



Publication Year	2019
Acceptance in OA	2020-12-30T09:18:01Z
Title	Discovery and progenitor constraints on the Type Ia supernova 2013gy
Authors	Holmbo, S., Stritzinger, M. D., Shappee, B. J., Tucker, M. A., Zheng, W., Ashall, C., Phillips, M. M., Contreras, C., Filippenko, A. V., Hoefflich, P., Huber, M., Piro, A. L., Wang, X. F., Zhang, J. -J., Anais, J., Baron, E., Burns, C. R., Campillay, A., Castellón, S., Corco, C., Hsiao, E. Y., Krisciunas, K., Morrell, N., Nielsen, M. T. B., Persson, S. E., Taddia, F., TOMASELLA, Lina, Zhang, T. -M., Zhao, X. -L.
Publisher's version (DOI)	10.1051/0004-6361/201834389
Handle	http://hdl.handle.net/20.500.12386/29330
Journal	ASTRONOMY & ASTROPHYSICS
Volume	627

Discovery and progenitor constraints on the Type Ia supernova 2013gy[★]

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Received 5 October 2018 / Accepted 30 January 2019

ABSTRACT

We present an early-phase g -band light curve and visual-wavelength spectra of the normal Type Ia supernova (SN) 2013gy. The light curve is constructed by determining the appropriate S -corrections to transform KAIT natural-system B - and V -band photometry and Carnegie Supernova Project natural-system g -band photometry to the Pan-STARRS1 g -band natural photometric system. A Markov chain Monte Carlo calculation provides a best-fit single power-law function to the first ten epochs of photometry described by an exponent of $2.16^{+0.06}_{-0.06}$ and a time of first light of MJD 56629.4 $^{+0.1}_{-0.1}$, which is $1.93^{+0.12}_{-0.13}$ days (i.e., <48 h) before the discovery date (2013 December 4.84 UT) and $-19.10^{+0.12}_{-0.13}$ days before the time of B -band maximum (MJD 56648.5 \pm 0.1). The estimate of the time of first light is consistent with the explosion time inferred from the evolution of the Si II λ 6355 Doppler velocity. Furthermore, discovery photometry and previous nondetection limits enable us to constrain the companion radius down to $R_c \leq 4 R_\odot$. In addition to our early-time constraints, we used a deep +235 day nebular-phase spectrum from *Magellan*/IMACS to place a stripped H-mass limit of $<0.018 M_\odot$. Combined, these limits effectively rule out H-rich nondegenerate companions.

Key words. supernovae: general – supernovae: individual SN 2013gy

1. Introduction

Today, Type Ia supernova (SN Ia) cosmology is limited by a subtle matrix of systematic errors. Although significant effort is being placed to reduce known systematics in order to improve upon the accuracy of measuring their peak luminosities (see Phillips et al. 2019), obtaining distances accurate to the percent

* The spectra are only available at the CDS via anonymous ftp to cdsarc.u-strasbg.fr (130.79.128.5) or via <http://cdsarc.u-strasbg.fr/viz-bin/qcat?J/A+A/627/A174>.

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level will likely require a better understanding of their progenitors and explosion physics. Although SNe Ia are generally agreed to originate from the thermonuclear disruption of a carbon-oxygen white dwarf (Hoyle & Fowler 1960), progenitor scenarios which give rise to SNe Ia and explosion processes involved remain critical questions to answer. With the growth of wide field surveys, these questions can be addressed better than ever before. In particular, the discovery of SNe Ia shortly after explosion can provide new information through their early light curves, such as shock cooling of the white dwarf surface (Piro et al. 2010; Rabinak et al. 2012), collision

with the companion star (Kasen 2010), evidence for shallow radioactive nickel (Piro & Morozova 2016; Maeda et al. 2018; Polin et al. 2019), and interaction with circumstellar material (Piro & Morozova 2016; Noebauer et al. 2016; Maeda et al. 2018; Tucker et al. 2019).

The current sample of SNe Ia with early-time observations exhibit at least two different morphologies (Stritzinger et al. 2018). In one case, the early light curves are well fit with a single power-law function. Objects exhibiting such a behavior include SN 2011fe (Zhang et al. 2016), KSN 2011b (Olling et al. 2015; Shappee et al. 2019), KSN 2012a (Olling et al. 2015; Shappee et al. 2019), ASASSN-14lp (Shappee et al. 2016), and SN 2015F (Im et al. 2015; Cartier et al. 2017). On the other hand, some SNe Ia exhibit a two-phase morphology. In such cases the first phase consists of a linear increase in flux typically extending over three to five days past the time of first light (t_{first}), followed by an abrupt increase in flux (power law) as the light curve rises to maximum brightness. Objects exhibiting a two-phase morphology include SN 2012cg, (Marion et al. 2016; Shappee et al. 2018), SN 2012fr (Contreras et al. 2018), SN 2013dy (Zheng et al. 2013; Pan et al. 2015), SN 2014J (Zheng et al. 2014; Siverd et al. 2015; Goobar et al. 2015), iPTF 16abc (Miller et al. 2018), and SN 2017cbv (Hosseinzadeh et al. 2017). A two-component power-law function (or broken power-law function; Zheng et al. 2017) provides superior fits to the early light curves of these objects and more accurate estimates of both t_{first} and the rise time (t_{rise}) to maximum (t_{max}).

Aside from four SNe Ia located in Kepler fields (Olling et al. 2015; Brown et al. 2018; Cornect et al. 2018), the current sample of objects studied at early times consists of low-cadence coverage, and it may be possible that a two-component emission structure has been misinterpreted as a single power law. For example, Foley et al. (2012) fit the early light curve of SN 2009ig with a single power law; however, when photometry computed from an open-filter discovery image is included, a two-component power law may indeed provide a superior fit (see Contreras et al. 2018). Taken together with the recent observations of MUSSES1604D (also known as SN 2016jhr; Jiang et al. 2017) a SN 2006bt-like event (Foley et al. 2010; Stritzinger et al. 2011) that exhibited clear departures from a single and double power-law rise, the early-phase observational parameter space of SNe Ia is far from fully explored.

Currently there are <20 SNe Ia in the literature discovered within approximately three days of t_{first} , and the addition of more objects to the sample is crucial to fully map out the early-time parameter space. Here we present an early g -band light curve of the normal SN 2013gy (also known as PS1-13ejo) based on photometry calibrated to one well-understood photometric system. Specifically, Pan-STARRS1 (PS1; Tonry et al. 2012; Chambers et al. 2016) first g -band detection images are combined with follow-up imaging obtained by the Carnegie Supernova Project-II (CSP-II; Phillips et al. 2019) and the Lick Observatory Supernova Search (LOSS; Filippenko et al. 2001). In addition to the light curve, a comprehensive set of visual-wavelength spectra is also presented.

Kim et al. (2013) announce the discovery of SN 2013gy in open-filter images obtained with the 0.76 m Katzman Automatic Imaging Telescope (KAIT) at Lick Observatory on 2013 December 6.33 (UT dates are used throughout this paper). An open-filter image taken the previous day provides a nondetection limit of 19.3 mag. SN 2013gy was also recovered in PS1 images obtained on 2013 December 5.34 with an apparent g -band (hereafter g_{PS1}) magnitude of 19.48 ± 0.09 . The previous PS1 nondetection image was obtained 5.95 days earlier on 2013

November 29.39 with a limiting g_{PS1} magnitude of 20.54. An optical spectrum obtained by Tomasella et al. (2013) on 2013 December 7.5 indicated the object was a young SN Ia.

With J2000 coordinates of $\alpha = 03^{\text{h}}42^{\text{m}}16^{\text{s}}.88$ and $\delta = -04^{\circ}43'18''52$, SN 2013gy was located $32''$ north and $11''$ east from the center of the SB(s)b host galaxy NGC 1418. A color image of NGC 1418 with the position of SN 2013gy indicated is provided in Fig. A.1. The redshift of NGC 1418 is $z = 0.014023$ (Catinella et al. 2005) and its Tully Fisher (TF) distance ranges from 43.8 Mpc (Springob et al. 2007) to 88.8 Mpc (Theureau et al. 2007). Given the high dispersion of the TF distances, in the following we adopt the redshift distance, which after correcting for a Virgo, Great Attractor, and Shapley infall model and adopting $H_0 = 73 \text{ km s}^{-1} \text{ Mpc}^{-1}$ corresponds to a distance of $55.9 \pm 3.9 \text{ Mpc}$ (i.e., $\mu = 33.75 \pm 0.15 \text{ mag}$). This is fully consistent with the distance derived from the comprehensive set of optical and near-infrared (NIR) light curves obtained by the CSP-II, which provides $\mu = 33.68 \pm 0.09 \text{ mag}$ (see Sect. 2.3).

2. Early-time observations of SN 2013gy

2.1. Photometry and visual-wavelength spectroscopy

Our early-phase g_{PS1} -band light curve of SN 2013gy is constructed using data from three different facilities. This includes two epochs of g_{PS1} separated by ~ 20 min, 52 epochs of g -band photometry obtained by the CSP-II (hereafter g_{CSP}) extending from -15.3 d^1 to $+59.4 \text{ d}$, and 37 epochs of B - and V -band photometry from KAIT extending from -16.2 d to $+53.6 \text{ d}$. To facilitate our analysis presented below, in the following the KAIT B - and V -band photometry is summed together (in flux space) and then multiplied by one half, creating what we refer to as $BV.5$ -band photometry (see below). Reduced g_{PS1} -band photometry was downloaded from the PS1 webpage², while CSP-II and KAIT images were reduced in the standard manner following the techniques described by Krisciunas et al. (2017) and Zheng et al. (2013), respectively. Natural-system photometry of the supernova was computed for each set of science images relative to local sequences of stars. These local sequences of stars were calibrated to the natural system of each setup using standard-star photometry converted to the natural systems through the use of color terms obtained from multiple observations of the standard fields (see Contreras et al. 2018, for details). The resulting g_{PS1} -, g_{CSP} -, and KAIT $BV.5$ -band photometry is listed in Table A.1, along with the S-corrections (see below) that transform their photometry to the g_{PS1} natural system.

Table A.2 contains the journal of spectroscopic observations. These include 18 epochs of unpublished early optical spectra, five published spectra (Graham et al. 2017), and a late-phase spectrum taken with the *Magellan Baade* telescope. The spectra were reduced following standard procedures (see, e.g., Hamuy et al. 2006), were color-matched to multiband photometry, and are plotted in the rest frame in Fig. A.2.

2.2. S-correction and a definitive g_{PS1} light curve

Here we describe how the g_{CSP} - and $BV.5$ -band photometry was transformed to the g_{PS1} natural system. We transform to the g_{PS1} natural system to avoid S-correcting the earliest

¹ In keeping with tradition, temporal phases are given with respect to the time of B -band maximum brightness ($t_{B,\text{max}}$) unless explicitly stated.

² <https://star.pst.qub.ac.uk/ps1threepi/psdb/candidate/1034216880044318600/> (Smartt et al. 2014).

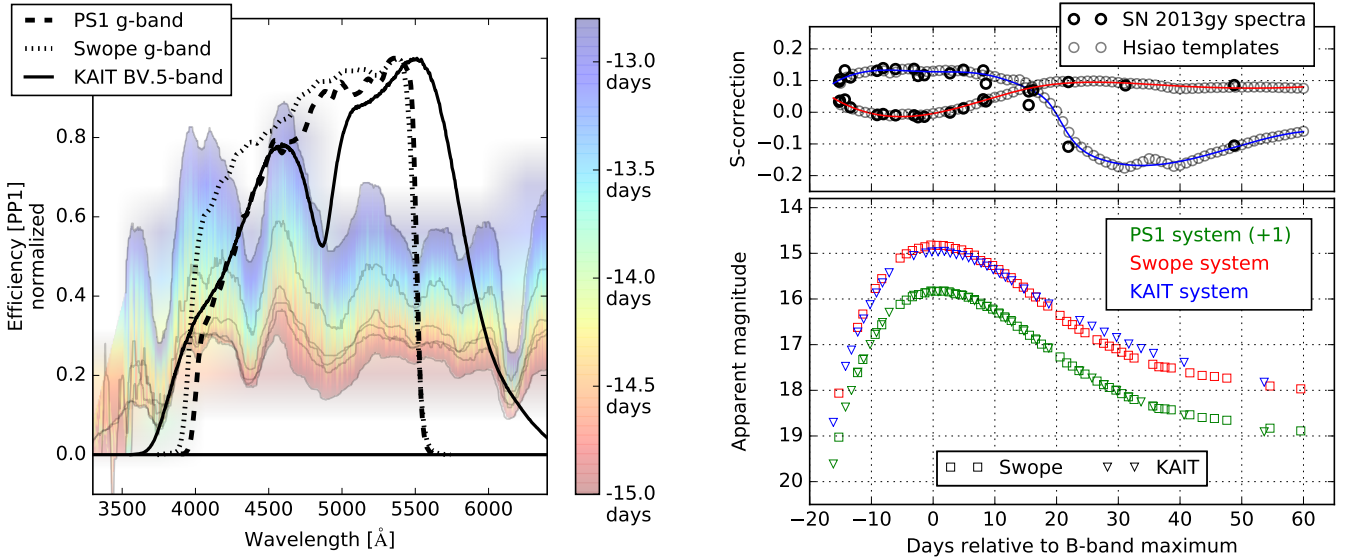


Fig. 1. System response functions and S-correction results. *Left*: comparison of the normalized g_{PS1} (dashed line), the g_{CSP} (dotted line), and $BV.5$ (solid line) system response functions. Also plotted are seven visual-wavelength spectra of SN 2013gy ranging between -16.1 d and -13.4 d as indicated by the color bar. *Right, top panel*: S-corrections as a function of phase using observed (dark circles) spectra and Hsiao et al. (2007) template spectra (light circles). The smoothing splines of those S-corrections (red and blue lines for g_{CSP} and $BV.5$, respectively) enable us to accurately transform g_{CSP} -band and $BV.5$ -band natural-system photometry to the g_{PS1} natural system at any phase. *Right, bottom panel*: g_{CSP} - (red squares) and $BV.5$ -band (blue triangles) natural-system photometry, and the S-corrected versions (green squares and green triangles) transformed to the g_{PS1} natural system (offset by +1 mag).

observations where the S-corrections are more uncertain. To compute accurate S-corrections requires system response functions and a spectrum, and/or a spectral template (Stritzinger et al. 2002). The left panel of Fig. 1 contains the g_{PS1} -, g_{CSP} -, and KAIT $BV.5$ -band response functions, along with early-phase optical spectra. The $BV.5$ response function is obtained by summing together the KAIT B - and V - response functions and multiplying the product by one half. The motivation for doing this is that the S-corrections computed for the KAIT B and V bands are found to be very similar (on the order of ± 0.6 mag), but with opposite signs. So, by combining their photometry into a $BV.5$ system, the required S-corrections to transform the KAIT photometry to the g_{PS1} system are minimized and found to lie between -0.1 to 0.2 mag out to +60 d.

To compute an S-correction at a specific epoch, the observed and template spectra are first color-matched to CSP-II broadband photometry. The S-correction that transforms g_{CSP} -band natural-system photometry to the g_{PS1} -band natural system is computed by taking the difference between synthetic photometry computed using the two different passbands and the spectrum – that is, $m(g_{\text{PS1}}) = m(g_{\text{CSP}}) - \Delta S$. Given that the spectral energy distribution of SNe Ia is time dependent, S-corrections are computed for each epoch and then a smoothing spline is derived to give S-corrections at all times. Imaging was obtained using both the observed spectra and the Hsiao et al. spectral template. The S-corrections transforming our g_{CSP} - and $BV.5$ -band photometry to the g_{PS1} natural system are plotted in Fig. 1 (right, top panel) and listed in Table A.1. The observed g_{CSP} and $BV.5$ photometry is also plotted (right, bottom panel) with and without the S-corrections.

2.3. Light-curve parameters and reddening

Basic light-curve parameters were computed using the light-curve fitter SNooPy (Burns et al. 2011), and key results are listed in Table A.3. SNooPy fits provide an estimate of the time and magnitude of peak brightness, the light-curve decline-

rate parameter $\Delta m_{15}(B)$, and the color-stretch parameter s_{BV} (Burns et al. 2014), as well as host-galaxy reddening. Using the SNooPy “max model” $t_{B,\text{max}}$ is found to have occurred on $\text{MJD} = 56648.5 \pm 0.1$ with $m_B = 14.699 \pm 0.013$ mag, and $\Delta m_{15}(B) = 1.234 \pm 0.060$ mag. Fitting the entire suite of CSP-II optical/NIR light curves³ with SNooPy’s “EBV method2”, we obtain an $s_{BV} = 0.892 \pm 0.05$ mag, a host-galaxy color excess of $E(B - V)_{\text{host}} = 0.106 \pm 0.004_{\text{stat}} \pm 0.060_{\text{sys}}$ mag, and a distance modulus $\mu = 33.68 \pm 0.09$ mag. Combining $E(B - V)_{\text{host}}$ with the Milky Way component of $E(B - V)_{\text{MW}} = 0.049$ mag (Schlafly & Finkbeiner 2011; Schlegel et al. 1998), we obtain $E(B - V)_{\text{tot}} = 0.155 \pm 0.060$ mag, which is adopted throughout. Adjusting m_B for reddening and K-corrections, and assuming our adopted distance to the host, we find that SN 2013gy reached a peak absolute B -band magnitude $M_B = -19.32 \pm 0.16$ mag.

3. Results

3.1. Early g_{PS1} light curve and constraints on t_{first}

Plotted in the left panel of Fig. 2 is the early-time g_{PS1} -band light curve of SN 2013gy. Overplotted on the light curve is the best-fit single power-law function (i.e., $f(t) = A(t - t_{\text{first}})^\alpha$), a single power-law function with an index of two, and a linear function. The right, top panel shows the residuals between the best-fit model and the photometry. Finally, the bottom right panel of Fig. 2 contains the estimated error in the fit parameters and their correlation. Fitting was performed using the Markov chain Monte Carlo (MCMC) Python package PyMC3⁴ (Salvatier et al. 2016) and the No-U-Turn Sampler (NUTS; Hoffman & Gelman 2011). When fitting we considered only data obtained prior to -10 d (Conley et al. 2006; Ganeshalingam et al. 2011), and α was a free parameter with priors defined using the results of

³ A detailed analysis of a comprehensive ultraviolet, optical, and NIR dataset of SN 2013gy will be presented in a forthcoming publication.

⁴ <https://pymc-devs.github.io/pymc3>

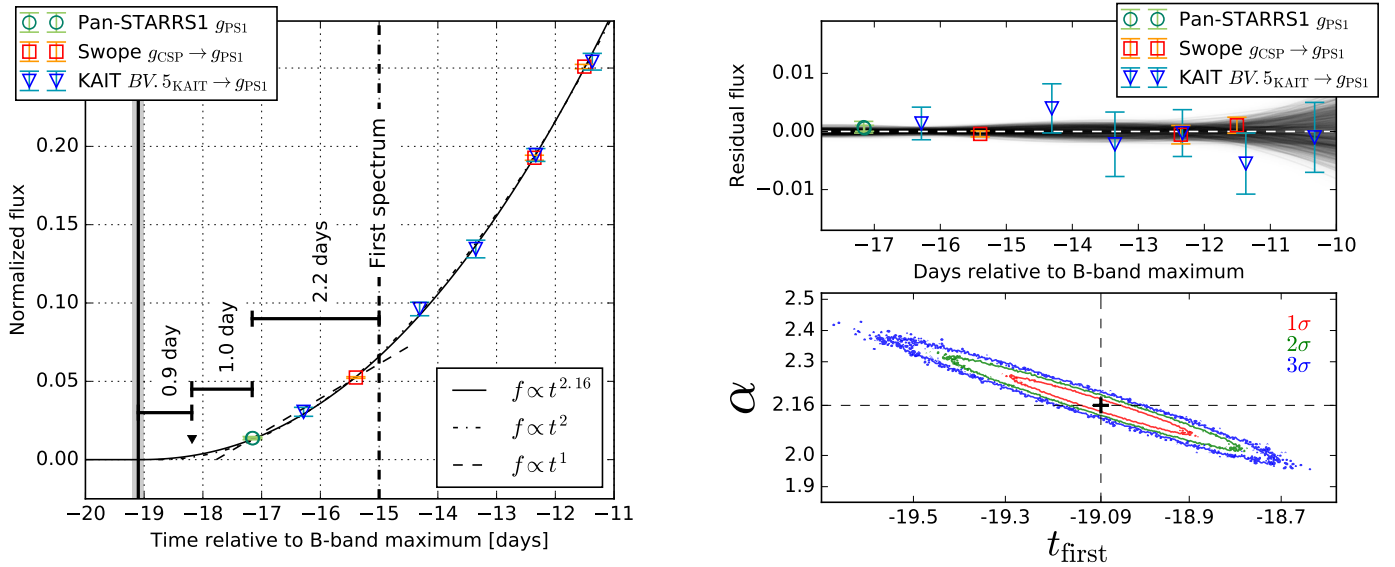


Fig. 2. *Left:* early-time g_{PS1} -band light curve of SN 2013gy based on photometry measured from PS1, CSP-II, and KAIT images, plotted along with the best-fit, single power-law function characterized by an index $\alpha = 2.16$ (solid curve), the “fireball” model ($\alpha = 2$, dash-dot curve), and a linear rise for the first ~ 3 d (dashed curve). The KAIT open-filter nondetection obtained 1 d prior to discovery is plotted as a black filled triangle. The solid vertical black line indicates t_{first} with its width corresponding to the 1σ uncertainty (gray region), and the vertical dashed line corresponds to the epoch of the first spectrum. *Right, top:* residuals between the best-fit model and the data (horizontal dashed white line); black lines correspond to models sampled from below (right, bottom panel), which contains 1σ , 2σ , and 3σ error ellipses indicating the uncertainties in the model fit parameters α and t_{first} .

Firth et al. (2015). The best-fit power law corresponds to $\alpha = 2.16^{+0.06}_{-0.06}$ and $t_{\text{first}} = 56629.4^{+0.1}_{-0.1}$ MJD. We note that from examination of the right, bottom panel of Fig. 2, a value of $\alpha = 2.0$ is found to be within 2σ of the best-fit value. Assuming $\alpha = 2.0$ would imply a t_{first} value that is 7 h later than our inferred best-fit value. We also note that because the first 3 epochs span more than three days and the uncertainty of the first KAIT measurement is relatively large, we cannot rule out a linear rise prior to ~ 15 d before $t_{B,\text{max}}$. The best fit to a linear rise implies a value of t_{first} that is 1.3 d later and only 14 h before the PS1 images. Comparison of the best-fit t_{rise} to the first PS1 detection indicates that SN 2013gy was discovered $1.93^{+0.12}_{-0.13}$ d past t_{first} , and when compared to $t_{B,\text{max}}$, one obtains $t_{\text{rise}} = -19.10^{+0.12}_{-0.13}$ days.

3.2. Spectroscopy and an estimate of $t_{\text{explosion}}$

Spectra of SN 2013gy taken around maximum light resemble those of normal SNe Ia. Pseudo-equivalent-width (pEW) measurements of the Si II $\lambda 5972$ and superior fit and implies $\lambda 6355$ features indicate that it is “core normal” (CN) according to Branch et al. (2006). Moreover, the position of the absorption minimum of the Si II $\lambda 6355$ feature in the -0.8 d spectrum indicates a Doppler velocity at maximum absorption of $-\nu_{\text{abs}} = 10\,180 \pm 90$ km s $^{-1}$, consistent with a normal object in the Wang et al. (2009) system.

Piro & Nakar (2013) suggest that an estimate of the explosion time ($t_{\text{explosion}}$) can be obtained through the evolution of Si II $\lambda 6355$. This is achieved by fitting the early-time velocity evolution with an appropriate power law. Following Piro & Nakar, we fit a $v \propto t^\beta$ power law with $\beta = -0.20$, -0.22 , and -0.24 . The resulting goodness-of-fit parameter χ^2 per degree of freedom ($\chi^2/\text{d.o.f.}$) is plotted in the top-left panel of Fig. 3 for each power-law index. A constant systematic error is added to get a best fit with $\chi^2/\text{d.o.f.} = 1$. We find that $\beta = -0.20$ provides a superior fit and implies $t_{\text{explosion}} = 56\,629.45^{+0.28}_{-0.31}$ MJD.

Comparison of t_{first} and $t_{\text{explosion}}$ reveals that SN 2013gy likely experienced minimal to no dark phase.

3.3. Constraints on interaction with a nondegenerate companion

Companion radius (R_c) limits are computed following the methodology of Shappee et al. (2018), which is based on comparing the observations to the analytical model predictions of Kasen (2010). Taking a conservative explosion time to be ~ 2 d before discovery, mass of the companion to be $1 M_\odot$ and assuming a favorable viewing angle of 15° , we compute R_c limits using both the discovery g_{PS1} photometry and the previous KAIT nondetection. Our limits on R_c are plotted vs. days relative to the epoch of discovery in the top-right panel of Fig. 3. Combined limits indicate $R_c \leq 4 R_\odot$ in the case of a favorable viewing angle. Adopting the TF distances as upper and lower bounds on the distance, results in $R_c \sim 3.4^{+8.5}_{-2.1} R_\odot$ for the estimated value of t_{first} . The same analysis has been done a number of times for other SNe, yielding possible companion radii limits ranging between a fraction of R_\odot and tens of R_\odot (Bloom et al. 2012; Foley et al. 2012; Silverman et al. 2012; Goobar et al. 2015; Olling et al. 2015; Im et al. 2015; Marion et al. 2016; Shappee et al. 2016, 2018; Hosseinzadeh et al. 2017; Miller et al. 2018).

3.4. Limits on the presence of companion material

Hydrodynamic simulations of ejecta-companion interactions indicate that $>0.15 M_\odot$ of material could be removed from hydrogen-rich (H-rich) Roche-lobe overflow (RLOF) companions via stripping and/or momentum-transfer or ablation and/or heating and is expected to have a velocity dispersion of $v \approx 1000$ km s $^{-1}$ (e.g., Marietta et al. 2000; Pan et al. 2012; Boehner et al. 2018). This unbound material should be visible

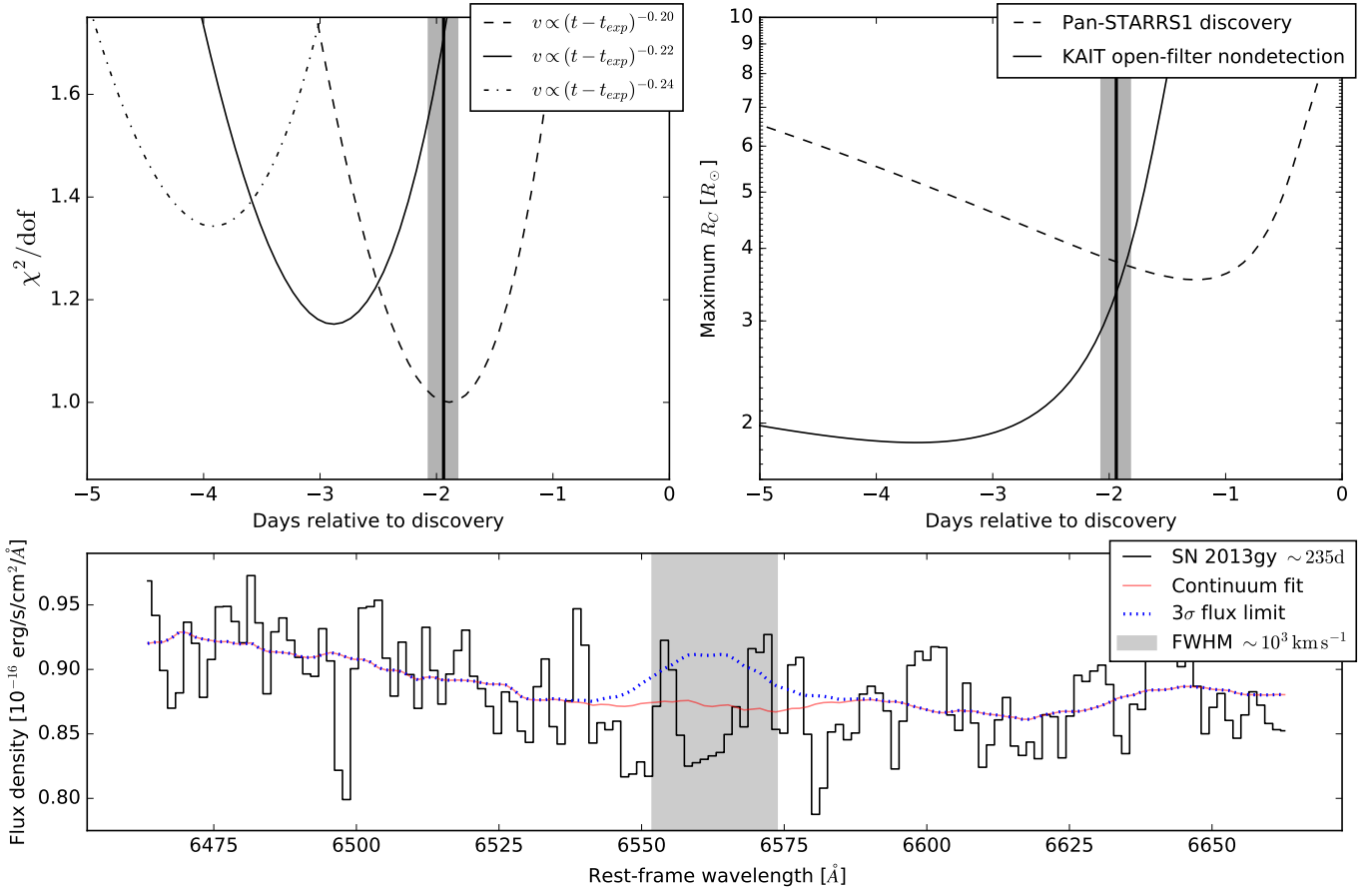


Fig. 3. *Left:* estimate of $t_{\text{explosion}}$ based upon functional fits to the velocity evolution of Si II $\lambda 6355$. *Right:* maximum allowed radius of a companion star as a function of explosion time as inferred from the discovery g_{PS1} photometry and the previous KAIT nondetection limit plotted vs. days since discovery. The vertical black line in the left and right panels corresponds to the best-fit t_{first} , accompanied by a shaded gray region corresponding to its 1σ uncertainty. *Bottom:* rest-frame-corrected nebular spectrum of SN 2013gy (black line), the continuum fit (red), and our 3σ H α flux limit (blue line). The shaded area extends over the expected location of H α ± 1000 km s $^{-1}$.

during the nebular phase – when the ejecta become optically thin – and lead to a prominent emission signature.

To place a limit on the H α emission we have followed conventional techniques in the literature (e.g., Leonard 2007; Shappee et al. 2013). Briefly, the late-phase (+235 d) spectrum of SN 2013gy was flux calibrated to match photometry computed from PS1 stars (Flewelling et al. 2016) located in an R -band acquisition image. Next, the continuum was fitted in the vicinity of H α and subtracted from the spectrum, yielding no statistical evidence for H α emission as demonstrated in the bottom panel of Fig. 3. Following Leonard (2007), we obtain a 3σ statistical flux limit of 4.5×10^{-16} ergs s $^{-1}$ cm $^{-2}$ \AA^{-1} . For our adopted distance and using the models of Botyánszki et al. (2018), this corresponds to an H-mass limit of $<0.018 M_{\odot}$. This limit is an order of magnitude lower than all estimates of unbound mass values for normal SNe Ia, effectively ruling out H-rich RLOF companions. This finding is consistent with similar constraints of a number of other SNe Ia discovered soon after explosion (Mattila et al. 2005; Leonard 2007; Lundqvist et al. 2013, 2015; Shappee et al. 2013, 2018; Maguire et al. 2016; Graham et al. 2017; Sand et al. 2018; Dimitriadis et al. 2019). This leaves only a small parameter space of subdwarfs and He stars whose signatures could remain undetected in the spectrum. We note that inspection of an unpublished early-phase NIR spectral time-series of SN 2013gy obtained by the CSP-II yields no indication of He lines, and unfortunately, no late-phase NIR spectrum was obtained.

4. Discussion

We have presented discovery and follow-up photometry of the normal SN Ia 2013gy. Photometry obtained from different facilities was carefully calibrated to the g_{PS1} system. The early rise of the light curve is well fit with a single power-law function, indicating that SN 2013gy was discovered within 48 h of t_{first} . The corresponding t_{first} value is consistent with a $t_{\text{explosion}}$ estimate based on the Doppler velocity evolution of the Si II $\lambda 6355$ feature, suggesting minimal to no dark phase for SN 2013gy. A short dark phase has been inferred from the study of several other supernovae including SN 2009ig, SN 2011fe, SN 2012cg (Piro & Nakar 2014), and iPTF 16abc (Miller et al. 2018). In contrast, the transitional iPTF 13ebh may have experienced an approximately four-day dark phase (see Hsiao et al. 2015).

Examination of the pEW values of the Si II $\lambda 5972$ and $\lambda 6355$ doublets shows that SN 2013gy is core normal (Branch et al. 2006, 2009), while its Si II $\lambda 6355$ Doppler velocity makes it a normal object according to the classification scheme of Wang et al. (2009). Moreover, as discussed by Stritzinger et al. (2018), the early intrinsic $(B - V)_0$ colors reveal it to be a red and rapidly evolving object, similar to SN 2011fe. In summary, SN 2013gy displays all of the characteristics of a normal SN Ia.

Making use of the early discovery of SN 2013gy, limits were placed on the radius of any companion star of $R \leq$

$4 M_{\odot}$. Combined with our stripped H-mass limit of $\lesssim 0.018 M_{\odot}$ obtained from a nebular-phase *Magellan* spectrum, our analysis effectively rules out H-rich RLOF companions. This result highlights that it is possible to constrain existing SN Ia progenitor models with the use of both early and late-phase observations.

Acknowledgements. A special thanks to the referee for detailed comments to the submitted manuscript. We thank P. Massey and M. Kasliwal for obtaining a spectrum each with the *Magellan* telescopes. S. Holmbo and M. D. Stritzinger are supported by a research grant (13261) from the VILLUM FONDEN. A. V. Filippenko's supernova group is grateful for financial assistance from USA National Science Foundation (NSF) grant AST-1211916, the TABASGO Foundation, the Christopher R. Redlich Fund, and the Miller Institute for Basic Research in Science (U. C. Berkeley). Research at Lick Observatory is partially supported by a generous gift from Google. P. Höflich acknowledges NSF grants AST-1715133 and AST-1613472. L. Tomasella is partially supported by the PRIN-INAF 2016 with the project "Towards the SKA and CTA era: discovery, localisation, and physics of transient sources" (P.I. M. Giroletti). The CSP-II has been funded by the USA NSF under grants AST-0306969, AST-0607438, AST-1008343, AST-1613426, AST-1613455, and AST-1613472, and in part by a Sapere Aude Level 2 grant funded by the Danish Agency for Science and Technology and Innovation (PI Stritzinger). J.-J. Zhang is supported by the National Natural Science Foundation of China (NSFC, grants 11403096 and 11773067), the Key Research Program of the CAS (grant KJZD-EW-M06), the Youth Innovation Promotion Association of the CAS (grant 2018081), and the CAS Light of West China Program. X. Wang is supported by the National Natural Science Foundation of China (NSFC grants 11325313 and 11633002), and the National Program on Key Research and Development Project (grant 2016YFA0400803). Based in part on observations made with the Nordic Optical Telescope (P49-017; PI Stritzinger), operated by the Nordic Optical Telescope Scientific Association at the Observatorio del Roque de los Muchachos, La Palma, Spain, of the Instituto de Astrofísica de Canarias. This paper includes data gathered with the 6.5 m *Magellan* telescopes located at Las Campanas Observatory, Chile. We acknowledge the support of the staff of Lijiang 2.4 m and the Xinglong 2.16 m telescopes. Partially based on observations made with the *Copernico* 1.8 m telescope (Asiago, Italy) operated by INAF Osservatorio Astronomico di Padova. This paper made use of discovery images obtained by the Pan-STARRS1 Survey (PS1), which was made possible through contributions of the Institute for Astronomy, the University of Hawaii, the Pan-STARRS Project Office, the Max-Planck Society and its participating institutes, the Max Planck Institute for Astronomy, Heidelberg and the Max Planck Institute for Extraterrestrial Physics, Garching, The Johns Hopkins University, Durham University, the University of Edinburgh, Queen's University Belfast, the Harvard-Smithsonian Center for Astrophysics, the Las Cumbres Observatory Global Telescope Network Incorporated, the National Central University of Taiwan, the Space Telescope Science Institute, the National Aeronautics and Space Administration (NASA) under grant NNX08AR22G issued through the Planetary Science Division of the NASA Science Mission Directorate, the NSF under grant AST-1238877, the University of Maryland, and Eotvos Lorand University (ELTE).

References

- Bloom, J. S., Kasen, D., Shen, K. J., et al. 2012, *ApJ*, 744, L17
 Boehner, P., Plewa, T., & Langer, N. 2018, *MNRAS*, 465, 2060
 Botyánszki, J., Kasen, D., & Plewa, T. 2018, *ApJ*, 852L, 6
 Branch, D., Dang, L. C., & Hall, N. 2006, *PASP*, 118, 560
 Branch, D., Dang, L. C., & Baron, E. 2009, *PASP*, 121, 238
 Brown, J. S., Stanek, K. Z., Vallely, P., et al. 2018, *ATel*, 11253, 00
 Burns, C. R., Stritzinger, M. D., Phillips, M. M., et al. 2011, *AJ*, 141, 19
 Burns, C. R., Stritzinger, M. D., Phillips, M. M., et al. 2014, *ApJ*, 789, 32
 Cartier, R., Sullivan, M., Firth, R. E., et al. 2017, *MNRAS*, 464, 4476
 Catinella, B., Haynes, M. P., & Giovanelli, R. 2005, *AJ*, 130, 1037
 Chambers, K. C., Magnier, E. A., & Metcalfe, N. 2016, ArXiv e-prints [arXiv:1612.05560]
 Conley, A., Howell, D. A., Howes, A., et al. 2006, *AJ*, 132, 1707
 Contreras, C., Phillips, M. M., Burns, C. R., et al. 2018, *ApJ*, 859, 24
 Cornect, R., Brimacombe, J., Stone, G., et al. 2018, *ATel*, 11259
 Dimitriadis, G., Rojas-Bravo, C., Kilpatrick, C. D., et al. 2019, *ApJ*, 870, L14
 Filippenko, A. V., Li, W. D., Treffers, R. R., & Modjaz, M. 2001, in *Small Telescope Astronomy on Global Scales*, eds. W. P. Chen, C. Lemme, & B. Paczyński (San Francisco: ASP), 246, 121
 Firth, R. E., Sullivan, M., Gal-Yam, A., et al. 2015, *MNRAS*, 446, 3895
 Flewelling, H. A., Magnier, E. A., & Chambers, K. C. 2016, ArXiv e-prints [arXiv:1612.05243]
 Foley, R. J., Narayan, G., Challis, P. J., et al. 2010, *ApJ*, 708, 1748
 Foley, R. J., Challis, P. J., Filippenko, A. V., et al. 2012, *ApJ*, 744, 38
 Ganeshalingam, M., Li, W., & Filippenko, A. V. 2011, *MNRAS*, 416, 2607
 Goobar, A., Kromer, M., Siverd, R., et al. 2015, *ApJ*, 799, 106
 Graham, M. L., Kumar, S., Hosseinzadeh, G., et al. 2017, *MNRAS*, 472, 3437
 Hamuy, M., Folatelli, G., Morrell, N. I., et al. 2006, *PASP*, 118, 2
 Hoffman, M. D., & Gelman, A. 2011, ArXiv e-prints [arXiv:1111.4246]
 Hosseinzadeh, G., Sand, D. J., Valenti, S., et al. 2017, *ApJ*, 845, L11
 Hoyle, F., & Fowler, W. A. 1960, *ApJ*, 132, 565
 Hsiao, E. Y., Conley, A., Howell, D. A., et al. 2007, *ApJ*, 663, 1187
 Hsiao, E. Y., Burns, C. R., Contreras, C., et al. 2015, *A&A*, 578, A9
 Im, M., Choi, C., Yoon, S.-C., et al. 2015, *ApJS*, 221, 22
 Jiang, J., Doi, M., Maeda, K., et al. 2017, *Nature*, 550, 80
 Kasen, D. 2010, *ApJ*, 708, 1025
 Kim, H., Zheng, W. K., Li, W., et al. 2013, *CBET*, 3743, 1
 Krisciunas, K., Contreras, C., Burns, C. R., et al. 2017, *AJ*, 154, 211
 Leonard, D. C. 2007, *ApJ*, 670, 1275
 Lundqvist, P., Mattila, S., Sollerman, J., et al. 2013, *MNRAS*, 435, 329
 Lundqvist, P., Nyholm, A., Taddia, F., et al. 2015, *A&A*, 577, A39
 Maeda, K., Jiang, J.-A., Shigejima, T., & Doi, M. 2018, *ApJ*, 861, 78
 Maguire, K., Taubenberger, S., Sullivan, M., & Mazzali, P. A. 2016, *MNRAS*, 457, 3254
 Marietta, E., Burrows, A., & Fryxell, B. 2000, *ApJS*, 128, 615
 Marion, G. H., Brown, P. J., Vinkó, J., et al. 2016, *ApJ*, 820, 92
 Mattila, S., Lundqvist, P., Sollerman, J., et al. 2005, *A&A*, 443, 649
 Miller, A. A., Cao, Y., Piro, A. L., et al. 2018, *ApJ*, 852, 100
 Noebauer, U. M., Taubenberger, S., Blinnikov, S., Sorokina, E., & Hillebrandt, W. 2016, *MNRAS*, 463, 2972
 Olling, R. P., Mushotzky, R., Shaya, E. J., et al. 2015, *Nature*, 521, 332
 Pan, K.-C., Ricker, P. M., & Taam, R. E. 2012, *ApJ*, 760, 21
 Pan, Y.-C., Foley, R. J., Kromer, M., et al. 2015, *MNRAS*, 452, 4307
 Phillips, M. M., Contreras, C., Hsiao, E. Y., et al. 2019, *PASP*, 131, 014001
 Piro, A. L., & Morozova, V. S. 2016, *ApJ*, 826, 96
 Piro, A. L., & Nakar, E. 2013, *ApJ*, 769, 67
 Piro, A. L., & Nakar, E. 2014, *ApJ*, 784, 85
 Piro, A. L., Chang, P., & Weinberg, N. N. 2010, *ApJ*, 708, 598
 Polin, A., Nugent, P., & Kasen, D. 2019, *ApJ*, 873, 1
 Rabinak, I., Livne, E., & Waxman, E. 2012, *ApJ*, 757, 35
 Salvatier, J., Wiecki, T. V., & Fonnesbeck, C. 2016, *Astrophysics Source Code Library* [record ascl:1610.016]
 Sand, D. J., Graham, M. L., Botyánszki, J., et al. 2018, *ApJ*, 863, 24
 Schlafly, E. F., & Finkbeiner, D. P. 2011, *ApJ*, 737, 103
 Schlegel, D. J., Finkbeiner, D. P., & Davis, M. 1998, *ApJ*, 500, 525
 Shappee, B. J., Stanek, K. Z., Pogge, R. W., & Garnavich, P. M. 2013, *ApJ*, 762, L5
 Shappee, B. J., Piro, A. L., Holoién, T. W.-S., et al. 2016, *ApJ*, 826, 144
 Shappee, B. J., Piro, A. L., Stanek, K. Z., et al. 2018, *ApJ*, 855, 6
 Shappee, B. J., Holoién, T. W.-S., Drout, M. R., et al. 2019, *ApJ*, 870, 13
 Silverman, J. M., Ganeshalingam, M., Cenko, S. B., et al. 2012, *ApJ*, 756, L7
 Siverd, R. J., Goobar, A., Stassun, K. G., et al. 2015, *ApJ*, 799, 105
 Smartt, S., Smith, K. W., & Wright, D. 2014, *ATel*, 5850
 Springob, C. M., Masters, K. L., Haynes, M. P., Giovanelli, R., & Marinoni, C. 2007, *ApJS*, 172, 599
 Stritzinger, M., Hamuy, M., Suntzeff, N. B., et al. 2002, *AJ*, 124, 2100
 Stritzinger, M. D., Phillips, M. M., Boldt, L. N., et al. 2011, *AJ*, 142, 156
 Stritzinger, M. D., Shappee, B. J., Piro, A., et al. 2018, *ApJ*, 864, L35
 Theureau, G., Hanski, M. O., Courdreaux, N., Hallet, N., & Martin, J.-M. 2007, *A&A*, 465, 71
 Tomasella, L., Benetti, S., & Tartaglia, L. 2013, *ATel*, 5638
 Tonry, J. L., Stubbs, C. W., Lykke, K. R., et al. 2012, *ApJ*, 750, 99
 Tucker, M. A., Shappee, B. J., & Wisniewski, J. P. 2019, *ApJ*, 872, L22
 Wang, X., Filippenko, A. V., Ganeshalingam, M., et al. 2009, *ApJ*, 699, L139
 Zhang, K., Wang, X., Zhang, J.-J., et al. 2016, *ApJ*, 820, 67
 Zheng, W., Silverman, J., Filippenko, A. V., et al. 2013, *ApJ*, 778, L15
 Zheng, W., Shivvers, I., Filippenko, A. V., et al. 2014, *ApJ*, 783, L24
 Zheng, W., Kelly, P. L., & Filippenko, A. V. 2017, *ApJ*, 848, 66

Appendix A: Additional material

Table A.1. continued.

Table A.1. Photometry of SN 2013gy and S-corrections.

MJD	Phase ^(a)	Magnitude	S-correction	g_{PS1}
<i>g_{PS1}</i> band				
56631.34	-17.16	19.483 ± 0.079	...	19.483 ± 0.079
56631.35	-17.15	19.482 ± 0.040	...	19.482 ± 0.040
<i>g_{CSP}</i> band				
56633.22	-15.28	18.062 ± 0.011	0.036	18.026 ± 0.011
56636.26	-12.24	16.622 ± 0.009	0.010	16.612 ± 0.009
56637.11	-11.39	16.330 ± 0.006	0.004	16.326 ± 0.006
56639.11	-9.39	15.769 ± 0.008	-0.006	15.775 ± 0.008
56640.17	-8.33	15.555 ± 0.007	-0.009	15.564 ± 0.007
56643.15	-5.35	15.104 ± 0.008	-0.014	15.118 ± 0.008
56644.14	-4.36	15.011 ± 0.007	-0.014	15.025 ± 0.007
56645.13	-3.37	14.938 ± 0.007	-0.013	14.951 ± 0.007
56646.12	-2.38	14.890 ± 0.009	-0.011	14.901 ± 0.009
56647.16	-1.34	14.857 ± 0.007	-0.009	14.866 ± 0.007
56648.13	-0.37	14.819 ± 0.007	-0.006	14.825 ± 0.007
56649.11	0.61	14.826 ± 0.006	-0.002	14.828 ± 0.006
56650.15	1.65	14.830 ± 0.007	0.002	14.828 ± 0.007
56651.10	2.60	14.850 ± 0.008	0.006	14.844 ± 0.008
56652.10	3.60	14.895 ± 0.009	0.011	14.884 ± 0.009
56653.13	4.63	14.925 ± 0.005	0.016	14.909 ± 0.005
56654.18	5.68	14.982 ± 0.006	0.022	14.960 ± 0.006
56655.19	6.69	15.040 ± 0.009	0.028	15.012 ± 0.009
56656.20	7.70	15.159 ± 0.009	0.034	15.125 ± 0.009
56657.21	8.71	15.202 ± 0.010	0.039	15.163 ± 0.010
56658.14	9.64	15.266 ± 0.008	0.045	15.221 ± 0.008
56659.15	10.65	15.357 ± 0.006	0.051	15.306 ± 0.006
56660.11	11.61	15.448 ± 0.008	0.056	15.392 ± 0.008
56661.16	12.66	15.534 ± 0.007	0.061	15.473 ± 0.007
56662.12	13.62	15.663 ± 0.007	0.066	15.597 ± 0.007
56663.11	14.61	15.741 ± 0.007	0.071	15.670 ± 0.007
56664.10	15.60	15.870 ± 0.010	0.075	15.795 ± 0.010
56665.08	16.58	15.955 ± 0.008	0.079	15.876 ± 0.008
56666.06	17.56	16.071 ± 0.011	0.083	15.988 ± 0.011
56667.07	18.57	16.154 ± 0.010	0.086	16.068 ± 0.010
56669.06	20.56	16.359 ± 0.013	0.090	16.269 ± 0.013
56670.14	21.64	16.491 ± 0.013	0.092	16.399 ± 0.013
56671.06	22.56	16.562 ± 0.011	0.093	16.469 ± 0.011
56672.12	23.62	16.662 ± 0.009	0.094	16.568 ± 0.009
56673.09	24.59	16.739 ± 0.009	0.095	16.644 ± 0.009
56675.05	26.55	16.893 ± 0.008	0.095	16.798 ± 0.008
56676.11	27.61	16.970 ± 0.007	0.095	16.875 ± 0.007
56677.04	28.54	17.033 ± 0.009	0.094	16.939 ± 0.009
56678.11	29.61	17.089 ± 0.009	0.094	16.995 ± 0.009
56679.05	30.55	17.166 ± 0.006	0.093	17.073 ± 0.006
56680.12	31.62	17.233 ± 0.008	0.092	17.141 ± 0.008
56681.05	32.55	17.293 ± 0.008	0.091	17.202 ± 0.008
56684.09	35.59	17.430 ± 0.008	0.088	17.342 ± 0.008
56685.07	36.57	17.484 ± 0.008	0.086	17.398 ± 0.008
56686.05	37.55	17.492 ± 0.010	0.085	17.407 ± 0.010
56687.06	38.56	17.511 ± 0.009	0.084	17.427 ± 0.009
<i>BV.5</i> band				
56690.08	41.58	17.617 ± 0.010	0.081	17.536 ± 0.010
56692.08	43.58	17.674 ± 0.011	0.079	17.595 ± 0.011
56694.06	45.56	17.696 ± 0.009	0.078	17.618 ± 0.009
56696.07	47.57	17.733 ± 0.012	0.077	17.656 ± 0.012
56703.07	54.57	17.906 ± 0.019	0.077	17.829 ± 0.019
56708.06	59.56	17.968 ± 0.011	0.080	17.888 ± 0.011
56632.33	-16.17	18.703 ± 0.108	0.090	18.613 ± 0.108
56634.31	-14.19	17.477 ± 0.053	0.109	17.367 ± 0.053
56635.26	-13.24	17.120 ± 0.050	0.116	17.003 ± 0.050
56636.28	-12.22	16.725 ± 0.025	0.122	16.602 ± 0.025
56637.25	-11.25	16.439 ± 0.025	0.126	16.312 ± 0.025
56638.29	-10.21	16.125 ± 0.022	0.130	15.995 ± 0.022
56639.32	-9.18	15.868 ± 0.022	0.132	15.737 ± 0.022
56640.29	-8.21	15.647 ± 0.029	0.133	15.514 ± 0.029

MJD	Phase ^(a)	Magnitude	S-correction	g_{PS1}
56641.38	-7.12	15.426 ± 0.025	0.133	15.293 ± 0.025
56645.36	-3.14	15.049 ± 0.019	0.130	14.920 ± 0.019
56647.32	-1.18	14.994 ± 0.014	0.128	14.865 ± 0.014
56648.31	-0.19	14.969 ± 0.011	0.128	14.841 ± 0.011
56649.30	0.80	14.966 ± 0.011	0.128	14.839 ± 0.011
56650.32	1.82	14.958 ± 0.014	0.128	14.830 ± 0.014
56651.27	2.77	14.996 ± 0.011	0.128	14.868 ± 0.011
56652.26	3.76	15.028 ± 0.014	0.127	14.901 ± 0.014
56653.28	4.78	15.051 ± 0.014	0.127	14.924 ± 0.014
56654.26	5.76	15.098 ± 0.011	0.126	14.972 ± 0.011
56655.21	6.71	15.160 ± 0.011	0.124	15.036 ± 0.011
56656.28	7.78	15.214 ± 0.015	0.122	15.093 ± 0.015
56657.23	8.73	15.280 ± 0.015	0.118	15.162 ± 0.015
56658.24	9.74	15.358 ± 0.015	0.114	15.244 ± 0.015
56659.25	10.75	15.435 ± 0.015	0.109	15.326 ± 0.015
56660.21	11.71	15.523 ± 0.025	0.104	15.419 ± 0.025
56661.23	12.73	15.601 ± 0.032	0.097	15.504 ± 0.032
56663.22	14.72	15.782 ± 0.017	0.081	15.701 ± 0.017
56665.23	16.73	15.959 ± 0.023	0.063	15.896 ± 0.023
56667.24	18.74	16.110 ± 0.031	0.032	16.078 ± 0.031
56672.25	23.75	16.470 ± 0.045	-0.110	16.580 ± 0.045
56674.24	25.74	16.596 ± 0.029	-0.136	16.732 ± 0.029
56676.23	27.73	16.727 ± 0.030	-0.151	16.878 ± 0.030
56678.22	29.72	16.845 ± 0.028	-0.161	17.006 ± 0.028
56680.20	31.70	16.986 ± 0.033	-0.166	17.152 ± 0.033
56682.19	33.69	17.086 ± 0.030	-0.168	17.254 ± 0.030
56684.19	35.69	17.192 ± 0.032	-0.167	17.358 ± 0.032
56689.16	40.66	17.389 ± 0.054	-0.153	17.542 ± 0.054
56702.14	53.64	17.824 ± 0.111	-0.085	17.909 ± 0.111

Notes. ^(a)Temporal phase given in days with respect to the time of *B*-band maximum light on MJD = 56648.5 ± 0.10, i.e., 2013 December 22.

Table A.2. Journal of spectroscopic observations.

Date	MJD	Phase ^(a)	Telescope ^(b)	Instrument	$-\nu_{\text{abs, Si II } \lambda 6355}$ (km s ⁻¹)
2013 Dec 07.5	56633.5	-15.0	FTS	FLOYDS ^c	14 030 ⁺¹⁹⁸ ₋₁₀₄
2013 Dec 07.7	56633.7	-14.8	LJT	YFOSC	13 709 ⁺²⁷⁴ ₋₁₅₁
2013 Dec 07.8	56633.8	-14.7	<i>Copernico</i>	AFOSC	13 732 ⁺⁵¹⁰ ₋₁₃₂
2013 Dec 08.5	56634.5	-14.0	FTS	FLOYDS ^c	13 756 ⁺⁹⁰ ₋₇₁
2013 Dec 08.8	56634.8	-13.7	LJT	YFOSC	13 322 ⁺³²⁶ ₋₁₄₂
2013 Dec 09.7	56635.7	-12.8	LJT	YFOSC	13 067 ⁺¹⁷⁰ ₋₆₁
2013 Dec 14.0	56640.0	-8.5	<i>Magellan Clay</i>	MagE	11 509 ⁺⁷⁶ ₋₃₃
2013 Dec 15.4	56641.4	-7.1	FTS	FLOYDS ^c	11 174 ⁺¹⁴⁶ ₋₆₆
2013 Dec 17.4	56643.4	-5.1	FTS	FLOYDS ^c	10 645 ⁺²⁸ ₋₂₈
2013 Dec 20.5	56646.5	-2.0	FTS	FLOYDS ^c	10 673 ⁺²²² ₋₂₄
2013 Dec 20.7	56646.7	-1.8	LJT	YFOSC	10 631 ⁺⁵² ₋₆₆
2013 Dec 21.7	56647.7	-0.8	LJT	YFOSC	10 178 ⁺⁹⁰ ₋₄₇
2013 Dec 25.8	56651.8	+3.3	NOT	ALFOSC	10 149 ⁺⁹⁰ ₋₇₁
2013 Dec 27.7	56653.7	+5.2	XLT	BFOSC	9862 ⁺⁴¹⁵ ₋₈₅
2013 Dec 28.3	56654.3	+5.8	Keck	LRIS	9913 ⁺⁷¹ ₋₁₄
2013 Dec 31.2	56657.2	+8.7	<i>Magellan Baade</i>	IMACS	9734 ⁺⁹⁴ ₋₅₇
2013 Dec 31.7	56657.7	+9.2	LJT	YFOSC	9682 ⁺⁴⁷ ₋₅₇
2014 Jan 03.6	56660.6	+12.1	XLT	BFOSC	10 971 ⁺⁸³⁶ ₋₈₅₀
2014 Jan 07.6	56664.6	+16.1	LJT	YFOSC	9267 ⁺¹⁷⁰ ₋₉₄
2014 Jan 08.3	56665.3	+16.8	<i>du Pont</i>	WFCCD	9059 ⁺⁹⁹ ₋₉₉
2014 Jan 14.0	56671.0	+22.5	NOT	ALFOSC	8705 ⁺³⁰² ₋₁₂₃
2014 Jan 23.3	56680.3	+31.8	<i>Shane</i>	Kast	8134 ⁺⁹⁴ ₋₃₃
2014 Feb 09.9	56697.9	+49.4	NOT	ALFOSC	7487 ⁺⁴³⁹ ₋₄₈₂
2014 Aug 14.5	56883.5	+235.0	<i>Magellan Baade</i>	IMACS	...

Notes. ^(a)Temporal phase given in days with respect to the time of *B*-band maximum light on MJD = 56648.5 ± 0.10, i.e., 2013 December 22. ^(b)FTS = 2.0 m Faulkes Telescope South, LJT = 2.4 m Li-Jiang Telescope, *Copernico* = INAF Osservatorio Astronomico di Padova 1.8 m Telescope, NOT = 2.5 m Nordic Optical Telescope, XLT = 2.6 m Xing-Long Telescope. ^(c)Spectra were published by [Graham et al. \(2017\)](#).

Table A.3. Key parameters of SN 2013gy.

$t_{B\text{max}}$ [MJD]	$\Delta m_{15}(B)$ [mag]	s_{BV}	$E(B-V)_{\text{tot}}$ [mag]	μ ^(a) [mag]	M_B [mag]	t_{first} [MJD]	α	t_{rise} [days]
56648.5 ± 0.10	1.234 ± 0.060	0.892 ± 0.05	0.106 ± 0.060	33.75 ± 0.15	-19.32 ± 0.16	56629.4 ± 0.13	2.16 ^{+0.06} _{-0.06}	19.10 ^{+0.12} _{-0.13}

Notes. ^(a)Based on red-shift assuming $H_0 = 73$ km s⁻¹ and adopting a correction based on a Virgo, Great Attractor and Shapley infall model.

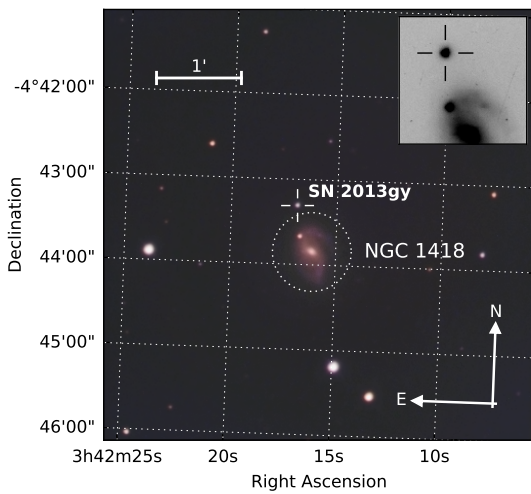


Fig. A.1. Composite image of NGC 1418 (north up and east left) constructed from multiband images obtained with the Swope telescope. The inset contains a close-up view of the position of SN 2013gy, highlighting its position in the outer region of the host.

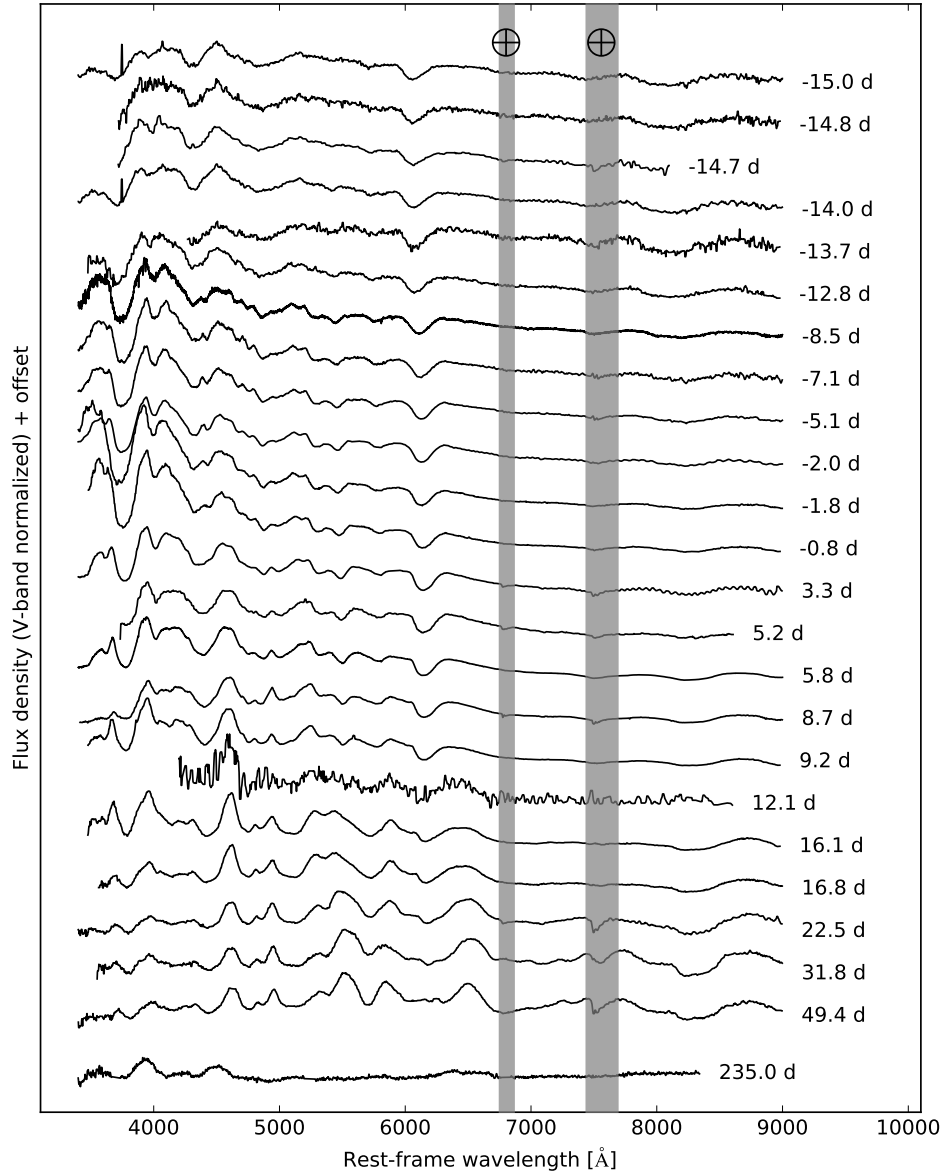


Fig. A.2. Spectroscopic time series of SN 2013gy used to compute S-corrections, enabling us to transform natural-system g_{CSP} - and $BV.5$ -band photometry to the g_{PS1} natural system.