



<b>Publication Year</b>	2019
<b>Acceptance in OA</b>	2020-11-25T09:19:14Z
<b>Title</b>	Experimental characterization of modal noise in multimode fibers for astronomical spectrometers
<b>Authors</b>	OLIVA, Ernesto, RAINER, Monica, TOZZI, Andrea, SANNA, Nicoletta, Iuzzolino, M., BRUCALASSI, ANNA
<b>Publisher's version (DOI)</b>	10.1051/0004-6361/201936610
<b>Handle</b>	<a href="http://hdl.handle.net/20.500.12386/28526">http://hdl.handle.net/20.500.12386/28526</a>
<b>Journal</b>	ASTRONOMY & ASTROPHYSICS
<b>Volume</b>	632

# Experimental characterization of modal noise in multimode fibers for astronomical spectrometers

E. Oliva<sup>1</sup>, M. Rainer<sup>1</sup>, A. Tozzi<sup>1</sup>, N. Sanna<sup>1</sup>, M. Iuzzolino<sup>2</sup>, and A. Brucalassi<sup>1</sup>

<sup>1</sup> INAF – Osservatorio Astrofisico di Arcetri, Largo E. Fermi 5, 50125 Firenze, Italy  
e-mail: [oliva@arcetri.astro.it](mailto:oliva@arcetri.astro.it)

<sup>2</sup> Officina Stellare S.P.A., via delle Tecnica 8789, 36030 Sarcedo, VI, Italy

Received 30 August 2019 / Accepted 10 October 2019

## ABSTRACT

**Context.** High resolution spectroscopy at high signal-to-noise ratios (S/Ns) is one of the key techniques of the quantitative study of the atmospheres of extrasolar planets. Observations at near-infrared wavelengths with fiber-fed spectrographs coupled to extremely large telescopes are particularly important to tackle the ultimate goal of detecting biosignatures in rocky planets.

**Aims.** To achieve high S/Ns in fiber-fed spectrographs, the systematic noise effects introduced by the fibers must be properly understood and mitigated. In this paper we concentrate on the effects of modal noise in multimode fibers.

**Methods.** Starting from our puzzling on-sky experience with the GIANO-TNG spectrometer we set up an infrared high resolution spectrometer in our laboratory and used this instrument to characterize the modal noise generated in fibers of different types (circular and octagonal) and sizes. Our experiment includes two conventional scrambling systems for fibers: a mechanical agitator and an optical double scrambler.

**Results.** We find that the strength of the modal noise primarily depends on how the fiber is illuminated. It dramatically increases when the fiber is under-illuminated, either in the near field or in the far field. The modal noise is similar in circular and octagonal fibers. The Fourier spectrum of the noise decreases exponentially with frequency; i.e., the modal noise is not white but favors broad spectral features. Using the optical double scrambler has no effect on modal noise. The mechanical agitator has effects that vary between different types of fibers and input illuminations. In some cases this agitator has virtually no effect. In other cases, it mitigates the modal noise, but flattens the noise spectrum in Fourier space; i.e., the mechanical agitator preferentially filters the broad spectral features.

**Conclusions.** Our results show that modal noise is frustratingly insensitive to the use of octagonal fibers and optical double scramblers; i.e., the conventional systems used to improve the performances of spectrographs fed via unevenly illuminated fibers. Fiber agitation may help in some cases, but its effect has to be verified on a case-by-case basis. More generally, our results indicate that the design of the fiber link feeding a spectrograph should be coupled with laboratory measurements that reproduce, as closely as possible, the conditions expected at the telescope.

**Key words.** instrumentation: spectrographs – techniques: miscellaneous – techniques: spectroscopic

## 1. Introduction

Quantitative spectroscopy of the atmospheres of extrasolar planets is one of the key science cases of modern astrophysics. High resolution spectrographs coupled to extremely large telescopes are particularly important to achieve the ultimate goal of detecting biosignatures in the atmospheres of earth-like planets (e.g., [Maiolino et al. 2013](#)). There are two main techniques: transit spectroscopy, which is based on a time series of high resolution stellar spectra taken during the transit of a planet (e.g., [Brogi et al. 2012](#)); and direct light detection, which is based on simultaneous high resolution spectra taken through an integral field unit (IFU) that includes the star and surrounding exoplanet(s). The contrast between the planet and the stellar signals can be boosted using adaptive optics correction (e.g., [Snellen et al. 2015](#)). In either case the spectrum of the exoplanet atmosphere consists of a large number of narrow ( $\approx$  few  $\text{km s}^{-1}$ ) lines with amplitudes orders of magnitude weaker than the signal from the star. For earth-like planets orbiting  $M$  dwarfs the amplitude of the planet spectrum is only a few  $\times 10^{-5}$  of the stellar continuum. Cross-correlation techniques are used to identify the exoplanet spectrum taking advantage of the Doppler shift introduced by the orbital motion of the planet.

These observations are normally simulated assuming that the spectral data are only affected by detector-related noises (readout and dark-current noise) and photon noise. Recently, [Hawker & Parry \(2019\)](#) investigated the effects systematic extra noise introduced by a non-perfect correction of the telluric features. Other systematic effects introduced by the instrument are evaluated based on practical experience with existing spectrometers. No detailed attempt is made to identify the technical aspects of the instruments that are critical, useful, or just redundant for this type of measurements.

In this paper we concentrate on the effects of modal noise in spectrographs fed via multimode fibers. This phenomenon is related to the interference between the light modes that propagate through a fiber. Inside a spectrometer, the interference pattern translates into a non-flat, wavy spectral distribution. Noise is generated whenever the interference pattern changes between different exposure; i.e., when the wavy pattern cannot be corrected using flat-field calibrations.

From the theoretical point of view the available literature on modal noise can be classified into two main groups. The first group does not consider how the fiber-input is illuminated. For example, [Grupp \(2003\)](#), [Chen et al. \(2006\)](#), and [Lemke et al. \(2011\)](#) modeled the effect of modal noise

introduced by vignetting occurring in the spectrometer, downstream of the fiber. These authors unanimously concluded that the effect can be averaged out and effectively mitigated by continuously agitating or stretching the fiber. The second group mainly concentrates on how the fiber is illuminated. The most remarkable example is that of [Corbett & Allington-Smith \(2006\)](#) who argued that the abovementioned theories must be called into question because of the changes in the statistics of modal noise with arbitrary illumination of the fiber. These authors also mentioned that fiber agitation may have different effects, depending on how the fiber is illuminated. Other instructive examples are those of [Mahadevan et al. \(2014\)](#) and [Halverson et al. \(2014\)](#) who studied the best strategies to illuminate the fibers from the calibration sources, while [Blind et al. \(2017\)](#) studied the effects of a nonuniform illumination of the fiber carrying the stellar light.

From the experimental point of view there are many papers that concentrate on the importance of mechanical agitation of the fiber, from the earliest work of [Baudrand & Walker \(2001\)](#) to the comprehensive characterization of different electromechanical devices of [Petersburg et al. \(2018\)](#). Mechanical agitators are often included in the fiber links for the calibration lights; while they are only rarely used for the fibers carrying the stellar light from the telescope. The spectrographs HARPS and ESPRESSO of the European Southern Observatory (ESO) are the most remarkable examples of spectrometers that do not shake the science fibers. A similar approach is also foreseen for the Near Infrared Planet Searcher (NIRPS) under construction for the ESO 3.6 m telescope (see [Blind et al. 2017](#)). Other types of “static scramblers” (i.e., octagonal fibers and optical double scramblers) are favored in these instruments, even though their effects on modal noise are virtually unknown. Curiously enough, [Mahadevan et al. \(2014\)](#) argued that these static devices have no useful effect on modal noise generated in fibers carrying coherent light from calibration sources.

In this paper we address the questions of under which circumstances modal noise is important, how modal noise affects observations that require very high signal-to-noise ratios (S/Ns), and which technical solutions are useful to mitigate the effect of modal noise. We start describing in Sect. 2 our puzzling experience with data collected with the infrared spectrometer GIANO ([Origlia et al. 2014](#)) at the Telescopio Nazionale *Galileo* (TNG) of the Roque de Los Muchachos Observatory in La Palma. We then describe the laboratory spectrometer that we specifically built to study the phenomenon (Sect. 3). The results are presented in Sect. 4. In Sect. 5 we discuss the results and draw our conclusions.

## 2. Modal noise in GIANO spectra with fibers

The spectrograph GIANO is the high resolution ( $R \approx 50\,000$ ) infrared (950–2450 nm) spectrometer of TNG ([Origlia et al. 2014](#)). The GIANO instrument was designed and built for direct feeding of light at a dedicated focal station. In 2012, after a reorganization of the focal stations at the TNG, the instrument was provisionally commissioned and used in the GIANO-A configuration: with the spectrometer positioned on the rotating building and fed via a pair of  $\varnothing 85\,\mu\text{m}$  fibers connected to the only available focal station ([Tozzi et al. 2014](#)). In 2016 the spectrometer was eventually moved to the originally foreseen configuration (called GIANO-B), in which the instrument is directly fed from the telescope (i.e., no fibers). This configuration can also be used for simultaneous observations with HARPS-N: the so-called GIARPS mode ([Tozzi et al. 2016](#)).

The fiber link used in the provisional GIANO-A configuration could not be fully optimized for astronomical observations. One of the main limitations was modal noise. This was already evident in the very first acquisitions of calibration frames: flats taken at different times showed variable, wavy spectral features. To cure this problem we introduced a mechanical agitator in the fiber link carrying the light to the spectrometer. This solved the problem of the calibration light: comparisons of various groups of flat exposures always yielded noise values compatible with the theoretical photon noise, even for S/Ns exceeding 3000. However, much to our surprise, the mechanical agitator did not improve the quality of the stellar spectra, even though the scientific and calibration lights went through the same fibers and mechanical agitator. All our attempts to improve the situation by modifying the agitation mechanisms were unsuccessful.

The effect of modal noise is best visible in data aimed at detecting exoplanet features during transits. The observations consist of a continuous series of spectra of the same star with the instrument in the same configuration ([Snellen et al. 2010](#)). The spectra are organized in 2D frames with the wavelength on the  $X$ -axis and time on the  $Y$ -axis. The “Step 1” images are made by combining the normalized spectra: they include the lines intrinsic to the star (stable over time) and the telluric absorption features, whose depths vary with time and air mass. Both families of lines are removed (to a first approximation) by a linear fit with air mass of each wavelength channel: this generates the “Step 2” frames.

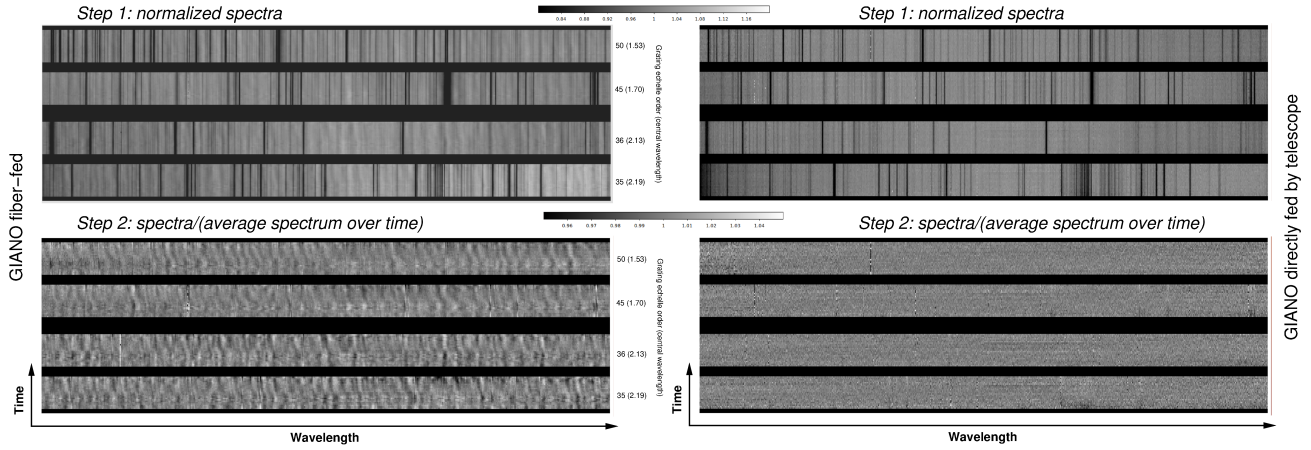
The data shown in Fig. 1 are an instructive comparison of such spectral series acquired using the GIANO spectrometer with and without fibers. The irregular, wavy pattern in the data taken with fibers is the effect of modal noise. The peak-to-valley amplitude of these features is about 10% of the continuum level. In spite of this large noise, these data were successfully used by [Brogi et al. \(2018\)](#) to detect water in the HD 189733b. Data with a similar pattern of modal noise were used to detect methane in HD 102195b ([Guilluy et al. 2019](#)). These results are encouraging because they demonstrate that the signal of the exoplanet atmosphere can be detected even in situations in which the modal noise strongly limits the final quality of the data. More specifically, these results indicate that modal noise primarily produces broad features, while it does not seem to affect the measurement of narrow (unresolved) lines such as those produced in the atmosphere of exoplanets.

## 3. Characterizing modal noise in the laboratory

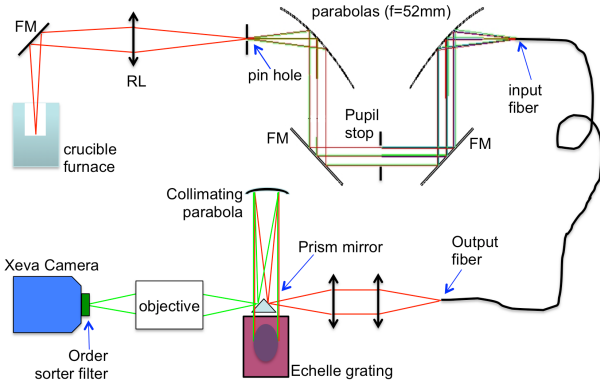
Given the puzzling results obtained at the telescope, we decided to investigate the origin and effects of modal noise using a dedicated spectrometer in the laboratory. The instrumental setup is summarized in Fig. 2; see also [Iuzzolino et al. \(2014, 2015\)](#). The light source is a commercial furnace at 1150 °C. This somewhat inconvenient device is the only light source that guarantees a sufficiently bright and uniform illumination both in the near field and far field.

The light is injected into the fiber via a 1:1 optical relay with a field of view of 2 mm and a focal aperture of  $F/2$ ; i.e., adequate for a complete filling of all the fiber modes. The near-field illumination of the fiber can be limited by inserting a field stop (pinhole) at the entrance, while the far-field illumination can be regulated using a pupil stop (iris) in the collimated beam. Both stops can be moved laterally to modify the illumination pattern of the fiber input.

The fiber exit feeds a high resolution spectrometer that uses the spare grating of GIANO: a R2 echelle with 23.2 lines  $\text{mm}^{-1}$ .



**Fig. 1.** Times-series spectra of stars collected with the GIANO spectrometer. *Left- and right-hand panels:* data collected with and without fiber feeding, respectively. Modal noise causes the wavy pattern clearly visible in the data taken with fibers.



**Fig. 2.** Schematic layout of the spectrometer used for the characterization of the modal noise.

The grating angle is adjusted to obtain a central wavelength of 1600 nm at order 48. The light in the other orders is blocked by a narrowband filter positioned in front of the detector. The detector is a commercial InGaAs camera with  $320 \times 256$  pixels of 30 microns size. The spectral dispersion is  $0.032 \text{ nm pixel}^{-1}$ .

Optionally, the fiber can be directed through a mechanical agitator that vigorously shakes a 30 cm section of the fiber at about 20 Hz. The other optional subsystem is an optical double scrambler that couples the light between two identical fibers, reverting the near and far fields. This device was originally designed for the  $\varnothing 85$  microns fibers of GIANO-A.

The spectra were collected with three types of fibers as follows: first, fused silica fibers with circular core  $\varnothing 85 \mu\text{m}$  (from the same batch used for the WEAVE instrument); second, fused silica fibers with octagonal core with  $67 \mu\text{m}$  pitch (from the stock catalog of Ceramoptec); and third, fused silica fibers with circular core  $\varnothing 50 \mu\text{m}$  (from the catalog of Thorlabs). In all cases, the fibers are mounted in 8 m long patches with commercial SMA connectors. Two such patches were alternatively used for each type of fiber. The results obtained with the two patches are indistinguishable.

#### 4. Results from laboratory spectra

The laboratory spectra were organized in continuous time series of 1200 exposures each with 0.5 s integration time. The spectra were analyzed using a technique similar to that employed for

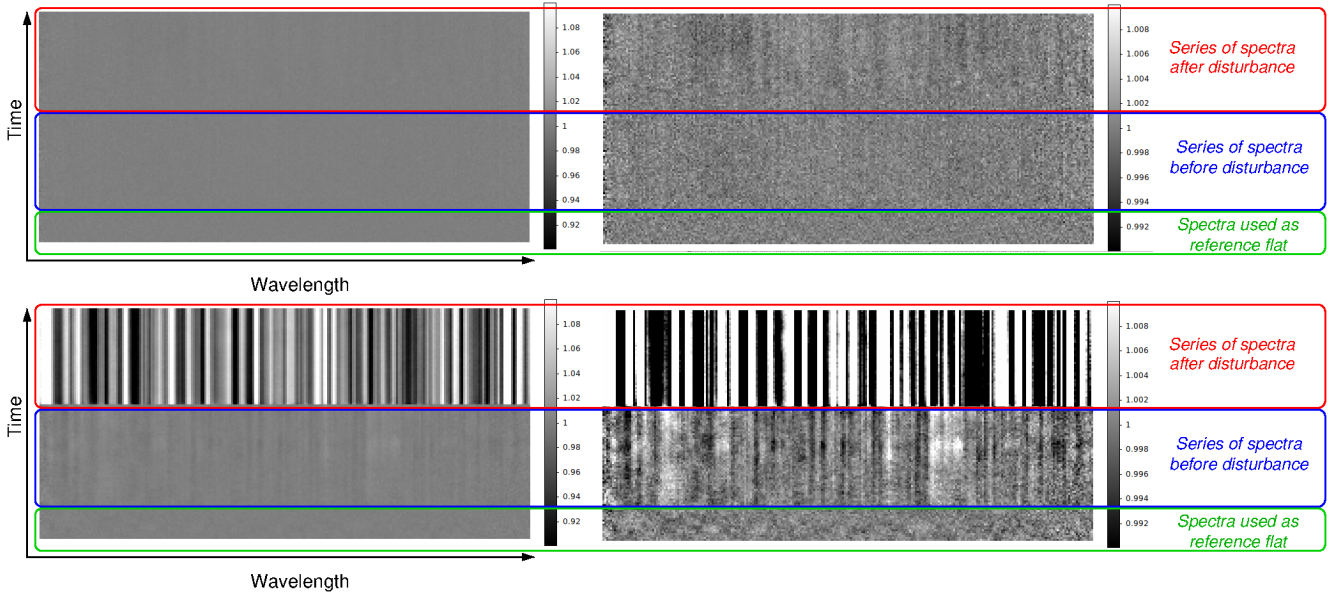
the transits of exoplanets. The average of the first 200 spectra of the series were used as reference flat. All the spectra of the series were divided by the reference flat and renormalized (i.e., divided by their median value). Therefore, they are equivalent to the Step 2 data described in Sect. 2 and Fig. 1.

We took the first 700 spectra with the system in a undisturbed configuration, thereby avoiding in particular any movement of the fiber. At this point, we intentionally generated a quick disturbance in the system. Therefore, the effect of disturbance was recorded in the remaining 500 spectra of the series. There are three types of disturbance as follows:

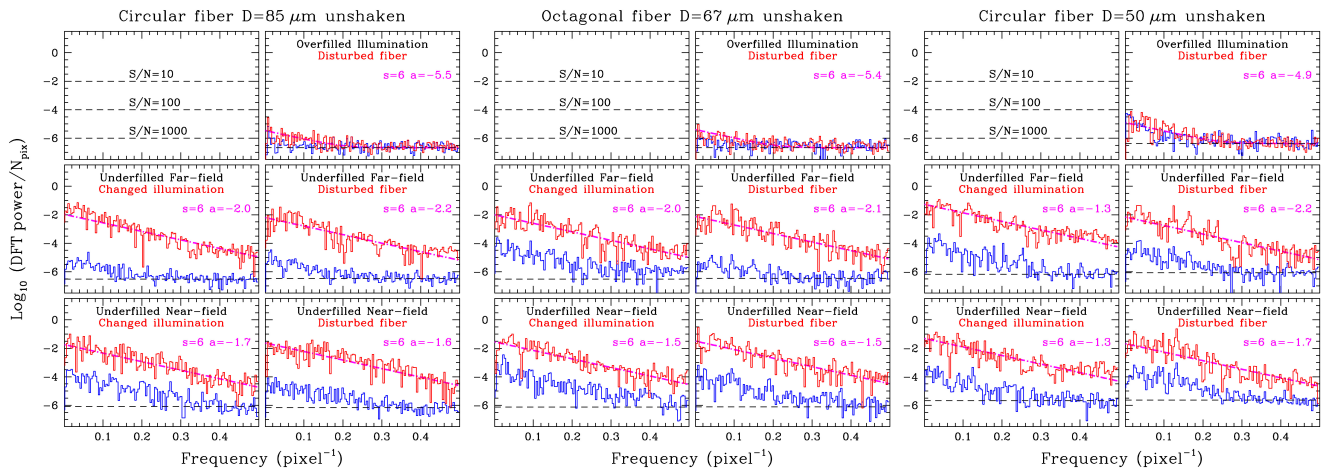
1. Change of input illumination with the fiber under-illuminated in the near field. A pinhole of 10 microns is used as field stop. The disturbance consists of laterally moving the pinhole by 20 microns.
2. Change of input illumination with the fiber under-illuminated in the far-field. A pupil stop on the collimated beam limits the focal aperture to  $F/15$ . The disturbance consists of laterally moving the pupil stop by twice its diameter.
3. Mechanical disturbance to the fiber. This consists of gently lifting and laying back a portion of the fiber on the optical bench. This disturbance is applied for any type of fiber illumination; i.e., under-illuminated and overfilled.

Figure 3 shows representative examples of such 2D spectra taken under different conditions. The top spectrum is taken with overfilled illumination of the fiber. This spectrum is dominated by white noise and only shows a very weak pattern of extra noise after the disturbance. The bottom spectrum is taken with the fiber under-illuminated in the near field. This spectrum is characterized by a very prominent modal noise after the disturbance and also shows a weak pattern of modal noise in the spectra taken before the disturbance. Most probably, this is caused by tiny thermo-mechanical instabilities within the fiber. In all spectra the modal noise seems to be dominated by low spatial frequencies, i.e., relatively broad features.

A quantitative analysis of the modal noise can be conveniently performed by computing the discrete Fourier transform (DFT) of the extracted 1D spectra before and after the disturbance. Representative results are shown in Figs. 4–6. These plots show the variation of the logarithm of the average DFT power with spatial frequency ( $\text{pixel}^{-1}$ ). In such a diagram, spectra dominated by white noise would appear as horizontal lines with constant value equal to the inverse square of the S/Ns, as indicated in the top left insert of the figures.



**Fig. 3.** Example of 2D times-series spectra with different illuminations of the fiber. *Top panel:* uniform illumination (overfilled). *Bottom panel:* under-filled illumination. *Left- and right-hand panels:* same data with different gray scale limits:  $\pm 10\%$  and  $\pm 1\%$ , respectively.



**Fig. 4.** Discrete Fourier transform spectral analysis of data taken with different fibers without mechanical agitator. See text (Sect. 4) for details

We first concentrate on the spectra shown in the bottom left panel of Fig. 4. These spectra refer to measurements using a  $\varnothing 85 \mu\text{m}$  circular fiber without mechanical agitation. The fiber is under-illuminated in the near field and is disturbed by moving the input illumination. The horizontal (black) dashed line at the bottom is the white noise level expected from the spectrometer alone. This is the limiting noise of the measurement and is given by

$$\text{rms}_{\text{spectrometer}} = \frac{\sqrt{2 N_{\text{frames}} (n_e + \text{RON}^2 + n_{\text{dark}})}}{N_{\text{frames}} n_e} \quad (1)$$

where  $n_e$  is the number of photoelectrons collected in each frame, RON is the readout noise,  $n_{\text{dark}}$  is the number of electrons produced by the detector dark current in each frame, and  $N_{\text{frames}}$  are the number of exposures summed to obtain the 1D spectrum. The factor of 2 is included to take into account the noise introduced by dividing for the flat-field, which also consists of the sum of  $N_{\text{frames}}$ . We verified that the limiting noise of the spectra scaled as predicted by the above relationship down to values of  $\text{rms}_{\text{spectrometer}} \approx 10^{-4}$ , equivalent to S/Ns of  $10^4$ . These

measurements were performed without fibers, i.e., with the spectrometer directly illuminated through a pinhole. The bottom (blue) histogram is the measured DFT in the undisturbed part of the spectral series. The top (red) histogram is the measured DFT in the disturbed part. The latter is fitted using an exponential function as follows:

$$\text{DFT}_{\text{total}} = \left( \text{DFT}_{\text{modal-noise}}^{-2} + \text{rms}_{\text{spectrometer}}^{-4} \right)^{-1/2} \quad (2)$$

$$\text{DFT}_{\text{modal-noise}} = 10^a \cdot 10^{-s\nu} \quad (3)$$

where  $\nu$  is the frequency ( $\text{pixel}^{-1}$ ). The fitted strength ( $a$ ), its slope ( $s$ ), and the function  $\text{DFT}_{\text{total}}$  are displayed in magenta.

The other panels of Fig. 4 show the same type of data for various types of fibers, illumination and disturbances. The plotted spectra are randomly selected among a large set of data taken under similar conditions. The numerical results are report in Table 1. All these data yield very similar results that can be summarized as follows.

– Modal noise is very weak and barely measurable when the illumination of the fiber is overfilled both in the near field and far

**Table 1.** Average values of strength ( $a$ ) and slope ( $s$ , see Eq. (2)) of the DFT spectrum of modal noise generated under different conditions.

Fiber	Double scrambler	Shaker	Under-fill	Disturbance	$a$	$s$
Circular $\varnothing 85 \mu\text{m}$	No	No	Near-field	Changed illumination	-2.0	5.9
Circular $\varnothing 85 \mu\text{m}$	No	No	Near-field	Moved fiber	-2.1	6.1
Circular $\varnothing 85 \mu\text{m}$	No	No	Far-field	Changed illumination	-1.7	6.0
Circular $\varnothing 85 \mu\text{m}$	No	No	Far-field	Moved fiber	-1.7	5.9
Circular $\varnothing 85 \mu\text{m}$	No	Yes	Near-field	Changed illumination	-3.6	4.0
Circular $\varnothing 85 \mu\text{m}$	No	Yes	Near-field	Moved fiber	-4.0	3.9
Circular $\varnothing 85 \mu\text{m}$	No	Yes	Far-field	Changed illumination	-3.8	4.1
Circular $\varnothing 85 \mu\text{m}$	No	Yes	Far-field	Moved fiber	-4.1	3.9
Circular $\varnothing 85 \mu\text{m}$	Yes	No	Near-field	Changed illumination	-1.9	6.0
Circular $\varnothing 85 \mu\text{m}$	Yes	No	Near-field	Moved fiber	-2.1	5.9
Circular $\varnothing 85 \mu\text{m}$	Yes	No	Far-field	Changed illumination	-1.6	6.1
Circular $\varnothing 85 \mu\text{m}$	Yes	No	Far-field	Moved fiber	-1.7	6.0
Circular $\varnothing 50 \mu\text{m}$	No	No	Near-field	Changed illumination	-1.5	6.0
Circular $\varnothing 50 \mu\text{m}$	No	No	Near-field	Moved fiber	-2.0	5.8
Circular $\varnothing 50 \mu\text{m}$	No	No	Far-field	Changed illumination	-1.4	6.1
Circular $\varnothing 50 \mu\text{m}$	No	No	Far-field	Moved fiber	-1.8	5.9
Circular $\varnothing 50 \mu\text{m}$	No	Yes	Near-field	Changed illumination	-2.0	5.9
Circular $\varnothing 50 \mu\text{m}$	No	Yes	Near-field	Moved fiber	-2.3	6.0
Circular $\varnothing 50 \mu\text{m}$	No	Yes	Far-field	Changed illumination	-1.7	5.8
Circular $\varnothing 50 \mu\text{m}$	No	Yes	Far-field	Moved fiber	-2.0	5.9
Octagonal $67 \mu\text{m}$	No	No	Near-field	Changed illumination	-2.1	6.1
Octagonal $67 \mu\text{m}$	No	No	Near-field	Moved fiber	-2.0	5.8
Octagonal $67 \mu\text{m}$	No	No	Far-field	Changed illumination	-1.6	5.9
Octagonal $67 \mu\text{m}$	No	No	Far-field	Moved fiber	-1.5	6.0
Octagonal $67 \mu\text{m}$	No	Yes	Near-field	Changed illumination	-2.4	4.0
Octagonal $67 \mu\text{m}$	No	Yes	Near-field	Moved fiber	-2.6	4.1
Octagonal $67 \mu\text{m}$	No	Yes	Far-field	Changed illumination	-2.4	4.0
Octagonal $67 \mu\text{m}$	No	Yes	Far-field	Moved fiber	-2.7	4.1

**Notes.** Typical scatters within each group of data are  $\pm 0.3$  for  $a$  and  $\pm 0.1$  for  $s$ . See Sect. 4 for details.

field (top right panels). Indeed, the ideal condition of overfilling is very difficult to achieve in practice; in our case it required using a furnace as light source (see Sect. 3).

– The strength of modal noise (parameter  $a$  of Eq. (3)) is similar in all the spectral series with under-illuminated fibers. The average value is  $a = -1.8 \pm 0.3$  and is the same for all fibers types and measurement setups. In particular we do not find any significant difference between octagonal and circular fibers.

– The slope of the fit to the modal noise (parameter  $s$  of Eq. (3)) is remarkably similar in all the measurements. Its average value is  $s = 6.0 \pm 0.1$  without any significant difference between different types of fibers, illumination and disturbance.

Representative spectra obtained using the mechanical agitator of the fiber are shown in Fig. 5 and the derived parameters are listed in Table 1. The results strongly depend on the type of fiber.

– In the spectra taken with the  $\varnothing 85 \mu\text{m}$  fibers, the strength of the modal noise decreases by about two orders of magnitudes. The slope of the DFT distribution becomes shallower with  $s = 4.0 \pm 0.1$ .

– In the spectra taken with the  $67 \mu\text{m}$  octagonal fibers, the strength of the modal noise decreases by only 0.7 dex and its slope decreases to  $s = 4.0 \pm 0.1$ , i.e., the same flattening seen with the larger circular fiber.

– In the spectra taken with the  $\varnothing 50 \mu\text{m}$  fibers, the average strength and the slope of the modal noise are similar to those observed without the mechanical agitator.

Noticeably, any decrease of modal noise is coupled with a flattening of its DFT noise spectrum. In other words, mechanical

agitation (when it works) preferentially filters the broadest features of the modal noise.

Representative spectra obtained with an optical double scrambler are shown in Fig. 6 and the derived parameters are listed in Table 1. The strength and slope of the noise spectrum are identical within the errors to the measurements without this device (Fig 4). This indicates that the double scrambler has no effect on the modal noise.

## 5. Discussion and conclusions

To the best of our knowledge, the data presented in this paper are the first systematic attempt to characterize modal noise in the laboratory using multimode fibers feeding a high resolution spectrometer similar to those designed (or built) for astronomical observations at large telescopes. To emphasize the effects, the data are collected using unusually long wavelengths, relatively small fiber sizes, and somewhat extreme conditions for the input illumination of the fiber. More specifically, the number of propagating modes in our experiments vary between  $\approx 10$  and 700, while most fiber-fed astronomical spectrographs work with several thousands modes. The most important results are as follows.

First, modal noise primarily depends on how the fiber is illuminated. This is in line to what postulated by [Corbett & Allington-Smith \(2006\)](#).

Second, octagonal fibers and optical double scramblers do not seem to have any effect on modal noise. This is in line with what was already suggested by [Mahadevan et al. \(2014\)](#).

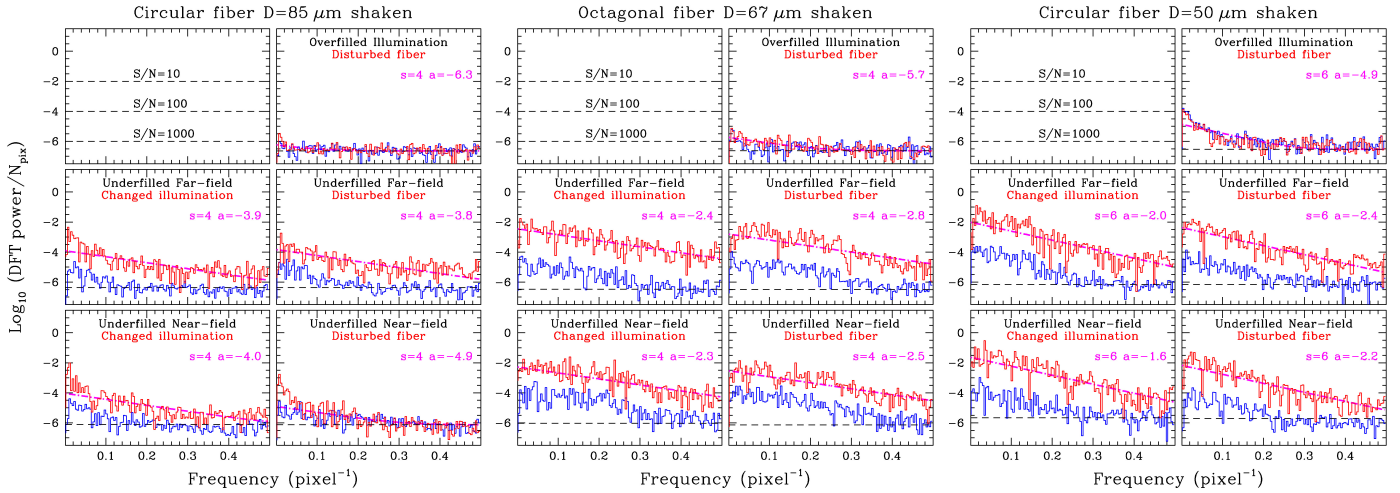


Fig. 5. Same as Fig. 4 for spectra with mechanical agitation of the fiber.

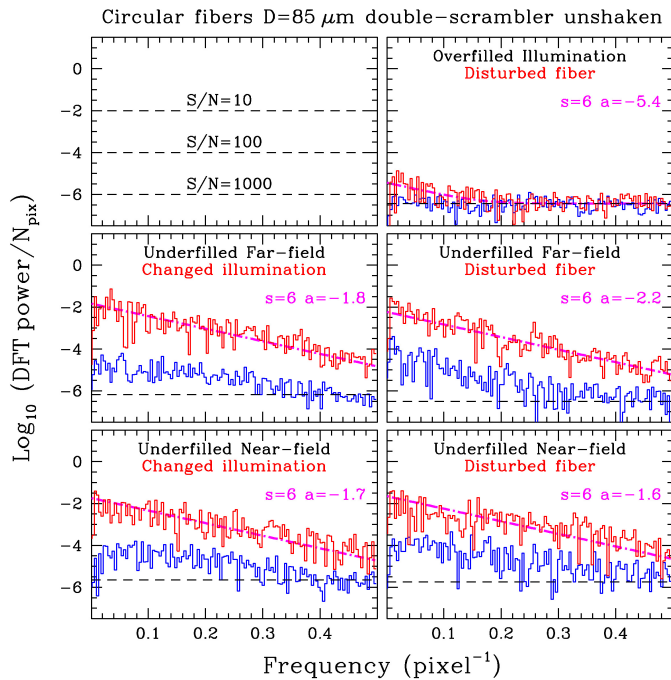


Fig. 6. Same as Fig. 4 for spectra with the optical double scrambler inserted.

It implies that well-known technical solutions, employed with great success in current instruments optimized for accurate radial velocity measurements, are not necessarily useful in spectrographs that work in a parameter space in which modal noise is much more critical.

Third, mechanical agitation of the fiber may help to decrease the strength of the modal noise. However, shaking has a remarkably different effects on fibers of different sizes and types. Moreover, when it works it flattens the DFT noise spectrum; i.e., it preferentially filters broad spectral features. Therefore, any trade-off study on mechanical agitators should include a quantitative analysis of the DFT noise spectrum.

Fourth, modal noise can be decreased below the required S/Ns by uniformly illuminating the fiber both in the near field and far field. However, this is a very difficult condition

to obtain in practice even for calibration sources (see, e.g., Mahadevan et al. 2014). For the science fibers coupled to the telescope we have to find a reasonable compromise between the total throughput and the modal noise generated by the various effects that contribute to under-illumination (e.g., telescope central obscuration, and point spread function variations). This task should ideally be performed in the laboratory using a realistic simulator of the telescope, fiber link, and spectrometer.

*Acknowledgements.* We thank the anonymous referee for helpful comments and suggestions. We are grateful to Isabelle Guinourd for providing us with the  $\varnothing 85 \mu\text{m}$  fibers. This research has been partly supported by INAF through the TECNO-INAF 2011 and “FRONTIERA-2016” competitive projects. Based on observations made with the GIANO spectrometer at the Italian Telescopio Nazionale Galileo (TNG) operated on the island of La Palma by the Fundación Galileo Galilei of the INAF at the Spanish Observatorio del Roque de los Muchachos of the Instituto de Astrofísica de Canarias.

## References

- Baudrand, J., & Walker, G. A. H. 2001, *PASP*, **113**, 851  
 Blind, N., Conod, U., & Wildi, F. 2017, Proc. AO4ELT5 Conf. [arXiv:1711.00835]  
 Brogi, M., Snellen, I. A. G., de Kok, R. J., et al. 2012, *Nature*, **486**, 502  
 Brogi, M., Giacobbe, P., Guilluy, G., et al. 2018, *A&A*, **615**, A16  
 Chen, C.-H., Reynolds, R. O., & Kost, A. 2006, *Appl. Opt.*, **45**, 519  
 Corbett, J., & Allington-Smith, J. 2006, *SPIE Conf. Ser.*, **6269**, 62693N  
 Grupp, F. 2003, *A&A*, **412**, 897  
 Guilluy, G., Sozzetti, A., Brogi, M., et al. 2019, *A&A*, **625**, A107  
 Halverson, S., Mahadevan, S., Ramsey, L., et al. 2014, *SPIE Conf. Ser.*, **9147**, 91477Z  
 Hawker, G. A., & Parry, I. R. 2019, *MNRAS*, **484**, 4855  
 Iuzzolino, M., Tozzi, A., Sanna, N., Zangrilli, L., & Oliva, E. 2014, *SPIE Conf. Ser.*, **9147**, 914766  
 Iuzzolino, M., Sanna, N., Tozzi, A., & Oliva, E. 2015, *Mem. Soc. Astron. Ital.*, **86**, 490  
 Lemke, U., Corbett, J., Allington-Smith, J., & Murray, G. 2011, *MNRAS*, **417**, 689  
 Mahadevan, S., Halverson, S., Ramsey, L., & Venditti, N. 2014, *ApJ*, **786**, 18  
 Maiolino, R., Haehnelt, M., Murphy, M. T., et al. 2013, ArXiv e-prints [arXiv:1310.3163]  
 Origlia, L., Oliva, E., Baffa, C., et al. 2014, *SPIE Conf. Ser.*, **9147**, 91471E  
 Petersburg, R. R., McCracken, T. M., Eggerman, D., et al. 2018, *ApJ*, **853**, 181  
 Snellen, I. A. G., de Kok, R. J., de Mooij, E. J. W., & Albrecht, S. 2010, *Nature*, **465**, 1049  
 Snellen, I., de Kok, R., Birkby, J. L., et al. 2015, *A&A*, **576**, A59  
 Tozzi, A., Oliva, E., Origlia, L., et al. 2014, *SPIE Conf. Ser.*, **9147**, 91479N  
 Tozzi, A., Oliva, E., Iuzzolino, M., et al. 2016, *SPIE Conf. Ser.*, **9908**, 99086C