



<b>Publication Year</b>	2018
<b>Acceptance in OA</b>	2020-12-09T15:41:21Z
<b>Title</b>	Optical polarisation variability of radio-loud narrow-line Seyfert 1 galaxies. Search for long rotations of the polarisation plane
<b>Authors</b>	Angelakis, E., Kiehlmann, S., Myserlis, I., Blinov, D., Eggen, J., Itoh, R., MARCHILI, Nicola, Zensus, J. A.
<b>Publisher's version (DOI)</b>	10.1051/0004-6361/201832890
<b>Handle</b>	<a href="http://hdl.handle.net/20.500.12386/28754">http://hdl.handle.net/20.500.12386/28754</a>
<b>Journal</b>	ASTRONOMY & ASTROPHYSICS
<b>Volume</b>	618

# Optical polarisation variability of radio-loud narrow-line Seyfert 1 galaxies<sup>★</sup>

## Search for long rotations of the polarisation plane

E. Angelakis<sup>1</sup>, S. Kiehlmann<sup>2</sup>, I. Myserlis<sup>1</sup>, D. Blinov<sup>3</sup>, J. Eggen<sup>4,5</sup>, R. Itoh<sup>6</sup>, N. Marchili<sup>7</sup>, and J. A. Zensus<sup>1</sup>

<sup>1</sup> Max-Planck-Institut für Radioastronomie, Auf dem Hügel 69, 53121 Bonn, Germany  
e-mail: [eangelakis@mpi-fr-bonn.mpg.de](mailto:eangelakis@mpi-fr-bonn.mpg.de)

<sup>2</sup> Owens Valley Radio Observatory, California Institute of Technology, Pasadena, CA 91125, USA

<sup>3</sup> Department of Physics and Institute of Theoretical & Computational Physics, University of Crete, 710 03 Heraklion, Greece

<sup>4</sup> Center for Research and Exploration in Space Science, NASA Goddard Space Flight Center, Greenbelt, MD 20771, USA

<sup>5</sup> Department of Astronomy, University of Maryland, College Park, MD 20742, USA

<sup>6</sup> Department of Physics, School of Science, Tokyo Institute of Technology, 2-12-1 Ohokayama, Meguro, Tokyo 152-8551, Japan

<sup>7</sup> Istituto di Astrofisica e Planetologia Spaziali Via Fosso del Cavaliere 100, 00133 Rome, Italy

Received 23 February 2018 / Accepted 5 July 2018

### ABSTRACT

**Context.** Narrow-line Seyfert 1 galaxies (NLSy1s) constitute the active galactic nuclei subclass associated with systematically lower black hole masses. A few radio-loud NLSy1s have been detected in MeV–GeV energy bands by *Fermi*, and evidence that blazar-like jets are operating also in radio-loud NLSy1s, has been accumulated.

**Aims.** We wish to quantify the temporal behaviour of the optical polarisation, fraction, and angle for a selected sample of radio-loud NLSy1s. We also search for rotations of the polarisation plane similar to those commonly observed in blazars.

**Methods.** We have conducted *R*-band optical linear polarisation monitoring of a sample of ten radio-loud NLSy1 galaxies; five of them have previously been detected by *Fermi*. The dataset obtained with our pivoting instrument, the RoboPol polarimeter of the Skinakas observatory, has been complemented with observations from the KANATA, Perkins, and Steward observatories. When evidence for long rotations of the polarisation plane was found (at least three consecutive measurements covering at least 90°), we carried out numerical simulations to assess the probability that they are caused by intrinsically evolving electric vector position angles (EVPAs) instead of observational noise.

**Results.** Even our moderately sampled sources show clear indications of variability in both polarisation fraction and angle. For the four best-sampled objects in our sample we find multiple periods of significant polarisation angle variability. Several of these events qualify as long rotations. In the two best-sampled cases, namely J1505+0326 and J0324+3410, we find indications for three long rotations of the polarisation angle. We show that although noise can induce the observed behaviour, it is much more likely that the apparent rotation is indeed caused by intrinsic evolution of the EVPA. To our knowledge, this is the very first detection of such events in this class of sources. In the largest dataset (J0324+3410), we find that the EVPA concentrates around a direction that is at 49.3° to the 15 GHz radio jet, implying a projected magnetic field at an angle of 40.7° to that axis.

**Conclusions.** We assess the probability that pure measurement uncertainties are the reason behind the observed long rotations of the polarisation plane. We conclude that although this is not improbable, it is much more likely that intrinsic rotations are responsible for the observed phenomenology. We conclude, however, that much better sampled and larger datasets of larger source samples are necessary to constrain the physical mechanism(s) that generate long EVPA rotations in NLSy1s.

**Key words.** galaxies: active – galaxies: Seyfert – polarization – techniques: polarimetric – methods: numerical – methods: statistical

## 1. Introduction

The term narrow-line Seyfert 1 galaxies (hereafter NLSy1s) signifies the subset of active galactic nuclei whose broad Balmer emission line has a width of  $FWHM(H\beta) \lesssim 2000 \text{ km s}^{-1}$ , and which have weak forbidden lines with  $[O III]\lambda 5007/H\beta < 3$  (Osterbrock & Pogge 1985; Goodrich 1989; Zhou et al. 2006). They are thus associated with black hole masses in the range  $10^6$ – $10^8 M_\odot$  (e.g. Xu et al. 2012; Foschini et al. 2015), which

is lower than those of powerful radio galaxies, which typically exceed  $10^8 M_\odot$ . Assuming these estimates to be free of biases (for claims of the opposite, see Marconi et al. 2008; Baldi et al. 2016), the detection of jet GeV emission (Abdo et al. 2009a,b; D’Ammando et al. 2012, 2015) and jet radio emission (e.g. Foschini et al. 2012; Angelakis et al. 2015; Lähteenmäki et al. 2017) from radio-loud (RL)<sup>1</sup> NLSy1s challenges the current understanding of relativistic jet formation, in which powerful relativistic jets are preferentially found in elliptical galaxies with nuclear black hole masses beyond  $10^8 M_\odot$  (for a review of the main arguments, see e.g. D’Ammando et al. 2017).

<sup>★</sup> Table 2 and the polarisation curves shown in figures 1, 8, 12, 14 are only available at the CDS via anonymous ftp to [cdsarc.u-strasbg.fr](http://cdsarc.u-strasbg.fr) (130.79.128.5) or via <http://cdsarc.u-strasbg.fr/viz-bin/qcat?J/A+A/618/A92>

<sup>1</sup> The radio loudness  $R$  is defined as the ratio of the 6 cm flux to the optical flux at 4400 Å (Kellermann et al. 1989).

Angelakis et al. (2015) presented a comprehensive study of the radio emission of the four RL NLSy1s detected by *Fermi*. The dataset covered the band from 2.64 GHz to 142.33 GHz at ten frequencies and with a cadence of less than 30 days (for three of the four sources). Despite the generally lower fluxes, all sources showed the typical characteristics seen in blazars: intense variability accompanied by dramatic spectral evolution, which is indicative of shocks operating in a plasma outflow. We computed limiting values of the brightness temperature and inferred rather moderate Doppler factors, implying the presence of mildly relativistic jets. The computed jet powers appeared comparable to the least energetic blazars, the BL Lac objects. In conclusion, the sources showed all the typical characteristics of blazars, only scaled to lower intensities. Fuhrmann et al. (2016) focused on the dynamics of the 15 GHz jet of 1H0323+342. Our analysis revealed superluminal components that are indicative of a relativistic jet, from which we inferred a viewing angle smaller than  $9^\circ$ , which confirmed the “aligned jet” scenario.

We currently focus on the optical polarisation of RL NLSy1s. Eggen (2012) and Eggen et al. (2013), among the first studies on the subject, reported that PMN J0948+0022 showed significant and variable polarisation in optical bands. The same source was observed by Itoh et al. (2013), who found minute timescale optical polarisation variability. During this “pulse”, the polarisation exceeded 30%, while interestingly, the polarisation angle, that is, the electric vector position angle (EVPA), appeared to be unchanged. The authors interpreted their findings as evidence of synchrotron emission radiated from a compact region of highly ordered magnetic field. In the case of J0849+5108, on the other hand, Maune et al. (2014) observed rapid intra-night variability in polarisation degree and angle during a major broadband outburst event that lasted for roughly five days. More recently, Itoh et al. (2014) studied 1H0323+342. They reported that the EVPA remained roughly parallel to the jet orientation, implying a magnetic field transverse to the jet axis.

Long rotations of the optical polarisation plane have been found in blazars (e.g. Kikuchi et al. 1988; Marscher et al. 2008, 2010; Abdo et al. 2010; Blinov et al. 2016a). Models that have been put forth to interpret the observations include physical rotation of emission elements on a helical trajectory (Marscher et al. 2008), propagation in large-scale bent jet (Abdo et al. 2010), turbulent plasma processes resulting in random walks (Marscher 2014), or light travel-time effects within an axisymmetric emission region (Zhang et al. 2015). Interestingly, it has been argued that these physical processes are likely associated with increased episodic gamma-ray activity (Blinov et al. 2018).

Beyond the potential of using polarisation monitoring to probe the physical processes at the emission site, the variability of the EVPA in particular can further our understanding of the conditions present during the high-energy jet emission production. In this context, we wish to (a) quantify the variability of the *R*-band optical polarisation fraction and angle for a selected sample of RL NLSy1 galaxies, (b) examine whether long rotations of the polarisation plane occur in RL NLSy1s, (c) parametrise these rotations and examine their association with the high-energy activity, and (d) ultimately understand the physical mechanisms producing them.

Here we present a study of a sample of ten RL NLSy1s; five of them have been detected by *Fermi* (cf. Sect. 2 and Table 1). When the datasets allowed, we studied the variability of both polarisation parameters and searched for EVPA rotation candidates. For the two best-sampled cases, we assessed the probability of these events being driven by intrinsic EVPA variability rather

than observational noise. This distinction was accomplished by conducting exhaustive simulations. In the following, we emphasise both our findings and the method for assessing the probability itself. In these two cases we indeed find evidence that the rotations are intrinsic to each source. This is the first time that such events are reported for RL NLSy1s.

## 2. Source sample and dataset

The selection of our sample has been based mostly on the radio loudness and observability of the sources from the Skinakas telescope (i.e. optical magnitude and position). It includes five of the eight sources that have been reported to radiate significant emission in the MeV–GeV energy range (Abdo et al. 2009a,b; D’Ammando et al. 2012, 2015; Yao et al. 2015; Liao et al. 2015). It also includes another five RL sources, which together with the first five, make up a total of ten targets. All the sources with at least one data point from our monitoring are listed in Table 1. Median values and ranges of the polarisation parameters for all the sources discussed here are shown in Table 3.

### 2.1. RoboPol dataset

The RoboPol<sup>2</sup> dataset is the basis for this study. The instrument is mounted on the 1.3 m telescope of the Skinakas observatory (Papamastorakis 2007) and has been monitoring our sample in the *R* band. All the details of the measurement techniques and the instrument characteristics are discussed in King et al. (2014) and Angelakis et al. (2016), where post-measurement quality criteria are discussed in depth.

### 2.2. KANATA dataset

The KANATA observations were conducted with the 1.5 m telescope of Higashi-Hiroshima Observatory. The polarimetry was performed with the HOWPol polarimeter (Kawabata et al. 2008). The observing cycle includes successive exposures at four position angles of a half-wave plate at 0, 45, 22.5, and 67.5°. The instrumental polarisation (peaking at ~4%) was modelled and removed before further analysis. The residual uncertainties are estimated from large numbers of unpolarised standard stars and are smaller than 0.5%.

### 2.3. Perkins dataset

The Perkins dataset was obtained with a Johnson *R* filter using the PRISM instrument on the 1.8 m Perkins Telescope of the Lowell Observatory, which also includes a rotating half-wave plate polarimeter. The observing cycle included exposures with the half-wave plate at 0, 45, 90, and 135°. The averages of two to four such cycles were used as the final measurement. Instrumental offsets of the EVPA and percent polarisation (usually less than 1%) were determined by observing in-field polarised and unpolarised standard stars (Schmidt et al. 1992).

### 2.4. Steward Observatory dataset

The Steward Observatory data have also been obtained at *R* band, and they have been retrieved from the online archive<sup>3</sup>. The data acquisition and reduction is described in Smith et al. (2009).

<sup>2</sup> <http://robopol.org>

<sup>3</sup> <http://james.as.arizona.edu/~psmith/Fermi/>

**Table 1.** List of sources in our sample and their relevant parameters.

ID	Survey ID	Redshift	$M_{\text{BH}}$	$R$	Notes
J0324+3410	1H 0323+342	0.062900 <sup>1</sup>	$2 - 3.4 \times 10^7$ A,O,P	318 <sup>N</sup>	Fermi detected <sup>5</sup>
J0849+5108	SBS 0846+513	0.584701 <sup>2</sup>	$0.8 - 9.8 \times 10^7$ B,C,D	1445 <sup>G</sup>	Fermi detected <sup>6</sup>
J0948+0022	PMN J0948+0022	0.585102 <sup>2</sup>	$0.2 - 8.1 \times 10^8$ E,F,10	355 <sup>G,10</sup>	Fermi detected <sup>7</sup>
J1305+5116	WISE J130522.75+511640.3	0.787552 <sup>2</sup>	$3.2 \times 10^8$ G	223 <sup>G</sup>	Optical spec. indicates strong outflow <sup>11</sup>
J1505+0326	PKS 1502+036	0.407882 <sup>2</sup>	$0.04 - 2 \times 10^8$ G,H,5,1,11	1549 <sup>G</sup>	Fermi detected <sup>5</sup>
J1548+3511	HB89 1546+353	0.479014 <sup>2</sup>	$7.9 \times 10^7$ G	692 <sup>G</sup>	Evidence for past radio variability
J1628+4007	RX J16290+4007	0.272486 <sup>2</sup>	$3.5 \times 10^7$ K,10	29 <sup>M,10</sup>	High optical and radio variability <sup>10</sup>
J1633+4718	RX J1633.3+4718	0.116030 <sup>4</sup>	$3 \times 10^6$ J	166 <sup>G</sup>	Evidence for past radio variability
J1644+2619	FBQS J1644+2619	0.145000 <sup>3</sup>	$2.1 \times 10^8$ L	447 <sup>M</sup>	Fermi detected <sup>8</sup>
J1722+5654	SDSS J172206.02+565451.6	0.425967 <sup>2</sup>	$2.5 - 3.3 \times 10^7$ G,9	234 <sup>G,9</sup>	Evidence for high-amplitude optical variability <sup>9</sup>

**Notes.** Columns: (1) Source identifier, (2) survey identifier, (3) redshift, (4) black hole mass, (5) radio loudness. (1) Zhou et al. (2007); (2) Hewett & Wild (2010); (3) Foschini et al. (2015); (4) Oh et al. (2015); (5) Abdo et al. (2009b); (6) D’Ammando et al. (2012); (7) Abdo et al. (2009a); (8) D’Ammando et al. (2015); (9) Komossa et al. (2006a); (10) Komossa et al. (2006b); (11) Komossa et al. (2016); (A) Landt et al. (2017); (B) Zhou et al. (2005); (C) Shen et al. (2011); (D) Paliya et al. (2016); (E) Zhou et al. (2003); (F) Abdo et al. (2009a); (G) Yuan et al. (2008); (H) Paliya & Stalin (2016); (I) Calderone et al. (2013); (J) Yuan et al. (2010); (K) Foschini et al. (2015); (L) D’Ammando et al. (2017); (M) Doi et al. (2016); (N) Foschini (2011); (O) Wang et al. (2016); (P) Kynoch et al. (2018).

### 3. Rice bias treatment

The functional dependence of the polarisation fraction  $p$  on the normalised Stokes parameters  $q$  and  $u$  introduces a bias in its determination from repeated observations in the presence of noise. The effect becomes particularly important at low signal-to-noise ratios (S/Ns). Here we study the temporal behaviour of the EVPA from observations of already moderate sampling. To avoid unnecessary data loss that could be imposed by using only high-significance data, we used all available observations after we corrected for the Rice bias.

We assumed a fixed polarisation vector with real amplitude  $p_0$  at an angle of  $\chi_0$ , observed in the presence of experimental noise. Unless the S/N is high, none of the observed polarisation parameters  $p$  and  $\chi$  determined from repeated observations will follow a normal distribution, even though  $\chi$  will populate a distribution symmetric around  $\chi_0$ .

Gaussian noise causes  $q$  and  $u$  to be normally distributed about their true values  $q_0$  and  $u_0$ , respectively, and with equal uncertainties  $\sigma_q = \sigma_u = \sigma_{q_0} = \sigma_{u_0}$ , which are also equal to the uncertainties in  $p$ ,  $\sigma_p$ . However, the probability of measuring polarisation in the range  $[p, p + dp]$  independent of polarisation angle (integrating over all angles), as was first demonstrated by Serkowski (1958), is given by the Rice distribution (Rice 1945),

$$F(p|p_0)dp = \frac{p}{\sigma_p} \exp\left[-\frac{p^2 + p_0^2}{2\sigma_p^2}\right] I_0\left(\frac{pp_0}{\sigma_p^2}\right) \frac{dp}{\sigma_p}, \quad (1)$$

with  $I_0$  the zeroth-order modified Bessel function. The asymmetry of Eq. (1) with respect to  $p$  and  $p_0$  is the cause for the observed polarisation bias especially at low S/Ns. At high S/Ns, the Rice distribution tends to a normal distribution with a mean around the true value of polarisation  $p_0$  and a spread  $\sigma_p$ . Vinokur (1965), Simmons & Stewart (1985) and Naghizadeh-Khouei & Clarke (1993) have investigated the distributions of the observed amplitudes and the angles from repeated observation and thorough descriptions of the problem can be found in Wardle & Kronberg (1974), Vaillancourt (2006) and Clarke (2010).

We adopted the same approach as in Pavlidou et al. (2014) to determine the polarisation amplitude. As a best-guess of  $p_0$ ,

we took the approximation of the maximum-likelihood estimator  $\hat{p}$  given by Vaillancourt (2006):

$$\hat{p} = \begin{cases} 0 & \text{for } p/\sigma_p < \sqrt{2} \\ \sqrt{p^2 - \sigma_p^2} & \text{for } p/\sigma_p \geq \sqrt{2} \end{cases}. \quad (2)$$

The uncertainty in the debiased polarisation fraction is set to that of the observed  $p$ ,  $\sigma_p$  as long as it is bounded at zero.

Because we are interested in the occurrence of long rotations of the polarisation plane, it is particularly important to assess the uncertainty in the polarisation angle even in cases of low S/N. To do this, we adopted the approach presented by Naghizadeh-Khouei & Clarke (1993). We solved their Eq. (4) for  $\sigma_\theta$  with the S/N of the debiased value ( $\hat{p}/\sigma_p$ ).  $\sigma_\theta$  was then taken as the uncertainty in the angle  $\sigma_\chi$ .

### 4. Long rotations of the polarisation plane

In Sect. 3 we discussed the treatment of the polarisation fraction. Here we clarify the conventions and terminology used for the EVPA.

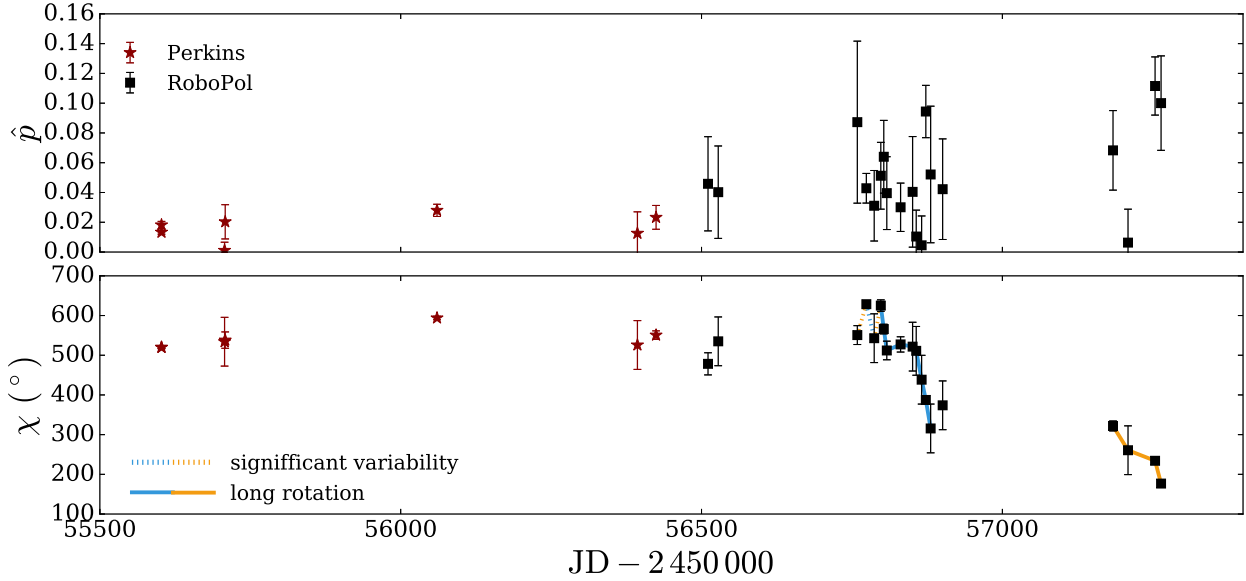
The EVPAs were initially computed from the observed  $q$  and  $u$  as

$$\chi_{\text{obs}} = \frac{1}{2} \times \arctan\left(\frac{u}{q}\right) \quad (3)$$

and hence carry the inherent “ $n \times \pi$ ” ambiguity. For each  $(q, u)$  pair, we chose the solution of Eq. (3) for which the difference from the previous data point is smaller than  $90^\circ$ .

The resulting data points were then termed “adjusted” and are the values we use in the following discussion. Thus, the adjustment is made under the assumption of minimum variability between adjacent data points.

Phases of significant EVPA variability are defined as sequences of data points over which the adjusted polarisation angle  $\chi$  changes significantly between consecutive observations ( $|\Delta\chi| > \sqrt{\sigma_{\chi_i}^2 + \sigma_{\chi_{i+1}}^2}$ ). Such periods are marked in the plots of Fig. 1 by dotted coloured lines. Lines of the same colour connect data points that show an overall trend in the same direction. Over such periods, insignificant changes in the opposite direction are



**Fig. 1.** J1505+0326: Debiased polarisation fraction  $\hat{p}$  and adjusted EVPA ( $\chi$ ) as a function of time. The coloured lines mark periods of significant monotonous (within the uncertainties) EVPA evolution. Solid lines mark periods of long rotations (i.e. at least three sequential data points and an angle larger than  $90^\circ$ ). Blue and orange connecting lines are used alternately for clarity.

allowed. This approach is described in detail in Kiehlmann et al. (2016). We defined such periods as *long* EVPA rotations (i.e. long rotations of the polarisation plane), when they

1. consisted of at least three data points, and
2. exceeded  $90^\circ$ .

Such periods are marked with solid coloured lines. A rotation was terminated when significant variability showed a change in sign.

## 5. Analysis of the polarisation fraction and angle variability

Here, we study the temporal behaviour of the optical polarisation (fraction and angle) for sources whose datasets are large enough. We list all the RoboPol measurements of the remaining sources in Table 2. Median values and ranges of the polarisation parameters for all the sources discussed here are shown in Table 3.

### 5.1. J1505+0326

We start with the source J1505+0326 because (a) it shows clearly discernible events that facilitate implementing our approach, and (b) because our analysis shows that it is the best candidate to have undergone an intrinsic long rotation of the polarisation plane.

Figure 1 presents the RoboPol and Perkins observations of the optical linear polarisation parameters  $\hat{p}$  and  $\chi$ .

The coloured lines there (dotted or solid) mark five periods of significant, continuous EVPA variability that we have detected. The absolute rotation angles of those events in order of occurrence are  $77.9$ ,  $85.7$ ,  $82.2$ ,  $309.5$ , and  $145.1^\circ$ .

The first three events consist of only two consecutive data points (dotted lines). The last two, however, consist of at least three sequential data points and exceed the limit of  $90^\circ$ ; they thus qualify as long-rotation candidates (solid lines).

The polarisation fraction spreads around a median of around  $0.04$  with a standard deviation  $\sim 0.03$  (Table 3). Figure 2 shows its cumulative distribution function for all the measurements

(dotted black line), during phases of rotation (solid blue line), and during non-rotating phases (dashed blue line). The median,  $\hat{p}$ , of the non-rotating phases alone is  $0.022$ , while that over rotating phases is  $0.043$ . Despite the difference, a two-sample Kolmogorov–Smirnov (KS) test, however, gave no indication for different parent populations, which in turn prohibits any conclusion about the behaviour of the polarisation during the rotations.

In the following we focus on the largest long rotation and assess the probability that it is driven by an intrinsically rotating polarisation plane.

#### 5.1.1. J1505+0326: Long rotation

Figure 3 zooms in on the largest of the potential rotations (MJD  $56790$ – $56880$ ). Its change in angle  $\Delta\chi$  is  $-309.5^\circ$ , and it lasts for approximately 83 days, yielding a mean rotation rate of  $-3.7 \text{ deg d}^{-1}$ . However, the combination of sparse sampling and large uncertainties in the angle makes the estimate of the direction of a rotation highly uncertain, which in turn makes the detection of the rotation itself uncertain.

As we discussed in Sect. 4, for each pair  $(q, u)$ , we chose the solution of Eq. (3) for which the absolute difference,  $|\Delta\chi|$ , from the previous angle is less than  $90^\circ$ . This condition controls the direction of the EVPA evolution. However, the uncertainty associated with each angle computation must also be accounted for when this condition is checked. If the sum of the absolute difference  $|\Delta\chi|$  between two consecutive data points and the uncertainty in that difference  $\sigma_{\Delta\chi}$  exceeds  $90^\circ$ , that is,  $90^\circ \leq |\Delta\chi| + \sigma_{\Delta\chi}$ , the direction of the rotation becomes uncertain because both solutions of Eq. (3),  $\chi$  and  $\chi + \pi$  could be valid<sup>4</sup>.

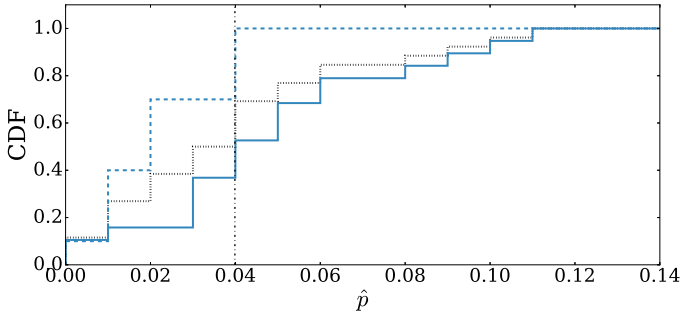
With the exception of the earliest measurement (left-most point), each angle measurement (solid symbols) in Fig. 3 is paired with its  $180^\circ$  conjugate (empty symbols).

<sup>4</sup> Clearly, in the absence of physical constraints, any solution of the form  $n \times \pi$  is equally valid. The choice of the smallest step is justified by the assumption of minimum variability.

**Table 3.** Integrated polarisation characteristics.

Source	N	$\langle \hat{p} \rangle$	$\sigma_p$ ( $^\circ$ )	$\langle \chi \rangle$ ( $^\circ$ )	$\chi_{\min}$ ( $^\circ$ )	$\chi_{\max}$ ( $^\circ$ )
J0324	115	0.012	0.016	-6.7	-89.1	+87.0
J0849	15	0.100	0.078	33.3	+6.8	+62.6
J0948	30	0.024	0.028	9.0	-83.0	+79.7
J1505	26	0.040	0.030	-1.9	-61.7	+88.6
J1305+5116	2	0.010	0.002	-19.3	-29.6	-8.9
J1548+3511	3	0.021	0.024	-10.8	-32.9	+47.9
J1628+4007	2	0.000	...	...	...	...
J1633+4718	4	0.024	0.004	-4.4	-8.1	+3.7
J1644+2619	4	0.022	0.015	-28.4	-34.8	-16.8
J1722+5654	2	0.000	...	...	...	...

**Notes.** Columns: (1) Source ID, (2) number of available measurements, (3) median debiased polarisation fraction, (4) spread of polarisation fraction, (5) median polarisation angle, (6), (7) minimum and maximum polarisation angle in the range  $[-90^\circ, 90^\circ]$ .



**Fig. 2.** J1505+0326: Distribution of the debiased polarisation fraction  $p$ . The dot-dashed vertical line marks the median of the distribution. The blue dashed line shows the distribution of  $\hat{p}$  during the non-rotating phases and the solid line shows this during the rotating phases. The black dotted line corresponds to all the measurements.

The critical steps with  $90^\circ \leq |\Delta\chi| + \sigma_{\Delta\chi}$  are shown in red. Clearly, their number prevents us from reliably telling the direction that the EVPA follows, making the detection of the rotation uncertain.

The uncertainty in the detected rotation can also be shown by examining the effect of the uncertainties in  $q$  and  $u$  on the rotation angle. For simplicity, we assumed that the measured  $q$  and  $u$  are the means of the Gaussian distributed fractional Stokes parameters which is equivalent to saying that they describe the “real” intrinsic behaviour of the source.

We then added Gaussian noise based on their uncertainties, recalculated the EVPA curve, and compared its parameters with those of the observed one.

In Fig. 4 we show the distribution of rotation angles,  $\Delta\chi$ , for a total of  $10^4$  simulated light curves. On the basis of our assumptions, the probability of detecting a rotation with an absolute angle  $|\Delta\chi|$  within  $1\sigma$  of the observed value is approximately 0.22. For larger rotations ( $|\Delta\chi| \geq 309.5^\circ$ ) the probability is around 0.081.

From this we conclude that we cannot be confident about the intrinsic evolution of the EVPA. Even if we knew the intrinsic variability, the limited sampling and the measurement uncertainties would allow a vast range of possible EVPA curves that would result in varying changes of the EVPA (i.e.  $\Delta\chi$ ).

Subsequently, the previous test can determine the most likely observation, but it cannot tell us anything about the intrinsic variability. For example, although the bin with the highest probability appears at about  $-140^\circ$ , this does not imply that the intrinsic EVPA rotation covers, most probably,  $140^\circ$ .

Finally, the data points in Fig. 3 could be aligned with roughly the same rotation rate if 180-degree shifts were chosen accordingly instead of obeying the convention of smallest change between consecutive measurements. This would result in a rotation  $360^\circ$  larger than shown in Fig. 3.

Clearly, the measurement of an intrinsic rotation is limited by the sparse sampling and the  $180^\circ$  ambiguity. In order to assess the reliability of the observed event, we took two steps.

1. We first estimated the probability that the measurement uncertainties induce a fake rotation in the absence of a real one, and
2. We then estimated the likelihood of an intrinsic rotation given the observed data.

The following simulations used exactly the same time sampling as the data and thus are affected by the  $180^\circ$  ambiguity in the same way.

### 5.1.2. Is the observed rotation an artefact of noise?

Here we assess the probability of the observational noise inducing the apparent rotation in the absence of an intrinsic rotation; that is, assuming  $d\chi_{\text{intr}}/dt = 0 \text{ deg d}^{-1}$ .

For simplicity, we set  $q$  to the mean polarisation fraction in our simulations during the observed rotation  $\bar{p}_{\text{rot}}$  and  $u$  to zero, which results in  $\chi = 0^\circ$ . Subsequently, we added Gaussian noise  $\mathcal{N}$  to these values according to the estimated uncertainties:

$$q = \bar{p}_{\text{rot}} + \mathcal{N}(0, \sigma_q) \quad (4)$$

$$u = \mathcal{N}(0, \sigma_u). \quad (5)$$

$\mathcal{N}(0, \sigma)$  denotes that the noise is centred at 0. We ran  $10^4$  simulations. For each run the same algorithm as used for the observed data was used to identify full rotations. We define as “full rotation” in our simulations a rotation that consists of as many data points as the observed long rotations. The probability of finding a full rotation is

$$P(\text{full rotation} | d\chi_{\text{intr}}/dt = 0) = 2.7 \times 10^{-2}. \quad (6)$$

We also find that

$$P(|\Delta\chi_{\text{intr}}| \geq 309.5^\circ | d\chi_{\text{intr}}/dt = 0) = 10^{-3} \quad (7)$$

and

$$P(\text{full rotation}; |\Delta\chi_{\text{intr}}| \geq 309.5^\circ | d\chi_{\text{intr}}/dt = 0) = 6 \times 10^{-4} \quad (8)$$

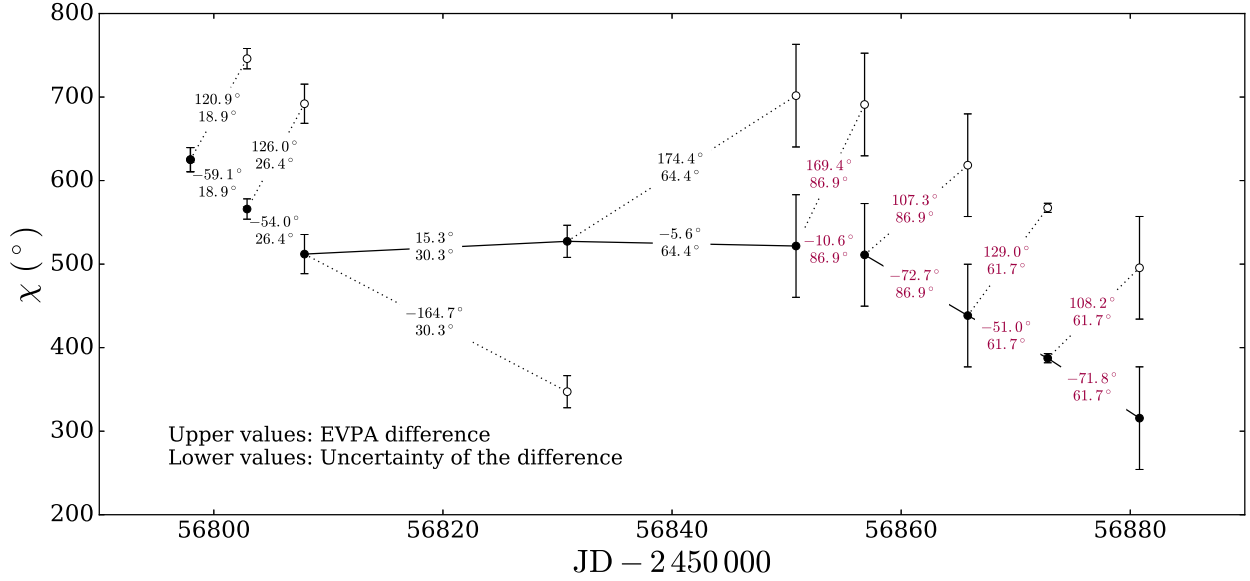
In Fig. 5 we show the results of the simulations.

This exercise shows that although it is not impossible that the observed event is merely an artefact of noise, it is fairly improbable. Thus, there must be intrinsic variability even if we cannot be sure of its exact nature.

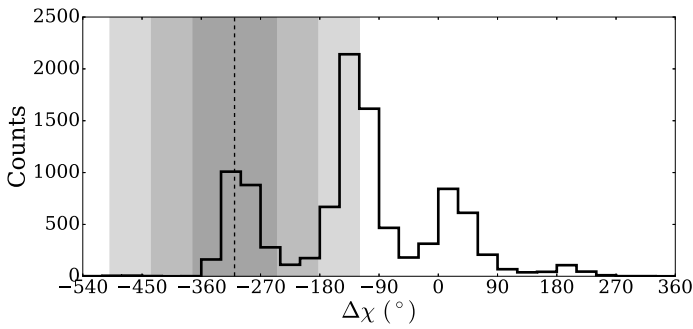
### 5.1.3. Most probable parameters of the intrinsic event

Having shown that intrinsic variability seems much more likely to be driving the observed EVPA behaviour, we wish to estimate the most probable parameters of the potential intrinsic rotation.

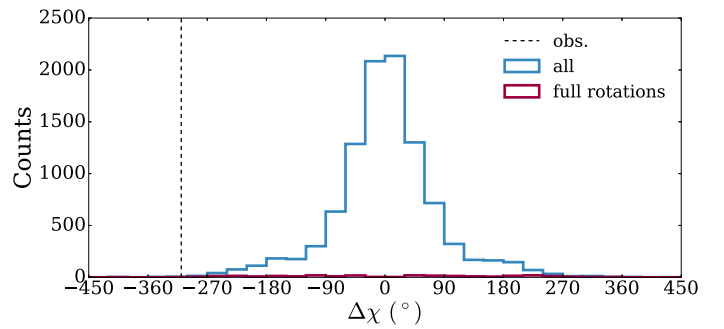
Our analysis relies on the assumption of a constant intrinsic rotation rate  $d\chi_{\text{intr}}/dt$  as well as a constant polarisation fraction



**Fig. 3.** J1505+0326: Apparent long rotation. With the exception of the left-most point, each angle measurement (solid symbols) is paired with its  $180^\circ$  conjugate (empty symbols). The upper values note EVPA differences ( $\Delta\chi$ ) and the lower values their uncertainties ( $\sigma_{\Delta\chi}$ ). Red marks highlight points in which the uncertainty  $\sigma_{\Delta\chi}$  in  $\Delta\chi$  is so large that both solutions of Eq. (3),  $\chi$  and  $\chi + \pi$ , could be valid, making the direction of the rotation uncertain.



**Fig. 4.** J1505+0326: Distribution of  $\Delta\chi$  in simulated EVPA curves. The grey areas mark the 1, 2, and  $3\sigma$  intervals, while the dashed line is the observed rotation of  $309.5^\circ$ . The most probable value (peak of the solid distribution) is at about  $-140$  deg.



**Fig. 5.** J1505+0326: Distribution of  $\Delta\chi$  in the simulated EVPA curves where we assume the absence of intrinsic variability and the operation solely of noise. The dashed line is the observed rotation angle of  $309.5^\circ$ .

(during the rotation). The rotation is simulated in  $q$ - $u$  space by adding Gaussian noise to  $q$  and  $u$ .

We tested a range of rotation rates  $d\chi_{\text{intr}}/dt = \{-12., -11.5, \dots, +0.5\}$  in units of  $\text{deg d}^{-1}$ . For each rate we ran  $25 \times 10^3$  simulations and computed the probability of

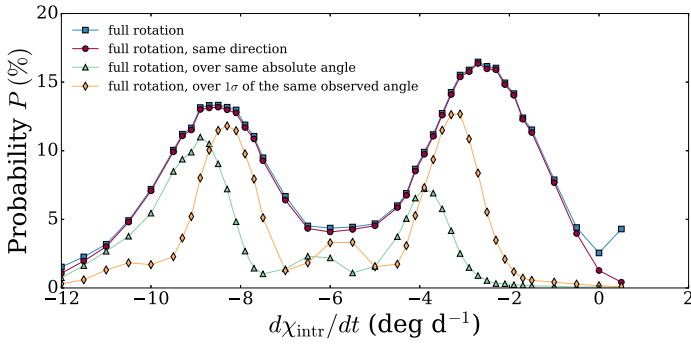
1. observing a full rotation, that is, over the entire period we simulated (i.e. including all data points); this is shown with blue squares in Fig. 6;
2. observing a full rotation in the same direction as in the data (in this case, negative derivative in EVPA); this is shown with red circles in Fig. 6;
3. observing a full rotation over an angle at least as large (in absolute terms) as the observed one ( $|\Delta\chi_{\text{sim}}| \geq |\Delta\chi_{\text{obs}}|$  i.e.  $\Delta\chi_{\text{sim}} \leq \Delta\chi_{\text{obs}}$ ); this is shown with green triangles in Fig. 6;
4. observing a full rotation with an angle within the  $1\sigma$  range of the rotation angle observed in the data ( $\Delta\chi_{\text{obs}} - \sigma_{\Delta\chi, \text{obs}} \leq \Delta\chi_{\text{sim}} \leq \Delta\chi_{\text{obs}} + \sigma_{\Delta\chi, \text{obs}}$ ); this is shown with orange diamonds in Fig. 6.

Figure 6 shows the resulting probability distributions. For a full rotation over angles at least as large as the observed one,

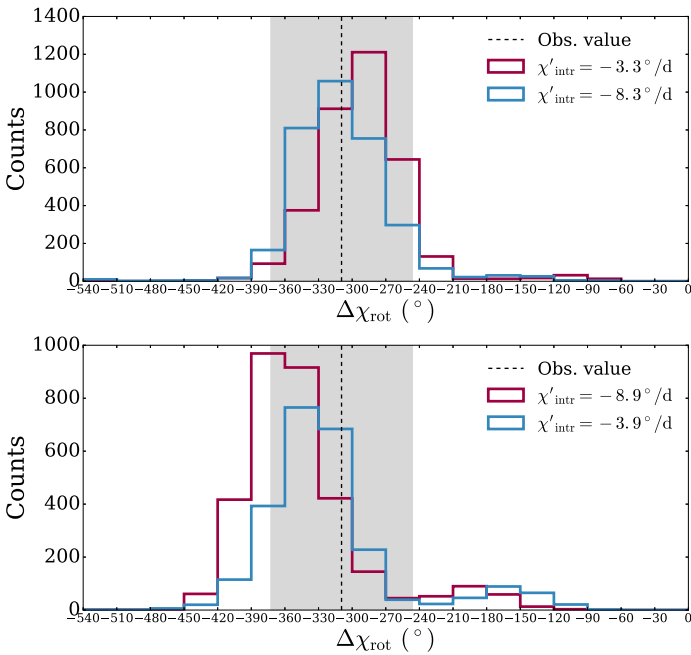
the most likely intrinsic rotation rate is  $-8.9 \pm 0.1 \text{ deg d}^{-1}$  with a corresponding probability of 0.11 (green diamonds). The second most probable rate is  $-3.9 \pm 0.1 \text{ deg d}^{-1}$  with a probability of 0.068. For a full rotation over an angle within  $1\sigma$  of the observed one (orange diamonds), the most probable intrinsic rotation rate is found to be  $-3.1 \pm 0.1 \text{ deg d}^{-1}$  with a probability of 0.129, while the second most likely one is  $-8.3 \pm 0.1 \text{ deg d}^{-1}$  with a probability of 0.119.

In Fig. 7 we show the distribution of rotation angle at the most likely intrinsic rotation rates based on the  $1\sigma$  criterion (upper panel) and on the basis of the extreme-span criterion (lower panel, case 3 of Sect. 5.1.3). The observed rotation angle (dashed vertical line) is consistent with an intrinsically constant rotation.

Considering all this and the conclusions of Sect. 5.1.2, we realise that it is much more likely that an intrinsic EVPA rotation (with the addition of pseudo-variability introduced by the uncertainties) is causing the observed event. Assuming a constant rotation rate also shows that the observed rotation is more likely the result of an intrinsic event than that of pure noise.



**Fig. 6.** J1505+0326: Probability distributions assuming an intrinsic EVPA rotation with a constant rate.



**Fig. 7.** J1505+0326: Distribution of rotation angles at the most likely intrinsic rotation rates based on the  $1\sigma$  (upper panel) and the extreme-span criterion (lower panel, case 3 of Sect. 5.1.3). The dashed line marks the observed value, and the grey area shows the  $1\sigma$  uncertainty.

## 5.2. J0324+3410

In Fig. 8 we show the debiased polarisation fraction  $\hat{p}$  and angle  $\chi$  as a function of time. The dataset includes RoboPol, KANATA, Perkins, and Steward observatory measurements. Both  $\chi$  and  $p$  show significant variability.

The polarisation fraction  $\hat{p}$  spreads around a median of 0.012 with a standard deviation of 0.016 (see Fig. 9 and Table 3). The distribution of the EVPA confined in the  $[-90^\circ, 90^\circ]$  range is shown in Fig. 10. It is distributed rather narrowly around a median of  $-6.7^\circ$  (with a standard deviation of  $40.2^\circ$ ). This preferred direction is at an angle of  $49.3^\circ$  with the 15 GHz radio jet axis, which is found to be remarkably stable at a position angle of  $124^\circ$  (Fuhrmann et al. 2016).

In total, we detected 28 apparent rotations of the polarisation plane with rotation angle  $\Delta\chi$ , ranging from approximately  $19^\circ$  to  $402^\circ$  (Fig. 11). Ten of these qualify as long rotations as they include at least three measurements and exceed  $90^\circ$ . In Fig. 9 we show the distribution of  $\hat{p}$  during phases of rotation and of non-rotation separately. In the former case, the median  $\hat{p}$  is 0.014 and in the latter, it is only 0.007. This indication that the polarisation fraction centres around different values in these two phases is not supported by a two-sample KS test.

As in the case of 1505+0326, in the following we concentrate on the rotation candidates. Figure 12 focuses on the area where the two main potential rotations occur. The largest occurs at around MJD 56640.5–56672.4 over  $402 \pm 87^\circ$ , corresponding to a mean rate of approximately  $13 \text{ deg d}^{-1}$ . The second largest event occurs around MJD 56595.6–56633.5 with an angle of  $349 \pm 66^\circ$  and a mean rate of  $9 \text{ deg d}^{-1}$ .

### 5.2.1. Largest potential rotation

Figure 13 (upper panel) demonstrates the uncertainty associated with the evolution of the measured EVPA. All steps are critical, and we cannot be certain of the direction the EVPA intrinsically takes at any point in its evolution. All this makes the very detection of the rotation uncertain.

Following the approach presented in Sect. 5.1.2, we examined whether the uncertainties in  $q$  and  $u$  alone can cause the observed rotation in the absence of an intrinsic rotation. We assumed again that the measured  $q$  and  $u$  are correct estimates of the means of the Gaussian-distributed Stokes parameters. After running  $10^4$  simulations, we find that the probability of finding one full rotation (passing over all points) is  $\sim 2 \times 10^{-2}$  (Eq. (6)), while that of finding a full rotation with absolute angle larger than observed is only  $8 \times 10^{-4}$  (Eq. (8)). Hence, although it is not impossible that the observed event is an artefact of noise while the EVPA remains intrinsically unchanged, it is rather unlikely. The associated probability is only  $\sim 10^{-3}$ . It is then possible that the EVPA indeed undergoes an intrinsic variability event.

To estimate the most probable parameters of the intrinsic EVPA variability, we assumed (as in Sect. 5.1.3) a constant intrinsic EVPA rotation rate and constancy of the polarisation fraction during the intrinsic rotation.

After  $2.5 \times 10^4$  iterations, we find that the most likely intrinsic rotation rate for a full rotation with an angle at least as large as the observed one is  $19 \pm 0.5 \text{ deg d}^{-1}$  (the probability is about  $10^{-2}$ ). The most probable rotation rate for a full rotation with an angle within  $1\sigma$  of the observed one is  $10 \pm 0.25 \text{ deg d}^{-1}$  (the probability is about 0.033).

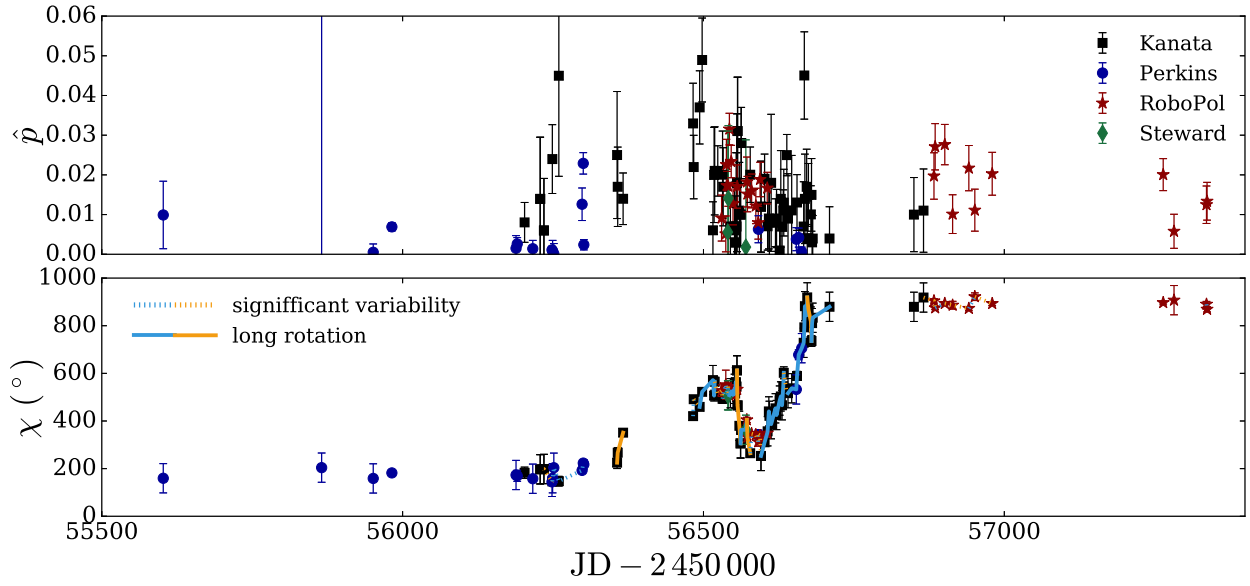
These probabilities are indeed low, but they are higher than those for the pure-noise scenario, indicating that intrinsic variability is more likely. Nevertheless, the low probability indicates that the simple assumption of a constant rotation rate is not likely in any case.

### 5.2.2. Second largest potential rotation

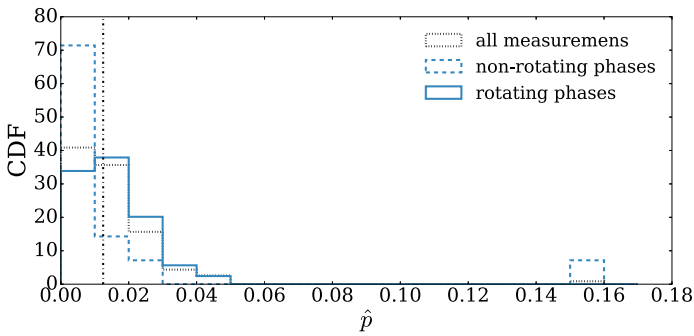
In the lower panel of Fig. 13 we show the second largest potential rotation. The probability that this event is the mere result of noise is as low as  $\sim 10^{-3}$ .

We repeated the analysis we presented earlier. After  $2.5 \times 10^4$  simulated EVPA curves, we find that for a full rotation with an angle at least as large as the observed one, the most probable intrinsic rate is  $10 \pm 0.25 \text{ deg d}^{-1}$  ( $P = 0.01$ ). For a full rotation over an angle within  $1\sigma$  of the observed one, the most probable rate is  $9.5 \pm 0.25 \text{ deg d}^{-1}$  ( $P = 0.008$ ).

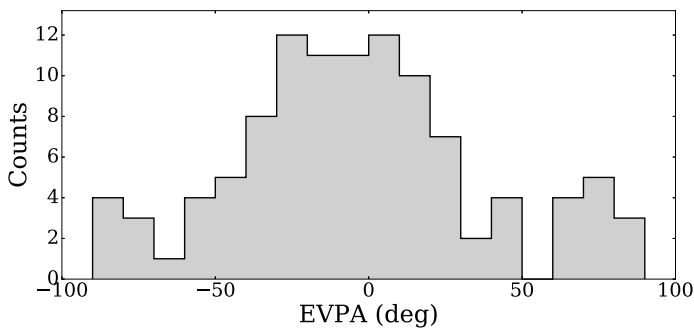
These probabilities are low and comparable to those of the noise artefact hypothesis. An intrinsic rotation of constant rate is only marginally more likely than noise. A realistic scenario would be that the observed behaviour results from the



**Fig. 8.** J0324+3410: Polarisation variability curve. *Upper panel:* Debiased polarisation fraction over time. *Lower panel:* EVPA over time. The coloured lines mark periods of monotonous (within the uncertainties) EVPA evolution.

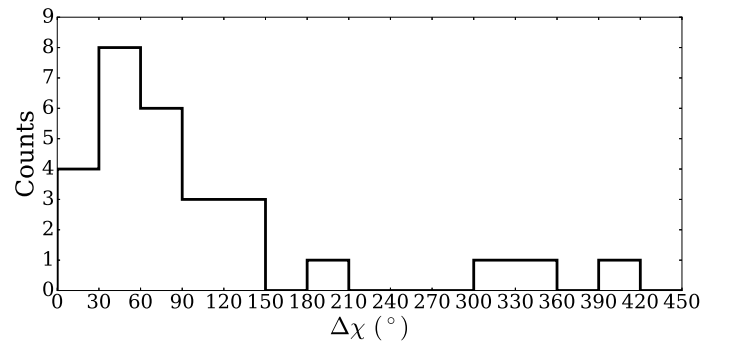


**Fig. 9.** J0324+3410: Distribution of the debiased polarisation fraction  $\hat{p}$ . The dot-dashed vertical line marks the median of the distribution. The blue dashed line shows the distribution of  $\hat{p}$  during the non-rotating phases, and the solid line the distribution during the rotating phases. The black dotted line corresponds to all the measurements.



**Fig. 10.** J0324+3410: Distribution of the observed polarisation angles in the range  $[-90^\circ, 90^\circ]$ .

combination of intrinsic variability and observational noise. The noise, however, makes the observed rotation angle an inadequate indicator of the intrinsic behaviour. Hence, although there may be intrinsic variability, we cannot recover it due to the noisy data.



**Fig. 11.** J0324+3410: Distribution of rotation angles  $\Delta\chi$  for all the apparent rotations.

### 5.3. J0849+5108 and J0948+0022: Polarisation variability

In Fig. 14 we show the  $\chi$  and  $\hat{p}$  datasets for J0849+5108 (*upper panel*) and J0948+0022 (*lower panel*) as a function of time. As shown in Fig. 14 and summarised in Table 3, the limited dataset does not allow sound quantification of the variability characteristics for either of the polarisation parameters. Intense variability is, nevertheless, clearly visible for both  $\hat{p}$  and  $\chi$ .

For J0849+5108 the median  $\langle\hat{p}\rangle$  is around 0.1 with a standard deviation 0.078. Concerning the angle, the available dataset revealed a total of nine rotations, none of which exceeded  $90^\circ$ . Clearly, despite the clear signs of variability, the data sparseness prevents any understanding of the intrinsic nature of the variability.

In the case of J0948+0022, the slightly richer dataset (KANATA, Perkins, RoboPol, and Steward) reveals the occurrence of 11 rotations, 2 of which over angles beyond  $90^\circ$ . For the largest rotation, the EVPA changed by  $268^\circ$  (Fig. 15). Although this EVPA curve is fairly reliable, the steps  $\chi_i - \chi_{i+1}$  between adjacent data points are very large and close to  $90^\circ$ , making this curve unsuitable for further analysis. Subsequently, nothing can be said as to whether the source underwent intrinsic EVPA rotations or not. For both sources, better sampled datasets are necessary. As in previous sections, in Fig. 16 we show the cumulative

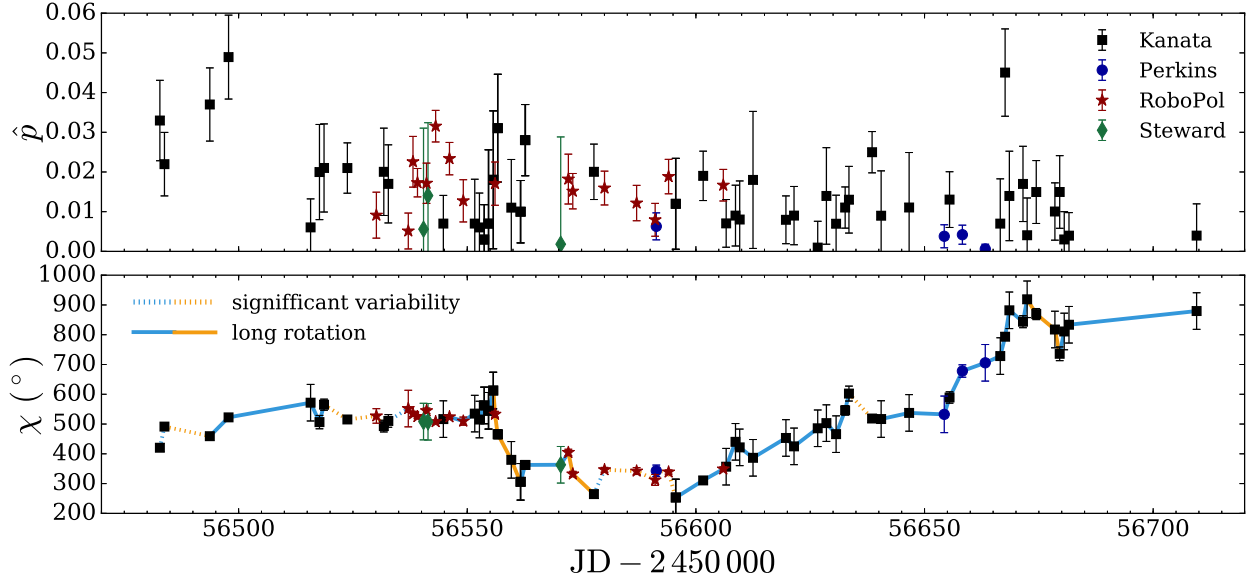


Fig. 12. J0324+3410: Zoom on the region of the largest potential rotations.

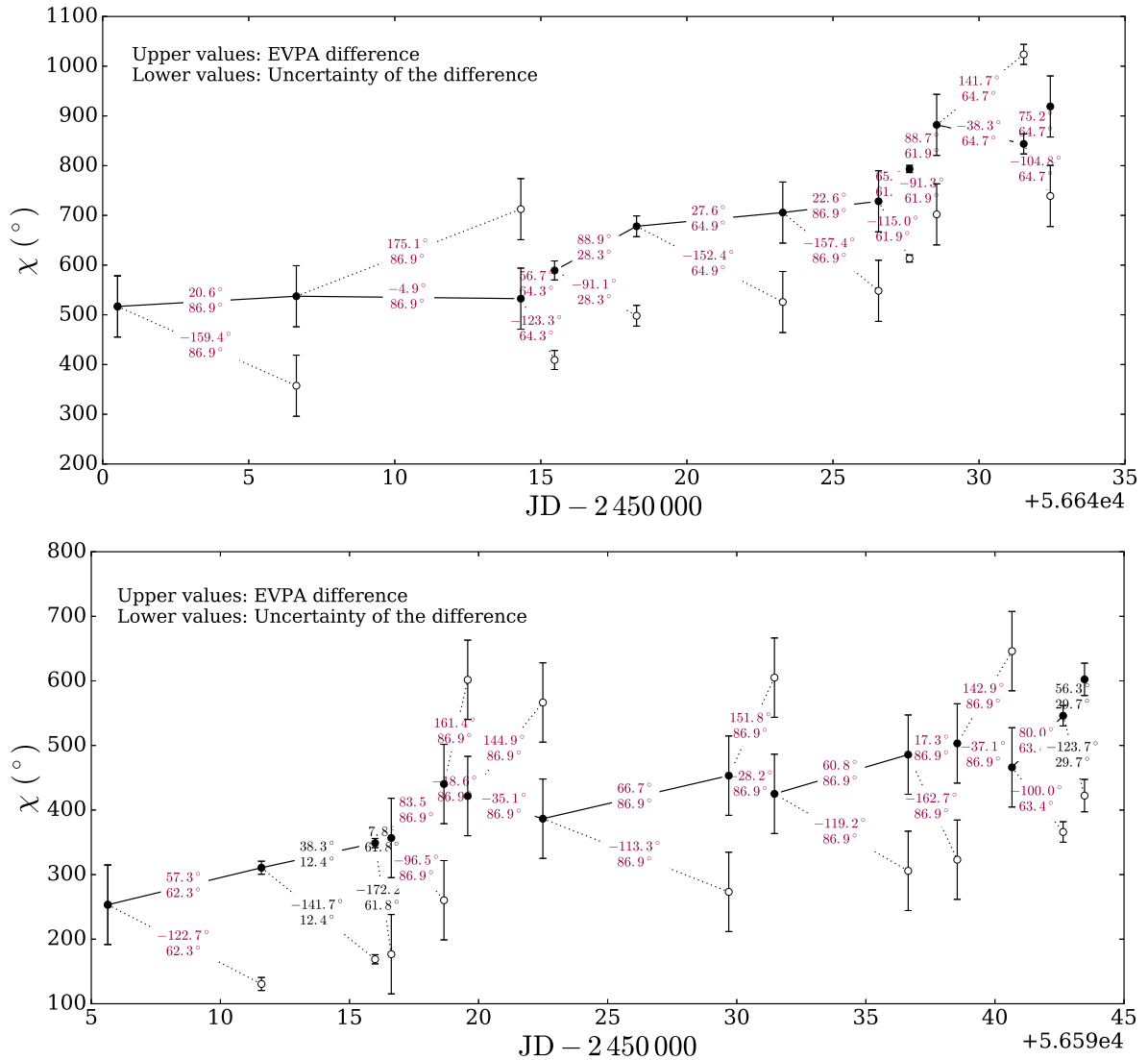
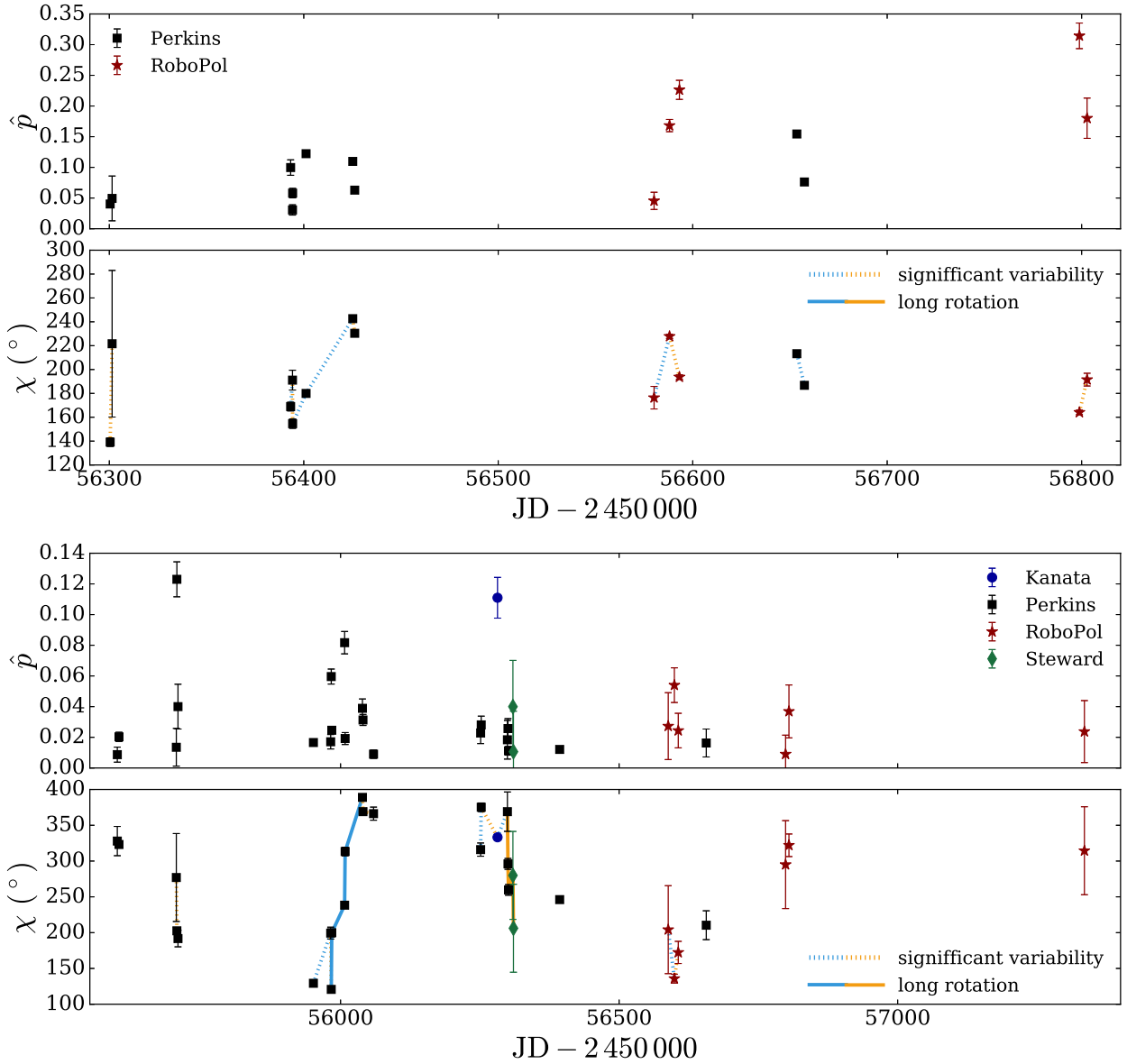


Fig. 13. J0324+3410: Two main potential EVPA rotations: the largest (*upper* panel) and the second largest potential (*lower* panel). The uncertainty in the observed events caused by sparse sampling and large measurement uncertainties becomes apparent.



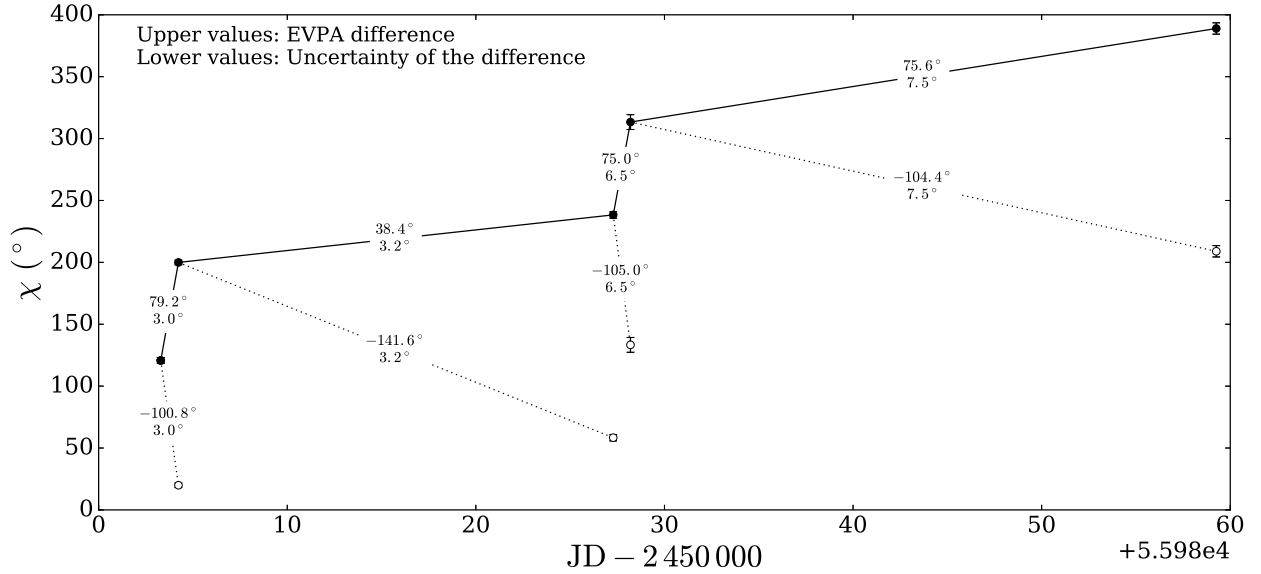
**Fig. 14.** Polarisation fraction and EVPA as a function of time for J0849+5108 (*upper panel*) and J0948+0022 (*lower panel*). The coloured lines mark periods of significant, monotonous (within the uncertainties) EVPA evolution. Dotted lines mark periods of significant evolution, while solid lines show periods of long rotations.

distribution function of  $\hat{p}$  separately for all measurements and for phases with rotation candidates. The median  $\langle\hat{p}\rangle$  is around 0.024 with a spread of 0.028. During phases of rotation, the median is 0.028, and over non-rotating phases, it drops to 0.016. As in previous cases, however, a two-sample KS test did not support the hypothesis of a different behaviour in the two activity states.

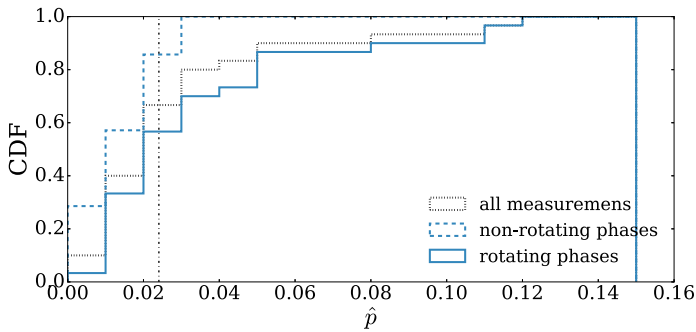
## 6. Discussion

As a result of the scarcity of high-cadence datasets, the current understanding of the optical polarisation variability of RL NLSy1s is incomplete and has been gained through selected case studies rather than a systematic population study. PMN J0948+0022, for example, was found to have variable optical polarisation (degree and angle, Eggen 2012; Eggen et al. 2013) even on minute timescales (Itoh et al. 2013). Itoh and collaborators reported that the polarisation briefly exceeded

30%, while the EVPA remained unchanged. Maune et al. (2014) reported that J0849+5108 also showed rapid intra-night variability both in degree of polarisation and angle. 1H 0323+342 was found by Itoh et al. (2014) to possess an EVPA that remained constantly parallel to the jet orientation. All these findings indicate how unknown the optical polarisation from such systems and its temporal behaviour remains. Our study attempts to overcome this barrier by examining a larger sample that would potentially allow us to extract general conclusions and further compare them with blazars. Larger systematic studies are, however, needed to further clarify (a) the confirmation of EVPA rotation events in RL NLSy1s and later a comparison of their parameters to those in blazars, (b) the behaviour of the polarisation degree during rotations of the EVPA (Blinov et al. 2016b), and (c) the association of rotation events with the GeV energy band activity (Blinov et al. 2018). Below we discuss some of the understanding we have gained with this work, which seems to indicate similarities to the behaviour seen in blazars.



**Fig. 15.** J0948+0022: Main potential EVPA rotation. The uncertainty in the observed events caused by sparse sampling and large uncertainties is obvious.



**Fig. 16.** J0948+0022: Distribution of the debiased polarisation fraction  $\hat{p}$ . The dot-dashed vertical line marks the median of the distribution. The blue dashed line shows the distribution of  $\hat{p}$  during the non-rotating phases, and the solid line shows this during the rotating phases. The black dotted line corresponds to all the measurements.

The polarisation variability. As we have shown, the optical polarisation parameters of RL NLSy1 galaxies show clear signs of variability. Both the polarisation fraction  $\hat{p}$  and angle  $\chi$  show phenomenologies similar to those seen in blazars, indicating similar processes in the two classes. It would be premature to conclude that the observed variability of RL NLSy1s implies turbulent processes until datasets with much higher cadence of a much larger sample of sources have been considered. The “blazars-like” character of these sources is established anew, however, and it is important for future studies to focus on examining the fundamental differences and similarities between the two classes.

The polarisation angle variability. The very detection of significant variability of the EVPA represents a particularly important finding. Earlier studies have indicated that in selected cases, the polarisation angle remained stable even during phases of intense polarisation fraction variability. This stability was interpreted as an indication for a high degree of magnetic field organisation in the regions where the radiation is produced. Our analysis establishes that the variability of the polarisation angle is common in all our sources with at least moderate sampling.

Preferred orientation of the polarisation plane. In the one case with a sufficiently large dataset (namely J0324+3410), we examined the distribution of the polarisation plane orientations. This was done by confining the polarisation orientations within the range  $[-90^\circ, 90^\circ]$ . The EVPA of this source is not oriented randomly. It instead shows a concentration (with some breadth) around a preferred direction of  $-6.7^\circ$ . This orientation is at  $49.3^\circ$  to the position angle of the 15 GHz radio jet. The behaviour of preferential orientation of the EVPA resembles the high synchrotron peak frequency sources discussed by Angelakis et al. (2016, 2017).

Under the assumption that the optical emission is thin, this configuration implies a projected net magnetic field that is oriented at  $40.7^\circ$  to the radio jet<sup>5</sup>. This misalignment could be an indication of a combination of poloidal and toroidal field components. A more systematic approach, of course, would require the careful consideration of projection effects, the location and size of the emission region (potentially very different for the optical and radio emission), and the effects of relativistic aberration (Lyutikov et al. 2005). The variability accompanying the preferential orientation, on the other hand, could be understood in terms of an additional turbulent magnetic field component (e.g. Marscher 2014), a regular modulation of the emission region location (e.g. Marscher et al. 2008), and relativistic delays and projection effects (Zhang et al. 2014). Such configurations are not un-realistic. It should be noted, however, that the optical emission likely originates from a more compact region much closer to the jet base when interpreting these observations.

The detection of long polarisation plane rotations. In the general framework of polarisation variability, we have searched for long rotations of the polarisation plane and assessed the probability of such events being driven by intrinsic rotations of the polarisation plane. As we have shown, the events detected (even for the best sampled EVPA curves) have to be viewed with caution. The probability that they are the mere result of noise is not null. Our analysis clearly shows, however, that the probability of an intrinsic event driving the apparent behaviour is significantly

<sup>5</sup> In the optically thin regime of synchrotron emission, the projected magnetic field is perpendicular to the observed EVPA.

higher, especially in the case of J1505+0326. We emphasise the importance of this analysis for the assessment of these probabilities, which we consider a part of our work equally as important as the detection of the rotation events themselves. Denser datasets of larger samples are necessary to prove that the occurrence of such events is a general characteristic of RL NLSy1s.

The fractional polarisation during rotations. As suggested by Blinov et al. (2016b), EVPA rotations in blazars seem to be associated with lower fractional polarisation in a statistical sense. This is a rather mild effect, but it may provide a diagnostic for the rotation mechanisms (Kiehlmann et al. 2017). Although our dataset is insufficient to test this, for the two largest datasets (J0324+3410 and J1505+0326), the polarisation tends to be marginally higher during the rotation phases. A two-sample KS tests did not provide any evidence that the distributions of  $\hat{p}$  in the two activity phases are really different, however. This ambiguity will be studied in a future publication.

Physical interpretation. Clearly, the current dataset cannot shed light on the physical interpretation of the observed variability. We cannot tell whether it is the physical rotation of the emission element on a helical trajectory, the macroscopic properties of the jet, turbulent processes resulting in random walks, light travel-time effects, or any other process that causes EVPA rotations. Much longer and better sampled light curves of larger samples are necessary to prove that the apparent EVPA rotation is caused by intrinsic variability. Furthermore, studying whether EVPA variability is correlated with gamma-ray flaring might also provide insights into the physical processes driving these rotations. Blinov et al. (2018) reported that no EVPA rotations have been detected in blazars that are not associated with some activity in the *Fermi* energy bands. This is of fundamental importance to understanding the mechanism behind the long rotations of the polarisation plane. Investigating whether this is true for RL NLSy1s would be a natural next step.

## 7. Conclusions

We have conducted optical polarisation monitoring of a sample of ten RL NLSy1 galaxies, five of which have been found to radiate significant MeV–GeV emission. Our main goal was to quantify the variability of the two polarisation parameters, the debiased fraction  $\hat{p}$  and angle  $\chi$ . We further examined whether long rotations of the EVPA (similar to those found in blazars), are present. Our main conclusions are listed below.

1. All cases with adequately large datasets for both  $\hat{p}$  and  $\chi$  show significant variability.
2. For the four GeV emitting sources in our sample that also had dense and long enough datasets, we find significant variability in the EVPA. For the remaining sources, either the sparseness of the datasets or the noise do not allow such studies.
3. In the case J0324+3410, we find that the EVPA spreads around a preferred orientation that is at an angle of 49.3° to the 15 GHz radio jet. Hence the projected magnetic field is at an angle of 40.7° to the jet axis.
4. In two of those cases, namely J1505+0326 and J0324+3410, we have found evidence for the presence of intrinsic EVPA rotations. Careful numerical simulations have been conducted to assess the probability that these events are driven by intrinsic variability.
5. For the three largest apparent rotations, we have assessed the likeliness that the observed rotations are the result of pure observational noise in the absence of an intrinsic event. We showed that although measurement uncertainties may indeed induce such behaviours, it is rather unlikely. It appears much more likely that the observed variability is indeed driven by intrinsic rotation at a constant rate.
6. For the two main candidate events, we estimated the most probable parameters for the intrinsic rotation on the basis of the constant rate assumption. We conclude that a linear trend of the intrinsic rotation is more likely than a non-varying EVPA. Most likely, however, a more complex situation appears more realistic. Relaxing this condition would make the probability of an intrinsic event causing the observed rotation even higher.
7. For the best-sampled cases, we examined the behaviour of the polarisation fraction during rotation and non-rotation periods. A two-sample KS tests indicates no significant difference between them.
8. Our analyses show that more observations are clearly needed for further concluding on all the topics discussed here. Although there is evidence for long rotations of the optical polarisation plane, higher cadence data of larger samples are needed.

*Acknowledgements.* The authors wish to thank the anonymous referee for the very careful examination of the manuscript and the constructive comments, S. Komossa for selecting the non-*Fermi*-detected RL NLSy1s of our sample based on the criteria given in Table 1, the internal MPIfR referee N. MacDonald for the careful multiple reading of the manuscript and the detailed comments, V. Karamanavis for the project planing, discussions, and comments on the manuscript, and V. Pavlidou for the discussions and comments as well as the support to the Skinakas observing proposal, and Dr. J. Maune for acquisition of a large portion of the Perkins data. The RoboPol project is a collaboration between Caltech in the USA, MPIfR in Germany, the Toruń Centre for Astronomy in Poland, the University of Crete/FORTH in Greece, and IUCAA in India. Data acquired with the Perkins telescope was funded by the PEGA RPE grant at Georgia State University. Data from the Steward Observatory spectropolarimetric monitoring project were used. This program is supported by Fermi Guest Investigator grants NNX08AW56G, NNX09AU10G, NNX12AO93G, and NNX15AU81G.

## References

- Abdo, A. A., Ackermann, M., Ajello, M., et al. 2009a, *ApJ*, 699, 976  
 Abdo, A. A., Ackermann, M., Ajello, M., et al. 2009b, *ApJ*, 707, L142  
 Abdo, A. A., Ackermann, M., Ajello, M., et al. 2010, *Nature*, 463, 919  
 Angelakis, E., Fuhrmann, L., Marchili, N., et al. 2015, *A&A*, 575, A55  
 Angelakis, E., Hovatta, T., Blinov, D., et al. 2016, *MNRAS*, 463, 3365  
 Angelakis, E., Blinov, D., Böttcher, M., et al. 2017, ArXiv e-prints [arXiv:1711.04824]  
 Baldi, R. D., Capetti, A., Robinson, A., Laor, A., & Behar, E. 2016, *MNRAS*, 458, L69  
 Blinov, D., Pavlidou, V., Papadakis, I., et al. 2016a, *MNRAS*, 462, 1775  
 Blinov, D., Pavlidou, V., Papadakis, I. E., et al. 2016b, *MNRAS*, 457, 2252  
 Blinov, D., Pavlidou, V., Papadakis, I., et al. 2018, *MNRAS*, 474, 1296  
 Calderone, G., Ghisellini, G., Colpi, M., & Dotti, M. 2013, *MNRAS*, 431, 210  
 Clarke, D. 2010, *Stellar Polarimetry*  
 D’Ammando, F., Orienti, M., Finke, J., et al. 2012, *MNRAS*, 426, 317  
 D’Ammando, F., Orienti, M., Larsson, J., & Giroletti, M. 2015, *MNRAS*, 452, 520  
 D’Ammando, F., Acosta-Pulido, J. A., Capetti, A., et al. 2017, *MNRAS*, 469, L11  
 Doi, A., Oyama, T., Kono, Y., et al. 2016, *PASJ*, 68, 73  
 Eggen, J. R. 2012, in *Am. Astron. Soc. Meet. Abstr.*, 219, 243.29  
 Eggen, J. R., Miller, H. R., & Maune, J. D. 2013, *ApJ*, 773, 85  
 Foschini, L. 2011, in *Narrow-Line Seyfert 1 Galaxies and their Place in the Universe*  
 Foschini, L., Angelakis, E., Fuhrmann, L., et al. 2012, *A&A*, 548, A106  
 Foschini, L., Berton, M., Caccianiga, A., et al. 2015, *A&A*, 575, A13  
 Fuhrmann, L., Karamanavis, V., Komossa, S., et al. 2016, *Res. Astron. Astrophys.*, 16, 176  
 Goodrich, R. W. 1989, *ApJ*, 342, 224  
 Hewett, P. C., & Wild, V. 2010, *MNRAS*, 405, 2302  
 Itoh, R., Tanaka, Y. T., Fukazawa, Y., et al. 2013, *ApJ*, 775, L26  
 Itoh, R., Tanaka, Y. T., Akitaya, H., et al. 2014, *PASJ*, 66, 108  
 Kawabata, K. S., Nagae, O., Chiyonobu, S., et al. 2008, in *Ground-based and Airborne Instrumentation for Astronomy II*, Proc. SPIE, 7014, 70144L

- Kellermann, K. I., Sramek, R., Schmidt, M., Shaffer, D. B., & Green, R. 1989, *AJ*, **98**, 1195
- Kiehlmann, S., Savolainen, T., Jorstad, S. G., et al. 2016, *A&A*, **592**, C1
- Kiehlmann, S., Blinov, D., Pearson, T. J., & Liodakis, I. 2017, *MNRAS*, **472**, 3589
- Kikuchi, S., Mikami, Y., Inoue, M., Tabara, H., & Kato, T. 1988, *A&A*, **190**, L8
- King, O. G., Blinov, D., Ramaprakash, A. N., et al. 2014, *MNRAS*, **442**, 1706
- Komossa, S., Voges, W., Adorf, H.-M., et al. 2006a, *ApJ*, **639**, 710
- Komossa, S., Voges, W., Xu, D., et al. 2006b, *AJ*, **132**, 531
- Komossa, S., Xu, D., & Zensus, J. A. 2016, in *Star Clusters and Black Holes in Galaxies across Cosmic Time*, eds. Y. Meiron, S. Li, F.-K. Liu, & R. Spurzem, *IAU Symp.*, **312**, 63
- Kynoch, D., Landt, H., Ward, M. J., et al. 2018, *MNRAS*, **475**, 404
- Lähteenmäki, A., Järvelä, E., Hovatta, T., et al. 2017, *A&A*, **603**, A100
- Landt, H., Ward, M. J., Baloković, M., et al. 2017, *MNRAS*, **464**, 2565
- Liao, N.-H., Liang, Y.-F., Weng, S.-S., Gu, M.-F., & Fan, Y.-Z. 2015, *ArXiv e-prints* [arXiv:1510.05584]
- Lyutikov, M., Pariev, V. I., & Gabuzda, D. C. 2005, *MNRAS*, **360**, 869
- Marconi, A., Axon, D. J., Maiolino, R., et al. 2008, *ApJ*, **678**, 693
- Marscher, A. P. 2014, *ApJ*, **780**, 87
- Marscher, A. P., Jorstad, S. G., D’Arcangelo, F. D., et al. 2008, *Nature*, **452**, 966
- Marscher, A. P., Jorstad, S. G., Larionov, V. M., et al. 2010, *ApJ*, **710**, L126
- Maune, J. D., Eggen, J. R., Miller, H. R., et al. 2014, *ApJ*, **794**, 93
- Naghizadeh-Khouei, J., & Clarke, D. 1993, *A&A*, **274**, 968
- Oh, K., Yi, S. K., Schawinski, K., et al. 2015, *ApJS*, **219**, 1
- Osterbrock, D. E., & Pogge, R. 1985, *ApJ*, **297**, 166
- Paliya, V. S., & Stalin, C. S. 2016, *ApJ*, **820**, 52
- Paliya, V. S., Rajput, B., Stalin, C. S., & Pandey, S. B. 2016, *ApJ*, **819**, 121
- Papamastorakis, Y. 2007, *Ipparchos*, **2**, 14
- Pavlidou, V., Angelakis, E., Myserlis, I., et al. 2014, *MNRAS*, **442**, 1693
- Rice, S. O. 1945, *Bell Systems Tech. J.*, **24**, 46
- Schmidt, G. D., Elston, R., & Lupie, O. L. 1992, *AJ*, **104**, 1563
- Serkowski, K. 1958, *Acta Astron.*, **8**, 135
- Shen, Y., Richards, G. T., Strauss, M. A., et al. 2011, *ApJS*, **194**, 45
- Simmons, J. F. L., & Stewart, B. G. 1985, *A&A*, **142**, 100
- Smith, P. S., Montiel, E., Rightley, S., et al. 2009, *Fermi Symposium*, eConf Proceedings C091122, [arXiv:0912.3621]
- Vaillancourt, J. E. 2006, *PASP*, **118**, 1340
- Vinokur, M. 1965, *Ann. Astrophys.*, **28**, 412
- Wang, F., Du, P., Hu, C., et al. 2016, *ApJ*, **824**, 149
- Wardle, J. F. C., & Kronberg, P. P. 1974, *ApJ*, **194**, 249
- Xu, D., Komossa, S., Zhou, H., et al. 2012, *AJ*, **143**, 83
- Yao, S., Yuan, W., Zhou, H., et al. 2015, *MNRAS*, **454**, L16
- Yuan, W., Zhou, H. Y., Komossa, S., et al. 2008, *ApJ*, **685**, 801
- Yuan, W., Liu, B. F., Zhou, H., & Wang, T. G. 2010, *ApJ*, **723**, 508
- Zhang, H., Chen, X., & Böttcher, M. 2014, *ApJ*, **789**, 66
- Zhang, H., Chen, X., Böttcher, M., Guo, F., & Li, H. 2015, *ApJ*, **804**, 58
- Zhou, H.-Y., Wang, T.-G., Dong, X.-B., Zhou, Y.-Y., & Li, C. 2003, *ApJ*, **584**, 147
- Zhou, H.-Y., Wang, T.-G., Dong, X.-B., Li, C., & Zhang, X.-G. 2005, *Chin. J. Astron. Astrophys.*, **5**, 41
- Zhou, H., Wang, T., Yuan, W., et al. 2006, *ApJS*, **166**, 128
- Zhou, H., Wang, T., Yuan, W., et al. 2007, *ApJ*, **658**, L13