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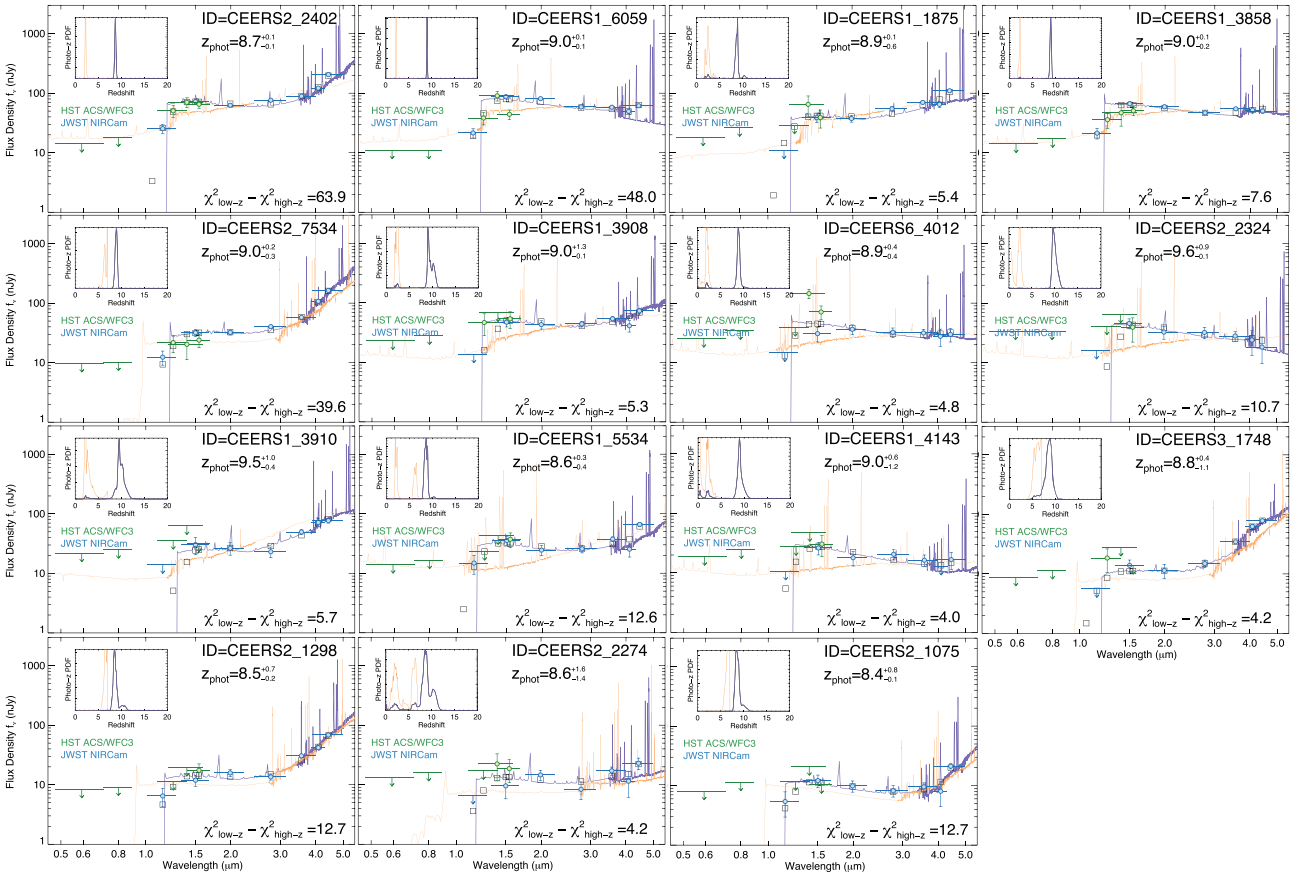

Figure 18. The SEDs and $\mathcal{P}(z)$s of the $z \sim 9$ sample.

Appendix C
Sources Removed from the Sample

Figure 19 shows two objects removed from our sample after remeasuring the colors in smaller apertures due to the stretching of their Kron apertures by nearby bright sources. Figure 20 shows cutout images for the 36 sources removed as spurious sources following the visual inspection described in Section 4.3. We list the coordinates for all removed sources in Table 6.

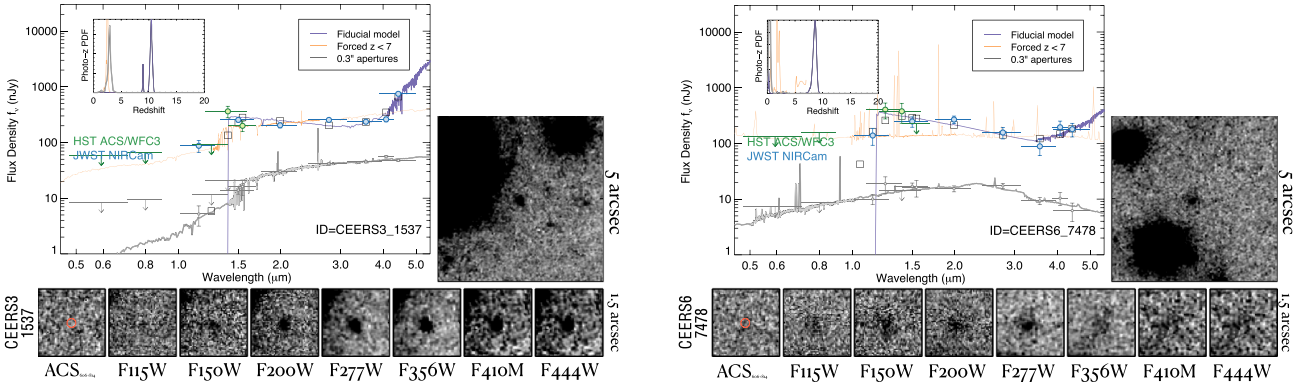

Figure 19. Two sources that had incorrectly drawn Kron apertures due to the presence of nearby bright sources. The large plot shows the SED, with the Kron NIRCam (HST) photometry in blue (green). The best-fit EAZY model to these data is shown in purple (with squares denoting the model bandpass-averaged fluxes), with the best-fit $z < 7$ solution shown in orange. The small gray circles show the fluxes measured in $0\farcs3$ diameter apertures, with the gray line showing the EAZY fit to these small-aperture fluxes. The $P(z)$s for all three EAZY runs are shown in the top left. The large image shows a $5'' \times 5''$ cutout around each source, highlighting the nearby neighbor responsible for stretching the Kron aperture, leading to much brighter Kron fluxes. The small images show a $1\farcs5$ region around each source in the seven NIRCam bands (and the stacked ACS dropout bands), with the red circle denoting a $0\farcs3$ diameter region around the source. The $P(z)$ from the colors measured in this small circular aperture prefers lower redshifts in both cases; thus, these sources were removed from the sample.

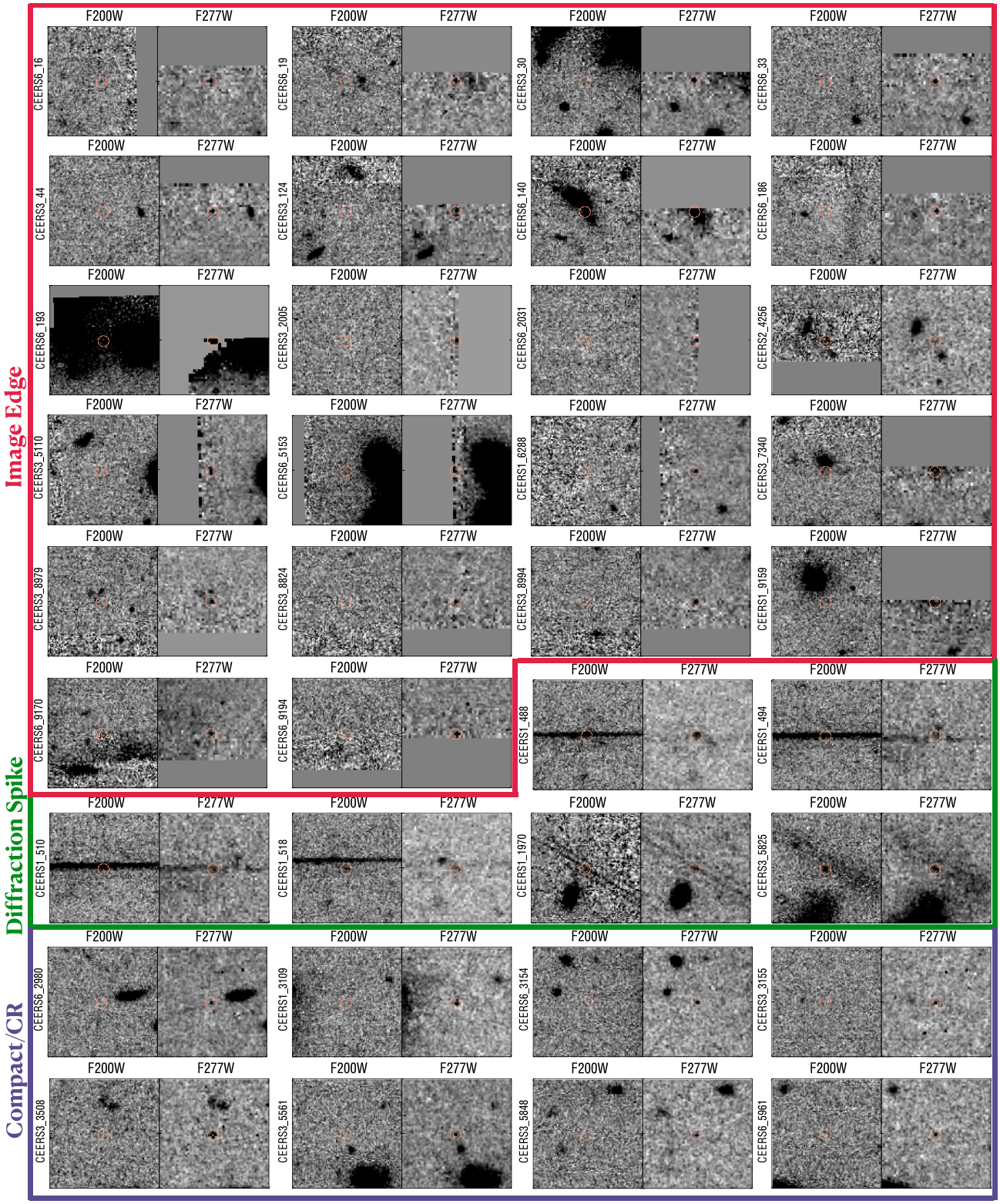

Figure 20. This compilation shows $3'' \times 3''$ cutout images of the 36 sources identified as spurious from our visual inspection of the initial list of candidates. The majority of these spurious detections come from very near the image edges, which are easily identifiable (and can be automated in the future). Six objects are obvious diffraction spikes. The remaining eight sources are very compact and boxy and visible only in the LW channels. We conclude that these are highly likely to be unflagged cosmic rays, which will be better flagged in future reductions (M. Bagley et al. 2022, in preparation). We note that seven out of eight are in the CEERS3 or 6 pointings, which had fewer images with longer exposure times than CEERS1 and 2, leading to more numerous cosmic-ray hits.

Table 6
List of Removed Sources

ID	R.A.	Decl.	Field	Reason
488	215.005247	53.017772	1	Diffraction spike
494	215.006368	53.018553	1	Diffraction spike
510	215.006574	53.018658	1	Diffraction spike
518	215.005202	53.017651	1	Diffraction spike
1970	214.945032	52.966214	1	Diffraction spike
3109	214.937432	52.953783	1	CR residual
6288	214.985433	52.968347	1	Image edge
9159	214.947769	52.980439	1	Image edge
4256	214.878918	52.904385	2	Image edge
30	214.735947	52.832000	3	Image edge
44	214.768557	52.855112	3	Image edge
124	214.798821	52.875759	3	Image edge
1537	214.750580	52.829452	3	Bad Kron
2005	214.797283	52.862516	3	Image edge
3155	214.758502	52.829340	3	CR residual
3508	214.775758	52.838390	3	CR residual
5110	214.798858	52.844780	3	Image edge
5561	214.779455	52.827964	3	CR residual
5825	214.775453	52.823445	3	Diffraction spike
5848	214.780253	52.826894	3	CR residual
7340	214.749582	52.841873	3	Image edge
8824	214.811070	52.829804	3	CR residual
8979	214.834667	52.845385	3	CR residual
8994	214.861852	52.864635	3	CR residual
16	214.806895	52.826457	6	Image edge
19	214.788359	52.813215	6	Image edge
33	214.820973	52.836319	6	Image edge
140	214.864192	52.866051	6	Image edge
186	214.871337	52.870820	6	Image edge
193	214.878573	52.875993	6	Image edge
2031	214.849667	52.843711	6	Image edge
2980	214.811286	52.810462	6	CR residual
3154	214.811199	52.809295	6	CR residual
5153	214.851208	52.826106	6	Image edge
5961	214.832642	52.808095	6	CR residual
7478	214.851426	52.812204	6	Bad Kron
9170	214.914199	52.845844	6	Image edge
9194	214.897088	52.833510	6	Image edge

Note. ID and coordinates of sources shown in Figures 19 and 20.

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