Testing the ability of the ExoMars 2018 payload to document geological context and potential habitability on Mars

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Abstract

The future ExoMars rover mission (ESA/Roscosmos), to be launched in 2018, will investigate the habitability of the Martian surface and near subsurface, and search for traces of past life in the form of textural biosignatures and organic molecules. In support of this mission, a selection of relevant Mars analogue materials has been characterised and stored in the International Space Analogue Rockstore (ISAR), hosted in Orléans, France. Two ISAR samples were analysed by prototypes of the ExoMars rover instruments used for petrographic study. The objective was to determine whether a full interpretation of the rocks could be achieved on the basis of the data obtained by the ExoMars visible-IR imager and spectrometer (MicrOmega), the close-up imager (CLUPI), the drill infrared spectrometer (Ma_Miss) and the Raman spectrometer (RLS), first separately then in their entirety. In order to not influence the initial instrumental interpretation, the samples were sent to the different teams without any additional information. This first step was called the "Blind Test" phase. The data obtained by the instruments were then complemented with photography of the relevant outcrops (as would be available during the ExoMars mission) before being presented to two geologists tasked with the interpretation. The context data and photography of the outcrops and of the samples were sufficient for the geologists to identify the rocks. This initial identification was crucial for the subsequent, iterative interpretation of the spectroscopic data. The data from the different spectrometers was, thus, cross-calibrated against the photographic interpretations and against each other. In this way, important mineralogical details, such as evidence of aqueous alteration of the rocks, provided relevant information concerning potential habitable conditions. The final conclusion from this test is that, when processed together, the ExoMars payload instruments produce complementary data allowing reliable interpretation of the

155 geological context and potential for habitable environments. This background information is
156 fundamental for the analysis and interpretation of organics in the processed Martian rocks.
157
158
159 **1. Introduction**
160
161 The ExoMars rover mission (ESA/Roscosmos) will be launched in 2018 (*ESA, 2013*).
162 The science objectives of the mission are as follows: 1) to investigate the habitability of the
163 landing site; 2) to determine whether the materials present are compatible with the
164 preservation of potential traces of life; and 3) to search for traces of past or present life,
165 including prebiotic and abiotic organics (*ESA, 2013*). To achieve these objectives, the mission
166 will have to investigate the geological diversity of rocks at the landing site. The ExoMars
167 rover consists of a suite of complementary instruments for observation and analysis. A
168 panoramic camera PanCam and the close-up imager CLUPI will be used to obtain context,
169 structural, and textural information from the kilometre- to the sub-millimetre-scale, while the
170 ISEM (long range infrared spectrometer), mounted on the mast, will determine the target
171 soil/rock bulk mineralogical composition. A drill will obtain samples down to two meters
172 depth. The drill is equipped with an IR spectrometer "Mars Multispectral Imager for
173 Subsurface Studies" (Ma_Miss; *Coradini et al., 2001*) for determining down hole mineralogy
174 (Fig. 1). The samples will be delivered to the internal laboratory where they will be crushed.
175 The mineralogy will be investigated with the visual and infrared (IR) imaging spectrometer
176 MicrOmega and the Raman Laser Spectrometer (RLS), which can also detect the presence of
177 organic matter. More detailed investigation of the organics will be made by the instrument
178 Mars Organic Molecule Analyser (MOMA), consisting of laser desorption mass spectrometer
179 (LDMS) and gas chromatograph mass spectrometery).

In support of this mission, *Bost et al. (2013)* developed a collection of Mars analogue rocks and minerals collectively known as the International Space Analogue Rockstore (ISAR, www.isar.cnrs-orleans.fr), stored at the CNRS in Orléans, France. ISAR contains well-characterised samples (sedimentary, volcanic and magmatic in origin) available for testing and calibrating space instruments. Currently, several teams use this collection to obtain scientific reference data on minerals and rocks.

However, to date, there has been no comprehensive test of the ExoMars geological suite of instruments using representative Mars analogue samples. Such studies are essential for adequately preparing future *in situ* investigations and to develop protocols based on the cross-correlation of the data. In this contribution, we describe a test of the ExoMars geological instrument suite consisting of two phases. In the first phase, called the "Blind test", the ExoMars instrument teams CLUPI, Ma_Miss, MicrOmega, and RLS were requested by the ISAR group to perform mission representative measurements on two rocks selected from the ISAR collection. They were given small sample aliquots without additional context information and without data from the other instruments. For interpretation, each instrument team could only rely on its own measurements. In the second phase, the complete set of ExoMars instrument results, together with aerial images —comparable to Mars orbital images— and photographs of the outcrops from which the samples had been obtained in the field—mimicking PanCam photographs—, were submitted to geologists having no previous knowledge of the rocks used in this exercise. The aim was, first, to identify the rocks, then the information from each instrument was cross-correlated in order to re-evaluate the initial interpretation. Finally, the cross-calibrated data were used to fully characterise the samples,

not simply in terms of rock type, but also for any features, such as alteration in the presence of water, that could provide information on potential habitability.

2. Methodology used for the Blind Test

This Blind Test was designed to address the geological capabilities of the ExoMars payload, so a procedure similar to that defined for the ExoMars mission was used.

2.1 Sample selection and characterisation during the ExoMars mission.

The characterisation of rocks and the selection of samples during the ExoMars rover mission will be made using a specific protocol. The panoramic instruments, including PanCam (including wide-angle and high resolution cameras) and ISEM, will scan and analyse the panorama and identify potentially interesting targets. This information will be employed to decide which target to approach for further investigation. PanCam will be used to study outcrops, rocks, and soils, while detailed images of these materials will be made by the close-up imager CLUPI —accommodated on an external wall of the drill box—which has several viewing modes. CLUPI will observe rock textures in macroscopic mode to understand their nature and characterise potentially visible morphological biosignatures. The synergetic combination of PanCam and CLUPI will provide powerful, nested imaging capabilities from the panoramic to the submillimetric scale. This data is important for interpretation of rock type, mode of formation, and habitability potential, as well as for identifying lithologies that could potentially preserve traces of past life.

Samples will be collected with a drill tool hosting the Ma_Miss instrument (Fig. 1). Ma_Miss can perform spectral measurements to identify subsurface minerals as the drill

moves in the borehole. PanCam and CLUPI will also be used to study the powdered fines produced during drilling, as well as the samples collected by the drill at high resolution, prior to their delivery to the analytical laboratory.

Once a sample has reached the analytical laboratory, it is crushed by the Sample Preparation and Distribution System (SPDS) to a particle size of approximately 250µm and delivered to a carrousel for IR and Raman spectrometry (MicrOmega and RLS) investigations and for analysis using the Mars Organic Molecule Analyser (MOMA) laser desorption mass spectrometry (LDMS) and gas chromatograph mass spectrometer (GCMS).

The test described here was designed to address the geological capabilities of the ExoMars payload and not the organic analyses that form part of a separate study. Furthermore, the PanCam and ISEM systems were not used because the test concentrated on the *in situ* measurements, although field photographs obtained with a commercial camera served as substitute PanCam data.

2.2 Sample selection and preparation for the Blind Test

Two samples from the ISAR collection were chosen for their analogy with Martian rocks and their pertinence for astrobiology. These samples, labelled “sample A” and “sample B”, were previously fully characterised in the laboratory using XRD, Raman spectroscopy, IR spectroscopy, optical microscopy, Scanning Electron Microscopy (SEM) and Electron Microprobe (EMP) in the framework of the ISAR project (*Bost et al., 2013*). Powdered samples with a grain size of ~ 250 µm, similar to those provided by the rock crusher in the ExoMars SPDS (*Lopez-Reyes et al., 2012; Foucher et al., 2013*), were sent to the RLS and MicrOmega teams, without any images of the original samples. Both samples were sent as

hand specimens to the CLUPI team, and as small slabs and powders to the Ma_Miss team. The teams only knew that the samples were representative of Martian rocks and that they could potentially contain biosignatures. It is important to note that, when the Blind Test phase was carried out (it started in 01/2012), these samples and the associated data were not available on the ISAR website and the related publication (*Bost et al., 2013*, submitted in 11/2012) had not yet been published. Consequently, neither the instruments teams nor the geologists had access to information related to the samples before this exercise.

3. ExoMars instruments used for the blind test

3.1. Close Up Imager (CLUPI)

The CLUPI (Close-Up Imager) is an instrument developed by the Space Exploration Institute (Space-X) in Neuchatel, Switzerland. It is a miniaturized, low-mass, low-power, efficient and highly adaptive imaging system, composed of a colour image sensor (2652 x 1768 x 3 pixels), an optics with focus mechanism and processing electronic boards. The camera is capable of focusing on any target at distances from 10 cm to infinity. The functionality of z-stacking is also implemented in order to increase the scientific return.

CLUPI is positioned on the rover's drill box and replaces the geologist's hand lens. It's scientific objectives during the ExoMars mission are: (1) Geological environment survey: determination and characterization of surface and rock types present in the immediate surroundings of the rover, study of the physical properties of the surface (e.g. compaction state) by inspection of the rover tracks, contribution to the selection of sites for detailed investigations (drilling); (2) Close-up outcrop observation: study at high resolution (down to 7

μm/px at 10 cm distance) of the texture, structure, and morphology of outcrops, surface rocks and particles, as well as potential biofabrics, colour variations and possible layering; (3) Drilling area observation: characterization of site before drilling; (4) Drilling operation observation: to provide information about the ejected fines as they are produced, potential colour changes indicative of geological variation with depth, and the mechanical behaviour of the drilled surface in contact with the drill tip; (5) Drilled core sample observation: to allow comparison of the extracted sample with the sampling area and visual examination of the texture and physical properties of the particles; (6) Drill hole observation: provide information on the surface state after drilling, the amount of ejected fines, their colour, and their physical properties.

The two Blind Test samples were imaged using a CLUPI analogue camera (Sigma SD15) with the same colour image sensor as the real CLUPI (but 20° field of view optics instead of real CLUPI's 14° FoV). The samples were deposited on a Martian surface analogue composed of Permian redbeds (red sandstones from Weitenau, southern Germany), and illuminated with a Sun simulator (Fig. 2a). Images of each sample were acquired from two rover-representative working distances, 54 cm (real CLUPI equivalent 76 cm) and 25 cm (real CLUPI equivalent 35 cm), with a pixel resolution of 60 μm and 28 μm, respectively.

3.2 Mars Multispectral Imager for Subsurface Studies (Ma_Miss)

The Mars Multispectral Imager for Subsurface Studies (Ma_Miss) is developed by the Institute for Space Astrophysics and Planetology in Rome, Italy (*Coradini et al. 2001*). Ma_Miss is a miniaturized near-infrared imaging spectrometer in the range 0.4-2.2 μm with 20-nm spectral sampling. It is positioned in the drill tool a few centimeters above the drill tip,

where a sapphire window (characterized by high transparence and hardness) protects the Ma_Miss optical head, permitting observation of the borehole wall.

The Ma_Miss optical head performs the double task of illuminating the borehole wall with a spot of approximately 1-mm diameter and, collecting the scattered light coming from a 0.1-mm diameter region of the target, Ma_Miss can acquire spectral data from the walls of the drilled borehole. It is also capable of making spectral images of the borehole by using the drill rotation and translation movements. This instrument can obtain downhole images of the excavated borehole wall, performing acquisitions at different depths during vertical translation—in principle, from 0 to 2-m depth. Ma_Miss can also create so-called ring images, performing acquisitions during rotation of the drill at a fixed depth. The Ma_Miss breadboard used for this test consists of the following subsystems: a 5W illumination lamp, an optical fiber illumination bundle, an optical head that focuses the light the sample and recollects the scattered light, a sapphire window (interface between optical head and the environment) and a collecting optical fibre (Fig. 2b).

The Blind Test samples were analysed using a preliminary version of the breadboard setup, interfaced with a commercial FieldSpec spectrophotometer. Reflectance spectra were acquired in the spectral range 0.35–1.8 μm. The sample slabs were placed directly on the sample holder. The plane surface (cut rock surface, representative of an abraded rock surface and the external surface of the drill-hole) was oriented perpendicularly to the optical axis of the spectrometer. Several spectra were acquired in different positions on the samples, simulating the Ma_Miss stratigraphic column acquisitions.

3.3 MicrOmega

MicrOmega consists of a visible light microscope and a near infrared imaging spectrometer (*Pilorget, 2012; Pilorget and Bibring, 2013*) (Fig. 2c). The instrument acquires monochromatic images with a high resolution of 20 µm x 20 µm per pixel at wavelengths between 0.5 µm and 0.9 µm and with a continuous and high spectral sampling from 0.9 µm to 2.5 µm (now up to 3.5 µm). In this way MicrOmega acquires the entire spectrum in a spectral domain for each pixel, thus enabling it to identify the composition of the samples at their grain scale.

The Blind Test samples powders were analysed using the MicrOmega breadboard. *In situ* reflectance hyperspectral spectra were acquired on samples 7.4 x 5.9 mm^2 in size with a spatial sampling of 23µm. The samples were illuminated by a monochromator with an Acousto Optical Tunable Filter (AOTF) in the range of 0.9-2.5 µm, oriented about 20° with respect to the sample surface. The 320 x 256 pixel-infrared detector (MCT Mars SW1 (Sofradir)) is sensitive in the 0.85-2.5 µm spectral region. The focal plane was cooled down to 190°K.

3.4. Raman Laser Spectrometer (RLS)

The Raman Laser Spectrometer (RLS) is developed at the Associated Unit of the University of Valladolid-CSIC-Center of Astrobiology (UVa-CAB), in Spain (*Rull et al., 2011a, b*). The RLS is accommodated in the ExoMars rover's Analytical Laboratory Drawer (ALD). In automatic mode, the RLS can perform raster analysis of at least 20 points (and up to 40) of the powdered samples, using a 50-µm spot size and an irradiance level of 0.6–1.2 kW/cm^2 with a 532 nm continuous wave, green laser.

The powdered Blind Test samples were analyzed using an RLS ExoMars simulator that includes an SPDS in order to perform measurements under the operation conditions

353 imposed by the rover-based operation (Fig. 2d; *Lopez-Reyes et al., 2014, Rull et al., 2011a*).
354 This system provides automatic flattening of the powdered samples, autofocus at each
355 measurement point, and optical images of the samples. In addition, the system autonomously
356 optimizes the acquisition parameters (integration time, number of accumulations, etc.) at each
357 spot with the aid of appropriate algorithms (*Lopez-Reyes et al., 2014*). Thirty spots per sample
358 were acquired in automatic mode. The spectra were pre-processed to remove the baseline and
359 instrument artifacts. Mineral identification based on specific Raman lines was performed
360 using the spectral database developed at the University of Valladolid (*Hermosilla, 2012*).
361
362 **4. Results of the Blind Test phase**
363
364 4.1. Sample A
365
366 4.1.1. CLUPI
367
368 CLUPI images were acquired on rough surfaces, as well as on a fresh cut face
369 (simulating an abraded surface). Note that a similar system is not planned for the ExoMars
370 rover mission. The external, uncleaned surface of the rock is characterized by alternating
371 white and grey layers ranging from mm to cm in thickness. The surface of the rock appears to
372 be a fracture surface and is coated with a whitish-orange-coloured alteration product (Figs.
373 4a-1 and a-2). The sharp angles observable on the fracture surface suggest that the rock is
374 brittle. The cleaned surface of the rock provides a better view of the layering (Figs. 3a-3 and
375 a-4). The layers are generally parallel to each other although the basal layer exhibits gently-
376 inclined internal laminae and its surface of the latter layer includes some 0.5 to 1 cm-sized,
377 rounded protrusions having a whitish cortex and a clear, orange-coloured internal component.

The laminated nature of the rock suggests a sedimentary origin and the protuberances on the surface of the lower layer may be either mineral precipitations or detrital inclusions.

4.1.2 Ma_Miss

Spectral images of the cut rock surface were collected using the Ma_Miss breadboard and documented alternating dark and light albedo layers suggestive of a sedimentary rock. Several different spots were acquired on the sample, both on the dark and bright layers, simulating the stratigraphic column of the borehole. Also the powder was measured. The dark and bright layers of the slab have corresponding spectra with very similar shapes and absorption features, although they show different levels of reflectance and spectral contrast.

The spectra acquired on both samples (slab and powder) are characterized by the OH^- absorption at 1.4 μm (Fig.3b), indicating the presence of a water-containing mineral. On the slab, the 1.4-μm band is larger and deeper for the high albedo layers than for the dark layers. This could be due to real differences in the H_2O or OH content, or just due to a reduced spectral contrast on the darker region (Fig.3b). A strong negative slope characterizes the spectra acquired in the bright region; the spectra of the dark layers show a smaller blue slope. Spectra of both layers show a clear crystal field (C.F.) absorption at 1 μm, whose wings extend beyond 1.2 μm, likely due to Fe^{2+} absorption (*Burns, 1993; Hunt, 1977; Gaffey, 1985*). The interpretation of this 1-μm absorption is not unambiguous because the iron responsible of the absorption could be present in silicates, oxides, sulfides, or carbonates. The succession of bright and dark spectra with very similar spectral shapes is also suggestive of a layered (sedimentary) structure.

4.1.3. MicrOmega

403

404 The IR spectrum (Fig.3c) obtained with the MicrOmega breadboard on the sample A
405 powder shows absorption features at 1.4 μm and 2.2 μm. There is also a very weak absorption
406 feature at 1.9 μm. These features reflect the presence of H_2O and OH^- in the minerals with
407 which they are associated. The preliminary identification is a match with the spectrum of
408 kaolinite, a group of white clays (Fe-poor) containing aluminum.

409

410 4.1.4. RLS

411

412 The thirty Raman spectra acquired on the sample powder permit identification of
413 quartz, anatase, calcite, muscovite and disordered carbonaceous matter (Fig. 3d). The main
414 rock-forming mineral is quartz.

415

416

417 4.2 Sample B

418

419 4.2.1. CLUPI

420

421 CLUPI photographed the rough and cut surfaces of Sample B. The rough surface is
422 brownish in color and characterized by a criss-cross network of indentations (Fig. 4a-1 and a-
423 2). The cut surface shows that the criss-cross network, resembling buff-coloured acicular
424 structures, infilled veins, cracks, or crystals, continues into the rock (Fig. 5a-3 and a-4). The
425 brown surface colour and buff-coloured acicular structures are restricted to the outer portion
426 of the rock, which contrasts with the uniformly grey colour of the internal portion of the rock

427 in which the acicular texture is still faintly visible. This contrast indicates significant
428 weathering of the outer part of the rock.
429
430 4.2.2 Ma_Miss
431
432 Both the rough and freshly cut sample surfaces were observed. As in the previous
433 sample, acquisition of Ma_Miss spectra at various points on the slabbed sample simulates the
434 acquisition of data "downhole" in the drill column. Spectra were obtained both in the "bright
435 region" (where the buff-coloured acicular structures occur) and in the "dark (grey) region" on
436 the flat, cut rock surface (Fig.4b). The two regions are characterized by the presence of OH;
437 the 1.4-μm band in the bright region spectra is larger and deeper than the corresponding band
438 in the dark region. The spectrum of the dark region appears flat without evident absorption
439 features. The bright region is characterized by a deep absorption near 1.0 μm, due to Fe^{2+}
440 (*Burns, 1993*) and by an absorption at 0.7 μm, likely due to Fe^{2+}-Fe^{3+} intervalence charge
441 transfer (IVCT, although electronic processes due to transition elements such as Ni, Co, Cr,
442 Fe, Mn, Ti can occur in certain minerals in this region of the spectrum*; Burns, 1993*).
443 The absorption bands of spectra taken in the bright region are suggestive of the
444 presence of mafic silicates, iron oxides, and hydrates indicative of possibly extensively altered
445 mafic or ultramafic rocks with a higher concentration of hydrated mineral phases in the
446 brighter region (the water OH band suggests alteration of silicates). The spectrum acquired on
447 powder substantially shows the same absorption bands as the rock sample, *i.e.* the iron
448 electronic transitions at 0.7-1.0 μm, the OH absorption at 1.4 μm.
449
450 4.2.3. MicrOmega
451

452 The IR spectrum obtained from the powdered sample with the MicrOmega breadboard
453 show absorption features at about 0.97, 1.43, 1.65, 1.88, 1.95, and 23.3 μm. They are
454 interpreted to reflect the presence of a mineralogical assemblage composed of saponite,
455 serpentine and forsterite (the magnesium end-member of olivine) (Fig. 4c).
456
457 4.2.4 RLS
458
459 The Raman analysis of the powdered sample B provided generally fluorescent spectra
460 with weak and very broad bands (Fig. 4d). This is consistent with a very low degree of
461 crystallinity. The main bands can be assigned to a combination of magnetite and talc. Other
462 small bands may also be assigned to clay minerals but more precise identification was not
463 possible. Brucite is not compatible with the observed spectra; antigorite and/or lizardite
464 (serpentine minerals) have also low probability. The first is characterised by a strong band at
465 1041 cm^{-1} that is not observed, while the bands of the second occur at higher wavenumbers
466 than observed. However, the presence of a chloritoid cannot be totally ruled out. In some
467 cases, the main band at 668 cm^{-1} shows a shoulder near 600 cm^{-1} that is consistent with the
468 symmetrical chain vibration of chloritoid.
469
470
471 **5. Results of the geological interpretation phase**
472
473 Two geologists specialised in geochemistry (C.R.) and petrology (N.L.B.) interpreted
474 jointly the bulk observational and analytical data for each sample to identify the rock type. In
475 order to use the same kind of data set as would be available during the future ExoMars rover
476 mission outcrop images (Figure 5) corresponding to PanCam images were provided to

complement the CLUPI images and the Ma_Miss , MicrOmega and RLS spectra. Outcrops images were obtained using commercial cameras (Olympus E410 camera, with a 10.00 Megapixel resolution, for sample A and Olympus OM1 camera for sample B).

The data interpretation followed a typical strategy, starting with the geological context (here given by the satellite observation), following by the optical outcrop and sample observation (here PanCam and CLUPI images) and finishing with the compositional data (here given by the Ma_Miss, MicrOmega and RLS spectroscopic data). Although the identification of the analogue rocks based on the optical images is a routine matter for geologists, this study underlined the importance of obtaining a maximum of information by cross-correlating the data in order to improve and/or re-evaluate the interpretation made by each instrument separately.

5.1. Data interpretation of sample A.

The outcrop images show that the rock is massive but highly layered (Fig. 5a). It is thoroughly crosscut by numerous fractures, which suggests that it is hard and brittle. The mm- to cm-wide layers consist of alternating grey-white to dark blue-grey beds. The upper and lower boundaries of the beds are generally linear, sometimes wavy, and could be interpreted as sedimentary features. Some beds are irregular in thickness because of pinching (either due to tectonic boudinage or sedimentary features?). Some dark centimetric, rounded to angular features disrupt the bedding. The massive, competent aspect of the outcrop is compatible with siliceous beds (cherts or quartzites) or marbles (metamorphosed carbonate beds). Evaporites are doubtful, given the brittle character of the outcrop. The alternation of beds with different colours in an apparently homogeneously competent material may either be accounted for by

502 different grain size in beds with similar composition, or by mineralogical differences. If the
503 rocks are siliceous (*i.e.* chert), the darker beds could be finer grained (light would diffuse at
504 grain boundaries), whereas the lighter levels could be coarser-grained. Another possibility is
505 that the darker beds display films of carbonaceous matter at grain boundaries, or contain fine
506 oxide or sulfide grains.
507
508 CLUPI and Ma_Miss observations confirm the sedimentary origin of the sample (Fig.
509 3a and b). The grey-white amorphous layer in the lower part of the sample looks like silica
510 gel, which would support the hypothesis of a chert. The lobated surface of some beds could be
511 fine sedimentary structures that have been preserved as they were rapidly covered by the
512 overlying sediment. Although the rock exhibits a massive appearance suggestive of chert, the
513 clear Fe^{2+} absorption seen by Ma_Miss (and MicrOmega) at 1.0 μm is typical of igneous rock
514 silicates. This suggests that the rock was originally volcanic in nature. The low albedo of the
515 dark region may then possibly be due to the presence of vitrified material (such as a glass).
516 The dichotomy in albedo and spectral characteristics observed in the different spots indicate a
517 stratified structure typical of sedimentary rocks. Moreover, the presence of OH absorptions
518 suggests that the volcanic material making up the rock was altered in the presence of water.
519 The detection of kaolinite by MicrOmega (Fig. 3c) is very surprising and does not fit
520 with the optical observations of the hard, brittle character of the outcropping rock. It is
521 concluded that, if kaolinite is present in the sample, it is more likely a very minor phase than a
522 major component of the sample.
523 The RLS data (Fig. 3d) are in more direct accordance with the optical data. In
524 particular, the detection of quartz as a major constituent fits well with the previous
525 interpretation of a chert rock type. The presence of carbonaceous matter is also consistent
526 with the suggestion that the rock is of sedimentary origin, the carbonaceous compounds being

more specifically associated with the darker layers. Since calcite is detected in only a few analyses, it is interpreted as only a minor rock component. The systematic detection of anatase and the small amounts of muscovite may indicate a detrital volcanic origin of the sediment in interaction with hydrothermal processes. Water was involved in the formation of this sediment. The Raman data help to eliminate the hypothesis of a banded. The kaolinite hypothesized by MicrOmega, the muscovite proposed by RLS, and the hydrated components identified by Ma_Miss can all be associated with dioctahedral smectites.

Finally, the interpretation that the rock is a banded chert with anatase (common in chert) and a small amount of dioctahedral mica is the most likely. The lighter beds are siliceous, whereas the darker ones could contain carbonaceous matter and/or anatase. Quartz is the main constituent as shown by Raman analyses. Note that quartz cannot be detected by IR spectroscopy in the spectral range used by MicrOmega and Ma_Miss . In order to explain the detection of kaolinite, which is not in accordance with a chert, the IR data interpretation was revised. It is concluded that the spectrum most probably corresponds to muscovite, in accordance with the Raman data and consistent with the fact that the spectra of kaolinite and muscovite are relatively similar in the 0.9 - 2.5 μm spectral range.

5.2. Data interpretation of sample B.

The surface of the rocks at the outcrop appears dark green to red, suggesting they are iron-rich and partly oxidized (Fig. 5b).

The outcrop photographs show that the rock is massive and characterised by what appear to be cracks. Its twisted structure evokes corodate basalt and in this regard, the rugged surface could correspond to a scoriaceous lava. The sample surface shows a reddish stain

551 (patina?) which is compatible with lava. Such a patina could also characterize a peridotite
552 (mantle rock) exposed to weathering. The white dots could be calcite or plagioclase.

553 The CLUPI optical observations confirm the red patina on the sample surface, which
554 evokes a weathered volcanic rock. The acicular texture is characteristic of the spinifex texture
555 of komatiites (Fig. 4a).

556 The Ma_Miss spectra are in accordance with the a volcanic origin of the rock (Fig.
557 4b), *i.e.* clear crystal field absorption due to Fe^{2+} at 1.0 μm indicative of iron silicate-bearing
558 phases of ultramafic/mafic rocks. Moreover, the OH^- band suggests mineral alteration,
559 indicative of possibly extensively altered mafic or ultramafic rocks, with higher a
560 concentration of hydrated mineral phases in the brighter region (serpentine group minerals,
561 olivine/pyroxene alteration products, for example antigorite, see *Clark et al., 1990*). The
562 spectral signals pointing to Fe^{2+} - Fe^{3+} IVCT transitions at 0.7 μm and to Fe^{2+} C.F. transition
563 at ~1.0 μm, together with the OH^- absorption bands are quite consistent with the fact that such
564 primitive mafic rocks are unstable in the present day oxidised and hydrated surface
565 environment. Again, the dark colour of the sample is compatible with a volcanic rock (basalt)
566 or a peridotite. The exposed surface of the sample has in its centre a massive, light-brown
567 structure whose periphery has a brecciated structure (Fig 5b). The light-brown central
568 structure is surrounded by a dense network of dark linear structures, which could correspond
569 to fractures or, more likely, to skeletal olivine crystals in exhibiting spinifex texture.

570 On the fresh cut surface, the rock appears dark and very massive. It is finely grained
571 and probably basaltic in nature. White dots on the surface, seen also in the outcrop, could be
572 either plagioclase or calcite. Yellowish greenish, narrow, specular phases occur in the upper
573 part of the sample. Two interpretations can be proposed:

- The surface of the rock may represent a fracture plane along which the rock is altered. The yellow crystals would then result from the alteration of a mineral such as olivine along this plane.
- The structure on top of the rock represents magmatic layering. The yellow skeletal crystals are rooted on a planar surface perpendicular to the observed rock section (a magmatic floor) and grow perpendicularly or obliquely to this plane. This evokes skeletal crystal growth from a supercooled, layered magma of low viscosity. Given the probable basaltic nature of the sample, the yellow acicular crystals are probably olivine forming a spinifex texture.

The olivine, serpentine and saponite (a trioctahedral smectite) detected by IR spectroscopy are quite consistent with a mafic rock (e.g. an olivine-bearing basalt) that has been hydrothermally altered to serpentine and saponite (Fig. 4c).

In the final analysis, particular aspects, such as outcrop structure, the macroscopic aspect of the lava, and IR data, favour the interpretation of anolivine-bearing basalt. Additional features allow the rock to be characterized as ultramafic. These include primary magmatic features, such as the spinifex texture of a mineral identified as olivine, possible magmatic layering, and the Mg- and Fe^{2+}-rich character of the rock. Moreover, the presence of serpentine with iron oxidation-related features on the surface of sample B demonstrates that the rock has been aqueously altered. However, some index minerals of mafic magmas, such as pyroxene and plagioclase, were not detected by spectroscopy, although the white dots visible on the rock surface could be plagioclase. Antigorite was detected in the Ma-Miss IR spectra and perhaps in the Raman spectrum. Magnetite was detected by Raman and is typically formed during serpentinization (alteration) of mafic rocks.

6. Full characterization of the samples

In this section, we present the complementary analytical data obtained for the two samples in the framework of the ISAR collection using a large range of laboratory techniques and laboratory instrumentation. Detailed data are also available on the ISAR website: www.isar.cnrs-orleans.fr.

6.1 Sample A

Sample A is the sample 00AU05 of the ISAR collection (*Bost et al., 2013*). It is a silicified volcanic sediment (chert) from the 3.446 Gy-old Kitty's Gap Chert in the Panorama formation of the Warrawoona Group, Pilbara craton, Australia (*de Vries et al., 2004; Westall et al., 2006*). Although its main constituent is now microcrystalline quartz (SiO_2), optical microscopy of thin sections of the rock shows that it consists of volcanic clasts that have been altered to muscovite and anatase and then largely replaced by silica of seawater and hydrothermal origin (confirmed by μ-Raman spectroscopy and mapping). Structures observable at outcrop scale (layering) and textures observable at the microscopic scale indicate that the rock represents volcanic sediments that were deposited in a very shallow marine environment; such as an infilling tidal channel (*de Vries et al., 2004*). The traces of carbonaceous matter identified by Raman are related to the presence of fossilized (silicified) microbial colonies (*Westall et al., 2006*; *Westall et al., 2011*). Concentrated on the surfaces of volcanic grains and in the pore spaces between the volcanic grains, these colonies most likely represent relatively simple microorganisms, such as chemolithotrophs that obtain their energy,

624 nutrients and carbon from inorganic sources. This sample is, thus, particularly relevant in
625 terms of the search for life on Mars since these volcanoclastic sediments were deposited in a
626 shallow water aqueous environmental setting that would have been relatively common in the
627 Noachian period on Mars. The simple, chemotrophic life forms that they contain could
628 therefore hypothetically reflect the kinds of simple life that may have occurred on Noachian
629 Mars (*Westall et al. 2011; Westall et al., 2013*). Moreover, during the Noachian,
630 hydrothermal processes associated with impacts and volcanic activity were likely to have
631 been important on Mars (e.g. *Schwenzer and Kring, 2009*), and the precipitation of silica and
632 subsequent silicification of igneous and sedimentary rocks and any life forms that might be
633 associated is therefore possible. Silica has only recently been detected on Mars (*Bish et al.*
634 *2013; Blake et al., 2013*), possibly due to technical limitations since quartz has no IR signal in
635 the spectral range used for Martian exploration.
636
637 6.2 Sample B
638
639 Sample B is the sample 10ZA09 of the ISAR collection (*Bost et al., 2013*). It is a
640 weathered komatiite from the type locality on the Komatii River in the Barberton Greenstone
641 belt, in South Africa (*Bost et al., 2013*). The main constituents are olivine, antigorite, micas
642 and clays, as well as traces of hematite, magnetite and talc. Some trace of carbon
643 (carbonaceous matter) is also observed in this sample. Volcanic rocks, in particular basalts,
644 are very common on the martian surface (e.g. *McSween et al. 2009*). Although, they are richer
645 in Fe and Mg than present-day terrestrial volcanics, many volcanic rocks dating back to the
646 Archaean epoch were also richer in Fe and Mg, especially the komatiites. The possible
647 presence of komatiite-like rocks from the Noachian epoch on Mars has been evoked by *Nna-*
648 *Mvondo and Martinez-Frias (2007*). It is also interesting to note that a recent experiment to

649 produce artificial basalts with a martian composition surprisingly created spinifex-like
650 textures (*Bost et al., 2012; Chevrel et al., 2013*).
651
652 **7. Discussion and conclusions**
653
654 The results collected during the Blind Test are compared to the ISAR data in Table 1.
655 There are only a few differences between the analyses made by the Exomars breadboard
656 instruments (and a CLUPI-like camera) and those made by standard laboratory instruments.
657 For sample A, calcite was detected by the RLS while it was not observed during the
658 characterization made for the ISAR collection (although *Orberger et al. (2006*) previously
659 detected traces of Ca-Mg-carbonates in this sample). This is due to the large area of analysis
660 (50 μm) of the RLS compared to spot analyses made by laboratory instruments. On the other
661 hand, goethite and rutile, present in the sample, were not detected by the ExoMars
662 instruments. For sample B, the ExoMars instruments did not detect phlogopite, hematite and
663 dolomite.
664
665 Although a trained geologist can identify rock type from observation, it is clear that
666 cross-correlation between data from different instruments, both observational and analytical,
667 is essential to fully characterise unknown rock types, as demonstrated by this study. The
668 iterative approach documented here, refining initial observational and analytical
669 interpretations through comparison with data obtained by other methods, demonstrates the
670 force of this interactive process and the complementarity of the ExoMars geological
671 instrument suite.
672

This study thus confirms the ability of the ExoMars geological instruments to carry out high quality analyses. The panoramic (field camera) and smaller-scale (CLUPI) images of the geological context provided by the cameras and the mineralogical information obtained with the RLS, Ma_Miss and MicrOmega instruments are each necessary and suitably complementary. The trained geologists were able to determine rock type from the variety of details obtained from orbit and from the field/hand specimen images. This preliminary identification was very helpful for interpreting the spectral data. The cross-calibrated spectral data were essential for the subsequent mineralogical interpretation (Table 2), in particular for determining the presence of water-bearing mineral species, important for understanding deposition/weathering/alteration signatures that have a bearing on microbial-scale habitability and the potential for preserving past traces of life.

These results allow a number of important conclusions to be drawn for future Mars and general planetary *in situ* missions:

1. Cross-correlation of data obtained with a complementary suite of observational and analytical instruments, evaluated by trained geologists is essential for the full characterisation of the rocks.
2. While the use of pure minerals for space instrument calibration is useful during the development phase of the instruments, preparation for an *in situ* mission using a suite of complementary instruments requires cross-testing with suitable analogue rocks exhibiting heterogeneous structures, textures, and mineralogy.
3. Interpretation of the data is best made by the multidisciplinary mission team, including geologists, spectroscopists, geochemists, and engineers (evaluation of the microbial-scale habitability and eventual biosignatures needs also to include relevant expertise).

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Figures caption:

Figure 1: Sketch of the ExoMars rover and the instruments locations. Credit: ESA.

Figure 2: (a) CLUPI analogue lab setup. (b) Ma_Miss breadboard setup photograph. (c) MicrOmega breadboard setup photograph (modified after *Pilorget and Bibring, 2013*) (d) RLS ExoMars simulator.

Figure 3: **Sample A analyses**. (a) CLUPI images of the rough surface and cut face from 54 cm (real CLUPI equivalent 76 cm) (*1* and *2*) and 25 cm (real CLUPI equivalent 35 cm) (*3* and *4*). The scale bar is 2cm. (b) Ma_Miss image of the cut slab and VNIR reflectance spectra acquired on the slab in the dark lower albedo layers (red and blue) and in the bright higher albedo layers (black, cyan and pink). The colored dots give the positions of spot analyses on the rock. Spectra have been shifted along the Y-axis for clarity. This series of acquisitions simulated the analysis of a stratigraphic column. (c) MicrOmega IR spectrum of the powder. (d) RLS Raman spectra of the powdered sample with mineralogical assignation.

Figure 4: **Sample B analyses**. (a) CLUPI images of the rough surface and cut face from 76 cm working distance (*i* and *ii* respectively) and at 35 cm working distance (*iii* and *iv*). The scale bar is 2cm. (b) Ma_Miss image of the slab and VNIR reflectance spectra acquired on the slab in the dark lower albedo layers (red) and in the bright, higher albedo layers (containing buff-colored acicular features) (black and green). The colour of the spectral lines corresponds to locations marked with the same colour on the rock surface. (c) MicrOmega IR spectrum of the powder. (d) RLS Raman spectra of the powder with Raman mineralogical assignation.

822 **Figure 5**: Outcrop photographs of samples A and sample B (Fig. 5a and 5b, respectively)

823