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The Fornax Deep Survey with VST

X. The assembly history of the bright galaxies and intra-group light in the Fornax A subgroup

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ABSTRACT

Context. We present the study of the south-west group in the Fornax cluster centred on the brightest group galaxy (BGG) Fornax A, which was observed as part of the Fornax Deep Survey (FDS). This includes the analysis of the bright group members ($m_B < 16$ mag) and the intra-group light (IGL).

Aims. The main objective of this work is to investigate the assembly history of the Fornax A group and to compare its physical quantities as a function of the environment to that of the Fornax cluster core.

Methods. For all galaxies, we extracted the azimuthally averaged surface brightness profiles in three optical bands (g, r, i) by modelling the galaxy’s isophotes. We derived their colour ($g - i$) profiles, total magnitude, effective radius in all respective bands, stellar mass, and the break radius in the r -band. The long integration time and large covered area of the FDS allowed us to also estimate the amount of IGL.

Results. The majority of galaxies in the Fornax A group are late-type galaxies (LTGs), spanning a range of stellar mass of $8 < \log(M_* M_\odot) < 10.5$. Six out of nine LTGs show a Type III (up-bending) break in their light profiles, which is either suggestive of strangulation halting star formation in their outskirts or their HI-richness causing enhanced star formation in their outer-discs. Overall, we do not find any correlations between their physical properties and their group-centric distance. The estimated luminosity of the IGL is $6 \pm 2 \times 10^{10} L_\odot$ in the g -band, which corresponds to about 16% of the total light in the group.

Conclusions. The Fornax A group appears to be in an early-stage of assembly with respect to the cluster core. The environment of the Fornax A group is not as dense as that of the cluster core, with all galaxies except the BGG showing similar morphology, comparable colours and stellar masses, and Type III disc-breaks, without any clear trend in these properties with group-centric distances. The low amount of IGL is also consistent with this picture, since there were no significant gravitational interactions between galaxies that modified the galaxies’ structure and contributed to the build-up of the IGL. The main contribution to the IGL is from the minor merging in the outskirts of the BGG NGC 1316 and, probably, the disrupted dwarf galaxies close to the group centre.

Key words. surveys – galaxies: groups: individual: Fornax A – galaxies: spiral – galaxies: structure – galaxies: evolution – galaxies: photometry

1. Introduction

Galaxies tend to gather in gravitationally bound systems during their evolution, which are seen as dense knots in the cosmic web filaments. In the Local Universe, it has been found that about half of the population of galaxies are found in groups and clusters

(Karachentsev 2005). According to the hierarchical structure formation, these galaxy clusters are formed from the merging of smaller structures such as galaxy groups (e.g. Blumenthal et al. 1984; Sommer-Larsen 2006; Rudick et al. 2009). The terminology “galaxy group” and “galaxy cluster” depends on various factors, for example, the galaxy number density, the virial mass,

X-ray luminosity, and velocity dispersion. Groups of galaxies or poor clusters have masses in the range from 10^{13} to $10^{14} M_{\odot}$ (Bower & Balogh 2004; Mamon et al. 2007), and rich clusters are more massive $>10^{14} M_{\odot}$ (e.g. Bahcall 1996). In addition to this, the intra-cluster medium (ICM) of rich clusters is hotter (X-ray temperature = 2–14 keV) than that of poor clusters and groups, and the radial velocity dispersion of galaxies in rich clusters is higher than that of groups (e.g. Bahcall 1988, 1996).

The definition of these gravitationally bound systems range from binary galaxies to superclusters, all of which are important in the studies of environmental mechanisms. Several surveys have shown variations in observed properties of galaxies in different environments, that is, from the field to the core of clusters in the Local Universe (e.g. Lewis et al. 2002; Kauffmann et al. 2004; Peng et al. 2010; Grootes et al. 2017; Barsanti et al. 2018) and also at higher redshifts, up to $z \sim 1$ (e.g. Poggianti et al. 2008; Lubin et al. 2009; Cucciati et al. 2010, 2017; Cooper et al. 2012; Muzzin et al. 2012; Paulino-Afonso et al. 2018; Guglielmo et al. 2019). Such transitions in galaxy properties typically occur around cluster or group scales of ~ 2 – 3 virial radii (Gómez et al. 2003). One such observed property is the decrease in the star formation rate in galaxies of a given morphological type in groups and clusters that is visible out to $2r_{200}$ (Lewis et al. 2002). The mechanisms causing such changes are proposed to depend on the local density of the environment and, hence, they are different in groups and clusters.

The physical processes which are prominent in dense environments of clusters are ram-pressure stripping (Gunn & Gott 1972) and harassment (Moore et al. 1996). Ram-pressure stripping is not prevalent in poor clusters and galaxy groups as it depends on the velocity of the galaxy and the hot ICM. Mergers and strangulation are more common in groups of galaxies (e.g. Barnes 1985; Zabludoff & Mulchaey 1998). Major mergers can significantly alter the morphology of spiral galaxies by destroying their discs and therefore producing a major remnant that resembles an elliptical galaxy (Dekel et al. 2005). Sometimes, the disc can survive gas-rich mergers albeit resulting in the build-up of a galactic bulge (e.g. Stewart et al. 2009; Hopkins et al. 2010). On the other hand, disc strangulation is a result of the tidal forces in the group and cluster that suppresses the infall of gas onto the discs of galaxies which in turn reduces star formation (Larson et al. 1980). The aforementioned processes that occur in groups of galaxies are most effective when the velocity dispersion of the group is similar to the internal velocity (rotation or dispersion) of its galaxy members (Barnes 1985; Zabludoff & Mulchaey 1998).

During the infall of galaxy groups which then form a cluster, the stripped material of galaxies settles to constitute the intra-group light (IGL) or the intra-cluster light (ICL; Napolitano et al. 2003; Fujita 2004; Contini et al. 2014). Therefore, the ICL and IGL are fossil records of past interactions and mergers as their growth and evolution over time provide a link to the assembly state of the cluster and group (Mihos 2015, and references therein).

Overall, high density environments provide a plethora of information for galaxy formation and evolution. There have been several studies on the effect of the environment on galaxy structures and properties in the Local Universe (e.g. Dressler 1980; Pastoriza et al. 1994; de Mello et al. 2002; Lewis et al. 2002; Gómez et al. 2003; Bower & Balogh 2004; Boselli et al. 2006; Mamon et al. 2007; Huertas-Company et al. 2013; Tully et al. 2014; Deeley et al. 2017; Finn et al. 2018, and references therein) and at higher redshifts, up to $z \sim 1$ (e.g. Cooper et al.

2006; van der Wel et al. 2007; Tasca et al. 2009; Nantais et al. 2013; Kelkar et al. 2015; Allen et al. 2016; Krywult et al. 2017; Paulino-Afonso et al. 2019, and references therein). With the advent of deep photometry (e.g. Zibetti et al. 2005), exploring faint regions of galaxy outskirts, ICL and IGL have gained importance over the last years (e.g. Rudick et al. 2009; Mihos et al. 2017; Iodice et al. 2016, 2017a,b; Spavone et al. 2017, 2018, and references therein).

As such, the Fornax cluster has been an environment of interest. It is the second most massive galaxy concentration within 20 Mpc, after the Virgo cluster, with a virial mass of $M_{\text{vir}} = 7 \times 10^{13} M_{\odot}$ (Drinkwater et al. 2001a). The Fornax cluster has been extensively studied by several surveys in multiple wavelengths (e.g. Martin & GALEX 2005; Scharf et al. 2005; Jordán et al. 2007; Davies et al. 2013; Muñoz et al. 2015; Su et al. 2017; Eigenthaler et al. 2018; Zabel et al. 2019; Serra et al. 2019).

The Fornax Deep Survey (FDS) with the VLT Survey Telescope (VST, Venhola et al. 2017; Iodice et al. 2019a) is one of the deepest and widest data sets mapping the Fornax cluster out to the virial radius (~ 0.7 Mpc). With FDS we (i) mapped the surface brightness around the brightest cluster members NGC 1399 and NGC 1316 out to an unprecedented distance of about ~ 200 kpc ($R \sim 6R_e$) and down to $\mu_g \approx 29$ – 31 mag arcsec $^{-2}$ (Iodice et al. 2016, 2017b); (ii) traced the spatial distribution of candidate globular clusters (GCs) inside ~ 0.5 deg 2 of the cluster core (D’Abrusco et al. 2016; Cantiello et al. 2018); (iii) studied the galaxy outskirts, detected the ICL and faint ($\mu_g \approx 28$ – 30 mag arcsec $^{-2}$) features in the intra-cluster region in the core of the cluster (Iodice et al. 2016, 2017a, 2019a; Raj et al. 2019) and in the outskirts of NGC 1316 (Iodice et al. 2017b); (iv) provided the largest size and magnitude limited catalogue of dwarf galaxies in the cluster (Venhola et al. 2017, 2018, 2019). The analysis of the deep images from FDS suggests that the Fornax cluster is not completely relaxed inside the virial radius. The bulk of the gravitational interactions between galaxies takes place in the west-northwest core region, where most of the bright early-type galaxies (ETGs) are located and where the intra-cluster baryons (i.e. diffuse light, GCs, and PNe) are found (Pota et al. 2018; Spiniello et al. 2018; Iodice et al. 2019a).

This paper focuses on the south-west (SW) Fornax A group using the FDS data to study the bright group members ($m_B < 16$ mag) and their physical properties as a function of the environment and, thus, compare them with those observed for late-type galaxies (LTGs) in the Fornax cluster. Maddox et al. (2019) presented a spectroscopic census of all the galaxies in the Fornax environment, covered by FDS which also includes background objects. We use this data to study and compare the kinematics of the assembly history of the bright galaxies in the Fornax A subgroup to those of the core.

The outline of this paper is as follows. In Sect. 2, we give a brief summary of FDS. In Sect. 3, we describe our methods of analysis along with the estimated parameters. In Sect. 4, we explain the results based on the surface photometry analysis. In Sect. 5, we address the assembly history of the bright galaxies in Fornax A. In Sect. 6, we discuss the formation history of the IGL in the Fornax A group. In Sect. 7, we compare the environments of the Fornax cluster core and Fornax subgroup. In Sect. 8, we summarise our results. A brief description of each of the galaxies presented in this work is given in Appendix A and their corresponding images (g -band), surface brightness profiles, ($g-i$) colour maps and colour profiles are shown in Appendix B. Methodologies of some of this work are presented in Appendix C.

2. The Fornax Deep Survey

The Fornax Deep Survey is a multi-band (u, g, r, i) survey, and a part of the Guaranteed Time Observation surveys FOCUS (P.I. R. Peletier) and VEGAS (P.I. E. Iodice, Capaccioli et al. 2015). The observations were acquired with the ESO VLT Survey Telescope (VST), located at Cerro Paranal, which is a 2.6 m optical telescope (Schipani et al. 2012) equipped with the wide field camera OmegaCam (Kuijken 2011), that covers $1 \times 1 \text{ deg}^2$ field of view. FDS data sample 26 deg^2 of the Fornax cluster, which also includes the in-falling Fornax A subgroup (Iodice et al. 2017b). Detailed explanations of the observing strategy and reduction are given by Iodice et al. (2016, 2017a) and Venhola et al. (2018). The FWHM (arcsec) and depth (mag arcsec⁻²) of the FDS fields in each filter (g, r, i), used in this work, are given in Table 1 (see also Venhola et al. 2018).

Catalogues and data products provided by FDS were presented by Venhola et al. (2018) on dwarf galaxies in the Fornax cluster, Iodice et al. (2019a) on photometry of ETGs, and Raj et al. (2019) on LTGs in the central 9 deg^2 . The g -band mosaic covering $3.6 \times 2 \text{ deg}^2$ around NGC 1316 and the SW side of the Fornax cluster with named circles indicating the galaxies studied in this work, is shown in Fig. 1. In this image, we also mark all galaxies taken from the catalogue by Venhola et al. (2019) that are part of the Fornax A subgroup for the purpose of group-environmental analysis. The galaxies marked in Fig. 1 are grouped into giant LTGs ($m_B < 16 \text{ mag}$), dwarf elliptical (dE), and dwarf late-type (dI) galaxies. The sample of spiral galaxies ($m_B < 16 \text{ mag}$) has been selected from Ferguson (1989) and are listed in Table 2. In this study, we adopt a distance of $D = 20.8 \pm 0.5 \text{ Mpc}$ for NGC 1316 (Cantiello et al. 2013).

3. Analysis

In this section, we briefly describe the procedures to derive galaxy parameters. We adopt the methods presented by Iodice et al. (2019a) and Raj et al. (2019). Results for individual galaxies of the sample are given in Appendix A.

3.1. Isophote fitting

The azimuthally averaged surface brightness (SB) profiles of each source were extracted in three bands (g, r, i) using the IRAF task ELLIPSE¹ (Jedrzejewski 1987). As preliminary steps to the isophote fitting, we masked all bright sources (e.g. galaxies, stars) that are close to the galaxies in the sample. In some cases, for galaxies that are close in projection to each other (NGC 1326A and NGC 1326B; NGC 1316 and NGC 1317), we adopted an iterative process by modelling the neighbouring source using the BMODEL task in IRAF. This model is created from the results of the isophotal analysis (ELLIPSE) and then subtracted from the original image. In the fitting routine, we adopted the k-sigma clipping algorithm for cleaning deviant sample points at each annulus².

¹ IRAF is distributed by the National Optical Astronomy Observatories, which are operated by the Association of Universities for Research in Astronomy, Inc., under cooperative agreement with the National Science Foundation

² Overall, the manual masking, 3σ clipping and model-subtraction improve the isophotal fitting, ensuring the rejection of unmasked sources and defects.

Isophotes are fitted in elliptical annuli, starting from the centre³ of the galaxy to the edge of the FDS field ($\sim 0.5 \text{ deg}$). Since most of the galaxies have dust absorption and/or irregular shape close to their centre, we keep the geometrical centre fixed, while the position angle and ellipticity are free parameters. According to Iodice et al. (2016), the intensity profiles are then used to derive (i) the limiting radius (R_{lim}) that corresponds to the outer-most annulus, where the light of the source blends into the average background level (it is the residual after subtracting the sky frame resulting in a value close to zero); (ii) derive the residual sky background of each source from their outer annuli in three respective bands. The limiting radius and the residual background levels are locally estimated for each galaxy in the sample. The background levels are computed at $R \geq R_{\text{lim}}$ as the average counts out to the edge of the image.

3.2. Products: Total magnitude, effective radius, colours, and stellar mass

For each galaxy of the sample, we derived the “total magnitude”, “effective radius” R_e , average “ $g-i$ ” and “ $g-r$ ” colours following the method presented by Iodice et al. (2019a) and Raj et al. (2019). All parameters are given in Table 3. Firstly, the intensity profiles from the output of the isophote fit are used to derive azimuthally averaged SB profiles. The residual background level estimated at $R \geq R_{\text{lim}}$ of each band for all galaxies is subtracted. We also took into account the uncertainties on the photometric calibration and sky subtraction in estimating the errors on magnitudes (see Capaccioli et al. 2015; Iodice et al. 2016, 2019a). For all galaxies in the sample, we derived the $g-i$ colour profiles from the SB profiles, and colour maps from the images, which are shown in Appendix B. We then derived total magnitudes and effective radii in the g, r, i bands of all galaxies in our sample using the growth curve from the isophote fits. We estimated the stellar mass M_* using the empirical relation from Taylor et al. (2011) which assumes a Chabrier IMF, that is, $\log_{10} \frac{M_*}{M_\odot} = 1.15 + 0.70(g-i) - 0.4M_i$. Here M_i is the absolute magnitude in i -band, $g-i$ is the average colour derived in the previous step. We then estimated the stellar mass-to-light ratio M_*/L_i with 1σ accuracy of 0.1 dex (see Taylor et al. 2011).

3.3. Estimation of the IGL

The IGL is estimated from the mosaic in the g and r bands as their surface brightness limit is fainter than that of the i -band (see Table 1). The brightest stars and galaxies surrounding NGC 1316 are modelled using the IRAF tasks ELLIPSE and BMODEL as explained in Sect. 3.1, and then subtracted from the image. The resultant residual image is used to derive the radial profiles and thus 2D model of NGC 1316, which is also subtracted from the image. The final residual image is used to estimate the IGL using the IRAF task PHOT, which computes the integrated flux inside a specified aperture. The uncertainties on the estimate of the IGL take into account (1) the error on the background estimate ($\approx 10\%$ in g , $\approx 50\%$ in r); (2) the poissonian error on the flux ($\approx 26\%$ in g and r) and (3) the error on the photometric calibration ($\approx 0.1\%$ in g and r).

According to Iodice et al. (2016), the stellar envelope of NGC 1316 starts to dominate at $R > R_{\text{tr}} = 5.5'$, where R_{tr}

³ In the dust-free galaxies, the centre corresponds to the intensity peak. If the galaxy light is strongly affected by the dust absorption, we adopted the geometric centre of the isophotes.

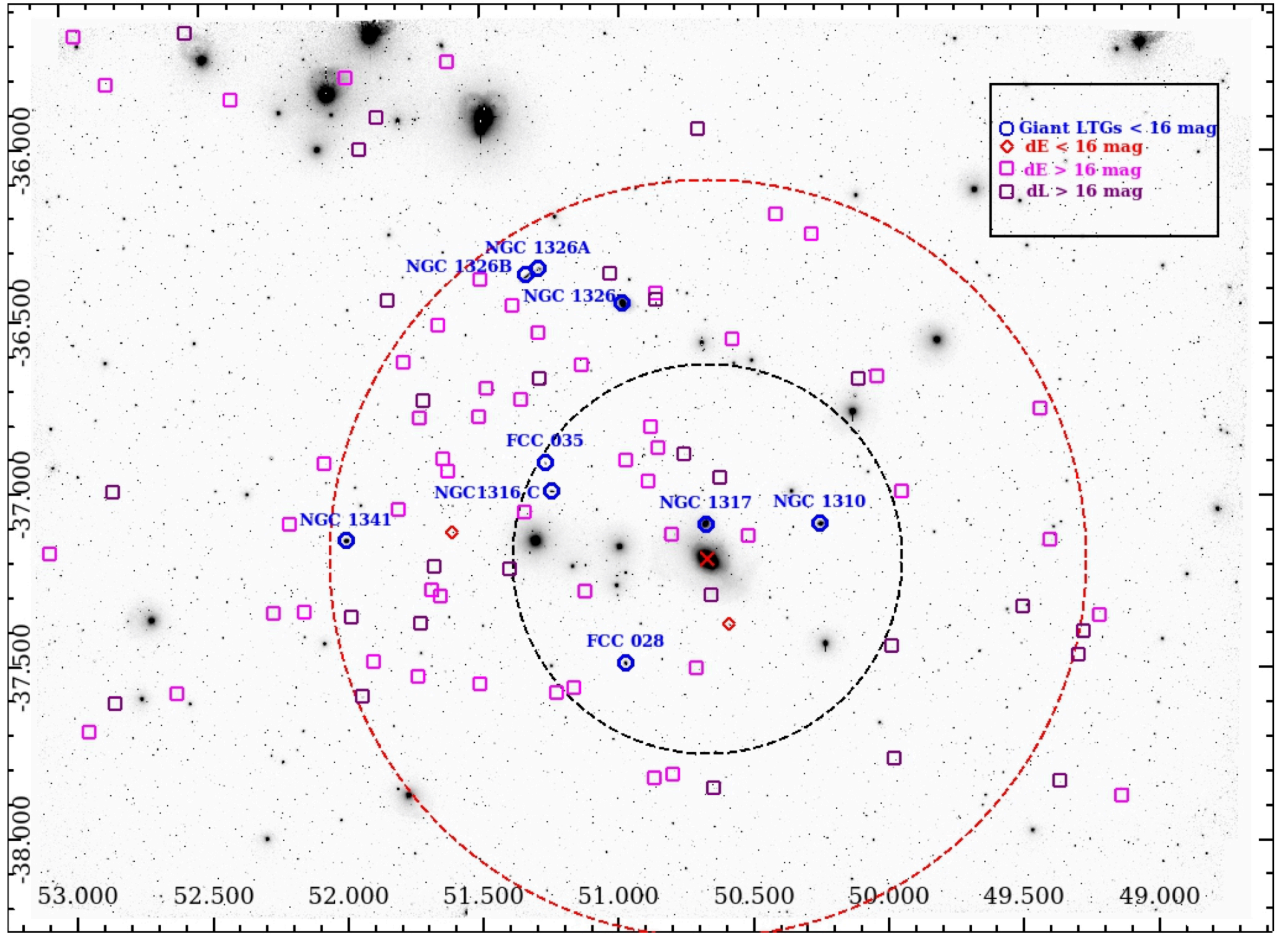


Fig. 1. VST mosaic in g -band of the Fornax A subgroup covering $3 \times 2 \text{ deg}^2$. The red dashed circle is an approximate measure of the virial radius of the group, which is $\sim 1 \text{ deg}^2$ (Drinkwater et al. 2001a). The black dashed circle ($r = 0.58 \text{ deg}$) is the region where the IGL is estimated. All galaxies part of the Fornax cluster and Fornax A subgroup, taken from Venhola et al. (2019) are indicated. The red cross indicates NGC 1316. Bright galaxies ($m_B < 16 \text{ mag}$) are represented with blue circles. Red diamond represents dE ($m_B < 16 \text{ mag}$), pink squares represent dE ($m_B > 16 \text{ mag}$), and purple squares represent dl ($m_B > 16 \text{ mag}$).

Table 1. Image quality of FDS fields used in this work.

Field	g band		r band		i band	
	$FWHM$ [arcsec]	Depth [$\frac{\text{mag}}{\text{arcsec}^2}$]	$FWHM$ [arcsec]	Depth [$\frac{\text{mag}}{\text{arcsec}^2}$]	$FWHM$ [arcsec]	Depth [$\frac{\text{mag}}{\text{arcsec}^2}$]
(1)	(2)		(3)		(4)	
F22	1.04 ± 0.07	26.52	0.81 ± 0.05	25.90	0.85 ± 0.07	25.16
F25	1.11 ± 0.10	26.63	0.77 ± 0.06	25.84	0.85 ± 0.07	25.11
F26	0.93 ± 0.07	25.89	0.81 ± 0.05	25.96	0.91 ± 0.07	25.06
F27	1.07 ± 0.10	26.39	0.78 ± 0.06	25.63	0.89 ± 0.10	24.88
F28	1.08 ± 0.14	26.31	0.79 ± 0.06	25.57	0.92 ± 0.09	24.89

Notes. Column 1 – FDS fields ; Cols. 2 to 4 – average seeing and rms of the FWHM, and surface brightness corresponding to 1σ S/N per arcsec² for g, r, i bands. The PSF FWHM and depth of each filter in all fields are adopted from Venhola et al. (2018).

is the transition radius⁴. Taking this into account, we masked this region of NGC 1316. The IGL is estimated in a region

⁴ The transition radius R_{tr} is defined as the distance from the centre of the galaxy to the region where an inflection (evident in their SB profiles) in the stellar halo occurs (Cooper et al. 2010, 2013).

$\sim 33.5 \text{ arcmin}^2$ surrounding NGC 1316, which covers nearly 60% of the group including the bright galaxies. In Appendix C, we show the final g -band image used in estimating the IGL of the Fornax A subgroup. In this image, we also mark the loops that are part of the stellar envelope (see also Iodice et al. 2017b). Through our deep photometry of this group, we also find a new

Table 2. LTGs brighter than $m_B < 16$ mag in the Fornax A subgroup.

FCC name	α	δ	Morph. type	Radial velocity	m_B	FDS field	Alternative names
(1)	h:m:s	d:m:s	(4)	km s ⁻¹	mag	(7)	(8)
FCC013	03 21 03.4	-37 06 06	SBcII	1805	12.55	F28	NGC 1310, ESO 357-G 019
FCC022	03 22 44.3	-37 06 13	Sa pec	1941	11.91	F26	NGC 1317, ESO 357-G 023
FCC028	03 23 54.2	-37 30 33	SmIII	1405	13.88	F26	ESO 301-IG 011
FCC029	03 23 56.4	-36 27 53	SBa(r)	1360	12.20	F25	NGC 1326, ESO 357-G 026
FCC033	03 24 58.4	-37 00 34	SdIII pec	1800	14.32	F26	NGC 1316C, ESO 357-G 027
FCC035	03 25 04.2	-36 55 39	SmIV	1800	15.3	F26	
FCC037	03 25 08.5	-36 21 50	SBcIII	1831	13.77	F25	NGC 1326A, ESO 357-IG 028
FCC039	03 25 20.3	-36 23 06	SdIII	999	13.59	F25	NGC 1326B, ESO 357-G 029
FCC062	03 27 58.4	-37 09 00	SbcII	1876	13.21	F22	NGC 1341, ESO 358-G 008

Notes. Column 1 – Fornax A subgroup members from [Ferguson \(1989\)](#); Cols. 2 and 3 – Right Ascension and Declination (J2000); Col. 4 – Morphological type from [Ferguson \(1989\)](#); Col. 5 – Heliocentric radial velocity obtained from NED; Col. 6 – Total magnitude in B band as given in NED (The NASA/IPAC Extragalactic Database (NED) is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration.); Col. 7 – Location in the FDS Field; Col. 8 – Alternative catalogue names.

Table 3. Derived parameters of LTGs in the Fornax A subgroup.

Object	m_g	$g-r$	$g-i$	Re_r	D_{core}	M_i	M_*	M/L
(1)	mag	mag	mag	arcsec	deg	mag	$10^{10} M_\odot$	(9)
FCC013 (NGC 1310)	12.48 ± 0.04	0.57 ± 0.1	0.77 ± 0.02	27.6 ± 0.2	0.25	-19.74	0.47	0.76
FCC022 (NGC 1317)	11.15 ± 0.01	0.77 ± 0.02	0.99 ± 0.02	35.36 ± 0.03	0.02	-20.98	1.71	0.88
FCC028	13.33 ± 0.02	0.57 ± 0.04	0.98 ± 0.05	22.1 ± 0.5	0.28	-19.07	0.29	0.87
FCC029 (NGC 1326)	10.62 ± 0.02	0.62 ± 0.04	1.01 ± 0.04	48.2 ± 1	0.68	-21.53	2.94	0.91
FCC033 (NGC 1316C)	13.81 ± 0.03	0.69 ± 0.1	0.97 ± 0.04	22.6 ± 0.5	0.37	-18.30	0.14	0.85
FCC035	14.79 ± 0.03	0.17 ± 0.1	0.18 ± 0.1	17.1 ± 0.4	0.44	-17.39	0.017	0.24
FCC037 (NGC 1326A)	13.55 ± 0.03	0.51 ± 0.1	0.95 ± 0.1	41.0 ± 0.8	0.87	-18.54	0.17	0.83
FCC039 (NGC 1326B)	12.87 ± 0.1	0.32 ± 0.1	0.73 ± 0.1	53.03 ± 2	0.86	-19.00	0.18	0.58
FCC062 (NGC 1341)	12.37 ± 0.02	0.54 ± 0.04	1.01 ± 0.04	20.6 ± 0.4	0.92	-19.72	0.55	0.91

Notes. Column 1 – Fornax A subgroup members from [Ferguson \(1989\)](#); Col. 2 – Total magnitude in g -band derived from the isophote fit. Values were corrected for the Galactic extinction, using the absorption coefficient by [Schlegel et al. \(1998\)](#); Cols. 3 and 4 – Average $g-r$ and $g-i$ colours; Col. 5 – Effective radius in r -band, in arcsec; Col. 6 – Projected distance from the galaxy centre in degree, that is, from NGC 1316 (FCC021); Col. 7 – Absolute magnitude in i -band, derived using the distance modulus from NED and [Tully et al. \(2009\)](#); Cols. 8 and 9 – Stellar mass and mass-to-light (M/L) in i -band.

patch of diffuse light in the SW part of NGC 1316, which also contributes to the IGL (see Appendix C).

The total magnitude⁵ of the IGL in the Fornax A group, within an area of ~ 33.5 arcmin² around NGC 1316, is $m_g = 10.1 \pm 0.3$ mag and $m_r = 9.2 \pm 0.3$ mag in the g and r bands, respectively. Therefore, the total luminosity is $6 \pm 2 \times 10^{10} L_\odot$ in the g band and $7 \pm 2 \times 10^{10} L_\odot$ in the r band. As pointed out earlier, the IGL estimate includes the stellar envelope of NGC 1316. The fraction of the IGL luminosity to the (i) luminosity of the brightest group galaxy (BGG) NGC 1316 is $L_{\text{IGL},g}/L_{\text{BGG},g} \sim 19\%$; (ii) total luminosity of all group members (including dwarf galaxies) inside the virial radius is $L_{\text{IGL},g}/L_{\text{group},g} \sim 16 \pm 4\%$. The total luminosity of NGC 1316 ($L_{\text{BGG},g} = 3.11 \times 10^{11} L_\odot$) is given by [Iodice et al. \(2017b\)](#). The total luminosity for dwarf galaxies ($9 \times 10^8 L_\odot$ in g -band) is provided by [Venholo et al. \(2019\)](#). The total luminosity for all bright LTGs in the sample studied in this work are derived from magnitudes listed in Table 3.

⁵ Values were corrected for Galactic extinction using the absorption coefficient by [Schlegel et al. \(1998\)](#).

The contribution to the IGL (in g -band) of the luminous loops in the stellar envelope of NGC 1316 (loops L1, L2, L9, SW loop L5) is 22%. According to [Iodice et al. \(2017b\)](#), the total luminosity from these loops is $1.3 \times 10^{10} L_\odot$ in the g -band. We stress that the IGL estimation is a lower limit as it is not possible to estimate the diffuse component which is in the line-of-sight of masked objects.

3.4. Products: Break radius, disc scale-length, and disc colour

We used the algorithm developed by [Raj et al. \(2019\)](#) to derive the break radius B_r , inner and outer scale-lengths (h_{in} and h_{out})⁶ from the deconvolved SB profiles in r -band. The ($g-i$) colours of the scale-lengths (h_{in} and h_{out}) are also derived from deconvolved SB profiles. The deconvolution method is briefly described in the following.

⁶ The disc scale-length refers to the region of the disc before and after the break radius.

Table 4. Parameters of LTGs with a disc break.

Object	B_r		σ_{nB_r}	μ_{B_r}	h_{in}		h_{out}		Break type	Hubble type (T)
	[arcsec]	[kpc]	[arcsec]	$[\frac{\text{mag}}{\text{arcsec}^2}]$	[arcsec]	$g - i$ [mag]	[arcsec]	$g - i$ [mag]		
(1)	(2)	(3)	(3)	(4)	(5)	(5)	(6)	(6)	(7)	(8)
FCC013(NGC 1310)	76.8	7.76	3.7	25.01	54.2	0.83 ± 0.03	43.7	1.27 ± 0.1	III	5
FCC022(NGC 1317)	–	–	–	–	–	–	–	–	I	1
FCC028	38.8	3.91	1.6	23.60	24.0	0.83 ± 0.1	60.8	0.87 ± 0.1	III	9
FCC029(NGC 1326)	53.6	5.41	1.5	21.82	24.76	0.92 ± 0.07	159.8	0.68 ± 0.03	III	1
FCC033(NGC 1316C)	–	–	–	–	–	–	–	–	I	7
FCC035	27.6	2.8	0.3	23.94	25.64	-0.57 ± 0.02	34.2	-0.93 ± 0.1	III	9
FCC037(NGC 1326A)	–	–	–	–	–	–	–	–	I	5
FCC039(NGC 1326B)	95.5	9.64	8.2	25.01	94.9	0.52 ± 0.02	37.0	0.88 ± 0.2	III	7
FCC062(NGC 1341)	52.3	5.28	2.3	24.03	38.9	0.78 ± 0.02	68.2	0.86 ± 0.03	III	4

Notes. Column 1 – LTGs with a disc break; Col. 2 – Break radius in units of arcsec and kpc (1 arcsec = 0.101 kpc); Col. 3 – Standard deviation of the break radii from $(n + 1)^3$ combinations; Col. 4 – Surface brightness at the break radius; Col. 5 – Inner scale-length in units of arcsec, and average $g - i$ colour for h_{in} ; Col. 6 – Outer scale-length in units of arcsec, and average $g - i$ colour for h_{out} ; Col. 7 – Profile classification; Col. 8 – Hubble type T taken from Drinkwater et al. (2001b).

In the process of detecting disc-breaks, we made sure to account for the effect of the point spread function (PSF). This is to avoid false detection of disc-breaks as a result of the scattered light from a galaxy’s bright core in the regions around them. We first measured the PSF out to the galaxy’s disc and then the PSF is characterised from the VST images so that the broadening effect of the seeing on the galaxies is taken into account (Capaccioli et al. 2015). Raj et al. (2019) show how the effect of the PSF can affect the SB profiles of galaxies of various sizes (see also Capaccioli et al. 2015). We then deconvolved galaxies with the PSF, using the Lucy-Richardson algorithm (Lucy 1974; Richardson 1972). We briefly explain the main steps of the algorithm used to derive the break radius.

1. The disc regions are determined from the 1D multi-component (bulge and disc) fits explained in Appendix C (the multi-component best fits are shown in Fig. C.2 and their corresponding best fit parameters are listed in Table C.1). The starting point is the intersection between the bulge and disc components. The regions are provided to the algorithm to define inner disc limit $range_{in}$ and outer disc limit $range_{out}$, within R_{lim} (see Table C.2).
2. The algorithm produces $(n + 1)^2$ best linear fits by varying $range_{in}$ and $range_{out}$, and $(n + 1)^3$ estimates of the break radius, where $n = 2$ in both cases.
3. The median of the $(n + 1)^3$ intersecting points is chosen as the final break radius B_r , and the standard deviation of the $(n + 1)^3$ estimates is given as sigma σ_{nB_r} on the estimation of the break radius.
4. Through iterations⁷ of the previous steps, a minimal standard deviation of the $(n + 1)^3$ break radii is obtained which in turn provides the best selection of the disc scale-lengths h_{in} and h_{out} .
5. The average $(g - i)$ colours of the inner and outer discs are derived for the regions h_{in} and h_{out} .

The break radius (in arcsec and kpc) of six galaxies along with σ_{nB_r} , surface brightness at the break μ_{B_r} , and average $(g - i)$ colour for the components h_{in} and h_{out} are listed in Table 4. Fitting results are shown in Fig. 2.

⁷ The number of iterations varies for each galaxy with a minimum of five until a minimal standard deviation of the $(n + 1)^3$ break radii is obtained.

4. Results

4.1. Morphology and colour analysis of galaxies in the Fornax A subgroup

In this section, we analyse the morphology, structure, colours, and stellar mass, as a function of group-centric distance, that is from BGG NGC 1316. This choice is motivated by making a direct comparison with the segregation of the galaxies’ properties found in the cluster core (Iodice et al. 2019a; Raj et al. 2019), as discussed in Sect. 7.

The Fornax A subgroup is populated with bright (<16 mag) LTGs and only one giant peculiar ETG, that is, NGC 1316 as its central galaxy. Most of the bright galaxies ($>75\%$) in the sample are located in the north-east (NE) side of the Fornax A group. In Fig. 3 (left panel), we plot the morphological stage parameter T as a function of projected distances from the group centre. This shows that the sample of LTGs is heterogeneous, similar to that of the LTGs in the Fornax cluster core (see right panel of Fig. 3 and Raj et al. 2019). Overall, we do not find any morphological segregation of galaxies in the group.

All LTGs in the Fornax A group have stellar masses in the range $8 < \log(M_*)[M_\odot] < 10.5$, uniformly distributed inside the group (see top-left panel of Fig. 4). As expected, from the colour-magnitude relation, the fainter galaxies ($M_i \leq -18$ mag) in the group have bluer colours, $g - i \leq 0.5$ mag (see lower left panel of Fig. 4).

ETGs and LTGs follow a different size-mass relation (see Shen et al. 2003) and this is likely the reason for the difference in the size-mass relation driven by the massive galaxy NGC 1316 and a less-pronounced relation observed for the bright spiral galaxies in the Fornax A subgroup (Fig. 4: top right panel). We do not find any correlation of effective radius as a function of the projected distance from NGC 1316 (Fig. 4: lower right panel). The effective radii of all bright spiral galaxies in the Fornax A subgroup are $R_e \leq 70$ arcsec (~ 7.07 kpc).

As found for the other properties derived for the LTGs in Fornax A group (i.e. morphology and stellar mass), the colour distribution also appears to be quite uniform inside the group, without any evident colour segregation with the group-centric distance (see Fig. 5). A slight, but not significant trend (Pearson’s correlation coefficient and the correlation significance are -0.41 and 0.26) might be observed in $(g - r)$ colours as

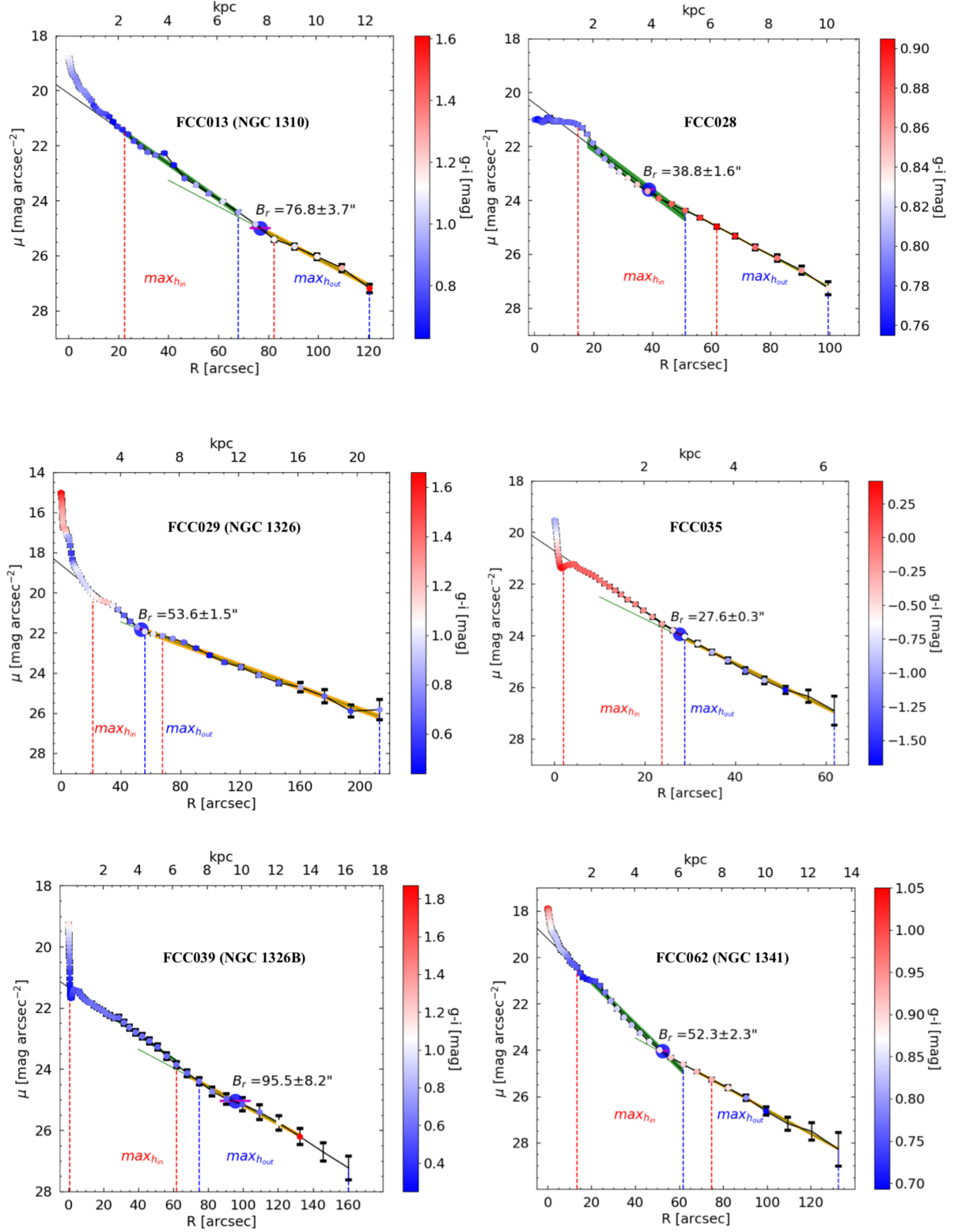


Fig. 2. Surface brightness profiles of galaxies with Type III disc break, with $(g - i)$ colour as a third parameter. In each plot, the break radius (B_r) is marked at the intersecting point of the linear fits between h_{in} and h_{out} , with σ_{nB_r} as the median of $(n + 1)^3$ combinations of best fits on the inner and outer scale-lengths. The vertical dashed lines ($max_{h_{in}}$ and $max_{h_{out}}$) indicate the regions for $max_{range_{in}} = range_{in} \pm n$ and $max_{range_{out}} = range_{out} \pm n$ where the algorithm produces $(n + 1)^2$ linear least square fits. The shaded regions on h_{in} and h_{out} indicate the median of the rms of the residuals for $(n + 1)^2$ linear least square fits.

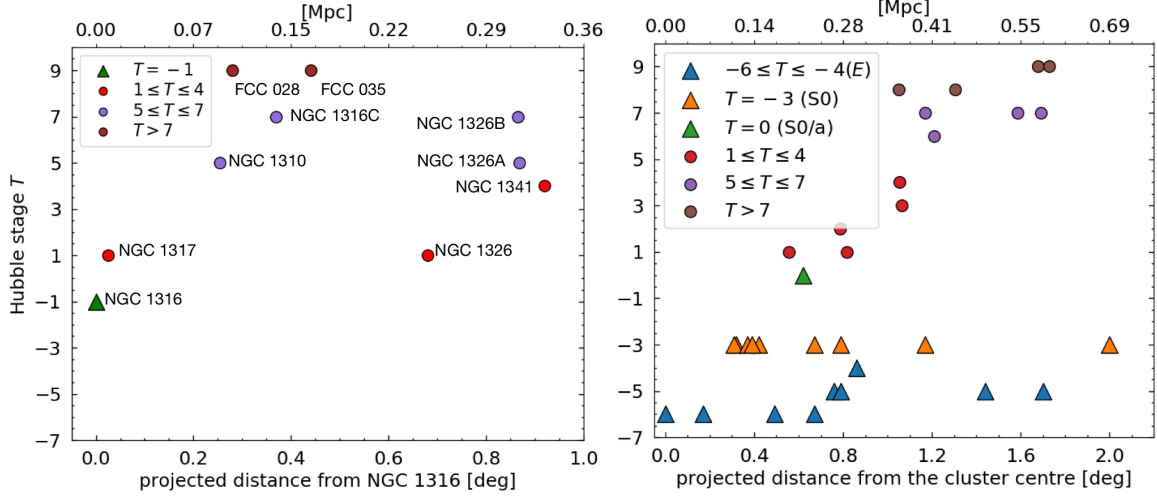


Fig. 3. Hubble type T as a function of group-centric (*left panel*) and cluster-centric distance (*right panel*) for bright galaxies in the Fornax A subgroup (*left panel*) and Fornax cluster core (*right panel*). Data for the cluster core galaxies are taken from Iodice et al. (2019a) for ETGs and Raj et al. (2019) for LTGs.

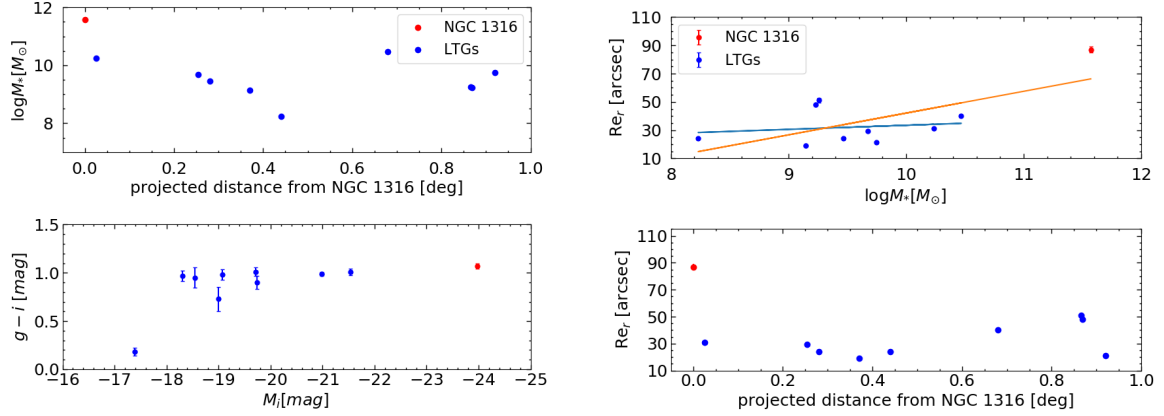


Fig. 4. *Left panels:* stellar mass as function of projected distance from NGC 1316 (*top panel*) and colour-magnitude relation (*lower panel*) for bright spiral galaxies inside the virial radius of the Fornax A subgroup. *Right panels:* size-mass relation (*top panel*) and effective radius as a function of projected distance from NGC 1316 (*lower panel*) for all the bright spiral galaxies in the Fornax A subgroup. A least square fit to the data is shown in orange (including NGC 1316) and blue (excluding NGC 1316).

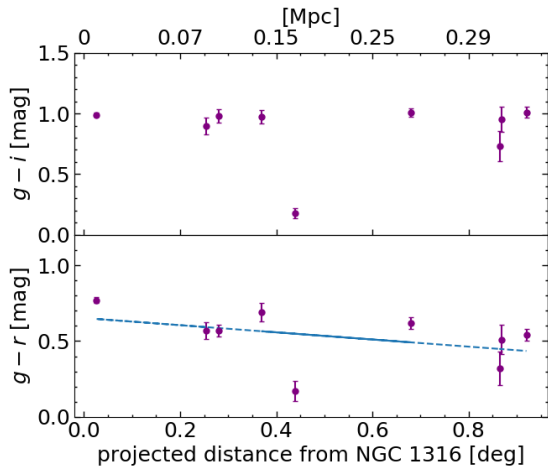


Fig. 5. $g - i$ (*top panel*) and $g - r$ (*lower panel*) colours as a function of projected distance from NGC 1316 in degrees and Mpc. A least square fit to the $g - r$ data is shown. Slope of the fit is 0.23 ± 0.15 .

a function of projected distance, whereas, on average, most ($\sim 90\%$) of the LTGs in the Fornax A subgroup have colours in the range $0.7 \leq g - i \leq 1$ mag. The very blue outlier FCC035 ($g - i = 0.17 \pm 0.1$ mag) has a nuclear star cluster in its central region and this contributes to its blue ($g - i$) colour.

4.2. Analysis of the discs

Similar to our analysis in Raj et al. (2019) for the galaxies in the virial radius of the Fornax cluster, we study the discs of the galaxies in the group as well. Our results of the disc-breaks show that six galaxies have Type III (up-bending) profiles. Even considering that this is a small sample, we examine the main correlations of the disc-breaks with other galaxy properties (e.g. stellar mass, total magnitude) inside the Fornax A group. Results are also compared with those derived for the LTGs in the Fornax cluster core by Raj et al. (2019).

We do not find any disc-break trends with regard to the stellar mass (see top left panel of Fig. 6). The same can be said for (i) total magnitude as a function of break radius normalised to the

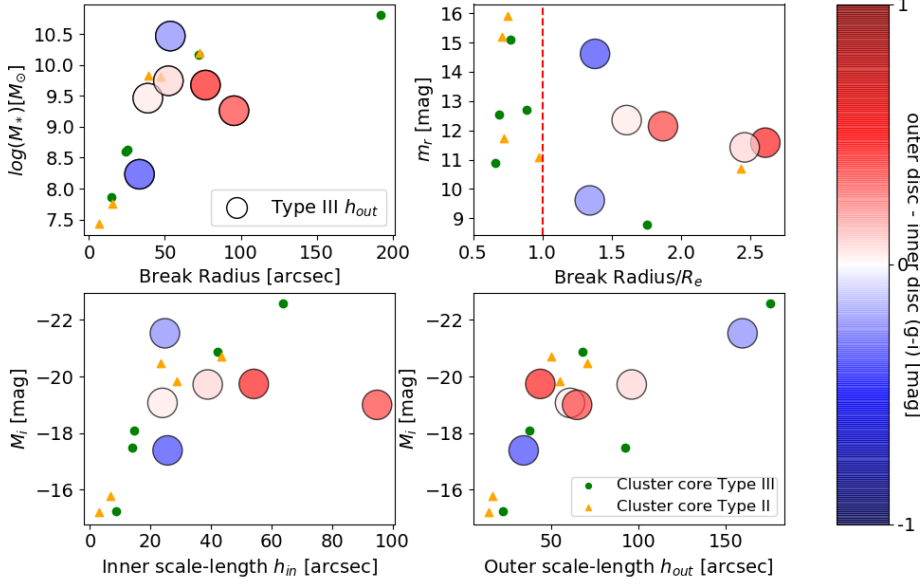


Fig. 6. Analysis of galaxies with a break in their surface brightness profiles. Circles represent Type III, with average $(g - i)$ colour of $h_{\text{out}} - h_{\text{in}}$ as colour map. Orange triangles and green circles represent Type II and Type III galaxies in the Fornax cluster core (data taken from Raj et al. 2019). *Top panel left:* stellar mass as a function of break radius. *Top panel right:* total magnitude as a function of break radius normalised to effective radius. The red dashed line indicates 80% of the Fornax cluster core galaxies with disc-breaks. *Bottom panels:* absolute magnitude in i -band as a function of inner scale-length h_{in} (left) and outer scale-length h_{out} (right).

effective radius in r (top right panel); (ii) absolute magnitude as a function of inner-scale length, h_{in} (lower left panel in Fig. 6). A less-pronounced relation between the outer scale-lengths and absolute magnitude is observed for the LTGs in the Fornax A subgroup (lower right panel in Fig. 6). Galaxies in the Fornax A group have disc-breaks located at different scale lengths, irrespective of their effective radius and total magnitude, unlike the galaxies in the cluster core (Raj et al. 2019) where there were observed trends such as, increasing stellar mass with B_r (see also cluster core LTGs in Fig. 6). For all six galaxies in the Fornax A group, the break radius is located beyond $1R_e$.

Except for NGC 1326 ($g - i_{h_{\text{out}} - h_{\text{in}}} = -0.24$ mag) and FCC 035 ($g - i_{h_{\text{out}} - h_{\text{in}}} = -0.36$ mag), the other four galaxies show redder outer discs. FCC 035 is a very late-type galaxy ($T = 9$) and is the faintest galaxy with minor disturbances in its outskirts (see Fig. B.6). NGC 1326 ($T = 1$) is the brightest in the LTG sample (see Fig. B.4) with a double bar (Erwin et al. 2005).

5. Assembly history of the bright LTGs in the Fornax A subgroup

In this section, we combine the main results for the brightest galaxies in the Fornax A group (morphology and disc structure) in the two-dimensional projected phase-space (PPS) to trace its assembly history. Recently, studies concerning the location of galaxies in phase-space have gained importance (e.g. Mahajan et al. 2011; Hernández-Fernández et al. 2014; Muzzin et al. 2014; Rhee et al. 2017). These studies show the connection between a galaxy’s phase-space location and its effect on star formation and stellar mass. This helps to further understand the effect of the environment on the evolution of galaxies, following on the morphology-density relation (Dressler 1980).

We used the published catalogue of spectroscopic redshifts by Maddox et al. (2019) to plot the phase-space diagram for 15 galaxies part of the Fornax A subgroup, including dwarfs. This is shown in Fig. 7. Maddox et al. (2019) use a sample of 40 galaxies within a radius of 0.7 Mpc and find the central velocity of the Fornax A subgroup to be 1778 km s^{-1} with a velocity dispersion of 204 km s^{-1} , and estimate the dynamical mass

of the Fornax A subgroup to be $1.6 \times 10^{13} M_{\odot}$. According to Maddox et al. (2019), we adopted 1.05 deg ($\sim 0.38 \text{ Mpc}$) as the virial radius of the Fornax A subgroup.

In Fig. 7, we plot the velocities (normalised to the group velocity dispersion) as a function of the projected distance (normalised to the virial radius). Along with this, we also plot the trajectory curve for the escape velocity⁸ (v_{esc}). According to Rhee et al. (2017), in the PPS we have also included the regions populated by the galaxies that merged into the cluster potential at different infalling times. They are the “ancient infallers” ($t < 8 \text{ Gyr}$), “intermediate infallers” ($t < 4\text{--}7 \text{ Gyr}$), and “recent infallers” ($t < 1 \text{ Gyr}$). Rhee et al. (2017) explain the origin of these regions by comparing them to the trajectory such that galaxies falling for the first time are close to the escape velocity curve, and as they pass through the pericentre, they start gaining velocity. Over time, after passing through the apocentre and many more pericentres, they eventually end up in the region of ancient infallers.

The PPS for the Fornax A group shows that the majority of the group members are located in the region of the recent and intermediate infallers. This would suggest that most galaxies are still assembling into the Fornax A group. The ancient infallers are the central and massive galaxy NGC 1316 and the close (in projection) barred spiral galaxy NGC 1317 (see also Fig. B.2). Most of the bright spiral galaxies in the Fornax A group (NGC 1310, NGC 1326A, FCC 035, and NGC 1341) seem to be intermediate infallers. Except for FCC035, galaxies in this region have average colour $0.9 < g - i < 1 \text{ mag}$. Although the Fornax A group is not as dense as the Fornax cluster, the group potential should somehow impact the outskirts of the discs. This is what we mainly observe for NGC 1316C, FCC 028, NGC 1326 and NGC 1326B (see Figs. B.3, B.5, B.7 and B.8) that belong to the region of recent infallers. NGC 1326, the most massive LTG in the sample, also shows plenty of star-forming regions.

Six of the galaxies in the region of intermediate and recent infallers also show a Type III disc-break. We explain the possible mechanisms for the formation of this break in Sect. 7. Finally,

⁸ The escape velocity $v_{\text{esc}} = \sqrt{\frac{2GM_{\text{vir}}}{R_{\text{vir}}} K(s)}$, where $K(s) = g_c \frac{\ln(1+cs)}{s}$, $s = \frac{R_{3D}}{R_{\text{proj}}}$ and $g_c = [\ln(1+c) - \frac{c}{1+c}]^{-1}$.

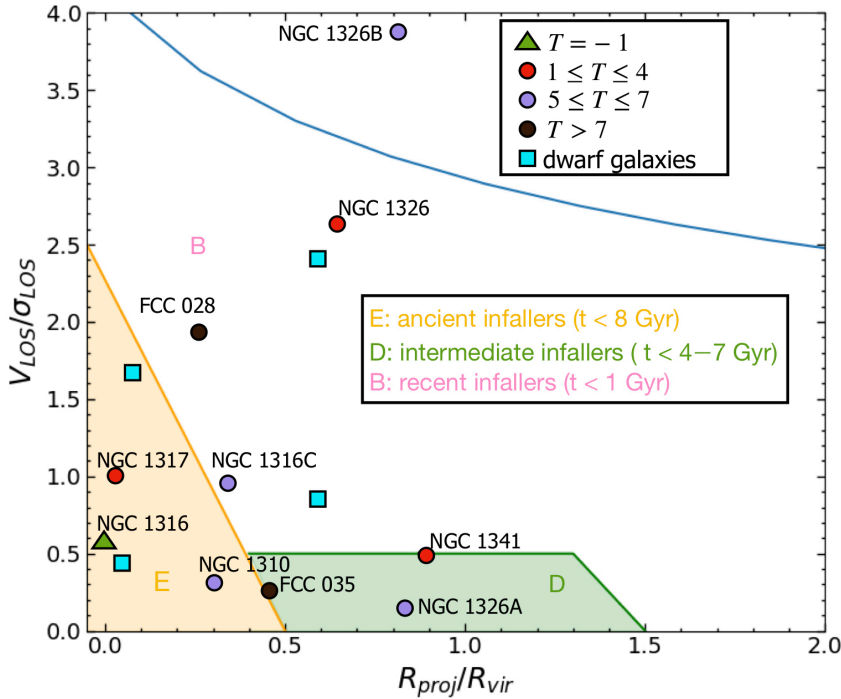


Fig. 7. Projected phase-space diagram of galaxies in the Fornax A subgroup within a group-centric distance of $1.05 \text{ deg} \sim 0.38 \text{ Mpc}$. The blue trajectory corresponds to the escape velocity of the group. Regions E (orange), D (green), B (pink) represent ancient, intermediate, and recent infallers as defined by Rhee et al. (2017). The bright spiral galaxies are indicated with coloured circles based on their morphological type T (red, purple, brown), NGC 1316 is indicated with a green triangle ($T = -1$) and dwarf galaxies are indicated with blue squares.

according to the PPS, the interacting system NGC 1326A and NGC 1326B may not be interacting, and that NGC 1326B is experiencing a fair amount of tidal interaction through its first infall, which would explain the star-forming regions in its outerdisc.

6. The build-up of the IGL in the Fornax A subgroup

In this section, we summarise relevant results from this work and those from previous works on the Fornax A group, which help to understand the formation history of its IGL. We then discuss how predictions from models fit in well with our results from observations.

The IGL in the Fornax A group is estimated within a region $\sim 35 \text{ arcmin}^2$ ($\sim 0.5 \text{ deg}$), which covers most of the group extension (see Sect. 3.3). Inside this radius, Venhola et al. (2019) find a significant drop in the number of dwarf galaxies ($\sim 50\%$). This might be an indication that the IGL in the Fornax A group was assembled by past merger of NGC 1316 (see Iodice et al. 2017b; Serra et al. 2019) and disruption of satellite galaxies on infall, which also contributed to the build-up of the extended stellar envelope around NGC 1316. This hypothesis is further supported by comparing the IGL fraction in the Fornax A subgroup with theoretical predictions from Contini et al. (2014) (see Fig. 8). According to Contini et al. (2014), galaxies which get destroyed while falling into the cluster or group make the accreted component of the ICL or IGL. Since the IGL in the Fornax A subgroup fits well in both regions from tidal disruption and merger channels, this suggests that at the present epoch, both mechanisms contributed to the build-up of this component.

In Fig. 8 we also compare the IGL derived for the Fornax A group with values obtained for other groups of galaxies, from observations of comparable depth and with the same image analysis. In particular, we plot the fraction of IGL to the total luminosity of the bright member galaxies as a function of the BGG stellar mass (in the range 10^{11} – $10^{12} M_{\odot}$) for four groups:

Fornax A (NGC 1316), NGC 5018 (Spavone et al. 2018), NGC 1533 (Cattapan et al. 2019), and IC 1459 (Iodice et al. 2020). Compared with theoretical predictions, the IGL fraction in each group reflects their assembly history. The main contribution to the IGL in the NGC 5018 group comes from the extended tidal tail that is tracing an ongoing merger (Spavone et al. 2018). In the Fornax A group, as well as in the NGC 1533 triplet (Cattapan et al. 2019), that have a comparable IGL fraction, mergers would have partly contributed to the IGL, and signs of past mergers are confined in the BGG’s envelope (in the form of loops and shells). As pointed out by Iodice et al. (2020), given the low IGL fraction ($\sim 2\%$) and the structure and properties of galaxy members, the IC 1459 group is still in an early mass assembly phase.

The connection between the IGL fraction and evolution is also consistent with the different ratios of ETG-to-LTG galaxies. In the Fornax A group, $\text{ETG/LTG} = 0.11$, since it has only one peculiar ETG (NGC 1316) and nine bright LTGs. NGC 5018 and NGC 1533 groups have $\text{ETG/LTG} = 1.6$, while IC 1549 $\text{ETGs/LTGs} = 0.3$. Although the Fornax A subgroup has the lowest ratio among the groups discussed above, a major fraction of the IGL with respect to the IC 1459 group, which has comparable ETG-to-LTG ratio, comes from the past merger of NGC 1316 and disrupted dwarf galaxies. In the IC 1459, no signs of interactions are found with a galaxy of comparable mass (Iodice et al. 2020).

7. The tale of the Fornax environment: The cluster core versus the infalling group

It is known that galaxies are redder and brighter in dense environments than in low-density environments (Faber 1973; Oemler 1974; Visvanathan & Sandage 1977; Butcher & Oemler 1984). The Fornax environment, comprising the cluster core and the Fornax A group, can depict “pre-processing” of galaxies, in particular, how LTGs evolve to the observed ETGs in the cluster

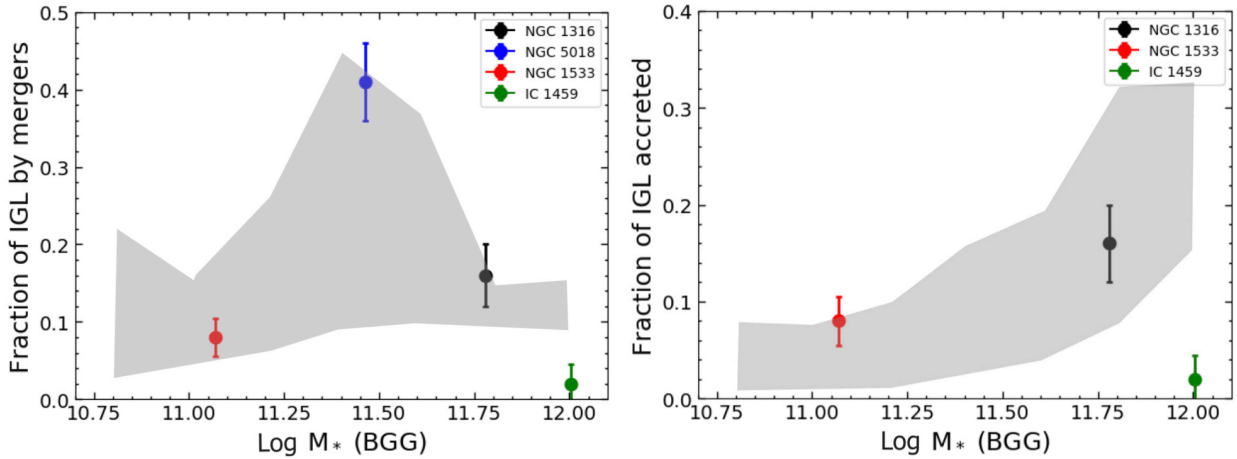


Fig. 8. Fraction of IGL by mergers (*left panel*) and fraction of IGL accreted (*right panel*), as a function of the BGG stellar mass for four groups: NGC 1316, NGC 5018 (Spavone et al. 2018), NGC 1533 (Cattapan et al. 2019), IC 1459 (Iodice et al. 2020). The shaded areas represent the 20th and 80th percentile of the “model disruption” described by Contini et al. (2014).

core. From literature, the proposed physical processes that galaxies undergo in groups and subclusters are strangulation and merging. In this section, we compare the properties of the group to that of the cluster core thus showing the different evolutionary stages for these environments.

Galaxy populations: The Fornax cluster core has 18 bright ($m_B < 15$ mag) ETGs (Iodice et al. 2019a) and 13 bright ($m_B < 16.6$ mag) LTGs (Raj et al. 2019). The total number of dwarf galaxies in the cluster is 564 (Venhola et al. 2018, 2019). However, the Fornax A group has nine bright ($m_B < 16$ mag) LTGs and one peculiar ETG, that is the central galaxy NGC 1316. The total number of dwarf galaxies is ~ 60 within $R_{\text{vir}} = 1.0$ deg (Venhola et al. 2019). Overall, the number of galaxies in the Fornax cluster core is nearly ten times more than that of the Fornax A group.

Morphological segregation: Nearly 70% of the ETGs ($T \leq -2$) are located within 0.8 deg of the cluster core (see Fig. 3). This region, where the bulk of the X-ray emission is also found, correspond to the high-density region of the cluster (Iodice et al. 2019a). Beyond this region, that is $D_{\text{core}} > 1$ deg, 60% of the LTGs are located (see Raj et al. 2019). These results are consistent with the morphology-density relation (Dressler 1980). We do not observe the same for the galaxies in the Fornax A group (see Fig. 3). This suggests that the environment is not as dense as the Fornax cluster core, where the late-type star-forming galaxies are slowly transforming into passively evolved cluster members (e.g. Oemler 1974). The absence of ETGs in the Fornax A group and the morphological segregation confirm that the group is at a different stage of evolution.

Colour segregation: For galaxies in the cluster core, a strong colour segregation as a function of the cluster-centric radius is observed, where the redder and massive ETGs populate the high-density region of the cluster (Iodice et al. 2019a). Most of the blue LTGs are found in the low-density region of the cluster. The Fornax A group, dominated by LTGs, does not show any colour segregation (see Fig. 5). The LTGs are about the same in $(g - i)$ colours as the LTGs in Fornax core (see left panel of Fig. 9). Figure 9 (right panel) also shows that there are LTGs with lower stellar mass in the Fornax core ($7 < \log(M_*)[M_\odot] < 8$). Hence the overall stellar mass range of LTGs, both in the core and the group is $7 < \log(M_*)[M_\odot] < 11$.

Structure of LTGs: The LTGs in the cluster core exhibit very different substructures that are well-correlated with their locations in the cluster core (see Raj et al. 2019). As for the LTGs in Fornax A group, their substructures are almost similar showing regular star-forming regions in their spiral arms. Although we find that *Hubble* type T is heterogeneous in both samples, LTGs in the core show more irregularities in their structures (e.g. lopsidedness, tidal tails, boxy discs) in comparison to those in Fornax A group, where almost none show any such irregularities (except for NGC 1327B). This suggests that the environment of the cluster core has played a major role in the structure-transformation of LTGs, while galaxies in the Fornax A subgroup are undergoing minor changes or have not undergone any changes yet.

Disc-breaks of LTGs: We find that many of the galaxies in the Fornax core have undergone pre-processing, star formation quenching in their outskirts during their time in the cluster (Zabel et al. 2019; Raj et al. 2019; Iodice et al. 2019b). This is also clearly seen through their disc-breaks, that is, the break radius of most (80%) of these galaxies is located within $1R_e$. Beyond this radius, there are no substructures (e.g. spiral arms) or star-forming regions. However, most of the disc-breaks seen in the group are of Type III, in contrast to those in the cluster core. Moreover, the break radius of galaxies in the Fornax A subgroup is located beyond $1R_e$. This suggests that the environmental processes in the group have not severely altered the structures of the galaxies therein.

Mergers and strangulation are common phenomena in the group environment (Barnes 1985; Zabludoff & Mulchaey 1998; Hashimoto & Oemler 2000), but this assumption is not enough to understand the origin of Type III disc-breaks alone. While they have been many explanations for their occurrence (e.g. Bakos et al. 2008; Younger et al. 2007; Laine et al. 2016; Watkins et al. 2019), we state a few which depend on the HI content, such that if the gas is not replenished, star formation is eventually halted. It is also important to consider internal mechanisms such as the formation of bars, which can also form a Type III disc-break (e.g. Erwin et al. 2005; Pohlen & Trujillo 2006). Wang et al. (2018) showed that Type III disc-breaks were prevalent in HI-rich galaxies with low-spin. The in-situ star formation from gas accretion builds the outer-disc of these Type III galaxies. Alternatively, enhanced star formation in the

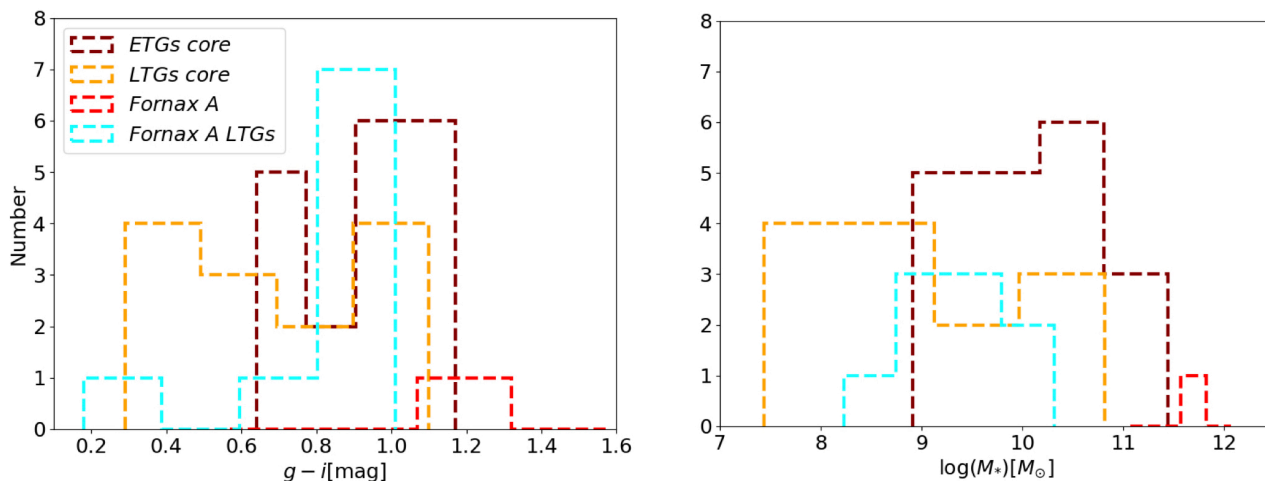


Fig. 9. Histogram of the colour ($g-i$) distribution (*left panel*) and of the stellar mass (*right panel*) for the bright galaxies in the Fornax core and Fornax A subgroup.

inner-discs of galaxies can cause steeper inner-discs than outer-discs (Hunter & Elmegreen 2006). Gas in outer-discs are vulnerable to the tidal forces in a dense environment. In this case, the gas that falls inwards can cause star-bursts in the inner central regions (e.g. Barnes & Hernquist 1991, 1996; Hopkins et al. 2013, and references therein). If LTGs in the Fornax A group are undergoing strangulation, their outer-discs would not have cold gas to fuel star formation. HI data are required to confirm if Type III disc-breaks in LTGs part of the Fornax A group, have formed as a consequence of HI-richness or the contrary.

X-ray emission: The main contributions to the X-ray emission in Fornax A group are provided by the centre of the BGG, NGC 1316, and by two of the most luminous stellar loops in the outskirts (L1 in the south and L2 on the east of the galaxy, Mackie & Fabbiano 1998). In addition, there is an elongated region of X-ray emission to the NW of NGC 1316, corresponding to the region of the plume (Iodice et al. 2017b). The total luminosity of the X-ray emission is $3.1 \times 10^{40} \text{ erg s}^{-1}$ (Kim & Fabbiano 2003) is an order of magnitude fainter than the total X-ray luminosity in the Fornax core which is $5 \times 10^{41} \text{ erg s}^{-1}$ (Jones et al. 1997). Moreover, the X-ray emission in the core of the cluster, which peaks on the BCG NGC 1399 is much more extended, reaching $\sim 0.4 R_{\text{vir}} \sim 0.3 \text{ Mpc}$. Most of the brightest and massive galaxies in the core are located inside the X-ray emission and all of them are ETGs (Iodice et al. 2019a). Therefore, most of the processes linked to the galaxy harassment that happened in this region of the cluster, and that shaped most of the observed galaxy properties, cannot be taken into account in the Fornax A group. This is further consistent with the homogeneous population of LTGs found in this group and with the overall picture that the Fornax A group is not as evolved as the cluster core.

8. Summary and conclusions

With FDS data, we have analysed the structure of the brightest ($m_B < 16 \text{ mag}$) LTGs in Fornax A SW group of the Fornax cluster. We have also taken advantage of the depth of the data and derived the IGL amount in the Fornax A group. In the study of disc-breaks of LTGs in different environments, we have compared the break radius of LTGs in the cluster core to those of the group. Since the Fornax cluster has been studied

by several surveys, we have made the best use of available public data, in this case, the spectroscopic redshift compilation by Maddox et al. (2019) to look at the group members in phase-space. Thus, for the first time, we give a complete picture of the environment of the Fornax A group, and compare it to that of the Fornax cluster core. We summarise the results of this work as follows:

1. In contrast to what is observed in the core of the Fornax cluster, in the Fornax A group, there is no observed morphological, mass and colour segregation of galaxies. LTGs in the Fornax A group have stellar masses in the range $8 < \log(M_*)[M_\odot] < 10.5$, while there are less massive LTGs in the Fornax cluster core ($7 < \log(M_*)[M_\odot] < 11$).
2. We detect only Type III disc-breaks in the SB profiles of galaxies in the Fornax A group. Of them, four galaxies have outer discs redder than their inner discs, and two galaxies have bluer outer discs. On the contrary, in the Fornax core, we detected all types of disc-breaks, suggesting a variety of mechanisms and interactions with the environment (Raj et al. 2019). The formation of the breaks in the Fornax A group needs more investigation of their HI content, in order to investigate whether HI-rich galaxies can cause enhanced star formation in the galaxy's outer-disc, or whether these galaxies are experiencing strangulation which halts star formation in the outer-discs.
3. The break radii of six galaxies in the Fornax A subgroup are located beyond $1R_e$. This is contrary to the results of the galaxies in the Fornax cluster core where the break radii of galaxies (80% of the sample) are located within $1R_e$. This suggests that the environmental mechanisms in the Fornax A group are not as strong as those in the cluster core, where the latter has altered the structures of the galaxies' disc.
4. Through phase-space analysis, we find that most (80%) of the group members are intermediate and recent infallers, consistent with a younger assembly epoch with respect to the Fornax core, where the PPS is equally populated by ancient infallers (that dominate the high-density region of the cluster) and by intermediate and recent infallers that are symmetrically distributed around the cluster core in the low-density region (Iodice et al. 2019b).
5. The luminosity of the IGL in the Fornax A group is estimated to be $6 \pm 2 \times 10^{10} L_\odot$ in g . Compared to the total

luminosity of the group, the fraction of the IGL is $\sim 16\%$. Our estimates agree with model predictions from Contini et al. (2014). Thus, we address the formation history of the IGL in the Fornax A group which is mainly through the past merger of NGC 1316 and the disruption of dwarf galaxies in the same region. This is further supported by the significant drop in the number of dwarf galaxies ($\sim 50\%$) towards the group centre found by Venhola et al. (2019).

According to the above results and the whole analysis performed in this work, the Fornax A group appears to be in an early-stage of assembly with respect to the cluster core. The environment of the Fornax A subgroup is not as dense as that of the cluster core, with all galaxies except the BGG showing similar morphology, comparable colours and stellar masses, and Type III disc-breaks, without any clear trend of these properties with group-centric distances. The low amount of IGL is also consistent with this picture, since there were no significant gravitational interactions between galaxies that modified the galaxy structure (forming tidal tails or stellar streams) and contributed to the build-up of the IGL. The main contribution to the IGL is from the minor merging in the outskirts of the BGG NGC 1316 and, probably, the disrupted dwarf galaxies close to the group centre.

As future perspectives involve the analysis of the star formation and, therefore, the comparison of it with the detected disc-breaks, we will study the $H\alpha$ distribution of the brightest LTGs in the Fornax A group (Raj et al., in prep.). A comparison of the overall structural parameters of galaxies in the cluster core and the group is currently analysed by Su et al. (in prep.). All of the mentioned work is part of FDS, thus providing a detailed analysis of the Fornax environment and galaxies therein.

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Appendix A: Bright LTGs in the Fornax A subgroup

In this section, we briefly describe the galaxies analysed in this work.

A.1. FCC013

Also known as NGC 1310, is a barred-spiral galaxy (SBcII) with morphological type $T = 5$, located at a projected distance $D_{\text{core}} = 0.25$ deg from NGC 1316. The HI disc of this galaxy is more extended than its stellar disc, with $M_{\text{HI}} = 48.0 \pm 0.9 \times 10^7 M_{\odot}$ (Serra et al. 2019). The galaxy has an effective radius $R_{e,r} = 27.6''$ and star-forming spiral arms. We detect a Type III disc-break in its SB profile and, hence, the break radius is 7.8 kpc. The outer disc appears redder than the inner disc with average $(g - i)_{h_{\text{out}}-h_{\text{in}}} = 0.44$ mag. The origin of the Type III break in this galaxy is likely due to its high HI content.

A.2. FCC022

Also known as NGC 1317, is a spiral galaxy (SA pec) with morphological type $T = 1$, located at closest projected distance from NGC 1316, that is, $D_{\text{core}} = 0.025$ deg from NGC 1316. The HI disc of the galaxy is confined to its stellar disc and has HI mass $M_{\text{HI}} = 26.4 \pm 0.6 \times 10^7 M_{\odot}$ (Horellou et al. 2001; Serra et al. 2019), while in the central part of this galaxy, that is, at $R \sim 0.5'$, a star-forming ring is visible. It appears to be an early-type spiral, with almost no star-forming regions beyond $R \sim 0.5'$, which can be seen in its SB image (Appendix B.2). This can be justified with the results from phase-space analysis that FCC022 has passed through the pericentre and lost its neutral gas through interactions with the intra-group medium (see also Serra et al. 2019).

A.3. FCC028

FCC028 is a spiral galaxy (SmIII) with morphological type $T = 9$, located at projected distance $D_{\text{core}} = 0.28$ deg from NGC 1316. From literature, the galaxy has been associated to a ring galaxy (Few & Madore 1986; Corwin et al. 1985; Lauberts 1982). The ring can be seen as star forming regions in the SB image B.3, which has likely originated from a collision (Maddox et al. 2019). The HI disc of the galaxy is more extended than its stellar disc and has HI mass $M_{\text{HI}} = 13.7 \pm 0.6 \times 10^7 M_{\odot}$ (Serra et al. 2019). We detect a Type III disc-break in its SB profile and its break radius is at 3.9 kpc. The outer disc of this galaxy is redder than its inner disc with an average $(g - i)_{h_{\text{out}}-h_{\text{in}}} = 0.04$ mag.

A.4. FCC029

Also known as NGC 1326, is a barred-spiral galaxy (SBa(r)) with morphological type $T = 1$ is located at a projected distance $D_{\text{core}} = 0.68$ deg from NGC 1316. Erwin (2004) describe it as a double barred galaxy. It is the most luminous, massive galaxy in our sample of LTGs with $M_i = -21.53$ and $M_* = 2.94 \times 10^{10} M_{\odot}$. We detect a Type III disc-break with a break radius = 5.4 kpc. The outer disc is bluer than its inner disc

with $(g - i)_{h_{\text{out}}-h_{\text{in}}} = -0.24$ mag. This can be associated to the star-forming regions in its outskirts.

A.5. FCC033

Also known as NGC 1316C, is a spiral galaxy (SdIII pec) with morphological type $T = 7$, located at a projected distance $D_{\text{core}} = 0.37$ deg from NGC 1316. The spiral arms are within $R = 10''$. The effective radius of this galaxy is $R_{e,g} = 22.6 \pm 0.5''$ and has a faint corona (Corwin et al. 1985).

A.6. FCC035

FCC 035 is a spiral galaxy (SmIV) with morphological type $T = 9$, and is located at a projected distance $D_{\text{core}} = 0.44$ deg from NGC 1316. It is the faintest galaxy in our sample $M_i = -17.39$ and $M_* = 0.017 \times 10^{10} M_{\odot}$. We detect a Type III disc break in its SB profile with break radius = 2.8 kpc. The outer disc is bluer than its inner disc with $(g - i)_{h_{\text{out}}-h_{\text{in}}} = -0.36$ mag. The $H\alpha$ (Raj et al., in prep.) distribution in the galaxy is concentrated in a star-forming region located in its spiral arms, which explains the blue $(g - i)$ colours. This can also be associated to its very-late morphological type and irregularities in its spiral structure.

A.7. FCC037

Also known as NGC 1326A, is a barred-spiral galaxy (SBcIII) with morphological type $T = 5$, located at a projected distance $D_{\text{core}} = 0.87$ deg. This galaxy has been studied as an interacting system along with FCC 039 (NGC 1327B). However, from our analysis in PPS (data taken from Maddox et al. 2019), we find the velocity difference between these galaxies is large and thus, they might not be interacting.

A.8. FCC039

Also known as NGC 1327B, is a spiral galaxy (SdIII) with morphological type $T = 7$, located at a projected distance $D_{\text{core}} = 0.86$ deg from NGC 1316. This galaxy has irregular spiral (dusty) arms, and disruptions in outer disc as a result of its recent interaction with the tidal field of the Fornax environment. We detect a Type III disc-break in its surface brightness profile and a break radius = 9.6 kpc. As mentioned in Appendix A.7, this galaxy might not be interacting with FCC 037 (NGC 1327A). However, it can be a fly-by thus showing faint tidal tails in its outskirts. The outer disc of this galaxy is redder than its inner disc with $(g - i)_{h_{\text{out}}-h_{\text{in}}} = 0.36$ mag.

A.9. FCC062

Also known as NGC 1341, is a barred-spiral galaxy (SbcII) with morphological type $T = 4$, located at furthest distance from NGC 1316, that is, $D_{\text{core}} = 0.92$ deg. We detect a Type III disc-break in its SB profile with disc break = 5.28 kpc. The outer disc of this galaxy is redder than its inner disc with spiral arms, Type III disc break with $(g - i)_{h_{\text{out}}-h_{\text{in}}} = 0.08$ mag.