



<b>Publication Year</b>	2021
<b>Acceptance in OA</b>	2022-03-24T16:06:37Z
<b>Title</b>	Cepheid Metallicity in the Leavitt Law (C-MetaLL) survey: I. HARPS-N@TNG spectroscopy of 47 Classical Cepheid and 1 BL Her variables
<b>Authors</b>	RIPEPI, Vincenzo, CATANZARO, Giovanni, MOLINARO, Roberto, GATTO, MASSIMILIANO, DE SOMMA, GIULIA, MARCONI, Marcella, M. Romaniello, LECCIA, Silvio, MUSELLA, ILARIA, E. Trentin, CLEMENTINI, Gisella, TESTA, Vincenzo, CUSANO, FELICE, J. Storm
<b>Publisher's version (DOI)</b>	10.1093/mnras/stab2460
<b>Handle</b>	<a href="http://hdl.handle.net/20.500.12386/31897">http://hdl.handle.net/20.500.12386/31897</a>
<b>Journal</b>	MONTHLY NOTICES OF THE ROYAL ASTRONOMICAL SOCIETY
<b>Volume</b>	508

# Cepheid metallicity in the Leavitt law (C-metall) survey – I. HARPS-N@TNG spectroscopy of 47 classical Cepheids and 1 BL Her variables

V. Ripepi<sup>1</sup>,<sup>1</sup>★ G. Catanzaro<sup>2</sup>,<sup>2</sup>★ R. Molinaro,<sup>1</sup> M. Gatto<sup>1,3</sup>, G. De Somma<sup>1,3,4</sup>, M. Marconi<sup>1</sup>,<sup>1</sup> M. Romaniello,<sup>5</sup> S. Leccia,<sup>1</sup> I. Musella<sup>1</sup>, E. Trentin,<sup>1,6</sup> G. Clementini,<sup>7</sup> V. Testa,<sup>8</sup> F. Cusano<sup>7</sup> and J. Storm<sup>1,6</sup>

<sup>1</sup>INAF-Osservatorio Astronomico di Capodimonte, Salita Moiarriello 16, I-80131 Naples, Italy

<sup>2</sup>INAF-Osservatorio Astrofisico di Catania, Via S. Sofia 78, I-95123 Catania, Italy

<sup>3</sup>Dipartimento di Fisica ‘E. Pancini’, Università di Napoli ‘Federico II’, Via Cinthia, I-80126 Napoli, Italy

<sup>4</sup>Istituto Nazionale di Fisica Nucleare (INFN)-Sez. di Napoli, Via Cinthia, I-80126 Napoli, Italy

<sup>5</sup>European Southern Observatory (ESO), Karl-Schwarzschild-Str., D-85748 Garching, Germany

<sup>6</sup>Leibniz-Institut für Astrophysik Potsdam (AIP), An der Sternwarte 16, D-14482 Potsdam, Germany

<sup>7</sup>INAF-Osservatorio di Astrofisica e Scienza dello Spazio, Via Gobetti 93/3, I-40129 Bologna, Italy

<sup>8</sup>INAF-Osservatorio Astronomico di Roma, via Frascati 33, I-00078 Monte Porzio Catone, Italy

Accepted 2021 August 24. Received 2021 August 23; in original form 2021 August 1

## ABSTRACT

Classical Cepheids (DCEPs) are the most important primary indicators of the extragalactic distance scale. Establishing the dependence on metallicity of their period–luminosity and period–Wesenheit (*PLZ/PWZ*) relations has deep consequences on the calibration of secondary distance indicators that lead to the final estimate of the Hubble constant ( $H_0$ ). We collected high-resolution spectroscopy for 47 DCEPs plus 1 BL Her variables with HARPS-N@TNG and derived accurate atmospheric parameters, radial velocities, and metal abundances. We measured spectral lines for 29 species and characterized their chemical abundances, finding very good agreement with previous results. We re-determined the ephemerides for the program stars and measured their intensity-averaged magnitudes in the *V*, *I*, *J*, *H*, *K<sub>s</sub>* bands. We complemented our sample with literature data and used the *Gaia* Early Data Release 3 (EDR3) to investigate the *PLZ/PWZ* relations for Galactic DCEPs in a variety of filter combinations. We find that the solution without any metallicity term is ruled out at more than the  $5\sigma$  level. Our best estimate for the metallicity dependence of the intercept of the *PLK<sub>s</sub>*, *PWJK<sub>s</sub>*, *PWVK<sub>s</sub>*, and *PWHVI* relations with three parameters is  $-0.456 \pm 0.099$ ,  $-0.465 \pm 0.071$ ,  $-0.459 \pm 0.107$ , and  $-0.366 \pm 0.089$  mag dex<sup>-1</sup>, respectively. These values are significantly larger than the recent literature. The present data are still inconclusive to establish whether or not also the slope of the relevant relationships depends on metallicity. Applying a correction to the standard zero-point offset of the *Gaia* parallaxes has the same effect of reducing by  $\sim 22$  per cent the size of the metallicity dependence on the intercept of the *PLZ/PWZ* relations.

**Key words:** stars: abundances – stars: distances – stars: fundamental parameters – stars: variables: Cepheids – distance scale.

## 1 INTRODUCTION

Classical Cepheids (DCEPs) represent a fundamental step of the extra-galactic distance ladder thanks to the period–luminosity (PL) and period–Wesenheit (PW) relations that holds for these objects (see e.g. Leavitt & Pickering 1912; Madore 1982; Caputo, Marconi & Musella 2000; Riess et al. 2016). These relations are usually calibrated by means of geometric methods such as trigonometric parallaxes, eclipsing binaries and water masers, and, in turn, used to calibrate secondary distance indicators, including the Type Ia supernovae (SNIa), which are sufficiently intrinsically bright to allow measuring the distances of far galaxies in the unperturbed Hubble flow. The Hubble constant ( $H_0$ ) is then estimated by measuring the slope of the relation between the distance to galaxies in the Hubble

flow and their recession velocities. This three-step path has long been known as the distance ladder, and used in the last decades to estimate the  $H_0$  value (e.g. Sandage et al. 2006; Freedman et al. 2012; Riess et al. 2016).

The actual value of  $H_0$  has become a matter of hot debate in the last years, as its value estimated by means of the distance ladder (see e.g. Riess 2020, and references therein) is significantly discrepant with respect to the  $H_0$  value estimated by the Planck Cosmic Microwave Background (CMB) project adopting the flat  $\Lambda$  Cold Dark Matter ( $\Lambda$ CDM) model (e.g. Freedman 2017; Verde, Treu & Riess 2019, and references therein). This occurrence did not change even with the recent Early Data Release 3 of the *Gaia* mission (EDR3 *Gaia* Collaboration 2016, 2021). Indeed, the latest estimate by the group led by A. Riess is  $H_0 = 73.2 \pm 1.3$  km s<sup>-1</sup> Mpc<sup>-1</sup> (Riess et al. 2021), in tension by  $4.2\sigma$  with the Planck flat  $\Lambda$ CMB value of  $H_0 = 67.4 \pm 0.5$  km s<sup>-1</sup> Mpc<sup>-1</sup> (Planck Collaboration 2020).

\* E-mail: [vincenzo.riepi@inaf.it](mailto:vincenzo.riepi@inaf.it) (VR); [giovanni.catanzaro@inaf.it](mailto:giovanni.catanzaro@inaf.it) (GC)

To better quantify and characterize this tension it is mandatory to investigate the remaining systematics concerning the various paths to estimate  $H_0$ . In particular, it is critical to reduce both the random and the systematic errors, in each rung of the distance ladder calibration. In this way it will be possible to enhance the accuracy of the estimated  $H_0$  value and, in turn, to put rigorous constraints on the different cosmological models.

In this context, an enduring source of uncertainty in the cosmic distance scale is represented by the dependence on the metallicity of the Cepheid PL/PW relations which are the standard candles used to calibrate the secondary distance indicators such as the SNIa. Indeed, the metal content of DCEPs is thought to affect both the slope and intercept of their PL/PW relations, especially in the optical filters, whereas this dependence is expected to be mild for various band combinations involving near-infrared (NIR) filters (see e.g. Fiorentino et al. 2007; Ngeow et al. 2012; Di Criscienzo et al. 2013; Fiorentino, Musella & Marconi 2013; Gieren et al. 2018, and references therein). However, the metallicity dependence should be considered to mitigate the systematic effects in the definition of the cosmic distance ladder (Romaniello et al. 2008; Bono et al. 2010, and references therein).

Previous estimates for the metallicity dependence of the PL/PW relations published in the literature differ significantly from each other (e.g. Macri et al. 2006; Romaniello et al. 2008; Bono et al. 2010; Freedman & Madore 2011; Shappee & Stanek 2011; Pejcha & Kochanek 2012; Groenewegen 2013; Kodric et al. 2013; Fausnaugh et al. 2015; Riess et al. 2016). These diverse results are mainly caused by the significant uncertainties still affecting metal abundance measures for DCEPs beyond the Milky Way (MW). The situation has completely changed with the advent of the Gaia mission, whose astrometric measures allow us to measure locally, in the MW the metallicity dependence of the PL/PW relations.

The first such an attempt was carried out by Groenewegen (2018, G18 hereafter) who presented an analysis of the metallicity dependence of Galactic DCEPs PL/PW relations in the NIR based on Gaia data release 2 (DR2 Gaia Collaboration 2016, 2018) parallaxes. This relatively large sample of Galactic DCEPs (more than 400) also has NIR photometry and individual [Fe/H] measurements obtained from high-resolution spectroscopy. As result, G18 did not find a metallicity dependence of the intercept of the NIR PL/PW better than  $1\sigma$  significance, mainly due to the still too imprecise DR2 parallaxes and uncertain Gaia zero-point offset (ZPO) of the parallaxes themselves. A similar result was obtained by Ripepi et al. (2019) in the Gaia band-passes. More recently, Ripepi et al. (2020a) analysed the metallicity dependence of the NIR PL/PW relations for Galactic DCEPs using a sample similar to that by G18. They found indications of a dependence of not only the intercept of the NIR PL/PW relations, but, for the first time, also of the slopes of these relations. However, the statistical significance of these findings was inconclusive.

More recently, the SHOES team (Riess et al. 2016, 2021) used a sample of 75 DCEPs with accurate *Hubble Space Telescope* (HST) photometry (Riess et al. 2018) and Gaia EDR3 parallaxes, to derive a PW relation in the ( $H$ ,  $V$ ,  $I$ ) bands taking into account a metallicity term on the intercept, equal to  $-0.18 \pm 0.13$ , or  $-0.20 \pm 0.13$  mag dex $^{-1}$  (hence slightly more significant than  $1\sigma$ ) and in agreement with the value used by the previous works of the group, i.e.  $-0.17$  mag dex $^{-1}$ . The large error is likely due to the small sample adopted for that study and the modest metallicity interval spanned by the adopted DCEP sample. A similar result for the intercept metallicity dependence was obtained by Breuval et al. (2021) which combined three samples of DCEPs: MW, Large Magellanic Cloud (LMC) and Small Magellanic Cloud (SMC) to estimate the metallicity term of the

PL relation. To this aim they used the accurate geometric LMC and SMC distances based on late-type detached eclipsing binary systems (DEBs) recently published by Pietrzyński et al. (2019) and Graczyk et al. (2020), respectively. In the MW, they adopted DCEPs parallaxes from the EDR3. Breuval et al. (2021) reported a metallicity effect of  $-0.221 \pm 0.051$  mag dex $^{-1}$  on the intercept of the PL in the  $K_s$  band. This metallicity effect estimate is not entirely based on the DCEP properties but mixes DCEPs and DEBs as standard candles. Both the above-quoted works assumed that the metallicity affects only the intercept of the DCEPs PL/PW relations, but this is a hypothesis not corroborated by the theoretical results as mentioned above.

In this context, we are carrying out a long term program dubbed ‘Cepheid-Metallicity in the Leavitt Law’ (C-MetaLL) devoted to the direct measure of the metallicity dependence of the DCEPs PL/PW relations using only MW pulsators with accurate metallicities from high-resolution spectroscopy, optical and NIR photometry, and Gaia parallaxes. The main scope of this project is to enlarge by more than 50 per cent the number of DCEPs with metallicities from high-resolution spectroscopy and NIR photometry than currently available in the literature. To reach our goal, we are gathering high-resolution spectra of Galactic DCEPs with HARPS-N@TNG,<sup>1</sup> UVES@VLT<sup>2</sup> and PEPSI@LBT,<sup>3</sup> as well as NIR photometry with the REM telescope.<sup>4</sup> In particular, we aim at increasing the metallicity range of observed DCEPs towards the metal-poor ([Fe/H] <  $-0.3$ – $-0.4$  dex) regime where only a dozen DCEPs exist with literature measures (see Section 5).

In this paper we present the first results of the C-MetaLL project, presenting high-resolution spectroscopy for a sample of 47 DCEPs and 1 BL Her (BLHER in the following)<sup>5</sup> observed with HARPS-N@TNG. This sample is complemented by literature data to estimate new PL-Metallicity (PLZ) and PW-Metallicity (PWZ) relations.

The paper is organized as follows: in Section 2, we describe observations and data analysis; in Section 3, we derive atmospheric parameters for the program stars; in Section 4, we characterise the program stars from the photometric point of view while in Section 5 we merge our sample with literature data; in Section 6, we discuss the astrometric properties of our and the literature samples; in Section 7, we derive the PLZ/PWZ relations; in Section 8, we discuss our results. Finally, in Section 9 we summarize the main findings of this work.

## 2 OBSERVATIONS AND DATA REDUCTION

Multiphase spectroscopic observations for our targets were obtained at the 3.5m Telescopio Nazionale Galileo (TNG) equipped with the HARPS-N instrument, during four different periods between summer 2019 and fall 2020 (TNG AOT 39 to AOT 42). HARPS-N features an echelle spectrograph covering the wavelength range between 3830 and 6930 Å, with a spectral resolution  $R = 115\,000$ . We collected at least three spectra for each source. The signal-to-noise ratio (SNR) varies from 50 to 100 at  $\lambda = 5000$  Å. Coordinates of our targets and the observation log are provided in Table 1.

<sup>1</sup>High Accuracy Radial velocity Planet Searcher – North at Telescopio Nazionale Galileo: <http://www.tng.iac.es/instruments/harps/>

<sup>2</sup>Ultraviolet and Visual Echelle Spectrograph at Very Large Telescope: <https://www.eso.org/sci/facilities/paranal/instruments/uves.html>

<sup>3</sup>Potsdam Echelle Polarimetric and Spectroscopic Instrument at Large Binocular Telescope: <https://pepsi.aip.de/>

<sup>4</sup>Rapid Eye Mount telescope: <https://www.eso.org/public/italy/teles-instr/lasilla/rem/>

<sup>5</sup>As we show in the following, star ASAS\_J162326–0941.0, formerly classified as DCEP in the literature, turned out to be a BLHER variable.

**Table 1.** Properties of the program stars. The different columns provide: (1) name of the targets according to the SIMBAD database (Wenger et al. 2000); (2) right ascension; (3) declination; (4) exposure time (in seconds) of the acquired spectra.

ID (1)	RA (J2000) (2)	Dec (J2000) (3)	$T_{\text{exp}}$ (s) (4)
ASASSN J180946.70– 182238.2	18:09:46.68	–18:22:38.2	3 × 1800
ASAS J052610+1151.3	05:26:09.64	+11:51:13.2	3 × 1800
ASAS J061022+1438.6	06:10:22.23	+14:38:40.2	3 × 1200
ASAS J063519+2117.8	06:35:19.18	+21:17:48.1	1 × 3000+2 × 3600
ASAS J065413+0756.5	06:54:13.53	+07:56:29.1	1 × 3000+2 × 3600
ASAS J070832– 1454.5	07:08:31.69	–14:54:26.8	3 × 1200
ASAS J070911– 1217.2	07:09:10.80	–12:17:09.2	3 × 1800
ASAS J072424– 0751.3	07:24:24.52	–07:51:19.6	3 × 1800
ASAS J074401– 1707.8	07:44:00.61	–17:07:44.0	1 × 3000+2 × 3600
ASAS J074412– 1704.9	07:44:11.92	–17:04:51.6	3 × 1800
ASAS J162326– 0941.0	16:23:26.38	–09:40:59.3	3 × 1200
ASAS J180342– 2211.0	18:03:41.74	–22:10:58.7	3 × 1500
ASAS J182714– 1507.1	18:27:13.45	–15:07:04.7	3 × 1800
ASAS J183347– 0448.6	18:33:46.84	–04:48:32.9	3 × 1200
ASAS J183652– 0907.1	18:36:52.35	–09:07:04.8	1 × 1800+2 × 2700
ASAS J183904– 1049.3	18:39:04.03	–10:49:21.2	1 × 1200+2 × 1800
ASAS J192007+1247.7	19:20:06.96	+12:47:43.0	3 × 1500
ASAS J192310+1351.4	19:23:10.24	+13:51:24.0	3 × 2700
BD+59 12	00:12:40.0	+60:13:35.1	3 × 1800
CF Cam	03:35:12.01	+58:17:40.9	3 × 3600
DR2 468646563398354176	04:06:26.84	+56:22:57.4	3 × 2400
DR2 514736269771300224	02:15:31.46	+63:31:04.0	3 × 2400
HD 160473	17:41:08.26	–23:28:27.5	3 × 1500
HD 344787	19:43:29.39	+23:10:40.6	3 × 1500
HO Vul	19:58:53.51	+24:23:42.2	3 × 3000
OGLE-GD-CEP-0066	06:47:26.46	+04:11:27.3	1 × 3000+2 × 3600
OGLE-GD-CEP-0104	07:25:41.60	–19:53:48.0	3 × 3600
OO Pup	07:33:38.80	–16:19:01.6	3 × 2700
OP Pup	07:39:21.49	–17:20:49.1	3 × 2700
OR Cam	03:34:23.74	+58:24:50.0	3 × 2100
TX Sct	18:28:14.50	–11:15:08.8	3 × 1800
V1495 Aql	18:53:17.88	–00:06:27.3	3 × 2400
V1496 Aql	18:54:59.53	–00:04:36.4	3 × 1200
V1788 Cyg	20:42:37.18	+38:27:25.4	3 × 3600
V2475 Cyg	20:13:56.21	+35:19:41.4	3 × 3600
V355 Sge	19:35:44.75	+18:56:42.6	3 × 1800
V363 Cas	00:15:14.33	+60:20:25.7	3 × 1800
V371 Gem	06:10:19.36	+24:01:15.1	3 × 1800
V383 Cyg	20:28:58.16	+34:08:06.4	3 × 2700
V389 Sct	18:27:35.88	–11:16:42.9	3 × 2100
V536 Ser	18:11:30.59	–15:55:33.5	3 × 1800
V5567 Sgr	18:21:05.53	–18:27:19.7	3 × 1200
V598 Per	03:11:07.34	+55:30:29.7	3 × 3600
V824 Cas	00:22:31.30	+63:01:59.1	3 × 2400
V912 Aql	18:48:16.32	+00:49:03.6	3 × 1800
V914 Mon	06:45:53.32	+10:03:41.3	3 × 1800
V946 Cas	02:44:19.39	+64:45:57.5	3 × 3000
V966 Mon	06:27:34.78	+09:49:49.7	3 × 3600
X Sct	18:31:19.74	–13:06:29.4	3 × 1200
ZTF J000234.99+650517.9	00:02:34.99	+65:05:17.9	3 × 2400

Reduction of all the spectra, which included bias subtraction, spectrum extraction, flat fielding, and wavelength calibration, was done using the HARPS-N reduction pipeline. Radial velocities were measured by cross-correlating each spectrum with a synthetic template. The cross-correlation was performed using the IRAF task *FXCOR* and excluding Balmer lines as well as wavelength ranges with telluric lines. The IRAF package *RVCORRECT* was used to determine the heliocentric velocity by correcting the spectra for the Earth’s motion.

### 3 ABUNDANCE ANALYSIS

#### 3.1 Stellar parameters

To measure the iron abundance and, more in general, the chemical pattern of our targets, we first need to estimate the main atmospheric parameters such as effective temperature ( $T_{\text{eff}}$ ), surface gravity ( $\log g$ ), microturbulent ( $\xi$ ), and line broadening parameter ( $v_{\text{br}}$ ), which represents the combined effects of macroturbulence and

rotational velocity (though, in the case of DCEPs, it is essentially dominated by macroturbulence).

An effective temperature for each spectrum of our targets was estimated by using the line depth ratios (LDRs) method (Kovtyukh & Gorlova 2000). LDRs have the advantage of being sensitive to temperature variations, but not to abundances and interstellar reddening. Typically, we measured about 32 LDRs in each spectrum.

An iterative procedure was then applied to determine microturbulent velocities, iron abundances and gravities. The  $\xi$  values were estimated by demanding the slope of  $[\text{Fe}/\text{H}]$  as a function of equivalent width (EW) be null, that is, the iron abundance does not depend on EWs. For this purpose we used a sample of 145 Fe I spectral lines, extracted from the line list published by Romaniello et al. (2008). The iron content was estimated by converting the measured EWs into abundances through the WIDTH9 code (Kurucz & Avrett 1981), after generating the appropriate model atmosphere with ATLAS9. At this stage, we neglected the  $\log g$ , as the Fe I lines are insensitive to this parameter. EWs were measured using an *IDL*<sup>6</sup> semi-automatic custom routine which allowed us to minimize errors in the continuum estimation on the wings of the spectral lines. Then, we estimated the surface gravity by imposing the ionization balance between Fe I and Fe II lines. For the Fe II we used a list of 24 lines extracted from the compilation by Romaniello et al. (2008).

The atmospheric parameters estimated as described above are summarized in Table 2. They were used as input values for the abundance analysis presented in the next section. In particular, for the iron abundance we used the weighted average of the values derived from the individual spectra available for each target.

### 3.2 Abundances

The stellar parameters in Table 2 were used to compute stellar atmosphere models and synthetic spectra to analyse the spectroscopic observations of our targets. We applied a spectral synthesis technique to avoid problems arising from spectral line blending caused by line broadening. Synthetic spectra were generated in three steps: (i) plane-parallel local thermodynamic equilibrium (LTE) atmosphere models were computed using the ATLAS9 code (Kurucz 1993a, b); (ii) stellar spectra were synthesized by using SYNTHE (Kurucz & Avrett 1981); (iii) the synthetic spectra were convoluted for instrumental and line broadening. This was evaluated by matching the synthetic line profiles to a selected set of the observed metal lines.

In our targets we detected spectral lines for 29 different chemical species. For all elements we performed the following analysis: we divided the observed spectra into intervals, 25 or 50 Å wide, and derived the abundances in each interval by performing a  $\chi^2$  minimization of the differences between the observed and synthetic spectra. The minimization algorithm was written in *IDL* language, using the *amoeba* routine. We adopted lists of spectral lines and atomic parameters from Castelli & Hubrig (2004), who have updated the original parameters of Kurucz & Bell (1995). When necessary we also checked the NIST data base (Kramida et al. 2019).

The derived abundances are listed in Table 3. We considered several sources of errors in our abundances. First, we evaluated the expected errors caused by variations in the fundamental stellar parameters of ( $\delta T_{\text{eff}} = \pm 150$  K,  $\delta \log g = \pm 0.2$  dex, and  $\delta \xi = \pm 0.3$  km s<sup>-1</sup>). According to our simulations, the uncertainties in

the atmospheric parameters contribute  $\pm 0.1$  dex to the total error budgets. Total errors were evaluated by summing in quadrature this value to the standard deviations obtained from the average abundances.

Fig. 1 shows the derived chemical elements distributions in the form of histograms, where bins have been fixed to 0.15 dex (that is representative of the experimental errors). A Gaussian fit has been overimposed on each histogram. In the upper right corner of each box, we also report the centre of the Gaussian and its full width at half-maximum (FWHM). All the abundances are referred to the solar value (Grevesse et al. 2010). In the following we will comment on the distribution of the chemical abundance for each element or group of elements.

(i) *C(N)O*: carbon was measured through the following spectral lines:  $\lambda\lambda$  4770.027, 4932.049, 5052.144, and 5380.325 Å, the average abundance is about 0.23 dex under the solar value. For oxygen, whose abundance has been inferred via the forbidden line [O I] 6300.304 Å and the O I triplet at 6155.8 Å, we obtained a distribution practically centred on the solar value. Unfortunately, in our spectral range no nitrogen lines have been detected.

(ii) *Sodium*: we measured four Na I lines in each star:  $\lambda\lambda$  5682.647, 5688.217, 6154.230, and 6160.753 Å. As expected, sodium is on average quite overabundant, its average value is  $\approx 0.3$  dex over the corresponding solar value.

(iii) *Aluminum*: we detected two aluminum lines in our spectra, namely: Al I  $\lambda\lambda$  6696.018 and 6698.667 Å. The obtained distribution is centred around 0.1 dex.

(iv)  *$\alpha$ -elements: Mg, Si, S, Ca, and Ti*: magnesium (inferred via the two spectral lines of Mg I at  $\lambda\lambda$  4702.991 and 5528.405 Å), silicon, sulfur (via the S II line  $\lambda$  6757.153 Å), calcium (three red lines generated by the Ca I,  $\lambda\lambda$  6462.567, 6493.781, and 6499.650 Å) and titanium (we detected ten lines of Ti I as well as ten of Ti II) show abundances consistent with or slightly over than solar values.

(v) *Scandium*: three Sc II lines were detected in our spectra, namely,  $\lambda\lambda$  5031.021, 5239.813, and 6245.621 Å. They lead to a distribution centred at 0.3 dex.

(vi) *Vanadium*: five lines were measured:  $\lambda\lambda$  4379.246, 5698.482, 5703.569, 6090.194, and 6243.088 Å. Vanadium is slight overabundant in our sample of stars.

(vii) *Iron peak elements: Cr, Mn, Fe, Co, Ni*: all the iron peak elements show distributions centred on standard (Mn, Fe, and Ni) or slightly over (Cr and Co) solar abundances.

(viii) *Copper*: we measured two spectral lines (Cu I 5218.201 and 5782.170 Å) obtaining an average abundance in agreement with the Sun.

(ix) *Zinc*: using the Zn I spectral lines at  $\lambda\lambda$  4680.134, 4722.157, and 6362.338 Å, we derived a distribution of abundances centred on about  $-0.20$  dex.

(x) *Heavier elements: Sr, Y, Zr* solar abundance was derived for strontium (using the Sr II lines  $\lambda\lambda$  4077.709 and 4215.519 Å), while for yttrium (Y II lines  $\lambda\lambda$  4883.682, 4900.120, and 5087.418 Å) and for zirconium (Zr II lines  $\lambda\lambda$  4379.742 and 6114.853 Å) slight and moderate overabundances have been obtained, respectively.

(xi) *Barium*: abundances of barium (inferred by using five Ba II spectral lines at  $\lambda\lambda$  4554.033, 4934.100, 5853.625, 6141.713, and 6496.897 Å) are spread out over a wide interval of values, ranging from approximately from  $\approx -1.0$  dex to  $\approx +1.0$  dex.

(xii) *Rare earth elements: La, Ce, Pr, Nd, Sm, Eu, Gd*: in our sample, lanthanum (La II  $\lambda\lambda$  4921.775 and 5290.818 Å), cerium (Ce II  $\lambda$  4562.359 Å), neodymium (Nd II  $\lambda\lambda$  4959.119, 4989.950, 5092.794, and 5293.163 Å), samarium (Sm II  $\lambda\lambda$  4537.941, 4577.688,

<sup>6</sup>IDL (Interactive Data Language) is a registered trademark of Harris Geospatial Solutions

**Table 2.** Atmospheric parameters. For each spectrum of our targets we report: identification (column 1); heliocentric Julian date at mid exposure (column 2), pulsation phase (column 3), effective temperature (column 4), gravity (column 5), microturbulent, broadening parameter, and radial velocity (columns 6–8).

ID	HJD 2400000.+	$\phi$	$T_{\text{eff}}$ (K)	$\log g$	$\xi$ (km s <sup>-1</sup> )	$v_{\text{br}}$ (km s <sup>-1</sup> )	$v_{\text{rad}}$ (km s <sup>-1</sup> )
ASASSN J180946.70 –182238.2	58648.6242	0.872	5860 ± 170	1.9 ± 0.3	3.8 ± 0.3	15 ± 2	31.7 ± 0.3
	58662.5124	0.777	5480 ± 200	1.6 ± 0.3	4.3 ± 0.3	18 ± 3	51.2 ± 0.6
	58677.5251	0.836	5750 ± 240	1.5 ± 0.4	3.8 ± 0.2	18 ± 2	41.0 ± 0.3
ASAS J052610+1151.3	58809.5630	0.366	5650 ± 220	1.2 ± 0.1	2.5 ± 0.3	13 ± 1	41.3 ± 0.1
	58814.4975	0.532	5590 ± 120	1.3 ± 0.1	2.6 ± 0.3	13 ± 1	50.8 ± 0.1
	58837.4555	0.957	6390 ± 310	1.9 ± 0.1	2.8 ± 0.3	13 ± 1	25.2 ± 0.2
ASAS J061022+1438.6	58809.6053	0.608	5860 ± 140	1.0 ± 0.1	2.6 ± 0.3	13 ± 1	56.6 ± 0.1
	58814.5464	0.629	5900 ± 150	1.1 ± 0.1	2.7 ± 0.3	13 ± 1	56.9 ± 0.1
	58837.4976	0.373	5860 ± 170	0.9 ± 0.1	2.7 ± 0.3	13 ± 1	51.2 ± 0.1
ASAS J063519+2117.8	59152.6212	0.418	6140 ± 180	2.0 ± 0.1	3.7 ± 0.3	15 ± 2	48.3 ± 0.2
	59171.5498	0.934	6650 ± 250	2.3 ± 0.1	3.6 ± 0.2	16 ± 2	26.6 ± 0.2
	59202.7366	0.555	6040 ± 210	1.5 ± 0.1	3.5 ± 0.3	18 ± 1	52.0 ± 0.2
ASAS J065413+0756.5	59152.6940	0.408	6250 ± 290	2.1 ± 0.1	3.3 ± 0.3	13 ± 1	42.5 ± 0.1
	59172.6137	0.032	6540 ± 200	1.9 ± 0.1	3.3 ± 0.2	12 ± 2	20.8 ± 0.2
	59260.3476	0.887	6480 ± 260	2.1 ± 0.1	3.3 ± 0.2	14 ± 1	26.6 ± 0.2
ASAS J070832–1454.5	58809.6666	0.392	5970 ± 190	1.0 ± 0.1	2.4 ± 0.3	15 ± 1	61.4 ± 0.1
	58814.6358	0.170	6130 ± 250	1.2 ± 0.1	2.6 ± 0.3	15 ± 1	56.8 ± 0.1
	58837.5954	0.764	6060 ± 120	1.2 ± 0.1	2.4 ± 0.3	15 ± 1	64.5 ± 0.2
ASAS J070911–1217.2	58809.6873	0.955	6440 ± 240	1.9 ± 0.1	2.5 ± 0.3	18 ± 1	63.4 ± 0.2
	58814.6554	0.015	6410 ± 170	1.8 ± 0.1	2.5 ± 0.3	18 ± 1	60.7 ± 0.2
	58837.6146	0.532	6170 ± 180	1.8 ± 0.1	2.6 ± 0.3	18 ± 1	76.2 ± 0.1
ASAS J072424–0751.3	58809.7098	0.027	6390 ± 210	2.0 ± 0.1	2.6 ± 0.3	22 ± 1	63.3 ± 0.3
	58814.6787	0.426	6030 ± 280	1.7 ± 0.1	2.7 ± 0.3	22 ± 1	78.8 ± 0.3
	58837.6373	0.510	6060 ± 130	1.7 ± 0.1	2.5 ± 0.3	22 ± 1	81.2 ± 0.4
ASAS J074401–1707.8	59152.7297	0.853	6180 ± 250	1.5 ± 0.1	3.2 ± 0.2	10 ± 1	108.1 ± 0.1
	59172.7003	0.743	6530 ± 310	1.3 ± 0.1	3.2 ± 0.2	12 ± 1	94.9 ± 0.2
	59264.4987	0.818	6440 ± 150	1.4 ± 0.1	3.3 ± 0.2	14 ± 2	97.4 ± 0.2
ASAS J074412–1704.9	58809.7553	0.709	5580 ± 240	0.7 ± 0.1	1.8 ± 0.9	19 ± 1	89.7 ± 0.2
	58814.7347	0.768	5810 ± 190	1.3 ± 0.1	2.0 ± 0.9	19 ± 1	86.6 ± 0.2
	58837.6921	0.651	5550 ± 140	0.6 ± 0.1	1.7 ± 0.8	19 ± 1	90.8 ± 0.2
ASAS J162326–0941.0	58648.4844	0.675	6050 ± 170	1.8 ± 0.4	2.9 ± 0.1	11 ± 1	48.2 ± 0.2
	58662.4401	0.801	6170 ± 230	1.8 ± 0.5	2.9 ± 0.1	10 ± 1	40.2 ± 0.2
	58677.4015	0.656	6080 ± 170	2.0 ± 0.4	2.8 ± 0.1	11 ± 1	49.5 ± 0.2
ASAS J180342–2211.0	59029.5474	0.453	4820 ± 180	1.5 ± 0.1	4.8 ± 0.7	13 ± 5	13.4 ± 0.1
	59046.5040	0.851	4930 ± 180	1.6 ± 0.1	4.5 ± 0.2	20 ± 3	27.7 ± 1.9
ASAS J182714–1507.1	59029.5667	0.283	5600 ± 260	0.7 ± 0.1	2.3 ± 0.3	17 ± 1	20.8 ± 0.1
	59046.5280	0.342	5550 ± 200	1.2 ± 0.1	2.6 ± 0.2	17 ± 1	24.1 ± 0.1
ASAS J183347–0448.6	58648.5014	0.918	6450 ± 170	1.7 ± 0.5	3.0 ± 0.4	14 ± 1	–8.1 ± 0.3
	58662.5939	0.461	6120 ± 180	1.5 ± 0.5	2.8 ± 0.4	13 ± 1	–1.4 ± 0.2
	58677.4421	0.247	6280 ± 180	1.1 ± 0.6	2.4 ± 0.1	14 ± 1	–8.4 ± 0.2
ASAS J183652–0907.1	58648.5855	0.409	5990 ± 130	2.2 ± 0.3	2.5 ± 0.3	10 ± 1	–3.9 ± 0.2
	58662.6195	0.827	6250 ± 160	2.0 ± 0.5	2.8 ± 0.3	15 ± 1	–6.8 ± 0.2
	58677.6301	0.623	6090 ± 180	2.7 ± 0.4	2.5 ± 0.3	13 ± 1	–0.1 ± 0.2
ASAS J183904–1049.3	58648.6433	0.132	6240 ± 240	1.8 ± 0.5	2.5 ± 0.3	20 ± 1	–6.9 ± 0.4
	58662.6544	0.714	6140 ± 160	1.6 ± 0.5	2.6 ± 0.3	20 ± 1	5.8 ± 0.3
	58677.5867	0.597	6090 ± 150	1.7 ± 0.4	2.7 ± 0.3	20 ± 1	9.3 ± 0.3
ASAS J192007+1247.7	59029.47931	0.619	5600 ± 130	0.7 ± 0.1	2.4 ± 0.1	18 ± 2	7.1 ± 0.3
	59046.45261	0.586	5500 ± 220	0.7 ± 0.1	2.2 ± 0.1	18 ± 2	9.8 ± 0.1
	59053.65998	0.421	5250 ± 125	0.8 ± 0.1	2.6 ± 0.2	14 ± 1	6.2 ± 0.1
ASAS J192310+1351.4	59029.49690	0.234	6250 ± 130	1.1 ± 0.1	2.1 ± 0.2	12 ± 2	–2.1 ± 0.1
	59053.57714	0.334	6250 ± 220	1.5 ± 0.1	2.1 ± 0.2	11 ± 1	1.0 ± 0.1

**Table 2** – *continued*

ID	HJD 2400000.+	$\phi$	$T_{\text{eff}}$ (K)	$\log g$	$\xi$ ( $\text{km s}^{-1}$ )	$v_{\text{br}}$ ( $\text{km s}^{-1}$ )	$v_{\text{rad}}$ ( $\text{km s}^{-1}$ )
BD+59 12	59152.34090	0.772	$6350 \pm 250$	$1.5 \pm 0.1$	$3.5 \pm 0.3$	$24 \pm 1$	$-28.8 \pm 0.1$
	59167.44631	0.088	$6380 \pm 320$	$1.1 \pm 0.1$	$3.3 \pm 0.2$	$22 \pm 2$	$-44.8 \pm 0.3$
	59200.33040	0.660	$6270 \pm 180$	$1.5 \pm 0.1$	$3.6 \pm 0.4$	$22 \pm 2$	$-31.0 \pm 0.1$
CF Cam	59152.50388	0.274	$6060 \pm 150$	$1.7 \pm 0.1$	$3.2 \pm 0.4$	$10 \pm 1$	$-39.3 \pm 0.1$
	59215.53402	0.953	$6400 \pm 154$	$1.9 \pm 0.1$	$3.3 \pm 0.2$	$12 \pm 2$	$-45.2 \pm 0.1$
	59259.32798	0.595	$5585 \pm 140$	$0.8 \pm 0.1$	$3.0 \pm 0.2$	$11 \pm 1$	$-26.5 \pm 0.1$
DR 2468646563398354176	59152.54715	0.053	$6430 \pm 200$	$0.0 \pm 0.1$	$3.1 \pm 0.2$	$11 \pm 1$	$29.1 \pm 0.1$
	59170.59596	0.088	$6670 \pm 230$	$0.5 \pm 0.1$	$3.3 \pm 0.2$	$11 \pm 2$	$30.0 \pm 0.1$
	59242.45968	0.114	$6520 \pm 210$	$1.2 \pm 0.1$	$3.1 \pm 0.4$	$10 \pm 1$	$30.5 \pm 0.1$
DR 2514736269771300224	59152.39425	0.876	$6450 \pm 280$	$1.4 \pm 0.1$	$3.7 \pm 0.4$	$14 \pm 1$	$-66.3 \pm 0.1$
	59170.30991	0.307	$6050 \pm 190$	$1.1 \pm 0.1$	$3.2 \pm 0.2$	$9 \pm 1$	$-63.4 \pm 0.1$
	59185.62302	0.949	$6470 \pm 280$	$1.3 \pm 0.1$	$3.2 \pm 0.2$	$13 \pm 1$	$-69.6 \pm 0.1$
HD 160473	59029.53160	0.732	$6150 \pm 150$	$1.0 \pm 0.1$	$2.3 \pm 0.2$	$14 \pm 1$	$5.1 \pm 0.1$
	59046.48443	0.217	$6000 \pm 80$	$0.9 \pm 0.1$	$2.5 \pm 0.2$	$12 \pm 1$	$2.9 \pm 0.1$
HO Vul	59029.61249	0.211	$5950 \pm 290$	$1.9 \pm 0.1$	$2.7 \pm 0.2$	$19 \pm 2$	$-12.0 \pm 0.1$
	59046.72463	0.250	$5950 \pm 190$	$1.9 \pm 0.1$	$2.9 \pm 0.5$	$19 \pm 2$	$-9.2 \pm 0.1$
	59053.68819	0.486	$5500 \pm 140$	$1.3 \pm 0.1$	$2.6 \pm 0.3$	$19 \pm 2$	$5.8 \pm 0.2$
	59058.66606	0.370	$5640 \pm 250$	$1.3 \pm 0.1$	$2.5 \pm 0.3$	$20 \pm 2$	$-1.3 \pm 0.2$
OGLE-GD-CEP-0066	59152.65764	0.186	$6460 \pm 200$	$1.4 \pm 0.1$	$3.2 \pm 0.2$	$11 \pm 1$	$40.2 \pm 0.1$
	59172.65751	0.835	$6470 \pm 280$	$1.6 \pm 0.1$	$3.3 \pm 0.2$	$12 \pm 1$	$53.1 \pm 0.2$
	59259.37833	0.342	$6410 \pm 175$	$1.8 \pm 0.1$	$3.2 \pm 0.2$	$9 \pm 1$	$46.1 \pm 0.1$
OGLE-GD-CEP-0104	59197.70252	0.799	$6320 \pm 210$	$1.6 \pm 0.1$	$3.4 \pm 0.2$	$22 \pm 1$	$87.2 \pm 0.2$
	59242.53786	0.387	$6010 \pm 220$	$1.3 \pm 0.1$	$3.1 \pm 0.4$	$16 \pm 1$	$89.5 \pm 0.2$
	59264.45418	0.451	$6130 \pm 225$	$1.4 \pm 0.1$	$3.0 \pm 0.3$	$16 \pm 1$	$93.0 \pm 0.2$
OO Pup	59242.58199	0.065	$5820 \pm 210$	$0.7 \pm 0.1$	$3.8 \pm 0.4$	$14 \pm 1$	$81.2 \pm 0.1$
	59258.52041	0.516	$5205 \pm 160$	$0.9 \pm 0.1$	$4.2 \pm 0.5$	$15 \pm 2$	$109.3 \pm 0.1$
	59259.42289	0.598	$5460 \pm 200$	$0.9 \pm 0.1$	$4.1 \pm 0.5$	$18 \pm 2$	$111.2 \pm 0.1$
	59263.50787	0.970	$5880 \pm 145$	$0.9 \pm 0.1$	$3.6 \pm 0.4$	$11 \pm 2$	$85.1 \pm 0.1$
	59264.54255	0.065	$5900 \pm 275$	$1.2 \pm 0.1$	$3.8 \pm 0.4$	$14 \pm 1$	$81.1 \pm 0.1$
OP Pup	58809.7328	0.710	$6230 \pm 160$	$1.6 \pm 0.1$	$2.1 \pm 0.3$	$11 \pm 1$	$76.0 \pm 0.2$
	58814.7059	0.624	$6200 \pm 230$	$1.8 \pm 0.1$	$2.2 \pm 0.3$	$11 \pm 1$	$76.0 \pm 0.1$
	58837.6642	0.458	$6030 \pm 200$	$1.4 \pm 0.1$	$2.1 \pm 0.3$	$11 \pm 1$	$70.0 \pm 0.1$
OR Cam	58809.5398	0.272	$6000 \pm 200$	$1.3 \pm 0.1$	$2.2 \pm 0.3$	$15 \pm 1$	$-44.1 \pm 0.1$
	58814.5224	0.624	$5960 \pm 140$	$1.5 \pm 0.1$	$2.5 \pm 0.3$	$15 \pm 1$	$-34.5 \pm 0.1$
	58837.4260	0.839	$6180 \pm 270$	$1.6 \pm 0.1$	$2.6 \pm 0.3$	$15 \pm 1$	$-48.7 \pm 0.1$
TX Sct	58648.5236	0.949	$5730 \pm 170$	$1.4 \pm 0.3$	$3.5 \pm 0.3$	$14 \pm 1$	$18.0 \pm 0.2$
	58662.5583	0.525	$4800 \pm 150$	$1.1 \pm 0.1$	$4.1 \pm 0.4$	$13 \pm 1$	$45.9 \pm 0.6$
	58677.5638	0.142	$5200 \pm 170$	$1.0 \pm 0.3$	$3.5 \pm 0.3$	$12 \pm 1$	$18.3 \pm 0.5$
V1495 Aql	59029.4584	0.113	$5750 \pm 200$	$0.4 \pm 0.1$	$2.6 \pm 0.2$	$14 \pm 1$	$37.1 \pm 0.1$
	59046.5736	0.058	$6000 \pm 190$	$0.2 \pm 0.1$	$2.5 \pm 0.2$	$13 \pm 1$	$36.2 \pm 0.1$
	59053.6351	0.861	$6000 \pm 230$	$0.7 \pm 0.1$	$2.7 \pm 0.2$	$13 \pm 1$	$33.8 \pm 0.1$
V1496 Aql	58648.6816	0.350	$5200 \pm 140$	$0.2 \pm 0.3$	$4.4 \pm 0.5$	$13 \pm 2$	$78.4 \pm 0.3$
	58662.6680	0.564	$5050 \pm 150$	$0.1 \pm 0.4$	$4.2 \pm 0.5$	$14 \pm 2$	$87.0 \pm 0.5$
	58677.4830	0.792	$5230 \pm 240$	$0.3 \pm 0.4$	$4.3 \pm 0.6$	$16 \pm 3$	$93.8 \pm 0.7$
V1788 Cyg	59029.6752	0.412	$4950 \pm 150$	$0.9 \pm 0.1$	$4.5 \pm 0.3$	$10 \pm 2$	$-13.1 \pm 0.1$
	59046.6059	0.613	$4800 \pm 120$	$1.0 \pm 0.1$	$5.0 \pm 0.5$	$10 \pm 2$	$4.0 \pm 0.2$
	59053.5364	0.105	$5650 \pm 190$	$1.1 \pm 0.1$	$2.9 \pm 0.2$	$16 \pm 1$	$-40.7 \pm 0.2$
	59066.7076	0.040	$6000 \pm 230$	$1.3 \pm 0.1$	$3.1 \pm 0.2$	$18 \pm 2$	$-40.5 \pm 0.2$
V2475 Cyg	59029.6422	0.926	$6180 \pm 130$	$1.4 \pm 0.1$	$3.6 \pm 0.3$	$13 \pm 2$	$-12.3 \pm 0.1$
	59053.4596	0.987	$6450 \pm 190$	$0.2 \pm 0.1$	$2.5 \pm 0.2$	$13 \pm 1$	$-16.2 \pm 0.3$
	59102.3996	0.222	$5700 \pm 250$	$0.9 \pm 0.1$	$2.9 \pm 0.3$	$12 \pm 1$	$-19.0 \pm 0.1$
V355 Sge	58648.7036	0.579	$4730 \pm 140$	$0.9 \pm 0.2$	$4.1 \pm 0.3$	$14 \pm 2$	$46.0 \pm 1.0$
	58662.6872	0.015	$5340 \pm 120$	$0.7 \pm 0.3$	$3.4 \pm 0.2$	$15 \pm 1$	$16.4 \pm 0.4$
	58677.4521	0.475	$4730 \pm 100$	$0.7 \pm 0.1$	$3.6 \pm 0.5$	$12 \pm 2$	$40.6 \pm 0.6$

Table 2 – continued

ID	HJD 2400000.+	$\phi$	$T_{\text{eff}}$ (K)	$\log g$	$\xi$ (km s <sup>-1</sup> )	$v_{\text{br}}$ (km s <sup>-1</sup> )	$v_{\text{rad}}$ (km s <sup>-1</sup> )
V371 Gem	58809.5859	0.154	6500 ± 200	1.9 ± 0.1	2.2 ± 0.3	10 ± 1	21.5 ± 0.2
	58814.5662	0.484	6150 ± 160	1.6 ± 0.1	2.3 ± 0.3	10 ± 1	33.6 ± 0.1
	58837.4785	0.203	6420 ± 300	1.7 ± 0.1	2.4 ± 0.3	10 ± 1	23.7 ± 0.1
V383 Cyg	59029.7032	0.402	5550 ± 270	1.3 ± 0.1	2.9 ± 0.5	18 ± 2	-9.3 ± 0.1
	59053.4980	0.090	5950 ± 280	1.6 ± 0.1	2.9 ± 0.2	18 ± 2	-15.3 ± 0.1
	59067.6913	0.098	6270 ± 170	1.6 ± 0.1	2.9 ± 0.2	16 ± 1	-32.1 ± 0.1
V389 Sct	59029.5895	0.997	6300 ± 300	0.3 ± 0.1	2.2 ± 0.2	11 ± 1	25.6 ± 0.1
	59046.5486	0.467	6000 ± 230	1.0 ± 0.1	2.4 ± 0.2	10 ± 1	39.3 ± 0.1
	59053.6068	0.328	6000 ± 230	0.8 ± 0.1	2.4 ± 0.2	10 ± 1	33.2 ± 0.1
V536 Ser	58648.5635	0.034	5740 ± 100	1.0 ± 0.4	4.2 ± 0.3	12 ± 1	1.2 ± 0.2
	58662.4887	0.579	5340 ± 120	1.2 ± 0.3	3.7 ± 0.1	12 ± 2	24.1 ± 0.5
	58677.5026	0.246	4730 ± 100	0.6 ± 0.1	3.4 ± 0.9	12 ± 1	40.6 ± 0.5
V598 Per	59152.4608	0.752	5540 ± 140	2.4 ± 0.1	4.4 ± 0.5	16 ± 2	-2.8 ± 0.1
	59200.4233	0.214	6040 ± 200	1.1 ± 0.1	3.2 ± 0.3	13 ± 1	-33.8 ± 0.1
	59242.4884	0.635	5510 ± 140	1.0 ± 0.1	3.5 ± 0.4	14 ± 1	-11.6 ± 0.1
V824 Cas	59152.3638	0.427	5970 ± 180	0.8 ± 0.1	3.4 ± 0.2	11 ± 1	-68.6 ± 0.1
	59167.4690	0.250	6160 ± 190	0.9 ± 0.1	3.5 ± 0.2	11 ± 1	-76.0 ± 0.1
	59200.3533	0.396	6020 ± 220	0.8 ± 0.1	3.5 ± 0.2	10 ± 1	-69.9 ± 0.1
V912 Aql	58648.6623	0.452	5620 ± 150	2.0 ± 0.2	2.9 ± 0.3	8 ± 1	14.4 ± 0.2
	58662.4650	0.589	5570 ± 100	1.6 ± 0.3	3.3 ± 0.3	10 ± 1	21.8 ± 0.2
	58677.4225	0.988	6500 ± 200	1.6 ± 0.5	2.5 ± 0.3	14 ± 1	-5.5 ± 0.2
V914 Mon	58809.6472	0.069	6130 ± 170	1.4 ± 0.1	2.5 ± 0.3	9 ± 1	44.5 ± 0.1
	58814.6125	0.921	6210 ± 220	1.3 ± 0.1	2.5 ± 0.3	9 ± 1	38.5 ± 0.2
	58837.5770	0.488	6360 ± 260	1.7 ± 0.1	2.5 ± 0.3	9 ± 1	45.9 ± 0.2
V946 Cas	59152.4241	0.531	5610 ± 270	0.6 ± 0.1	3.6 ± 0.5	12 ± 1	-4.3 ± 0.1
	59170.3837	0.768	5820 ± 240	1.1 ± 0.1	4.0 ± 0.3	18 ± 2	0.7 ± 0.1
	59200.3861	0.846	6150 ± 230	1.7 ± 0.1	4.2 ± 0.4	18 ± 2	-7.6 ± 0.1
V966 Mon	59187.7141	0.404	5580 ± 150	1.2 ± 0.1	2.5 ± 0.2	13 ± 2	51.4 ± 0.1
	59216.6039	0.453	5530 ± 230	1.4 ± 0.1	2.7 ± 0.2	13 ± 2	53.4 ± 0.1
	59242.6160	0.808	5570 ± 240	0.8 ± 0.1	2.7 ± 0.4	15 ± 3	50.2 ± 0.2
V5567 Sgr	58648.6052	0.103	5670 ± 80	1.5 ± 0.4	3.4 ± 0.4	12 ± 1	-1.4 ± 0.2
	58662.5328	0.529	5400 ± 160	1.7 ± 0.3	3.7 ± 0.4	15 ± 2	24.8 ± 0.2
	58677.5445	0.067	5700 ± 140	2.1 ± 0.3	3.2 ± 0.3	8 ± 1	0.1 ± 0.1
X Sct	58648.5420	0.690	5750 ± 200	2.1 ± 0.3	3.2 ± 0.3	18 ± 2	28.5 ± 0.3
	58662.5769	0.033	6700 ± 240	2.1 ± 0.4	2.8 ± 0.3	15 ± 2	-8.2 ± 0.3
	58677.6049	0.613	5700 ± 220	2.1 ± 0.3	3.0 ± 0.3	17 ± 1	24.5 ± 0.2
ZTF J000234.99+650517.9	59152.3112	0.676	5360 ± 210	0.3 ± 0.1	4.3 ± 0.8	16 ± 1	-56.7 ± 0.1
	59167.4165	0.058	5400 ± 230	0.7 ± 0.1	4.1 ± 0.5	13 ± 1	-70.8 ± 0.1
	59200.3003	0.066	5480 ± 200	0.6 ± 0.1	4.1 ± 0.8	13 ± 1	-71.0 ± 0.1

4642.228, and 4704.400 Å), europium (Eu II  $\lambda$  6645.114 Å) and gadolinium (Gd II  $\lambda$  5092.249 Å) show almost solar value. On the contrary, praseodymium (Pr II  $\lambda$  5219.045 Å) has a distribution of abundances centred on  $\approx -0.45$  dex.

### 3.2.1 Comparison with literature

Abundances from high-resolution spectroscopy are available in the literature for two among our targets: V5567 Sgr and X Sct, from the studies of Martin et al. (2015) and Genovali et al. (2015), respectively. A comparison performed for the elements in common is shown in Fig. 2, where we have adopted as  $\sigma_{\text{TOT}}$  the sum in quadrature between

errors derived in this paper and those reported by the other authors. Almost all of our abundances are in agreement, within the errors, with the literature estimates.

A large data set of abundances derived for 435 DCEPs (based on an analysis of 1127 spectra), was more recently published by Luck (2018). This author derived abundance gradients using *Gaia* DR2 parallax data and drew the reader's attention to a possible anticorrelation between [O/Fe] and iron abundances, i.e. an increase of iron corresponding to a decrease of [O/Fe]. In Fig. 3, we extended and compared our results with those of Luck (2018) to all chemical species in common.

We confirm the anticorrelation found for oxygen. In general our data are in agreement with those of Luck (2018), even if some

**Table 3.** Abundances expressed in solar terms (Grevesse et al. 2010) for the chemical elements detected in our targets.

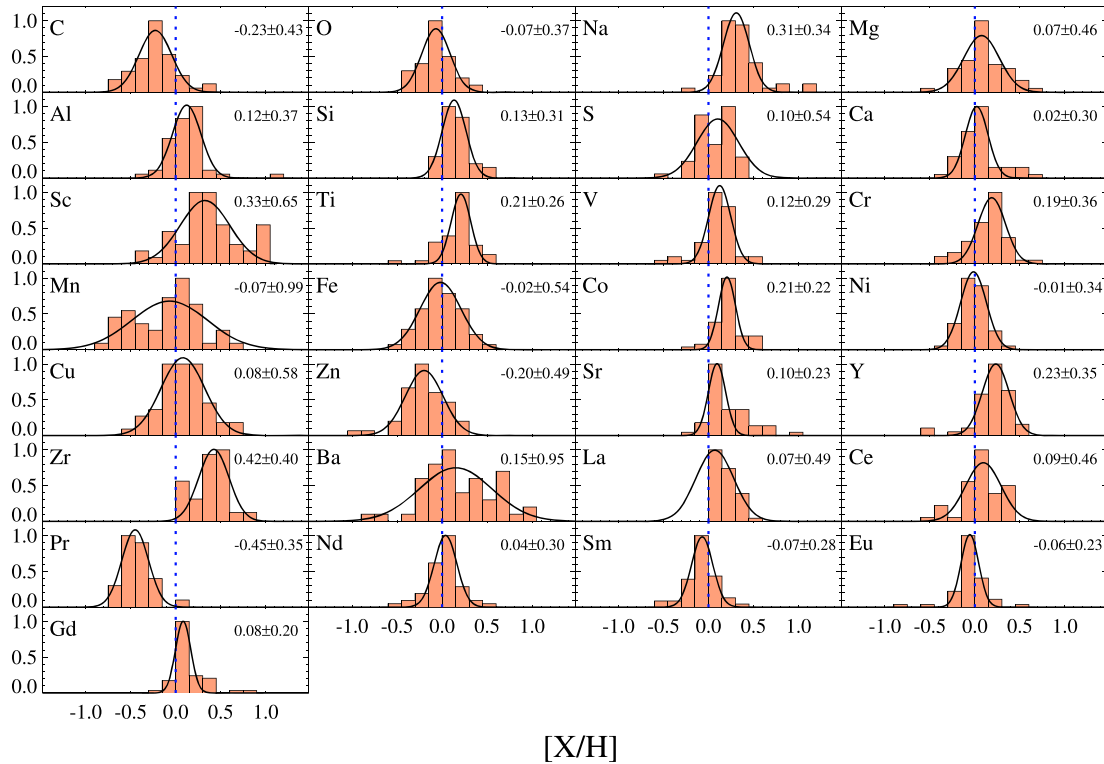
ID	[C/H]	[O/H]	[Na/H]	[Mg/H]	[Al/H]	[Si/H]	[S/H]	[Ca/H]	[Sc/H]	[Ti/H]	[V/H]
ASASSN J180946.70	-0.14 ± 0.12	-0.03 ± 0.10	0.58 ± 0.15	0.04 ± 0.12	0.23 ± 0.20	0.19 ± 0.10	0.26 ± 0.05	0.09 ± 0.15	0.34 ± 0.09	0.17 ± 0.15	0.17 ± 0.11
ASAS 1052610+1151.3	-0.26 ± 0.18	-0.13 ± 0.09	0.29 ± 0.13	0.21 ± 0.13	-0.08 ± 0.10	-0.04 ± 0.16	0.10 ± 0.15	0.08 ± 0.14	0.17 ± 0.13	0.09 ± 0.19	0.14 ± 0.16
ASAS 4061022+1438.6	-0.27 ± 0.10	-0.09 ± 0.15	0.33 ± 0.15	0.36 ± 0.13	0.10 ± 0.10	0.08 ± 0.15	0.10 ± 0.08	-0.01 ± 0.16	0.47 ± 0.18	0.21 ± 0.16	0.22 ± 0.12
ASAS 1063519+2117.8	-0.43 ± 0.18	-0.21 ± 0.15	0.23 ± 0.07	-0.21 ± 0.18	0.18 ± 0.13	0.23 ± 0.17	-0.04 ± 0.14	0.20 ± 0.12	0.20 ± 0.18	0.18 ± 0.11	0.50 ± 0.11
ASAS 4065413+0756.5	-0.45 ± 0.19	-0.01 ± 0.10	0.36 ± 0.13	-0.07 ± 0.12	-0.07 ± 0.13	0.10 ± 0.17	-0.19 ± 0.12	-0.16 ± 0.09	-0.18 ± 0.46	0.06 ± 0.20	0.06 ± 0.20
ASAS 1070832-1454.5	-0.11 ± 0.19	-0.34 ± 0.22	0.25 ± 0.15	0.23 ± 0.10	0.15 ± 0.11	0.01 ± 0.16	0.14 ± 0.14	0.02 ± 0.15	0.37 ± 0.15	0.17 ± 0.10	0.17 ± 0.18
ASAS 1070911-1217.2	-0.17 ± 0.12	-0.19 ± 0.10	0.12 ± 0.07	0.09 ± 0.14	-0.11 ± 0.14	0.00 ± 0.10	0.20 ± 0.16	-0.08 ± 0.16	0.44 ± 0.18	0.16 ± 0.16	0.14 ± 0.14
ASAS 1072424-0751.3	-0.69 ± 0.15	-0.10 ± 0.10	0.27 ± 0.10	0.36 ± 0.15	-0.21 ± 0.12	-0.04 ± 0.12	-0.14 ± 0.13	0.06 ± 0.16	0.58 ± 0.20	0.17 ± 0.18	0.14 ± 0.17
ASAS 1074401-1707.8	-0.46 ± 0.15	-0.20 ± 0.15	0.17 ± 0.10	-0.18 ± 0.18	0.12 ± 0.13	0.06 ± 0.18	-0.14 ± 0.12	-0.22 ± 0.05	-0.06 ± 0.17	-0.07 ± 0.14	0.25 ± 0.11
ASAS 1074412-1704.9	-0.26 ± 0.12	-0.34 ± 0.15	0.15 ± 0.12	0.15 ± 0.12	-0.21 ± 0.18	0.01 ± 0.18	0.27 ± 0.09	0.01 ± 0.07	0.55 ± 0.18	0.19 ± 0.13	0.21 ± 0.09
ASAS 1162326-0941.0	0.42 ± 0.05	-0.09 ± 0.05	0.56 ± 0.10	-0.04 ± 0.10	0.23 ± 0.10	0.19 ± 0.11	0.31 ± 0.15	-0.09 ± 0.10	0.36 ± 0.13	0.16 ± 0.15	0.04 ± 0.10
ASAS 1180342-2211.0	-	0.40 ± 0.17	1.06 ± 0.06	0.02 ± 0.25	1.08 ± 0.13	0.52 ± 0.13	0.25 ± 0.15	0.18 ± 0.06	0.63 ± 0.13	0.36 ± 0.19	0.24 ± 0.14
ASAS 1182714-1507.1	-0.02 ± 0.20	0.15 ± 0.15	0.88 ± 0.12	0.35 ± 0.18	0.25 ± 0.13	0.40 ± 0.12	0.25 ± 0.13	0.59 ± 0.05	1.02 ± 0.24	0.36 ± 0.12	-0.04 ± 0.19
ASAS 1183347-0448.6	-0.20 ± 0.12	-0.04 ± 0.10	0.49 ± 0.13	0.26 ± 0.09	0.23 ± 0.10	0.19 ± 0.13	0.25 ± 0.05	0.03 ± 0.10	0.35 ± 0.15	0.23 ± 0.14	0.14 ± 0.12
ASAS 1183652-0907.1	-0.32 ± 0.10	-0.32 ± 0.10	0.42 ± 0.12	0.26 ± 0.12	0.10 ± 0.10	0.27 ± 0.14	0.26 ± 0.05	0.04 ± 0.10	0.22 ± 0.09	0.21 ± 0.10	0.14 ± 0.13
ASAS 1183904-1049.3	-0.24 ± 0.14	-0.07 ± 0.10	0.38 ± 0.08	0.14 ± 0.12	0.07 ± 0.10	0.12 ± 0.15	0.25 ± 0.13	0.01 ± 0.10	0.31 ± 0.09	0.21 ± 0.15	0.14 ± 0.09
ASAS 1192007+1247.7	0.02 ± 0.15	-0.07 ± 0.14	0.31 ± 0.15	0.56 ± 0.16	0.12 ± 0.13	0.24 ± 0.13	-0.14 ± 0.15	0.58 ± 0.13	1.03 ± 0.25	0.31 ± 0.18	-0.03 ± 0.21
ASAS 1192310+1351.4	-0.39 ± 0.08	0.11 ± 0.11	0.25 ± 0.15	0.22 ± 0.18	0.06 ± 0.13	0.16 ± 0.09	-0.04 ± 0.15	0.12 ± 0.10	0.55 ± 0.37	0.26 ± 0.20	0.12 ± 0.10
BD+59 12	-0.40 ± 0.12	-0.01 ± 0.10	0.25 ± 0.08	-0.19 ± 0.20	0.25 ± 0.13	0.18 ± 0.18	-0.04 ± 0.15	-0.09 ± 0.06	0.05 ± 0.07	0.05 ± 0.10	0.26 ± 0.09
CF Cam	-0.25 ± 0.13	-0.20 ± 0.11	0.19 ± 0.08	-0.28 ± 0.18	0.00 ± 0.13	0.20 ± 0.16	-0.19 ± 0.12	-0.11 ± 0.04	0.24 ± 0.20	0.29 ± 0.18	0.42 ± 0.08
DR 246864656398354176	-0.62 ± 0.13	-0.23 ± 0.15	0.12 ± 0.05	-0.21 ± 0.27	0.25 ± 0.13	0.07 ± 0.14	-0.14 ± 0.15	-0.25 ± 0.10	-0.43 ± 0.17	-0.17 ± 0.16	0.23 ± 0.12
DR 2514736269771300224	-0.27 ± 0.15	0.24 ± 0.32	0.33 ± 0.10	0.01 ± 0.13	0.18 ± 0.13	0.21 ± 0.12	0.06 ± 0.15	-0.01 ± 0.07	0.16 ± 0.14	0.15 ± 0.11	0.19 ± 0.12
HD 160473	-0.17 ± 0.09	0.05 ± 0.12	0.41 ± 0.14	0.19 ± 0.22	0.18 ± 0.13	0.24 ± 0.18	0.15 ± 0.11	0.30 ± 0.08	0.95 ± 0.20	0.25 ± 0.19	0.01 ± 0.14
HO Vul	0.38 ± 0.10	0.11 ± 0.12	0.41 ± 0.13	0.51 ± 0.22	0.06 ± 0.13	0.35 ± 0.18	0.15 ± 0.15	0.47 ± 0.10	0.96 ± 0.21	0.30 ± 0.16	0.10 ± 0.10
OGLE-GD-CEP-0066	-0.46 ± 0.11	-0.01 ± 0.15	0.25 ± 0.08	-0.12 ± 0.14	-0.13 ± 0.13	0.09 ± 0.15	-0.14 ± 0.15	-0.14 ± 0.08	-0.06 ± 0.13	-0.09 ± 0.11	0.26 ± 0.09
OGLE-GD-CEP-0104	-0.47 ± 0.07	-0.40 ± 0.15	0.22 ± 0.15	-0.09 ± 0.26	0.12 ± 0.13	0.07 ± 0.18	-0.24 ± 0.10	-0.08 ± 0.08	-0.01 ± 0.08	-0.02 ± 0.20	0.09 ± 0.21
OO Pup	-0.40 ± 0.12	-0.40 ± 0.12	0.26 ± 0.20	0.04 ± 0.14	0.12 ± 0.13	0.03 ± 0.19	-0.14 ± 0.12	0.03 ± 0.05	0.17 ± 0.16	0.05 ± 0.10	0.08 ± 0.08
OP Pup	-0.31 ± 0.09	-0.17 ± 0.12	0.21 ± 0.13	0.10 ± 0.12	-0.05 ± 0.13	0.05 ± 0.13	-0.14 ± 0.12	-0.12 ± 0.12	0.44 ± 0.14	0.17 ± 0.17	-0.01 ± 0.10
OR Cam	-0.17 ± 0.11	0.00 ± 0.10	0.26 ± 0.17	0.04 ± 0.14	-0.03 ± 0.13	0.03 ± 0.13	0.19 ± 0.15	-0.07 ± 0.12	0.24 ± 0.20	0.13 ± 0.19	0.21 ± 0.14
TX Set	-0.23 ± 0.10	-0.19 ± 0.10	0.66 ± 0.14	-0.04 ± 0.10	0.18 ± 0.05	0.17 ± 0.07	0.31 ± 0.10	0.10 ± 0.10	0.30 ± 0.11	0.18 ± 0.08	0.14 ± 0.14
V1495 Aql	0.10 ± 0.14	0.05 ± 0.15	1.06 ± 0.15	0.71 ± 0.25	0.56 ± 0.13	0.48 ± 0.10	0.35 ± 0.11	0.65 ± 0.04	0.99 ± 0.25	0.57 ± 0.18	0.44 ± 0.20
V1496 Aql	0.12 ± 0.09	0.31 ± 0.10	0.76 ± 0.10	0.14 ± 0.12	0.22 ± 0.10	0.39 ± 0.15	0.41 ± 0.10	0.28 ± 0.10	0.68 ± 0.15	0.26 ± 0.15	0.04 ± 0.11
V1788 Cyg	0.16 ± 0.07	-0.24 ± 0.15	0.44 ± 0.03	0.40 ± 0.16	0.18 ± 0.13	0.50 ± 0.13	0.05 ± 0.15	0.44 ± 0.05	1.05 ± 0.23	0.53 ± 0.20	0.24 ± 0.10
V2475 Cyg	-0.10 ± 0.10	-0.07 ± 0.07	0.53 ± 0.14	0.57 ± 0.22	0.37 ± 0.13	0.39 ± 0.10	0.25 ± 0.15	0.30 ± 0.08	0.90 ± 0.38	0.46 ± 0.15	0.52 ± 0.15
V355 Sge	-0.33 ± 0.11	-0.11 ± 0.10	0.46 ± 0.15	0.06 ± 0.10	0.21 ± 0.05	0.17 ± 0.11	0.16 ± 0.10	0.13 ± 0.10	0.58 ± 0.15	0.18 ± 0.15	0.19 ± 0.16
V371 Gem	-0.29 ± 0.08	-0.03 ± 0.13	0.33 ± 0.11	0.10 ± 0.13	-0.08 ± 0.11	0.04 ± 0.12	0.14 ± 0.15	0.07 ± 0.15	0.07 ± 0.15	0.14 ± 0.13	0.10 ± 0.05
V383 Cyg	-0.09 ± 0.14	-0.01 ± 0.17	0.13 ± 0.15	0.41 ± 0.18	0.03 ± 0.13	0.20 ± 0.17	-0.04 ± 0.15	0.09 ± 0.16	0.75 ± 0.35	0.34 ± 0.19	0.25 ± 0.13
V389 Ser	-0.27 ± 0.08	0.21 ± 0.15	0.53 ± 0.13	0.41 ± 0.18	0.18 ± 0.13	0.26 ± 0.18	0.15 ± 0.15	0.24 ± 0.07	0.65 ± 0.26	0.39 ± 0.17	0.29 ± 0.15
V536 Ser	-0.14 ± 0.05	-0.14 ± 0.10	0.60 ± 0.06	0.08 ± 0.05	0.25 ± 0.05	0.19 ± 0.10	0.35 ± 0.05	0.13 ± 0.10	0.31 ± 0.07	0.21 ± 0.10	0.11 ± 0.13
V5567 Sgr	-0.15 ± 0.12	0.10 ± 0.10	0.25 ± 0.10	-0.06 ± 0.07	0.10 ± 0.06	0.13 ± 0.10	0.19 ± 0.08	0.00 ± 0.07	0.11 ± 0.05	0.14 ± 0.12	0.04 ± 0.09
V598 Per	-0.31 ± 0.11	0.05 ± 0.14	0.39 ± 0.20	-0.04 ± 0.25	0.37 ± 0.13	0.18 ± 0.14	-0.04 ± 0.16	0.03 ± 0.05	0.31 ± 0.23	0.13 ± 0.20	0.13 ± 0.21
V824 Cas	-0.39 ± 0.08	0.18 ± 0.10	0.45 ± 0.17	0.09 ± 0.24	0.25 ± 0.13	0.14 ± 0.19	-0.04 ± 0.12	0.02 ± 0.04	0.26 ± 0.07	0.11 ± 0.10	0.07 ± 0.15
V1612 Aql	-0.10 ± 0.10	-0.05 ± 0.04	0.32 ± 0.10	0.01 ± 0.12	0.20 ± 0.06	0.14 ± 0.10	0.24 ± 0.13	0.04 ± 0.10	0.17 ± 0.07	0.23 ± 0.15	0.09 ± 0.13
V1614 Mon	-0.18 ± 0.10	-0.09 ± 0.15	0.24 ± 0.17	-0.13 ± 0.15	-0.05 ± 0.07	0.06 ± 0.18	0.09 ± 0.09	-0.17 ± 0.17	0.13 ± 0.10	-0.02 ± 0.18	0.16 ± 0.13
V946 Cas	-0.65 ± 0.08	-0.15 ± 0.15	0.22 ± 0.15	-0.20 ± 0.18	-0.07 ± 0.13	0.04 ± 0.15	-0.54 ± 0.15	-0.22 ± 0.05	-0.10 ± 0.16	-0.15 ± 0.19	0.29 ± 0.12
V966 Mon	-0.49 ± 0.13	-0.39 ± 0.15	0.00 ± 0.08	-0.53 ± 0.18	-0.13 ± 0.13	-0.08 ± 0.13	-0.19 ± 0.15	-0.38 ± 0.10	-0.37 ± 0.28	-0.46 ± 0.08	-0.42 ± 0.17
X Set	-0.14 ± 0.12	-0.04 ± 0.10	0.30 ± 0.18	0.06 ± 0.13	0.09 ± 0.10	0.09 ± 0.15	0.16 ± 0.10	-0.09 ± 0.10	0.47 ± 0.10	0.22 ± 0.15	0.14 ± 0.15
ZTF J000034+650517.9	-0.30 ± 0.19	0.11 ± 0.15	0.36 ± 0.23	-0.12 ± 0.17	0.00 ± 0.13	0.11 ± 0.15	-0.14 ± 0.15	0.06 ± 0.05	0.28 ± 0.20	-0.02 ± 0.18	0.33 ± 0.15
ASASSN J180946.70	0.23 ± 0.16	0.20 ± 0.15	0.15 ± 0.11	0.10 ± 0.10	0.12 ± 0.10	0.24 ± 0.10	0.17 ± 0.13	0.12 ± 0.09	0.30 ± 0.19	0.40 ± 0.40	0.08 ± 0.15

Table 3 – continued

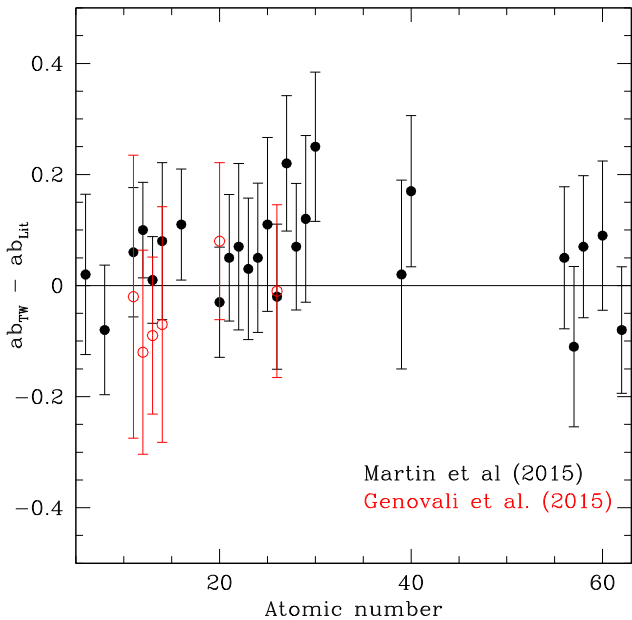
ID	[C/H]	[O/H]	[Na/H]	[Mg/H]	[Al/H]	[Si/H]	[S/H]	[Ca/H]	[Sc/H]	[Ti/H]	[V/H]
ASAS 1052610+1151.3	0.12 ± 0.18	0.01 ± 0.16	-0.06 ± 0.10	0.12 ± 0.09	-0.08 ± 0.20	-0.07 ± 0.10	0.06 ± 0.14	-0.10 ± 0.15	0.30 ± 0.15	0.34 ± 0.14	1.04 ± 0.07
ASAS 1061022+1438.6	0.15 ± 0.10	0.12 ± 0.19	0.01 ± 0.11	0.17 ± 0.10	0.02 ± 0.18	0.13 ± 0.10	-0.10 ± 0.13	0.10 ± 0.12	0.32 ± 0.13	0.50 ± 0.14	0.49 ± 0.40
ASAS 1063519+2117.8	0.09 ± 0.13	-0.51 ± 0.08	-0.18 ± 0.15	0.53 ± 0.26	-0.04 ± 0.16	0.27 ± 0.15	-0.31 ± 0.15	0.44 ± 0.21	0.04 ± 0.12	0.56 ± 0.15	0.10 ± 0.16
ASAS 1065413+0756.5	0.08 ± 0.10	-0.64 ± 0.14	-0.27 ± 0.15	0.15 ± 0.15	-0.12 ± 0.20	0.11 ± 0.13	-0.43 ± 0.11	0.30 ± 0.19	0.23 ± 0.12	0.30 ± 0.04	-0.23 ± 0.16
ASAS 1070832-1454.5	0.06 ± 0.10	0.24 ± 0.12	-0.11 ± 0.11	0.20 ± 0.10	-0.07 ± 0.10	0.16 ± 0.12	-0.19 ± 0.11	0.09 ± 0.11	0.18 ± 0.12	0.53 ± 0.16	0.66 ± 0.10
ASAS 1070911-1217.2	0.10 ± 0.13	0.32 ± 0.11	-0.11 ± 0.11	-0.08 ± 0.15	-0.08 ± 0.15	0.23 ± 0.10	-0.26 ± 0.12	0.07 ± 0.10	0.17 ± 0.14	0.40 ± 0.13	0.90 ± 0.15
ASAS 1072424-0751.3	0.15 ± 0.19	-0.06 ± 0.16	-0.04 ± 0.11	0.19 ± 0.14	-0.03 ± 0.13	0.15 ± 0.07	0.03 ± 0.10	0.17 ± 0.12	0.19 ± 0.08	0.74 ± 0.12	0.65 ± 0.13
ASAS 1074401-1707.8	-0.20 ± 0.11	-0.74 ± 0.06	-0.41 ± 0.15	0.19 ± 0.16	-0.30 ± 0.13	0.02 ± 0.15	-0.57 ± 0.11	0.25 ± 0.21	-0.10 ± 0.17	0.30 ± 0.15	-0.22 ± 0.22
ASAS 1074412-1704.9	0.04 ± 0.11	-0.02 ± 0.12	-0.18 ± 0.10	0.25 ± 0.14	-0.17 ± 0.12	0.24 ± 0.10	-0.25 ± 0.07	0.07 ± 0.15	0.00 ± 0.07	0.54 ± 0.06	0.71 ± 0.19
ASAS 1162326-0941.0	0.15 ± 0.15	0.05 ± 0.10	-0.03 ± 0.11	0.10 ± 0.10	0.12 ± 0.15	0.32 ± 0.10	0.11 ± 0.03	0.10 ± 0.15	0.30 ± 0.15	0.42 ± 0.12	0.26 ± 0.15
ASAS 1180342-2211.0	0.37 ± 0.15	-0.31 ± 0.10	0.10 ± 0.15	0.27 ± 0.23	0.14 ± 0.12	-0.06 ± 0.15	0.22 ± 0.11	0.68 ± 0.07	0.42 ± 0.10	0.78 ± 0.18	-0.05 ± 0.09
ASAS 1182714-1507.1	0.33 ± 0.11	0.52 ± 0.09	0.36 ± 0.14	0.30 ± 0.23	0.23 ± 0.11	0.14 ± 0.10	0.08 ± 0.19	0.20 ± 0.10	0.22 ± 0.05	0.30 ± 0.04	-0.71 ± 0.10
ASAS 1183347-0448.6	0.27 ± 0.09	-0.05 ± 0.13	0.20 ± 0.11	0.22 ± 0.11	0.11 ± 0.14	0.38 ± 0.10	-0.04 ± 0.10	0.08 ± 0.10	0.22 ± 0.07	0.60 ± 0.18	0.41 ± 0.15
ASAS 1183904-1049.3	0.19 ± 0.17	-0.02 ± 0.15	0.04 ± 0.11	0.22 ± 0.15	0.00 ± 0.09	0.13 ± 0.10	-0.16 ± 0.08	0.07 ± 0.10	0.25 ± 0.14	0.48 ± 0.10	0.61 ± 0.15
ASAS 1192007+1247.7	0.29 ± 0.13	0.49 ± 0.07	0.31 ± 0.14	0.07 ± 0.32	0.00 ± 0.15	-0.04 ± 0.09	0.06 ± 0.23	0.30 ± 0.15	0.54 ± 0.24	0.00 ± 0.04	0.10 ± 0.24
ASAS 1192310+1351.4	0.17 ± 0.12	-0.24 ± 0.20	0.03 ± 0.15	0.11 ± 0.16	-0.08 ± 0.15	-0.50 ± 0.15	-0.27 ± 0.08	0.57 ± 0.17	0.14 ± 0.06	0.10 ± 0.09	-0.15 ± 0.16
BD+59 12	-0.01 ± 0.15	-0.54 ± 0.05	-0.10 ± 0.14	0.46 ± 0.18	-0.13 ± 0.13	0.27 ± 0.15	-0.42 ± 0.14	0.22 ± 0.09	0.36 ± 0.06	0.43 ± 0.15	-0.30 ± 0.29
CF Cam	0.32 ± 0.13	-0.40 ± 0.16	-0.17 ± 0.15	0.57 ± 0.10	-0.05 ± 0.14	0.65 ± 0.15	-0.37 ± 0.10	-0.20 ± 0.07	0.17 ± 0.04	0.60 ± 0.29	0.42 ± 0.16
DR 246864656398354176	-0.38 ± 0.17	-0.80 ± 0.07	-0.45 ± 0.11	0.32 ± 0.23	-0.19 ± 0.19	0.15 ± 0.12	-0.57 ± 0.16	-0.07 ± 0.36	-0.50 ± 0.10	0.40 ± 0.04	-0.16 ± 0.30
DR 2514736269771300224	0.09 ± 0.13	-0.32 ± 0.05	-0.09 ± 0.15	0.38 ± 0.23	-0.04 ± 0.10	0.71 ± 0.15	-0.33 ± 0.14	0.39 ± 0.15	0.06 ± 0.31	0.30 ± 0.16	0.07 ± 0.24
HD 160473	0.24 ± 0.19	0.21 ± 0.04	0.23 ± 0.15	0.21 ± 0.28	0.12 ± 0.08	-0.04 ± 0.27	-0.09 ± 0.12	0.40 ± 0.15	0.29 ± 0.10	0.30 ± 0.04	0.49 ± 0.18
HO Vul	0.32 ± 0.13	0.46 ± 0.04	0.20 ± 0.15	0.18 ± 0.21	0.11 ± 0.10	-0.09 ± 0.12	-0.28 ± 0.23	0.55 ± 0.07	0.35 ± 0.05	0.10 ± 0.15	0.33 ± 0.20
OGLE-GD-CEP-0066	0.03 ± 0.18	-0.53 ± 0.06	-0.25 ± 0.15	0.23 ± 0.20	-0.10 ± 0.14	0.15 ± 0.15	-0.43 ± 0.10	0.20 ± 0.23	0.23 ± 0.19	0.15 ± 0.05	0.03 ± 0.18
OGLE-GD-CEP-0104	-0.07 ± 0.08	-0.75 ± 0.10	-0.27 ± 0.13	0.03 ± 0.17	-0.25 ± 0.13	-0.10 ± 0.15	-1.03 ± 0.15	0.30 ± 0.07	-0.20 ± 0.14	0.52 ± 0.04	-0.16 ± 0.22
OO Pup	-0.12 ± 0.14	-0.51 ± 0.12	-0.18 ± 0.15	0.15 ± 0.12	-0.20 ± 0.15	-0.10 ± 0.15	-0.44 ± 0.11	0.10 ± 0.20	-0.10 ± 0.15	0.05 ± 0.04	-0.11 ± 0.28
OP Pup	0.11 ± 0.19	0.06 ± 0.20	-0.11 ± 0.13	0.22 ± 0.14	-0.08 ± 0.15	-0.17 ± 0.20	-0.27 ± 0.12	0.09 ± 0.15	0.21 ± 0.10	0.50 ± 0.11	0.41 ± 0.20
OR Cam	0.19 ± 0.19	-0.01 ± 0.16	-0.06 ± 0.09	0.17 ± 0.12	-0.03 ± 0.18	0.48 ± 0.11	-0.16 ± 0.10	0.17 ± 0.15	0.27 ± 0.08	0.59 ± 0.11	0.47 ± 0.17
TX Set	0.32 ± 0.06	0.21 ± 0.10	0.11 ± 0.11	0.22 ± 0.09	0.02 ± 0.09	0.32 ± 0.10	-0.03 ± 0.10	0.09 ± 0.15	0.30 ± 0.10	0.30 ± 0.12	0.31 ± 0.15
V1495 Aql	0.47 ± 0.10	0.64 ± 0.10	0.55 ± 0.10	0.50 ± 0.25	0.38 ± 0.19	0.35 ± 0.12	0.24 ± 0.14	0.29 ± 0.15	0.45 ± 0.09	0.52 ± 0.04	0.06 ± 0.26
V1496 Aql	0.37 ± 0.11	0.17 ± 0.08	-0.07 ± 0.11	0.32 ± 0.13	0.13 ± 0.11	0.08 ± 0.10	0.06 ± 0.07	0.11 ± 0.15	0.40 ± 0.13	0.48 ± 0.10	0.00 ± 0.15
V1788 Cyg	0.61 ± 0.13	0.02 ± 0.08	0.37 ± 0.15	0.48 ± 0.20	0.24 ± 0.17	0.36 ± 0.15	-0.23 ± 0.42	0.55 ± 0.07	0.50 ± 0.15	-	0.71 ± 0.14
V2475 Cyg	0.34 ± 0.12	-0.03 ± 0.14	0.25 ± 0.15	0.17 ± 0.20	0.20 ± 0.15	-0.23 ± 0.15	-0.09 ± 0.22	0.10 ± 0.15	0.29 ± 0.06	0.10 ± 0.09	-0.02 ± 0.10
V355 Sge	0.23 ± 0.15	0.00 ± 0.14	0.03 ± 0.11	0.22 ± 0.10	0.02 ± 0.18	0.17 ± 0.15	-0.04 ± 0.14	0.12 ± 0.10	0.49 ± 0.19	0.30 ± 0.15	0.41 ± 0.10
V371 Gem	0.06 ± 0.19	0.06 ± 0.17	-0.06 ± 0.13	0.23 ± 0.12	-0.09 ± 0.09	0.11 ± 0.17	-0.19 ± 0.19	0.07 ± 0.15	0.19 ± 0.17	0.54 ± 0.16	0.67 ± 0.06
V383 Cyg	0.27 ± 0.13	-0.08 ± 0.20	0.17 ± 0.15	0.42 ± 0.36	0.06 ± 0.12	-0.37 ± 0.12	-0.27 ± 0.13	0.74 ± 0.10	0.11 ± 0.06	0.15 ± 0.04	-0.08 ± 0.10
V389 Ser	0.21 ± 0.12	-0.06 ± 0.20	0.22 ± 0.15	0.27 ± 0.13	0.17 ± 0.19	-0.01 ± 0.31	-0.15 ± 0.08	0.61 ± 0.17	0.29 ± 0.18	0.77 ± 0.04	0.10 ± 0.16
V536 Ser	0.27 ± 0.15	0.11 ± 0.20	0.17 ± 0.15	0.27 ± 0.23	0.02 ± 0.11	0.27 ± 0.15	-0.38 ± 0.15	0.39 ± 0.18	0.07 ± 0.44	0.32 ± 0.11	0.00 ± 0.10
V567 Sgr	0.10 ± 0.10	0.00 ± 0.14	0.05 ± 0.11	0.17 ± 0.10	0.07 ± 0.09	-0.02 ± 0.12	0.04 ± 0.10	0.11 ± 0.10	0.27 ± 0.15	0.32 ± 0.11	0.00 ± 0.10
V598 Per	0.12 ± 0.12	0.24 ± 0.12	0.07 ± 0.15	0.19 ± 0.20	0.02 ± 0.11	0.27 ± 0.15	-0.37 ± 0.15	0.39 ± 0.18	0.04 ± 0.27	0.15 ± 0.04	0.07 ± 0.20
V824 Cas	0.16 ± 0.18	-0.28 ± 0.06	-0.08 ± 0.15	0.27 ± 0.23	-0.02 ± 0.11	0.46 ± 0.14	-0.37 ± 0.15	-0.01 ± 0.11	0.04 ± 0.27	0.02 ± 0.10	0.00 ± 0.10
V612 Aql	0.22 ± 0.18	0.15 ± 0.11	0.14 ± 0.11	0.10 ± 0.10	0.01 ± 0.18	0.08 ± 0.08	-0.16 ± 0.15	0.12 ± 0.11	0.30 ± 0.14	0.39 ± 0.15	0.28 ± 0.15
V614 Mon	0.01 ± 0.12	-0.45 ± 0.14	-0.16 ± 0.10	0.11 ± 0.14	-0.05 ± 0.20	0.11 ± 0.07	-0.29 ± 0.12	0.07 ± 0.19	0.12 ± 0.20	0.39 ± 0.14	0.17 ± 0.11
V946 Cas	0.30 ± 0.10	-0.75 ± 0.10	-0.43 ± 0.12	-0.06 ± 0.20	-0.24 ± 0.17	-0.17 ± 0.11	-0.46 ± 0.13	-0.05 ± 0.19	0.04 ± 0.15	0.02 ± 0.04	-0.17 ± 0.09
V966 Mon	0.33 ± 0.17	-0.59 ± 0.14	-0.60 ± 0.15	-0.22 ± 0.26	-0.37 ± 0.10	-0.35 ± 0.13	-0.82 ± 0.15	0.99 ± 0.20	-0.60 ± 0.10	0.02 ± 0.04	-0.86 ± 0.12
X Xct	0.07 ± 0.15	0.07 ± 0.15	0.50 ± 0.11	0.27 ± 0.10	0.01 ± 0.15	-0.06 ± 0.10	-0.16 ± 0.10	0.10 ± 0.04	0.35 ± 0.15	0.54 ± 0.15	0.06 ± 0.20
ZTF J0002234+34517.9	-0.14 ± 0.14	-0.51 ± 0.50	-0.15 ± 0.12	0.00 ± 0.20	-0.21 ± 0.14	0.08 ± 0.15	-0.41 ± 0.11	0.42 ± 0.24	0.23 ± 0.15	0.52 ± 0.04	-0.36 ± 0.28

Table 3 – continued

ID	[La/H]	[Ce/H]	[Pr/H]	[Nd/H]	[Sm/H]	[Eu/H]	[Gd/H]
ASASSN J180946.70	0.18 ± 0.06	0.41 ± 0.15	–	0.12 ± 0.16	0.00 ± 0.20	–0.01 ± 0.04	0.37 ± 0.25
ASAS 052610+1151.3	0.28 ± 0.24	0.32 ± 0.12	–0.33 ± 0.31	0.14 ± 0.10	0.07 ± 0.16	–0.04 ± 0.14	0.03 ± 0.06
ASAS 061022+1438.6	0.07 ± 0.08	0.16 ± 0.14	–	0.04 ± 0.15	–0.18 ± 0.08	–0.08 ± 0.02	0.00 ± 0.10
ASAS 063519+2117.8	0.23 ± 0.23	0.41 ± 0.13	–0.33 ± 0.12	0.34 ± 0.13	–0.02 ± 0.23	–0.05 ± 0.30	0.30 ± 0.14
ASAS 0656413+0756.5	0.08 ± 0.13	–0.07 ± 0.15	–	–0.04 ± 0.17	–0.25 ± 0.18	–0.27 ± 0.13	0.20 ± 0.10
ASAS 070832–1454.5	0.18 ± 0.04	0.16 ± 0.14	–0.60 ± 0.05	0.09 ± 0.12	–0.06 ± 0.05	–0.03 ± 0.04	0.03 ± 0.06
ASAS 070911–1217.2	0.41 ± 0.12	0.26 ± 0.15	–	0.18 ± 0.11	0.02 ± 0.13	0.18 ± 0.07	0.10 ± 0.10
ASAS 072424–0751.3	0.17 ± 0.15	0.06 ± 0.20	–	0.14 ± 0.10	–0.08 ± 0.12	0.05 ± 0.13	0.10 ± 0.10
ASAS 074401–1707.8	0.09 ± 0.20	–0.15 ± 0.04	–0.30 ± 0.17	–0.17 ± 0.13	–0.12 ± 0.13	–0.15 ± 0.15	0.10 ± 0.10
ASAS 074412–1704.9	0.19 ± 0.19	0.25 ± 0.13	–	0.04 ± 0.15	–0.50 ± 0.26	–0.49 ± 0.21	–0.27 ± 0.32
ASAS J162326–0941.0	0.23 ± 0.08	0.10 ± 0.15	–0.50 ± 0.10	–0.07 ± 0.18	–0.05 ± 0.05	–0.13 ± 0.13	0.90 ± 0.10
ASAS J180342–2211.0	0.57 ± 0.10	0.13 ± 0.08	–0.65 ± 1.06	0.28 ± 0.04	0.38 ± 0.04	0.56 ± 0.05	0.60 ± 0.71
ASAS J182714–1507.1	0.07 ± 0.22	–0.32 ± 0.15	–0.47 ± 0.23	–0.12 ± 0.19	–0.40 ± 0.14	–0.12 ± 0.11	–0.05 ± 0.07
ASAS J183347–0448.6	0.26 ± 0.07	0.04 ± 0.15	–	–0.01 ± 0.16	–0.05 ± 0.05	0.03 ± 0.06	0.40 ± 0.26
ASAS J183652–0907.1	0.15 ± 0.05	0.10 ± 0.15	–	0.02 ± 0.16	–0.17 ± 0.06	–0.11 ± 0.07	0.10 ± 0.10
ASAS J183904–1049.3	0.14 ± 0.13	0.05 ± 0.15	–	–0.22 ± 0.17	–0.15 ± 0.05	–0.03 ± 0.02	0.23 ± 0.23
ASAS J192007+1247.7	0.08 ± 0.18	–0.07 ± 0.15	–	0.15 ± 0.13	–0.03 ± 0.25	–0.04 ± 0.17	–0.03 ± 0.15
ASAS J192310+1351.4	0.12 ± 0.12	0.14 ± 0.09	–0.45 ± 0.07	0.05 ± 0.15	–0.10 ± 0.07	–0.15 ± 0.14	0.25 ± 0.21
BD+59 12	0.11 ± 0.07	–0.27 ± 0.04	–0.33 ± 0.21	–0.13 ± 0.06	–0.23 ± 0.12	–0.05 ± 0.05	0.33 ± 0.21
CF Cam	0.32 ± 0.21	0.23 ± 0.15	–0.17 ± 0.21	0.38 ± 0.13	0.07 ± 0.21	0.03 ± 0.13	0.10 ± 0.10
DR 2468646593398354176	0.18 ± 0.33	–0.60 ± 0.15	–0.53 ± 0.42	–0.14 ± 0.08	–0.28 ± 0.23	–0.28 ± 0.25	–0.05 ± 0.05
DR 2514736269771300224	0.19 ± 0.10	0.10 ± 0.15	–0.50 ± 0.06	–0.07 ± 0.08	–0.10 ± 0.09	0.00 ± 0.13	0.10 ± 0.10
HD 160473	0.14 ± 0.04	–0.07 ± 0.15	–	0.00 ± 0.19	–0.10 ± 0.07	–0.04 ± 0.05	0.10 ± 0.10
HO Vul	0.39 ± 0.20	–0.15 ± 0.04	–	0.06 ± 0.07	0.00 ± 0.17	0.20 ± 0.11	0.27 ± 0.29
OGLE-GD-CEP-0066	0.31 ± 0.12	0.10 ± 0.04	–0.23 ± 0.35	–0.15 ± 0.12	–0.03 ± 0.12	–0.03 ± 0.18	0.23 ± 0.23
OGLE-GD-CEP-0104	0.10 ± 0.15	–0.07 ± 0.15	–	–0.35 ± 0.18	–0.22 ± 0.18	–0.09 ± 0.17	0.03 ± 0.06
OO Pup	0.18 ± 0.13	–0.40 ± 0.04	–0.66 ± 0.31	–0.17 ± 0.20	–0.02 ± 0.04	–0.03 ± 0.13	0.04 ± 0.13
OP Pup	0.18 ± 0.12	0.36 ± 0.12	–0.33 ± 0.12	0.14 ± 0.11	–0.10 ± 0.10	–0.07 ± 0.13	0.10 ± 0.10
OR Cam	0.20 ± 0.04	0.31 ± 0.14	–	0.09 ± 0.12	–0.12 ± 0.08	0.00 ± 0.06	0.10 ± 0.10
TX Set	0.02 ± 0.26	0.10 ± 0.15	–	0.14 ± 0.20	0.00 ± 0.16	0.00 ± 0.18	0.10 ± 0.10
V1495 Aql	0.02 ± 0.17	0.01 ± 0.15	–	0.04 ± 0.14	–0.37 ± 0.12	0.06 ± 0.04	0.00 ± 0.10
V1496 Aql	0.07 ± 0.11	0.10 ± 0.15	–	0.15 ± 0.18	–0.14 ± 0.09	0.12 ± 0.08	0.07 ± 0.06
V1788 Cyg	0.34 ± 0.10	0.40 ± 0.04	–0.63 ± 0.35	0.53 ± 0.14	0.17 ± 0.12	0.18 ± 0.08	0.10 ± 0.10
V2475 Cyg	0.33 ± 0.17	0.10 ± 0.04	–0.33 ± 0.21	–0.05 ± 0.19	0.15 ± 0.31	0.12 ± 0.19	0.10 ± 0.10
V355 Sge	0.10 ± 0.17	0.11 ± 0.15	–0.53 ± 0.31	0.18 ± 0.20	0.17 ± 0.15	0.00 ± 0.08	0.40 ± 0.36
V371 Gem	0.25 ± 0.09	0.16 ± 0.13	–0.33 ± 0.12	0.24 ± 0.15	–0.03 ± 0.06	–0.03 ± 0.15	0.10 ± 0.10
V383 Cyg	0.14 ± 0.07	0.10 ± 0.04	–	–0.10 ± 0.04	–0.07 ± 0.23	–0.10 ± 0.05	0.20 ± 0.23
V389 Ser	0.12 ± 0.15	0.14 ± 0.09	–0.57 ± 0.06	0.04 ± 0.17	–0.20 ± 0.17	–0.06 ± 0.12	0.00 ± 0.17
V536 Ser	0.14 ± 0.16	–0.04 ± 0.10	–	–0.07 ± 0.12	0.00 ± 0.17	–0.02 ± 0.12	0.30 ± 0.35
V5567 Sgr	0.21 ± 0.12	0.31 ± 0.10	–0.55 ± 0.07	0.15 ± 0.10	–0.05 ± 0.07	–0.16 ± 0.05	0.05 ± 0.07
V598 Per	0.14 ± 0.10	0.07 ± 0.15	–0.37 ± 0.32	0.03 ± 0.16	0.00 ± 0.35	0.40 ± 0.31	0.10 ± 0.10
V824 Cas	0.23 ± 0.08	0.15 ± 0.10	–	0.07 ± 0.14	–0.07 ± 0.03	0.50 ± 0.05	0.10 ± 0.10
V912 Aql	0.08 ± 0.11	0.31 ± 0.15	–0.53 ± 0.23	0.24 ± 0.20	0.10 ± 0.10	0.17 ± 0.10	0.13 ± 0.15
V914 Mon	0.18 ± 0.17	0.32 ± 0.15	–0.60 ± 0.70	–0.06 ± 0.15	–0.08 ± 0.13	–0.06 ± 0.04	0.10 ± 0.10
V946 Cas	0.34 ± 0.20	0.15 ± 0.04	–0.40 ± 0.10	0.06 ± 0.11	0.08 ± 0.20	–0.12 ± 0.24	0.07 ± 0.06
V966 Mon	0.42 ± 0.16	–0.40 ± 0.10	–	–0.60 ± 0.13	–0.49 ± 0.04	–0.90 ± 0.44	0.13 ± 0.12
X Set	0.12 ± 0.34	0.10 ± 0.20	–	–0.16 ± 0.21	0.30 ± 0.26	–0.26 ± 0.27	–0.03 ± 0.15
ZTF J000234+340000	0.12 ± 0.10	–0.40 ± 0.04	0.00 ± 0.10	–0.40 ± 0.15	–0.28 ± 0.06	–0.23 ± 0.06	0.00 ± 0.10



**Figure 1.** Histograms of the abundance distribution for each chemical element. In each panel, a Gaussian fit of the abundances has been overplotted (solid line) and the blue dotted line indicates the solar value. For clarity, histograms have been normalized to the unity. The centre of the Gaussian and its FWHM are also reported in the top-right corner of each panel.



**Figure 2.** Comparison between the abundances derived in this study and those available in the literature for two of our targets. Filled black circles refer to V5567 Sgr (Martin et al. 2015) and open red circles to X Sct (Genovali et al. 2015).

elements such as scandium, titanium, zinc, strontium, and zirconium are confined at the lower edge of Luck’s distribution. The coefficients of the linear fit plotted in Fig 3 are reported in Table A1 as well as the scatter of the data.

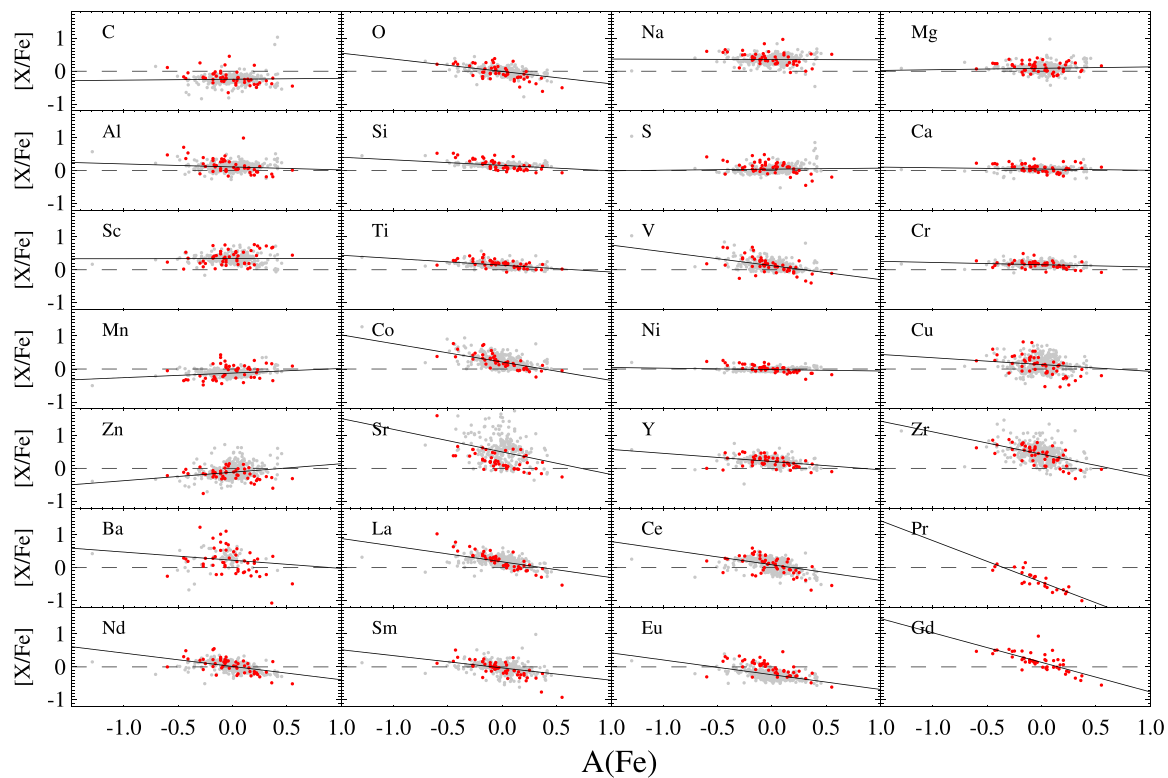
We also compared the distribution of our abundances in the context of the Galactic metallicity gradient. In Fig. 4, we reported the results by Luck (2018) with overimposed the abundances of our DCEP sample. Our abundances are in agreement with the general trend. The coefficients of the linear fit plotted in Fig 4 are reported in Table A2 as well as the scatter of the data.

#### 4 COMPLEMENTARY PHOTOMETRY AND REDDENING

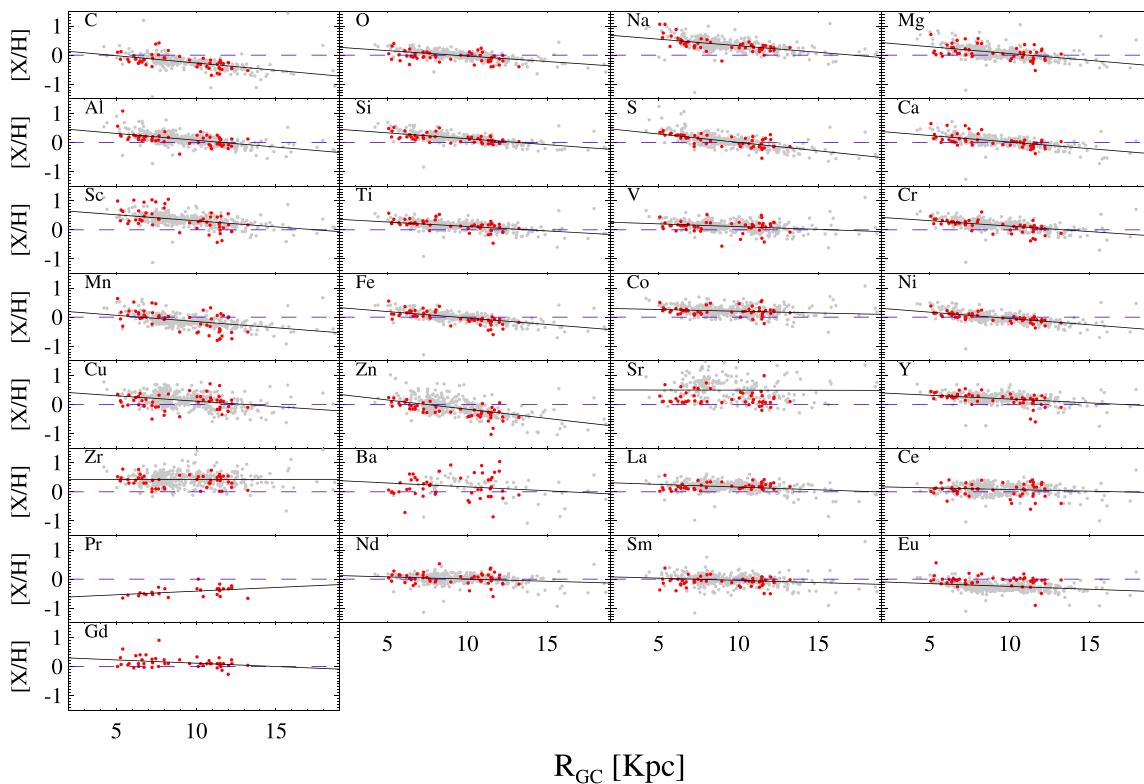
We collected multiband photometry available in the literature for our program stars, in order to accurately determine the pulsation period, time of maximum light (epoch), reddening  $E(B - V)$  and intensity averaged magnitudes of each target. These quantities are needed to construct the PL/PW relations. References for the literature photometry are provided in the last columns of Table 4. Our use of the different photometric data sets is discussed in the following sections.

##### 4.1 Period and epoch of maximum light estimates from optical photometry

An accurate determination of the period and time of maximum light is crucial to correctly estimate the average magnitude of the program stars, particularly in the  $J$ ,  $H$ ,  $K_s$  bands, as only single epoch NIR photometry is often available in the literature (see Section 4.2). Periods are available in the literature for almost all our targets. However, more accurate period (and related error) estimates can be obtained by merging the photometry of different surveys, in this way spanning much longer time intervals than it is covered by each single data set. Johnson  $V$ -band photometry spanning more than 6000 d is available for the majority of our stars by combining the ASAS (All



**Figure 3.** Comparison of the abundances derived in this work expressed as  $[X/Fe]$  (red filled dots) with those derived for 435 Galactic DCEPs by Luck (2018) (light grey symbols). The solid line is the best fit computed considering the literature and our sample of stars (coefficients of the line fit are reported in Table A1), while the dashed line represents the solar standard abundance. All abundances are plotted as a function of the iron abundance.



**Figure 4.** Galactic metallicity gradient for our sample of DECPs (red filled circles) compared with the gradient traced by 435 Galactic DCEPs studied in Luck (2018) (light grey symbols). The solid line is the best fit computed considering the literature and our sample of stars (coefficients of the line fit are reported in Table A2), while the dashed line represents the solar standard abundance.

**Table 4.** Results of the photometric analysis of the targets. The meaning of the different columns is: (1) stars' name; (2) mode of pulsation; (3) and (4) period of pulsation and relative error; (5) and (6) epoch of maximum light and relative error; (7) and (8) pulsation amplitude in the V band and relative error; (9)–(22) intensity-averaged magnitude and relative errors in the *B*, *V*, *I*, *J*, *H*, *Ks* bands, respectively; (23) literature sources for the optical photometry.

Star	Mode	P	P <sub>ERR</sub>	E <sup>MAX</sup>	E <sup>MAX</sup> <sub>ERR</sub>	ampV	ampV <sub>ERR</sub>	B	B <sub>ERR</sub>	V	V <sub>ERR</sub>
(1)	(2)	(d)	(d)	(d)	(d)	(mag)	(mag)	(mag)	(mag)	(mag)	(mag)
ASASSN J180946.70-182238.2	DCEP_F	7.28954	1.3e-04	57053.15	2.0e-02	0.692	0.010	–	–	11.456	0.010
ASAS J052610+1151.3	DCEP_F	4.232016	1.2e-05	51964.8488	1.0e-04	0.604	0.010	12.046	0.019	10.946	0.010
ASAS J061022+1438.6	DCEP_IO	4.83807	2.0e-05	52604.26	7.0e-02	0.258	0.010	–	–	9.880	0.010
ASAS J063519+2117.8	DCEP_IO	1.512365	2.0e-06	52617.06	2.0e-02	0.320	0.010	–	–	11.453	0.010
ASAS J065413+0756.5	DCEP_IO	0.7774053	4.0e-07	52547.541	6.0e-03	0.302	0.010	–	–	11.362	0.010
ASAS J070832-1454.5	DCEP_IO	6.38771	4.0e-05	51844.56	1.3e-01	0.144	0.010	10.809	0.016	9.696	0.010
ASAS J070911-1217.2	DCEP_IO	2.412282	9.0e-06	51860.011	2.0e-03	0.280	0.010	11.544	0.019	10.603	0.010
ASAS J072424-0751.3	DCEP_IO	2.071284	4.0e-06	51862.56755	1.0e-05	0.224	0.010	11.092	0.020	10.352	0.010
ASAS J074401-1707.8	DCEP_IO	2.629941	3.0e-06	51859.00045	2.0e-05	0.445	0.010	–	–	11.510	0.010
ASAS J074412-1704.9	DCEP_IO	4.7013	3.0e-05	51853.20	3.0e-02	0.267	0.010	–	–	10.196	0.010
ASAS J162326-0941.0	DCEP_IO	1.378236	3.0e-06	51931.41005	1.0e-05	0.261	0.010	10.607	0.010	9.817	0.010
ASAS J180342-2211.0	DCEP_F	42.6272	1.2e-03	51806.225	1.4e-02	0.785	0.010	–	–	10.907	0.010
ASAS J182714-1507.1	DCEP_F	5.54565	2.0e-05	51946.2010	3.0e-04	0.642	0.010	–	–	10.692	0.010
ASAS J183347-0448.6	DCEP_IO	3.102301	8.0e-06	51953.99	2.0e-02	0.300	0.010	12.354	0.021	10.846	0.010
ASAS J183652-0907.1	DCEP_IO	2.589889	7.0e-06	51955.25	1.4e-02	0.236	0.010	12.973	0.013	11.600	0.010
ASAS J183904-1049.3	DCEP_IO	3.05774	1.0e-05	51951.79	5.0e-02	0.228	0.010	11.913	0.037	10.590	0.010
ASAS J192007+1247.7	DCEP_F	8.62797	4.0e-05	52673.95390	4.0e-05	0.423	0.010	–	–	10.359	0.010
ASAS J192310+1351.4	DCEP_IO	2.972999	5.0e-06	52702.25806	7.0e-05	0.395	0.010	–	–	10.956	0.010
BD+59_12	DCEP_IO	2.8417	7.0e-05	56998.98	4.0e-02	0.200	0.010	–	–	10.217	0.010
CF Cam	DCEP_F	9.4361	3.0e-04	57007.925	2.0e-03	0.573	0.010	–	–	12.118	0.010
DR2 468646563398354176	DCEP_IO	2.99108	5.0e-05	56998.81	3.0e-02	0.374	0.010	–	–	11.245	0.010
DR2 514736269771300224	DCEP_IO	3.29896	3.0e-05	56998.5814	2.0e-04	0.355	0.010	–	–	10.789	0.010
HD 160473	DCEP_IO	3.779702	1.4e-05	51928.4838	2.0e-04	0.180	0.010	10.680	0.005	9.490	0.010
HO Vul	DCEP_F	5.630938	1.2e-05	52716.14412	8.0e-05	0.718	0.010	–	–	11.653	0.010
OGLE-GD-CEP-0066	DCEP_IO	2.312167	4.0e-06	51972.95	3.0e-02	0.395	0.010	–	–	11.504	0.010
OGLE-GD-CEP-0104	DCEP_IO	2.177743	2.0e-06	51863.502	7.0e-03	0.336	0.010	–	–	11.508	0.010
OO Pup	DCEP_F	10.98279	3.0e-05	51839.4660	3.0e-04	0.587	0.010	–	–	11.568	0.010
OP Pup	DCEP_IO	2.598824	3.0e-06	51858.62722	1.0e-05	0.439	0.010	–	–	11.022	0.010
OR Cam	DCEP_IO	3.6852	2.0e-04	57024.90	6.0e-02	0.286	0.010	–	–	10.402	0.010
TX Sct	DCEP_F	24.3522	2.0e-04	51879.853	2.0e-03	0.873	0.010	–	–	12.422	0.010
V1495 Aql	DCEP_F	8.79653	4.0e-05	51938.46232	1.0e-05	0.629	0.010	13.771	0.034	11.932	0.010
V1496 Aql	DCEP_F	65.83	7.0e-02	56474.6	1.0e-01	0.543	0.010	12.449	0.028	10.226	0.010
V1788 Cyg	DCEP_F	14.09383	8.0e-05	52582.99	1.2e-02	0.753	0.010	–	–	12.709	0.010
V2475 Cyg	DCEP_F	11.5559	3.0e-04	57042.879	3.0e-03	0.954	0.010	–	–	12.549	0.010
V355 Sge	DCEP_F	32.071	6.0e-06	52721.69	2.0e-02	0.660	0.012	14.966	0.010	12.280	0.010
V371 Gem	DCEP_IO	2.137457	9.0e-06	52614.91	3.0e-02	0.381	0.010	–	–	10.753	0.010
V383 Cyg	DCEP_F	4.612297	6.0e-06	46974.62802	1.0e-05	0.547	0.010	12.512	0.010	10.883	0.010
V389 Sct	DCEP_IO	3.793899	6.0e-06	51950.18481	6.0e-05	0.379	0.010	–	–	11.324	0.010
V536 Ser	DCEP_F	9.00985	5.0e-05	51917.9	1.1e-01	0.514	0.010	–	–	11.780	0.010
V5567 Sgr	DCEP_F	9.76289	5.0e-05	51920.97	7.0e-02	0.558	0.010	–	–	10.560	0.010
V598 Per	DCEP_F	5.66819	4.0e-05	56988.6154	6.0e-04	0.809	0.010	–	–	12.396	0.010
V824 Cas	DCEP_IO	5.35066	2.0e-05	50294.73783	1.0e-05	0.315	0.010	12.535	0.015	11.224	0.010
V912 Aql	DCEP_F	4.4004	7.0e-06	48552.15605	1.0e-05	0.707	0.010	13.431	0.013	11.359	0.010
V914 Mon	DCEP_IO	2.680631	1.5e-05	52542.15	4.0e-02	0.393	0.010	–	–	10.762	0.010
V946 Cas	DCEP_IO	4.238789	1.0e-05	52715.69	3.0e-02	0.480	0.010	–	–	11.544	0.010
V966 Mon	DCEP_IO/2O	1.068037	2.0e-06	52523.80	2.0e-02	0.281	0.010	–	–	11.728	0.010
X Sct	DCEP_F	4.198051	6.0e-06	51949.75476	1.0e-05	0.843	0.010	11.169	0.020	10.009	0.010
ZTF J000234.99+650517.9	DCEP_F	10.9323	6.0e-04	57002.19	7.0e-02	0.340	0.010	–	–	11.519	0.010

Star	<i>I</i>	<i>I</i> <sub>ERR</sub>	<i>J</i>	<i>J</i> <sub>ERR</sub>	<i>H</i>	<i>H</i> <sub>ERR</sub>	<i>K<sub>s</sub></i>	<i>K<sub>s</sub></i> <sub>ERR</sub>	<i>E(B – V)</i>	<i>E(B – V)</i> <sub>ERR</sub>	Source
(13)	(mag)	(mag)	(mag)	(mag)	(mag)	(mag)	(mag)	(mag)	(mag)	(mag)	(24)
ASASSN J180946.70–182238.2	–	–	7.421	0.031	6.727	0.051	6.392	0.023	1.113	0.075	a
ASAS J052610+1151.3	9.706	0.019	8.724	0.020	8.334	0.033	8.118	0.020	0.506	0.036	a,b,c,d
ASAS J061022+1438.6	–	–	7.643	0.030	7.222	0.047	7.129	0.029	0.400	0.075	a,c
ASAS J063519+2117.8	–	–	9.997	0.029	9.676	0.028	9.578	0.026	0.234	0.075	a,c
ASAS J065413+0756.5	–	–	10.060	0.029	9.732	0.031	9.604	0.029	0.322	0.075	a,c
ASAS J070832–1454.5	8.349	0.015	7.224	0.029	6.797	0.048	6.636	0.026	0.434	0.034	a,c,e
ASAS J070911–1217.2	9.474	0.015	8.476	0.028	8.111	0.039	7.922	0.033	0.405	0.039	a,c,e
ASAS J072424–0751.3	9.511	0.026	8.844	0.036	8.531	0.043	8.440	0.030	0.216	0.036	a,c,e
ASAS J074401–1707.8	–	–	10.078	0.031	9.758	0.030	9.651	0.028	0.141	0.075	a,c
ASAS J074412–1704.9	–	–	8.778	0.031	8.324	0.041	8.209	0.045	0.166	0.075	c
ASAS J162326–0941.0	8.964	0.014	8.351	0.039	8.042	0.051	7.931	0.031	0.293	0.034	a,c,e
ASAS J180342–2211.0	–	–	5.779	0.021	4.966	0.042	4.484	0.033	1.492	0.075	a,c
ASAS J182714–1507.1	–	–	7.850	0.020	7.308	0.020	7.124	0.024	0.706	0.075	a,c
ASAS J183347–0448.6	8.979	0.019	7.449	0.033	7.005	0.035	6.765	0.029	0.955	0.036	a,c,e
ASAS J183652–0907.1	9.560	0.010	8.249	0.028	7.685	0.029	7.445	0.031	0.983	0.034	a,c,e
ASAS J183904–1049.3	8.925	0.040	7.793	0.029	7.282	0.027	7.087	0.026	0.779	0.040	a,c,e
ASAS J192007+1247.7	–	–	6.404	0.021	5.905	0.038	5.574	0.017	1.139	0.075	a,c

**Table 4** – *continued*

Star	Mode	P	P <sub>ERR</sub>	E <sup>MAX</sup>	E <sub>ERR</sub> <sup>MAX</sup>	ampV	ampV <sub>ERR</sub>	B	B <sub>ERR</sub>	V	V <sub>ERR</sub>
(1)	(2)	(d)	(d)	(d)	(d)	(mag)	(mag)	(mag)	(mag)	(mag)	(mag)
ASAS J192310+1351.4	–	–	7.416	0.030	6.903	0.039	6.689	0.026	1.001	0.075	a,c
BD+59_12	–	–	8.149	0.037	7.759	0.029	7.620	0.030	0.419	0.075	a
CF Cam	–	–	8.445	0.022	7.819	0.024	7.543	0.024	1.003	0.075	a
DR2 468646563398354176	–	–	9.027	0.032	8.487	0.029	8.323	0.028	0.572	0.075	a
DR2 514736269771300224	–	–	8.486	0.035	8.038	0.028	7.864	0.029	0.513	0.075	a
HD 160473	7.901	0.010	6.756	0.030	6.254	0.045	6.048	0.027	0.651	0.033	a,c,e
HO Vul	–	–	7.825	0.019	7.217	0.029	6.996	0.026	1.188	0.075	a,c
OGLE-GD-CEP-0066	10.522	0.010	9.645	0.029	9.198	0.029	9.035	0.029	0.375	0.049	a,c,f
OGLE-GD-CEP-0104	10.390	0.010	9.791	0.033	9.364	0.030	9.188	0.029	0.397	0.046	a,c,f
OO Pup	–	–	8.813	0.030	8.260	0.055	8.039	0.023	0.608	0.075	a,c,i
OP Pup	–	–	9.046	0.031	8.697	0.030	8.606	0.033	0.294	0.075	a,c
OR Cam	–	–	7.014	0.028	6.455	0.026	6.211	0.026	1.035	0.075	a
TX Sct	–	–	7.004	0.029	6.049	0.027	5.638	0.021	1.635	0.075	c
V1495 Aql	9.765	0.017	8.230	0.026	7.571	0.053	7.302	0.024	1.127	0.044	a,c,g
V1496 Aql	7.756	0.019	5.951	0.023	5.302	0.029	4.918	0.023	1.161	0.037	a,c,g,h
V1788 Cyg	–	–	7.639	0.039	6.722	0.026	6.402	0.017	1.553	0.075	a,i
V2475 Cyg	–	–	8.287	0.023	7.502	0.024	7.173	0.020	1.249	0.075	a
V355 Sge	9.291	0.011	7.049	0.020	6.165	0.016	5.819	0.016	1.671	0.034	a,c,g
V371 Gem	–	–	8.774	0.028	8.484	0.029	8.358	0.027	0.328	0.075	c
V383 Cyg	8.964	0.072	7.479	0.020	6.900	0.031	6.672	0.016	1.024	0.038	a,i,j,k
V389 Sct	–	–	7.575	0.056	7.016	0.051	6.737	0.026	1.040	0.075	a,c,i
V536 Ser	–	–	7.782	0.023	7.091	0.040	6.775	0.020	1.249	0.075	a,c
V5567 Sgr	–	–	7.356	0.027	6.734	0.053	6.439	0.023	0.836	0.075	a,c
V598 Per	–	–	8.354	0.024	7.636	0.031	7.383	0.017	1.246	0.075	a
V824 Cas	–	–	8.109	0.030	7.658	0.047	7.397	0.028	0.675	0.044	a,e
V912 Aql	8.879	0.012	7.118	0.026	6.448	0.042	6.149	0.018	1.473	0.034	a,c,g,j
V914 Mon	–	–	8.981	0.027	8.573	0.033	8.430	0.031	0.330	0.075	c
V946 Cas	–	–	8.563	0.044	8.075	0.034	7.899	0.035	0.709	0.075	a,i
V966 Mon	–	–	9.443	0.031	9.069	0.031	8.920	0.030	0.575	0.075	a,c
X Sct	8.622	0.020	7.457	0.030	6.978	0.057	6.769	0.027	0.591	0.036	c,g
ZTF J000234.99+650517.9	–	–	7.812	0.026	7.119	0.026	6.891	0.017	1.056	0.075	a

Note. a, ASASSN (Shappee et al. 2014; Kochanek et al. 2017); b, Berdnikov et al. (2011); c, ASAS (Pojmanski 2002); d, Schmidt, Rogalla & Thacker-Lynn (2011); e, Berdnikov et al. (2009); f, Udalski et al. (2018); g, Berdnikov et al. (2015); h, Berdnikov et al. (2004); i, Alfonso-Garzón et al. (2012); j, Schmidt et al. (2005); k, Berdnikov (2008).

Sky Automated Survey Pojmanski 2002) and ASASSN (Shappee et al. 2014; Kochanek et al. 2017) surveys, which we complemented with our own observations used to reproduce the abundances of other literature data (see Table 4), when possible. Small zero-point differences on the order of hundredths of magnitude are present among different catalogues. They were fixed as follows: when the Johnson *B*, *V*, *I* photometry by Berdnikov et al. (2004, 2009, 2011, 2015) was available, we used these studies as reference (to keep the information on colour); in all other cases our reference was the ASAS catalogue.

A new determination of the period for all our targets was made by applying to the *V*-band time-series data the same multistep procedure as the one devised for the processing and specific characterization of the DCEPs and RR Lyrae stars observed by the *Gaia* mission (Clementini et al. 2016, 2019). That is, a first estimate of the pulsation frequency was obtained with the Lomb–Scargle (LS) algorithm (Lomb 1976; Scargle 1982), by selecting the frequency with the highest power peak in the periodogram. Then, the observations were folded using the selected frequency  $\nu$  and modeled by means of a truncated Fourier series:

$$V(\phi_k) = A_0 + \sum_{i=1}^{n_{\max}} A_i \cdot \sin(i2\pi\nu\phi_k + \Phi_i), \quad (1)$$

where  $\phi_k$  and  $V_k$  are the phases and the magnitudes of the  $k$ th observation and the sum runs over the maximum number of harmonics of the model. To obtain the  $A_0$ ,  $A_i$ , and  $\Phi_i$  values we first run a linear least-squares fitting to the data (since we assumed the frequency from the LS procedure). Starting from the linear estimate of  $A_0$ ,  $A_i$ ,

and  $\Phi_i$ , we then refined the value of the  $\nu^{lim}$  frequency by fitting the following truncated Fourier series:

$$V(t_k) = A_0 + \sum_{i=1}^{n_{\max}} A_i \cdot \sin(i2\pi\nu t_k + \Phi_i), \quad (2)$$

where the frequency  $\nu$  was now left free to vary. In this way, the least-square fit became nonlinear and was treated by means of a standard Levenberg–Marquardt algorithm (Marquardt 1963), thus obtaining more accurate pulsation frequencies, amplitudes and phases. To estimate the uncertainty on the derived parameters, we performed 1000 bootstrap simulations, where the time-series data were resampled and the nonlinear fitting procedure was repeated for every simulation. The robust standard deviation ( $\sigma_{\text{robust}} = 1.4826 \cdot \text{MAD}$ ) of the resulting frequency, amplitudes, and phase was adopted as an estimate of the uncertainty in these quantities. Finally, the time of maximum light of each target was calculated as the Heliocentric Julian day (HJD) of the brightest value of the *V*-band light curve which was closer to the HJD of the first observations minus three times the final pulsation period (inverse of the best-fitting frequency).

The *V*-band light curves of our 48 program stars obtained by folding the literature photometry according to the new periods and epochs of maximum light derived in our study are shown in Fig. B1.

The procedure to estimate the pulsation period and the epoch of our targets naturally provides also their *V*-band intensity-averaged magnitudes and peak-to-peak amplitudes along with the associated errors. All these quantities are presented in Table 4.

In addition to the *V*-band photometry, time-series in other bands of interest for the distance scale, i.e. Johnson–Cousins *B*, *I* bands also

available for several targets (see Table 4). We adopted the periods derived from the  $V$  band photometry to model the light curves and derive the intensity-averaged  $B$ ,  $I$  magnitudes, which are provided in Table 4.

## 4.2 NIR photometry

Time series  $JHK_s$  photometry is available in the literature only for X Sct (Monson & Pierce 2011) among our targets. For the remaining stars we used the single-epoch photometry from 2MASS (Skrutskie et al. 2006),<sup>7</sup> along with a template fitting procedure as implemented by Soszyński, Gieren & Pietrzyński (2005). A crucial point for the template fitting technique is the accuracy of the available periods and epochs. It is also very important that only a few pulsation cycles separate the epochs of the variable stars and the 2MASS observations, in order to minimize errors in phasing the NIR photometry. We verified that phasing errors  $[(P_{\text{err}} * n_{\text{cycles}})/P]$  are negligible for all our stars (the maximum value being  $\sim 0.01$  for ASASSN J180946.70-182238.2).

The other quantity required by the template procedure is the  $V$ -band amplitude, which we have derived for all our targets as described in Section 4.1. We calculated intensity averaged  $JHK_s$  magnitudes for the target stars, assigning an uncertainty equal to the sum in quadrature of the 2MASS uncertainty on the epoch photometry plus 0.02 mag for all the stars, to take into account the average uncertainty of the template fitting procedure (Soszyński et al. 2005, estimate an average error of 0.03 mag for their method. This value already includes the uncertainty of the single-epoch photometry. To avoid summing that latter contribution twice we adopted an average uncertainty of 0.02 mag for the template fitting procedure). Final values for the  $J$ ,  $H$  and  $K_s$  mean magnitudes are summarized in Table 4.

## 4.3 Reddening

Obtaining correct values of reddening for DCEPs is mandatory to construct and use the PL relations (the PW ones are reddening-free by definition). Many different methods have been developed in the past decades to face this problem (see e.g. Fernie 1967; Dean, Warren & Cousins 1978; Pel 1978; Tammann, Sandage & Reindl 2003; Kovtyukh et al. 2008; Andrievsky et al. 2012; Kashuba et al. 2016; Turner 2016, and references therein).

Reddening estimates are present in the literature only for V5567 Sgr ( $E(B - V) = 0.960 \pm 0.096$  mag,<sup>8</sup> Martin et al. 2015) and X Sct ( $E(B - V) = 0.581 \pm 0.030$  mag,<sup>9</sup> Fernie et al. 1995), among our 48 targets. For the remaining objects, the reddening was inferred from the period–colour (PC) relations, which are known to offer a reliable and straightforward method, widely used in the literature (see e.g. Tammann et al. 2003).

The most commonly used PC relations relying on the  $(B - V)$  and  $(V - I)$  colours were published by Tammann et al. (2003). We decided to determine our own relations by using the moderately larger DCEP sample by Ngeow et al. (2012) whose  $E(B - V)$  values are mainly based on the compilation of Fernie et al. (1995), multiplied by 0.95 as

in Tammann et al. (2003), so that our results will be basically compatible with those Authors. The sample by Ngeow (2012) consists of 349 usable objects, including first overtone pulsators, whose periods were fundamentalised using the relation  $P_F = P_{1O}/(0.716 - 0.027 \log P_{1O})$ , where  $P_F$  and  $P_{1O}$  are the fundamental and first overtone DCEPs (DCEP\_F and DCEP\_1O) periods, respectively (Feast & Catchpole 1997). Finally, we performed a least-square fit to the data adopting a  $\sigma$ -clipping technique with a  $2.5\sigma$  threshold. The resulting PCs are:

$$(B - V)_0 = (0.395 \pm 0.010)(\log P - 1) + (0.716 \pm 0.003) \quad (3)$$

$$(V - I)_0 = (0.318 \pm 0.012)(\log P - 1) + (0.757 \pm 0.003) \quad (4)$$

where the colours were dereddened using the Cardelli, Clayton & Mathis (1989) law with a ratio of total to selective absorption of  $R_V = 3.1$ . Equations (3) and (4) have rms = 0.051 and 0.054 mag, respectively, i.e. are less dispersed than to the corresponding relations by Tammann et al. (2003). Note also the use of the pivoting period at 10 d to reduce the correlation between slopes and intercepts. These equations compare with the analogous relations by Tammann et al. (2003) as follows:  $\Delta(B - V)_0 = 0.029 \log P - 0.04$ ;  $\Delta(V - I)_0 = 0.062 \log P - 0.058$ , where  $\Delta$  stands for ‘this work – Tammann et al. (2003)’. It can be seen that the results given in the two cases are compatible within the uncertainties, hence in the following we will use our relations.

As shown in Table 4,  $B$  and  $I$  photometry is missing for 32 and 31 targets, respectively. We can estimate the  $(V - I)$  colour of these sources (as well as of the whole data set) using the *Gaia* EDR3 magnitudes after properly converted them into  $(V - I)$ . To this aim, we used the Ngeow (2012) sample, along with the *Gaia* DR2 intensity-averaged  $G$ ,  $G_{BP}$  and  $G_{RP}$  magnitudes (Clementini et al. 2019; Ripepi et al. 2019) to derive the following conversion equation<sup>10</sup>:

$$(V - I)_G = G - G_{RP} + (0.5383 \pm 0.003) + (0.5492 \pm 0.006)\alpha + (0.1126 \pm 0.014)\alpha^2, \quad (5)$$

where  $(V - I)_G$  is the colour estimated from the *Gaia* photometry and  $\alpha = [(G_{BP} - G_{RP}) - 1.5]$ . The rms of equation (5) is 0.022 mag.

In the end, for 15 DCEPs with available  $B$ ,  $V$ ,  $I$ , we were able to obtain three estimates of the  $E(B - V)$ , using the  $(B - V)$ , the  $(V - I)$  and the  $(V - I)_G$  colours, separately. For additional two sources we calculated two reddening values, depending on the availability of the  $(B - V)$  and  $(V - I)$  colours, respectively. For these 17 sources we adopted a weighted mean of the three (two) measures as our best estimate of  $E(B - V)$ . For the remaining 31 stars, the  $E(B - V)$  reddening could be estimated only from the  $(V - I)_G$  colour, hence carries errors about twice than for the previous 12 sources. The individual reddening values are listed in Table 4.

<sup>10</sup>The correct average magnitude of a DCEP is normally calculated by integrating the light curve in intensity and then transformed back into magnitude. This is the procedure applied to derive the mean magnitudes of Cepheids and RR Lyrae stars published in the ‘vari’ tables of *Gaia* (Clementini et al. 2016, 2019). The *Gaia* magnitudes of the ‘gaia\_source’ general catalog are instead calculated through weighted arithmetic means and can differ from the ‘correct’ ones by several hundredths of magnitude (see e.g. Caputo, Marconi & Ripepi 1999). However, this effect becomes of the order of 1–2 per cent when using colours instead of magnitudes, because the ‘delta’ magnitude, albeit decreasingly large from the bluest to reddest filters, has the same sign in all bands, therefore when subtracting two magnitudes to calculate a colour, the value of ‘delta’ becomes very small, and generally negligible with respect to the uncertainty introduced by the use of the period–colour relation.

<sup>7</sup>We note that 47 over 48 targets have an ‘AAA’ quality flag in the 2MASS single-epoch photometry, indicating fully usable measurements, with the exception of ASAS J180342–2211.0 that has an ‘E’ flag in the  $H$  band. Therefore, for this star the  $H$  value is less accurate than for the others.

<sup>8</sup>Value scaled by 0.97, according to Groenewegen (2018)

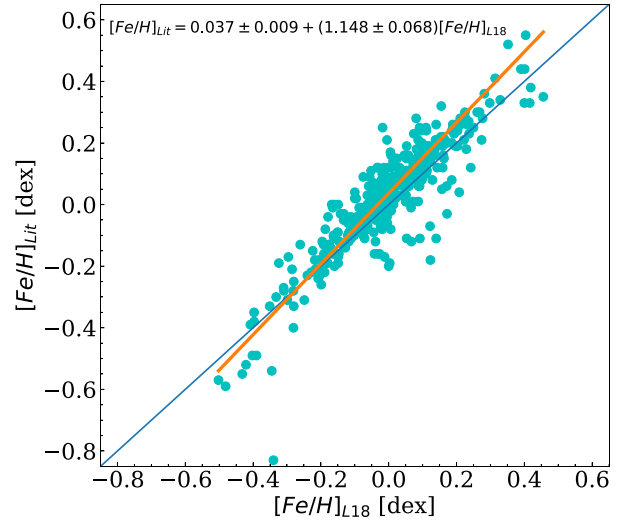
<sup>9</sup>Value scaled by 0.94, according to Groenewegen (2018)

## 5 LITERATURE SAMPLE

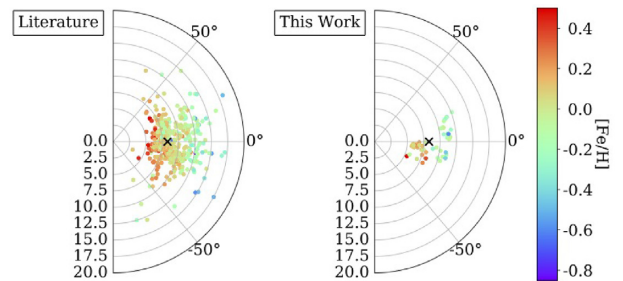
The main literature sources we adopt for our study is the updated catalog by Groenewegen (2018), from which we took the  $V$ ,  $J$ ,  $H$ ,  $K_s$  magnitudes as well as  $E(B-V)$  and  $[Fe/H]$  values on a uniform scale for 438 DCEPs.<sup>11</sup> This sample is complemented by  $V$ ,  $I$ ,  $J$ ,  $H$ ,  $K_s$  magnitudes,  $E(B-V)$  and  $[Fe/H]$  values compiled by Gaia Collaboration (2017). For a few stars we used photometry by van Leeuwen et al. (2007) and Genovali et al. (2014). For the  $V$ ,  $I$  magnitudes the main source was Ngeow (2012).<sup>12</sup> The  $J$ ,  $H$ ,  $K_s$  magnitudes listed by Groenewegen (2018) and van Leeuwen et al. (2007) were transformed in the 2MASS system (Skrutskie et al. 2006) using the transformations outlined by Groenewegen (2018, see section 2 for details). The  $J$ ,  $H$ ,  $K_s$  magnitudes obtained from the other sources were already in the 2MASS system. The uncertainties on the individual  $J$ ,  $H$ ,  $K_s$  magnitudes by Groenewegen (2018) were assigned according the criteria adopted in that paper. For the stars collected from other sources, we used the original individual uncertainties, or assigned the uncertainties in a similar way as in Groenewegen (2018). For the  $V$ ,  $I$  magnitudes we imposed a conservative uniform error of 0.02 mag for all the stars. The periods and the subclassification in F or IO or multimode DCEPs were taken by Ripepi et al. (2019), Ripepi et al. (2020a) and Skowron et al. (2019). The total sample includes 453 DCEPs, and precisely, 391 DCEP\_F, 43 DCEP\_IO, 17 DCEP\_F/IO, and 2 DCEP\_IO/2O. For the multimode objects we used the longest period.

An important ingredient of our analysis is represented by the iron abundance of the investigated DCEPs. The values listed by Groenewegen (2018) are based on a collection of literature values reported to a uniform scale by the Author. For comparison, we have also considered the work by Luck (2018) that has homogeneously estimated the iron abundance for 435 pulsators. We have cross matched his catalog with our sample, finding 414 matches. The remaining 21 non-matching objects have not been included as they are mainly stars suspected of being type II or anomalous Cepheids. Among the 414 matches, we have 357 DCEP\_F, 38 DCEP\_IO, 17 DCEP\_F/IO, and 2 DCEP\_IO/2O, respectively. A comparison between the iron abundance of our literature compilation and Luck (2018) sample is shown in Fig. 5. We note a fair agreement between the two sets of metallicity estimates (apart the very discrepant object GP Per), even if there is a measurable trend:  $[Fe/H]_{Lit} = 0.037 \pm 0.009 + (1.148 \pm 0.068)[Fe/H]_{L18}$ , where with  $[Fe/H]_{Lit}$  and  $[Fe/H]_{L18}$  we mean the samples built with our collection of different sources and the Luck (2018) one, respectively. Given the small but detectable differences in the two quoted metallicity scales, we decided to carry out our analysis using both samples.

The resulting merged catalogues were cross-matched with *Gaia* EDR3 to extract parallaxes, parallax errors, ruwe parameter of the sources which provides an indication of goodness of the astrometric solution,<sup>13</sup> as well as other parameters needed to calculate the individual zero-point corrections to the parallaxes (see Lindegren et al. 2021a, for full details). The approximate distribution in the Galaxy of the considered objects is shown in Fig. 6, where we plot



**Figure 5.** Comparison between the iron abundances from different literature sources and Luck (2018). The light blue and orange lines are the bisector and the fit to the data, respectively.



**Figure 6.** Position in the Galaxy of the Literature DCEPs (left panel) and those investigated in this work (right panel). The Sun is shown with a cross. The DCEPs symbols are colour-coded according to their metallicity.

separately the literature and our samples. The different symbols are colour coded according to the metallicity value.

## 6 GAIA EDR3 PARALLAXES

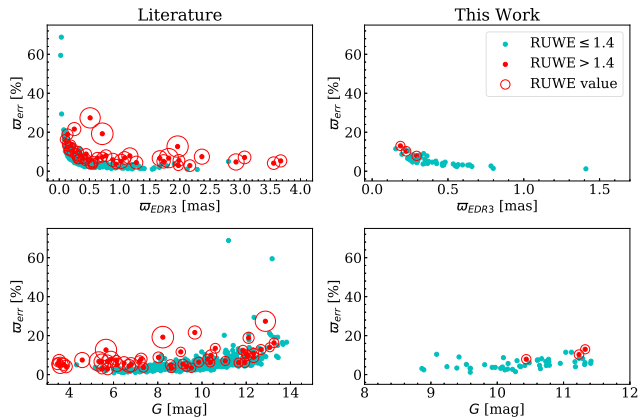
### 6.1 Parallax values for literature and present work samples.

The relative errors of the EDR3 parallaxes as a function of the parallax and  $G$  magnitude are displayed in the top and bottom panels of Fig. 7, respectively. In the figure the DCEPs are plotted with different colours depending on the ruwe value, using 1.4 as threshold (see, e.g. Lindegren et al. 2021b). In the figure, the size of the circles surrounding the objects with  $ruwe > 1.4$  is proportional to the ruwe value. The figure shows that: (i) the large majority of the investigated DCEPs has a relative precision on parallax better than 20 per cent; (ii) the trend of relative precision with parallax and magnitude is the expected one, as it increases for smaller parallaxes and larger magnitudes, respectively; (iii) the objects with large ruwe values do not follow the general trends, whereas many objects with  $1.4 < ruwe < 2$  do. This probably means that the parallaxes of these objects could be used; (iv) stars with  $G < 6$  mag have generally high ruwe values, as expected by the fact that the *Gaia* parallaxes for sources brighter than  $G \sim 6$  mag are still somewhat uncertain (Lindegren et al. 2021b).

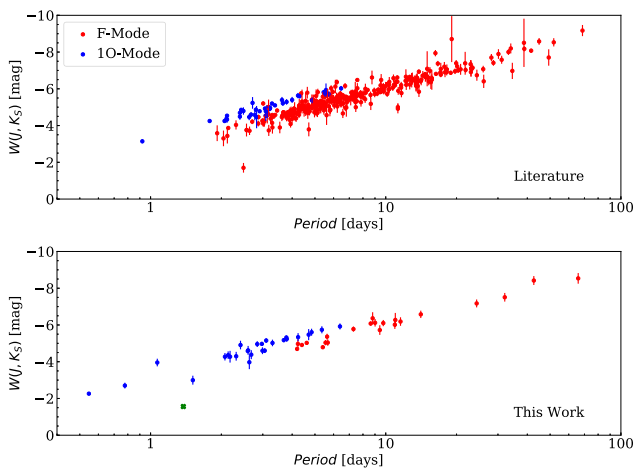
<sup>11</sup>The original catalog includes 452 objects, 14 sources were not considered because their classification as DCEPs is highly uncertain.

<sup>12</sup>When the  $V$  magnitudes was present in both Groenewegen (2018) and Ngeow (2012), we adopted these last work values to keep consistency with the  $J$  band.

<sup>13</sup>Section 14.1.2 of ‘Gaia Data Release 2 Documentation release 1.2’: <https://gea.esac.esa.int/archive/documentation/GDR2/>



**Figure 7.** Top panels: parallax versus relative parallax error ( per cent/100) for the literature (top-left) and our (top-right) samples, respectively. Red and green filled circles represent DCEPs with ruwe parameters below/above the nominal safety threshold, respectively (see labels). The size of the red circles is proportional to the same ruwe parameter. Bottom panels: as in the top panels, but in this case red and green filled circles represent DCEPs with  $G$  mag above/below the 6th magnitude (see labels). The size of the circles is always proportional to the ruwe.



**Figure 8.** PW relation in  $J, K_s$  bands for the literature and present work data sets (top and bottom panels, respectively). Red and blue symbols represent F and IO mode DCEPs, respectively. The green symbol in the bottom panel is for ASAS J162326–0941.0, one of our targets that turned out to be a BLHER star.

Considering the relatively large size of our samples, in the following we decided to be conservative and use the standard ruwe threshold = 1.4 and cleaned the samples from objects with ruwe value exceeding this limit. The resulting literature sample is thus composed of 373 F and 50 IO pulsators, respectively. Fig. 8 shows the  $PWJK_s$  for the literature sample (top panel) and for our sample (bottom panel). The figure shows that the large majority of the literature DCEPs have very precise parallaxes and both DCEP.F and DCEP.IO show rather tight relations. The only relevant exception is the star V447 Mon which appears significantly fainter than the rest of the DCEPs. Looking at its light curve on the ASAS-SN survey site (Shappee et al. 2014; Kochanek et al. 2017), there are little doubts that this star is a F-mode DCEP. Since both the Gaia astrometry and the 2MASS photometry of V447 Mon appear good, we do not have currently an explanation for its under-luminosity in the  $PWJK_s$  plane.

The tightness of the  $PWJK_s$  relation is even more evident for our sample, with a few exceptions. The star ASAS J162326–0941.0 is clearly underluminous with respect to the bulk of the other stars. Its parallax has high precision and the ruwe parameter is well below the 1.4 threshold. After inspecting its light curve shown in Fig. B1, we reckoned that ASAS J162326–0941.0 is a BLHER star misclassified as DCEP in the literature. We do not consider this star in the following analysis.

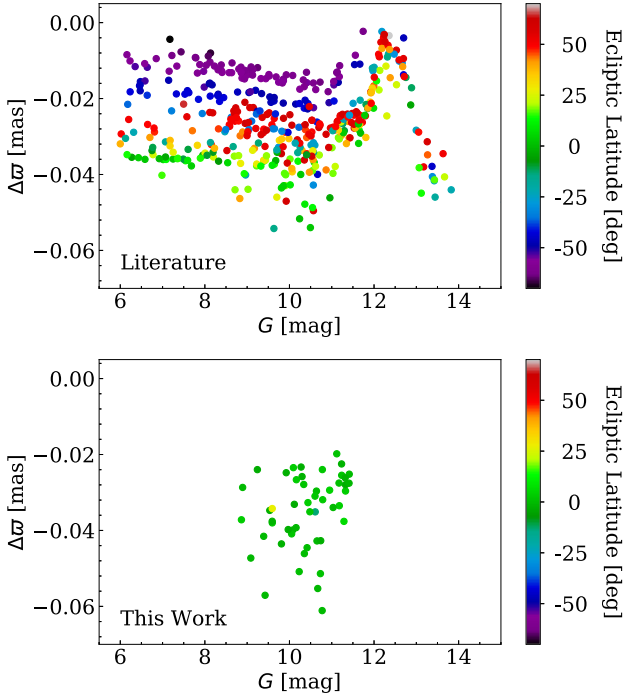
The other exceptions are V966 Mon (mixed mode 1O/2O pulsator with dominant period  $\sim 1.06$  d) and ASAS J063519+2117.8 (1O DCEP with period  $\sim 1.51$  d) which are about 0.5–0.6 mag too bright and too faint than expected, respectively. The ruwe parameter for both stars is well below 1.4 and also the photometry does not show obvious problems. The two stars appear well isolated on the sky, without any bright companion. We do not have a clear explanation for the deviation of these two stars apart from an undetected problem with astrometry.

It is important to note that Fig. 8 also shows the position of two additional DCEPs that were observed during the same observing run as the 48 stars discussed above, but whose spectra and photometry were published in separate papers, as they are both objects peculiar in different ways. Indeed, V363 Cas (mixed mode 1O/2O pulsator with dominant period  $\sim 0.546$  d) is a rare case of DCEP with Lithium in its spectra (Catanzaro et al. 2020) while HD 344787 (mixed mode F/IO pulsator with dominant period  $\sim 5.4$  d) is a Polaris analogue (Ripepi et al. 2021). Both stars are included in the following analysis.

## 6.2 EDR3 parallaxes zero-point correction

We adopted the correction for zero-point parallaxes developed by Lindegren et al. (2021a) and made the actual calculation through the python code accompanying that paper. As remarked by the Authors, the parallax zero-point offset depends in a complex way on (at least) the magnitude, colour, and ecliptic latitude of the target. This dependence in turn is different for the five- and six-parameter solutions in EDR3. As the corrections to the parallaxes add uncertainties on these values, we added in quadrature to each parallax error a fixed correction uncertainty of 0.013 mas. As explained in Lindegren et al. (2021a) the correction is only valid in the following intervals:  $G$  magnitude:  $6 < \text{phot\_g\_mean\_mag} < 21$  mag; Colour:  $1.1 < \text{nu\_eff\_used\_in\_astrometry} < 1.9$  for sources with 5 parameter solution;  $1.24 < \text{pseudocolour} < 1.72$  for 6 parameter solution sources (see Lindegren et al. 2021a, for details on these parameters). Beyond these values one can choose to both extrapolate the solution or neglect the correction. We decided to avoid using the extrapolated values and instead applied an average zero-point offset equal to the average of the correction for all the other DCEPs, i.e.  $0.025 \pm 0.011$  mas. As expected, the variables corrected with the fixed value all have  $G < 6$  mag. The corrections adopted here are displayed in Fig. 9 as a function of  $G$  magnitude and ecliptic latitude (see Lindegren et al. 2021a). It can be seen that the correction depends largely on the ecliptic latitude of the DCEP, but there is a minimum at  $G \sim 12.5$  mag, in agreement with Lindegren et al. (2021a).

After the publication of Lindegren et al. (2021a) individual parallax ZPO, several investigations appeared in the literature finding that these corrections are on average overestimated. Riess et al. (2021), based on the same sample of DCEPs adopted to estimate the metallicity dependence of the  $PWHVI$ , found an overcorrection of  $-14 \pm 6 \mu\text{as}$ . A similar result was obtained by Zinn (2021) from asteroseismology ( $-15 \pm 3 \mu\text{as}$  for stars with  $G > 10.8$  mag), while Gilligan et al. (2021) estimated an overcorrection of  $-10 \pm 7 \mu\text{as}$  from a study of RR Lyrae variables. More recently, Groenewegen



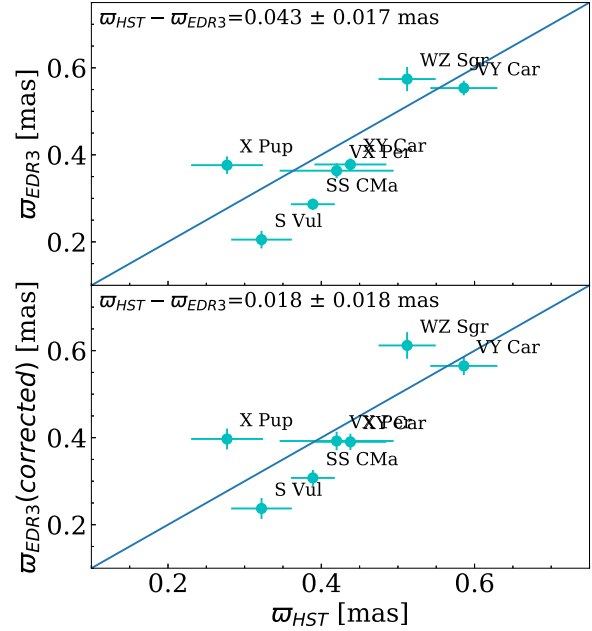
**Figure 9.** Corrections applied to the EDR3 parallaxes according to Lindegren et al. (2021a).

(2021) published individual parallax ZPO alternative to Lindegren et al. (2021a)’s but the data is still not available and the procedure rather complex. We will apply these new individual corrections in our future works. In this paper we shall consider the overcorrection by Riess et al. (2021). We expect that the results will be similar in the other cases.

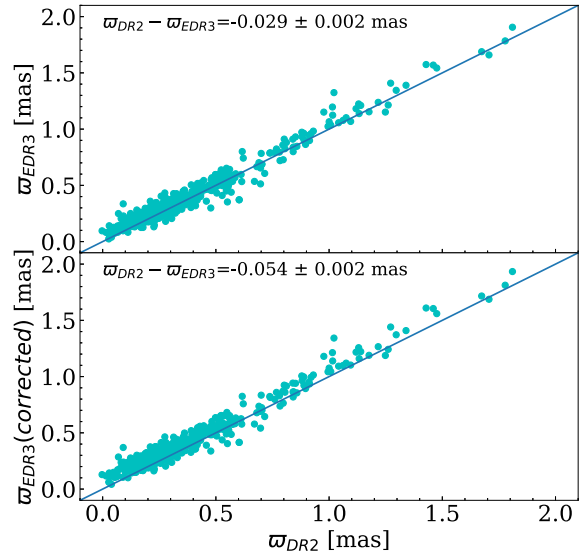
### 6.3 Comparison with DR2 and *HST* parallaxes

There are seven stars in our sample with independent parallaxes measured with the *HST* (Riess et al. 2018). The comparison between these values and those from EDR3 is reported in Fig. 10 where the top and bottom panels show the comparison with the original and ZPO corrected parallaxes, respectively. The weighted averages of the parallax differences in the two cases quoted above are reported in the figure. It can be seen that the uncorrected EDR3 parallaxes have underestimated values by  $\Delta\varpi = 0.043 \pm 0.017$  mas with respect to *HST* determinations. The zero-point offset is reduced to  $\Delta\varpi = 0.018 \pm 0.018$  mas when the corrected EDR3 values are considered. This result shows that the parallax correction is effective albeit not conclusive. However, we remark that the statistic is low and a few stars (namely X Pup and WZ Sgr) appear to have a different behaviour with respect to the others. Since all the stars have ruwe value smaller than 1.4 we have no reason to decide whether or not the EDR3 parallaxes of these stars are problematic.

The same comparison can be carried out for the DR2 parallaxes. The result is displayed in Fig. 11, where the top and bottom panels show again the comparison with uncorrected and corrected EDR3 parallaxes. As expected, the difference is larger in the case of corrected parallaxes which differ on average by  $\Delta\varpi = 0.054 \pm 0.002$  mas with respect to EDR3.



**Figure 10.** Comparison between the parallaxes in EDR3 and *HST*. Top and bottom panel show the difference without and with the individual ZPOs by Lindegren et al. (2021a), respectively.



**Figure 11.** Comparison between the parallaxes in EDR3 and DR2.

## 7 FITTING OF THE DATA

In this section we use the data described above to estimate the metallicity dependence of the DCEPs PL/PW relations. Based on the sample described in Section 5, we investigated the PL in the  $K_s$  ( $PLK_s$ ) band and the PW for the following three Wesenheit apparent magnitudes:  $w(J, K_s) = K_s - 0.69(J - K_s)$ ,  $w(V, K_s) = K_s - 0.13(V - K_s)$ , and  $w(H, V, I) = H - 0.386(V - I)$ , where the coefficients of these relations were taken from Cardelli et al. (1989) for the first two Wesenheit magnitudes and from Riess et al. (2021) for the last one. To determine the  $PLK_s$  we have obtained dereddened  $K_{s,0}$  magnitudes by adopting  $A_{K_s} = 0.34E(B - V)$  (as derived by

Cardelli et al. 1989, using  $R_V = 3.1$ ), where the  $E(B - V)$  values have been collected as described in Section 5.

To analyse the sample we followed the procedure adopted in Ripepi et al. (2020a). As in that work, to avoid any bias, we considered the whole sample, irrespective of the precision of individual parallaxes or of the presence of negative values for this parameter. To manage the negative parallaxes and to keep the symmetry in the parallax uncertainty, we used the astrometry-based luminosity (ABL, Feast & Catchpole 1997; Arenou & Luri 1999), which is defined as follows:  $ABL = 10^{0.2M} = \varpi 10^{0.2m-2}$ , where  $\varpi$  is the parallax, while  $M$  and  $m$  are the absolute and apparent generic magnitudes, respectively. Following Ripepi et al. (2020a), we can write  $M$  in the most general form, including the metallicity dependence on both the slope and intercept of the PL/PW relations:

$$M = \alpha + (\beta + \delta[\text{Fe}/\text{H}])(\log P - \log P_0) + \gamma[\text{Fe}/\text{H}], \quad (6)$$

where  $P$  and  $[\text{Fe}/\text{H}]$  are the observed period and iron abundance while  $P_0$  is a pivoting period. Therefore, in its most general form the ABL function we want to use is

$$ABL = \varpi 10^{0.2m-2} = 10^{0.2(\alpha+(\beta+\delta[\text{Fe}/\text{H}])(\log P - \log P_0)+\gamma[\text{Fe}/\text{H}])}, \quad (7)$$

where  $P$ ,  $[\text{Fe}/\text{H}]$ ,  $\varpi$  and the apparent magnitude/Wesenheit are the observables, while the unknowns are the four parameters of equation (6). Following Ripepi et al. (2020a), we considered PL/PW relations with different number of parameters, i.e. with no metallicity dependence (two parameters), with metallicity dependence only on the intercept (three parameters) and with metallicity dependence on both the slope and intercept (four parameters). As mentioned above, we considered the following four relations:  $PLK_s$ ,  $PWJK_s$ ,  $PWVK_s$ , and  $PWHVI$ . For each relation we calculated the solution with two, three and four parameters separately for the case of Lit and L18 metallicities, respectively. We also performed the ABL fitting to the data using two different approaches which are described in detail in the following sections.

## 7.1 Nonlinear fitting approach

A first approach to fit the equation (7) has been achieved by using the *nls* function from the R environment (R Core Team 2020). It allows to perform a nonlinear least-squares regression (NL-LSQs hereafter) based on the Gauss–Newton minimization algorithm (see e.g. Bates & Watts 1988). The different uncertainties on the ABL magnitudes have been taken into account by weighting the fit with standard  $w_i = 1/\sigma_i^2$  factors, where the  $\sigma_i$  term contains the contributions by observed magnitudes errors, reddening errors and parallax errors.

To estimate parameter errors a bootstrap approach has been applied. In particular, 100 simulated data sets have been obtained by resampling the original one, allowing repeated points. Then the fitting technique has been applied on all simulated samples and the coefficients of the equation (7) were recalculated 100 times. Finally a statistical analysis of the obtained coefficient distributions allowed to estimate their errors, defined as the robust standard deviations ( $std_{\text{robust}} = 1.4826 \cdot \text{std}$ ) of the quoted distributions.

Since the tested models have different number of parameters, we have also provided two estimators of the goodness of the models by calculating the Akaike information criterion (AIC; Akaike 2011) and the Bayesian information criterion (BIC; Schwarz 1978). They are defined respectively  $AIC = 2k - 2 \ln \mathcal{L}$  and  $BIC = k \cdot \ln N - 2 \ln \mathcal{L}$ , where  $k$  is the number of parameters in the fit and  $\ln \mathcal{L}$  is

given by:

$$\ln \mathcal{L} = -0.5 \sum_{i=1}^N \left[ \left( \frac{\text{res}_i}{s_i} \right)^2 + \ln (2\pi s_i^2) \right], \quad (8)$$

where  $N$  is the number of fitted points,  $\text{res}_i$  is the  $i$ th residual value, while the term  $s_i$  is the sum in quadrature of the  $i$ th uncertainty and the intrinsic scatter around the fitted relation. We have set the intrinsic scatter of the fitted relations equal to half the rms of the residuals, but checked that this choice does not affect the results significantly. The results of the fitting procedure are listed in Table 5.

## 7.2 Monte Carlo Markov Chain (MCMC) approach.

Complementary to the NL-LSQs approach, we performed an MCMC method, which is a widely used technique that preserves a high efficiency even when the parameter space is large. The idea is that, after a certain number of random steps (the length of the chain), moving towards a region of highest probability in the parameter space, the MCMC approximates the posterior distribution of a set of parameters. In this work we took advantage of the EMCEE package developed in Python (Foreman-Mackey et al. 2013),<sup>14</sup> by running 1000 chains each one with a size of 1000 steps (or walkers). The posterior distribution in the logarithmic form is defined, except for a normalization constant term, as

$$\log P(\theta|X) = \log P(X|\theta) + \log P(\theta), \quad (9)$$

where  $\theta$  represents the ensemble of parameters (i.e.  $\alpha$ ,  $\beta$ ,  $\gamma$ ,  $\delta$ ) and  $X$  is the observable, i.e. the left-hand side of equation (7).  $P(X|\theta)$  is the likelihood function, that we set, again in logarithmic form, as

$$\ln \mathcal{L} = \log P(X|\theta) = -0.5 \cdot \sum \left( \frac{X - \text{model}}{\Delta X} \right)^2, \quad (10)$$

where ‘model’ is the right-hand side of equation (7). Note that the minus sign indicates that the MCMC walks towards global minima. The function  $P(\theta)$  indicates our priors, that we set, in order to achieve a convergence in a finite number of walkers to  $-10 < \alpha < 0$ ;  $-6 < \beta < 0$ ;  $-2 < \gamma < 2$  and  $-5 < \delta < 5$ .

Our best estimation of each parameter is represented by the median of the posterior distribution defined in equation (9) and sampled via the MCMC. We set as uncertainties the 16th and 84th percentile. As discussed in the previous section the AIC and BIC are calculated also in this case to have an estimate of the goodness of each model. The results of the fitting procedure with the MCMC method are listed in Table 6. A comparison between the coefficient of the  $PLZ/PWZ$  relations calculated with the NL-LSQs and the MCMC methods are shown in Fig. 12. The agreement is excellent for all the coefficients, and this is particularly comforting, as the two approaches are completely different and independent of each other.

## 8 DISCUSSION

To test the goodness of our  $PLZ/PWZ$  relations, we apply them to the DCEPs in the LMC, derive its distance modulus and compare it to that geometrically derived (from eclipsing binaries) by Pietrzyński et al. (2019), i.e.  $\mu_{\text{LMC}} = 18.48 \pm 0.03$  mag (including systematic errors), which is currently considered one of the most accurate estimates in the literature.

To this aim, we collected literature photometry in the bands of interest of this work. More specifically, we considered about 4500

<sup>14</sup><https://emcee.readthedocs.io/en/stable/>

**Table 5.** Results of the fitting of the data with the NL-LSQs method. The ID identifies each different fit to the data;  $\alpha$ ,  $\beta$ ,  $\gamma$  and  $\delta$  are the coefficient of the *PLZ/PWZ* relations; n.MW is the number of MW DCEPs adopted in each fit; BIC and AIC are the goodness-of-fit parameters discussed in the text; Mode identifies the sample adopted in the fit: F means F-mode DCEPs only, F+IO means that both F- and IO-mode pulsators are used; Source shows the origin of the metallicities; Relation specifies the relationship used in the fit;  $\mu_{LMC}$  represents the distance modulus of the LMC obtained with the specific *PLZ* or *PWZ* relationship; n.LMC is the number of LMC DCEPs adopted to calculate the  $\mu_{LMC}$  value. The lines highlighted in grey represent our best solution for each *PLZ/PWZ* group.

ID	$\alpha$	$\beta$	$\gamma$	$\delta$	n.MW	BIC	AIC	Mode	Source	Relation	$\mu_{LMC}$	n.LMC
1	-5.819 ± 0.013	-3.160 ± 0.051			381	-1996	-2003	F	Lit	<i>PLK<sub>s</sub></i>	18.631 ± 0.011	2432
2	-5.760 ± 0.021	-3.069 ± 0.053	-0.530 ± 0.108		381	-2159	-2171	F	Lit	<i>PLK<sub>s</sub></i>	18.433 ± 0.010	2432
3	-5.761 ± 0.021	-3.070 ± 0.072	-0.527 ± 0.108	0.018 ± 0.402	381	-2154	-2170	F	Lit	<i>PLK<sub>s</sub></i>	18.432 ± 0.010	2432
4	-5.779 ± 0.016	-3.065 ± 0.045	-0.488 ± 0.138		349	-1894	-1906	F	L18	<i>PLK<sub>s</sub></i>	18.467 ± 0.011	2432
5	-5.782 ± 0.022	-3.077 ± 0.065	-0.464 ± 0.140	0.173 ± 0.582	349	-1895	-1910	F	L18	<i>PLK<sub>s</sub></i>	18.451 ± 0.010	2432
6	-5.826 ± 0.014	-3.145 ± 0.052			448	-2375	-2383	F+IO	Lit	<i>PLK<sub>s</sub></i>	18.634 ± 0.010	4396
7	-5.774 ± 0.015	-3.055 ± 0.050	-0.456 ± 0.099		448	-2504	-2516	F+IO	Lit	<i>PLK<sub>s</sub></i>	18.478 ± 0.009	4396
8	-5.770 ± 0.019	-3.046 ± 0.050	-0.493 ± 0.084	-0.129 ± 0.300	448	-2514	-2530	F+IO	Lit	<i>PLK<sub>s</sub></i>	18.488 ± 0.009	4396
9	-5.790 ± 0.019	-3.053 ± 0.040	-0.416 ± 0.156		411	-2236	-2248	F+IO	L18	<i>PLK<sub>s</sub></i>	18.508 ± 0.009	4396
10	-5.788 ± 0.022	-3.050 ± 0.053	-0.440 ± 0.147	-0.090 ± 0.388	411	-2229	-2245	F+IO	L18	<i>PLK<sub>s</sub></i>	18.514 ± 0.009	4396
11	-6.122 ± 0.014	-3.291 ± 0.064			381	-1903	-1911	F	Lit	<i>PWJK<sub>s</sub></i>	18.636 ± 0.011	2423
12	-6.061 ± 0.023	-3.197 ± 0.058	-0.554 ± 0.114		381	-2096	-2108	F	Lit	<i>PWJK<sub>s</sub></i>	18.429 ± 0.010	2423
13	-6.065 ± 0.020	-3.213 ± 0.112	-0.526 ± 0.109	0.166 ± 0.405	381	-2091	-2106	F	Lit	<i>PWJK<sub>s</sub></i>	18.415 ± 0.010	2423
14	-6.086 ± 0.029	-3.202 ± 0.047	-0.379 ± 0.287		349	-1804	-1815	F	L18	<i>PWJK<sub>s</sub></i>	18.510 ± 0.011	2423
15	-6.083 ± 0.021	-3.191 ± 0.097	-0.400 ± 0.199	-0.151 ± 0.992	349	-1786	-1801	F	L18	<i>PWJK<sub>s</sub></i>	18.524 ± 0.010	2423
16	-6.125 ± 0.016	-3.294 ± 0.048			448	-2325	-2333	F+IO	Lit	<i>PWJK<sub>s</sub></i>	18.633 ± 0.010	4360
17	-6.071 ± 0.018	-3.201 ± 0.041	-0.465 ± 0.071		448	-2448	-2460	F+IO	Lit	<i>PWJK<sub>s</sub></i>	18.473 ± 0.009	4360
18	-6.068 ± 0.020	-3.193 ± 0.056	-0.498 ± 0.093	-0.115 ± 0.271	448	-2444	-2460	F+IO	Lit	<i>PWJK<sub>s</sub></i>	18.482 ± 0.009	4360
19	-6.091 ± 0.022	-3.207 ± 0.039	-0.335 ± 0.151		411	-2162	-2174	F+IO	L18	<i>PWJK<sub>s</sub></i>	18.533 ± 0.009	4360
20	-6.089 ± 0.018	-3.201 ± 0.060	-0.371 ± 0.153	-0.133 ± 0.420	411	-2145	-2161	F+IO	L18	<i>PWJK<sub>s</sub></i>	18.544 ± 0.009	4360
21	-6.069 ± 0.012	-3.274 ± 0.044			369	-2040	-2048	F	Lit	<i>PWVK<sub>s</sub></i>	18.645 ± 0.012	2268
22	-6.008 ± 0.017	-3.182 ± 0.058	-0.540 ± 0.079		369	-2169	-2180	F	Lit	<i>PWVK<sub>s</sub></i>	18.443 ± 0.011	2268
23	-6.009 ± 0.022	-3.185 ± 0.058	-0.536 ± 0.121	0.025 ± 0.336	369	-2167	-2183	F	Lit	<i>PWVK<sub>s</sub></i>	18.441 ± 0.010	2268
24	-6.029 ± 0.017	-3.179 ± 0.047	-0.446 ± 0.186		338	-1925	-1936	F	L18	<i>PWVK<sub>s</sub></i>	18.496 ± 0.011	2268
25	-6.028 ± 0.018	-3.176 ± 0.055	-0.453 ± 0.154	-0.054 ± 0.660	338	-1918	-1933	F	L18	<i>PWVK<sub>s</sub></i>	18.501 ± 0.011	2268
26	-6.076 ± 0.015	-3.268 ± 0.043			434	-2406	-2414	F+IO	Lit	<i>PWVK<sub>s</sub></i>	18.647 ± 0.010	4108
27	-6.022 ± 0.022	-3.174 ± 0.049	-0.459 ± 0.107		434	-2502	-2514	F+IO	Lit	<i>PWVK<sub>s</sub></i>	18.490 ± 0.009	4108
28	-6.018 ± 0.015	-3.168 ± 0.047	-0.492 ± 0.103	-0.110 ± 0.257	434	-2499	-2516	F+IO	Lit	<i>PWVK<sub>s</sub></i>	18.497 ± 0.009	4108
29	-6.040 ± 0.018	-3.177 ± 0.047	-0.382 ± 0.125		398	-2233	-2245	F+IO	L18	<i>PWVK<sub>s</sub></i>	18.532 ± 0.009	4108
30	-6.038 ± 0.019	-3.174 ± 0.050	-0.410 ± 0.146	-0.098 ± 0.415	398	-2225	-2241	F+IO	L18	<i>PWVK<sub>s</sub></i>	18.539 ± 0.009	4108
31	-5.974 ± 0.015	-3.259 ± 0.044			276	-1738	-1745	F	Lit	<i>PWHVI</i>	18.558 ± 0.011	2000
32	-5.931 ± 0.016	-3.206 ± 0.064	-0.382 ± 0.085		276	-1795	-1806	F	Lit	<i>PWHVI</i>	18.410 ± 0.011	2000
33	-5.938 ± 0.017	-3.243 ± 0.072	-0.342 ± 0.093	0.329 ± 0.419	276	-1788	-1803	F	Lit	<i>PWHVI</i>	18.372 ± 0.010	2000
34	-5.946 ± 0.015	-3.199 ± 0.049	-0.383 ± 0.108		255	-1641	-1651	F	L18	<i>PWHVI</i>	18.428 ± 0.011	2000
35	-5.961 ± 0.019	-3.266 ± 0.064	-0.295 ± 0.123	0.924 ± 0.421	255	-1657	-1672	F	L18	<i>PWHVI</i>	18.322 ± 0.010	2000
36	-5.979 ± 0.011	-3.215 ± 0.050			317	-2007	-2015	F+IO	Lit	<i>PWHVI</i>	18.581 ± 0.010	2000
37	-5.937 ± 0.017	-3.172 ± 0.049	-0.366 ± 0.089		317	-2064	-2076	F+IO	Lit	<i>PWHVI</i>	18.435 ± 0.010	2000
38	-5.944 ± 0.019	-3.203 ± 0.065	-0.324 ± 0.085	0.261 ± 0.302	317	-2065	-2080	F+IO	Lit	<i>PWHVI</i>	18.409 ± 0.009	2000
39	-5.951 ± 0.014	-3.162 ± 0.059	-0.367 ± 0.084		292	-1897	-1908	F+IO	L18	<i>PWHVI</i>	18.453 ± 0.009	2000
40	-5.962 ± 0.019	-3.208 ± 0.063	-0.279 ± 0.116	0.611 ± 0.336	292	-1899	-1914	F+IO	L18	<i>PWHVI</i>	18.393 ± 0.010	2000

DCEPs in the LMC, having: *V*, *I* photometry by Udalski et al. (2018, OGLE IV survey); *J*, *K<sub>s</sub>* photometry (Ripepi et al. 2012, 2020b; Ripepi et al. in preparation); *H* (Inno et al. 2016). As for the periods, they were taken from the OGLE IV survey quoted above.

Then, we calculated the absolute  $M_{K_{0,s}}$  and Wesenheit magnitudes  $W(J, K_s)$ ,  $W(V, K_s)$ ,  $W(H, V, I)$  for each LMC DCEP adopting the coefficient of the *PLZ/PWZ* relations reported in Tables 5 and 6, using the periods by the OGLE IV survey quoted above and assuming  $[\text{Fe}/\text{H}]_{LMC} = -0.33$  dex (Romaniello et al. 2008).

Having apparent and absolute  $K_{0,s}$  and Wesenheit magnitudes, it is easy to obtain individual DM values for each LMC DCEP.F. The median of the distributions of these DMs is reported in Tables 5 and 6. The random error on the median is also provided and was calculated by summing in quadrature the standard deviation of the mean of the residuals obtained for the MW relations and the similar quantity for the LMC. The uncertainty is completely dominated by the MW relations, as the contribution of the LMC is negligible, given the large number of adopted pulsators in the LMC.

To better visualize the results we decided to calculate the following quantity:

$$n_\sigma = (\mu_{LMC}^{TW} - \mu_{LMC}^G) / \sigma, \quad (11)$$

where  $\mu_{LMC}^G$  and  $\mu_{LMC}^{TW}$  are the LMC DM as estimated by Pietrzyński et al. (2019) and in this work, respectively;  $\sigma$  is the uncertainty on the DM of the LMC obtained summing in quadrature the geometric uncertainty, equal to 0.03 mag with the random errors on  $\mu_{LMC}^{TW}$  reported in the Tables 5 and 6. Given that the value of  $\sigma$  is dominated by the uncertainty on  $\mu_{LMC}^G$ , we can approximately say that  $n_\sigma$  is the deviation of our results with respect to the geometric distance of the LMC, in units of the uncertainty on this last quantity. In Fig. 13, left- and right-hand panels show the value of  $n_\sigma$  for all the cases listed in Tables 5 and 6, respectively. The top and bottom panels report the cases adopting the Lit and L18 metallicity samples, respectively. An inspection of the figure allows us to draw some considerations:

**Table 6.** Fitting results from the MCMC approach. The meaning of the different columns is the same as in Table 5.

ID	$\alpha$	$\beta$	$\gamma$	$\delta$	n.MW	BIC	AIC	Mode	Source	Relation	$\mu_{LMC}$	n.LMC
1	$-5.824^{+0.006}_{-0.006}$	$-3.171^{+0.022}_{-0.022}$			381	-1613	-1621	F	Lit	$PLK_s$	$18.631 \pm 0.011$	2432
2	$-5.767^{+0.008}_{-0.008}$	$-3.085^{+0.023}_{-0.023}$	$-0.515^{+0.048}_{-0.049}$		381	-1717	-1729	F	Lit	$PLK_s$	$18.438 \pm 0.010$	2432
3	$-5.767^{+0.009}_{-0.009}$	$-3.086^{+0.029}_{-0.029}$	$-0.513^{+0.056}_{-0.056}$	$0.008^{+0.165}_{-0.164}$	381	-1711	-1727	F	Lit	$PLK_s$	$18.438 \pm 0.010$	2432
4	$-5.784^{+0.008}_{-0.007}$	$-3.075^{+0.024}_{-0.023}$	$-0.453^{+0.055}_{-0.055}$		349	-1569	-1581	F	L18	$PLK_s$	$18.480 \pm 0.011$	2432
5	$-5.785^{+0.008}_{-0.008}$	$-3.082^{+0.027}_{-0.027}$	$-0.440^{+0.061}_{-0.061}$	$0.094^{+0.204}_{-0.203}$	349	-1564	-1579	F	L18	$PLK_s$	$18.471 \pm 0.011$	2432
6	$-5.831^{+0.006}_{-0.006}$	$-3.152^{+0.019}_{-0.019}$			448	-1974	-1982	F+IO	Lit	$PLK_s$	$18.636 \pm 0.010$	4396
7	$-5.780^{+0.008}_{-0.008}$	$-3.068^{+0.021}_{-0.020}$	$-0.440^{+0.040}_{-0.041}$		448	-2083	-2095	F+IO	Lit	$PLK_s$	$18.482 \pm 0.009$	4396
8	$-5.776^{+0.008}_{-0.008}$	$-3.059^{+0.021}_{-0.021}$	$-0.479^{+0.051}_{-0.050}$	$-0.134^{+0.106}_{-0.104}$	448	-2078	-2095	F+IO	Lit	$PLK_s$	$18.493 \pm 0.009$	4396
9	$-5.794^{+0.007}_{-0.007}$	$-3.060^{+0.021}_{-0.021}$	$-0.393^{+0.046}_{-0.046}$		411	-1908	-1921	F+IO	L18	$PLK_s$	$18.516 \pm 0.009$	4396
10	$-5.792^{+0.007}_{-0.007}$	$-3.057^{+0.021}_{-0.021}$	$-0.417^{+0.056}_{-0.056}$	$-0.092^{+0.119}_{-0.118}$	411	-1903	-1919	F+IO	L18	$PLK_s$	$18.523 \pm 0.009$	4396
11	$-6.124^{+0.007}_{-0.007}$	$-3.303^{+0.023}_{-0.023}$			381	-1419	-1427	F	Lit	$PWJK_s$	$18.633 \pm 0.011$	2423
12	$-6.062^{+0.009}_{-0.008}$	$-3.210^{+0.024}_{-0.024}$	$-0.551^{+0.049}_{-0.050}$		381	-1532	-1544	F	Lit	$PWJK_s$	$18.426 \pm 0.010$	2423
13	$-6.067^{+0.009}_{-0.009}$	$-3.228^{+0.029}_{-0.029}$	$-0.523^{+0.056}_{-0.056}$	$0.171^{+0.170}_{-0.167}$	381	-1527	-1543	F	Lit	$PWJK_s$	$18.411 \pm 0.010$	2423
14	$-6.086^{+0.008}_{-0.008}$	$-3.208^{+0.024}_{-0.024}$	$-0.379^{+0.056}_{-0.056}$		349	-1392	-1404	F	L18	$PWJK_s$	$18.507 \pm 0.011$	2423
15	$-6.083^{+0.009}_{-0.008}$	$-3.197^{+0.028}_{-0.028}$	$-0.399^{+0.062}_{-0.062}$	$-0.164^{+0.208}_{-0.210}$	349	-1387	-1402	F	L18	$PWJK_s$	$18.524 \pm 0.010$	2423
16	$-6.126^{+0.006}_{-0.006}$	$-3.299^{+0.020}_{-0.020}$			448	-1841	-1849	F+IO	Lit	$PWJK_s$	$18.631 \pm 0.010$	4360
17	$-6.073^{+0.008}_{-0.008}$	$-3.212^{+0.021}_{-0.021}$	$-0.455^{+0.042}_{-0.042}$		448	-1948	-1961	F+IO	Lit	$PWJK_s$	$18.473 \pm 0.009$	4360
18	$-6.069^{+0.009}_{-0.009}$	$-3.201^{+0.023}_{-0.023}$	$-0.497^{+0.052}_{-0.052}$	$-0.151^{+0.112}_{-0.111}$	448	-1944	-1960	F+IO	Lit	$PWJK_s$	$18.485 \pm 0.009$	4360
19	$-6.090^{+0.007}_{-0.007}$	$-3.209^{+0.022}_{-0.022}$	$-0.330^{+0.048}_{-0.048}$		411	-1786	-1798	F+IO	L18	$PWJK_s$	$18.533 \pm 0.009$	4360
20	$-6.087^{+0.008}_{-0.008}$	$-3.200^{+0.022}_{-0.022}$	$-0.378^{+0.058}_{-0.058}$	$-0.192^{+0.129}_{-0.129}$	411	-1782	-1798	F+IO	L18	$PWJK_s$	$18.550 \pm 0.009$	4360
21	$-6.073^{+0.006}_{-0.006}$	$-3.284^{+0.022}_{-0.022}$			369	-1599	-1607	F	Lit	$PWVK_s$	$18.645 \pm 0.011$	2268
22	$-6.013^{+0.008}_{-0.008}$	$-3.196^{+0.023}_{-0.023}$	$-0.531^{+0.049}_{-0.049}$		369	-1709	-1720	F	Lit	$PWVK_s$	$18.445 \pm 0.011$	2268
23	$-6.014^{+0.009}_{-0.009}$	$-3.200^{+0.029}_{-0.029}$	$-0.526^{+0.056}_{-0.056}$	$0.034^{+0.166}_{-0.164}$	369	-1703	-1718	F	Lit	$PWVK_s$	$18.442 \pm 0.011$	2268
24	$-6.032^{+0.008}_{-0.008}$	$-3.187^{+0.024}_{-0.024}$	$-0.424^{+0.056}_{-0.056}$		338	-1570	-1582	F	L18	$PWVK_s$	$18.504 \pm 0.011$	2268
25	$-6.030^{+0.008}_{-0.008}$	$-3.179^{+0.027}_{-0.027}$	$-0.440^{+0.061}_{-0.060}$	$-0.125^{+0.205}_{-0.205}$	338	-1565	-1580	F	L18	$PWVK_s$	$18.516 \pm 0.011$	2268
26	$-6.079^{+0.006}_{-0.006}$	$-3.271^{+0.019}_{-0.019}$			434	-1964	-1973	F+IO	Lit	$PWVK_s$	$18.649 \pm 0.010$	4108
27	$-6.026^{+0.008}_{-0.008}$	$-3.185^{+0.020}_{-0.020}$	$-0.445^{+0.040}_{-0.041}$		434	-2076	-2088	F+IO	Lit	$PWVK_s$	$18.493 \pm 0.009$	4108
28	$-6.022^{+0.008}_{-0.008}$	$-3.176^{+0.021}_{-0.022}$	$-0.486^{+0.051}_{-0.051}$	$-0.141^{+0.101}_{-0.101}$	434	-2072	-2088	F+IO	Lit	$PWVK_s$	$18.504 \pm 0.009$	4108
29	$-6.042^{+0.007}_{-0.007}$	$-3.179^{+0.021}_{-0.021}$	$-0.365^{+0.045}_{-0.046}$		398	-1910	-1922	F+IO	L18	$PWVK_s$	$18.538 \pm 0.009$	4108
30	$-6.039^{+0.008}_{-0.008}$	$-3.174^{+0.021}_{-0.021}$	$-0.404^{+0.056}_{-0.056}$	$-0.143^{+0.117}_{-0.117}$	398	-1905	-1921	F+IO	L18	$PWVK_s$	$18.549 \pm 0.009$	4108
31	$-5.978^{+0.007}_{-0.007}$	$-3.278^{+0.025}_{-0.025}$			276	-1654	-1661	F	Lit	$PWHVI$	$18.554 \pm 0.011$	2000
32	$-5.934^{+0.009}_{-0.009}$	$-3.224^{+0.026}_{-0.026}$	$-0.391^{+0.058}_{-0.059}$		276	-1693	-1704	F	Lit	$PWHVI$	$18.403 \pm 0.011$	2000
33	$-5.942^{+0.011}_{-0.010}$	$-3.272^{+0.035}_{-0.035}$	$-0.344^{+0.062}_{-0.062}$	$0.419^{+0.202}_{-0.199}$	276	-1691	-1706	F	Lit	$PWHVI$	$18.352 \pm 0.011$	2000
34	$-5.949^{+0.008}_{-0.008}$	$-3.216^{+0.027}_{-0.027}$	$-0.373^{+0.067}_{-0.067}$		255	-1558	-1569	F	L18	$PWHVI$	$18.428 \pm 0.011$	2000
35	$-5.964^{+0.009}_{-0.009}$	$-3.285^{+0.032}_{-0.032}$	$-0.290^{+0.070}_{-0.070}$	$0.955^{+0.238}_{-0.235}$	255	-1569	-1583	F	L18	$PWHVI$	$18.315 \pm 0.010$	2000
36	$-5.983^{+0.007}_{-0.007}$	$-3.229^{+0.023}_{-0.023}$			317	-1899	-1907	F+IO	Lit	$PWHVI$	$18.579 \pm 0.010$	2000
37	$-5.942^{+0.009}_{-0.009}$	$-3.187^{+0.024}_{-0.024}$	$-0.367^{+0.052}_{-0.052}$		317	-1943	-1954	F+IO	Lit	$PWHVI$	$18.434 \pm 0.010$	2000
38	$-5.949^{+0.010}_{-0.010}$	$-3.223^{+0.033}_{-0.033}$	$-0.319^{+0.059}_{-0.059}$	$0.299^{+0.185}_{-0.184}$	317	-1940	-1955	F+IO	Lit	$PWHVI$	$18.403 \pm 0.009$	2000
39	$-5.955^{+0.008}_{-0.008}$	$-3.174^{+0.024}_{-0.024}$	$-0.356^{+0.060}_{-0.060}$		292	-1789	-1800	F+IO	L18	$PWHVI$	$18.456 \pm 0.009$	2000
40	$-5.966^{+0.009}_{-0.009}$	$-3.222^{+0.029}_{-0.029}$	$-0.273^{+0.066}_{-0.066}$	$0.617^{+0.211}_{-0.211}$	292	-1792	-1807	F+IO	L18	$PWHVI$	$18.393 \pm 0.010$	2000

(i) there is very little difference between the two approaches adopted for fitting the data. This is not surprising, as the coefficient of the  $PLZ/PWZ$  calculated with NL-LSQ and MCMC methods are very similar (see Fig. 12).

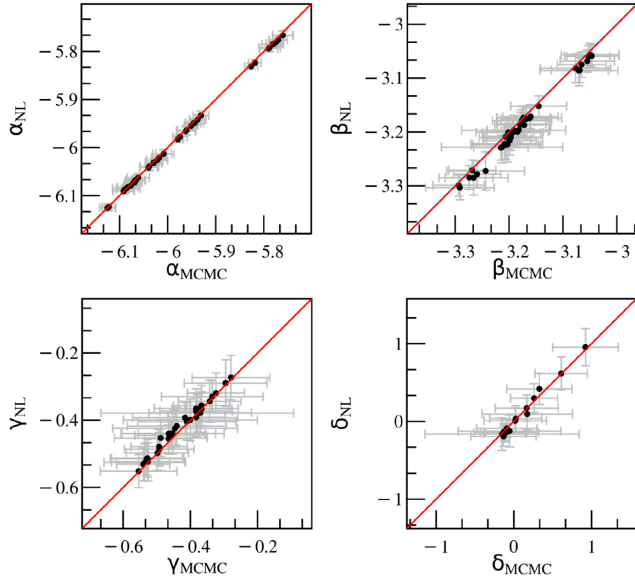
(ii) the  $PLK_s$ ,  $PWJK_s$ ,  $PWVK_s$  give always better distances to the LMC than the  $PWHVI$  relation in the NL-LSQ case, while the three-parameter case is as good as in the other bands for the L18 sample.

(iii) the two-parameter solution, i.e. without any metallicity term, is ruled out at more than the  $5\sigma$  level in all the cases except for the  $PWHVI$  relation ( $>3\sigma_s$ ).

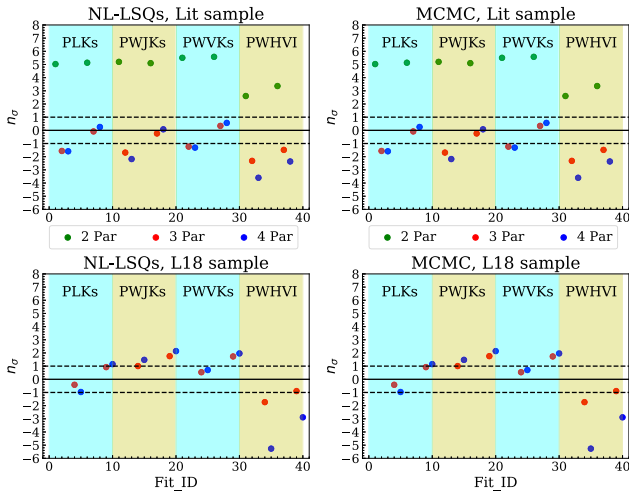
(iv) the fits with the complete sample, i.e. including both DCEP\_F and DCEP\_IO (fit ids 7, 8, 17, 18, 27, 28 in both Tables 5 and 6) provide better results in the case of the Lit sample, while they are equivalent in the case of the L18 sample.

Fig. 14 is very similar to Fig. 13, but this time on the y-axis we plotted the goodness of the fit AIC value. It is possible to see that the fit ids that provide the best LMC distances are also those with the lowest values of AIC, i.e. the best fits to the data. Also, the fits with four parameters, do not provide better AIC values with respect to the fits with three parameters.

As we have mentioned before, several authors have noticed that the individual parallax ZPO by Lindegren et al. (2021b) are overestimated in a measure that goes between 10 and 15  $\mu\text{s}$ . To verify the effect of this correction on our results, we decided to recalculate the distance of the LMC using the correction of 14  $\mu\text{s}$ , according to Riess et al. (2021) who calculated it in their derivation of the  $PWHVI$  relation for DCEPs. For simplicity, we considered only the case of the NL-LSQ fit with the Lit sample. The results



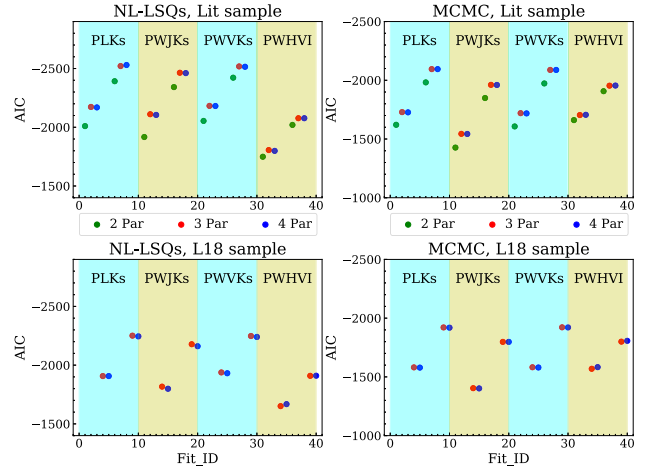
**Figure 12.** Comparison between the *PLZ/PWZ* relations coefficients calculated with the NL-LSQ and MCMC methods reported in Tables 5 and 6.



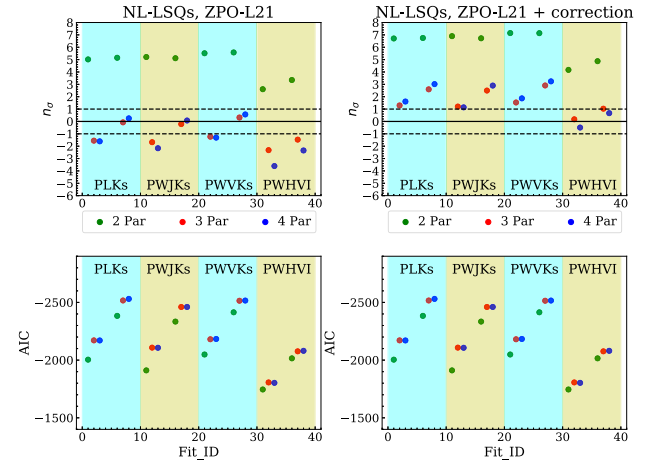
**Figure 13.** Number of  $\sigma$  from the geometric distance of the LMC for the 40 fittings to the data shown in each of Tables 5 and 6.

are shown in Table C1 which lists the same quantities of Table 5<sup>15</sup> and in Fig. 15, where the top and bottom panels represent the same quantities of Figs 13 and 14, respectively. An analysis of the figure reveals that the introduction of the correction worsens the results concerning the distance of the LMC for all the relations with the exception of the *PWHVI*, whose performance improves significantly, even if the corresponding AIC parameter is always worse than what obtained from the other *PLZ/PWZ* relations. To better understand this behaviour it is instructive to check how the coefficients of the *PLZ/PWZ* relations vary without/with the parallax ZPO corrections. This is shown in Fig. 16 for the same cases of Fig. 15 except for the two-parameter case which is scarcely informative. It can be seen that in all the considered cases, the introduction of the correction increases

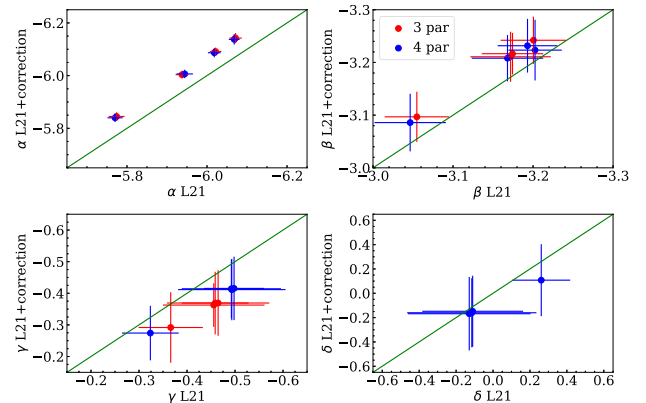
<sup>15</sup>We did not recalculate the parameters in the MCMC case because, as shown in Fig. 12, the NL and MCMC approaches provide almost identical results.



**Figure 14.** Significance of the fit (AIC) for the 40 fittings to the data shown in each of Tables 5 and 6.



**Figure 15.** Comparison between the nonlinear case fitting solution without (left-hand panels) and with the parallax ZPO by Riess et al. (2021) (right-hand panels). The top panels show the number of  $\sigma$ s from the geometric distance of the LMC, while the bottom panels display the AIC values relative to each case in the top panels.



**Figure 16.** Parameters of the *PLZ/PWZ* relations ( $\alpha$ ,  $\beta$  and  $\gamma$ ,  $\delta$ ) in the cases with three and four parameters (red and blue filled circles, respectively). The  $x$ - and  $y$ -axis show the results obtained in the NL-LSQs fitting solution without and with the parallax ZPO by Riess et al. (2021), respectively.

the absolute values of  $\alpha$  and  $\beta$  by  $\sim 1$  per cent, while  $\gamma$  is decreased (in absolute value) by  $\sim 18$ – $20$  per cent and  $\sim 13$ – $16$  per cent in the three- and four-parameter cases, respectively. The absolute value of  $\delta$  is decreased by  $\sim 30$  per cent. Since the  $\alpha$  and  $\beta$  coefficients remain practically unaffected, the obvious conclusion of this experiment is that the parallax ZPO correction and the metallicity terms are degenerate. The introduction of the ZPO correction is compensated by lower (in absolute sense) metallicity terms, both in the cases of three and four parameters. In the case of four parameters, the  $\gamma$  term decrease is smaller, since also the  $\delta$  coefficient decreases.

In all cases, without and with the introduction of the parallax ZPO correction, our results provide metallicity terms on the intercept that are larger than those found in the recent literature by Riess et al. (2021) and Breuval et al. (2021), even if the  $\gamma$  coefficient of the *PWHVI* relation with the correction is consistent within the errors with that estimated by Riess et al. (2021) that used the same *PWHVI* relation and the same parallax ZPO correction.

## 9 SUMMARY AND CONCLUSIONS

In this paper we have presented new time-resolved HARPS-N@TNG high-resolution ( $R = 115\,000$ ) spectroscopy (see Section 2) for a sample of 47 bright DCEPs plus 1 BLHER star<sup>16</sup> with good *Gaia* DR3 parallaxes. This is the first paper of a project devoted to significantly enlarge the number of MW DCEPs with reliable abundance measurements, which we will then use to carefully investigate the metallicity dependence of the PL/PW relations of these primary standard candles of the cosmic distance ladder.

In Section 3 we have carried out a detailed analysis in order to derive, from each spectrum, main atmospheric parameters such as the effective temperature, surface gravity, microturbulence, and radial velocity. We have performed a complete characterization of the chemical abundances of 29 species for which we have detected and measured spectral lines. We found on average, abundance consistent with the solar values for the  $\alpha$ - and *iron*-peaks elements, as well as for aluminum, scandium, vanadium, and copper. Among the  $\alpha$ -elements, sulfur is slightly underabundant. Solar values have been inferred even for heavy elements such as strontium, yttrium and zirconium and, in general, even for rare earths. Barium is spread out over a wide range ( $\approx 1$  dex) of overabundances. As expected for DCEPs, carbon and oxygen appear depleted with respect to the Sun and, on the contrary, sodium shows an average overabundance. In general, our abundances are in agreement with the general trend of those by Luck (2018), both for what concern the correlation with iron and the Galactic gradient. Additionally, we confirmed the anticorrelation between oxygen and iron.

In Section 4 we derived new, accurate estimates of the period and epoch of maximum light for all our program stars by combining *V*-band photometry from different sources available in the literature. The typical time baseline of the combined time-series data is larger than 15 yr for most of our targets. Periods and epochs inferred from the *V*-band data were used to fold the *B*, *I* data, whenever available in the literature, and derive the corresponding average magnitudes. Intensity-averaged magnitudes in the *J*, *H*, *K<sub>s</sub>* bands were computed from the 2MASS single epoch photometry by means of a template fitting technique. The average precision of the intensity-averaged magnitudes we obtained for our program stars is, for all bands, on the order of a few hundredths of magnitude. To estimate individual

reddening values for all our program stars, we calculated and used new, accurate PC relations in the (*B* – *V*) and (*V* – *I*) colours, which showed to be in very good agreement with previous results in the literature.

In Section 5 our sample of 47 DCEPs was complemented with DCEPs for which iron abundances from high-resolution spectroscopy are available in the literature. We considered two separate samples, one comprising a compilation of literature metallicities and the other based exclusively on the work by L18. The comparison between the two metallicity scales reveals that the L18 results give on average slightly lower metallicities, apart from the metal-poor regime ( $[\text{Fe}/\text{H}] \sim -0.3$ – $-0.4$  dex) where the trend is inverted.

In Section 6 we discussed the astrometric properties of our and literature samples, eliminating from the following analysis the stars with ruwe parameter larger than 1.4 and then correcting the *Gaia* parallax ZPO individually using the recipe by Lindgren et al. (2021b). The comparison between the ZPO-corrected *Gaia* and the *HST* parallaxes for the seven stars having such *HST* measures, reveals agreement at  $1\sigma$ , albeit with a rather large scatter. After removing the stars with poor astrometry, the total DCEPs sample adopted in this work, including the 47 stars presented here is of 448 and 411 stars for the Lit and L18 samples, respectively.

In Section 7 we faced the problem of deriving the *PLZ/PWZ* relations using the data presented in the previous sections. To this aim, we adopted the ABL formulation that allows to treat linearly the parallax. We calculated the *PLZ/PWZ* considering the following three cases: (i) no metallicity term (two parameters); (ii) with metallicity term on the intercept (three parameters); (iii) with the metallicity term on both the slope and the intercept (four parameters). We considered four different types of relations, involving NIR magnitudes: *PLK<sub>s</sub>*, *PWJK<sub>s</sub>*, *PWVK<sub>s</sub>* and *PWHVI*. The fitting procedure was carried out in two completely different ways: a standard NL-LSQs technique with a bootstrap routine to robustly calculate the errors and a Bayesian approach adopting an MCMC method to best estimate coefficients and errors. The fitting procedure has been carried out in a variety of cases: with 2,3,4 parameters, using the Lit or L18 metallicity samples and adopting separate DCEP\_F only or DCEP\_F+DCEP\_IO samples. In the end we have 40 cases that have been investigated with both the NL-LSQ and MCMC methods.

In Section 8 we tested the *PLZ/PWZ* relations calculated in this work against the geometric distance to the LMC estimated by Pietrzyński et al. (2019), and using the AIC/BIC parameters as goodness-of-fit estimators. To estimate the LMC distance with our *PLZ/PWZ* relations, we used literature data for  $\sim 4500$  DCEPs in the LMC and averaged out the single distance moduli obtained for each DCEP. We find that: (1) the two fitting methods adopted to derive the *PLZ/PWZ* relations provide results in excellent agreement each other; (2) the adoption of the Lit or the L18 metallicities gives consistent results; (3) a significant metallicity term is needed to derive LMC distances  $< 1\sigma$  from the geometric one. The best results are provided by the *PLK<sub>s</sub>*, *PWJK<sub>s</sub>* and *PWVK<sub>s</sub>* relations with three parameters, the Lit metallicities and the full sample (F+IO-mode pulsators), i.e. the solutions 7, 17, 27 of Table 5. In these cases we find a metallicity dependence of  $-0.456 \pm 0.099$ ,  $-0.465 \pm 0.071$ , and  $-0.459 \pm 0.107$  mag dex<sup>-1</sup>, respectively. These fits are also those with the best AIC/BIC values. The metal dependence for the *PWHVI* case is significantly lower:  $-0.366 \pm 0.089$  mag dex<sup>-1</sup>, with worse AIC/BIC values than the previous cases; (4) the fits with four parameters do not improve the results (AIC/BIC values equal to the three parameters case); (5) the inclusion of a correction by 14  $\mu\text{s}$  by Riess et al. (2021) to the *Gaia* parallaxes ZPO, improves significantly the results with the *PWHVI* relations, but provide  $> 1\sigma$

<sup>16</sup>ASAS J162326–0941.0, previously classified as DCEP is actually a T2CEP of BLHER type.

LMC distances for the other relationships; (6) the inclusion of the correction to the Gaia parallaxes ZPO causes a reduction by  $\sim 22$  per cent (in absolute terms) of the metallicity terms both in the three and four parameter cases. Since the resulting LMC distances are still compatible with the geometric one, this occurrence suggests that the inclusion of the parallax ZPO correction and the metallicity terms is largely degenerate. (7) In all the cases considered here, except the *PWHVI* relations with the ZPO correction, the metallicity terms we found are larger than those provided in the literature.

The fact that the ZPO correction and the metallicity terms appear degenerate means that it is important to have independent estimates of the parallax ZPO. In this sense we will test the recipe to calculate individual parallax ZPOs recently published in pre-print form by Groenewegen (2021). In any case, it seems that we are still not close to have a precision of 3–4  $\mu\text{s}$  on the parallax ZPO that is mandatory to derive the  $H_0$  constant with 1 per cent precision.

## ACKNOWLEDGEMENTS

We thank our anonymous referee for their comments that helped us to improve the manuscript. This work has been partially supported by INAF through the ‘Main Stream SSH program’ (1.05.01.86.28). This work has made use of data from the European Space Agency (ESA) mission *Gaia* (<https://www.cosmos.esa.int/gaia>), processed by the *Gaia* Data Processing and Analysis Consortium (DPAC, <https://www.cosmos.esa.int/web/gaia/dpac/consortium>). Funding for the DPAC has been provided by national institutions, in particular the institutions participating in the *Gaia* Multilateral Agreement. In particular, the Italian participation in DPAC has been supported by Istituto Nazionale di Astrofisica (INAF) and the Agenzia Spaziale Italiana (ASI) through grants I/037/08/0, I/058/10/0, 2014-025-R.0, and 2014-025-R.1.2015 to INAF (PI: M.G. Lattanzi). VR, MM, and GC acknowledge partial support from the project ‘MITiC: MIning The Cosmos Big Data and Innovative Italian Technology for Frontier Astrophysics and Cosmology’ (PI: B. Garilli). This publication makes use of VOSA, developed under the Spanish Virtual Observatory project supported by the Spanish MINECO through grant AyA2017-84089. VOSA has been partially updated by using funding from the European Union’s Horizon 2020 - Research and Innovation Framework Programme, under Grant Agreement No. 776403 (EXOPLANETS-A). This research has made use of the SIMBAD database, operated at CDS, Strasbourg, France. Based on data obtained from the ESO Science Archive Facility under request number 534967.

Based on observations made with the Italian Telescopio Nazionale Galileo (TNG) operated by the Fundación Galileo Galilei (FGG) of the Istituto Nazionale di Astrofisica (INAF) at the Observatorio del Roque de los Muchachos (La Palma, Canary Islands, Spain).

## DATA AVAILABILITY

The only proprietary data used in this paper are represented by the HARPS-N spectra. The reduced spectra are available upon request to the corresponding author (V. Ripepi).

## REFERENCES

Akaike H., 2011, in Lovric M., ed., *Akaike’s Information Criterion*. International Encyclopedia of Statistical Science. Springer, Berlin  
 Alfonso-Garzón J., Domingo A., Mas-Hesse J. M., Giménez A., 2012, *A&A*, 548, A79

Andrievsky S. M., Luck R. E., Kovtyukh V. V., Lepine J. R. D., 2012, *PASP*, 124, 934  
 Arenou F., Luri X., 1999, in Egret D., Heck A., eds, *ASP Conf. Ser. Vol. 167, Harmonizing Cosmic Distance Scales in a Post-Hipparcos Era*. Astron. Soc. Pac., San Francisco, p. 13  
 Bates D. M., Watts D. G., 1988, *Nonlinear Regression Analysis and Its Applications*. Wiley, New York  
 Berdnikov L. N., 2008, *yCat*, II/285  
 Berdnikov L. N., Samus N. N., Antipin S. V., Ezhkova O. V., Pastukhova E. N., Turner D. G., 2004, *PASP*, 116, 536  
 Berdnikov L. N., Kniazev A. Y., Kravtsov V. V., Pastukhova E. N., Turner D. G., 2009, *Astron. Lett.*, 35, 311  
 Berdnikov L. N., Kniazev A. Y., Sefako R., Kravtsov V. V., Pastukhova E. N., Zhuiko S. V., 2011, *Astron. Rep.*, 55, 816  
 Berdnikov L. N., Kniazev A. Y., Sefako R., Dambis A. K., Kravtsov V. V., Zhuiko S. V., 2015, *Astron. Lett.*, 41, 23  
 Bono G., Caputo F., Marconi M., Musella I., 2010, *ApJ*, 715, 277  
 Breuval L. et al., 2021, *ApJ*, 913, 38  
 Caputo F., Marconi M., Ripepi V., 1999, *ApJ*, 525, 784  
 Caputo F., Marconi M., Musella I., 2000, *A&A*, 354, 610  
 Cardelli J. A., Clayton G. C., Mathis J. S., 1989, *ApJ*, 345, 245  
 Castelli F., Hubrig S., 2004, *A&A*, 425, 263  
 Catanzaro G. et al., 2020, *A&A*, 639, L4  
 Clementini G. et al., 2016, *A&A*, 595, A133  
 Clementini G. et al., 2019, *A&A*, 622, A60  
 Dean J. F., Warren P. R., Cousins A. W. J., 1978, *MNRAS*, 183, 569  
 Di Francesco M., Marconi M., Musella I., Cignoni M., Ripepi V., 2013, *MNRAS*, 428, 212  
 Fausnaugh M. M., Kochanek C. S., Gerke J. R., Macri L. M., Riess A. G., Stanek K. Z., 2015, *MNRAS*, 450, 3597  
 Feast M. W., Catchpole R. M., 1997, *MNRAS*, 286, L1  
 Fernie J. D., 1967, *AJ*, 72, 1327  
 Fernie J. D., Evans N. R., Beattie B., Seager S., 1995, *Inf. Bull. Var. Stars*, 4148, 1  
 Fiorentino G., Marconi M., Musella I., Caputo F., 2007, *A&A*, 476, 863  
 Fiorentino G., Musella I., Marconi M., 2013, *MNRAS*, 434, 2866  
 Foreman-Mackey D., Hogg D. W., Lang D., Goodman J., 2013, *PASP*, 125, 306  
 Freedman W. L., 2017, *Nat. Astron.*, 1, 0121  
 Freedman W. L., Madore B. F., 2011, *ApJ*, 734, 46  
 Freedman W. L. et al., 2012, *ApJ*, 758, 24  
 Gaia Collaboration, 2016, *A&A*, 595, A1  
 Gaia Collaboration, 2017, *A&A*, 605, A79  
 Gaia Collaboration, 2018, *A&A*, 616, A1  
 Gaia Collaboration, 2021, *A&A*, 649, A1  
 Genovali K. et al., 2014, *A&A*, 566, A37  
 Genovali K. et al., 2015, *A&A*, 580, A17  
 Gieren W. et al., 2018, *A&A*, 620, A99  
 Gilligan C. K. et al., 2021, *MNRAS*, 503, 4719  
 Graczyk D. et al., 2020, *ApJ*, 904, 13  
 Grevesse N., Asplund M., Sauval A. J., Scott P., 2010, *Ap&SS*, 328, 179  
 Groenewegen M. A. T., 2013, *A&A*, 550, A70  
 Groenewegen M. A. T., 2018, *A&A*, 619, A8  
 Groenewegen M., 2021, preprint ([arXiv:2106.08128](https://arxiv.org/abs/2106.08128))  
 Inno L. et al., 2016, *ApJ*, 832, 176  
 Kashuba S. V. et al., 2016, *MNRAS*, 461, 839  
 Kochanek C. S. et al., 2017, *PASP*, 129, 104502  
 Kodric M. et al., 2013, *AJ*, 145, 106  
 Kovtyukh V. V., Gorlova N. I., 2000, *A&A*, 358, 587  
 Kovtyukh V. V., Soubiran C., Luck R. E., Turner D. G., Belik S. I., Andrievsky S. M., Chekhonadskikh F. A., 2008, *MNRAS*, 389, 1336  
 Kramida A., Ralchenko Yu., Reader J., NIST ASD Team, 2019, *NIST Atomic Spectra Database (version 5.7.1)*. National Institute of Standards and Technology, Gaithersburg, MD  
 Kurucz R. L., 1993a, in Dworetzky M. M., Castelli F., Faraggiana R., eds, *A new opacity-sampling model atmosphere program for arbitrary abundances*. In: *Peculiar versus normal phenomena in A-type and related stars*, IAU Colloquium 138. A.S.P. Conferences Series, Vol. 44, p. 87

- Kurucz R. L., 1993b, Kurucz CD-ROM 13: ATLAS9. SAO, Cambridge, USA
- Kurucz R. L., Avrett E. H., 1981, SAO Special Report, 391
- Kurucz R. L., Bell B., 1995, Kurucz CD-ROM No. 23
- Leavitt H. S., Pickering E. C., 1912, HarCi, 173
- Lindegren L. et al., 2021a, *A&A*, 649, A4
- Lindegren L. et al., 2021b, *A&A*, 649, A2
- Lomb N. R., 1976, *Ap&SS*, 39, 447
- Luck R. E., 2018, *AJ*, 156, 171
- Macri L. M., Stanek K. Z., Bersier D., Greenhill L. J., Reid M. J., 2006, *ApJ*, 652, 1133
- Madore B. F., 1982, *ApJ*, 253, 575
- Marquardt D. W., 1963, *SIAM J. Appl. Math.*, 11, 431
- Martin R. P., Andrievsky S. M., Kovtyukh V. V., Korotin S. A., Yegorova I. A., Saviane I., 2015, *MNRAS*, 449, 4071
- Monson A. J., Pierce M. J., 2011, *ApJS*, 193, 12
- Ngeow C.-C., 2012, *ApJ*, 747, 50
- Ngeow C.-C., Kanbur S. M., Bellinger E. P., Marconi M., Musella I., Cignoni M., Lin Y.-H., 2012, *Ap&SS*, 341, 105
- Pejcha O., Kochanek C. S., 2012, *ApJ*, 748, 107
- Pel J. W., 1978, *A&A*, 62, 75
- Pietrzyński G. et al., 2019, *Natur*, 567, 200
- Planck Collaboration VI, 2020, *A&A*, 641 A6
- Pojmanski G., 2002, *Acta Astron.*, 52, 397
- R Core Team, 2020, A Language and Environment for Statistical Computing. R Foundation for Statistical Computing, Vienna, Austria, available at <https://www.R-project.org/>
- Riess A. G., 2020, *Nat. Rev. Phys.*, 2, 10
- Riess A. G. et al., 2016, *ApJ*, 826, 56
- Riess A. G. et al., 2018, *ApJ*, 855, 136
- Riess A. G., Casertano S., Yuan W., Bowers J. B., Macri L., Zinn J. C., Scolnic D., 2021, *ApJ*, 908, L6
- Ripepi V. et al., 2012, *MNRAS*, 424, 1807
- Ripepi V., Molinaro R., Musella I., Marconi M., Leccia S., Eyer L., 2019, *A&A*, 625, A14
- Ripepi V. et al., 2020a, *A&A*, 642, A230
- Ripepi V., Molinaro R., Marconi M., Cioni M.-R. L., 2020b, *ASPC*, 529, 155
- Ripepi V. et al., 2021, *A&A*, 647, A111
- Romaniello M. et al., 2008, *A&A*, 488, 731
- Sandage A., Tammann G. A., Saha A., Reindl B., Macchetto F. D., Panagia N., 2006, *ApJ*, 653, 843
- Scargle J. D., 1982, *ApJ*, 263, 835
- Schmidt E. G., Johnston D., Langan S., Lee K. M., 2005, *AJ*, 130, 832
- Schmidt E. G., Rogalla D., Thacker-Lynn L., 2011, *AJ*, 141, 53
- Schwarz G., 1978, *Ann. Stat.*, 6, 461
- Shappee B. J., Stanek K. Z., 2011, *ApJ*, 733, 124
- Shappee B. J. et al., 2014, *ApJ*, 788, 48
- Skowron D. M. et al., 2019, *Science*, 365, 478
- Skrutskie M. F. et al., 2006, *AJ*, 131, 1163
- Soszyński I., Gieren W., Pietrzyński G., 2005, *PASP*, 117, 823
- Tammann G. A., Sandage A., Reindl B., 2003, *A&A*, 404, 423
- Turner D. G., 2016, *RMxAA*, 52, 223
- Udalski A. et al., 2018, *Acta Astron.*, 68, 315
- van Leeuwen F., Feast M. W., Whitelock P. A., Laney C. D., 2007, *MNRAS*, 379, 723
- Verde L., Treu T., Riess A. G., 2019, *Nat. Astron.*, 3, 891
- Wenger M. et al., 2000, *A&AS*, 143, 9
- Zinn J. C., 2021, *AJ*, 161, 214

## SUPPORTING INFORMATION

Supplementary data are available at *MNRAS* online.

**Appendix A:** Results of the fit of the trend of the abundance of the different elements as a function of the iron abundance and of the Galactocentric distance.

**Appendix B:** Light curves in V-band for the target stars.

**Appendix C:** Results for the *PLZIPWZ* relations in the case of adoption of a common Zero Point Offset (ZPO) for the Gaia parallaxes.

Please note: Oxford University Press is not responsible for the content or functionality of any supporting materials supplied by the authors. Any queries (other than missing material) should be directed to the corresponding author for the article.

This paper has been typeset from a  $\text{\TeX}/\text{\LaTeX}$  file prepared by the author.