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The ALPINE–ALMA [C II] Survey: Multiwavelength Ancillary Data and Basic Physical Measurements

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Abstract

We present the ancillary data and basic physical measurements for the galaxies in the ALMA Large Program to Investigate C⁺ at Early Times (ALPINE) survey—the first large multiwavelength survey that aims at characterizing the gas and dust properties of 118 main-sequence galaxies at redshifts $4.4 < z < 5.9$ via the measurement of [C II] emission at $158 \mu\text{m}$ (64% at $>3.5\sigma$) and the surrounding far-infrared continuum in conjunction with a wealth of optical and near-infrared data. We outline in detail the spectroscopic data and selection of the galaxies as well as the ground- and space-based imaging products. In addition, we provide several basic measurements including stellar masses, star formation rates (SFR), rest-frame ultra-violet (UV) luminosities, UV continuum slopes (β), and absorption line redshifts, as well as H α emission derived from *Spitzer* colors. We find that the ALPINE sample is representative of the $4 < z < 6$ galaxy population selected by photometric methods and only slightly biased toward bluer colors ($\Delta\beta \sim 0.2$). Using [C II] as tracer of the systemic redshift (confirmed for one galaxy at $z = 4.5$ out of 118 for which we obtained optical [O III] $\lambda 3727\text{\AA}$ emission), we confirm redshifted Ly α emission and blueshifted absorption lines similar to findings at lower redshifts. By stacking the rest-frame UV spectra in the [C II] rest frame, we find that the absorption lines in galaxies with high specific SFR are more blueshifted, which could be indicative of stronger winds and outflows.

Unified Astronomy Thesaurus concepts: Galaxy evolution (594); Galactic and extragalactic astronomy (563); Interstellar medium (846); Star formation (1569); Galaxy photometry (611); Stellar masses (1614); Interstellar medium wind (848); Galaxy winds (626); Galaxy processes (614); Stellar winds (1636); Far infrared astronomy (529)

1. Introduction

1.1. The Early Growth Phase in Galaxy Evolution

Galaxy evolution undergoes several important phases such as the ionization of neutral Hydrogen at redshifts $z > 6$ (also

known as the Epoch of Reionization) as well as a time of highest cosmic star formation rate (SFR) density at $z \sim 2\text{--}3$. The transition phase at $z = 4\text{--}6$ (a time roughly 0.9 to 1.5 billion years after the Big Bang), often referred to as the *early growth phase*, is currently in focus of many studies. This time

is of great interest for understanding galaxy evolution as it connects primordial galaxy formation during the epoch of reionization with mature galaxy growth at and after the peak of cosmic SFR density. During a time of only 600 Myr, the cosmic stellar mass density in the universe increased by one order of magnitude (Caputi et al. 2011; Davidzon et al. 2017), galaxies underwent a critical morphological transformation to build up their disk and bulge structures (Gnedin et al. 1999; Bournaud et al. 2007; Agertz et al. 2009), and their interstellar medium (ISM) became enriched with metal from sub-solar to solar amounts (Ando et al. 2007; Faisst et al. 2016b), while at the same time the dust attenuation of the UV light significantly increased (Finkelstein et al. 2012; Bouwens et al. 2015; Fudamoto et al. 2017; Popping et al. 2017; Cullen et al. 2018; Ma et al. 2019; Yamanaka & Yamada 2019). Furthermore, the most massive of these galaxies may become the first quiescent galaxies already at $z > 4$ (Glazebrook et al. 2017; Tanaka et al. 2019; Faisst et al. 2019; Stockmann et al. 2020; Valentino et al. 2020). All of this put together, makes the early growth phase an important puzzle piece to be studied in order to decipher how galaxies formed and evolved to become the galaxies (either star-forming or quiescent) that we observe in the local universe.

It is evident from studies at lower redshift that multi-wavelength observations are crucial for us to be able to form a coherent picture of galaxy evolution. To capture several important properties of galaxies, a *panchromatic* survey must comprise several spectroscopic and imaging data sets that cover a large fraction of the wavelength range of a galaxy's light emission, including (i) the rest-frame ultra-violet (UV) containing Ly α emission, as well as several absorption lines to study stellar winds and metallicity (Heckman et al. 1997; Maraston et al. 2009; Steidel et al. 2010; Faisst et al. 2016b), (ii) the rest-frame optical containing tracers of age (Balmer break) as well as important emission lines (e.g., H α) to quantify the star formation and gas metal properties (Kennicutt 1998; Kewley & Ellison 2008), and (iii) the far-infrared (FIR) continuum and several FIR emission lines (e.g., [C II] $\lambda 158 \mu\text{m}$ or [N II] $\lambda 205 \mu\text{m}$) that provide insights into the gas and dust properties of galaxies (De Looze et al. 2014; Pavesi et al. 2019).

Fortunately, the early growth phase at redshifts $z = 4\text{--}6$ is at the same time the highest-redshift epoch at which, using current technologies, such a panchromatic study can be carried out. The rest-frame UV part of the energy distribution at these redshifts has been probed in the past thanks to several large spectroscopic (Le Fèvre et al. 2015; Hasinger et al. 2018) and imaging (Capak et al. 2007; McCracken et al. 2012; Aihara et al. 2019) surveys from the ground as well as imaging surveys with the *Hubble Space Telescope* (*HST*; Scoville et al. 2007a; Grogin et al. 2011; Koekemoer et al. 2011). In addition, H α has been accessed successfully through observations with the *Spitzer Space Telescope* (Stark et al. 2013; de Barros et al. 2014; Smit et al. 2014; Faisst et al. 2016b, 2019; Rasappu et al. 2016; Smit et al. 2016; Lam et al. 2019). However, the FIR of $z > 4$ galaxies has only been probed sparsely in the past in less than a dozen galaxies using the Atacama Large (Sub-) Millimeter Array (ALMA; Riechers et al. 2014; Capak et al. 2015; Watson et al. 2015; Willott et al. 2015; Strandet et al. 2017; Carniani et al. 2018; Zavala et al. 2018b, 2018a; Casey et al. 2019; Jin et al. 2019) as well as some as part of Herschel surveys in lensed and unlensed fields (e.g., Egami et al. 2010; Casey et al. 2012, 2014; Combes et al. 2012). Commonly targeted by observations

with ALMA is singly ionized Carbon (C⁺) at $158 \mu\text{m}$, which is an important coolant for the gas in galaxies and is therefore broadly related to star formation activity and gas masses (Stacey et al. 1991; Carilli & Walter 2013; De Looze et al. 2014). The [C II] emission line is one of the strongest in the FIR and is in addition conveniently located in the ALMA Band 7 at redshifts $z = 4\text{--}6$ at one of the highest atmospheric transmissions compared to other FIR lines (see, e.g., Faisst et al. 2017). The origin of [C II] emission is still debated. In addition to photodissociation regions (PDRs) and the cold neutral medium, a significant fraction can also originate from ionized gas regions or CO-dark molecular clouds (Pineda et al. 2013; Vallini et al. 2015; Pavesi et al. 2016). Also, the increasing temperature of the Cosmic Microwave Background has an effect on the relation between [C II] and star formation (Ferrara et al. 2019). Both potentially complicates the interpretation of [C II] as SFR indicator at high redshifts. Similar to H α , [C II] traces the gas kinematics in a galaxy and is therefore an important component to quantify rotation- and dispersion-dominated systems as well as outflows (Jones et al. 2017; Pavesi et al. 2018; Kohandel et al. 2019; Ginolfi et al. 2020).

The FIR landscape has dramatically changed with the completion of the ALMA Large Program to Investigate C⁺ at Early Times (ALPINE; #2017.1.00428.L). ALPINE is laying the ground work for the exploration of gas and dust properties in 118 *main-sequence* star-forming galaxies in the early growth phase at $4.4 < z < 5.9$ and herewith started the first panchromatic survey of its kind at these redshifts (Figure 1).

1.2. ALPINE in a Nutshell

In the following, we summarize the scope of the ALPINE survey, we refer to Le Fèvre et al. (2019) for a broader overview of the program. ALPINE is a 69 hr large ALMA program started in Cycle 5 in 2018 May and completed during Cycle 6 in 2019 February. In total, 118 galaxies have been observed in Band 7 (covering [C II] emission at $158 \mu\text{m}$ and its nearby continuum) at a spatial resolution of $< 1''$ and with integration times ~ 30 minutes on-source depending on their predicted [C II] flux. The galaxies originate from two fields, namely the Cosmic Evolution Survey field (COSMOS; 105 galaxies, Scoville et al. 2007b) and the Extended Chandra Deep Field South (ECDFS; 13 galaxies, Giacconi et al. 2002). Due to gaps in the transition through the atmosphere, the galaxies are split into two different redshift ranges spanning $4.40 < z < 4.65$ and $5.05 < z < 5.90$ with medians of $\langle z \rangle = 4.5$ and 5.5 and galaxy numbers of 67 and 51, respectively. All galaxies are spectroscopically confirmed by either Ly α emission or rest-UV absorption lines and are selected to be brighter than an absolute UV magnitude of $M_{1500} = -20.2$. This limit is roughly equivalent to an SFR cut at $10 M_{\odot} \text{ yr}^{-1}$ and corresponds roughly to a limiting luminosity in [C II] emission of $L_{[\text{C II}]} = 1.2_{-0.9}^{+1.9} \times 10^8 L_{\odot}$ (assuming the relation derived by De Looze et al. 2014). Assuming a 3.5σ detection limit, the [C II] detection rate is 64%, and continuum emission is detected in 19% of the galaxies (see Figure 2).

The main science goals enabled by ALPINE are diverse and cover many crucial research topics at high redshifts:

1. Connecting [C II] to star formation at high redshifts,
2. Coherent study of the total SFR density at $z > 4$ including the contribution of dust-obscured star formation,

Layout of Current Data Products for ALPINE Galaxies

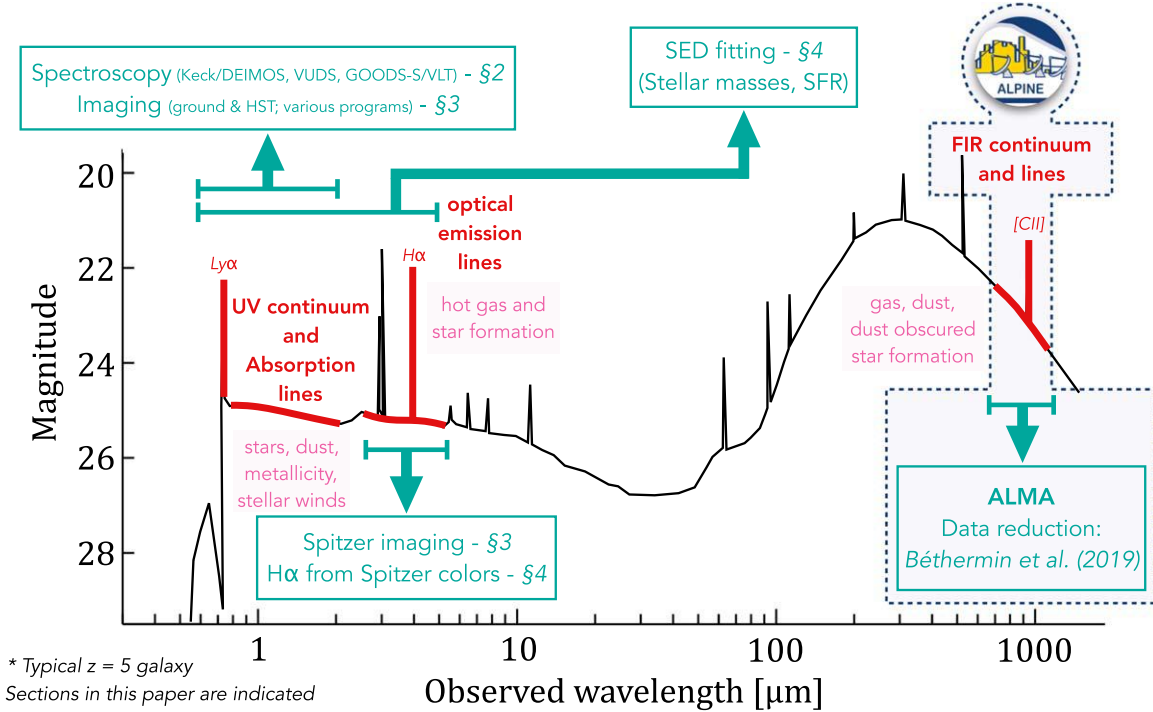


Figure 1. ALPINE builds the corner stone of a panchromatic survey at $z = 4-6$. The diagram shows the multiwavelength data products that are currently available for all of the ALPINE galaxies. The currently covered parts of the spectrum are indicated in red. The numbers link to sections in this paper where the data products and their analysis are explained in detail. The spectrum sketch is based on a typical $z = 5$ galaxy (adapted from Harikane et al. 2018).

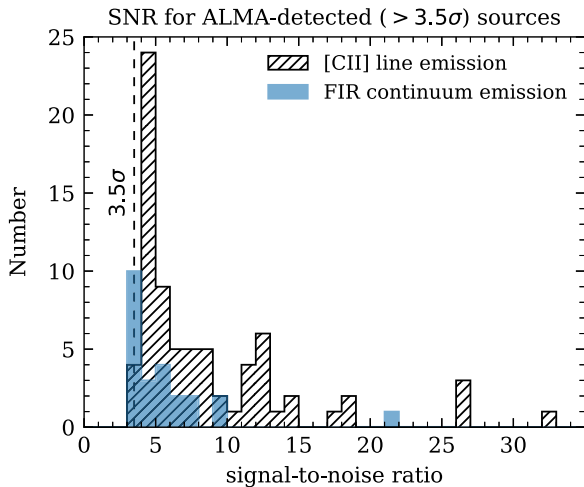


Figure 2. Signal-to-noise ratio of the ALMA-detected sources in the ALPINE sample. The different histograms show the numbers for [C II] and continuum detections above 3.5σ . For more information, see Le Fèvre et al. (2019) and Béthermin et al. (2020).

3. Study of gas dynamics and merger statistics from [C II] kinematics and quantification of UV-faint companion galaxies,
4. Study of gas fractions and dust properties at $z > 4$,
5. The first characterization of ISM properties using $L_{\text{FIR}}/L_{\text{UV}}$ and [C II]/FIR continuum diagnostics for a large sample at $z > 4$,
6. Quantifying outflows and feedback processes in $z > 4$ galaxies from [C II] line profiles.

Note that ALPINE provides at the same time the equivalent of a blind-survey of approximately 25 square-arcminutes. This enables us to estimate the obscured fraction of star formation (mostly below $z = 4$) by finding UV-faint galaxies with FIR continuum or [C II] emission. The serendipitous continuum sources and [C II] detections are discussed in detail in Béthermin et al. (2020) and F. Loiacono et al. (2020, in preparation). A more detailed description of these science goals can be found in our survey overview paper (Le Fèvre et al. 2019).

ALPINE is based on a rich set of ancillary data, which makes it the first panchromatic survey at these high redshifts including imaging and spectroscopic observations at FIR wavelengths (see Figure 1). The backbone for a successful selection of galaxies are rest-frame UV spectroscopic observations from the Keck telescope in Hawaii as well as the European Very Large Telescope (VLT) in Chile. These are complemented by ground-based imaging observations from rest-frame UV to optical, *HST* observations in the rest-frame UV, and *Spitzer* coverage above the Balmer break at rest-frame 4000 Å. The latter is crucial for the robust measurement of stellar masses at these redshifts (e.g., Faisst et al. 2016a).

For a survey overview of ALPINE see Le Fèvre et al. (2019) and for details on the data analysis see Béthermin et al. (2020). In this paper, we present these valuable ancillary data products and detail several basic measurements for the ALPINE galaxies. The outline of the paper is sketched in Figure 1. Specifically, in Section 2, we present the spectroscopic data and detail the spectroscopic selection of the ALPINE galaxies. In the same section, we also present stacked spectra and touch on velocity offsets between Ly α , [C II], and absorption line

redshifts. Section 3 is devoted to the photometric data products, which include ground- and space-based photometry. In Section 4.1, we detail the derivation of several galaxy properties from the observed photometry. These include stellar masses, SFRs, UV luminosities, UV continuum slopes, as well as H α emission derived from *Spitzer* colors. We conclude and summarize in Section 5. All presented data products are available in the online printed version of this paper.²⁸ The different catalogs and their columns are described in detail in the Appendix A. *HST* cutouts and rest-frame UV spectra for each of the ALPINE galaxies are shown in Appendix B.

Throughout the paper, we assume the Λ CDM cosmology with $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_\Lambda = 0.70$, and $\Omega_m = 0.30$. All magnitudes are given in the AB system (Oke 1974) and stellar masses and SFRs are normalized to a Chabrier (2003) initial mass function (IMF).

2. Spectroscopic Data and Selection

2.1. Spectroscopic Selection of ALPINE Galaxies

The ALPINE survey is only possible due to a spectroscopic pre-selection of galaxies from large spectroscopic surveys on COSMOS and ECDFS. This is because the ALMA frequency bands are narrow ($\sim 1000 \text{ km s}^{-1}$), and in order to observe [C II] emission, the redshift has to be known within a precision of $\sim 1000 \text{ km s}^{-1}$. The galaxy selection is refined to optimize the efficiency of the ALMA observations by creating groups of galaxies in spectral dimensions. Our sample also includes seven galaxies that were previously observed with ALMA by Riechers et al. (2014) and Capak et al. (2015). These are *HZ1*, *HZ2*, *HZ3*, *HZ4*, *HZ5*, *HZ6/LBG-1*, and *HZ8*, which correspond to the ALPINE galaxies *DC_536534*, *DC_417567*, *DC_683613*, *DC_494057*, *DC_845652*, *DC_848185*, and *DC_873321*, respectively. Furthermore, four galaxies from the VUDS survey (*vc_5101288969*, *vc_5100822662*, and *vc_510786441* in COSMOS and *ve_530029038* in ECDFS) are observed twice (resulting in a total number of 122 observations). The duplicate observations are used for quality assessment. Bethermin et al. (2020) describes the combination of these observations.

The rest-frame UV spectroscopic data from which the ALPINE sample is selected combine various large surveys on the COSMOS and ECDFS fields. Out of the 105 ALPINE galaxies on the COSMOS field, 84 are obtained by the large DEIMOS spectroscopic survey (Capak et al. 2004; Mallery et al. 2012; Hasinger et al. 2018) at the Keck telescope in Hawaii. The remaining spectra on the COSMOS field are obtained from the VIMOS Ultra Deep Survey (VUDS; Le Fèvre et al. 2015; Tasca et al. 2017) at the VLT in Chile. In total, six of the VUDS spectra are independently also observed as part of the Keck/DEIMOS survey (*vc_5100559223*, *vc_5100822662*, *vc_5101218326*, *vc_5101244930*, *vc_5101288969*, *vc_510786441*). The redshifts are consistent within 280 km s^{-1} , and we do not find any systematic offsets between the two observations (see also Section 2.4.1). Out of the 13 galaxies in the ECDFS field, 11 are obtained from spectroscopic observations with VIMOS (9) and FORS2 (2²⁹) at the VLT (Vanzella et al. 2007, 2008; Balestra et al. 2010), and two come from the *HST* grism survey *GRAPES* (Malhotra et al. 2005; Rhoads et al. 2009). The spectral resolution of the different data set varies between

$R \sim 100$ (ECDFS/*GRAPES* grism), $R \sim 180$ (ECDFS/VIMOS), $R \sim 230$ (COSMOS/VUDS), $R \sim 660$ (ECDFS/FORS2), and $R \sim 2500$ (COSMOS/DEIMOS).

Biases toward dust-poor star-forming galaxies with strong rest-frame UV emission lines (such as Ly α) can be common in purely spectroscopically selected samples. To minimize such biases as much as possible, the spectroscopically observed galaxies have been pre-selected through a variety of different selection methods. The largest fraction of galaxies in ALPINE is drawn from the Keck/DEIMOS and VUDS surveys on the COSMOS field. Both surveys include galaxies pre-selected in various ways, resulting in the most representative and inclusive spectroscopic high-redshift galaxy sample. Specifically, the VUDS survey combines predominantly a photometric redshift selection with a color-selected Lyman Break Galaxy (LBG) selection (Le Fèvre et al. 2015), known as the Lyman-break drop-out technique (see, e.g., Steidel et al. 1996; Dickinson 1998). The Keck/DEIMOS survey (providing 71% of the total ALPINE sample) consists of galaxies that are selected by narrowband surveys at $z \sim 4.5$ (7%) and $z \sim 5.7$ (27%), the drop-out technique (color selection) over the whole redshift range (49%), as well as purely by photometric redshifts (11%). In addition, four galaxies are selected by a $4.5 \mu\text{m}$ excess and one galaxy was pre-selected through X-ray emission using the *Chandra Observatory*. On the ECDFS field, the galaxies are mostly color-selected. Table 1 summarizes the different selections and corresponding numbers of galaxies and provides a complete list of references. We also list the numbers of galaxies with Ly α emission (76%) and weak Ly α emission or Ly α absorption ($\sim 24\%$). Note that the Keck/DEIMOS and VUDS samples have similar Ly α emission properties. However, note that above $z = 5$, the ALPINE sample is strongly dominated by narrowband-selected galaxies.

Figure 3 shows the distribution of redshifts of the ALPINE galaxies in the COSMOS and ECDFS fields. The colored histogram bars show stacked numbers of galaxies that are pre-selected by the different methods discussed above. The bins with galaxies in the ECDFS field are hatched. The narrowband-selected galaxies are prominent at $z \sim 5.7$ and represent the largest fraction of galaxies at $z > 5$ in ALPINE. On the other hand, the $z < 5$ sample consists mostly of color-selected galaxies. The VUDS galaxies are most represented at $z < 5$, while the DEIMOS spectra and the galaxies in ECDFS cover the whole redshift range.

Figure 4 shows the distribution of observed magnitudes as well as rest-frame 1500 \AA and $\sim 4000 \text{ \AA}$ luminosity of galaxies selected by the different methods. The photometry that is used is explained in detail in Section 3. The 1500 \AA rest-frame luminosity is derived from spectral energy distribution (SED) fitting (see Section 4.2 for details). The 4000 \AA rest-frame luminosity is derived directly from the UltraVISTA K_s and VLT K_s^V magnitude for galaxies in the COSMOS and GOODS-S field, respectively. The magnitudes and luminosities are not corrected for dust attenuation. Note that the K -band is rest-frame 3000 \AA at the highest redshifts ($z = 5.9$); hence at these redshifts, older and dustier galaxies would be biased to lower luminosities. As expected for spectroscopically selected galaxies, the ALPINE sample covers the brighter part of the galaxy magnitude and luminosity distribution. The different selection methods on their own are distributed differently in this parameter space. Most noticeably, the $z \sim 5.7$ narrowband-selected galaxies reside at the faintest luminosities, while the $4.5 \mu\text{m}$ continuum excess

²⁸ <https://caltech.box.com/s/eo4v0zaphnf3t0v1761pcos5gofmxx6>

²⁹ One of these galaxies, *ve_530029038*, has also been observed by the VUDS survey.

Table 1
Spectroscopy and Selection of ALPINE Galaxies

Survey	Selection	Number	Ref.
	COSMOS field (105 galaxies)		
Keck/DEIMOS ^a		84	1
	narrowband ($z \sim 4.5$) ^c	6	
	narrowband ($z \sim 5.7$) ^d	23	
	LBG (color) ^e	41	
	pure photo- z ^f	9	
	4.5 μm excess	4	
	X-ray (<i>Chandra</i>)	1	
	with Ly α emission	66	
	weak Ly α emission or absorption	18	
VUDS		21	2
	photo- z + LBG	21	
	[narrowband ($z \sim 4.5$)	3] ^b	
	[narrowband ($z \sim 5.7$)	1] ^b	
	[LBG (color)	1] ^b	
	[4.5 μm excess	1] ^b	
	with Ly α emission	16	
	weak Ly α emission or absorption	5	
		ECDFS field (13 galaxies)	
VLT GOODS-S		11	3
	primarily LBG (color)	11	
	total with Ly α emission	6	
	total without Ly α emission	5	
<i>HST</i> /GRAPES		2	4
	Grism (no a priori selection)	2	
	with Ly α emission	2	
	weak Ly α emission or absorption	0	

Notes.

^a For a detailed description of the selection criteria, we refer the reader to Mallery et al. (2012) and Hasinger et al. (2018).

^b Six of these galaxies are also observed as part of the Keck/DEIMOS survey (ref. 1). The corresponding number per selection from the Keck/DEIMOS program is given in square-brackets for those six galaxies.

^c Ly α emitters selected with *NB711*.

^d Ly α emitters selected with *NB814*.

^e Color-selected galaxies in B , g^+ , V , r^+ , and z^{++} using the criteria from Ouchi et al. (2004), Capak et al. (2004, 2011), Iwata et al. (2003), Hildebrandt et al. (2009).

^f Galaxies with a photometric redshift $z > 4$ with a probability of $> 50\%$ based on the Ilbert et al. (2010) photo- z catalog.

References: (1) Capak et al. (2004), Mallery et al. (2012), Hasinger et al. (2018), (2) Le Fèvre et al. (2015), (3) Vanzella et al. (2007, 2008), Balestra et al. (2010), (4) Malhotra et al. (2005), Rhoads et al. (2009).

selected galaxies are among the brightest. The *Chandra X-Ray Observatory*–detected galaxy *DC_845652* (green star) at $z = 5.3$ outshines all of the galaxies in UV luminosity.

All in all, although naturally biased to the brightest galaxies, this diverse selection function makes ALPINE an exemplary panchromatic survey that enables the study of a representative high- z galaxy sample at UV, optical, and FIR wavelengths.

2.2. Uniform Calibration of Spectra

All of the rest-frame UV spectra discussed in Section 2.1 are relative-flux corrected to remove sensitivity variations across the spectrograph as well as to correct atmospheric absorption features. However, not all of the spectra have been absolute-flux calibrated, which is important to measure absolute quantities such as their Ly α emission. Hence, we recalibrate the spectra using the Galactic extinction-corrected total broad-, intermediate-, and narrowband

photometry of the ALPINE galaxies (see Section 3 for details on the photometry). It turns out that the absolute-flux calibrated spectra are in excellent agreement (within better than 5% in flux) with our measured photometry, and the recalibration is not necessary in these cases. As the spectra come from different surveys, we convert them to a common format during the recalibration procedure.

To perform the absolute-flux calibration, we convolve each of the spectra with the transmission functions of the various optical broad-, intermediate-, and narrowband filters that exist on the COSMOS and ECDFS fields, respectively. On average, we use 4–9 filters for galaxies at $z < 5$ and 2–4 at $z > 5$. If the filter extends further than the spectrum, we extrapolate the spectrum by its medium continuum value. If the filter extends significantly beyond the spectrum ($> 50\%$), we do not consider the filter. We then compare the photometry obtained from the spectra to the total and Galactic extinction-corrected photometry discussed in Section 3, which allows us to obtain an average correction factor for each spectrum. We found that a single number for this correction per galaxy is enough for the calibration as the spectra already have been relative-flux calibrated. Since the uncalibrated spectra are mostly in units of counts, this correction is on the order of 10^{-21} for most galaxies. Our recalibration corrects for slit-losses and seeing variations. We also scale the variance in order to conserve the signal-to-noise ratio (S/N) of the spectrum. The final precision of our calibration is around 5%–10% in flux, which corresponds to the 1σ uncertainty in the photometry. Note that we do not consider undetected spectroscopic fluxes in this procedure; however, we use the constraints gained from the upper limits in the photometry for the calibration. Figure 5 shows three absolute-calibrated spectra at $z \sim 4.53$, $z \sim 4.56$ (with weak Ly α), and $z \sim 5.68$ to visualize our method. The filters that were used for the calibration are indicated in colors.

2.3. Stacked Spectra: Overview over Rest-frame UV Emission and Absorption Lines in ALPINE Galaxies

Figures 6 and 7 show stacks of different spectra. In order to create median-stacks of the spectra, we resample the spectra to a common wavelength grid and stack them in rest frame using their respective redshifts derived from [C II] or, if not available, from rest-UV absorption lines or Ly α emission. All stacks are subsequently binned to a resolution of 2 Å for visual purposes to emphasize the UV absorption features. To obtain a per-pixel uncertainty from the sky background for each stack (visualize by the gray line), we simply combine the inverse variances in the individual spectra in quadrature. The latter are the original inverse variance that we adjusted to the new normalization described in Section 2.2.

Panels 1a and 1b of Figure 6 compare the full stacked spectra of galaxies at $z < 5$ in COSMOS from observations with DEIMOS and as part of VUDS, as well as at $z > 5$ from observations in COSMOS from DEIMOS and in ECDFS from VIMOS and FORS2. In the former case, we adjust the resolution of the DEIMOS spectra ($R \sim 2500$) to that of the VUDS observations ($R \sim 230$) by applying a 1D Gaussian smoothing. The spectra are normalized to the median flux in the rest-frame wavelength range between 1300 and 1400 Å before stacking. For stacks of galaxies at $z > 5$, the rest-frame wavelength reaches up to 1500 Å, while for lower-redshift stacks, we show wavelengths up to rest frame 1700 Å. Several prominent spectral features are visible in the stacks in both redshift bins (indicated

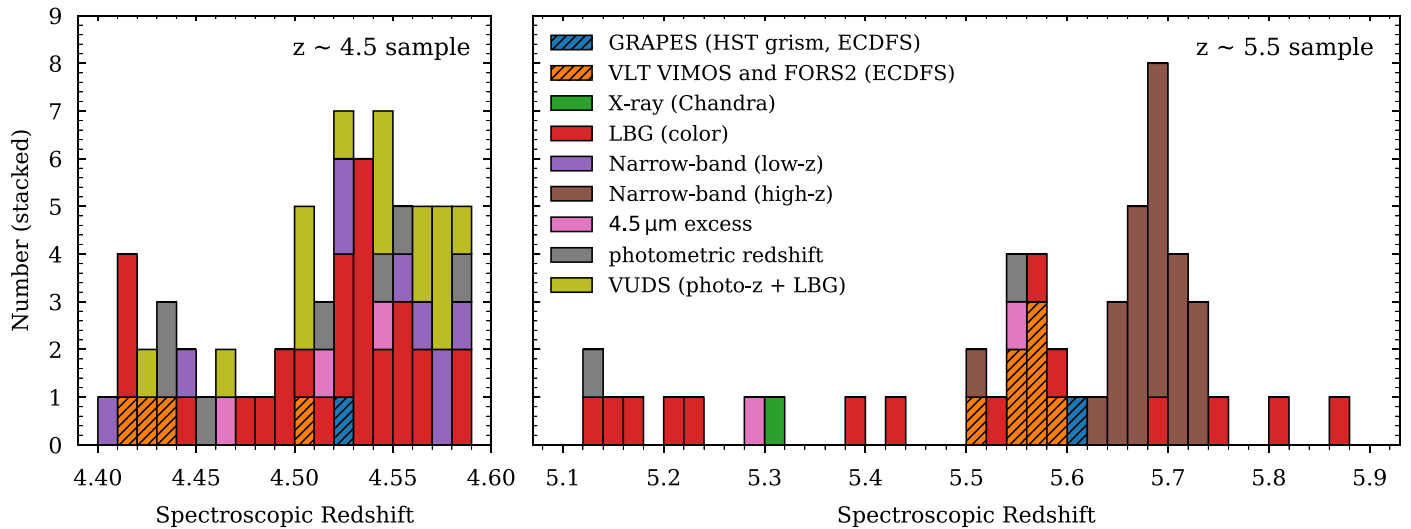


Figure 3. Redshift distribution of ALPINE galaxies. Each bar shows the stacked number of different selections per bin (see Table 1 and description in text). The bins with galaxies from the ECDFS field are hatched. The left and right panels show galaxies in the two different redshift bins.

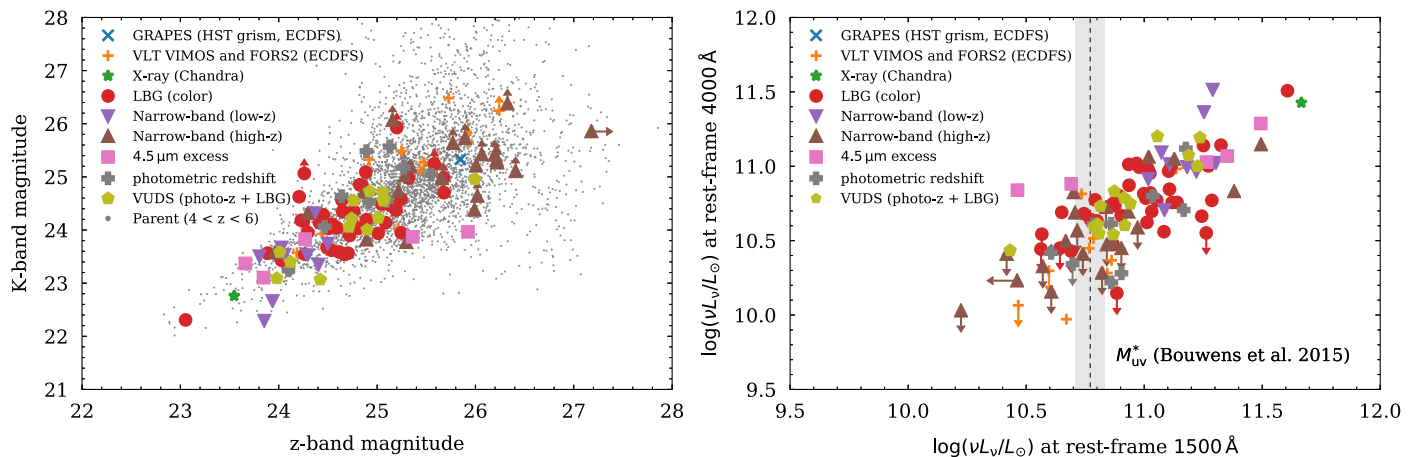


Figure 4. Comparison of observed (i.e., not corrected for dust) z -band and K -band magnitudes (left panel) and luminosities (right panel) for different selections listed in Table 1. The measurements on the parent sample in COSMOS at $4 < z < 6$ is shown in light gray. The color-coding is the same as in Figure 3. The arrows show 1σ upper limits. The gray area denotes the M_{UV}^* , the knee of the UV luminosity function, which corresponds to -21.1 ± 0.15 (or $\log(\nu L_\nu/L_\odot) \sim 10.77$) at $z = 5$ (Bouwens et al. 2015). The derivation of the photometry is described in detail in Section 3.

by gray bars). These include the $\text{Ly}\alpha$ emission line and N V at 1241 \AA and in addition UV absorption lines such as Si II at 1260 \AA , the Si III-[O I]-Si II complex at 1301 \AA , the two Si IV lines at 1398 \AA , as well as Si II, C IV, and He II at 1527 \AA , 1548 \AA , and 1640 \AA , respectively. Furthermore, we see indication of Fe II absorption at 1608 \AA in the COSMOS/DEIMOS spectra stack at $z < 5$. The depth of the UV absorption features are comparable for the different observations with the different instruments, verifying similar quality and little biases. However, note that the features in the ECDFS spectra are less pronounced due to the factor ~ 6 smaller number of spectra contributing to the stacks compared to the DEIMOS stacks. Panels 2a through 2e show the stacks for variously selected data sets below and above $z = 5$. The spectra are not normalized before stacking in these cases to provide a comparison of the absolute-flux values for the different redshifts and samples to the reader. The number of spectra per wavelength are shown on the top right for each panel. Note again that the number of high-redshift spectra drops toward redder wavelengths. This has to be

kept in mind when analyzing the spectral features in the stacks. Emission and absorption lines are indicated as in the other panels. As expected, the stacked spectra at higher redshifts are fainter, but still significant UV absorption features are present (see also Faisst et al. 2016b; Khusanova et al. 2019; Pahl et al. 2020).

Figure 7 shows stacked spectra in COSMOS observed with DEIMOS for the different selection categories (see also Table 1 and Figure 3). We split the LBG category in galaxies below and above $z = 5$. All of the spectra are smoothed with a Savitzky–Golay filter with size of 2 \AA for visualization purposes. The total number of spectra per stack is indicated in the upper left corner. All panels are scaled the same way to emphasize differences in brightness. The X-ray detected galaxy at $z = 5.3$ is UV bright compared to the other stacks and shows strong N V emission with overlaid Si II absorption as well as broad C IV emission. LBGs (i.e., color-selected galaxies) are preferentially fainter but of similar continuum brightness as narrowband-selected galaxies at $z \sim 4.5$. The latter show

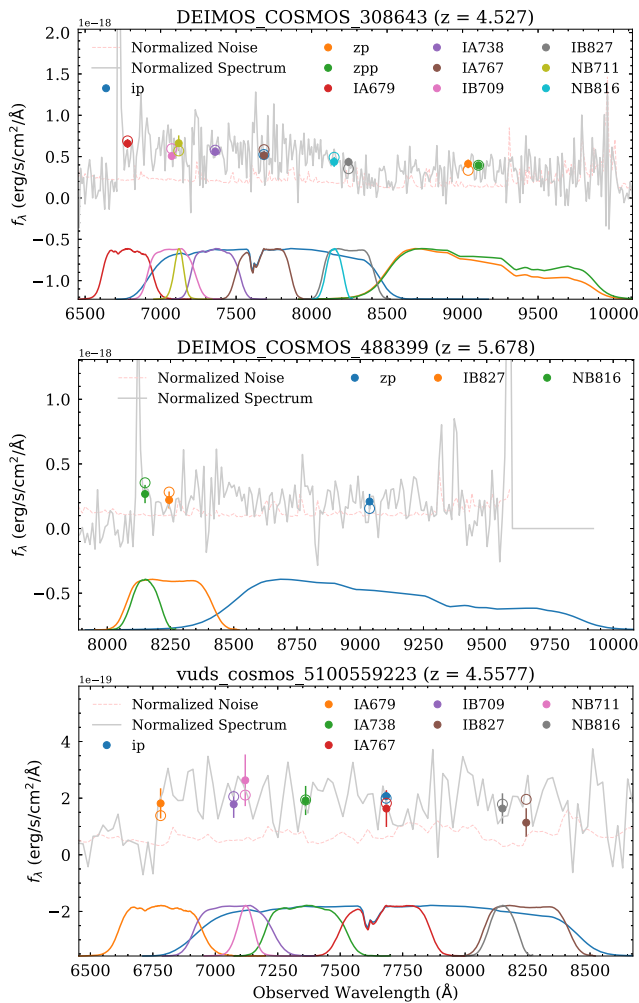


Figure 5. Absolute calibration of rest-frame UV spectra. Shown are three examples at $z = 4.53$, $z = 4.56$, and $z = 5.68$. The spectra are convolved by the filters, and the photometry (open circles) is compared to the total and Galactic extinction-corrected broad-, intermediate-, and narrowband photometry from catalogs (filled circles) described in Section 3.

significant C II, Si IV, and C IV absorption. As expected, narrowband-selected galaxies at $z \sim 5.7$ show strong Ly α emission and a faint continuum such that the S/N is too low to detect UV absorption features at great significance. The stack of galaxies selected by photometric redshifts shows to first-order similar properties as the LBGs. The $4.5 \mu\text{m}$ -excess continuum selected galaxies are, on average, the continuum brightest galaxies and show significant Ly α emission as well as absorption features.

2.4. Rest-UV Emission and Absorption Lines and Velocity Offsets

2.4.1. Measurements

We measure basic quantities from the individual rest-frame UV spectra. These include the redshift and equivalent width of Ly α emission as well as redshifts from various absorption lines.

The Ly α redshift ($z_{\text{Ly}\alpha}$) is based on the peak of the (asymmetric) Ly α emission to allow a direct comparison with models of Ly α radiative transfer (see, e.g., Hashimoto et al. 2015). The Ly α flux is measured by fitting a Gaussian to the

line and for measuring the equivalent width ($\sim f_{\text{line}}^{\text{tot}}/f_{\text{continuum}}$) the continuum redward of the Ly α line is used. These measurements are explained in more detail in Cassata et al. (2020).

The absorption redshifts are measured for each individual spectrum, if possible, using the lines Si II (1260.4 Å), [O I] (1302.2 Å),³⁰ C II (1334.5 Å), Si IV (1393.8 Å) and Si IV (1402.8 Å), Si II (1526.7 Å), and C IV (1549.5 Å).³¹ The first four are covered by observations in all galaxies, while the coverage of the latter depends on the redshift of the galaxy. Note that some of the above lines are predominantly formed in the ISM (low-ionization interstellar [IS] lines; Si II, [O I], C II, Si II), while others are formed in stellar winds (high-ionization wind lines; Si IV or C IV) and therefore can display strong velocity shifts (e.g., Castor & Lamers 1979; Leitherer et al. 2011). To increase the S/N of our measurements, we use all of the above lines to derive an absorption line redshift (referred to as $z_{\text{IS}+\text{wind}}$), but we compare the individual redshift from the IS (z_{IS}) and wind (z_{wind}) lines to investigate potential systematic differences. Before performing any measurements, we subtract a continuum model from each individual spectrum. The model is derived by fitting a fourth-order polynomial to the spectrum, which is smoothed by a 5 Å box kernel. We then fit the above absorption lines in five different rest-frame wavelength windows [1240 Å, 1280 Å], [1280 Å, 1320 Å], [1320 Å, 1350 Å], [1370 Å, 1420 Å], and [1500 Å, 1570 Å]. For the separate fit of the IS and wind lines, we split the last window into two ranges, namely [1510 Å, 1540 Å] and [1530 Å, 1570 Å] to separate the IS line Si II and the wind line C IV, respectively. The absorption lines can be significantly asymmetric due to stellar winds and the effect of optical depth. Fitting a single Gaussian to them could therefore bias the redshift measurements. Instead, we use the stacked spectrum of LBGs at $z \sim 3$ from Shapley et al. (2006) as a template, which we cross-correlate to the observed data within the wavelength range of a given window by χ^2 minimization. We let the redshift vary within a velocity range of $\pm 600 \text{ km s}^{-1}$ (corresponding to roughly 0.01 in redshift) around a prior absorption redshift, which is obtained by a manual cross-correlation of the same template to all possible absorption lines at once using the interacting redshift-fitting tool SpecPro³² (Masters & Capak 2011). We found that this approach significantly removes degeneracies in the fit and at the same time allows a visual inspection of all of the spectra to flag the ones with low S/N where no reasonable fit can be obtained.³³ For each galaxy, the so obtained $\chi^2(z)$ distribution is then converted into a probability density function $p(z)$ for each of the windows. These are combined, by choosing the necessary absorption lines, to a total probability $P(z)$ from which the final absorption line redshifts ($z_{\text{IS}+\text{wind}}$, z_{wind} , or z_{IS}) are derived. The errors on these redshifts are derived by repeating this measurement 200 times, thereby perturbing the fluxes according to a Gaussian error distribution with σ defined by the average flux noise of the continuum. Typical uncertainties are on the order of $\pm 100 \text{ km s}^{-1}$.

³⁰ Here, we refer to the [O I] absorption line complex consisting of Si III, C II, [O I], and Si II.

³¹ This absorption consists of two lines (1548.2 Å and 1550.8 Å) and here we give the average wavelength.

³² <http://specpro.caltech.edu>

³³ The value of this visual flag is -99 if the S/N is too low to obtain a redshift, and 1 and 2 for reliable and very reliable redshift measurements, respectively.

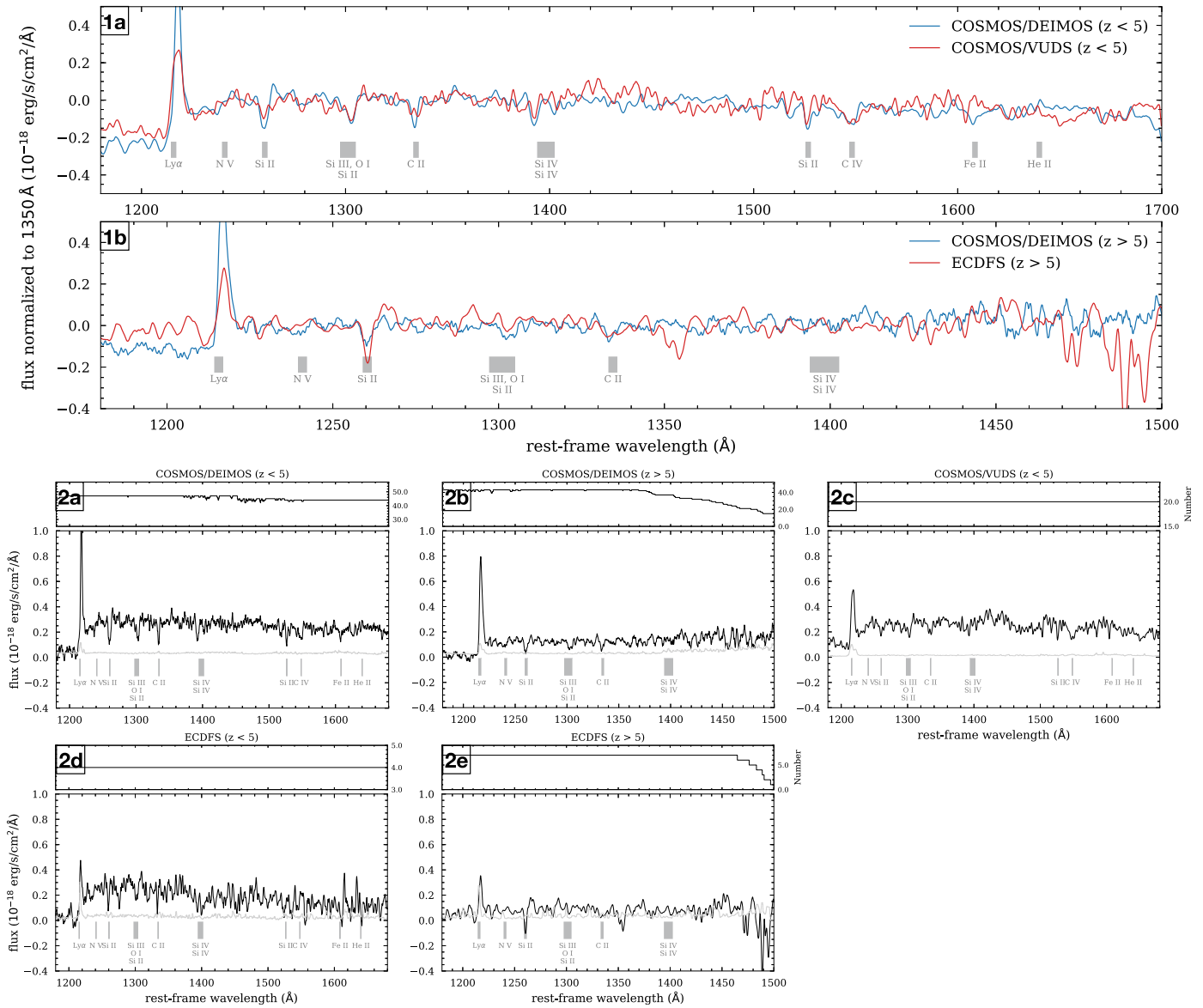


Figure 6. Examples of stacked ALPINE spectra. Panels 1a and 1b show stacked spectra at $z < 5$ (in COSMOS from DEIMOS observations and as part of the VUDS survey) and $z > 5$ (on COSMOS from DEIMOS and on ECDFS from VIMOS and FORS2 observations), respectively. The stacks are all normalized to the continuum between 1300 Å and 1400 Å and common emission and absorption features are indicated with gray bars. Note that the VUDS spectra have a lower native resolution ($R \sim 230$) compared to the DEIMOS observations ($R \sim 2500$); therefore, the latter have been degraded in resolution using a 1D Gaussian window function for visual comparison. Panels 2a through 2e show stacks at $z < 5$ and $z > 5$ for the different data sets. The number of spectra included per wavelength is shown on the top of each panel. The uncertainty in flux is indicated by the light gray line. The y-axis scale is the same such that the continuum brightness can be compared.

As mentioned in Section 2.1, six galaxies in COSMOS have been observed by the Keck/DEIMOS and VUDS spectroscopic surveys. Therefore there are two measurements for each of these galaxies. Specifically, for $vc_5100559223$, $vc_5100822662$, $vc_5101218326$, and $vc_5101244930$, the IS+wind redshift measurements agree within 200, 280, 70, and 110 km s⁻¹. These values are on the order of the measurement uncertainties. Note that while the VUDS slits are oriented east–west, the DEIMOS slits can be oriented north–south or in any other angle. This different orientation could also be responsible for the differences in velocity offsets. On the other hand, for $vc_5101288969$ and $vc_510786441$, we find significant differences of 1290 and 1010 km s⁻¹. A close inspection of the spectra shows that these are very low in S/N. Also, both have low visual quality flags (−99 and 1, indicating not robust measurements are possible),

and their redshifts are fit with less than three lines; thus, they should not be trusted. For all six spectra, we decided to prefer the VUDS observations because of their slightly better S/N at a cost of lower resolution.

2.4.2. Velocity Offsets with Respect to [C II] FIR Redshifts

The detection of [C II] by ALMA provides the systemic redshift of a galaxy. This enables us to study velocity offsets of rest-frame UV absorption lines and Ly α emission that will inform further about the properties of the ISM in these galaxies similarly to studies at lower redshifts using H α and C II] λ 1909 (e.g., Steidel et al. 2010; Marchi et al. 2019). Here, we give an overview of the velocity properties and compare them for galaxies with high and low specific SFRs.

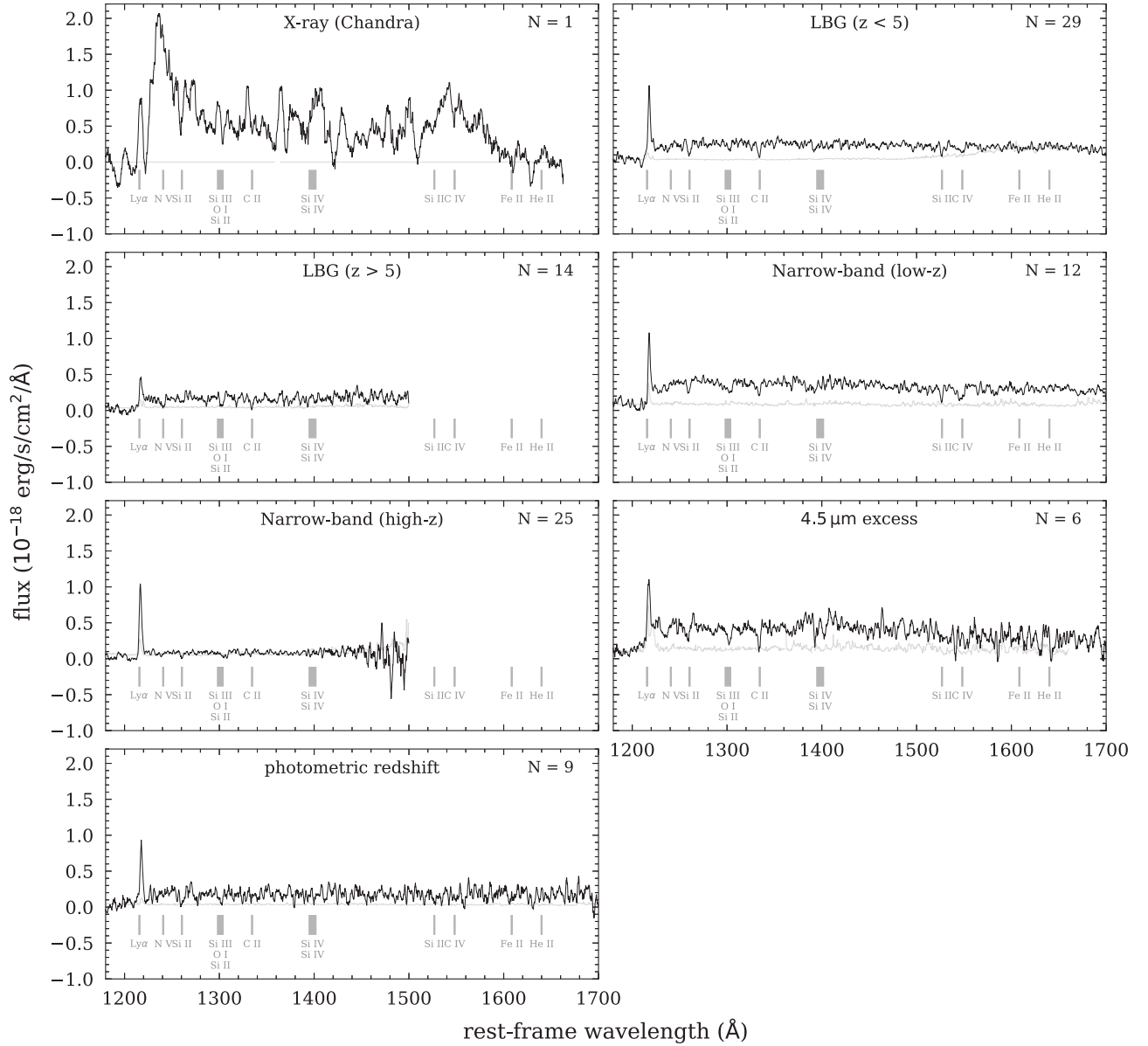


Figure 7. Stacked spectra in COSMOS for each of the selections discussed in Section 2.1 and listed in Table 1. Emission and absorption features are indicated by gray bars, and the number of spectra in the stack is shown on the upper right corner. We also show the X-ray-detected galaxy (*DC_845652*) at $z = 5.3$, which shows strong and broad N V and C IV emission. The uncertainty in flux is indicated by the light gray line.

In the following, we define the velocity difference for two redshifts (z_1 and z_2) as $\langle z_1 - z_2 \rangle \equiv \Delta v_{12} = c \times \left(\frac{1+z_1}{1+z_2} - 1 \right)$ where $c = 2.998 \times 10^5 \text{ km s}^{-1}$. The measurement of the [C II] redshifts are detailed in Bethermin et al. (2020). They are defined as the peak of a Gaussian fit to the [C II] line with spectral resolution of 25 km s^{-1} . The uncertainty of the redshift measurements was estimated by a Monte Carlo simulations with perturbed fluxes according to the error per spectral bin. The average uncertainty is roughly $50\text{--}60 \text{ km s}^{-1}$. For the absorption lines, we require that $z_{\text{IS+wind}}$ is measured from at least three absorption lines, and we only show galaxies that have not been flagged by our visual inspection with SpecPro as unreliable (flag -99). The average intrinsic measurement error per galaxy is $\pm 100 \text{ km s}^{-1}$. In relation to that, a systematic uncertainty of 0.5 \AA in the rest-frame wavelength of the absorption lines (e.g., due to calibration issues) turns into a velocity shift of $\sim 120 \text{ km s}^{-1}$.

Figure 8 shows stacked histograms of velocity differences. The number of galaxies used as well as the median of the distribution with scatter (*not* error on the median) are indicated as well. The upper panel compares the velocities measured from Ly α and the IS+wind absorption lines. We find a median offset on the order of $386_{-279}^{+257} \text{ km s}^{-1}$, which is consistent with other measurements at the same redshifts (see, e.g., Faisst et al. 2016b; Pahl et al. 2020) as well as at $z \sim 2\text{--}3$ (Steidel et al. 2010). The center and bottom panels compare the IS+wind and Ly α redshifts to the systemic redshift (here defined as the [C II] $\lambda 158 \mu\text{m}$ redshift, Bethermin et al. 2020). For the former, we find an offset of $-227_{-206}^{+168} \text{ km s}^{-1}$, and for the latter, we find $184_{-215}^{+201} \text{ km s}^{-1}$. These negative and positive velocity offsets can be related in a simple physical model involving the resonant scattering of Ly α photons and outflowing gas in the outskirts of galaxies (see detailed discussion in Steidel et al. 2010). The redshifted Ly α emission line (with respect to

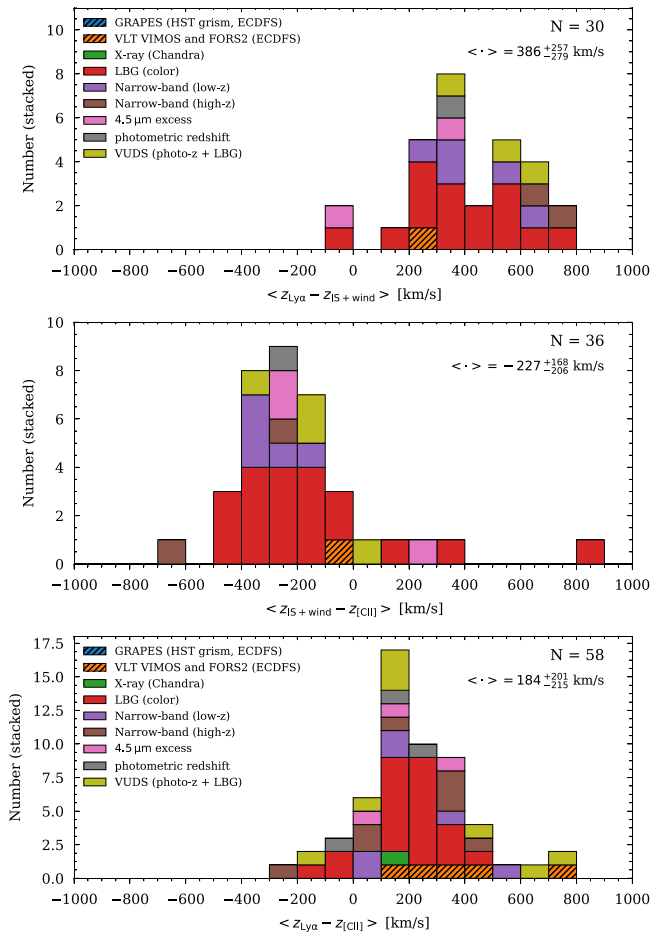


Figure 8. Stacked histograms of velocity offsets between redshifts derived from different spectral features. The number of galaxies and median of the distribution (including scatter) are indicated. Shown are the velocity offsets between $\text{Ly}\alpha$ emission and IS+wind absorption lines (*top panel*), as well as between $\text{Ly}\alpha$, IS+wind, and systemic redshift (*middle and bottom panel*). The latter two are discussed in detail in a forthcoming paper (Cassata et al. 2020). The average errors are on the order of $\pm 100 \text{ km s}^{-1}$, which corresponds to the size of the bins. We do not find any significant biases introduced by the different selection methods (color-coded as in previous figures).

systemic) can be explained by resonant scattering of the $\text{Ly}\alpha$ photons. Preferentially, redshifted $\text{Ly}\alpha$ photons scattered from the back of the galaxy can make it unscattered through the intervening gas inside the galaxy along the line of sight. The blueshift of IS absorption may depend on the outflow velocity of the absorbing gas as well as its covering fraction (or optical depth) inside the galaxy along the line of sight toward the observer. For a more in-depth discussion, we refer the reader to a companion paper by Cassata et al. (2020). Overall, we do not see a significant dependence of the velocity differences on the various selection techniques (color-coded in the figure).

Figure 9 compares the velocity offsets between IS (Si II, [O I], C II, Si II) and wind (Si IV, C IV) lines. We require that at least three IS lines and one wind line is measured. In addition, only galaxies that pass our visual classification (i.e., have flags other than -99 , see above) are used. Overall, we do not see any statistical difference between IS and wind lines, although there is a tail toward higher blueshifts in the case of wind lines. However, wind and outflows may be increased in galaxies with high and spatially dense star formation and young stellar populations. Therefore, we would expect different velocity shifts for the absorption lines with respect to the

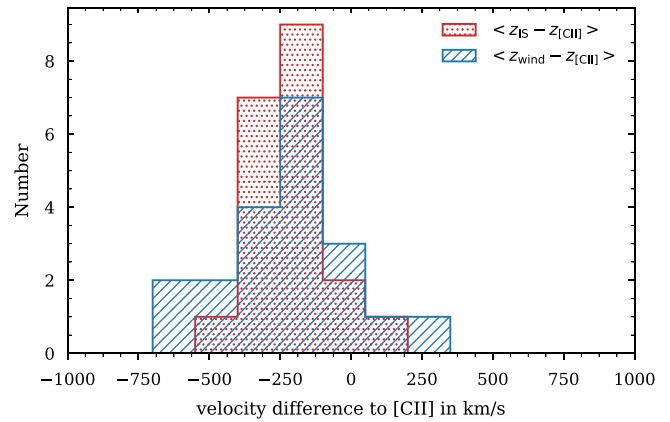


Figure 9. Histogram of velocity offset with respect to systemic (defined by the [C II] $\lambda 158 \mu\text{m}$ redshift) for IS (red; [O I], C II, and Si II) and stellar wind affected absorption lines (blue; Si IV and C IV). The average errors are on the order of $\sim 200 \text{ km s}^{-1}$, which corresponds to the size of the bins.

systemic redshift for highly star-forming galaxies. In Figure 10, we investigate this picture by stacking galaxies at the extreme ends of the sSFR distribution (we refer to Section 4 for details on the measurement of the physical properties of our galaxies), namely low ($< 4 \text{ Gyr}^{-1}$) and high ($> 5 \text{ Gyr}^{-1}$) sSFR, in their corresponding rest-frames defined by the systemic redshift (i.e., [C II] $\lambda 158 \mu\text{m}$ redshift). The sSFR is a good proxy of the star formation density in a galaxy as well as the age of the current stellar population (see, e.g., Cowie et al. 2011). We show the stacked spectra in five wavelength regions covering prominent absorption lines for each sSFR bin. The vertical dashed lines show the different absorption lines in the [C II] rest frame. First, we verify that the shifts between IS and wind lines are very similar for each sSFR bin (in concordance with Figure 8). However, it is intriguing that in the low sSFR stack, all absorption lines agree well with the [C II] redshift, while in the high sSFR stack, the lines are significantly blueshifted by $300\text{--}400 \text{ km s}^{-1}$. We also note that in the high sSFR stack, the C IV line shows a noticeable P-Cygni profile indicative of strong winds and outflows (Castor & Lamers 1979). These findings fit well into a picture of strong winds and outflows produced by the high star formation in these galaxies, which is also in line with recent results obtained through the stacking of ALPINE [C II] spectra (Ginolfi et al. 2020).

2.4.3. How Well Does [C II] Trace Systemic Redshift? Comparison to Optical [O II] Emission

The extended nature of [C II] may be indicative of its origin in the diffuse interstellar medium in addition to PDRs (Stacey et al. 1991; Gullberg et al. 2015; Vallini et al. 2015; Faisst et al. 2017). Moreover, recent work by Ginolfi et al. (2020) shows that [C II] emission is significantly affected by large-scale outflows caused by a high rate of star formation in these galaxies. However, as shown by the same study, the outflows seem to be symmetric, and therefore, we do not expect them to significantly change the centroid of the [C II] emission line.

During 2019 January 13–15, we were able to obtain a near-IR spectrum of one of our ALPINE galaxies (*DC_881725* at $z_{[\text{CII}]} = 4.5777$) using the Multi-Object Spectrometer For Infra-Red Exploration (MOSFIRE; McLean et al. 2010, 2012) at the 10 m Keck I telescope on Maunakea in Hawaii. The observations of a total on-source integration time of 24×3 minutes in *K* band ($1.92\text{--}2.40 \mu\text{m}$) were carried out

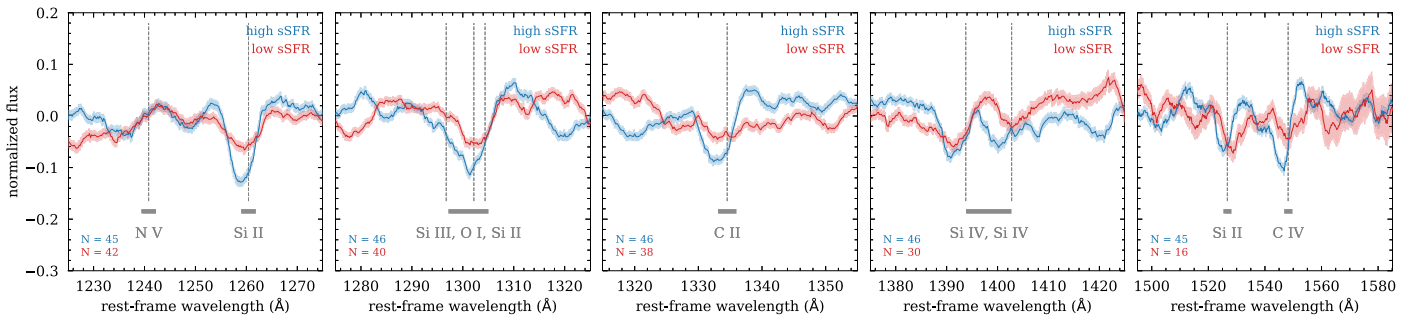


Figure 10. Stacked spectra (in C II systemic redshift) in two bins of sSFR (red: $< 4 \text{ Gyr}^{-1}$, blue: $> 5 \text{ Gyr}^{-1}$) for five wavelength regions covering prominent rest-UV absorption lines. The derivation of the sSFR for the ALPINE galaxies is detailed in Section 4.1. The average number of spectra in each bin is indicated together with the prominent absorption and emission lines. We note systematically stronger blueshifts of all absorption lines for the high sSFR stack. Particularly, note the strong blueshift of the high-ionization wind lines. The C IV lines in the high sSFR bin also show indication of a more pronounced P-Cygni profile, indicative of strong stellar winds and outflows in high sSFR galaxies. The 1σ uncertainties of the stacked spectra is indicated by the shaded regions.

under clear weather conditions with an excellent average seeing FWHM of $0''.3\text{--}0''.4$. We performed a standard data reduction using the MOSFIRE data reduction pipeline³⁴ (Version 2018). From the produced 2D spectrum and variance map, we extract the 1D spectrum at the spatial location of the galaxy using a weighted mean across ± 3.5 spatial pixels ($0''.18 \text{ px}^{-1}$).

We are able to detect the optical [O III] doublet (3727.09 \AA and 3729.88 \AA) at the spatial position of the galaxy at a level of $> 5\sigma$. Note that this is the first detection of optical [O III] in a galaxy with [C II] measurement from ALMA, which allows us, for the first time, to compare these two lines at these redshifts. In the bottom panel of Figure 11, we show the final spectrum in the rest frame of the [C II] emission. The width of $\sim 4 \text{ \AA}$ includes both [O III] lines. The theoretical rest-frame wavelength of the doublet is indicated by the black arrows. The position of the line agrees perfectly with the [C II] redshift derived from ALMA, indicating that FIR [C II] and optical [O III] trace the same systemic redshift. In addition, the top panel of the Figure shows the [C II] FIR line for comparison of the central wavelength of the lines. Note that the actual width of the [C II] line is more than 100 times larger than the the one of the optical [O III] line.

3. Photometry from Ground and Space

In this section, we summarize the ground- and space-based photometric data that are available for the ALPINE galaxies in the COSMOS (105 galaxies) and ECDFS (13 galaxies) fields. Although these fields differ in survey depth, reduction methods, and number and type of photometric filters used, we find that their overall photometric measurements are comparable within 1σ limits after their conversion to total magnitudes and the correction for the specific biases of each survey. Therefore, we can treat them separately to the first order for the matter of measuring various physical properties of the galaxies. The basis catalogs to which we match the ALPINE galaxies in the COSMOS and ECDFS field are the COSMOS2015³⁵ (Laigle et al. 2016) and the 3D-*HST*³⁶ (Brammer et al. 2012; Skelton et al. 2014) catalog, respectively. A summary of the different data available on the two fields including filter names, wavelengths, 3σ depths, and references to the measurements are given in Tables 2 and 3, respectively. In the following, we describe these data in more detail.

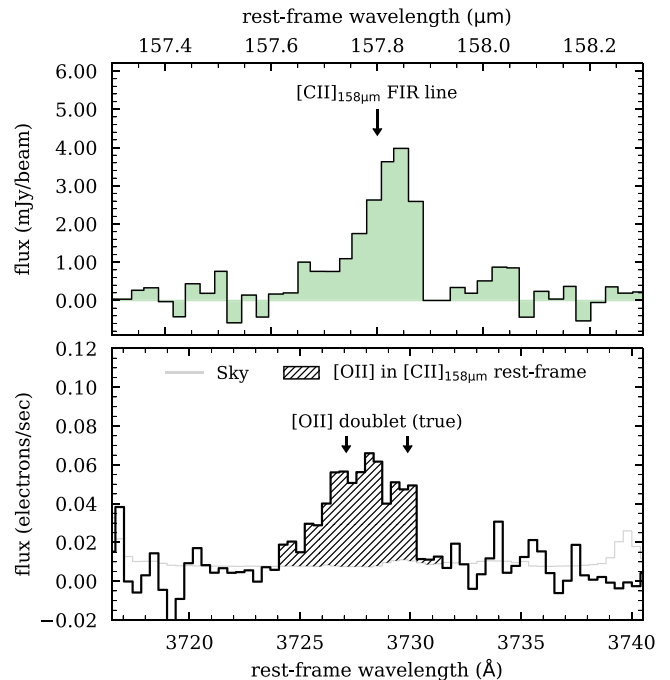


Figure 11. Comparison of optical and FIR emission lines in the case of galaxy *DC_881725*. Top panel: [C II] emission line at $158 \mu\text{m}$ observed with ALMA. Bottom panel: the black line shows the optical [O III] doublet at 3727 \AA observed with the MOSFIRE spectrograph on Keck in the rest-frame of the [C II] emission. The theoretical rest-frame wavelength of the doublet is indicated by the two arrows. This shows that there is no significant velocity offset between the [C II] FIR and optical [O III] emission; hence, this verifies the validity of [C II] as tracer of the systemic redshift of this galaxy.

3.1. Photometry on the ECDFS Field

The photometry for the galaxies on the ECDFS field is taken directly from the 3D-*HST* catalog, which provides ground-based observations as well as a wealth of data from *HST* imaging. The photometry (total fluxes and magnitudes) is corrected for Galactic extinction, PSF size as well as other biases; therefore, no further corrections are applied. The ALPINE galaxies are matched visually to the spatially closest 3D-*HST* counterpart using the *HST* WFC3/IR *F160W* image as reference. The spectroscopic redshifts match the photometric redshifts within their uncertainty ($\sim 0.1\text{--}0.2$), ensuring that we identified the correct counterpart.

The ground-based photometry available in ECDFS (including references) is listed in Table 2. Summarizing, this includes

³⁴ <https://keck-datareductionpipelines.github.io/MosfireDRP/>

³⁵ <http://cosmos.astro.caltech.edu/page/photom>

³⁶ <https://3dhst.research.yale.edu/Data.php>

Table 2
Photometry Available for Galaxies on the ECDFS Field

Observatory/Instrument	Filter	Central λ (Å)	3σ depth (mag)	Ref.
Ground-based				
MPG-ESO/WFI	<i>U</i> 38	3633.3	27.3	1
	<i>b</i>	4571.2	26.6	1
	<i>v</i>	5377.0	26.6	1
	<i>R_c</i>	6536.3	26.9	1
VLT/VIMOS	<i>I</i>	9920.2	26.5	1
	<i>U</i>	3720.5	28.6	2
VLT/ISAAC	<i>R</i>	6449.7	27.8	2
	<i>J^v</i>	12492.2	25.6	3
CFHT/WIRCam	<i>H^v</i>	16519.9	25.1	3
	<i>K_s^v</i>	21638.3	25.0	3
	<i>J_w</i>	12544.6	25.1	4
Subaru/Suprime-Cam	<i>K_s^w</i>	21590.4	24.5	4
	<i>IA</i> 427	4263.4	25.7	5
Subaru/Suprime-Cam	<i>IA</i> 445	4456.0	25.7	5
	<i>IA</i> 505	5062.5	25.8	5
	<i>IA</i> 527	5261.1	26.7	5
	<i>IA</i> 550	5512.0	26.0	5
	<i>IA</i> 574	5764.8	25.7	5
	<i>IA</i> 598	6000.0	26.6	5
	<i>IA</i> 624	6233.1	26.5	5
	<i>IA</i> 651	6502.0	26.7	5
	<i>IA</i> 679	6781.1	26.6	5
	<i>IA</i> 738	7371.0	26.5	5
	<i>IA</i> 767	7684.9	25.5	5
	<i>IA</i> 797	7981.0	25.2	5
	<i>IA</i> 856	8566.0	25.0	5
	Space-based			
<i>HST</i> /ACS	<i>F</i> 435W	4328.7	29.1	6
	<i>F</i> 606W	5924.8	29.1 (29.0)	6,7
	<i>F</i> 775W	7704.8	28.4	6
	<i>F</i> 814W	8058.2	28.9	7
<i>HST</i> /WFC3	<i>F</i> 850LP	9181.2	27.9	6
	<i>F</i> 125W	12516.3	27.6	8
	<i>F</i> 140W	13969.4	26.7	9
<i>Spitzer</i> /IRAC	<i>F</i> 160W	15391.1	27.7	8
	ch ₁	35634.3	~26.0 ^a	10
<i>Spitzer</i> /IRAC	ch ₂	45110.1	~26.0 ^a	10
	ch ₃	57593.4	24.4	11
	ch ₄	79594.9	24.3	11

Note.

^a The exposure time varies between 10 and 300 ks (corresponding to roughly 1.8 in magnitudes).

References: (1) Hildebrandt et al. (2006); Erben et al. (2005), (2) Nonino et al. (2009), (3) Wuyts et al. (2008); Retzlaff et al. (2010), (4) Hsieh et al. (2012), (5) Cardamone et al. (2010), (6) Giavalisco et al. (2004), (7) Koekemoer et al. (2011), (8) Grogin et al. (2011); Koekemoer et al. (2011), (9) Brammer et al. (2012); van Dokkum et al. (2013), (10) Ashby et al. (2013); Guo et al. (2013), (11) Dickinson et al. (2003).

the *U*38, *b*, *v*, *R_c*, and *I* broadband filters from the Wide Field Imager on the 2.2 meter MPG/ESO telescope, the *U* and *R* bands from VIMOS on the VLT, the near-IR filters *J^v*, *H^v*, and *K_s^v* from ISAAC on the VLT, *J_w* and *K_s^w* data taken by WIRCam on the CFHT, as well as 14 intermediate-band filter from the Suprime-Cam on the Subaru telescope. For galaxies at $z = 4.5$ and 5.5 , the Lyman-break falls roughly in the *v* and the *R_c*-band, and therefore, the galaxies are expected to be only faintly (or not at all) visible in these and blueward filters. On the other hand, the galaxies are bright at observed near-IR

Table 3
Photometry Available on the COSMOS Field

Observatory/Instrument	Filter	Central λ (Å)	3σ depth (mag) ^a	Ref.
Ground-based				
CFHT/MegaCam	<i>u</i> [*]	3823.3	26.6	1
Subaru/Suprime-Cam	<i>B</i>	4458.3	27.0	1, 2
	<i>V</i>	5477.8	26.2	1,2
	<i>r</i> ⁺	6288.7	26.5	1, 2
	<i>i</i> ⁺	7683.9	26.2	1, 2
	<i>z</i> ⁺⁺	9105.7	25.9	1, 2
	<i>IA</i> 427	4263.4	25.9	1, 2
	<i>IA</i> 464	4635.1	25.9	1, 2
	<i>IA</i> 484	4849.2	25.9	1, 2
	<i>IA</i> 505	5062.5	25.7	1, 2
	<i>IA</i> 527	5261.1	26.1	1, 2
Subaru/HSC	<i>IA</i> 574	5764.8	25.5	1, 2
	<i>IA</i> 624	6233.1	25.9	1, 2
	<i>IA</i> 679	6781.1	25.4	1, 2
	<i>IA</i> 709	7073.6	25.7	1, 2
	<i>IA</i> 738	7361.6	25.6	1, 2
	<i>IA</i> 767	7684.9	25.3	1, 2
	<i>IA</i> 827	8244.5	25.2	1, 2
	<i>NB</i> 711	7119.9	25.1	1, 2
	<i>NB</i> 816	8149.4	25.2	1, 2
	<i>Y_{HSC}</i>	9791.4	24.4	1
CFHT/WIRCam	<i>H_w</i>	16311.4	23.5	1
VISTA/VIRCAM	<i>K_s^w</i>	21590.4	23.4	1
	<i>Y</i>	10214.2	24.8 (25.3) ^b	1
	<i>J</i>	12534.6	24.7 (24.9) ^b	1
	<i>H</i>	16453.4	24.3 (24.6) ^b	1
	<i>K_s</i>	21539.9	24.0 (24.7) ^b	1
Space-based				
<i>HST</i> /ACS	<i>F</i> 435W	4328.7	— ^c	—
	<i>F</i> 475W	4792.3	— ^c	—
	<i>F</i> 606W	5924.8	— ^c	—
	<i>F</i> 814W	8058.2	29.2	3
<i>HST</i> /WFC3	<i>F</i> 850LP	9181.2	— ^c	—
	<i>F</i> 098M	9877.4	— ^c	—
	<i>F</i> 105W	10584.9	— ^c	—
	<i>F</i> 110W	11623.8	— ^b	—
	<i>F</i> 125W ^d	12516.3	27.6 ^{c,d}	4
	<i>F</i> 140W	13969.4	— ^c	—
<i>Spitzer</i> /IRAC	<i>F</i> 160W ^d	15391.1	27.5(25.0) ^{c,d,e}	4, 5
	ch ₁	35634.3	25.5	1, 6, 7
	ch ₂	45110.1	25.5	1, 6, 7
	ch ₃	57593.4	23.0	1, 7
ch ₄	79594.9	22.9	1, 7	

Notes.

^a Depth is measured in $3''$ aperture for ground-based photometry.

^b Depths of the ultra-deep area are given in parenthesis.

^c Ancillary pointings; therefore, depth varies depending on the specific observations.

^d Most of these data come from the CANDELS survey and their depth is indicated.

^e Some of the *F*160W data comes from the DASH survey (depth given in parenthesis).

References: (1) see Laigle et al. (2016), (2) Taniguchi et al. (2007, 2015) (3) Scoville et al. (2007b), Koekemoer et al. (2007), (4) Grogin et al. (2011), Koekemoer et al. (2011), (5) Momcheva et al. (2017), (6) Capak et al. (2012), Steinhardt et al. (2014), (7) Sanders et al. (2007).

wavelengths, i.e., filters redward of the *z* band (corresponding to roughly the *F*850LP filter). Figure 12 shows the stacked *F*850LP (*z*), *J^v* (*J*), and *K_s^v* (*K*) magnitude distributions of the ECDFS ALPINE galaxies split in $z < 5$ (hatched blue) and

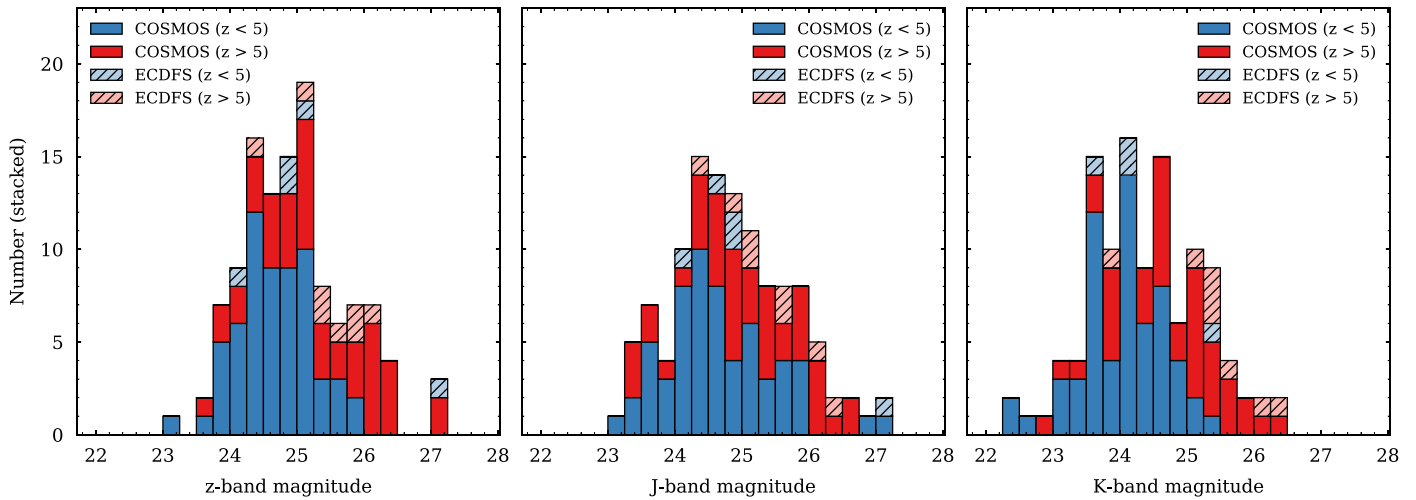


Figure 12. Stacked histograms of the magnitude distribution for the ALPINE galaxies in the COSMOS (solid) and ECDFS (hatched) fields. The blue and red color-coding indicates galaxies at $z < 5$ and $z > 5$, respectively. The magnitudes (from left to right) correspond to z^{++} , J , and K_s bands for COSMOS and $F850LP$, J^v , and K_s^v bands for ECDFS (see Tables 2 and 3 for more information on the filters).

$z > 5$ (hatched red). As expected, the latter sample occupies slightly fainter magnitudes.

The space-based photometry includes the four *Spitzer* bands at 3.6, 4.5, 5.8, and 8.0 μm . In addition, the public 3D-*HST* catalog includes a wealth of HST photometry. Specifically, it contains measurements in the ACS bands $F435W$, $F606W$, $F775W$, $F814W$, and $F850LP$ as well as in the WFC3/IR bands $F125W$, and $F160W$ bands for all 13 ALPINE galaxies. Only 10 galaxies have measurements in the WFC3/IR band $F140W$. The *HST* photometry is measured on PSF-matched images. As described in Skelton et al. (2014), the *Spitzer* and ground-based photometry are measured using the MOPHONGO (Labbé et al. 2006; Wuyts et al. 2007; Whitaker et al. 2011), which uses a high-resolution image (here the *HST* imaging) as spatial prior to estimate the contributions from neighboring blended sources in the lower resolution image. The different depths of these observations as well as references are listed in Table 2. A query of the Barbara A. Mikulski Archive for Space Telescopes (MAST³⁷) using the `mastquery` Python package³⁸ shows that in addition to the *HST* measurements contained in the 3D-*HST* catalog, four, ten, and two galaxies have coverage in the WFC3/IR bands $F098M$, $F105W$, and $F110W$, respectively. None of the galaxies has ACS $F475W$ coverage. These additional data that are not published in the 3D-*HST* catalog come from various other observation programs in and around the ECDFS field. We subsequently measure this additional photometry for all ALPINE galaxies in ECDFS using SExtractor (version 2.19.5, Bertin & Arnouts 1996) in different aperture sizes ($0''.7$ and $3''$) as well as auto magnitudes. For this, we first create a mosaic of all of the *HST* pointings that overlap with the ALPINE galaxies using the `AWS-drizzler`³⁹ tool that is part of the `grizli`⁴⁰ Python package (G. Brammer 2020, in preparation.). We use a $0''.06$ pixel scale and all *HST* images are registered to *Gaia* (see Section 3.3). SExtractor is run with relative THRESH_TYPE, and we set DETECT_MINAREA, DETECT_THRESH, DEBLEND_MINCONT, and DEBLEND_NTHRESH to 3, 1.5, 1.5, 0.001, and 64, respectively. If no

object is detected above the threshold (1.5σ) within $0''.7$ (roughly the ground-based seeing) of the original ALPINE coordinates, we consider the galaxy as undetected in a given band and replace its flux by a 1σ limit that is computed from the rms noise at the position of the galaxy. The photometry measured by SExtractor is subsequently corrected for galactic foreground extinction, which we assume to be constant for all galaxies in ECDFS at $E(B - V) = 0.0069$ mag. Figure 13 summarizes the *HST* data available for the galaxies in the ECDFS field. The blue squares show the layout of all of the *HST* pointings as of 2019 October, with darker shades of blue indicating more observations. The ALPINE galaxies at $z < 5$ and $z > 5$ are indicated with orange circles and squares, respectively.

3.2. Photometry on the COSMOS Field

Most of the ALPINE galaxies (105 out of 118) reside in the COSMOS field, and we match them to the latest photometric measurements presented in the COSMOS2015 catalog. The matching is again done on a galaxy-by-galaxy basis using the *HST*/ACS $F814W$ as well as UltraVISTA K_s images as references. We also match against the photometric redshifts given in the catalog in order to identify the correct counterpart.⁴¹

In Table 3, we list all of the photometric measurements available for the ALPINE galaxies on the COSMOS field. Summarizing, these include u^* -band observations from MegaCam on CFHT, the B , V , r^+ , i^+ , z^{++} as well as 12 intermediate-band and two narrowband filters from the Suprime-Cam on Subaru, the Y_{HSC} -band from the Hyper Suprime-Cam on Subaru as well as near-IR bands H_w and K_s^w from WIRC on CFHT and Y , J , H , and K_s from VIRCAM on the VISTA telescope. In addition, the galaxies are covered by the four *Spitzer* channels from 3.6 μm to 8.0 μm from the SPLASH survey⁴² (Capak et al. 2012; Steinhardt et al. 2014; Laigle et al. 2016). As described in (Laigle et al. 2016), the *Spitzer* photometry is measured using IRACLEAN (Hsieh et al. 2012), which uses positional

³⁷ <https://mast.stsci.edu>

³⁸ <https://github.com/gbrammer/mastquery>

³⁹ <https://github.com/grizli-project/grizli-aws>

⁴⁰ <https://github.com/gbrammer/grizli>

⁴¹ The photometric redshifts given in the COSMOS2015 catalog are consistent within 0.2 with our spectroscopic redshifts for more than 95% of all cases.

⁴² <http://splash.caltech.edu>

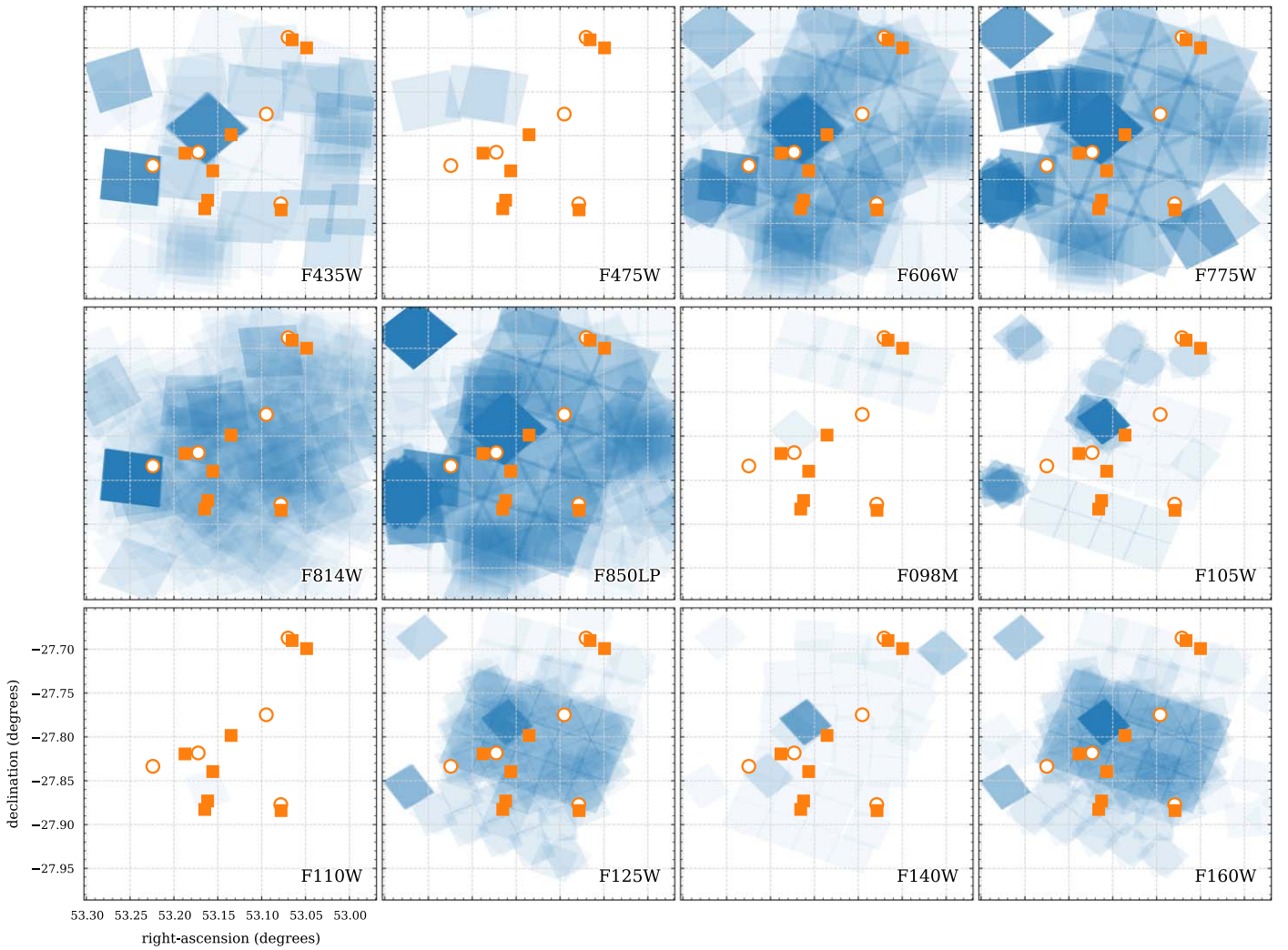


Figure 13. *HST* pointing footprints on the ECFDS field (blue) for different *HST* ACS and WFC3/IR filters. Darker colors mean more observations. The ALPINE galaxies are indicated by orange circles and squares for $z < 5$ and $z > 5$, respectively.

priors from higher-resolution imaging (in this case the $zYJKHK_s$ detection χ^2 -image) to deblend the photometry.

Contrary to the 3D-*HST* photometry catalog, the fluxes and magnitudes in the COSMOS2015 catalog are not total and not corrected for systematic biases and Galactic extinction. To perform these corrections, we follow the steps outlined in the appendix of Laigle et al. (2016). Specifically, we use the 3'' diameter aperture magnitudes (\mathcal{M}^3), which we correct for photometric (o_i , see their Equation (4)) and systematic offsets (s_f , see their Table 3) by applying

$$\mathcal{M}_{i,f}^{\text{tot,uncorr}} = \mathcal{M}_{i,f}^3 + o_i - s_f, \quad (1)$$

where i is the object identifier and f denotes the different filters. The total magnitudes are subsequently corrected for Galactic extinction by applying

$$\mathcal{M}_{i,f}^{\text{tot}} = \mathcal{M}_{i,f}^{\text{tot,uncorr}} - \text{EBV}_i \times F_f, \quad (2)$$

where EBV_i is the Galactic extinction from the Schlegel et al. (1998) maps on COSMOS for each object as given in the catalog and F_f are the extinction factors per filter given in Table 3 of Laigle et al. (2016).

Figure 12 shows the stacked magnitude distribution in z^{++} (z), J , and K_s (K) bands for the $z < 5$ (blue) and $z > 5$ (red) sub-samples. As expected, the high-redshift galaxies are fainter in all bands. In addition, we find the magnitude distributions between the galaxies in ECFDS and COSMOS to be similar. This indicates no major discrepancies in photometric (and, hence, physical) properties between the two samples.

In terms of *HST* imaging, all galaxies except one are observed in ACS *F814W* (Koekemoer et al. 2007; Scoville et al. 2007a). In addition to this, a MAST search shows that several galaxies are covered by other observing programs in the ACS bands *F435W* (3), *F475W* (5), *F606W* (21), and *F850LP* (5) as well as in the WFC3/IR bands, *F105W* (11), *F110W* (5), *F125W* (16), *F140W* (13), and *F160W* (53).⁴³ Note that the observations in *F160W* primarily come from the CANDELS survey (covering the central part of COSMOS; Grogin et al. 2011; Koekemoer et al. 2011) as well as the ‘‘drift and shift’’ (DASH; Momcheva et al. 2017) survey. While the CANDELS imaging is deep (>27.5 magnitudes at 3σ), the data from the DASH survey is much shallower (25.0 magnitudes at 3σ), and

⁴³ There is also data in the *F098W* filter, however, unfortunately no ALPINE galaxies are covered.

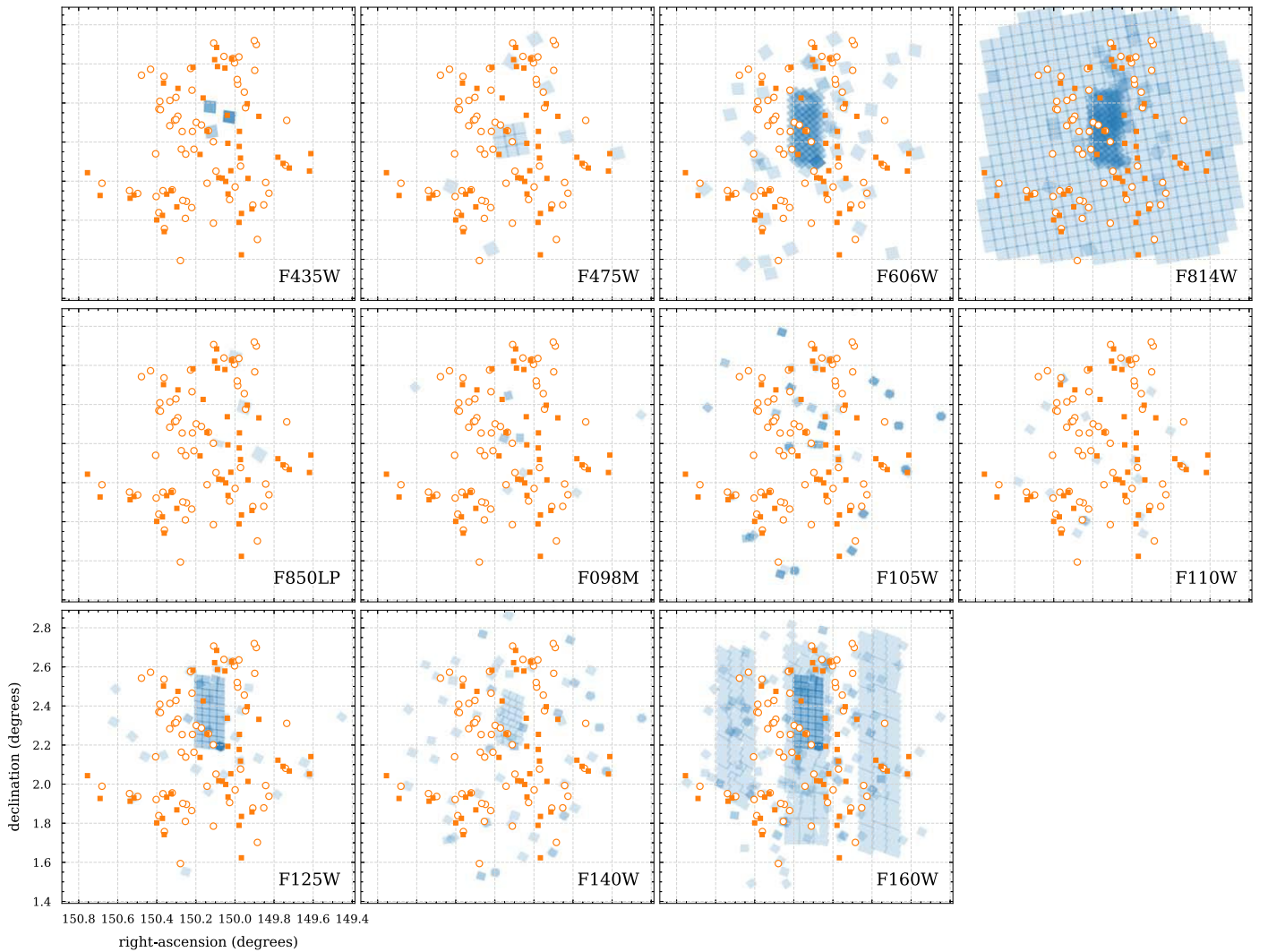


Figure 14. *HST* pointing footprints on the COSMOS field (blue) for different *HST* ACS and WFC3/IR filters. Darker colors mean more observations. The ALPINE galaxies are indicated by orange circles and squares for $z < 5$ and $z > 5$, respectively.

therefore, only half of the galaxies are detected. Furthermore, the spatial sampling in the latter does not allow for a detailed study of the structure of the galaxies. Figure 14 summarizes the *HST* pointings on the COSMOS field (blue). The CANDELS area in the center of the COSMOS field as well as the three DASH stripes are evident. The locations of the ALPINE galaxies are indicated with circles ($z < 5$) and squares ($z > 5$). The photometry available is summarized in Table 3 including depths (where applicable) and references.

In order to measure the photometry in these *HST* bands, we use `SExtractor` in the same way as described in Section 3.1. We first create a mosaic ($0''.06$ pixel sizes and registered to *Gaia*) using the `AWS-drizzler`. We subsequently measure the *HST* photometry of on each of the images for all ALPINE galaxies with coverage using `SExtractor` in apertures ($0''.7$ and $3''$) as well as auto magnitudes. If no object is detected above the set threshold level within $0''.7$ of the original ALPINE coordinates, we consider the galaxy as undetected in a given band. Similar to Equation (2), we correct the *HST* photometry for galactic foreground extinction using the Schlegel et al. (1998) extinction map. We compare these measurements to the ground-based photometry, by first performing a PSF matching by

smoothing the original *HST* images with a Gaussian kernel with FWHM of $0''.7$ and measure the photometry in a $3''$ aperture (as used in the COSMOS2015 catalog). We also compare the ground-based (i^+ , Y , J , H) with the *HST* ($F814W$, $F105W$, $F125W$, $F160W$) photometry in approximately matching filter bands. We find an overall agreement on average of 0.2 magnitudes (dark gray region) with an expected increase of scatter at fainter magnitudes. For the four brightest sources in H -band (< 23.4 AB), the ground-based measured photometry is systematically up to 0.5 magnitudes brighter in three out of four cases. Due to the low number of galaxies, it is difficult to investigate this statistically. A more detailed measurement of the *HST* photometry including deblending of specific sources such as merging or clumpy galaxies, will be provided in a forthcoming paper.

3.3. Astrometric Offsets between ALMA and Ancillary Data

Astrometric accuracy is crucial for many cases. First, accurate spatial offsets of light emission from rest-frame UV, optical, and sub-mm wavelengths reveal the properties of the interstellar medium such as the location and interplay of stars, dust, and gas. Second, the robustness of the identification of

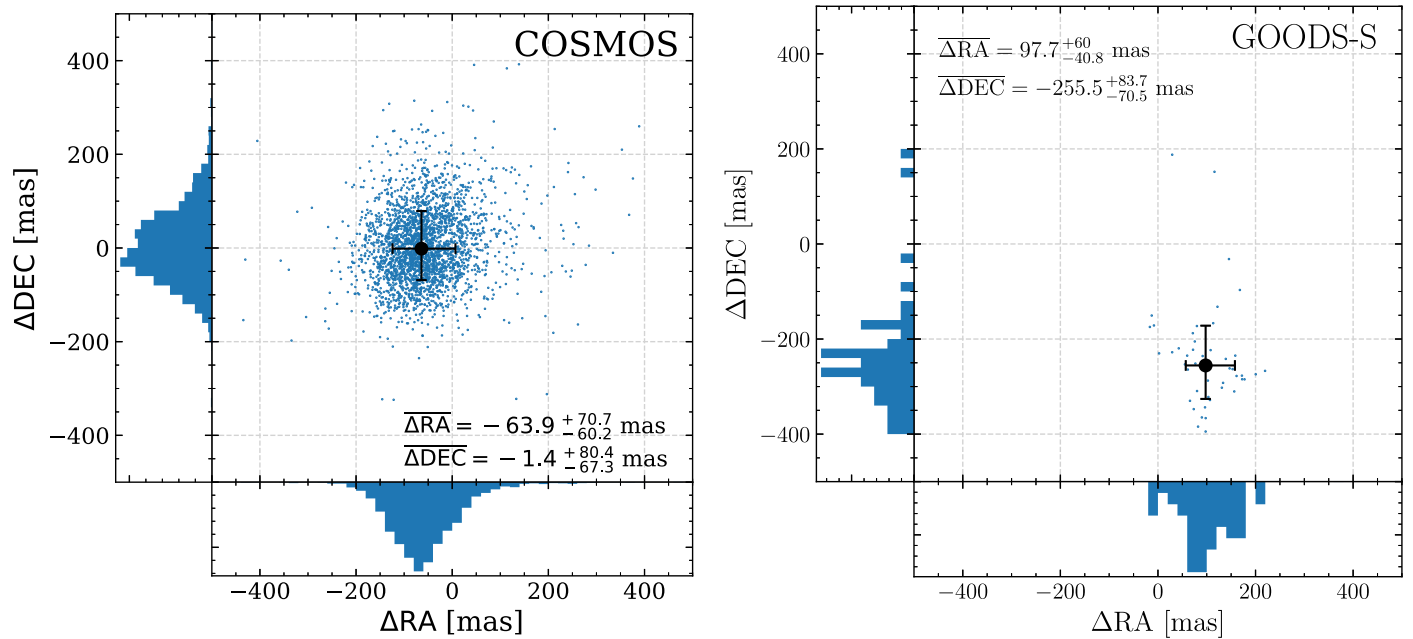


Figure 15. Scatter diagrams and histograms of offsets between the *Gaia* reference frame and the COSMOS2015 (left panel) and 3D-*HST* (ECDFS; right panel) catalogs. The offsets are in the sense “*Gaia*–COSMOS” and “*Gaia*–ECDFS,” respectively. For COSMOS, we find only a systematic offsets in R.A. direction of -64 mas. For ECDFS, the offsets are large in both directions. In addition to this, we measure a scatter in the astrometry of ~ 100 mas for both fields in both coordinates.

sub-mm counterparts needs a high astrometric accuracy of all involved data sets. As described in Sections 3.1 and 3.2, the *HST* images produced for the ALPINE galaxies are all aligned to *Gaia*. Our tests show that the average offsets are less than 15 mas in both R.A. and decl. with a scatter of no more than 30 mas (G. Brammer 2020, private communication). Unfortunately, the positional accuracy for current catalogs (such as 3D-*HST* or COSMOS2015 that are used here) are lower, and we expect significant offsets between those astrometric solution and ALMA data products. In the following, we characterize these offsets.

According to the technical handbook, ALMA observations are currently registered to the International Celestial Reference Frame (ICRF) to an accuracy better than ~ 5 mas. The ICRF is based on hundreds of extragalactic radio sources such as quasars distributed over the whole sky. The positional accuracy Δp of single observations can be estimated by

$$\Delta p = \frac{70000}{\nu \cdot B \cdot \sigma}, \quad (3)$$

where ν is the observed frequency in GHz, B is the maximum baseline length in kilometers, and σ is the S/N at the peak of emission. For ALPINE ($\nu = 330$ GHz and $B = 0.2$ km for C43-1), this leads to $\Delta p = 1060/\sigma$. The calibrators are detected well above 50σ , which leads to an absolute positional accuracy of ~ 20 mas or better.

To check the astrometric alignment of the photometric catalogs used here, we make use of the *Gaia* DR2 catalog (Mignard et al. 2018), which provides currently the most accurate absolute astrometry. As shown in Mignard & Klioner (2018), there are no significant offsets between this reference frame and the ICRF frame used by ALMA; hence, this test directly reveals potential differences in astrometry

between the 3D-*HST* and COSMOS2015 catalogs and our ALMA observations. Using the proper motion information of the stars from *Gaia*, we project their positions back in time to the year of calibration of the data products by using the equations

$$\begin{aligned} \Delta\alpha &= \cos(\delta) \cdot \mathcal{P}_\alpha \cdot (t_{\text{ref}} - t_{\text{gaia}}), \\ \Delta\delta &= \mathcal{P}_\delta \cdot (t_{\text{ref}} - t_{\text{gaia}}), \end{aligned} \quad (4)$$

where α and δ denote the R.A. and decl. (and \mathcal{P}_α and \mathcal{P}_δ their proper motion), t_{gaia} is the *Gaia* reference frame in years (here 2015.5), and t_{ref} is the reference frame of the calibration of the catalogs. To increase accuracy, we only include stars with a proper motion in both coordinates of less than 5 mas yr^{-1} in the following. Note that no parallax motion is included in the above formulae, which would result in less than 5 mas yr^{-1} astrometric shifts.

The COSMOS2015 catalog is calibrated to the MegaCam *i*-band data that was taken in 2004 (Laigle et al. 2016). For ECDFS, the catalog (3D-*HST*) is calibrated to the same reference system as the CANDELS *HST* images, namely ground-based *R*-band data taken in 2001 (Koekemoer et al. 2011; Skelton et al. 2014). However, the exact year is not important as we have selected stars with a relatively slow proper motion. In order to select stars on the COSMOS field, we use the ACS/F814W *SExtractor* catalog (Scoville et al. 2007a) and select sources with a magnitude brighter than 23 AB and *SExtractor* star/galaxy classifier value of > 0.8 . These stars are then matched to the COSMOS2015 catalog to obtain their position in that catalog. For ECDFS, we extract stars directly from the 3D-*HST* catalog by selecting sources brighter than 23 AB in F160W with `star_flag = 1` and a *SExtractor* star/galaxy classifier value larger than 0.8.

Other selections (e.g., different magnitude cuts) do not affect the following results. The *Gaia* stars are then matched to the star catalogs in COSMOS and ECDFS to obtain the astrometric offsets. In total, 47 and 2724 *Gaia* stars are used in ECDFS and COSMOS, respectively.

Figure 15 shows scatter plots and histograms comparing the position of the *Gaia* stars to the positions in the catalogs. We find significant systematic offsets in the astrometry in ECDFS of 98 mas in R.A. and -256 mas in decl.. These offsets are consistent with what was found in earlier studies (Dunlop et al. 2017; Franco et al. 2018; Whitaker et al. 2019). For COSMOS, we only find a significant offset in R.A. of -64 mas. In addition to that, there is a significant scatter in the astrometry on the order of 100 mas in both coordinates in both fields.

To compute the astrometric offset of individual ALPINE galaxies, we make use of the fact that the *HST* images are already aligned to the *Gaia* reference frame (see details in Sections 3.1 and 3.2). Specifically, we compute the offsets between the coordinates measured on the ACS/F814W images and the original coordinates from the 3D-*HST* or COSMOS2015 catalog. If no ACS F814W image is available or if the galaxy is not detected (which happens for redshifts $z > 5$), we use the deepest image red of F814W. If none are available (in four cases), we report the average offset as shown in Figure 15.

Note that the coordinates given in the final ancillary data catalog are not corrected for these offsets. However, we give the offsets for each galaxy in the columns `delta_RA` and `delta_DEC`, which can be *added* to the original coordinates to obtain *Gaia*-corrected R.A. and declination (see Appendix A).

4. Physical Properties

In this section, we detail measurements of various basic physical properties of the ALPINE galaxies that are based on their total, extinction-corrected photometry described in Section 3. These include physical quantities from SED fitting such as stellar masses, SFRs, ages, and dust attenuation (Section 4.1), and UV continuum slopes (Section 4.3), as well as quantities directly derived from the photometry such as UV magnitudes and luminosities (Section 4.2) and estimates of the $H\alpha$ luminosity and equivalent width from *Spitzer* colors (Section 4.4).

4.1. Stellar Mass and SFRs from SED Fitting

4.1.1. Fitting Method

For consistency and comparability with other studies on the COSMOS field, we choose the LePhare SED fitting code⁴⁴ (Arnouts et al. 1999; Ilbert et al. 2006) to derive stellar masses, SFRs, light-weighted stellar population ages, absolute magnitudes, optical dust reddening, and UV continuum slopes of the ALPINE galaxies.

Importantly, stellar masses, SFR, and sSFR values are computed from the marginalized probability distribution functions over all of the models and uncertainties are given in $\pm 1\sigma$ of this distribution.

We use a set of synthetic templates based on the Bruzual & Charlot (2003) stellar population library, which we tune to represent best galaxies at redshifts between $4 < z < 6$. In

detail, we use a series of different star formation histories (SFH) based on exponentially declining (with $\tau = 0.1, 0.3, 1.0$, and 3.0 Gyr), delayed⁴⁵ (with $\tau = 0.1, 0.5, 1.0$, and 3.0 Gyr), and constant star formation. We add dust attenuation corresponding to a stellar $E_s(B - V)$ from 0 to 0.5 spaced in steps of 0.05 assuming a Calzetti et al. (2000) dust attenuation law. To account for metallicity dependence, we use a solar (Z_\odot) and $0.2 Z_\odot$ metallicity. We also adopt a Chabrier (2003) IMF in the following. The model SEDs are generated for logarithmically spaced ages starting from 50 Myr to the age of the universe at the redshift of each galaxy. To each SED, various rest-frame UV and optical emission lines are added following the description in Ilbert et al. (2009) by using common conversions outlined in Kennicutt (1998). Specifically, the UV luminosity at 2300 \AA is converted to an SFR using the relation $\text{SFR}(M_\odot \text{ yr}^{-1}) = 1.4 \times 10^{-28} L_\nu (\text{erg s}^{-1} \text{ Hz}^{-1})$ and subsequently translated to an [O III] emission line flux using the relation $\text{SFR}(M_\odot \text{ yr}^{-1}) = (1.4 \pm 0.4) \times 10^{-41} L_{[\text{O III}]} (\text{erg s}^{-1})$. Other emission lines ($\text{Ly}\alpha$, [O III], $H\beta$, $H\alpha$) are derived by assuming specific ratios to [O III] that are calibrated by observations (see detailed description with references in Ilbert et al. 2009).

The models are fit to the photometry described in Section 3 (and listed in Tables 2 and 3). Specifically, for the galaxies in the ECDFS field, we use the ground-based observations in U , R , J^v , H^v , and K_s^v , as well as the intermediate-bands IA427, IA505, IA527, IA574, IA624, IA679, IA738, IA767, and IA856, and all four *Spitzer* channels ch_1 , ch_2 , ch_3 , and ch_4 . We also include observations in the *HST* filters F435W, F606W, F775W, F814W, F850LP, F125W, F140W, and F160W that are properly combined with the ground-based and *Spitzer* measurements in the 3D-*HST* catalog. For galaxies in the COSMOS field, we include the ground-based observations in u^* , B , V , r^+ , i^+ , z^{++} , Y , Y_{HSC} , J , H , H_w , K_s , and K_s^w , as well as the intermediate-bands IA427, IA464, IA484, IA505, IA527, IA574, IA624, IA679, IA709, IA738, IA767, and IA827. We use all four *Spitzer* channels but no *HST* observations as only the one filter (F814W) exists for all galaxies.

The fits are performed in flux density space (f_ν in Jansky), which has several advantages compared to magnitude space. Specifically, it allows a proper statistical treatment of limits in the data of all bands. While in the case of magnitude limits, an arbitrary significance level (e.g., 1σ , 3σ) has to be defined in order to use them in SED fitting, in flux density space no arbitrary choice needs to be made by the user as the limits manifest themselves only in the error bars of those bands where no flux is measured. This is important, as slightly different levels of significance set for the limits can have profound effects on the output parameters, especially in cases where limits are imposed in astrophysically important parts of the rest-frame spectra (e.g., the Balmer break). In addition, some SED codes, including LePhare, have difficulties dealing with limits (in the case of magnitudes) in a statistically consistent manner, and, instead, remove from consideration any models that slightly exceed these limits. Moving to flux density space for all fits alleviates these concerns.

Several studies have found that the photometric errors are generally underestimated (e.g., Ilbert et al. 2006, 2009, 2013; Skelton et al. 2014; Laigle et al. 2016). Artificially increasing the errors has been found to mitigate issues associated with

⁴⁴ <http://www.cfht.hawaii.edu/~arnouts/lephare.html>

⁴⁵ The delayed SFR is parameterized as $\psi(t) \propto \tau^{-2} t e^{-t/\tau}$ such that $\psi(t)$ is maximal at $t = \tau$.

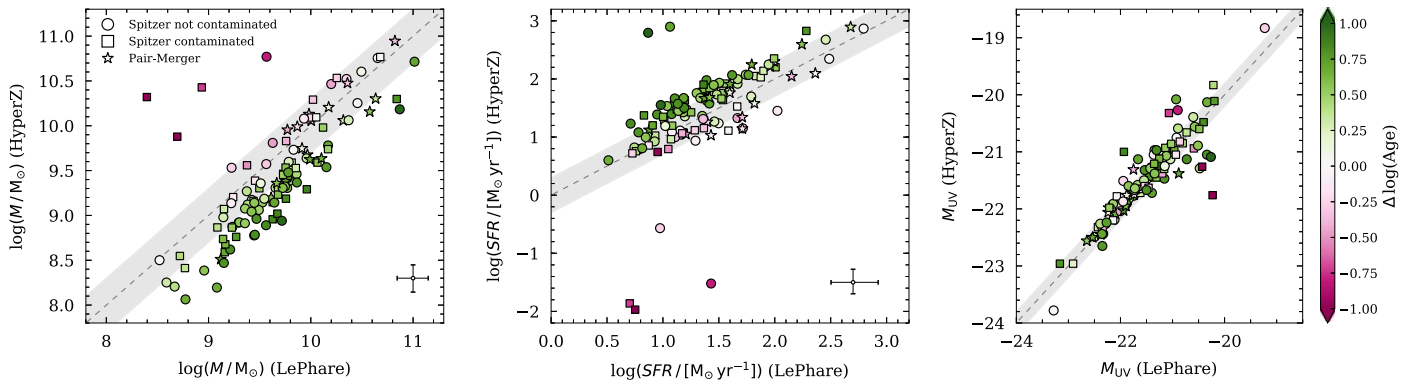


Figure 16. Comparison of stellar mass (left panel), SED-derived SFR (middle panel), and absolute rest-frame UV magnitude (right panel) measured by *LePhare* and *HyperZ*. The symbols are color-coded by the logarithmic difference in stellar population age (note that the scale is the same for all panels). The squares denote galaxies whose *Spitzer* photometry is blended with a nearby bright galaxy or star. The star-symbols denote mergers based on the classification in Le Fèvre et al. (2019). The 1–1 relation is shown as dashed line and a ± 0.3 margin is shown by the gray band. The systematic differences in stellar masses derived by *HyperZ* are likely caused by different implementations of rest-frame optical emission lines, which affect most of the measurements of stellar masses (based on rest-frame optical light).

poorly measured bands, poorly estimated zero-points, or inhomogeneous methods of measuring photometry. To avoid artificially small errors that would dominate the χ^2 budget of a fit of a given galaxy, we follow the prescription of Laigle et al. (2016) and rescale the official flux density errors by a factor of 1.1. In addition, following Ilbert et al. (2009), we correct for underestimated photometric errors due to varying PSF sizes between optical and *Spitzer* imaging by adding in quadrature the following systematic errors to the photometric error budget prior to fitting: 0.01 mag for all ground-based UV/optical broadband measurements; 0.05 mag for all ground-based UV/optical intermediate-bands and near-IR broadband measurements; 0.1 mag for *Spitzer* 3.6 and 4.5 μm and 0.3 mag for *Spitzer* 5.8 and 8.0 μm measurements.

The large PSF sizes of the *Spitzer* observations result in a large risk of blended photometry. This can lead to an overestimation of flux and, hence, to an overestimation of stellar masses.⁴⁶ A cleaner and manual deblending of the *Spitzer* photometry (e.g., using a position prior from *HST* imaging) is possible and will be pursued in a forthcoming paper. In the following, we flag galaxies that have a bright companion galaxy within a *Spitzer* 4.5 μm PSF FWHM ($2''/5$) based on the ground-based and *HST* imaging data. In total, one-third of the galaxies have contaminated *Spitzer* photometry to some degree. For 19% of the galaxies, the *Spitzer* photometry is severely contaminated and should not be trusted. We checked our simple flagging scheme against the contamination flags given in the 3D-*HST* catalog, which are based on the flux fraction that is overlapping with the galaxy for which the photometry is measured. For the 13 ALPINE galaxies in ECDFS, we find excellent agreement between both classifications. The *Spitzer* contamination flags are included in the ancillary data catalog (column `spitzer_cont`).

4.1.2. Systematic Uncertainties in Physical Properties from Modeling Assumptions

Depending on their exact methods, different SED-fitting codes may measure different physical properties even with the same photometry provided as input. In addition to more physical reasons (such as different assumptions on the stellar

population models and implementation of emission lines), the different treatment of undetected fluxes (or fitting in magnitude space), varying minimization techniques and weightings, and different scaling of the error of the input photometry can contribute to these discrepancies.

To investigate the amplitude of such differences, we compare the measurements from *LePhare* to a modified version of the *HyperZ* code (Bolzonella et al. 2000) that includes the effects of nebular emission (Schaerer & de Barros 2009), to estimate such systematic uncertainties in the fitted parameters. To minimize the degeneracy with other assumptions, we run *HyperZ* on the exact same photometry and with the exact same model SEDs (i.e., same metallicity, age, SFH, and dust attenuation law) as described in Section 4.1.1. Figure 16 compares the stellar mass (left panel), SED-derived SFR (middle panel), and absolute UV magnitude (right panel, see also Section 4.2) derived by the two codes. The symbols are color-coded by difference in stellar population age (positive values indicate younger ages derived by *HyperZ*). We find that the stellar masses derived by *HyperZ* are systematically smaller by 0.3–0.4 dex. Similarly, the SFRs are systematically larger by ~ 0.3 dex.

As expected, the absolute UV magnitudes are largely in agreement, as they are, to the first order, independent of the physical parameters and just represent a translation of the fitted UV flux.

The differences in stellar mass and SFR are due to the effect of nebular emission, which results in younger ages and larger emission line corrections of the intrinsic rest-frame optical continuum (observed by the *Spitzer* broad bands) in the *HyperZ* models, and therefore directly affects the stellar mass measurements (see, e.g., de Barros et al. 2014). This is confirmed by the fact that the stellar masses and SFR measured by the two codes agree well (within a factor of two) if the emission lines are turned off. The effect of nebular emission found by *HyperZ* may be overestimated for the ALPINE galaxies, e.g., if the emission lines are more strongly attenuated than the continuum (see Section 4.4). For consistency and comparability with other studies on the COSMOS field, we choose the *LePhare* fitting results as the default.

Next to the systematic offset discussed above, we also find four galaxies (*DC_472215*, *DC_503575*, *DC_722679*, *DC_790930*) whose stellar mass measurements are significantly discrepant, by

⁴⁶ Note that the stellar masses are primarily constrained through the *Spitzer* photometry that covers rest-frame wavelengths redward of the Balmer break.

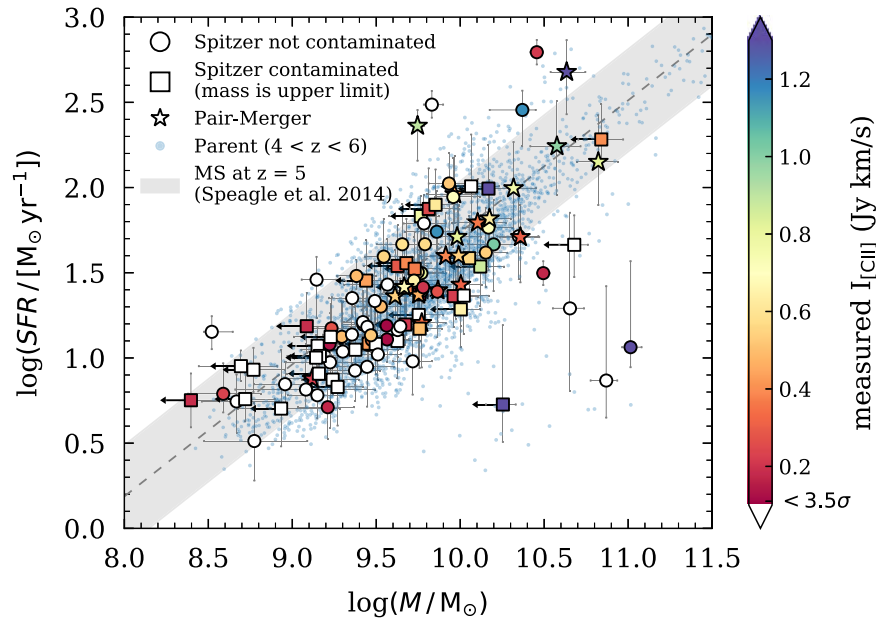


Figure 17. Relation between stellar mass and SFR (main sequence) of our ALPINE galaxies compared to all COSMOS galaxies at $4 < z < 6$ (blue points) and the main-sequence parameterization at $z = 5$ by Speagle et al. (2014) (gray band with ± 0.3 dex width). Galaxies with contaminated *Spitzer* photometry are marked with squares (their stellar mass is likely an upper limit) and mergers (classification by Le Fèvre et al. 2019) are shown as stars. The color denotes the [C II] flux in Jy km s^{-1} measured by ALPINE. Galaxies that are not detected at the 3.5σ level are shown with white face color.

more than one order of magnitude, between the two codes. Specifically, they are fitted with an old, low-SFR, and massive galaxy template with *HyperZ*, while a young, high-SFR, and low-mass galaxy template is preferred by *LePhare*. Three of these outliers have significantly contaminated *Spitzer* photometry (indicated by the squares). As a consequence, the apparent *Spitzer* fluxes and stellar masses are overestimated and the optical colors are artificially reddened, which makes their stellar masses largely unreliable. Furthermore, we artificially increased the errors of the *Spitzer* photometry in our *LePhare* measurements (see Section 4.1.1), hence they have smaller weights, which might reduce the effect of photometric contamination on the fit.

In addition, we investigate the effect of different dust reddening laws on the *LePhare* measurements. For this, we compare a Calzetti et al. reddening with a Small Magellanic Cloud (SMC; Prevot et al. 1984) reddening. The latter might be more suitable for metal-poor low-mass galaxies. Running *LePhare* with the same settings but adopting an SMC reddening curve, we find only small changes in the stellar mass and SFR measurements. Specifically, for the former, we find an average offset toward lower stellar masses in the case of SMC dust of 0.05 dex and a maximal offset of 0.2 dex. For latter, we find similar offsets toward lower SFRs in the case of SMC dust.

Finally, we note that the galaxy properties derived here are consistent within a factor of two with the ones published in the COSMOS2015 catalog (based on *photometric* redshifts). We conclude this from comparing galaxies with the same spectroscopic and photometric redshift within 0.1.

4.1.3. The ALPINE Galaxies on the $z = 5$ Main-sequence and [C II] Fluxes

Figure 17 summarizes the results of this section by showing the relation between stellar mass and SFR (the main sequence) of our ALPINE galaxies together with the [C II] emission (in

Jy km s^{-1}) measured by ALMA (in color). The measurements are compared to all galaxies (with photometric redshifts) at $4 < z < 6$ in the COSMOS2015 catalog (blue points) as well as the main-sequence parameterization by Speagle et al. (2014) at $z = 5$ (gray band with ± 0.3 dex width).

The comparison to the photometrically selected COSMOS parent sample indicates that the ALPINE sample is a fair representation of the overall population of star-forming $z > 4$ galaxies. The sample also includes at higher stellar masses galaxies that lie $2\text{--}3\sigma$ below the main sequence. Note that two of these galaxies at $\log(M/M_\odot) \sim 10.3$ and 10.7 have contaminated *Spitzer* photometry, and therefore, their stellar masses are upper limits. Other two galaxies at $\log(M/M_\odot) \sim 10.6$ and 10.9 do not show [C II] emission, which is expected if they are systems of low SFR below the main sequence. The galaxy with [C II] detection at $\log(M/M_\odot) \sim 10.1$ is classified as “extended dispersion dominated” by our [C II] morpho-kinematic classification (see Le Fèvre et al. 2019) and the *HST/ACS* imaging suggest a clumpy morphology. Although significantly below the main sequence, [C II] is still detected in that galaxy, perhaps indicative of dust-obscured star formation.

The color-coding of the points suggests a [C II] emission increase along the star-forming main sequence. Furthermore, the fraction of [C II] detected galaxies significantly drops below $\log(M/M_\odot) \sim 9.3$ or an SFR of less than $\sim 10 M_\odot \text{ yr}^{-1}$. This could be due to the effect of metallicity on the [C II] line strength by either indirectly a higher ionization state or directly through lower Carbon abundance (e.g., Narayanan et al. 2018). The relationship between star formation and [C II] emission will be studied in more detail in a forthcoming paper (Schaerer et al. 2020).

Note that there are galaxies with an SFR of less than $\sim 10 M_\odot \text{ yr}^{-1}$ in our sample, contrary to our initial selection. We emphasize that the initial selection was based on the observed absolute UV magnitude and not on any property

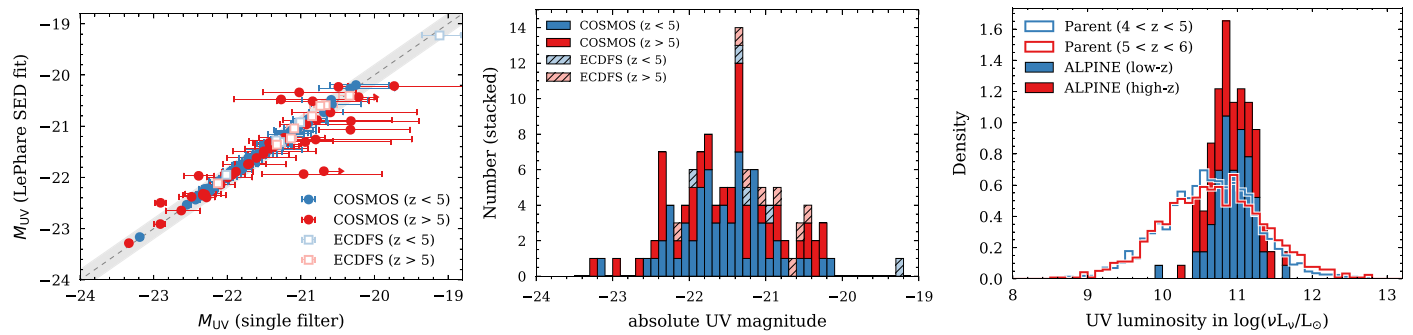


Figure 18. Left panel: comparison of absolute UV magnitudes derived from *LePhare* and from one single filter close to rest frame 1500 Å. The scatter at fainter magnitudes is an S/N effect in the latter measurement. The gray stripe shows ± 0.3 magnitudes around the one-to-one line. Middle panel: distribution of absolute UV magnitudes of the whole ALPINE sample. Right panel: the UV luminosity distribution of ALPINE galaxies in relation to a parent sample from the COSMOS2015 catalog selected in the same redshift range. ALPINE galaxies occupy brighter luminosity, caused by the selection in M_{UV} (see Section 2.1).

derived from SED fitting (such as SFR). This discrepancy is therefore expected within the uncertainty of measuring SFRs from SED fitting.

4.2. Measurement of UV Magnitudes and Luminosities

The UV luminosities and absolute UV magnitudes at rest frame 1500 Å (not dust corrected) are measured during the SED-fitting process with *LePhare* and are defined by the transmission curve of the Galaxy Evolution Explorer (GALEX) FUV filter (~ 1500 Å). As shown in Section 4.1.2, the absolute UV magnitudes measured by *LePhare* and *HyperZ* are in very good agreement.

We also compare these measurements to a more direct method by using the observed flux in a single filter that is closest to rest frame 1500 Å (Subaru z^{++} filter at $z < 5$ and UltraVISTA Y or Y_{HSC} band at $z > 5$ for galaxies in COSMOS and the *HST* filter *F850LP* for all of the galaxies in ECDFS). The left panel in Figure 18 shows a very good agreement between the two methods. The scatter is mainly due to the low S/N of the single-filter measurements of the second method (indicated by the large error bars). The scatter is enhanced for galaxies at $z > 5$ also due to the fact that the UltraVISTA Y -band observations are less deep than the z^{++} observations used for $z < 5$ galaxies. In the following, we will use the more robust absolute UV magnitude from *LePhare* as they depend less on the S/N of single observations.

The middle panel in Figure 18 shows the distribution of absolute UV magnitudes for galaxies in the COSMOS and ECDFS fields in two redshift bins as in the previous figures. The bulk of galaxies are between M_{UV} of -22.7 and -20.2 (consistent with the faint absolute UV magnitude limit of the survey, see Section 1.2). One of the galaxies in ECDFS (*CANDELS_GOODSS_37*) is significantly fainter ($M_{UV} = -19.2$). This galaxy has been added to the sample to fill in an empty frequency window. The measurement of the absolute UV magnitude from single band and SED fit agree and the galaxy is compact and isolated (i.e., no contamination in the photometry). The fit to its photometry with *LePhare* suggests a dust-free, low-mass ($\log(M/M_{\odot}) = 9.22$) galaxy that is forming stars at a rate typical for the main sequence ($\log(\text{SFR}/[M_{\odot} \text{ yr}^{-1}]) = 0.97$).

The right panel of Figure 18 compares the distribution of the UV luminosity of the ALPINE galaxies (stacked filled histogram for low and high redshift) to the same parent sample selected from the COSMOS2015 catalog and used in Figure 17 split in two redshift bins (empty histogram). As expected, the

absolute UV magnitude cut applied for the selection of the ALPINE sample (see Section 2.1) causes a bias toward the brighter end of the parent distribution.

4.3. UV Continuum Slopes

4.3.1. Method

The UV continuum slope (β , defined as $f_{\lambda} \propto \lambda^{\beta}$) generally correlates with the attenuation of stellar light by dust and is therefore an important tool to study the dust properties of galaxies especially at high redshifts (e.g., Meurer et al. 1999; Bouwens et al. 2009; Finkelstein et al. 2012). The UV continuum slope of a galaxy can be derived by various methods. Here, we compute β from the best-fit SEDs derived by *LePhare*. Compared to deriving the slopes directly from the observed photometry by a linear fit, this approach results in less-biased β measurements in the case of low-S/N observations (e.g., Finkelstein et al. 2012; Barisic et al. 2017). This is particularly true for the ALPINE galaxies, whose rest-frame UV continuum is predominantly covered by relatively shallow ground-based imaging. Deep *HST* coverage by a sufficient number of bands of this wavelength range is only available for a small fraction of the galaxies.

The β slopes are derived by a robust linear fit (to avoid the fit being affected by any absorption or emission lines) to the logarithmic slope of the best-fit *LePhare* SED in the wavelength range between 1300 Å and 2300 Å. To quantify uncertainties, we perturb the fluxes of each filter according to their individual errors assuming a Gaussian error distribution, refit the galaxies, and re-measure β from the resultant best-fit SED. For each galaxy, we repeat this procedure 1000 times to produce a probability density function. The uncertainties for β quoted here are the 1σ percentiles of this distribution and are, on average, on the order of $\Delta\beta = 0.2-0.3$.

4.3.2. Systematic Uncertainties and Dependencies on Dust Attenuation

In addition to photometric uncertainties, several model assumptions affect the measurement of β . We found that the two most important ones are the wavelength range over which β is fit and the assumed dust attenuation law.

We choose the wavelength range over which β is measured to be consistent with the definition of several other studies (Calzetti et al. 1994; Meurer et al. 1999; Finkelstein et al. 2012; Bouwens et al. 2014; Barisic et al. 2017; Fudamoto et al. 2017). In the upper panel of Figure 19, we demonstrate how the β measured for our