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Abstract: In this report, we show evidence for ripple and dune migration in Herschel Crater on Mars. We estimate an average dune migration of 0.8 meters and a minimum ripple migration of 1.1 meters in a time span of 3.7 Earth-Years.

These dunes and ripples are mainly shaped by prevailing winds coming from the north, however we also report the presence of secondary winds enhanced by the crater rim at regional scale and deflected by the dune topography at the dune scale.

These last are predicted by the Mars Regional Atmospheric Modeling System (MRAMS), an atmospheric mesoscale model, while the dominant flows from the north are underestimated. Modeled winds at the local scale refer that a multi directional wind regime is indicated as the first cause of the diverse set of ripples overlapping one on other.

For the first time, a survey integrating the assessment of dune and ripple migration is presented, showing how dune topography can influence the migration patterns of ripples and how underlying regional topography seems to control the rates of dune migration.

The migration patterns suggest that the prevailing winds from the north are locally-deflected winds (blowing from the NNW and from the NNE before deflection).

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A handwritten signature in black ink that reads "Marco Cardinale". The signature is written in a cursive, flowing style.

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1 **Present-day aeolian activity**
2 **in Herschel Crater, Mars.**

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47

48 **Abstract**

49 In this report, we show evidence for ripple and dune migration in Herschel Crater on Mars.

50 We estimate an average dune migration of 0.8 meters and a minimum ripple migration of 1.1 meters
51 in a time span of 3.7 Earth-Years.

52 These dunes and ripples are mainly shaped by prevailing winds coming from the north, however we
53 also report the presence of secondary winds enhanced by the crater rim at regional scale and
54 deflected by the dune topography at the dune scale.

55 These last are predicted by the Mars Regional Atmospheric Modeling System (MRAMS), an
56 atmospheric mesoscale model, while the dominant flows from the north are underestimated.

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58 cause of the diverse set of ripples overlapping one on other.

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60 showing how dune topography can influence the migration patterns of ripples and how underlying
61 regional topography seems to control the rates of dune migration.

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64

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71 **1. Introduction and study area**

72 The Martian surface has abundant active aeolian bedforms (Fenton, 2006; Bourke et al.,
73 2008) which have been recently observed to migrate in the current climatic setting (Silvestro et al.,
74 2010, 2011, 2013; Chojnacki et al., 2011; Hansen et al., 2011; Bridges et al., 2011, 2012, 2013;
75 Geissler et al., 2013; Sparavigna, 2013). New techniques that take advantage of the high resolution
76 of the HiRISE (High Resolution Imaging Science Experiment) data (McEwen et al., 2007) have
77 been recently applied to characterize small-scale aeolian bedforms on Mars. The migration rates of
78 ripples were computed using the Coregistration of Optically Sensed Images and Correlation (COSI-
79 Corr) software (Bridges et al., 2012), while ripple trends were automatically derived using the
80 Object-Based Ripple Analysis (OBRA) technique (Silvestro et al., 2011, 2013; Vaz and Silvestro,
81 2012, 2014).

82 The aim of this study is to use these two methods in combination to analyze dune and ripple
83 patterns and migration using a pair of overlapping HiRISE images in Herschel Crater, a 300 km
84 Noachian impact basin in the Mare Tyrrhenium region (MC22) (Figure 1). The dunes of Herschel are
85 of particular interest as they have been previously interpreted as ancient indurated aeolianites (due
86 to the grooved pattern visible on the dune slopes; Malin and Edgett, 2000). More recent images
87 from the HiRISE camera showed that such a pattern is formed by sand ripples which, together with
88 the dunes, are consistently migrating (Bridges et al., 2007, 2011, 2013;).

89 However, first evidences for sand motion in the Herschel Crater have been detected by (Cardinale et
90 al., 2012a).

91 In this work we compute ripple and dune migration rates and compare the migration
92 directions with the present-day winds simulated by the Mars Regional Atmospheric Modeling
93 System (MRAMS), a mesoscale atmospheric model (Michaels and Rafkin, 2008; Rafkin, 2001).

94 We use this wind model at two diverse grid scale to interpret the observed aeolian morphologies.

95 In this way, we test the capability of the wind model to predict the wind regime necessary for the
96 creation and the evolution of the Herschel dune fields; we show that the aeolian activity that is

97 shaping the dunes is not strictly unidirectional and that the topography of the crater is controlling
98 the wind flow at the dune scale.

99

100 **2. Methods**

101 We conducted a detailed geomorphological analysis of dunes and ripples in Herschel Crater
102 using a time series of HiRISE images and a stereo pair that was used to build a DTM with SOCET
103 SET (Mattson et al., 2011) (Figure 1, Table 1). Images S1 and T3 (Table 1) were orthorectified over
104 the DTM using the Co-registration of Optically Sensed Images and Correlation (COSI-Corr)
105 software package (Leprince et al., 2007; Bridges et al., 2011). Dune morphometric parameters
106 (slopes and aspect angles) were computed in ArcGIS and used to derive density stereoplots of the
107 slipface surface vectors, providing an approximation to the main sediment flux direction (Figure 2d)
108 (Silvestro et al., 2013). Ripple crestlines were mapped over the study dunes in the S1 image using
109 the OBRA (Object-Based Ripple Analysis) procedure introduced by Vaz and Silvestro (2014). This
110 technique is used to derive the main trends of the ripples, providing information about
111 wind/sediment interactions at smaller scales (Figure 3).

112 The lee fronts of the dune slip faces were manually digitized on the S1 and T3 images in
113 order to derive the dune migration rate and direction (Figure 4). In addition, we also evaluated the
114 spatial distribution of the dune migration azimuth (Figure 5). We then used COSI-Corr to track the
115 ripple displacement over the S1 and T3 images acquired 1359 Earth-days apart (Table 1). The result
116 is a ripple displacement map (Figure 6) from which we derived the average ripple migration rate
117 (Figures 7 and 8).

118 Finally, we estimated the potential timing of the sand-moving events by using the MRAMS
119 mesoscale atmospheric model (Michaels and Rafkin, 2008; Rafkin et al., 2001) on two diverse grid
120 scales. In the first one, the surface stress and wind direction have been modeled for one typical day
121 for each of the four seasons (at $L_s = 210^\circ$ (southern summer), $L_s = 300^\circ$ (southern spring), $L_s = 30^\circ$
122 (southern winter) and $L_s = 120^\circ$ (southern autumn) at a spatial resolution of $\sim 8 \text{ km} \times 8 \text{ km}$ (Figure

123 9). The output model state was recorded every twenty Martian-minutes for four typical sols.
124 In the second atmospheric simulation, we used a model with ~2km grid spacing to constrain the
125 local wind conditions in the Herschel Crater (Fig.10). The modeled winds have been sorted into 24
126 equal width direction bins over the dune field and each point may have a maximum 24 vectors and
127 these ones show the downwind direction of the wind.
128 The instantaneous model was recorded for 12 seasons at $L_s=0^\circ$ (southern winter), $L_s=30^\circ$ (southern
129 winter), $L_s=60^\circ$ (winter), $L_s=90^\circ$ (northern winter), $L_s=120^\circ$ (southern autumn), $L_s=150^\circ$ (autumn),
130 $L_s=180^\circ$ (northern autumn), $L_s=210^\circ$ (southern summer), $L_s=240^\circ$ (summer), $L_s=270^\circ$ (northern
131 summer), $L_s=300^\circ$ (southern spring) and $L_s=330^\circ$ (spring) by the MRAMS (Fig.11).

132

133 **3. Dune and ripple morphology**

134 The study dunes are located in a ~1200 km² dune field in the western floor of Herschel
135 crater and consist of barchans and barchanoids (Cardinale et al., 2012b). These dunes can be more
136 than 60 m tall and spaced ~200-800 m apart. Some of the dunes present an asymmetric structure,
137 with the slip face being elongated obliquely (Figures 2b and c). Visual assessment and stereonet
138 analysis reveal a high dispersion in the dune slipface orientation and slope values clustering at ~30°
139 (Figure 2d). This reflects the concave shape of the barchan and barchanoid slipfaces. Most of the
140 slip face vectors are oriented toward the south, trending between ~60° and ~300° with a main mode
141 located at ~125° (Figure 2d), reflecting the dune slipface asymmetry.

142 In Figure 3 we show some examples of the different types of ripple patterns in the area. On
143 the dune flanks ripples are spaced 2-4 meters apart and are two-dimensional with typical “Y”
144 junction terminations (Figure 3a), while on the top of the dunes the ripple pattern is more three-
145 dimensional, with diverse ripple sets overlapping (Figure 3b). Such a ripple arrangement probably
146 reflects the coupling between ripple straightness and slope described by Rubin (2012) and observed
147 in the field by Howard (1977). The complexity of the ripple pattern is shown in the rose diagram in
148 Figure 3c (showing the distribution of the crestline trends mapped automatically over the dunes in

149 the yellow box). The overall length-weighted circular distribution of the mapped ripple traces shows
150 different trends and two main modes at $\sim 45^\circ$ and $\sim 135^\circ$.

151 **4. Bedform migration**

152 **4.1. Dunes**

153 In Figure 4a we show the areal distribution of the average dune migration vectors for 211
154 dunes computed by comparing the pair S1 and T3 (Table 1) ($\Delta t = 1359$ Earth-days). Dunes which do
155 not have a clear slipface are excluded from the analysis. The distribution of the lee motion is neither
156 uniform nor unidirectional. A higher migration value is reported for the dunes located in the
157 northern dune field sector (1.2-2.2 (m)) with the dune displacements decreasing toward the south
158 (Figure 4a). Such a N-S migration trend can be attributed to the abrupt change in the roughness at
159 the dune field margin which triggers the development of an internal boundary layer that thickens
160 downwind (Jerolmack et al., 2012). Figure 4b highlights the north-south migration trend (left) and
161 shows that the area with larger migration values (between -14.70° and -14.75° in latitude)
162 corresponds to a drop in elevation that is well represented in the northern part of the Herschel dune
163 field (right). This suggests that the underlying large-scale topography also contributes to the shape
164 of the internal boundary layer.

165 On average, the dunes migrated 0.8 meters toward the SSE (Figure 5a shows the computed
166 average vector) giving a rate of migration of 0.45 meter/Mars-year (MY) (~ 0.2 meter/Earth-year or
167 m/EY), assuming that this value is constant from year to year. The measurements show high
168 directional dispersion ($\mu = 162^\circ \pm 38^\circ$) (Figure 5b), which might be partially due to the local
169 topography since the dunes are not migrating over a flat surface.

170 **4.2. Ripples**

171 In Figure 6 we show the ripple displacement map obtained with COSI-Corr. The map
172 reveals that significant movement occurred across the investigated area between March 2007 and
173 December 2010. In the northern area of the Herschel dune field the fastest ripples moved so far that
174 the correlation breaks down, that is, once the migration exceeds a distance at least equal to the

175 ripple wavelength (5.1 m) (Figure 6b). In the central and southern dune field sectors the ripple
176 displacement is smaller, so it can be traced (Figures 6a and c). The ripple migration rate also varies
177 with the height of the dunes, with the fastest migrating ripples located close to the dune crest of the
178 dunes (Figure 6c). This is the result of the linear relationship between height and ripple migration
179 also reported for the Nili Patera dunes (Bridges et al., 2012). During a period of 1359 Earth-days we
180 obtained an average vector for ripple migration of 1.1 meters and trending toward SSE (Figure 7a).
181 This gives a migration rate of 0.55 meters in one MY (~ 0.3 meters/EY). The measurements show a
182 high circular standard deviation (41.6°) and their directional trend is mainly bimodal with modes at
183 $\sim 175^\circ$ and $\sim 240^\circ$ (Figure 7b). The high directional dispersion of the ripple migration vectors is due
184 to the local dune topography which deflects the wind over the dunes as shown in Figure 8. In
185 particular, the secondary mode at $\sim 240^\circ$ is due to the ripple migration vectors in the lee of the dunes
186 (orange vectors in Figure 8).

187 **5. Modeled winds**

188 The atmospheric models (MRAMS simulations at 2 diverse grid scales) are used to evaluate
189 wind strength and direction, to explain the observed aeolian morphologies and dune changes (
190 Fenton et al., 2005; Hayward et al., 2007; Chojnacki et al., 2011).

191 Due to the complex pattern of the Herschel dune field, we suppose a multi directional wind regime
192 to be estimated by these simulations.

193 In Figures 9a and 9b (wind model at 8 km grid space) we show the daily maximum MRAMS ratio
194 (stress / threshold stress) vectors over the whole dune field. We define the stress ratio as the
195 aerodynamic surface stress divided by the minimum threshold aerodynamic stress calculated using
196 the expressions of Greeley and Iversen (1985). Dominant modeled wind direction is from the west
197 to the east (Figure 9c) with the strongest winds blowing close to the western crater rim (Figure 9a).
198 The predicted stress values are just above the (Kok 2010) threshold for sand saltation maintenance
199 (10% of the Greeley and Iversen (1985) saltation initiation threshold). In Figure 9b we show the
200 modeled wind directional variability. A general trend is visible with the winds being more uni-

201 directional close to the western crater rim (see the lower circular STD in this area and Table 2). In
202 Figure 9d we show the same data plotted by seasons with the important statistic parameters
203 summarized in Table 2.

204 The strongest winds blow at $L_s=30^\circ$ (southern winter) from the west to the east with a circular STD
205 of 34° . In the other seasons modeled winds are weaker and multi-directional ($CSTD>87^\circ$).
206 Dominant winds from the north to the south, matching the dune and ripple migration direction, are
207 not predicted by the model.

208 Output from MRAMS at 2km grid scale such as that of Fig. 10 is used to resolve topographically –
209 influenced wind flows not explained in the previous wind simulation.

210 Modeled wind strength and direction from twelve Martian sols are examined here (Fig.11).

211 Within the investigated dune field a predominant wind distribution is not visible; a wind direction
212 variability possibly induced by the crater topography is high in all the twelve studied seasons and
213 the strongest modeled winds are predicted to blow from N-NE in spring ($L_s=330^\circ$).

214 The modeled winds by this second MRAMS simulation highlights that prevailing winds from north,
215 matching the measured slip faces, are not predicted by this model.

216 According to this second simulation, we noticed that weaker winds are frequent in the Herschel
217 Crater such as other areas previously studied (Silvestro et. al.2012, 2013).

218 Even in this second model, the simulated shear stress values are just above the Kok threshold for
219 sand saltation maintenance with a maximum stress value of 0.30 for all the investigated seasons
220 (Fig.12).

221 According to the recent numerical models on the Martian sand saltation, the hypothesis on the
222 hysteresis phenomenon could show how after the initiation, the saltation can be sustained with
223 weaker winds on martian surface (Kok, 2010).

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227 **6. Discussion and conclusion**

228 Our results show that the ripples and the dunes in the Herschel Crater are mainly shaped by
229 winds blowing from the north to the south. However, the asymmetry in the dune form indicates that
230 the wind regime is not strictly uni-directional (Bourke M.C., 2009; Parteli et al., 2014).

231 In particular, following the model of Bagnold, (Bagnold, 1956), the influence of a local bimodal
232 wind regime (winds blowing from NNW and from NNE), should be the cause for the observed
233 asymmetry with the former probably being more frequent or stronger (Figure 2b-c).

234 The influence of more than one wind, the combination of dune collision, limb extension and
235 merging with downwind dunes (Bourke M.C., 2009) are also supported by the dunefield pattern
236 which is highly irregular and intricate (Bridges et al., 2007) and by the resultant bedform migration
237 directions.

238 The first MRAMS simulation, show modeled winds enhanced by the western crater rim
239 blowing to the east. The interaction of these flows with the dominant winds coming from the north
240 may be the cause of the observed dune morphology and migration direction. In the second MRAMS
241 simulation, the diverse combined wind flows may partially explain the intricate ripple pattern. The
242 ripple migration however, seems to be controlled by the local dune topography and any
243 extrapolation from local to dune field/regional scale has to be treated carefully.

244 The lack of dominant winds from the north in the MRAMS simulation can be attributed to the low
245 spatial and temporal coverage. A similar situation has also been reported by other workers
246 (Hayward et al., 2009; Silvestro et al., 2012) suggesting the importance of ground truth data when
247 deriving the wind regime of a certain area on Mars.

248 With the exception of the Nili Patera ripples, bedform migration rates in Herschel are
249 comparable to other areas on Mars (Bridges et al., 2012; Silvestro et al., 2011, 2013). However,
250 without continuous and long-term monitoring of Herschel and the other zones, this kind of
251 inference remains speculative.

252 In flat areas, changes in surface roughness can increase the boundary shear stress in the

253 upwind margins of dune fields (Jerolmack et al., 2012), producing spatial variations in the sediment
254 fluxes. In more complex terrains, like in the floor of Herschel crater, the observed relationship
255 between dune celerity and local topography (in particular the abrupt change in the roughness of
256 ~1km illustrated in the Fig.4) suggests that in addition to the roughness of the aeolian pattern, the
257 variations in the long wavelength topography (Pelletier et al., 2014) control the Herschel dunefield
258 properties.

259 Collectively, the combination of different methods of investigation helped to better decipher
260 the wind regime in Herschel Crater. At the dune field scale, the main winds from the north combine
261 with wind from the west enhanced by the western crater rim (the dune field is distant 28 km from
262 the westerly crater rim). At the dune scale, the topography of the dunes and the substrate
263 topography are controlling dune height, bedform migration rates and directions as described for the
264 White Sand dune field in New Mexico by Pelletier et al., 2014. The topography at regional and
265 local scale is indeed an important boundary condition that needs to be carefully addressed in order
266 to extract the best wind information from remote sensing images.

267

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371

372

373 **Supplemental table**

- 374 1. S1-S2-T3 HiRISE acquisition parameters.
- 375 2. Statistical parameters of the modeled winds divided by seasons.

376

377 **Captions**

- 378 1. a) Details of the study area, showing the distribution of the large dark dune fields within

379 Herschel Crater (MOLA shaded topography with Themis daytime mosaic). **b)** A perspective
380 view of the large dark dunes from HiRISE images PSP_002860_1650 and
381 ESP_020384_1650. A shaded relief map from MOLA data (top), showing the location of
382 Herschel Crater.

383 **2.** The Herschel Crater dune field (CTX image P05_002860_1650_XI_15S232W) slope map
384 derived from High Resolution Imaging Science Experiment digital terrain models (DTMs).
385 **b-c)** Barchan dunes with and without elongated horns in the northern and north-eastern area
386 of the dune field. The slope map suggests the presence of a slipface with a trend of 62° -
387 286° , denoting a predominant wind direction blowing from the northeast (HiRISE image
388 PSP_002860_1650). **d)** The lower hemisphere equal-area density stereoplot for all the
389 slipface surface vectors estimated from the HiRISE DTMs. The estimated dip angle is $\sim 30^\circ$

390 **3. a)** This inset represents the ripple length-weighted circular distribution (the mapped area
391 corresponds to the yellow window shown in Fig. 2). The mapped ripple population exhibits
392 a bimodal trend with two main modes trending 45° and 135° . **b)** An inset of a horn of a
393 dune, where along its flanks, the ripple crests are continuous and have a two dimensional
394 pattern. **c)** The pattern of the ripples superposing the dune's slopes is complex due to the
395 diverse wind flows blowing over the dunes.

396 **4. a)** Average dune lee front migration vectors overlain on the DTM. The colored vectors
397 represent the average displacement of the slipface lee fronts over three Earth-years (from
398 March 2007 to December 2010). **b-c)** The plots represent the statistics computed using 500
399 m moving windows for the migration and elevation at the base of the slipfaces. Note the
400 general decrease of the displacements when moving south, and the association of the area
401 with larger migrations (between -14.7° and -14.75°) with a drop in elevation.

402 **5. a)** This rose diagram shows the distribution of lee side migration azimuth. **b)** This rose
403 diagram shows the circular distribution of the lee side migration.

404 **6.** Ripple displacement map for the Herschel Crater dune field, derived from correlated

405 HiRISE images with high displacements shown with warmer colors. **b)** Fast ripples moved
406 so much that the correlation breaks down causing the observed fuzzy pattern. **c)** Area in
407 which the correlation starts to record the ripple migration.

408 **7.** Circular distribution of the migration vectors. **a)** Ripple migration mean vector. **b)** This rose
409 diagram shows the circular distribution of all ripple migration vectors. The secondary mode
410 at $\sim 240^\circ$ is due to the ripple migration vectors in the lee of the dunes (see Figure 8).

411 **8.** Daily maximum MRAMS stress (ratio) vectors computed from COSI-Corr vectors data.

412 **9.** MRAMS modeled wind stresses and directions. **a)** Daily maximum stress ratio vectors for
413 each of the 36 model nodes covering the dune field. The azimuth of the vectors corresponds
414 to the wind direction while the color is the ratio between the model aerodynamic surface
415 stress and the aerodynamic stress threshold for saltation initiation (Greeley and Iversen,
416 1985) **b)** Circular standard deviation associated with the mean vectors shown in a). **c)**
417 Circular distribution of the wind stress ratio vectors. **d)** Mean stress ratio vectors of the
418 modeled winds divided by seasons.

419 **10.** MRAMS winds in a GIS format over the studied dune field during the 12 investigated
420 seasons. Only winds $>10\%$ of the Greeley and Iversen saltation initiation threshold are
421 included in this plot. Each vector shows the downwind direction of the wind (the direction
422 of the wind is flowing toward). The length of each vector is proportional to the greatest
423 $\text{sfc_stress}/\text{sfc_stress_treshold}$ ratio at the point of the direction while the color shows the
424 relative frequency of each wind direction at each point (warmer colors correspond to more
425 winds blowing in that direction).

426 **11.** Each circular distribution contains the combined information of all 21 MRAMS higher
427 resolution points within the outline of one of the studied HiRISE images (see Fig.10) . the
428 radial direction in these plots is magnitude (of $\text{stress}/\text{stress_threshold}$), with the outer ring
429 being a value of 0.3 and the center of each plot being 0. Each bin is colored by the relative
430 frequency of each wind direction/magnitude (warmer colors correspond to more winds

431 blowing in that direction).

432 **12.** The circular distribution shows the sum of all the 12 investigated seasons.

Table1

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	<i>S1</i>	<i>S2</i>	<i>T3</i>
Product ID	PSP_002860_1650	PSP_003572_1650	ESP_020384_1650
Acquisition Date	07 March 2007	01 May 2007	01 Dicember 2010
Resolution	0.25m/pixel	0.25m/pixel	0.25m/pixel
Latitude (centered)	-14.807 degrees	-14.805 degrees	-14.813 degrees
Longitude (est)	127.888 degrees	127.890 degrees	127.897 degrees
Local Mars time	3.44 PM	3.22 PM	3.41 PM
Solar longitude	195.9 degrees, Northern Autumn	229.7 degrees, Northern Autumn	190.9 degrees, Northern Autumn
Solar incidence angle	56 degrees	49 degrees	55 degrees
Sub solar azimuth	7.3 degrees	351.4 degrees	9.7 degrees
Emission angle	2.4 degrees	27.3 degrees	4.5 degrees

STRESS RATIO

Solar Longitude (°)	Mean Azimuth (°)	Mean Stress Ratio (°)	Circular Standard Deviation
30	88,72726	0,03521729	33,93312
120	67,86047	0,005481146	86,73132
210	145,4258	0,005119130	103,0098
300	232,3831	0,003217317	120,6203

Figure1

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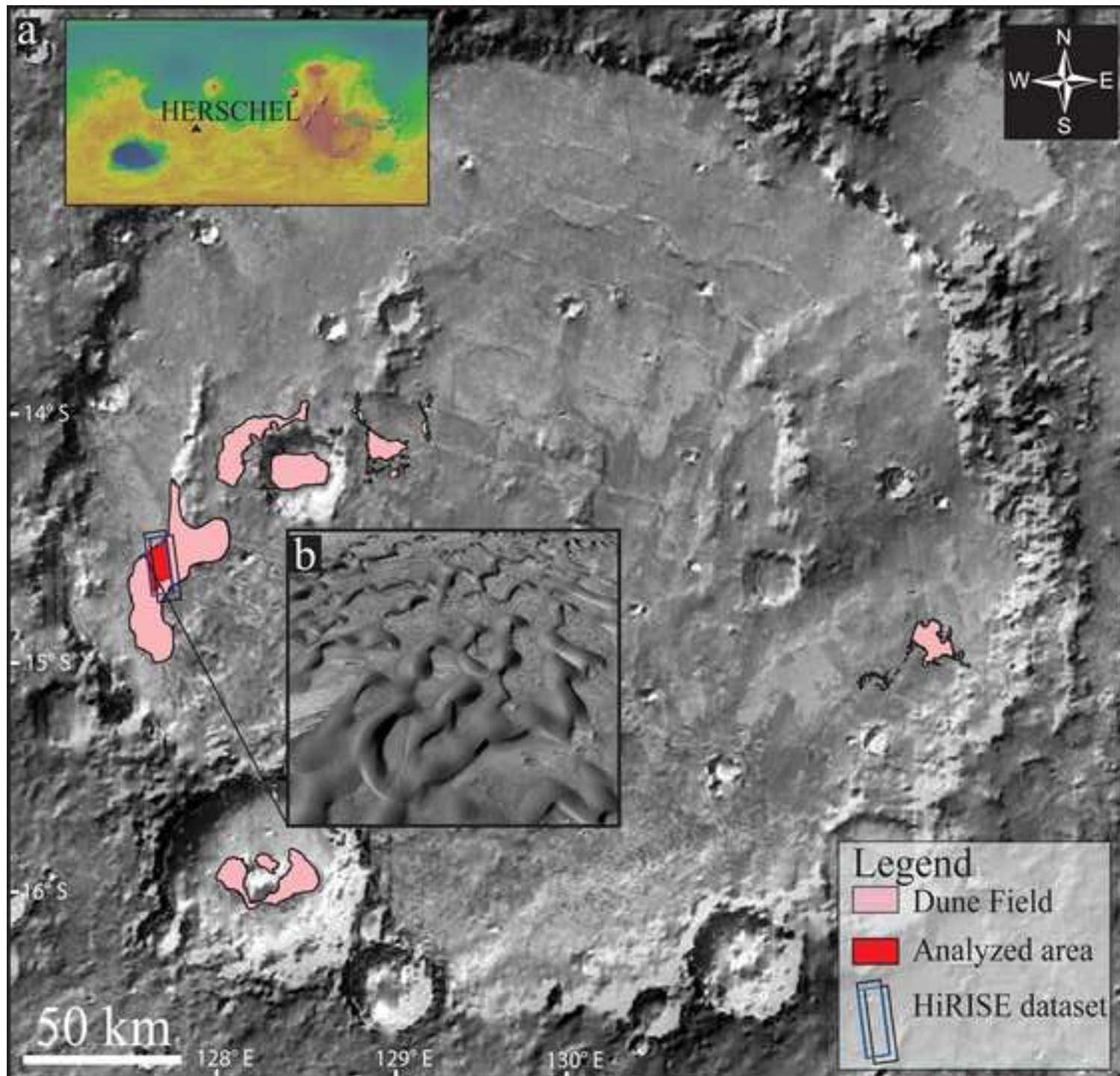


Figure2

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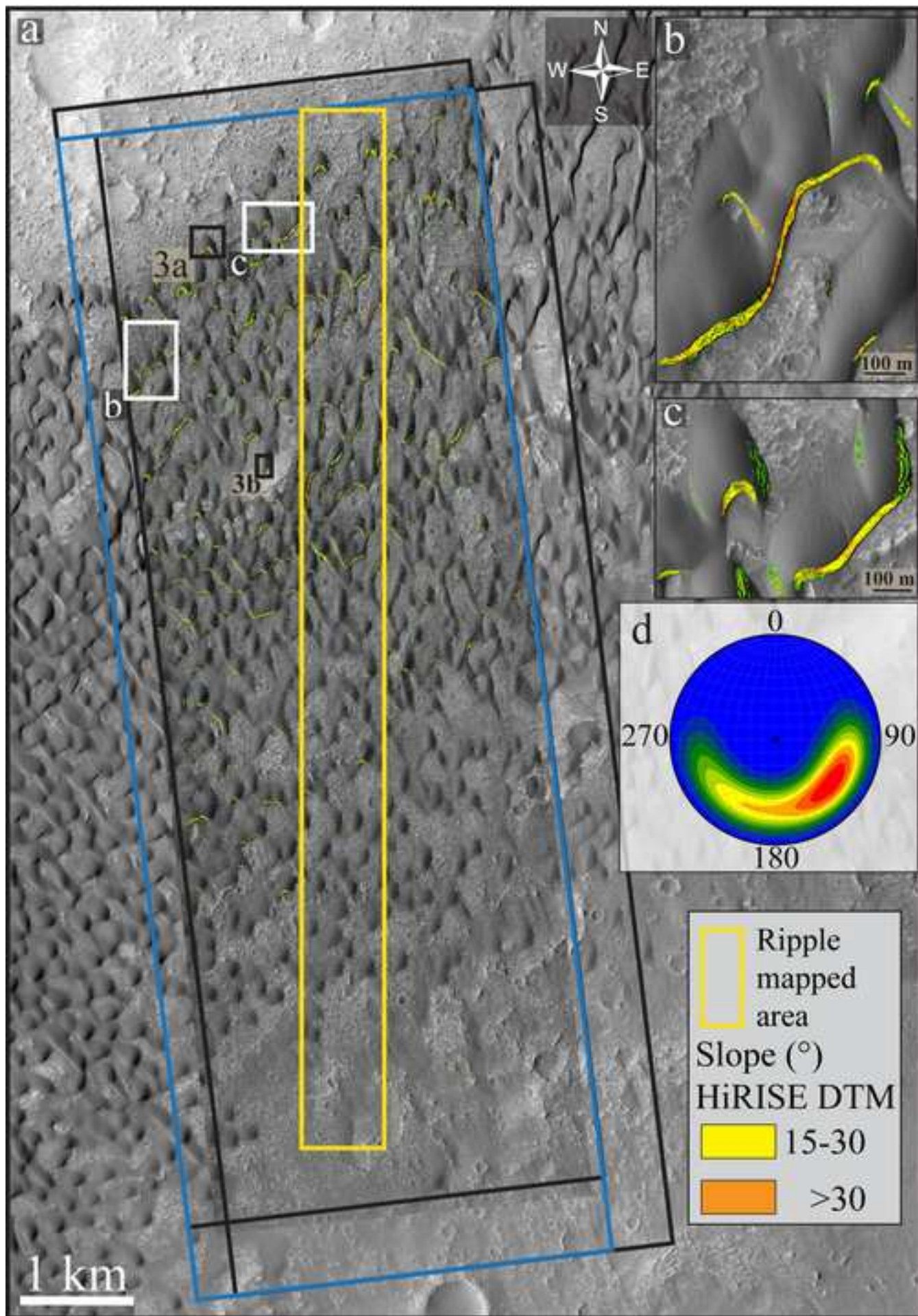


Figure3
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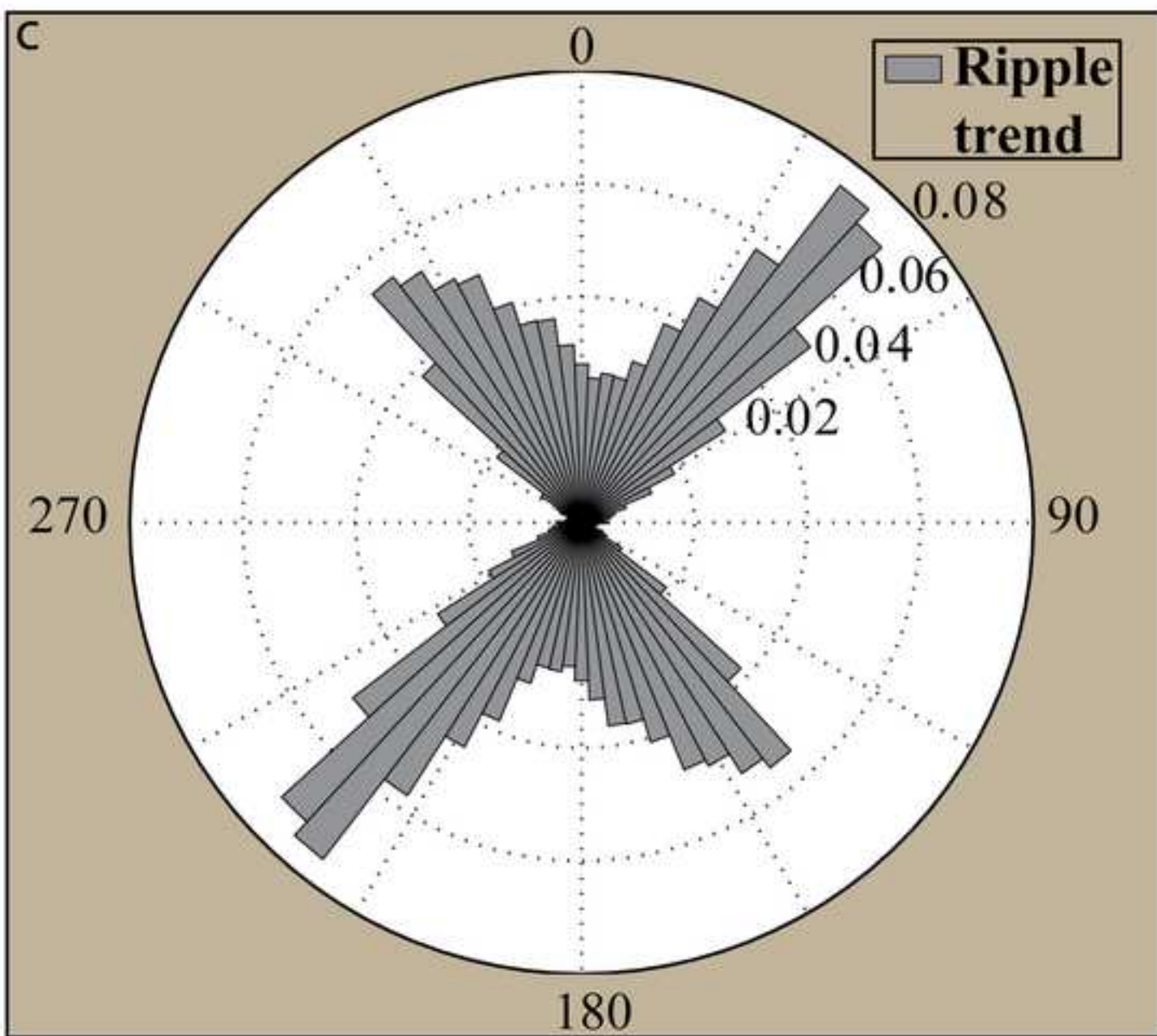
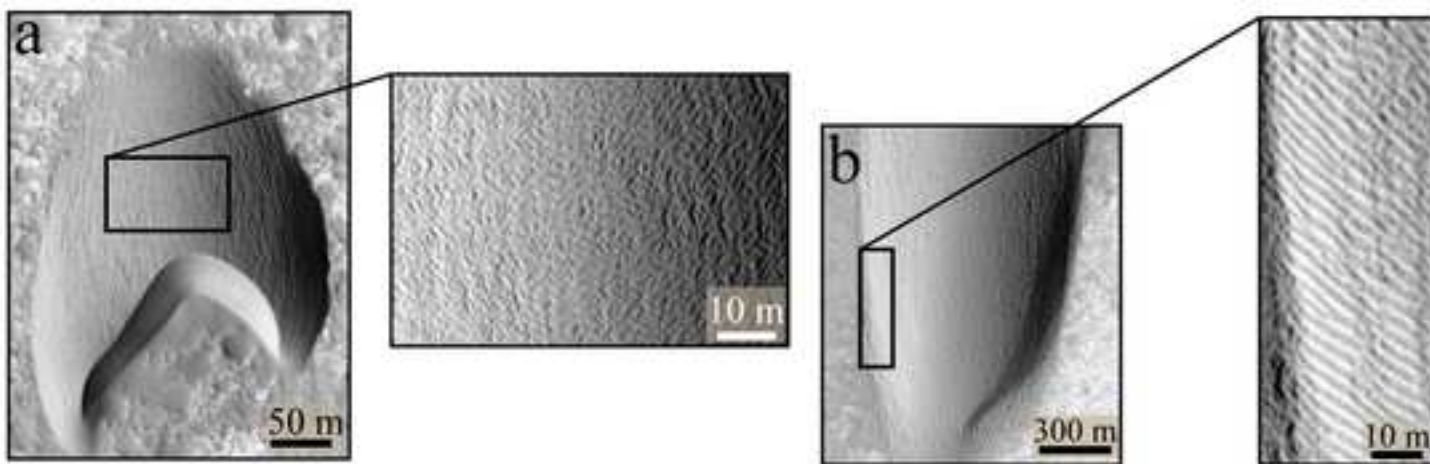


Figure4
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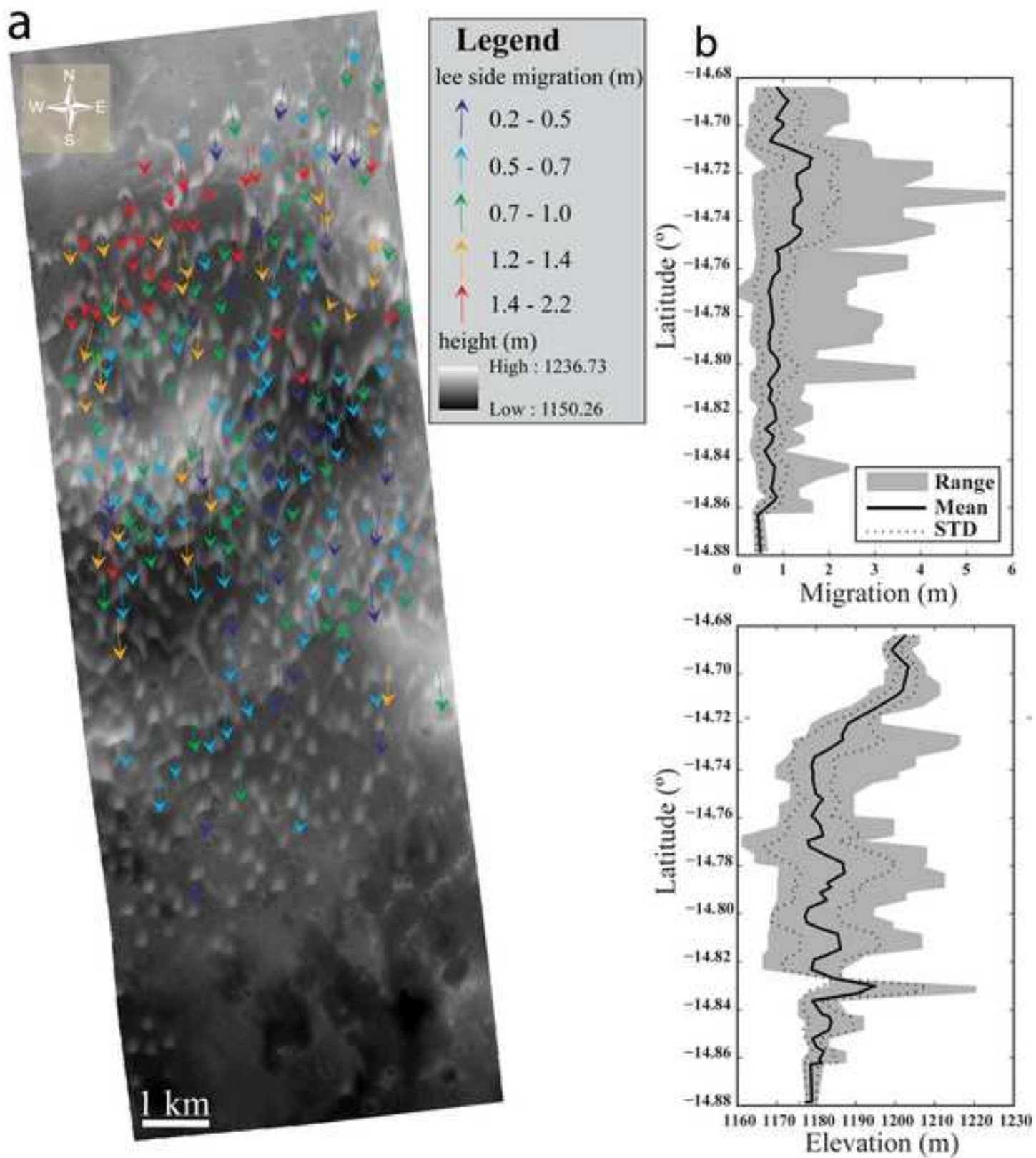


Figure5

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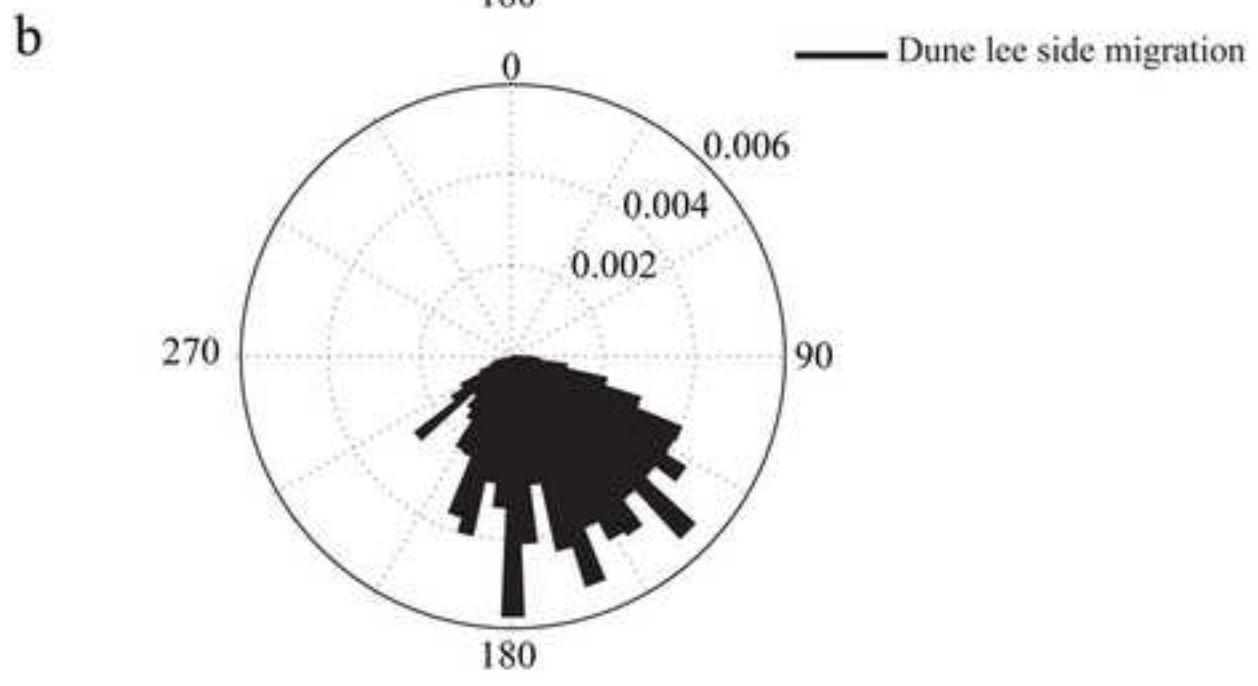
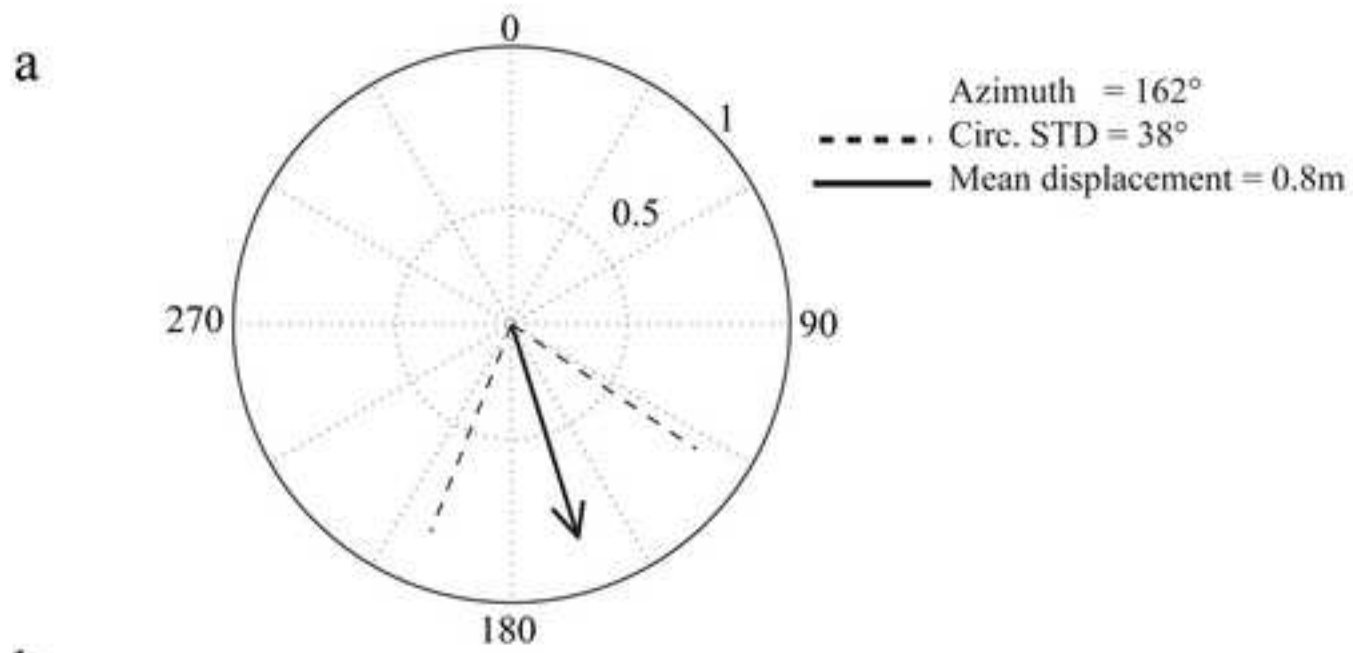


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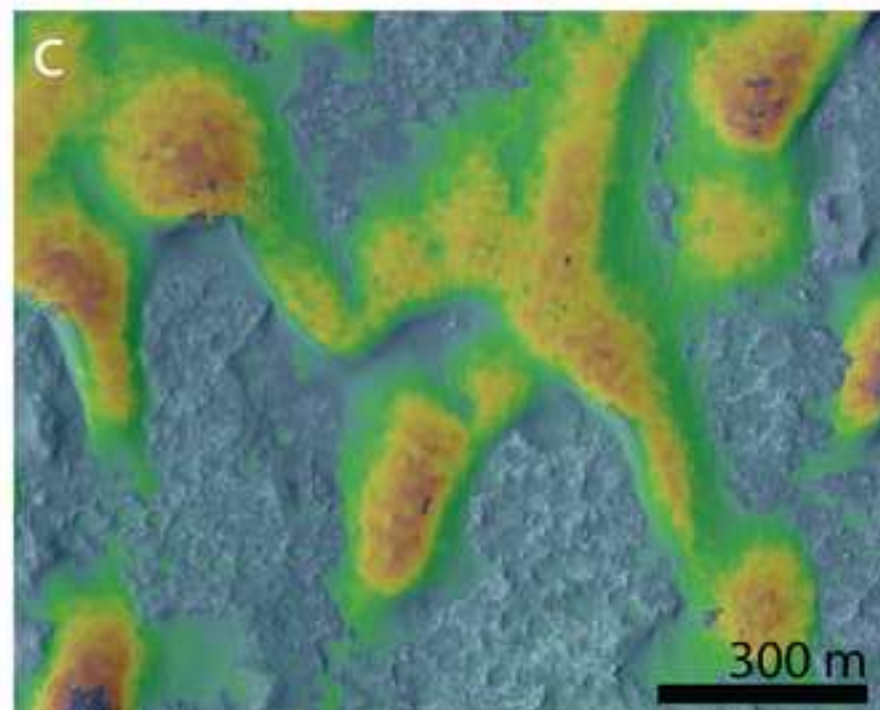
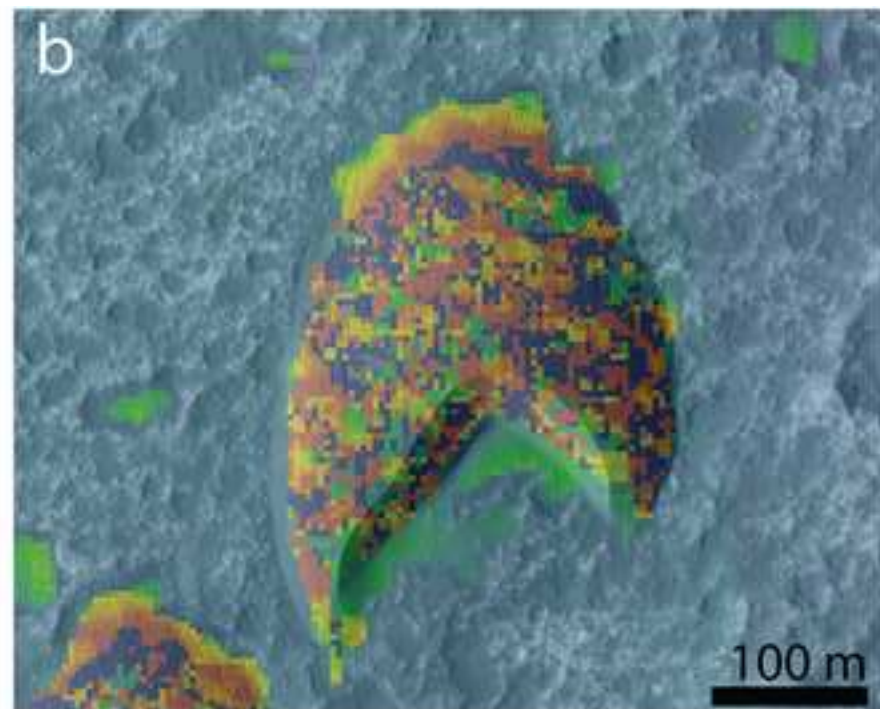
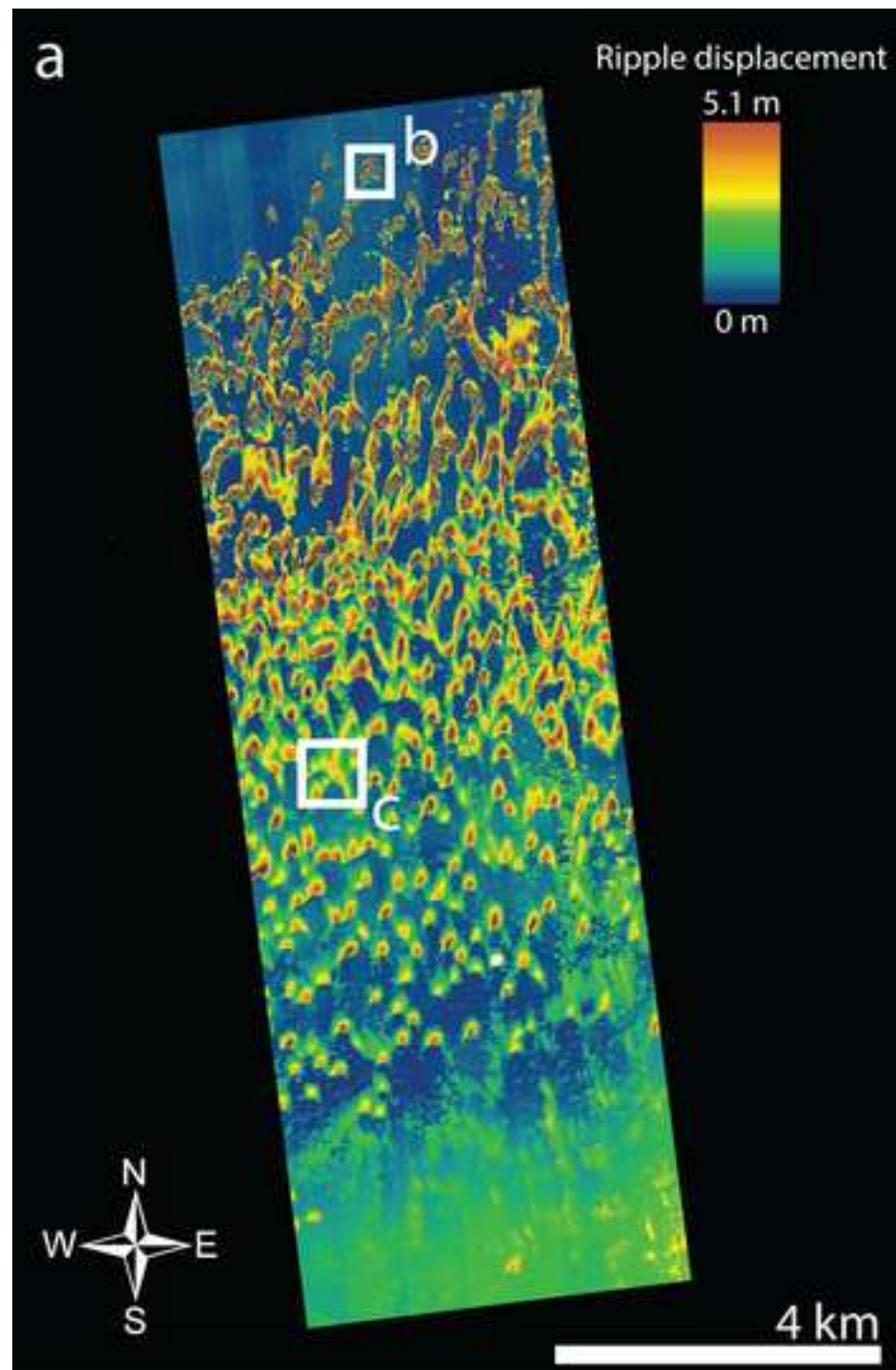


Figure7
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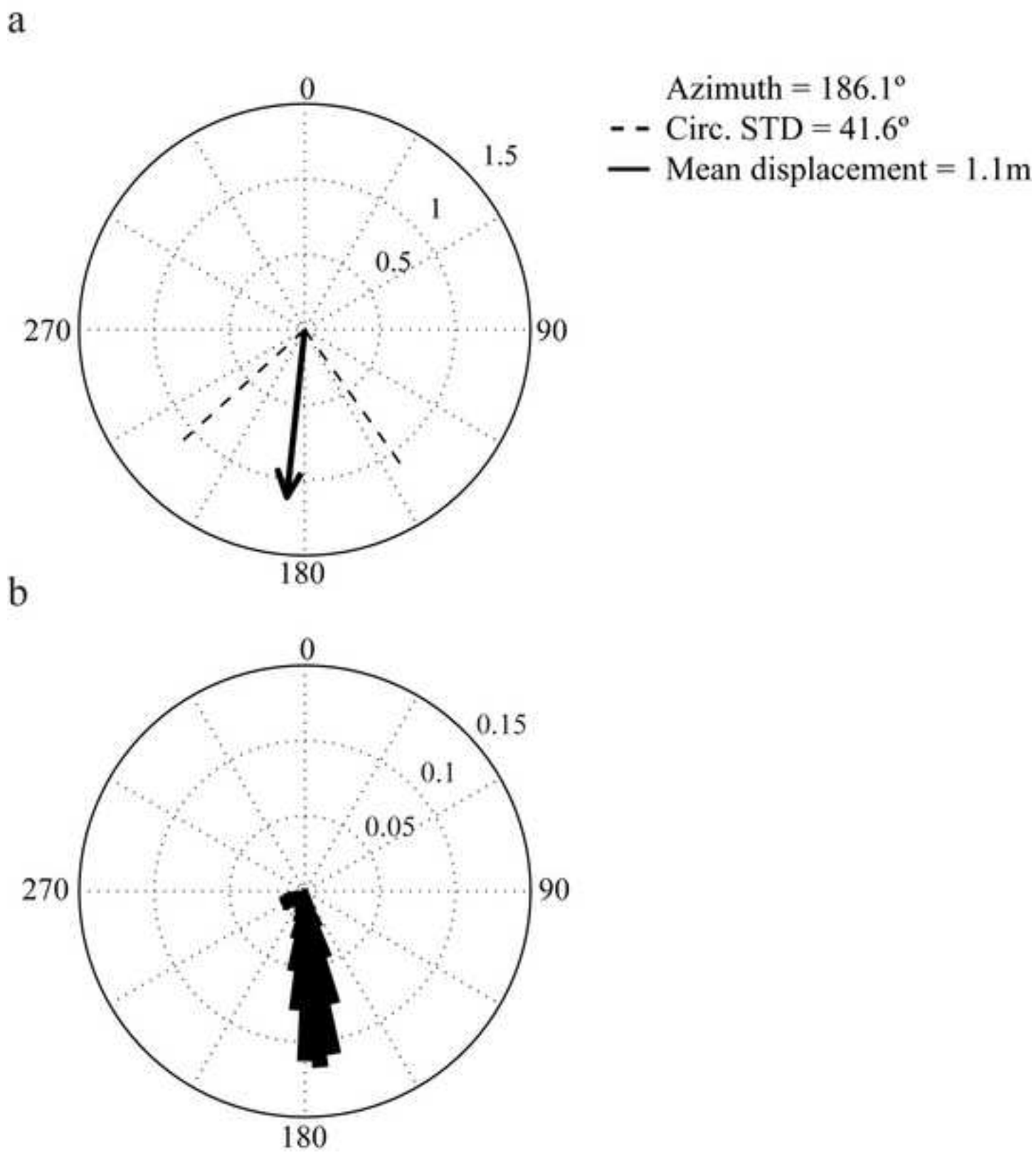


Figure8

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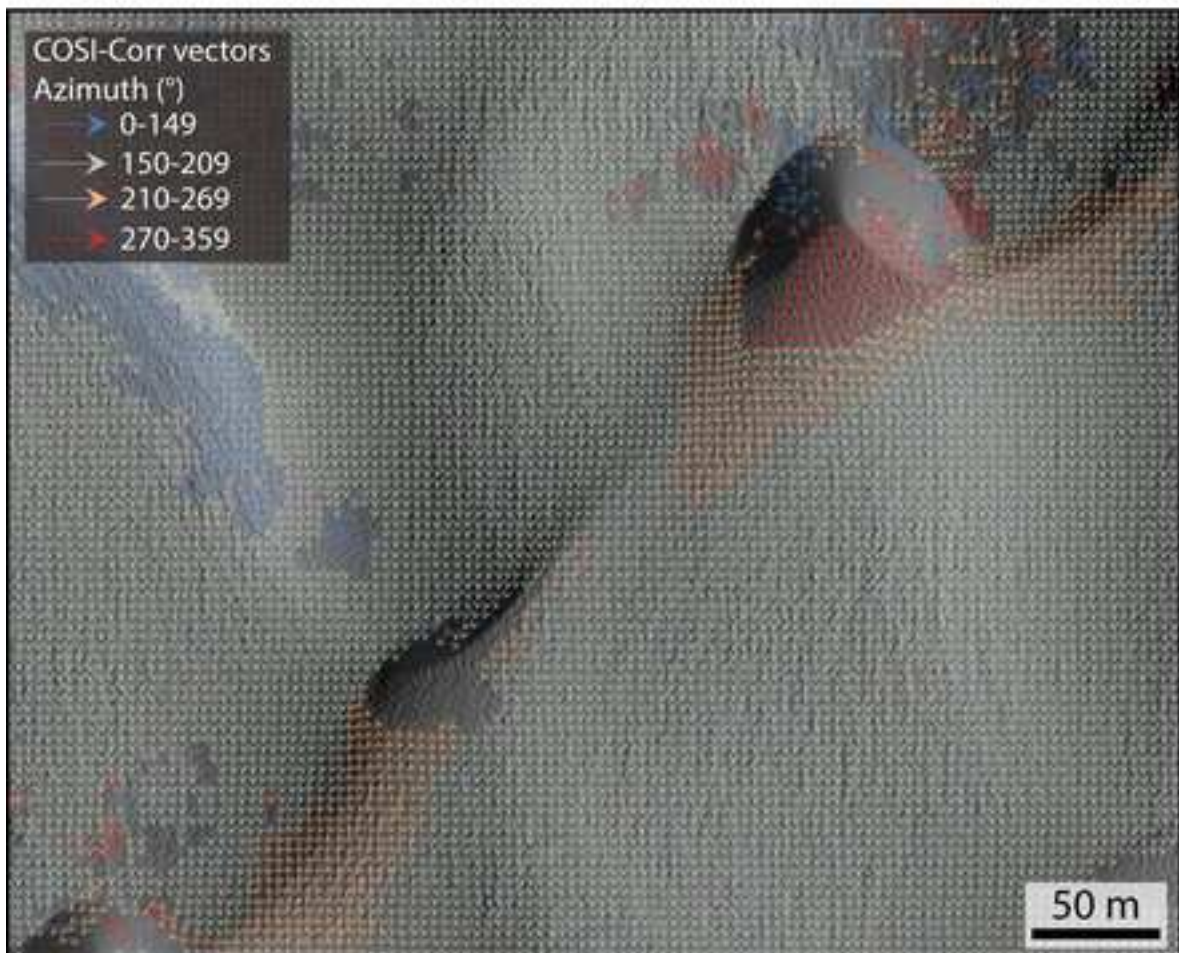
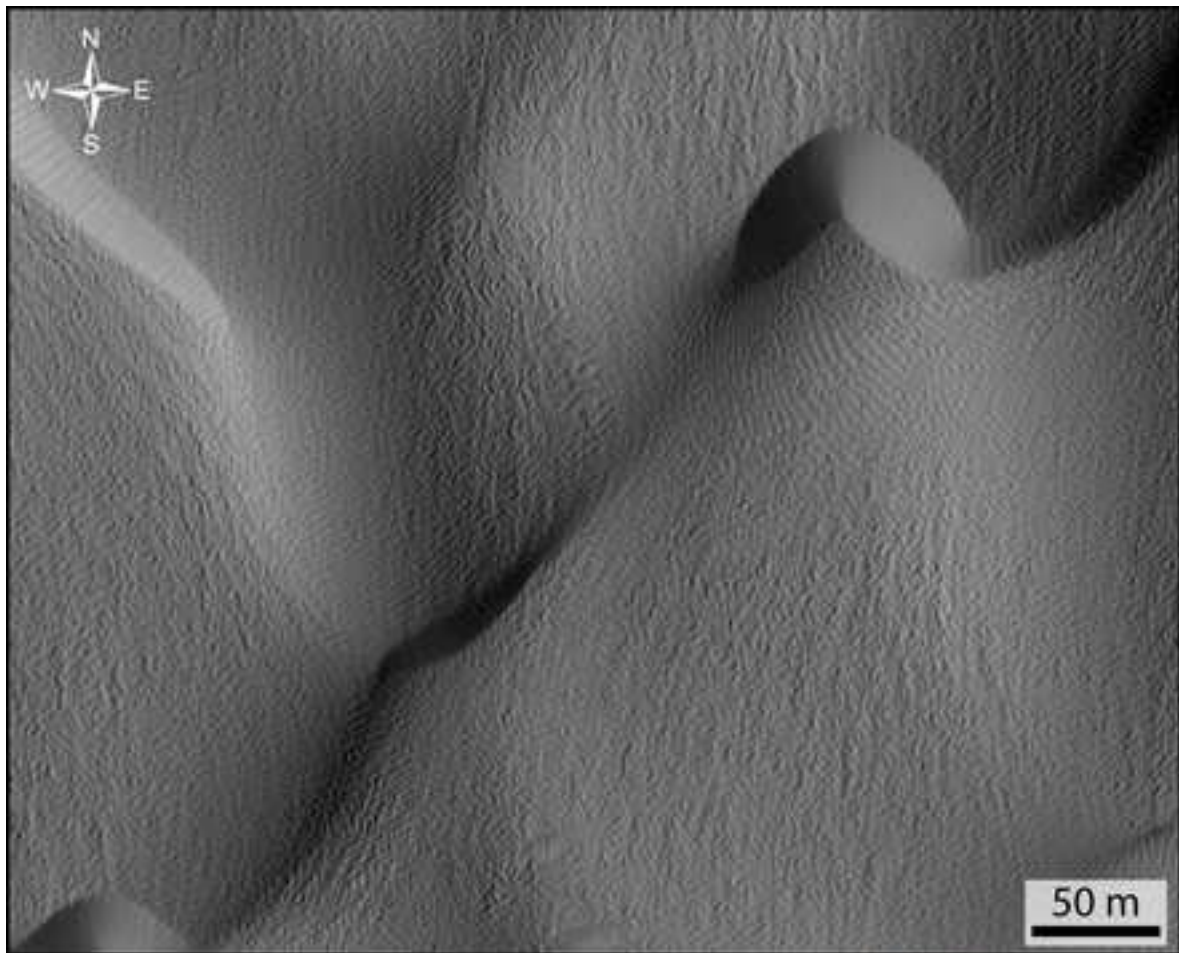
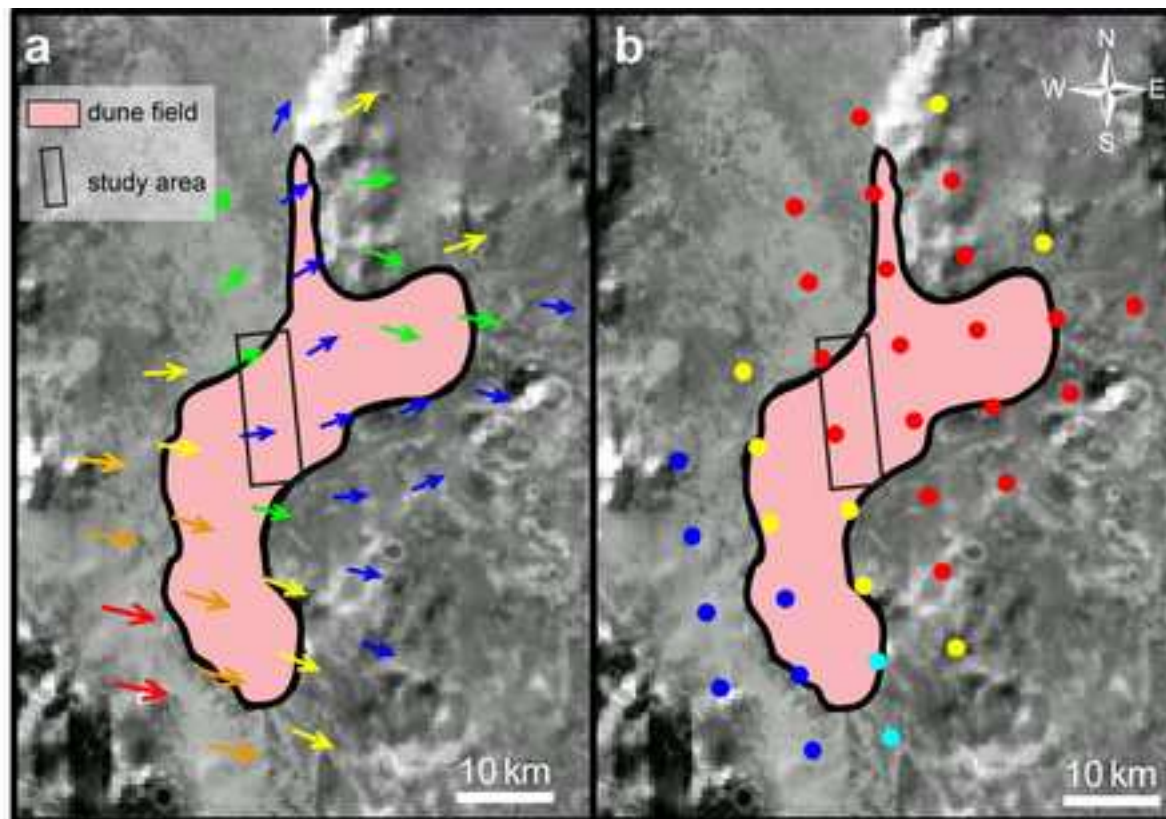
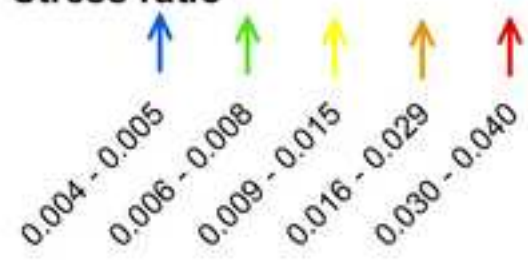


Figure9
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MRAMS

Stress ratio



Circ. st. dev (°)

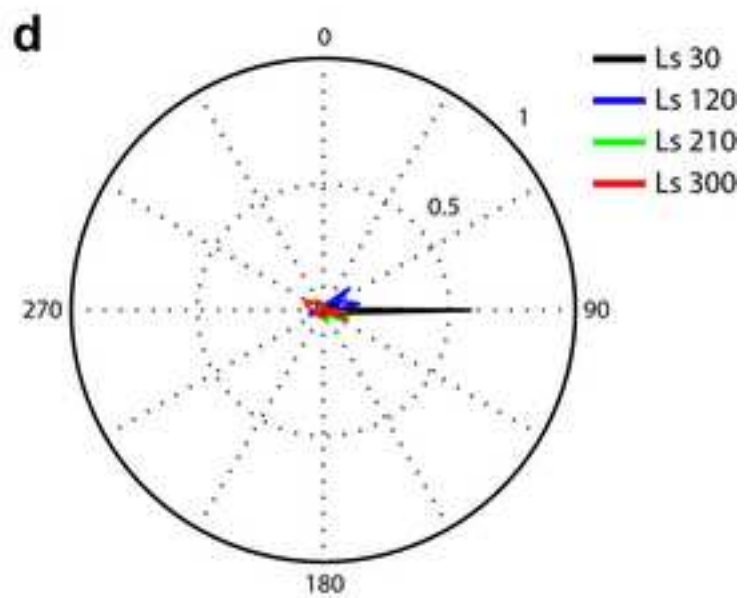
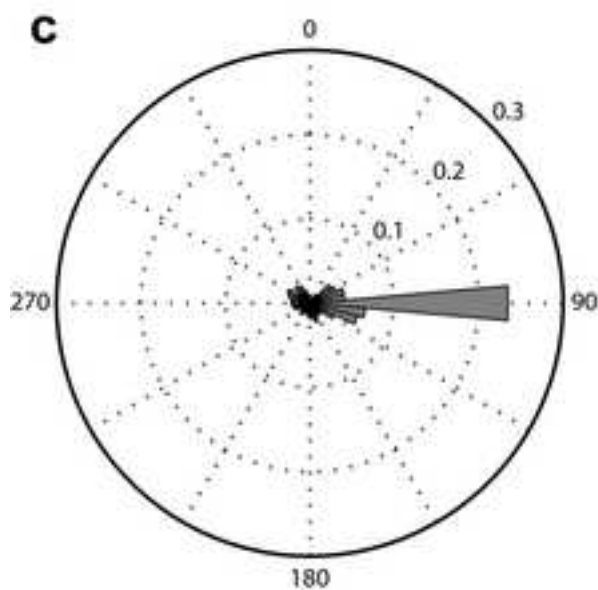
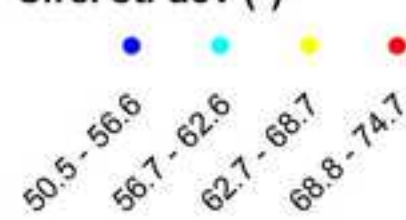


Figure10
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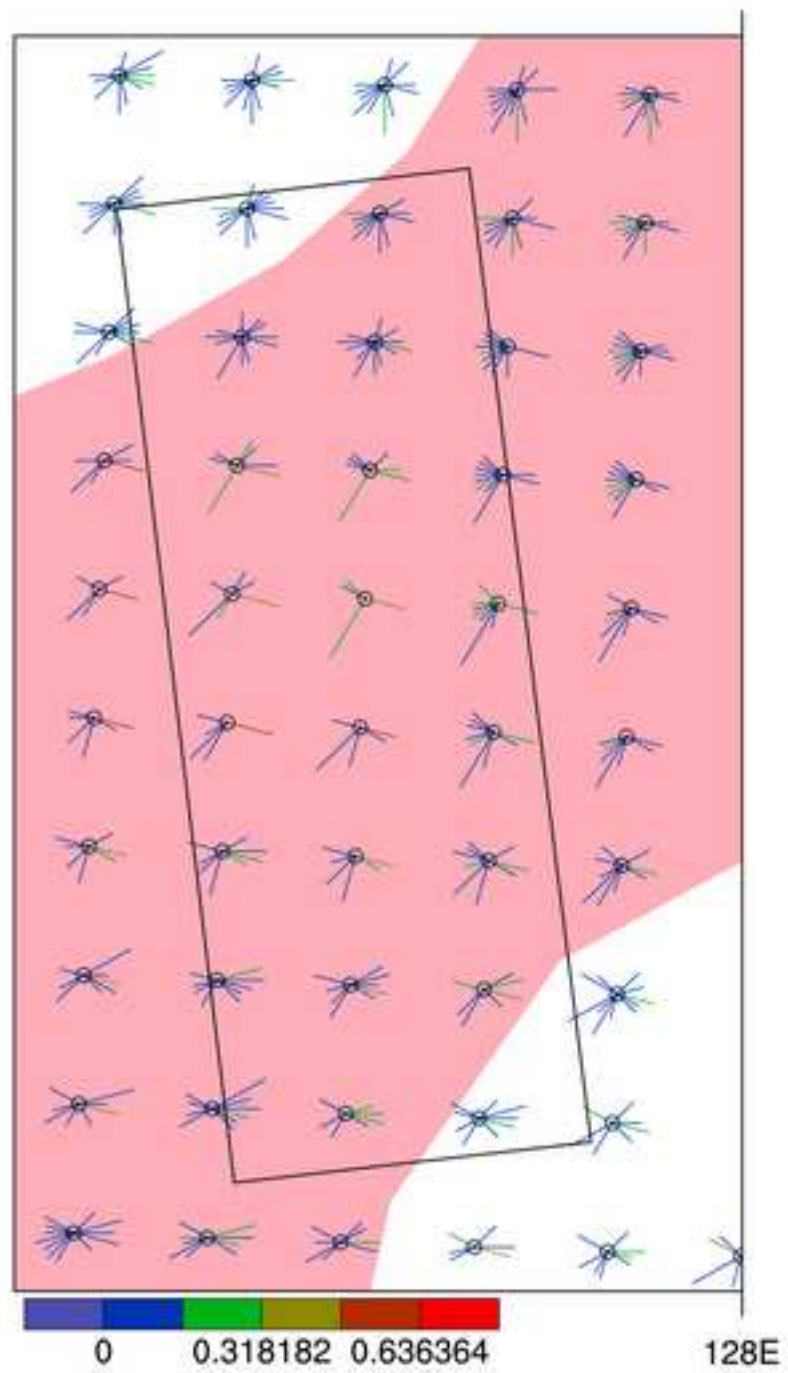


Figure11
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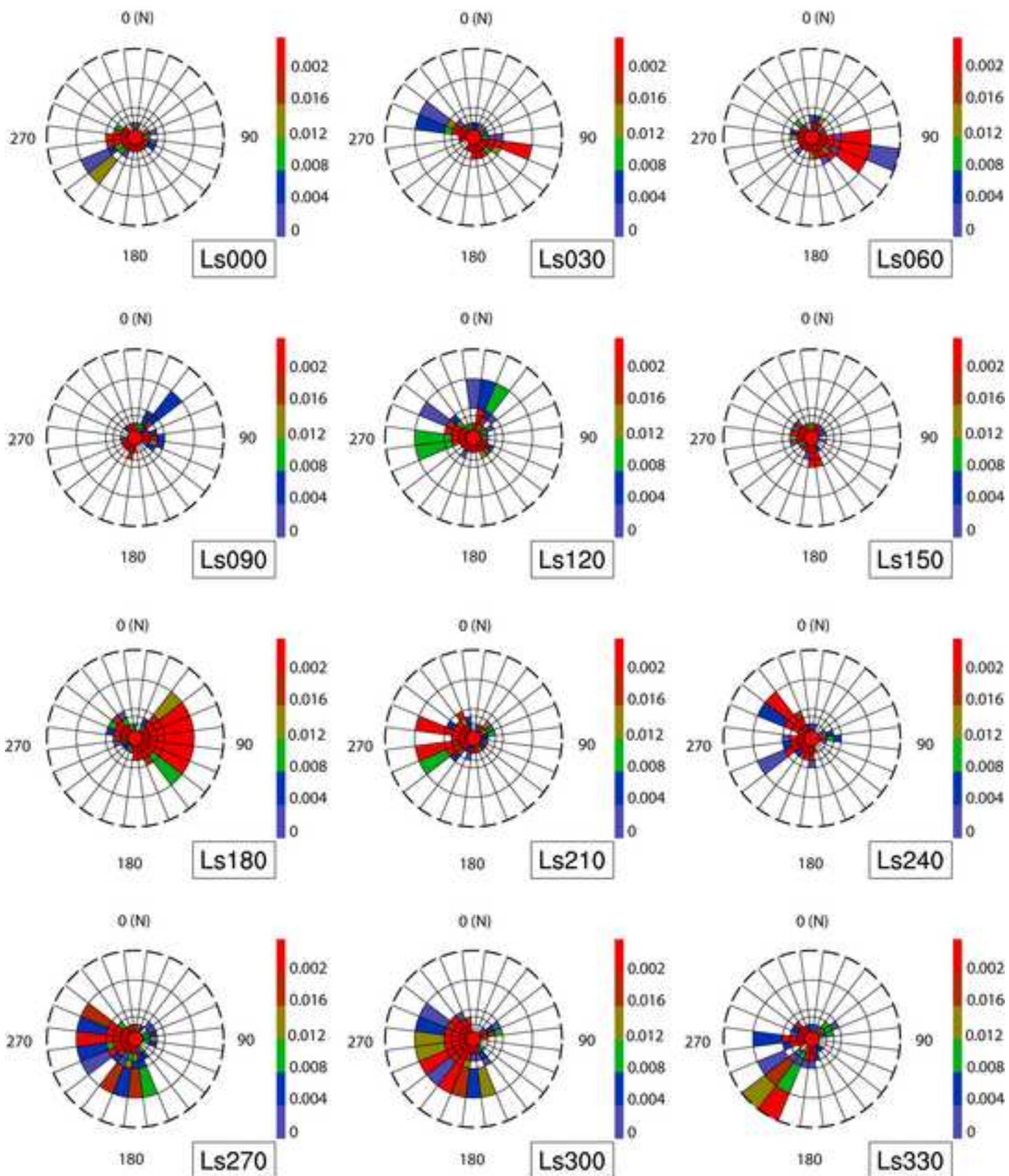


Figure12
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