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# A faint outburst of the accreting millisecond X-ray pulsar SAX J1748.9-2021 in NGC 6440

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## ABSTRACT

SAX J1748.9-2021 is an accreting X-ray millisecond pulsar observed in outburst five times since its discovery in 1998. In early October 2017, the source started its sixth outburst, which lasted only  $\sim 13$  days, significantly shorter than the typical 30 days duration of the previous outbursts. It reached a 0.3–70 keV unabsorbed peak luminosity of  $\sim 3 \times 10^{36}$  erg s<sup>-1</sup>. This is the weakest outburst ever reported for this source to date. We analysed almost simultaneous *XMM-Newton*, *NuSTAR*, and *INTEGRAL* observations taken during the decaying phase of its 2017 outburst. We found that the spectral properties of SAX J1748.9-2021 are consistent with an absorbed Comptonization plus a blackbody component. The former, characterized by an electron temperature of  $\sim 20$  keV, a photon index of  $\sim 1.6$ – $1.7$  keV, and seed photon temperature of 0.44 keV, can be associated to a hot corona or the accretion column, while the latter is more likely originating from the neutron star surface ( $kT_{bb} \sim 0.6$  keV,  $R_{bb} \sim 2.5$  km). These findings suggest that SAX J1748.9-2021 was observed in a *hard* spectral state, as it is typically the case for accreting millisecond pulsars in outburst.

**Key words:** accretion, accretion discs – X-rays: binaries – X-Rays: galaxies – X-rays: individuals: SAX J1748.9-2021.

## 1 INTRODUCTION

Accreting neutron stars (NS) with spin pulsations of the order of the millisecond ( $\sim 1.6$ – $10$  ms) and hosted in low-mass X-ray binary systems (LMXBs) belong to the class of the accreting X-ray millisecond pulsars (AMXPs), that counts 21 sources till date (e.g. Patruno, Haskell & Andersson 2017; Sanna et al. 2018; Strohmayer et al. 2018). It has been proven that AMXPs are formed through the so-called *recycling* scenario (see, e.g. Archibald et al. 2009; Papitto et al. 2013 b, for the observational evidences of such process). According to this scenario, the period of old pulsars in binary systems can be re-accelerated to a few milliseconds from the torque exerted by matter in the accretion disc formed during the outbursts. The properties of AMXPs suggest that NS magnetic fields of the order of  $10^8$  –  $10^9$  G (e.g. Hartman et al. 2008) are capable of driving a portion of the accreting matter from the accretion disc to the magnetic polar caps of the compact object, heating its surface. The photons produced by this process are responsible for the pulsating

X-ray emission (assuming that the NS magnetic and rotational axes are misaligned, as it is often the case). The timing analysis of the pulsations can provide information about the strength of the NS magnetic fields and the inner radius of the accretion discs (e.g. Burderi & King 1998; Burderi et al. 2006, 2007; Ghosh & Becker 2008; Patruno et al. 2009; Wilkinson et al. 2011 or di Salvo et al. 2008 for a review). The latter can also be inferred from either the detection of disc emission or, indirectly, from broad emission features (the most diffuse feature is the K-shell transition iron line at 6.4–7.0 keV) in the source X-ray spectrum. These are likely produced by reflection off of hard photons from the surface of the accretion disc and are affected by special and general relativistic effects caused by the fast rotation of the accretion disc close to the strong gravitational field of the compact object (e.g. Fabian et al. 1989; Papitto et al. 2009; Cackett et al. 2010; Di Salvo et al. 2018, submitted, and references therein).

The spectral properties of AMXPs have been largely investigated (see, e.g. Patruno & Watts 2012; Burderi & Di Salvo 2013; Campana & Di Salvo 2018, and references therein, for recent reviews) and showed that AMXPs are usually in *hard* states, characterized by the combination of a thermal soft component (temperatures  $< 2$  keV),

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a dominating Comptonizing component with electrons in an optically thin hot plasma with temperatures of 20–50 keV, and in some cases of a third thermal continuum component likely produced by the NS surface (e.g. Gilfanov et al. 1998; Falanga et al. 2005; Gierliński & Poutanen 2005; Patruno et al. 2009; Papitto et al. 2010; Papitto et al. 2013 a). It is believed that the first soft component is the accretion disc, while the hard component may be arising from either a hot corona or the accretion columns or the boundary layer (e.g. Popham & Sunyaev 2001).

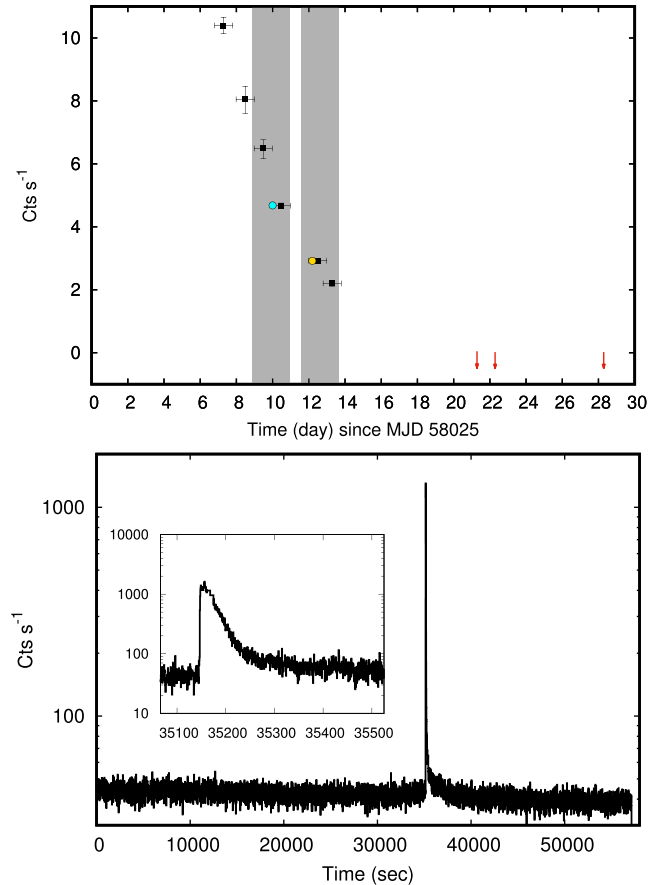
In this work, we focus on the AMXP SAX J1748.9-2021, discovered by *Beppo-SAX* in 1998 during its first recorded outburst. The source is located in the globular cluster NGC 6440 (in 't Zand et al. 1999), at a distance of  $\sim 8.5$  kpc and 0.6 kpc above the Galactic plane (Martins, Harvel & Miller 1980; Ortolani, Barbuy & Bica 1994; Kuulkers et al. 2003; Valenti, Ferraro & Origlia 2007). SAX J1748.9-2021 experienced further outbursts in 2001, 2005, 2010, and 2015 (in 't Zand et al. 1999; in 't Zand et al. 2001; Verbunt et al. 2000; Markwardt & Swank 2005; Patruno et al. 2009; Pintore et al. 2016; Sanna et al. 2016), with also a possible *a posteriori* associated outburst in 1971 (Markert et al. 1975). During quiescence, the companion star was detected (in 't Zand et al. 2001) and its mass was estimated to be in the range  $0.1\text{--}1 M_{\odot}$  (Altamirano et al. 2008). The binary system has a period of  $\sim 8.76$  h and a projected semi-major axis of  $\sim 0.4$  light-seconds (Altamirano et al. 2008; Patruno et al. 2009). SAX J1748.9-2021 showed intermittent pulsations at  $\sim 442.361$  Hz (Altamirano et al. 2008; Patruno et al. 2009; Sanna et al. 2016) and it is known to emit numerous type-I X-ray bursts (observed with RXTE and *XMM-Newton*; e.g. Galloway et al. 2008; Pintore et al. 2016). During the 2015 outburst, the source was observed in a soft state, with average spectral properties consistent with two soft thermal components plus a cold thermal Comptonized component ( $\sim 2$  keV) and an additional hard X-ray emission described by a power-law ( $\Gamma \sim 2.3$ , Pintore et al. 2016). These components were associated to the accretion disc, the NS surface, and a thermal Comptonized emission coming out of an optically thick plasma region, respectively. The origin of the high-energy tail was unclear, although a similar component has been detected in several other LMXBs during soft states (e.g. Di Salvo et al. 2000; Paizis et al. 2006; D'Ái et al. 2007; Piraino et al. 2007; Tarana et al. 2007).

In 2017, SAX J1748.9-2021 underwent its sixth outburst. The event was first detected by MAXI/GSC (Negoro et al. 2017) on September 29 and then observed by *Swift/XRT*, *INTEGRAL*, *XMM-Newton*, and *NuSTAR* (Bahramian et al. 2017; Di Gesu et al. 2017; Harita et al. 2017). Here, we present the study of the 2017 broadband spectral properties of SAX J1748.9-2021 using *XMM-Newton*, *NuSTAR*, *INTEGRAL*, and *Swift/XRT* observations.

## 2 DATA REDUCTION

We analysed *XMM-Newton* and *NuSTAR* DDT observations performed on October, 9 2017 (Obs.ID. 0795712201) and October, 11 2017 (Obs.ID. 90301320002), respectively. We also investigated all the *Swift/XRT* observations which monitored the source outburst (Fig. 1-top), as well as all the available *INTEGRAL* observation taken between October, 7 and 13.

**XMM-Newton** The *XMM-Newton* observation was taken in TIMING mode for a total exposure of  $\sim 56$  ks. We extracted EPIC-pn (Strüder et al. 2001) events only as the two EPIC-MOS cameras (Turner et al. 2001) were not pointing the source. We reduced the EPIC-pn data with SAS v15.0.0 (using the RDPHA corrections, e.g. Pintore et al. 2014), selecting single- and double-pixel



**Figure 1.** Top: background subtracted 0.3–10 keV *Swift-XRT* lightcurve (starting from the beginning of the outburst as reported by MAXI), where red arrows are  $3\sigma$  upper limits; the *XMM-Newton* and *NuSTAR* observations have been taken at MJD 58035 (cyan circle) and 58037 (orange circle), respectively. *INTEGRAL* observations were taken in the epochs inside the grey boxes. Bottom: background subtracted 0.3–10 keV *XMM/EPIC-pn* lightcurve; a type-I burst (zoomed in the inset) happened at  $\sim 35$  ks after the beginning of the observation.

events ( $\text{PATTERN} \leq 4$ ). We extracted source and background events from  $\text{RAWX}=[32:44]$  and  $\text{RAWX}=[3:5]$ , respectively. We verified that the background extracted in this region was not heavily contaminated by the source, comparing it with the background estimated from the EPIC-MOS data. During the observation, SAX J1748.9-2021 showed a type-I burst (Fig. 1-bottom).

As the average EPIC-pn count rate was  $\sim 30$  cts  $\text{s}^{-1}$ , it was well below the threshold for pile-up.<sup>1</sup> We extracted the average EPIC-pn spectrum and rebinned it with an oversample of three channels per energy resolution element using the *specgroup* task. The spectrum was analysed in the range 1.3–10 keV, to limit the EPIC-pn calibration uncertainties at low energies when operated in TIMING mode.

The RGS spectra, first and second order, were extracted using the standard *rgsproc* task. We created an average spectrum merging RGS1 and RGS2 data (using the tool *rgscombine* which combines, separately, the first- and second-order spectra). We rebinned the final spectrum with at least 100 counts per bin and we fitted it in the range 0.6–2.0 keV

<sup>1</sup><https://heasarc.gsfc.nasa.gov/docs/xmm/sas/USG/epicpileuptiming.html>

**NuSTAR** We followed the standard procedure based on NUSTARDAS (the *NuSTAR Data Analysis Software* v1.3.0) in the HEASOFT FTOOLS v6.23 to reduce the *NuSTAR* data and obtain cleaned event files. The total exposure time was 44 ks. We extracted source and background events from circular regions of radii 60'' and 90'', respectively. No type-I bursts were revealed during the observation. The FMPA and FMPB spectra were then grouped to have at least 100 counts per bin, respectively, and we fitted them in the range 3–70 keV.

**Swift** We used all the available 9 *Swift/XRT* observations collected between 2017, October 5 and 14, in both PC (the first and the last three observations of the monitoring) and WT mode (five observations between 7 and 11 October). We followed standard procedures to reduce the data and extract spectra and lightcurves.<sup>2</sup> We grouped the spectra to have at least 25 counts per bin and we fitted them in the range 0.3–10 keV. During the first *Swift* observation (taken in PC mode), pile-up was not negligible and we corrected for its effect<sup>3</sup> when we analysed this data.

**INTEGRAL** We also analysed the available *INTEGRAL* observations of the field around SAX J1748.9-2021 during the satellite revolutions 1871 and 1872. These covered the time span from 2017 October 7 at 21:39 to October 9 at 22:47 (UTC) and from 2017 October 10 at 13:29 to October 12 at 14:49 (UTC), respectively. We analysed all the publicly available *INTEGRAL* data and the data for which our group got data rights in AO14 by using version 10.2 of the Off-line Scientific Analysis software (OSA) distributed by the ISDC (Courvoisier et al. 2003). *INTEGRAL* observations are divided into ‘science windows’ (SCWs), i.e. pointings with typical durations of ~2-3 ks.

The source was within the JEM-X field of view (FoV; Lund et al. 2003) only for six SCWs during revolution 1872 and the total effective exposure time was of only 12.9 ks. As the source was located at the rim of the instrument FoV, the statistics of the data was too low to extract a meaningful spectrum. We used the JEM-X1 and JEM-X2 lightcurves with a time resolution of 2 s to search for type-I X-ray bursts, but no significant detections were found.

We used for IBIS/ISGRI (Ubertini et al. 2003; Lebrun et al. 2003) all SCWs where the source was located within 12 degree from the center of the FoV. This provided a total exposure of 58.4 ks and 44.6 ks in revolution 1871 and 1872, respectively. We extracted two ISGRI spectra, one for each revolution. Both these spectra could be well fit with a simple power-law with photon index of ~1.8. As no spectral variability could be measured between the two revolutions, we also extracted the average ISGRI spectrum by combining all data available. Finally, we fit the ISGRI spectrum in the energy band 20 keV – 100 keV.

## 3 RESULTS

### 3.1 Persistent emission

#### 3.1.1 9th October

We started our analysis with the *XMM-Newton* data, taken during the beginning of the decaying phase (9th October). Because the first part of the ISGRI observations was taken on the same epoch and no significant spectral variability was found in these data, in

order to have a broad-band spectrum we fitted simultaneously the EPIC-pn/RGS and ISGRI data.

On the basis of the spectral properties of the previous outburst, we initially adopted an absorbed NTHCOMP model (Zdziarski, Johnson & Magdziarz 1996; Życki, Done & Smith 1999) in XSPEC (v. 12.8.2; Arnaud 1996). For the absorption, we used the TBABS model with the abundances of Wilms, Allen & McCray (2000). We also added a multiplicative constant to take into account inter-calibration uncertainties between the different instruments and flux variations. We note that in the EPIC-pn spectrum there was a strong instrumental Au emission line (2.2 keV) and we fitted it with a gaussian in all the adopted models.

The TBABS × NTHCOMP fit did not provide acceptable results ( $\chi^2/dof = 969.12/764$ ), leaving several residuals especially in the band 1–10 keV.

Therefore, we added a soft component (a DISKBB model; Mitsuda et al. 1984) that significantly improved the best-fit  $\chi^2$  (899.87 for 763 dof). Its best-fit normalization, for a temperature of  $0.76 \pm 0.04$  keV and for an inclination angle of  $\sim 40^\circ$  (as found in Pintore et al. 2016), implies an implausible inner accretion disc radius of  $\sim 1.7 - 3.1$  km. Hence, this was more likely modelling the emission from the NS surface and we changed the DISKBB with a BBODYRAD model that, as expected, did not change significantly the best fit and gave a temperature of  $0.67 \pm 0.06$  keV and an emitting radius of  $2.4 \pm 0.2$  km.

The spectral residuals suggested to add also an emission line at the energy of the Iron K-shell transitions (6.4–6.9 keV), that we modelled with a gaussian. The Iron line has an energy of 6.5 keV, a width ( $\sigma$ ) of 0.1 keV (although the uncertainties are large), and an equivalent width (EW) of 0.015 keV, which we suggest may be due to the reflection off of hard photons from the inner regions of the accretion disc (e.g. Fabian et al. 1989). We then tried to substitute the GAUSSIAN model with a DISKLINE model (Fabian et al. 1989). Adopting the more recent and constrained spectral estimates reported in Pintore et al. (2016), we fixed the inclination angle to  $44^\circ$ , the emissivity index to  $-2.7$  and the outer disc radius at  $10^5$  gravitational radii: as expected, this model does not change significantly the best fit, giving an energy line of  $6.50 \pm 0.15$  keV. The inner disc radius was poorly constrained ( $R_{in} = 43^{+31}_{-43}$  km) and no robust indication can be inferred. We tentatively adopted also self-consistent reflection models as REFLIONX (Ross & Fabian 2005) or RFXCONV (Kolehmainen, Done & Díaz Trigo 2011) but they did not improve the fit and the spectral parameters could not be well constrained.

The final best-fit model (BBODYRAD + NTHCOMP + GAUSSIAN;  $\chi^2/dof = 875.85/774$ ) provided the best-fit parameters reported in Table 1 and shown in Fig. 2-left. We note that the nH is smaller than the one reported in Pintore et al. (2016). Such a discrepancy may be related to a possible variation of the local absorption during the latest outburst. Furthermore, we mention a caveat of our best-fit model as we found a degeneracy between the spectral parameters  $kT_{seed}$  and  $kT_{bb}$ , as acceptable fits can be also obtained with the BBODYRAD temperature smaller than the seed photon temperature. In the case the two temperatures were linked, the best-fit converges towards a mean temperature of  $0.53 \pm 0.2$  keV without a significant change in the other parameters but the statistical significance worsens.

Finally, we report that the averaged, unabsorbed 0.3–70 keV flux of SAX J1748.9-2021 was  $(3.6 \pm 0.03) \times 10^{-10}$  erg  $cm^{-2} s^{-1}$ , corresponding to a luminosity of  $L_x \sim 3 \times 10^{36}$  erg  $s^{-1}$  (for a distance of 8.5 kpc), i.e. ~1.6 per cent of the Eddington limit.

<sup>2</sup><http://www.swift.ac.uk/analysis/xrt/index.php>

<sup>3</sup>following the thread in <http://www.swift.ac.uk/analysis/xrt/pileup.php>

**Table 1.** Best fit spectral parameters obtained with the `BODYRAD` + `NTHCOMP` + `GAUSSIAN` model. Errors are at 90% for each parameter.

Model	Component	9th Oct. XMM+ISGRI	11th Oct. XRT+NuSTAR
TBABS	$n_H$ ( $10^{22}$ cm $^{-2}$ )	$0.51^{+0.02}_{-0.02}$	$0.62^{+0.14}_{-0.12}$
BBDYRAD	$kT_{bb}$ (keV)	$0.72^{+0.04}_{-0.06}$	$0.54^{+0.04}_{-0.04}$
	norm	$6.0^{+1.1}_{-1.3}$	$29^{+15}_{-9}$
NTHCOMP	$\Gamma$ (XMM)	$1.62^{+0.03}_{-0.04}$	$1.79^{+0.01}_{-0.01}$
	$kT_{seed}$ (keV)	$0.43^{+0.03}_{-0.03}$	= $kT_{bb}$
	$kT_e$ (keV)	$17^{+9}_{-4}$	$23^{+3}_{-4}$
	norm ( $10^{-2}$ )	$1.98^{+0.07}_{-0.07}$	$1.2^{+0.2}_{-0.2}$
GAUSSIAN	Energy (keV)	$6.54^{+0.08}_{-0.07}$	$6.54^*$
	$\sigma$ (keV)	$0.13^{+0.13}_{-0.09}$	$0.13^*$
	norm ( $10^{-5}$ )	$3.3^{+1.7}_{-1.3}$	$<2.1 \times 10^{-5}^{**}$
	EW (keV)	0.015	0.004
	$\chi^2/dof$	875.85/774	931.63/936

Notes: \* fixed to this value.

\*\* 90% upper limit on the normalization.

### 3.1.2 11th October

We then analysed the simultaneous *Swift*/*XRT* (Obs.ID: 00010341005) and *NuSTAR* spectra taken on the 11th of October. As done in the previous section, we described the broadband continuum with a `BBDYRAD` + `NTHCOMP` model, obtaining a statistically significant fit ( $\chi^2/dof = 931.24/935$ ). However, as the seed photons temperature was quite unconstrained and possibly consistent with the `BBDYRAD` temperature, we linked them in the fit (best-fit reported in Table 1 and shown in Fig. 2-right). We found a significant spectral variability in the photon index ( $\Gamma \sim 1.8$ ) and blackbody temperature ( $kT_{bb} \sim 0.55$  keV) with respect to the 9th of October observation. We calculated that the averaged, unabsorbed 0.3–70 keV flux of SAX J1748.9-2021 was  $(2.2 \pm 0.1) \times 10^{-10}$  erg cm $^{-2}$  s $^{-1}$ , i.e.  $\sim 40$  per cent lower than the flux measured with *XMM-Newton*+ISGRI. The flux on 11th October corresponds to a luminosity of  $L_x \sim 1.9 \times 10^{36}$  erg s $^{-1}$  (for a distance of 8.5 kpc), i.e. about 1 per cent of the Eddington limit.

Furthermore, we note that there is no evidence for an Iron emission line. Adding a gaussian line with energy and width fixed to those found in the previous observation, we estimated a 90 per cent upper limit on the line normalization of  $2.1 \times 10^{-5}$  photons cm $^{-2}$  s $^{-1}$  and on the equivalent width of 0.004 keV.

### 3.1.3 5th–14th October

Finally, we analysed the average spectrum of each *Swift*/*XRT* observation during the period 5th–14th October. We fitted them, in the range 0.3–10 keV, with a `TBABS`  $\times$  `NTHCOMP` model where we only fixed to 20 keV the electron temperature (as the fits were insensitive to this parameter). We note that there was no statistical need to add a soft component in any observation. The single `NTHCOMP` model provided very good results in all cases ( $\chi^2/dof \sim 1$ ), where we found that the inferred spectral parameters ( $n_H$ ,  $\Gamma$ , and  $kT_{seed}$ ) were all consistent within  $2\sigma$  with those reported in Table 1, implying no remarkable spectral variability.

## 3.2 Type-I burst

SAX J1748.9-2021 is a well known type-I X-ray burster and, during the 2017 outburst, we found at least two type-I bursts, one during the first *Swift*/*XRT* observation (at 19:32:14 UTC on October 5th) and one during the *XMM-Newton* observation (at 21:05:11 UTC on October 9th). In addition, *MAXI*/*GSC* detected a bright X-ray burst at 06:44 UTC on October 6th, which lasted for 10 s and consistent with the position of SAX J1748.9-2021 (Harita et al. 2017).

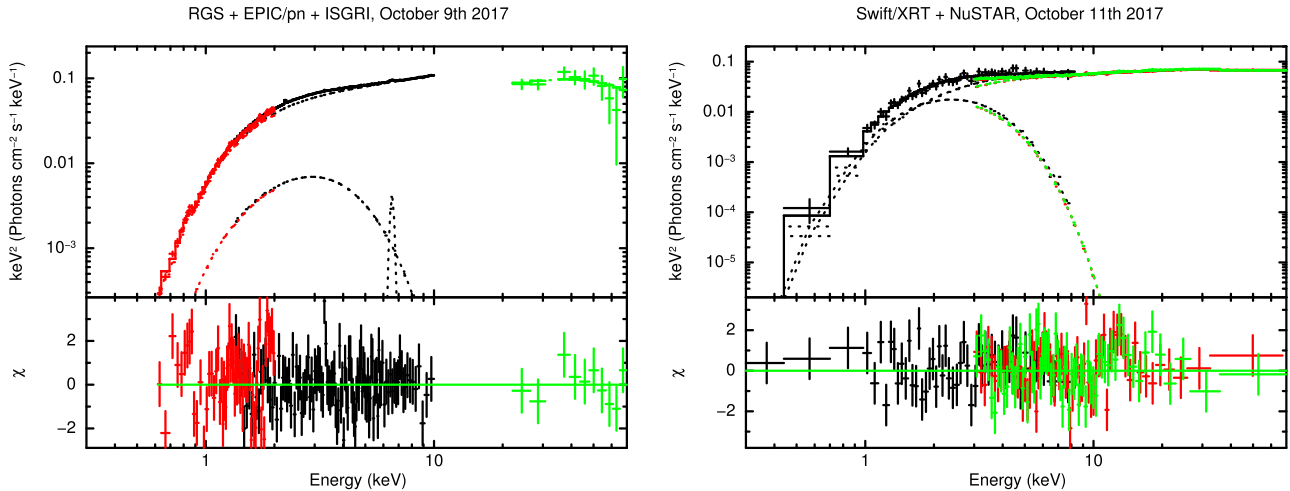
Because of the higher data quality, we focus on the burst observed in *XMM-Newton*. Modelling the burst lightcurve with an exponential function plus a constant, we found that the e-folding decay time is  $25 \pm 0.5$  s ( $1\sigma$  error), similar to those inferred from previous SAX J1748.9-2021 type-I bursts.

We then carried out a time-resolved spectroscopy, selecting a number of time intervals from the beginning of the burst to  $\sim 130$ s after it. For each interval, we extracted a spectrum (corrected for pile-up by excluding the central column `RAWX=37` from the data) and we fitted it with an absorbed `BBDYRAD` model (fixing  $n_H$  to  $0.51 \times 10^{22}$  cm $^{-2}$ ) in the range 1.0–10 keV, where we used as background the persistent emission pre-burst. The bolometric flux (estimated according to equation 3 in Galloway et al. 2008), the best-fit blackbody temperature, and the corresponding emitting radius are shown in Fig. 3. The profile shows that the flux increased, peaked at a constant value for  $\sim 10$ s, and then it decreased to the persistent emission level. A similar trend is observed for the temperature as well, which reached up to a maximum of 5 keV (although with large uncertainties). Instead, the radius of the emitting regions has a more complex trend that resembles a photospheric radius expansion (PRE) event, which is not unusual for this source (Galloway et al. 2008; Güver & Özel 2013). The mean bolometric flux registered at the peak (which lasted  $\sim 10$  s) with *XMM-Newton* was  $(50 \pm 4) \times 10^{-9}$  erg cm $^{-2}$  s $^{-1}$ .

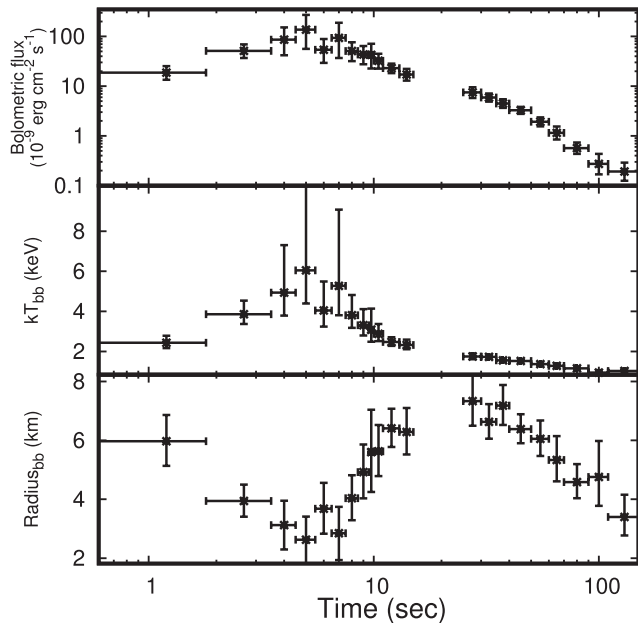
## 4 DISCUSSION

In this work, we present the 2017 outburst of SAX J1748.9-2021 which lasted for  $\sim 13$  days and with an exponential decay with e-folding decay time of  $\sim 4$  days. According to the *MAXI* lightcurve, SAX J1748.9-2021 reached the peak of the outburst between October, 4 and 5. Since the first *Swift*/*XRT* observation was taken on October, 5, we estimate that SAX J1748.9-2021 possibly reached at least an absorbed 0.3–70 keV peak flux of  $(7.0 \pm 0.7) \times 10^{-10}$  erg cm $^{-2}$  s $^{-1}$ , i.e. a factor of 2 higher than the flux measured during the latest *XMM-Newton* observation. This flux corresponds to a luminosity of  $L_x \sim 6.0 \times 10^{36}$  erg s $^{-1}$  (for a distance of 8.5 kpc), hence about 4 per cent of the Eddington luminosity.

Taking into account the marginal spectral variability during the decay, we found that the spectral properties of the 2017 outburst were characterized by a strong Comptonized emission with electron and seed photon temperatures of  $\sim 17 - 23$  keV and 0.4 keV, respectively, and photon index of 1.6–1.8. Such values correspond to an optical depth of  $\tau \sim 3 - 5$  (see equation A1 in Zdziarski et al. 1996). We estimated that, for a Compton parameter  $y \sim 3$ , a bolometric flux of  $3.4 \times 10^{-10}$  erg s $^{-1}$ , and a distance of 8.5 kpc, the radius of the region responsible for the seed photons is about 12 km (adopting the approach showed in the Discussion section of in Zand et al. 1999). It is therefore likely that the seed photons are produced in a region close to the NS, as the boundary layer or the NS surface, although we cannot rule-out contamination also from the accretion disc.



**Figure 2.** Left: unfolded ( $E^2f(E)$ ) *XMM-Newton*/EPIC-pn (black), RGS (red), and ISGRI (green) spectra fitted with an absorbed `BBODYRAD` + `NTHCOMP` + `GAUSSIAN` model. Right: *Swift*/XRT (black), *NuSTAR*/FMPA, and FMPB (red and green), fitted with an absorbed `BBODYRAD` + `NTHCOMP` model. Data have been rebinned for display purposes only.



**Figure 3.** Bolometric flux, blackbody temperature, and radius profiles of the type-I X-ray burst occurred during the *XMM-Newton* observation. The gap at  $\sim 12$ s– $25$ s is due to a temporary outage of the EPIC-pn instrument.

Furthermore, we found evidence of a significantly weaker blackbody component (although we remind here about the degeneracy between blackbody and seed photon temperature in the spectral fits discussed in Section 3.1.1) with a temperature of  $\sim 0.7$  keV and emitting radius of a few km, which origin is likely from the NS surface. This component carried  $\sim 5$  per cent and  $\sim 11$  per cent of the total source luminosity on October, 9 and October, 11, respectively.

On October, 9, it was also observed the presence of a relatively broad Iron line at 6.5 keV. We exclude that the broadening is either compatible with Compton down-scattering in a wind shell ejected at moderate relativistic velocities from the disc (Laurent & Titarchuk 2007) or Compton processes in an accretion disc corona above the accretion disc (White & Holt 1982; Kallman & White 1989; Vrtilik, Soker & Raymond 1993). Indeed, should these processes be

responsible for the broadening, the corresponding electron temperature can be obtained from the relation  $\Delta\epsilon/\epsilon = (4kT_e - \epsilon)/m_e c^2$  (where  $\Delta\epsilon$  is the line broadening and  $\epsilon$  is its energy). However, we estimated an electron temperature of  $\sim 4$  keV which is a factor of 4–5 lower than that found from our spectral analysis. Even if we consider the 90 per cent upper uncertainties on the broadening and energy line (i.e. 0.26 keV and 6.62 keV), we inferred an upper limit on the electron temperature of  $\sim 7$  keV, still well below our value. For this reason, we associate the line broadening to reflection processes from the surface of the accretion disc (see e.g. Fabian et al. 1989). Using the `DISKLINE` model, we could not constrain the dimension of the inner disc radius, and thus no definitive conclusion can be inferred on this parameter. Similar results are also obtained with self-consistent reflection models.

We remark that this is the sixth registered outburst of SAX J1748.9-2021, but its maximum peak luminosity is the weakest ever reported for this source to date (although *Swift* may underestimate the broad-band source flux). In fact, all the other five events were significantly brighter, with the 1998 one being the shortest in duration and the lowest in peak luminosity (Altamirano et al. 2008). However the latter is still at least a factor of 3 higher than the peak luminosity of the 2017 outburst. For the 1998 outburst, using *BeppoSAX* data, in 't Zand et al. (1999) found that the 0.1–100 keV spectrum of SAX J1748.9-2021 was consistent with a single thermal Comptonization model with electron temperature of 15 keV, optical depth in the range 3–6, and seed photon temperature of 0.6 keV (coming from a region of radius of  $\sim 13$  km). This indicates that the source was in a *hard* state. Based on our analysis, we suggest that SAX J1748.9-2021 was again in a *hard* state during the 2017 outburst (and possibly all along the decay as the *Swift*/XRT monitoring did not show any clear evidence of spectral variability) making the two outbursts very similar. Pulsations have been mainly observed during the hard states of SAX J1748.9-2021 and we confirm a pulse detection, that we will report in greater detail elsewhere (Sanna et al. in prep.).

The last outburst was hence clearly different from the 2015 one, when the source reached  $\sim 25$  per cent of the Eddington luminosity, and its spectral properties were consistent with a marked *soft* state (Pintore et al. 2016). In particular, it was characterized by the combination of four spectral continuum components, representing

the disc emission, the NS surface, a Comptonized region (possibly the boundary layer) and an additional non-thermal emission (described with a hard power-law), carrying  $\sim 20$  per cent, 25 per cent, 51 per cent, and 3 per cent of the total luminosity, respectively. In addition, the NS surface emission in 2015 had a temperature of  $\sim 1.1$  keV, about a factor of 2 higher than the temperature estimated in this work. In 2017, we also did not find any evidence of a hard power-law tail at high energy, confirming the *hard* state of the source, as such a component is generally seen only during *soft* states of LMXBs (e.g. Di Salvo et al. 2000, 2001; D’Amico et al. 2001; Iaria et al. 2004; Di Salvo et al. 2006; Paizis et al. 2006; D’Aí et al. 2007; Piraino et al. 2007; Tarana et al. 2007).

At least three type-I X-ray bursts were observed from SAX J1748.9-2021 with MAXI, *Swift/XRT*, and *XMM-Newton* in its 2017 outburst unlike the 1998 outburst where none of them were found. According to the linear relation between pre-burst count rate and elapsed time reported in Pintore et al. (2016) for the 2015 outburst, we predict that the expected recurrence time of the type-I bursts during the 2017 outburst would have been at least one every  $\sim 8 - 9$  ks, much shorter than the observed recurrence time. During the burst, the measured NS surface temperature increases up to  $\sim 5$  keV (although the uncertainties are quite large). On the other hand, the radius of the emitting area presents a more complex trend: in fact, it is  $\sim 6$  km at the beginning of the burst and it decreases below  $\sim 4$  km during the burst rising phase; at the maximum flux, the radius increased again to  $\sim 6$  km, remaining constant for at least 60 seconds. This behaviour is similar to the *radius expansion* process, i.e. the response of the NS outermost layers to a super-Eddington burst flux (see e.g. Kuulkers et al. 2003). In particular, the NS photosphere increases and expands and then contracts back to the NS surface. It was found that during the expansion/contraction phase the flux remains nearly constant to the Eddington limit, converting any excess into kinetic energy of the outflow (e.g. Ebisuzaki et al. 1983; Kato 1983; Paczynski & Proszynski 1986). We note that radius expansions are not unusual for SAX J1748.9-2021 (e.g. Güver, Özel & Psaltis 2012) although in the 2015 outburst they were not observed in any of the detected type-I X-ray bursts. Assuming that the mean flux at the peak of the burst reported here ( $5.0 \times 10^{-8}$  erg  $\text{cm}^{-2}$   $\text{s}^{-1}$ ) reached the Eddington limit and using the Eddington luminosity empirically obtained by Kuulkers et al. (2003) ( $3.8 \times 10^{38}$  erg  $\text{s}^{-1}$ ), we estimated a source distance of  $8.0 \pm 0.4$  kpc and  $5.2 \pm 0.4$  kpc (error at  $1\sigma$ ) for a pure helium ( $X = 0$ ) or a solar composition burst ( $X = 0.7$ ), respectively, which are well consistent with the previous distance estimates (Galloway et al. 2008).

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## REFERENCES

- Altamirano D., Casella P., Patruno A., Wijnands R., van der Klis M., 2008, *ApJ*, 674, L45
- Archibald A. M., Stairs I. H., Ransom S. M., Kaspi V. M. et al., 2009, *Science*, 324, 1411
- Arnaud K. A., 1996, in Jacoby G. H., Barnes J., eds., *Astronomical Data Analysis Software and Systems V*. Vol. 101 of ASP Conf. Ser., San Francisco CA, XSPEC: The First Ten Years. p. 17
- Bahramian A. et al., 2017, *The Astronomer’s Telegram*, 10826
- Burderi L., Di Salvo T., 2013, *Mem. Soc. Astron. Italiana*, 84, 117
- Burderi L., King A. R., 1998, *ApJ*, 505, L135
- Burderi L., Di Salvo T., Menna M. T., Riggio A., Papitto A., 2006, *ApJ*, 653, L133
- Burderi L., Di Salvo T., Lavagetto G., Menna M. T. e. a., 2007, *ApJ*, 657, 961
- Cackett E. M. et al., 2010, *ApJ*, 720, 205
- Campana S., Di Salvo T., 2018, Chapter of the book ‘Physics and Astrophysics of Neutron Stars’, *NewCompStar COST Action*, 1304
- Courvoisier T. J.-L., Walter R., Beckmann V., Dean A. J., Dubath P. e. a., 2003, *A&A*, 411, L53
- D’Aí A., Życki P., Di Salvo T., Iaria R., Lavagetto G., Robba N. R., 2007, *ApJ*, 667, 411
- D’Amico F., Heindl W. A., Rothschild R. E., Gruber D. E., 2001, *ApJ*, 547, L147
- Di Gesu L., Bozzo E., Kuulkers E., Bazzano A. e. a., 2017, *The Astronomer’s Telegram*, 10832
- Di Salvo T. et al., 2000, *ApJ*, 544, L119
- Di Salvo T., Robba N. R., Iaria R., Stella L., Burderi L., Israel G. L., 2001, *ApJ*, 554, 49
- Di Salvo T. et al., 2006, *ApJ*, 649, L91
- Di Salvo T., Burderi L., Riggio A., Papitto A., Menna M. T., 2008, in Yuan Y.-F., Li X.-D., Lai D., eds, *Astrophysics of Compact Objects Vol. 968 of American Institute of Physics Conference Series*, Timing of accreting millisecond pulsars. p. 233
- Ebisuzaki T., Hanawa T., Sugimoto D., 1983, *PASJ*, 35, 17
- Fabian A. C., Rees M. J., Stella L., White N. E., 1989, *MNRAS*, 238, 729
- Falanga M. et al., 2005, *A&A*, 444, 15
- Galloway D. K., Munro M. P., Hartman J. M., Psaltis D., Chakrabarty D., 2008, *ApJS*, 179, 360
- Ghosh P., Becker W., 2008, *Physics Today*
- Gierliński M., Poutanen J., 2005, *MNRAS*, 359, 1261
- Gilfanov M., Revnivtsev M., Sunyaev R., Churazov E., 1998, *A&A*, 338, L83
- Güver T., Özel F., 2013, *ApJ*, 765, L1
- Güver T., Özel F., Psaltis D., 2012, *ApJ*, 747, 77
- Harita S., Sugita S., Sakamaki A., Negoro H. e. a., 2017, *The Astronomer’s Telegram*, 10827
- Hartman J. M. et al., 2008, *ApJ*, 675, 1468
- Iaria R., Di Salvo T., Robba N. R., Burderi L., Stella L., Frontera F., van der Klis M., 2004, *ApJ*, 600, 358
- in ’t Zand J. J. M. et al., 1999, *A&A*, 345, 100
- in ’t Zand J. J. M., van Kerkwijk M. H., Pooley D., Verbunt F., Wijnands R., Lewin W. H. G., 2001, *ApJ*, 563, L41
- Kallman T., White N. E., 1989, *ApJ*, 341, 955
- Kato M., 1983, *PASJ*, 35, 33
- Kolehmainen M., Done C., Díaz Trigo M., 2011, *MNRAS*, 416, 311
- Kuulkers E., den Hartog P. R., in ’t Zand J. J. M., Verbunt F. W. M., Harris W. E., Cocchi M., 2003, *A&A*, 399, 663
- Laurent P., Titarchuk L., 2007, *ApJ*, 656, 1056
- Lebrun F., Leray J. P., Lavocat P., Crétole J., Arquès M., Blondel C. e. a., 2003, *A&A*, 411, L141
- Lund N., Budtz-Jørgensen C., Westergaard N. J., Brandt S., Rasmussen I. L. e. a., 2003, *A&A*, 411, L231
- Markert T. H., Backman D. E., Canizares C. R., Clark G. W., Levine A. M., 1975, *Nature*, 257, 32
- Markwardt C. B., Swank J. H., 2005, *The Astronomer’s Telegram*, 495, 1
- Martins D. H., Harvel C. A., Miller D. H., 1980, *AJ*, 85, 521

- Mitsuda K. et al., 1984, *PASJ*, 36, 741
- Negoro H., Ueno S., Tomida H., Ishikawa M. e. a., 2017, *The Astronomer's Telegram*, 10821
- Ortolani S., Barbay B., Bica E., 1994, *A&AS*, 108, 653
- Paczynski B., Proszynski M., 1986, *ApJ*, 302, 519
- Paizis A. et al., 2006, *A&A*, 459, 187
- Papitto A., Di Salvo T., D'Ai A., Iaria R., Burderi L., Riggio A., Menna M. T., Robba N. R., 2009, *A&A*, 493, L39
- Papitto A., Riggio A., di Salvo T., Burderi L., D'Ai A., Iaria R., Bozzo E., Menna M. T., 2010, *MNRAS*, 407, 2575
- Papitto A. et al., 2013a, *MNRAS*, 429, 3411
- Papitto A., Ferrigno C., Bozzo E., Rea N., Pavan L., Burderi L. et al., 2013b, *Nature*, 501, 517
- Patruno A., Watts A. L., 2012, preprint ([arXiv: 1206.2727](https://arxiv.org/abs/1206.2727))
- Patruno A., Rea N., Altamirano D., Linares M., Wijnands R., van der Klis M., 2009, *MNRAS*, 396, L51
- Patruno A., Altamirano D., Hessels J. W. T., Casella P., Wijnands R., van der Klis M., 2009, *ApJ*, 690, 1856
- Patruno A., Haskell B., Andersson N., 2017, *ApJ*, 850, 106
- Pintore F. et al., 2014, *MNRAS*, 445, 3745
- Pintore F. et al., 2016, *MNRAS*, 457, 2988
- Piraino S., Santangelo A., di Salvo T., Kaaret P., Horns D., Iaria R., Burderi L., 2007, *A&A*, 471, L17
- Popham R., Sunyaev R., 2001, *ApJ*, 547, 355
- Ross R. R., Fabian A. C., 2005, *MNRAS*, 358, 211
- Sanna A. et al., 2016, *MNRAS*, 459, 1340
- Sanna A., Bahramian A., Bozzo E., Heinke C. et al., 2018, *A&A*, 610, L2
- Strohmer T. E. et al., 2018, *The Astronomer's Telegram*, 11507
- Strüder L. et al., 2001, *A&A*, 365, L18
- Tarana A., Bazzano A., Ubertini P., Zdziarski A. A., 2007, *ApJ*, 654, 494
- Turner M. J. L. et al., 2001, *A&A*, 365, L27
- Ubertini P., Lebrun F., Di Cocco G., Bazzano A., Bird A. J., Broenstad K. e. a., 2003, *A&A*, 411, L131
- Valenti E., Ferraro F. R., Origlia L., 2007, *AJ*, 133, 1287
- Verbunt F., van Kerkwijk M. H., in't Zand J. J. M., Heise J., 2000, *A&A*, 359, 960
- Vrtilek S. D., Soker N., Raymond J. C., 1993, *ApJ*, 404, 696
- White N. E., Holt S. S., 1982, *ApJ*, 257, 318
- Wilkinson T., Patruno A., Watts A., Uttley P., 2011, *MNRAS*, 410, 1513
- Wilms J., Allen A., McCray R., 2000, *ApJ*, 542, 914
- Zdziarski A. A., Johnson W. N., Magdziarz P., 1996, *MNRAS*, 283, 193
- Życki P. T., Done C., Smith D. A., 1999, *MNRAS*, 309, 561

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