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Thermally altered subsurface material of asteroid (162173) Ryugu

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Analyses of meteorites and theoretical models indicate that some carbonaceous near-Earth asteroids may have been thermally altered due to radiative heating during close approaches to the Sun^{1–3}. However, the lack of direct measurements on the subsurface doesn't allow us to distinguish it from parent-body processes. In April 2019, the Hayabusa2 mission successfully completed an artificial impact experiment on the carbonaceous near-Earth asteroid (162173) Ryugu^{4,5}, which provided an opportunity to investigate exposed subsurface material and test potential effects of radiative heating. Here we report observations of Ryugu's subsurface material by the Near-Infrared Spectrometer (NIRS3) on the Hayabusa2 spacecraft. Reflectance spectra of excavated material exhibit a hydroxyl (OH) absorption feature that is slightly stronger and peak-shifted compared with that observed for the sur-

face, indicating that space weathering and/or radiative heating have caused subtle spectral changes in the uppermost surface. The strength and shape of the OH feature suggests that the subsurface material experienced heating above 300 °C, similar to the surface. In contrast, thermophysical modelling indicates that radiative heating cannot increase the temperature above 200 °C at the estimated excavation depth of 1 m, even at the smallest heliocentric distance possible for Ryugu. This supports the hypothesis that primary thermal alteration occurred on Ryugu's parent body.

On 5 April 2019, the Japan Aerospace Exploration Agency's (JAXA) Hayabusa2 spacecraft carried out an artificial impact experiment on the surface of Ryugu using the Small Carry-on Impactor (SCI)^{4,5}. The SCI module was separated from the spacecraft at 0.5 km altitude and fired a 2 kg copper projectile at a velocity of

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64 ~2 km s⁻¹ toward the target site on the equatorial region of Ryugu.
 65 On 25 April, Hayabusa2 approached the target site and acquired
 66 images using the Optical Navigation Camera Telescope (ONC-T),
 67 which showed that a semi-circular shaped crater of 17 m in diameter
 68 (hereafter referred to as the SCI crater) had been formed at 7.9° N,
 69 301.3° E (ref. 4). A heterogeneous expansion of ejecta was observed
 70 by the Deployable Camera (DCAM3) during the SCI experiment,
 71 and the analysis of the ONC-T images revealed that the excavated
 72 subsurface material was mainly distributed to the north side of the
 73 SCI crater⁴. The maximum depth of the SCI crater is ~2.7 m and it
 74 has been estimated that much of the subsurface material distributed
 75 on and around the crater would have been ejected from ~1 m
 76 depth⁴. The SCI experiment thus provides a unique opportunity to
 77 test the effects of surface modification by radiative heating^{1–3} and/
 78 or space weathering⁶.

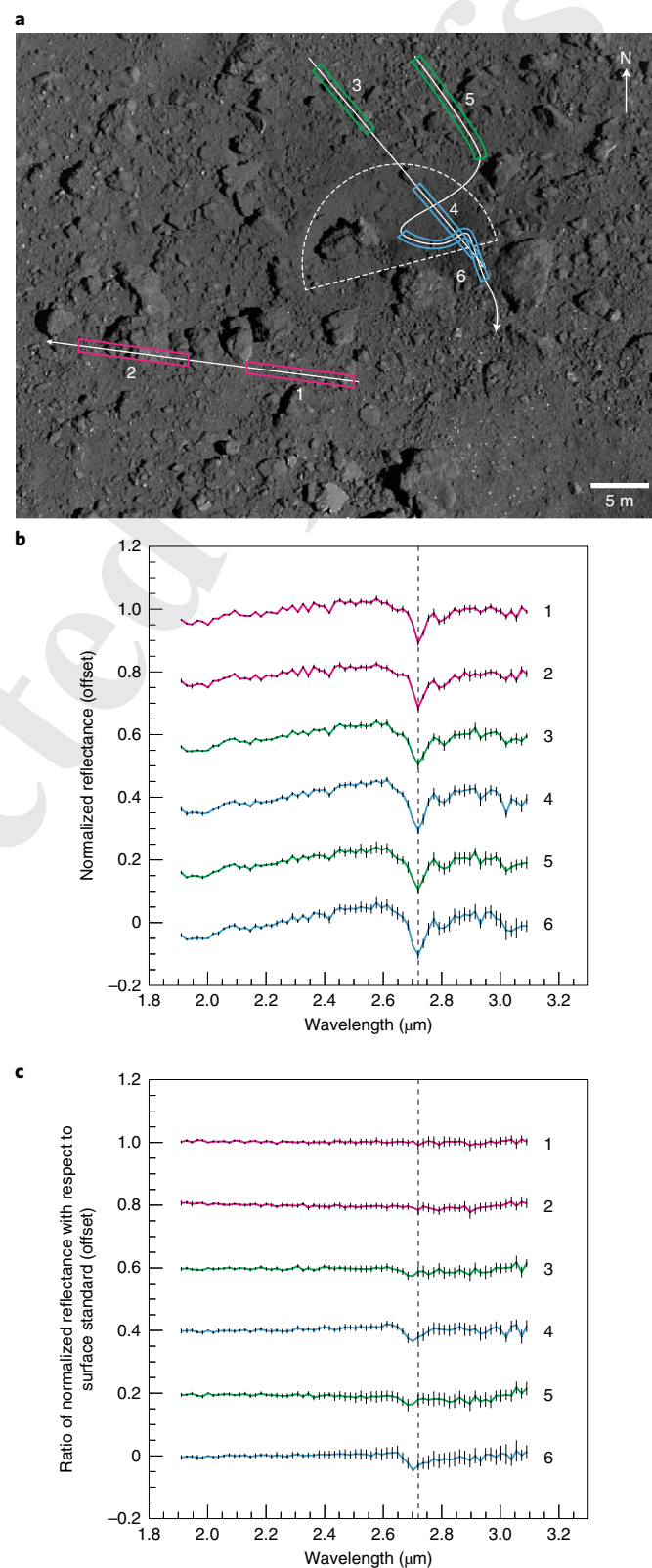
79 Before successfully acquiring a sample from ~20 m to the north-
 80 east of the SCI crater on 11 July 2019, the Hayabusa2 spacecraft
 81 performed three descent operations on different days to conduct
 82 a detailed survey of the impact site. During those descent opera-
 83 tions, the NIRS3 acquired continuous spectra over a wavelength
 84 range of 1.8 to 3.2 μm with a 0.1° field of view⁷. Two sites from
 85 each of the three descent observations were chosen to search for
 86 spectral differences between typical surface and excavated materi-
 87 als (see Methods). Figure 1 shows the location of the NIRS3 foot-
 88 prints and associated reflectance spectra for the six observations of
 89 the SCI crater region. The NIRS3 footprints of 16 May (sites 1 and
 90 2) pass through the region south of the crater, whereas those of 30
 91 May (sites 3 and 4) and 13 June (sites 5 and 6) span the region to
 92 the north and the crater interior. The distribution of crater ejecta as
 93 determined from the ONC-T and DCAM3 images reveals that sites
 94 1 and 2 and sites 3–6 correspond to ejecta-free and ejecta-rich areas,
 95 respectively⁴. All the spectra, including those over ejecta-rich areas,
 96 exhibit a weak and narrow OH feature with a reflectance minimum
 97 at 2.72 μm. This is similar to the spectral feature observed across
 98 the entire surface of Ryugu⁸ and indicates that the subsurface material
 99 in this region has been thermally altered to the same degree as the
 100 surface material.

101 Spectral differences between surface and subsurface were fur-
 102 ther evaluated by comparing the spectra observed for the SCI crater
 103 region with those from regions far from the impact site. Figure 1c
 104 shows the ratio of normalized reflectance spectra between the SCI
 105 crater region and a surface ‘standard’ (a region chosen to represent
 106 typical spectral properties of Ryugu; see Methods). The ratio spectra
 107 of sites 1 and 2 are almost flat, whereas those of sites 3 to 6 exhibit
 108 a subtle but clear feature at 2.7 μm. The presence of this feature in
 109 the ratio spectra indicates that the OH feature is slightly stronger
 110 and that the absorption peak is shifted toward a shorter wavelength
 111 within the spectral resolution of NIRS3 for the ejecta-rich sites com-
 112 pared with typical surface materials. Similar spectral differences
 113 have not previously been confirmed in any other surface materi-
 114 als on Ryugu, suggesting that these are intrinsic properties of the
 115

116
 117 **Fig. 1 | NIRS3 observations of the SCI crater region.** **a**, Context image
 118 taken by the ONC-T camera. The dashed semi-circle represents the rim
 119 of the SCI crater⁴. The arrows indicate the motion of NIRS3 footprints
 120 during three descent operations. **b**, NIRS3 spectra averaged over regions
 121 corresponding to the coloured boxes shown in **a**. The colour of each site
 122 represents its surface characteristics: magenta, green and cyan are the
 123 ejecta-free region, ejecta-rich region outside the crater and ejecta-rich
 124 region inside the crater, respectively. Error bars are 1σ within the boxes.
 125 Individual spectra of each site are shown in Extended Data Fig. 1. **c**, Ratios
 126 of the normalized spectra shown in **b** to the ones of the surface standard
 127 observed on the day before. The spectra are normalized and vertically
 128 shifted for clarity. The dashed vertical lines at 2.72 μm denote the peak
 129 wavelength of the OH absorption of the spectra in **b**.

subsurface material. Because the ejecta deposits have been found to
 span the sampling site⁹, the samples to be returned by Hayabusa2
 are expected to contain both surface and subsurface materials for
 comparison in laboratory.

Previous laboratory studies have shown that the peak wavelength
 of the OH absorption band varies with the Mg/Fe ratio of phyllo-



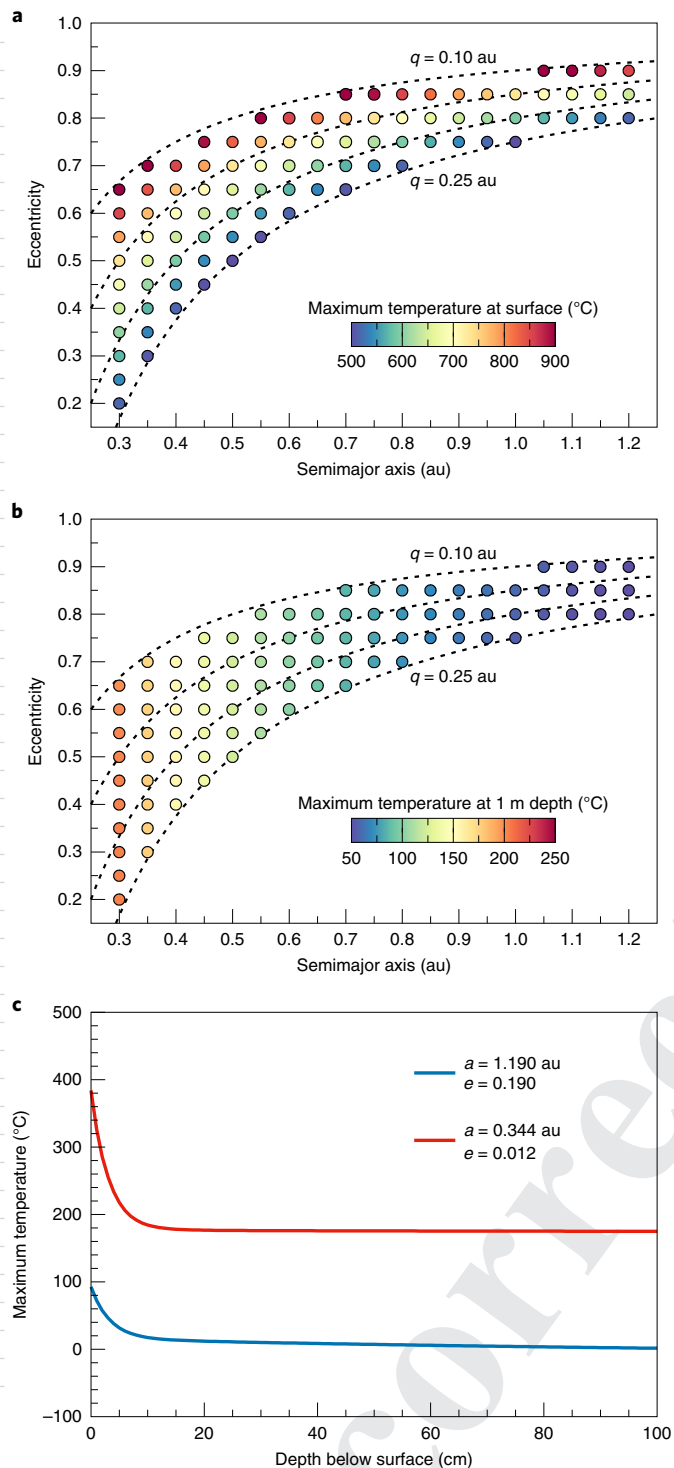


Fig. 2 | Maximum surface and subsurface temperatures at the SCI crater region. **a**, Maximum surface temperature for one revolution at a given Sun-approaching orbit. A grain density of $2,420 \text{ kg m}^{-3}$, a porosity of 41%, a thermal conductivity of $0.16 \text{ W m}^{-1} \text{ K}^{-1}$, and a temperature-dependent heat capacity³⁷ have been assumed. The dashed curves denote the perihelion distance q of 0.10, 0.15, 0.20 and 0.25 au. **b**, Same as **a** but for a depth of 1 m. **c**, Maximum temperature profile at the current orbit (blue) and the closest orbit to the Sun (red).

cate that the OH band position does not closely correlate with grain size and/or porosity¹², thus it is unlikely that intrinsic physical attributes can explain the spectral differences between surface and subsurface materials on Ryugu. Indeed, there is currently no evidence in the Thermal Infrared Imager (TIR) data to indicate differences in grain size and/or porosity for the SCI crater region¹³. Alternatively, space weathering can be considered as a likely cause of the observed spectral differences. Solar wind ion irradiation, a primary source of space weathering, affects the surface at depths of tens to hundreds of nanometres and preferentially sputters the volatile and lighter Mg compared with Fe^{14,15}. Recent laboratory experiments show that near-infrared spectra of irradiated phyllosilicates measured under vacuum conditions exhibit a similar shift in the OH band position as observed on Ryugu¹⁶. It is also possible that radiative heating may induce a change in the Mg/Fe ratio of phyllosilicates, but there is currently no available experimental data under relevant conditions to evaluate this possibility.

In spite of some differences from the surface, the overall spectral properties of the subsurface material on Ryugu still support a close similarity with thermally and/or shock metamorphosed carbonaceous chondrites. As previously discussed in ref. ⁸, the material on Ryugu was likely aqueously altered to a greater extent than indicated by the current weak OH feature, and the aqueously altered material was then heated to high enough temperatures to induce loss of H_2O but only partial dehydroxylation. This process seems to have affected materials that are currently within the upper 1 m of Ryugu. This raises questions as to the source of the heating: thermal alteration due to radiogenic and/or impact heating on the original parent body^{17–19}, or radiative heating during previous close encounters with the Sun^{20,21}? The orbits of near-Earth asteroids evolve chaotically due to a combination of close encounters with the terrestrial planets and resonances with the giant planets²². A theoretical study shows the possibility that Ryugu's perihelion distance may have been as small as 0.1 au in the course of its history, resulting in surface temperatures as high as $\sim 1,200 \text{ }^\circ\text{C}$ (ref. ²⁰). This suggests that radiative heating might be able to easily explain the thermal alteration of Ryugu's surface, and a recent study suggests this form of heating has likely influenced the spectral properties of Ryugu at visible wavelengths²¹. However, these studies do not consider thermal effects at the excavation depth of the SCI crater, thus we examine the range of temperatures that the subsurface material could have experienced by using a thermophysical model (see Methods).

Figure 2a and b show the maximum temperatures achieved at the surface and at 1 m depth, respectively, for one orbital revolution. We find that the maximum temperature of the surface increases with decreasing perihelion distance, whereas temperatures at 1 m depth increase with decreasing semimajor axis. The surface temperature reaches values of $700 \text{ }^\circ\text{C}$ at a perihelion distance of 0.15 au, and laboratory spectra of the heated Ivuna CI chondrite sample exhibit the complete loss of OH features above this temperature (Fig. 3). The persistence of an OH feature on Ryugu suggests that it never approached the Sun at a distance less than 0.15 au. In addition, dynamical models of near-Earth asteroids indicate that there is no object having an orbit entirely inside that of Mercury's orbit^{23–25}, and that their semimajor axis is never less than 0.344 au (ref. ²⁴). For that semimajor axis, the maximum temperature experienced at a depth of 1 m is estimated to be less than $200 \text{ }^\circ\text{C}$ (Fig. 2c). If impact gardening and mixing occur simultaneously with radiative heating, then materials that have experienced more heating might be able to exist in the subsurface. However, such mixing would act to homogenize the materials vertically and cannot explain the observed spectral differences between surface and subsurface materials. Thus, we can constrain the radiative heating of the material at the excavation depth of the SCI crater to below $200 \text{ }^\circ\text{C}$.

Although it is difficult to estimate the exact temperature to which materials on Ryugu have been heated, the presence of a weak and

silicates in carbonaceous chondrites^{10,11}. This would suggest that the subsurface material on Ryugu is more enriched in Mg–OH phyllosilicates relative to surface materials. Laboratory spectra also indi-

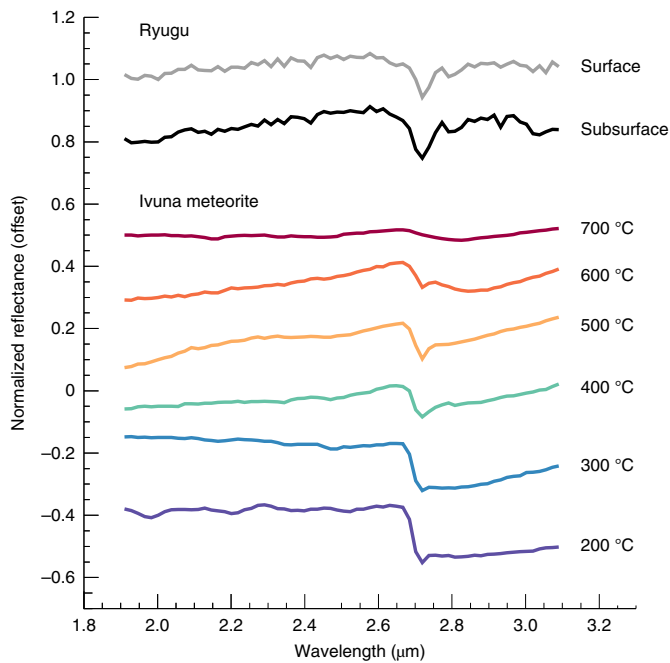


Fig. 3 | Ryugu's surface and subsurface spectra compared with laboratory spectra of heated Ivuna meteorite sample. The spectra of Ivuna meteorite taken from the Reflectance Experiment Laboratory (RELAB) database³⁸ are resampled with the same resolution as the Ryugu spectra. The spectra are normalized and vertically shifted for clarity.

narrow OH feature in Ryugu's spectra supports temperatures above 300 °C but below 700 °C, as demonstrated by the dehydration and partial dehydroxylation of Mg-rich phyllosilicates within this temperature range²⁶. It is also possible that a weak OH feature occurs due to incomplete hydration reactions on the parent body, but the OH band position of 2.72 μm is more indicative of Ryugu's material having experienced a high degree of aqueous alteration²⁷. We thus consider that the primary process of thermal alteration for material on Ryugu is more consistent with radiogenic and impact heating on the parent body from which it formed rather than radiative heating after re-accretion and formation. Extended Data Figs. 1–4.

Methods

Location of the NIRS3 footprints. The spacecraft trajectories for observations of the SCI crater region have not been well determined due to frequent delta-v manoeuvres. Therefore, we identified the location of NIRS3 footprints using the continuous series of images acquired by the ONC-T camera during the SCI crater characterization activities. The boresight of NIRS3 is co-registered with the ONC-T field of view and the pixels of the ONC-T image corresponding to the NIRS3 footprint have been precisely determined²⁸. To derive the footprints shown in Fig. 1a, we used 18 images taken from 02:36:15 to 02:40:39 on 16 May, 18 images taken from 02:35:44 to 02:40:07 on 30 May, and 12 images taken from 02:00:18 to 02:16:59 on 13 June.

Analysis of the NIRS3 spectra. Extended Data Fig. 2 shows the details of the NIRS3 spectra that we used for the analysis of the SCI crater region. The spectra of sites 3 and 4 and sites 5 and 6 were obtained under 9 °C and 10 °C higher detector temperatures, respectively, than the nominal value due to thermal flux from the surface of Ryugu. Because the responsivity of the detector changes with its temperature, we corrected that effect using additional calibration data obtained under the same temperature conditions. With the exception of this temperature correction the data processing is the same as previously published⁸.

To evaluate spectral differences between the surface and subsurface materials, we compared the spectra of the SCI crater region with those of regions not contaminated by the crater ejecta. Extended Data Fig. 3 shows the details of the NIRS3 spectra of regions that we chose as the surface standard. As shown in Extended Data Fig. 4, it has been found that a small residual of thermal correction appears between spectra of regions with different surface temperatures. To avoid such an artefact, we chose a region having a similar surface temperature to the SCI

crater region for the surface standard. Small variations in spectral slope among the spectra shown in Fig. 1c may be due to the thermal residual.

Thermophysical modelling. To investigate the effects of radiative heating on the SCI crater region, we set up a thermophysical model²⁹ that computes surface and subsurface temperatures on Ryugu with a Sun-approaching orbit. Taking into account the location of the SCI crater¹ and Ryugu's obliquity³⁰, we assume a single facet on the equator of a spherical object having the spin axis perpendicular to its orbital plane. Because the SCI crater is located at the equatorial region, even if Ryugu's past obliquity was different, the maximum temperature would not exceed that of the current perpendicular spin case considered here. The facet temperature T as a function of time t and depth z is determined by numerically solving the one-dimensional heat conduction equation:

$$\rho(1 - \phi)c \frac{\partial T}{\partial t} = \kappa \frac{\partial^2 T}{\partial z^2} \quad (1)$$

where ρ is the grain density, ϕ is the porosity, c is the specific heat capacity and κ is the thermal conductivity. Because Ryugu's surface is covered by decimetre- to metre-sized rocks without a fine-grained component³¹, we used the parameter values derived for a boulder on Ryugu from in situ measurements by the MASCOT lander³²: $\rho = 2,420 \text{ kg m}^{-3}$, $\phi = 41\%$, $\kappa = 0.16 \text{ W m}^{-1} \text{ K}^{-1}$ and a temperature-dependent heat capacity^{33,34}. These are consistent with remote sensing measurements^{33,34}.

The boundary condition at the surface is given by

$$\kappa \left(\frac{\partial T}{\partial z} \right)_{z=0} = \varepsilon \sigma T_{z=0}^4 - \frac{(1-A)S_{\odot}}{r^2} \max(0, \cos \theta) \quad (2)$$

where ε is the emissivity, σ is the Stefan–Boltzmann constant, A is the bond albedo, S_{\odot} is the solar constant at 1 au, r is the heliocentric distance in au, and θ is the angle between the surface normal and the solar vector. The position of Ryugu with respect to the Sun at a given time is computed using the Kepler's equation solution³⁵. We used a bond albedo of 0.0146 and an emissivity of 1.0 (ref. ³²), and assumed a constant rotational period of 7.6326 h (ref. ³⁰).

To ensure a sufficient depth for the seasonal temperature variations, we assumed an adiabatic boundary condition at 5 m depth:

$$\left(\frac{dT}{dz} \right)_{z=5 \text{ m}} = 0 \quad (3)$$

The current seasonal thermal skin depth of Ryugu is $l_s = \sqrt{\kappa P_{\text{orb}} / 2\pi\rho(1 - \phi)c} \approx 1.1 \text{ m}$, where P_{orb} is the orbital period, and it decreases with decreasing semimajor axis. We ran the model with a time step of 15 s and a depth step of 0.01 m. After 10 years integration, the results converged to temperature deviations of less than 1 K.

In the model, we used an emissivity of 1.0 as the nominal value considering the results of MASCOT³², but even in case of an emissivity of 0.9, the maximum temperatures of the surface and subsurface increase by only 15 K. In addition, we have also confirmed that the variation of maximum temperatures by the bond albedo uncertainty³⁶ is within 1 K.

Data availability

The data that support the plots within this paper and other findings of this study are available from the corresponding author upon reasonable request. The raw and calibrated NIRS3 data will be made available through the JAXA Data Archives and Transmission System (DARTS) website (<https://darts.isas.jaxa.jp/planet/project/hayabusa2/>).

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Author contributions

K.K. led the study, performed the data analysis and thermophysical modeling and wrote the manuscript. R.E.M. contributed to the interpretation of the results and assisted in the writing. T.I. led the development of NIRS3. M.A., M.O., S.M., M.M., L.R., Y.N., T.K. and T.A. contributed to the development and operation of NIRS3. L.R., C.P., D.L.D., E.P. and A.G. contributed to the data analysis. Y.T., T.N., T.H., M.M., L.R., M.A.B., R.B., C.P., F.P., D.L.D., F.V., D.T., E.P. and A.G. participated in the interpretation of the results. All authors participated in science data acquisition, mission planning, mission operations, or project management, and/or contributed to discussion of the results. The entire Hayabusa2 project team made this mission possible.

Competing interests

The authors declare no competing interests.

Additional information

Extended data is available for this paper at <https://doi.org/10.1038/s41550-020-01271-2>.

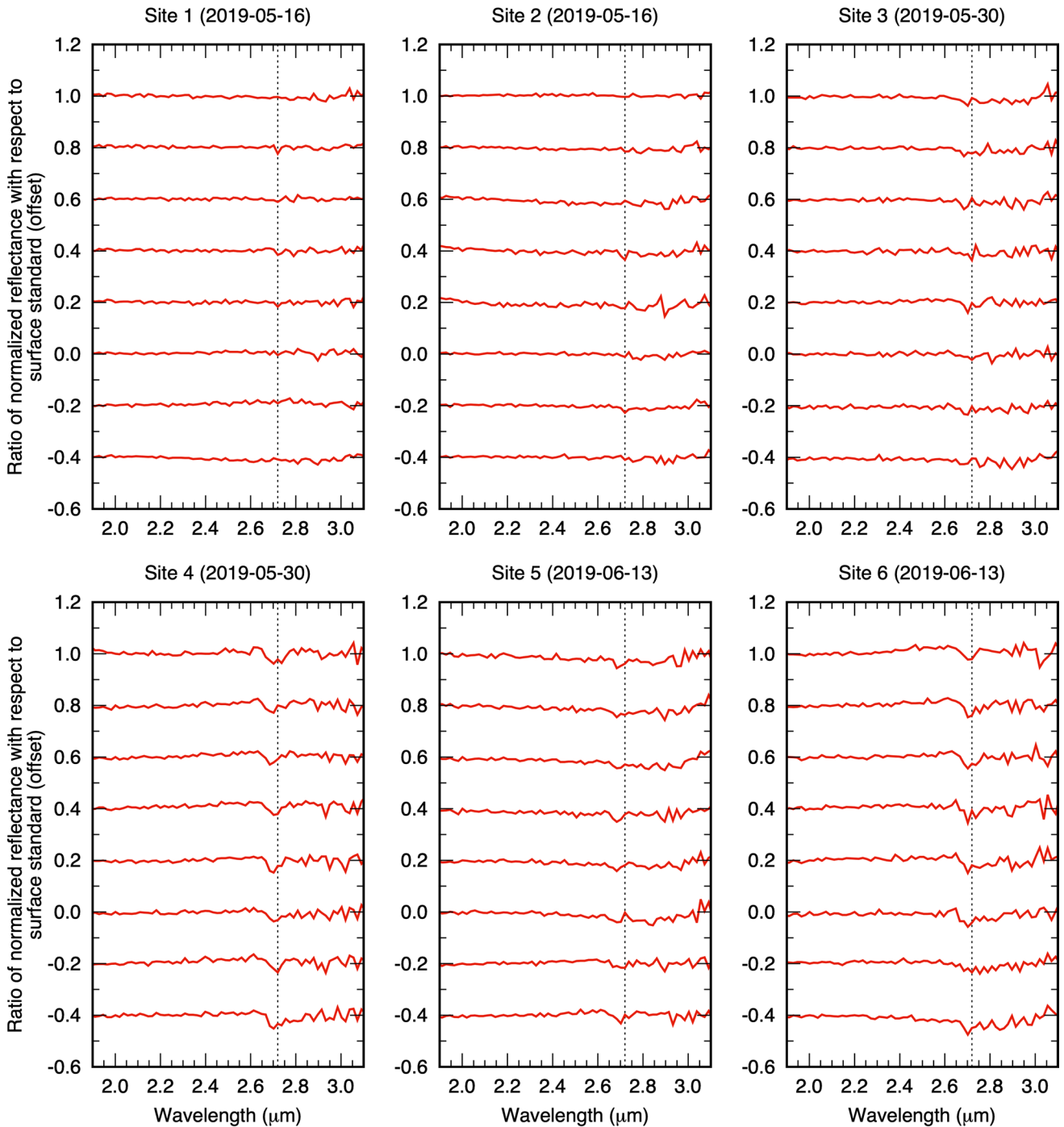
Correspondence and requests for materials should be addressed to K.K.

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Extended Data Fig. 1 | Individual spectra from each site used to derive the average spectra of the SCI crater region. The spectra are divided by the surface standard spectrum of the day before and vertically shifted for clarity.

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Site #	Observation date and time (UTC)	# of spectra	Solar distance (au)	Phase angle (°)	LIDAR range (m)	Surface temp. (°C)
1	2019-05-16 02:36:21 – 02:37:46	8	1.25	31.1	–	35 ⁺⁹ _{–9}
2	02:38:50 – 02:40:16	8	1.25	31.1	–	23 ⁺¹⁶ _{–15}
3	2019-05-30 02:35:50 – 02:37:16	8	1.21	33.7	252 – 277	41 ⁺⁷ _{–8}
4	02:38:30 – 02:39:56	8	1.21	33.7	298 – 325	50 ⁺⁷ _{–9}
5	2019-06-13 02:00:17 – 02:01:43	8	1.17	35.7	82 – 107	41 ⁺¹⁷ _{–12}
6	02:02:26 – 02:03:51	8	1.17	35.7	119 – 145	42 ⁺²⁰ _{–13}

Extended Data Fig. 2 | Details of observations of the SCI crater region. Add caption.

Observation date and time (UTC)	# of spectra	Longitude (°)	Latitude (°)	LIDAR range (km)	Surface temp. (°C)
2019-05-15 10:22:09 – 10:44:56	128	321 – 340	4 – 7	8.5 – 9.0	32 ⁺⁷ ₋₁₁
2019-05-29 06:58:34 – 07:21:21	128	112 – 136	-33 – -30	13.1 – 13.7	40 ⁺⁵ ₋₉
2019-06-12 02:46:54 – 03:09:41	128	314 – 327	-44 – -29	18.5 – 19.1	44 ⁺² ₋₁

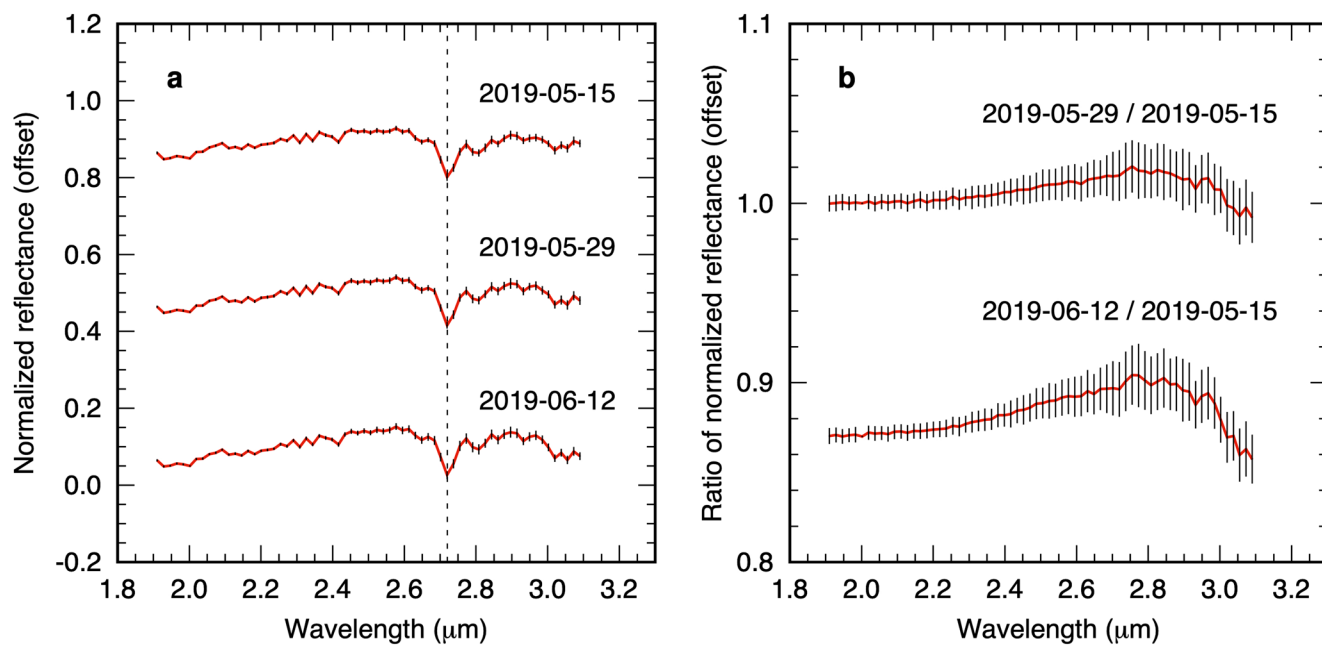
Extended Data Fig. 3 | Details of observations of the surface standard. Add caption.

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Extended Data Fig. 4 | NIR3 spectra of the surface standard. **a**, Spectra averaged over regions having the similar surface temperature to the SCI crater region. The details of these spectra are listed in Extended Data Fig. 3. **b**, Ratios between the normalized spectra shown in **a**. The non-flat shape of the ratio-spectra indicates the residual of thermal correction. The spectra are normalized and vertically shifted for clarity. Note that the vertical scale of **b** is much larger than that of **a** to show the curvature and uncertainties of the ratio-spectra.

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