MAGIC-II telescope, see [16] and [17]. A glass sheet is bent, at room temperature, by vacuum suction over a replication mold. To keep the shape, a reinforcing core in aluminum honeycomb and a second glass skin are glued together forming a stiff and lightweight sandwich structure. At the end of the polymerization a high reflective coating is deposited on the substrate. A detailed description of the principles over which the cold-shaping process is based is presented in [18] while an overview of the projects that so far have adopted this technological solution is presented in this conference proceeding [19]. This technology, with different implementations, is now being adopted as baseline for manufacturing the entire set of mirrors for CTA. The measurement results and expected optical performance of the telescope primary mirror are reported in [20].

The secondary mirror comes from a monolithic glass plate  $19 \text{ mm}$  thick that undergoes a thermal bending process. The bending is gravity assisted with a proprietary process developed by the FLABEG company ([www.flabeg.com](http://www.flabeg.com)). Details of the secondary mirror are reported in [21].



**Figure 3:** (left) Evolution of the angular resolution along the field of view. (right) Geometrical layout of the optical design implemented for the ASTRI SST-2M telescope.

### 3. THE ON-SITE INSTALLATION AND PRELIMINARY TESTS RESULTS

#### 3.1 The telescope installation

Figure 4 shows the main phases of the ASTRI SST-2M prototype on-site installation.

The assembly of the telescope has begun on August 22<sup>nd</sup> 2014 with the positioning of the mount assembly on the foundations (see Figure 4 top-left panel): the base, the fork and the electrical cabinets were pre-assembled at the factory, packed and shipped. Grouting has been concurrently executed. The perpendicularity has been measured by means of a spirit level with 2 arcsec resolution. The alignment tilt error is below 20 arcsec with respect to the gravity vector. The tensioning of the anchor bolts has been performed in two separate steps by means of a hydraulic jack up to the design value of 230 kN. The remaining structural parts of the telescope (the primary mirror dish, the quadrupode, the secondary mirror cell and the counterweights) have been installed, as per assembly plan, in the subsequent 5 days (see Figure 4 top-right panel).

The second phase was devoted to mirrors installation. Each segment of the primary mirror has been integrated to its active support structure and, using proper handling tools (see Figure 4 bottom-left panel), installed on the telescope (see Figure 4 bottom-right panel) as per procedure. Two days were necessary to complete the installation of the 18 segments of the primary mirror.

Three high-skilled mechanical technicians, one crane operator and an engineer coordinating the activities have carried out the opto-mechanical installation in about 10 days. The main supporting equipment was a mobile crane for lifting and positioning of the structural elements and a cherry picker to bring the technicians in place.

At last, the electrical routing and cabling of the entire telescope and auxiliary subsystems have been done. Two technicians and an engineer coordinating the activities have completed the work in 10 additional days.

At the end, the first complete integration on field of the ASTRI SST-2M prototype took more or less 30 days, including the materials on-site acceptance and inspection.



**Figure 4:** Highlights of the telescope opto-mechanical on-site installation. (top-left) Positioning and alignment of the mount assembly on the foundation (top-right) Completion of the telescope structure (bottom-left) Primary mirror segments and active optical support integration (bottom-right) Primary mirror installation.

### 3.2 Electro-mechanical and opto-mechanical tests

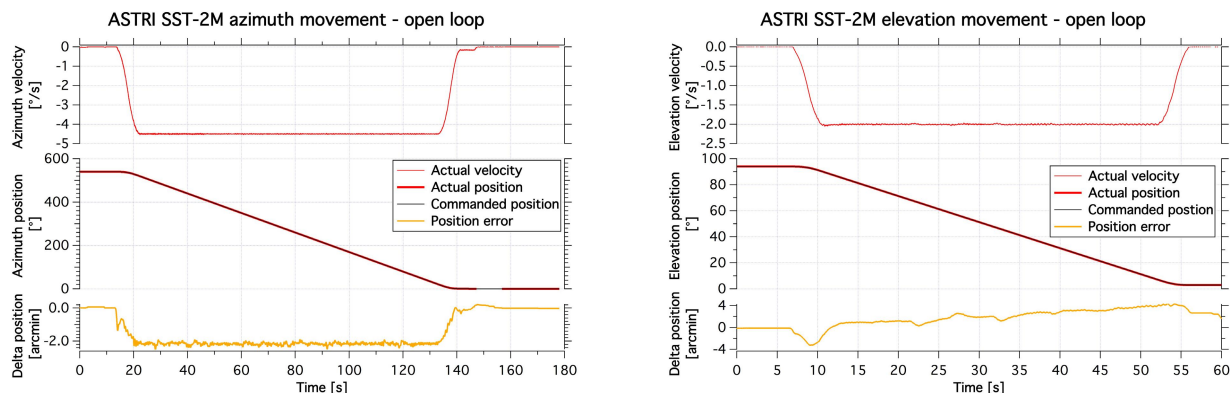
The electro-mechanical tests consist in a series of functional tests to assess the low level performance of the telescope. Tests concern for instance the electrical functionality of all the limit switches as well as their intervention on the telescope motion (interlock chain), the maximum velocities and accelerations achievable, the braking range upon emergency stop, the motors torque bias, the stow pins insertion and extraction, the axes motion range amplitude, the emergency movement, the telescope Eigen frequencies, the power consumption and the duration of the telescope control and safety subsystems under UPS.

To evaluate the axes motion a dedicated test has been performed. The telescope has been commanded to perform a full run at its maximum velocity until the limit switches intervention. The telescope shows a wide operation range both on azimuth and elevation. The total usable run during observation is 542.23° in azimuth and 91.78° in elevation. The chain of switches to limit the motion range is formed by a pre-limit to slow down the velocity, an operational limit and an emergency limit. Each axis has two chains, one for each direction of the motion. A summary of the positions of intervention of the limit switches is reported in Table 1.

During this test it has been possible to verify also the maximum velocities of the axes and the open-loop positioning error during the run. Figure 5 shows the axis performance moving at maximum velocity achieved (red slim line), the full motion range (actual in red and commanded in black) and the positioning error during the run (orange line) for both axes. The telescope is able to reach 4.5°/s in azimuth in such a way to complete a full turn in less than 80 s. Elevation reaches 2°/s and hence completes its run in less than 45 sec. The motion shows also a good positioning precision despite the open-loop configuration. In fact, the azimuth error is limited to about 2 arcmin while the elevation error remains below ±3.5 arcmin.

**Table 1:** Limit switches intervention position for the clock-wise (CW) and counter clock-wise (CCW) azimuth directions and for zenith (up-ward) and horizon (down-ward) elevation directions.

Azimuth		Elevation	
CW pre-limit	258.56°	Up-ward per limit	87.12°
CW operational limit	271.10°	Up-ward operational limit	91.22°
CW emergency limit	273.40°	Up-ward emergency limit	92.24°
CCW pre-limit	-261.04°	Down-ward per limit	3.77°
CCW operational limit	-271.13°	Down-ward operational limit	-0.56°
CCW emergency limit	-273.65°	Down-ward emergency limit	-1.30°
--	--	Down-ward bumpers	-1.12°

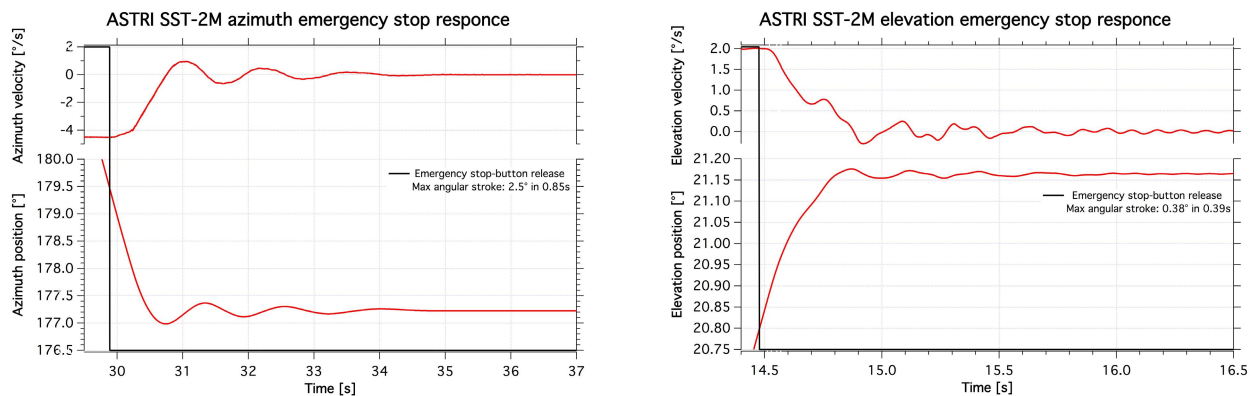


**Figure 5:** Full range movement for azimuth (left) and elevation (right) axes. The plots show the maximum velocities achieved (red slim line on top), the motion range (red and black lines) and the positioning errors during the run (orange line).

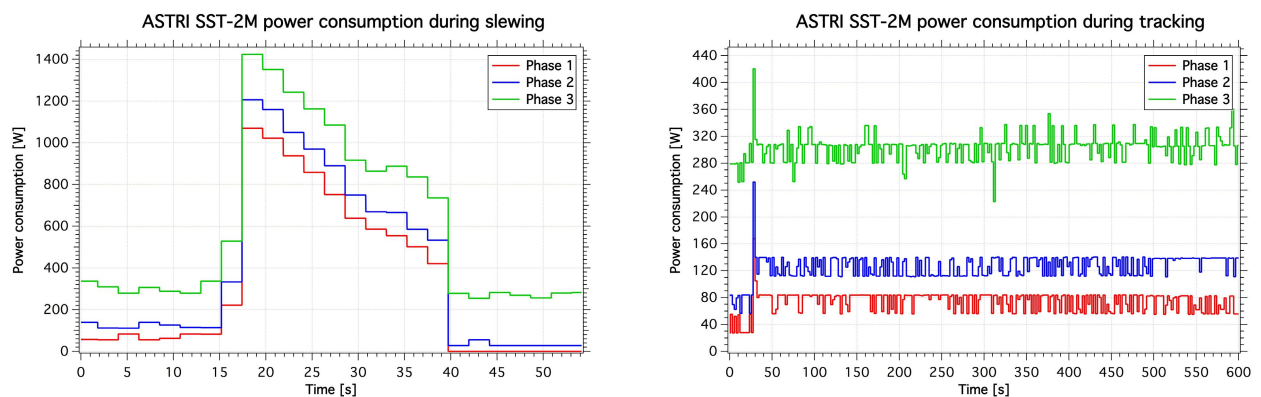
The emergency movement is done by means of an electrical drill connected to a second shaft of both axes motors. A test was performed in order to evaluate the maximum velocities achievable and hence the maximum time needed to park the telescope in case of an emergency situation that implies a failure of the motion control system. During the test, the motion of each axis has been verified on both directions. The results show a maximum velocity of  $1.5^{\circ}/s$  for the azimuth axis and  $0.4^{\circ}/s$  for the elevation. Less than 4 minutes are needed to drive the telescope to its parking position with the emergency toolkit.

The emergency braking range has been evaluated after the push of the red safety button during a fast slewing movement. The test has been repeated for both axes. The test results are shown in Figure 6. In azimuth the telescope brakes in  $2.3^{\circ}$  within  $0.61^{\circ}$  then few oscillations are triggered. These oscillations have a maximum amplitude of  $0.2^{\circ}$  and are completely dumped within 4 s. Similarly, in elevation the telescope brakes in  $0.37^{\circ}$  within  $0.34^{\circ}$  then few oscillations are also triggered. These oscillations have a maximum amplitude of  $0.01^{\circ}$  and are completely dumped in less than 1.5 s.

An estimation of the power consumption of the telescope has been also verified both during a fast slewing and during tracking. Figure 7 shows the cumulative consumption (of each phase of the power supply of the telescope) for the motors and drives system, the motion control and the safety units. The telescope requests a peak of power of about 3.8 kW during the fast slewing while it keeps its consumption to about 0.5 kW during the tracking phase.

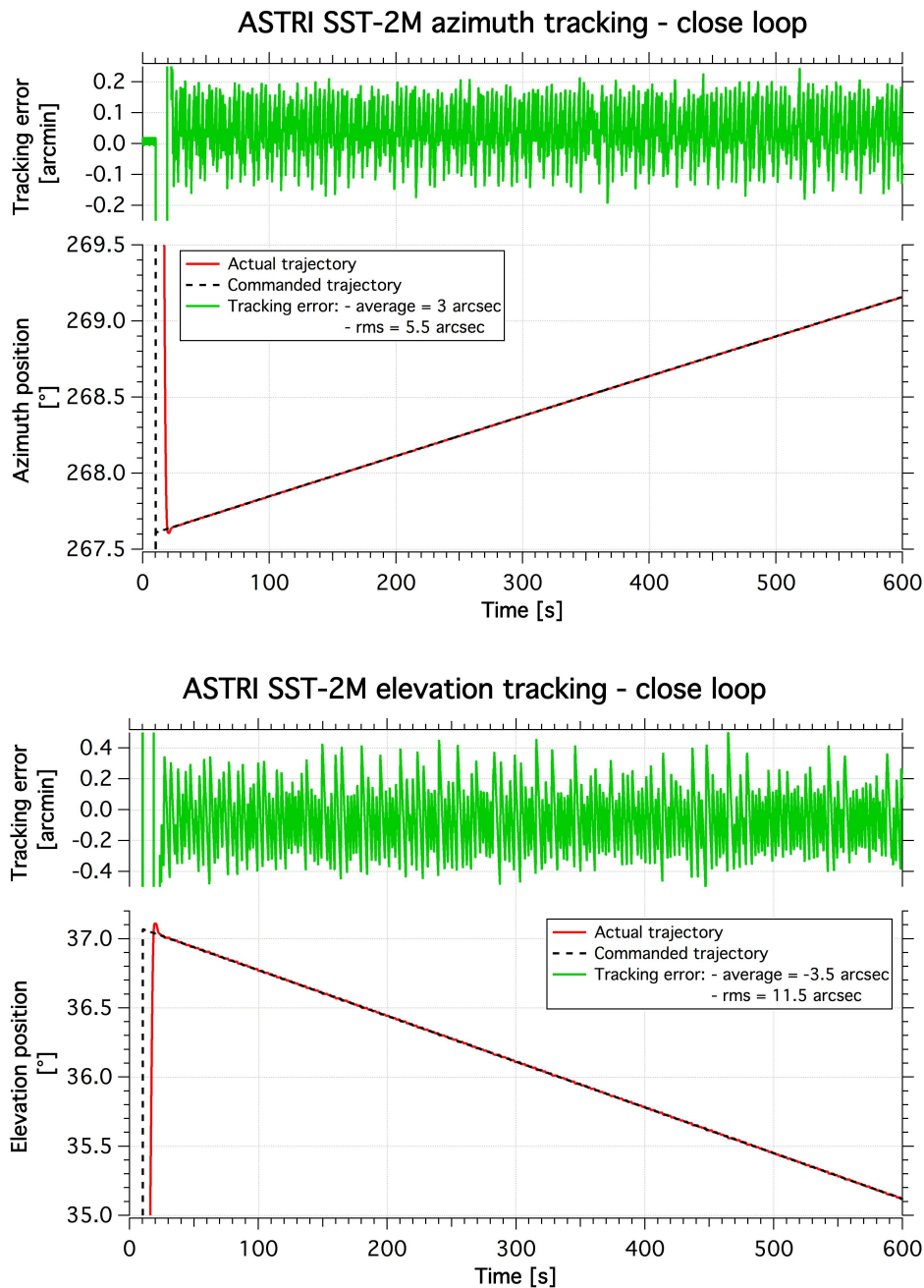


**Figure 6:** Behavior of the telescope upon an emergency stop. After the emergency-stop button intervention the telescope motion gets a complete stop within 1 s for both axis. Oscillations are triggered and quickly dumped. (left) Azimuth axis. (right) Elevation axis.



**Figure 7:** Power consumption of the ASTRI SST-2M telescope during a fast slewing (left) and tracking (right) operation.

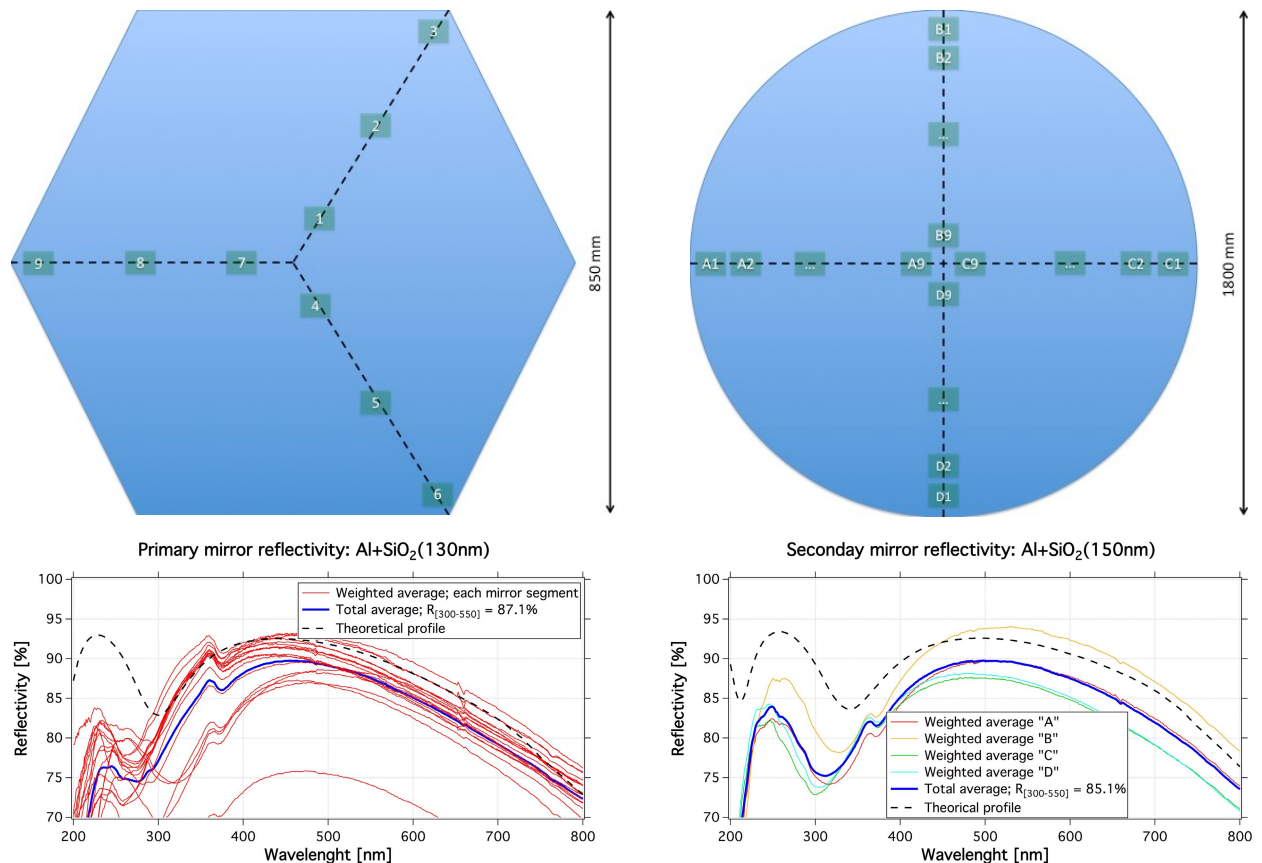
The opto-mechanical performance has been evaluated in terms of positioning and tracking errors. Concerning the tracking errors, the trajectory's coordinates generated by the telescope motion controller and the actual position of the telescope retrieved by the encoders reading during the run have been acquired. The tracking error has been evaluated as the rms differences of the two trajectories. The test has been repeated for a number of trajectories at different azimuth and elevation positions. A typical behavior is shown in Figure 8; the tracking error is of the order of 5.5 arcsec in azimuth and 11.5 arcsec in elevation over 10 minutes run.



**Figure 8:** Tracking error test. The telescope has been commanded to follow a number of trajectories. The commanded and actual paths (encoders reading) have been recorded (red and black dashed lines). The difference between the two (green line) is the tracking error: (left) azimuth tracking error and (right) elevation tracking error.

Optical tests so far performed at the telescope have concerned the monitoring of the eventual reflectivity degradation underwent by the primary and secondary mirrors coatings.

Local specular reflectivity has been measured with a spectrophotometer from Filmetrics, model F20-UVX equipped with an additional lamp model LS-DT2 Tungsten-Halogen and a CP-1 contact probe. Reflectance measurements have been taken in the spectral range 200-1700 nm and after 6 months of environmental ageing (winter season). Each segment of the primary mirror has been sampled over 9 points distributed on its surface. Instead, the secondary mirror was measured over 36 points. The measurement positions for both mirrors are shown in the top row of Figure 9. The bottom row shows the typical specular reflectivity curves measured at the telescope. Some differences with respect to the nominal profiles (black dashed lines are the simulated reflectivity curves) are detected. These can be ascribed both to deficiencies of deposition process and ageing due to environmental conditions. In any case, both coatings respect the reflectivity requirement imposed by the project.



**Figure 9:** Results of the reflectivity monitoring after 6 months of on-site ageing. (top-left) Measuring points on the primary mirror segments. (top-right) Measuring points on the secondary mirror. (bottom-left) Typical reflectivity curves for the primary mirror. (bottom-right) Typical reflectivity curves for the secondary mirror.

#### 4. CONCLUSIONS

This paper presents the on-site assembling and preliminary tests results of the ASTRI SST-2M telescope, prototype of the Small Size class of Telescopes with a Dual-mirror configuration for the Cherenkov Telescope Array.

The structural and mechanical designs of the telescope have been committed to experienced companies in the field of astronomical instrumentation. Same for the manufacturing of the electro-mechanical structure; it has been subcontracted to a consortium of top-level Italian companies.

All the structural and mechanical components of the telescope have been manufactured and preassembled on factory during the timeframe of February-August 2014. Shipping to the installation site was done in few days in August of the same year. The full assembly, including also the electrical cabling and the mirror installation, has been carried out in about three weeks of work between August and September 2014. The telescope structure has been officially inaugurated on September the 24<sup>th</sup>.

Movement (including axes close-loop tuning), safety and operation procedures have been developed and debugged during the following months as well as a number of electrical issues have been addressed.

Some tests involving the main functionalities of the telescope have been started. Preliminary results are reported in this proceeding. Even if some important tests have not yet been performed (such as the Eigen frequencies evaluation) the commissioning activities completed up to now yield a very positive behavior of the telescope. All the functionalities tested such as maximum axes range, maximum velocity, emergency braking range and emergency velocity movement are well within the requirements. Power consumption is inside the allocated budget and tracking accuracy seems well below the requirement with a good match with respect to the dynamic simulations.

Further tests and commissioning activities are ongoing and will last a few months.

## ACKNOWLEDGMENTS

This work was partially supported by the ASTRI "Flagship Project" financed by the Italian Ministry of Education, University, and Research (MIUR) and led by the Italian National Institute of Astrophysics (INAF). We acknowledge partial support by the MIUR Bando PRIN 2009 and TeChe.it 2014 Special Grants. We also acknowledge support from the Brazilian Funding Agency FAPESP (Grant 2013/10559-5) and from the South African Department of Science and Technology through Funding Agreement 0227/2014 for the South African Gamma-Ray Astronomy Programme.

We gratefully acknowledge support from the agencies and organizations listed under Funding Agencies at this website: <http://www.cta-observatory.org/>.

We also acknowledge the valuable collaboration of the consortium GEC composed by the companies EIE Group and Galbiati Group.

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