



<b>Publication Year</b>	2018
<b>Acceptance in OA</b>	2021-04-23T10:43:59Z
<b>Title</b>	Dust evolution in galaxy cluster simulations
<b>Authors</b>	GJERGO, EDA, GRANATO, Gian Luigi, MURANTE, Giuseppe, Ragone Figueroa, Cinthia, TORNATORE, Luca, BORGANI, STEFANO
<b>Publisher's version (DOI)</b>	10.1093/mnras/sty1564
<b>Handle</b>	<a href="http://hdl.handle.net/20.500.12386/30873">http://hdl.handle.net/20.500.12386/30873</a>
<b>Journal</b>	MONTHLY NOTICES OF THE ROYAL ASTRONOMICAL SOCIETY
<b>Volume</b>	479

# Dust evolution in galaxy cluster simulations

Eda Gjergo,<sup>1,2★</sup> Gian Luigi Granato,<sup>2,3</sup> Giuseppe Murante,<sup>2,3</sup>  
Cinthia Ragone-Figueroa,<sup>2,3</sup> Luca Tornatore<sup>2</sup> and Stefano Borgani<sup>1,2</sup>

<sup>1</sup>*Dipartimento di Fisica dell' Università di Trieste, Sezione di Astronomia, via Tiepolo 11, I-34131 Trieste, Italy*

<sup>2</sup>*INAF, Osservatorio Astronomico di Trieste, via Tiepolo 11, I-34131, Trieste, Italy*

<sup>3</sup>*Instituto de Astronomía Teórica y Experimental (IATE), Consejo Nacional de Investigaciones Científicas y Técnicas de la República Argentina (CONICET), Observatorio Astronómico, Universidad Nacional de Córdoba, Laprida 854, X5000BGR Córdoba, Argentina*

Accepted 2018 June 8. Received 2018 April 11; in original form 2018 May 24

## ABSTRACT

We implement a state-of-the-art treatment of the processes affecting the production and Interstellar Medium (ISM) evolution of carbonaceous and silicate dust grains within SPH simulations. We trace the dust grain size distribution by means of a two-size approximation. We test our method on zoom-in simulations of four massive ( $M_{200} \geq 3 \times 10^{14} M_{\odot}$ ) galaxy clusters. We predict that during the early stages of assembly of the cluster at  $z \gtrsim 3$ , where the star formation activity is at its maximum in our simulations, the proto-cluster regions are rich in dusty gas. Compared to the case in which only dust production in stellar ejecta is active, if we include processes occurring in the cold ISM, the dust content is enhanced by a factor 2–3. However, the dust properties in this stage turn out to be significantly different from those observationally derived for the *average* Milky Way dust, and commonly adopted in calculations of dust reprocessing. We show that these differences may have a strong impact on the predicted spectral energy distributions. At low redshift in star-forming regions our model reproduces reasonably well the trend of dust abundances over metallicity as observed in local galaxies. However we underproduce by a factor of 2–3 the total dust content of clusters estimated observationally at low redshift,  $z \lesssim 0.5$  using *IRAS*, *Planck*, and *Herschel* satellites data. This discrepancy does not subsist by assuming a lower sputtering efficiency, which erodes dust grains in the hot intracluster medium.

**Key words:** methods: numerical – ISM: evolution – galaxies: clusters: general – galaxies: clusters: intracluster medium – galaxies: evolution.

## 1 INTRODUCTION

A significant fraction of metals present in the interstellar medium (ISM) is *depleted* from the gas phase and locked into small solid particles, the *cosmic dust*. The size of these particles is distributed over a broad range, from a few tens of Å up to a few  $\mu\text{m}$ . In the Milky Way, about 50 per cent of the metal mass, or about 1 per cent of the ISM mass, is in dust. Theoretical works (e.g. Dwek 1998) predict the first percentage to be roughly constant, and as a consequence the second is approximately proportional to the ambient gas metallicity  $Z_{\text{gas}}$ . This estimate is confirmed by observations of metal-rich galaxies, but it is subject to a broad scatter. On the other hand, low-metallicity dwarf galaxies, having a low dust-to-gas (DtG) mass ratio, deviate substantially (e.g. Galametz et al. 2011; Rémy-Ruyer et al. 2014). As for the detailed chemical composition of dust, as a first approximation it is accepted that dust consists of two major

chemical classes: one carbon-based, and another named ‘astronomical silicates’, composed of essentially four elements, O, Si, Mg, and Fe. This is supported by depletion and dust-reprocessing studies (e.g. Draine 2003; Jenkins 2009, and references therein).

Despite this deceptively reassuring summary of cosmic dust properties, it is clear that the situation is much more complex (e.g. Jones 2013, and references therein). All the above mentioned properties of dust represent only the average at late cosmic times. However, observations show significant deviations, both from galaxy to galaxy in the local Universe, and within different environments of the Milky Way itself. Moreover, observations suggest substantial differences at early cosmic times. From a theoretical perspective, this is all but surprising. Indeed, dust grains constitute a living component of the ISM. Once dust grain seeds are produced, mostly in stellar outflows, they are subject to several evolutionary processes in the ambient gas, whose effectiveness depends on the physical and chemical gas conditions, as well as on the properties of the grains themselves. These processes, which we will describe in Section 2.4,

\* E-mail: gjergo@oats.inaf.it

alter the abundance, chemical composition, and size distribution of dust grains.

Most galaxy properties cannot be described accurately without accounting for dust. For instance, the chemical species that dust depletes are key ISM coolants. Moreover, dust surfaces catalyse the formation of molecular species such as  $H_2$  (e.g. Barlow & Silk 1976).  $H_2$  is the primary constituent of molecular clouds (MCs), the star formation sites. Among all dust effects, the best studied is the dust reprocessing of radiation emitted by astrophysical objects. Therefore, dust is of paramount importance in interpreting observations. Dust absorbs efficiently stellar (or AGN) UV and optical light. The absorbed energy is thermally re-emitted in the IR, mostly at  $\lambda >$  a few  $\mu\text{m}$ , with a peak at about 100  $\mu\text{m}$ . In the local Universe, dust contributes only to less than 1 per cent of the ISM mass. Despite being so scarce, it reprocesses about 30 per cent of all the photons emitted by stars and AGNs (Soifer & Neugebauer 1991). The reprocessed fraction increases as a function of the specific star formation activity of galaxies (Sanders & Mirabel 1996). This is because in star-forming objects, the primary radiative power originates from young stars, which are close or still embedded in their parent MCs (Silva et al. 1998; Granato et al. 2000; Charlot & Fall 2000). Thus, IR turns out to be a very good observational tracer for star formation (e.g. Kennicutt & Evans 2012). It is worth noticing that the absorption and scattering cross sections of grains depend not only on their composition but also, and strongly, on their radius  $a$ .<sup>1</sup> Therefore, reliable galactic SED models must take dust size distribution into account (Silva et al. 1998).

In this work, we implement within the GADGET-3 SPH code a state-of-the-art treatment of the processes affecting the production, evolution, and destruction of carbonaceous and silicate dust grains. We model the size distribution of dust grains by means of the two-size approximation developed by Hirashita (2015). His work demonstrates that it is possible to reproduce the broad results of a full grain-size treatment just considering only two well-chosen representative sizes. The computation of this solution is not very demanding and therefore it is well-suited for cosmological simulations. Moreover, the method can be generalized in the future to increase the number of grain sizes.

While the GADGET-3 code is suitable for simulations of galaxy formation, we apply it here to zoom-in simulations of two massive ( $\sim 10^{15} M_\odot$ ) and two smaller ( $\sim 5 \times 10^{14} M_\odot$ ) galaxy clusters. We are mainly interested in the high-redshift stages, where the star formation activity is at its maximum, and the proto-cluster regions are rich of cold dusty gas. The first aim of our dust evolution model is to couple it in the near future with post-processing radiative transfer computations. We plan to replace the uncertain assumptions on the dust content, chemical composition, and size distribution, with estimates derived from the computation of these properties in the simulated ISM. The presence of dust production and evolution will allow, at a second stage, to account for the role of dust as a catalyst for the formation of molecules, as well as the impact of gas heating and cooling due to dust (Montier & Giard 2004).

A few groups have successfully included some aspects of the evolution of dust content of the ISM within hydrodynamical simulations (e.g. Bekki 2013; McKinnon, Torrey & Vogelsberger 2016; McKinnon et al. 2017; Aoyama et al. 2017; Hou et al. 2017). Another possibility is to investigate the problem by means of post-processing computations (e.g. Zhukovska et al. 2016) or within semi-analytic models (Popping, Somerville & Galametz 2017). However, this

work provides a more comprehensive description of dust as we predict both the size distribution and the chemical composition of dust grains *at the same time*, self-consistently with a full treatment of the chemical evolution of the ISM. Moreover, while the focus of the aforementioned papers was on galactic disks, our work represents the first attempt to trace the evolution of dust component in simulations of galaxy clusters.

We follow Dwek (1998) in using a notation for metallicity that specifies when we refer to metals in the gas form  $Z_{\text{gas}}$ , metals in the solid form, i.e. dust  $Z_{\text{dust}}$ , or the sum of the two  $Z_{\text{tot}}$ . This is not a standard convention as in other works dealing with the evolution of the dust in the ISM, the symbol  $Z$  has been employed to refer to the *total* metal fraction (metallicity) of the ISM (e.g. Calura, Pipino & Matteucci 2008; Hirashita 2015), including both metals in gas and metals locked up in solid state grains.

This paper is organized as follows: in Section 2, we describe the model. We give a brief overview of our cosmological simulation set and we present our implementation of the dust evolution model. In Section 3, we describe and analyse the outcome of said dust model for a variety of global and internal properties. We also offer a preliminary data comparison. Section 4 reports the conclusions of this work.

## 2 MODEL

After implementing a treatment of dust evolution, in this work we rerun and analyse four zoom-in simulations for the formation of massive galaxy clusters, taken from a sample already presented in a number of papers by our group. We recall briefly in this section the main aspects of these simulations, while we refer the reader in particular to Ragone-Figueroa et al. (2013) and Ragone-Figueroa et al. (2018) for further numerical and technical details. The modifications introduced to take into account the evolution of dust in the ISM are fully described in Section 2.4.

### 2.1 The parent sample of simulated clusters

The full set of initial conditions describe 29 zoomed-in Lagrangian regions, selected from a parent DM-only and low-resolution simulation of a  $1 h^{-1} \text{Gpc}$  box. The regions are built to safely contain 24 dark matter haloes with  $M_{200} > 8 \times 10^{14} h^{-1} M_\odot$  and 5 smaller ones with  $M_{200} = [1 - 4] \times 10^{14} h^{-1} M_\odot$ .<sup>2</sup> We adopt the following cosmological parameters:  $\Omega_m = 0.24$ ,  $\Omega_b = 0.04$ ,  $n_s = 0.96$ ,  $\sigma_8 = 0.8$  and  $H_0 = 72 \text{ km s}^{-1} \text{Mpc}^{-1}$ .

These regions are re-simulated with a custom version of the GADGET-3 code (Springel 2005) to achieve better resolution and to include the baryonic physics. The mass resolution for the DM and gas particles is  $m_{\text{DM}} = 8.47 \times 10^8 h^{-1} M_\odot$  and  $m_{\text{gas}} = 1.53 \times 10^8 h^{-1} M_\odot$ , respectively. Note that the latter value is the *initial* mass of gas particles. However, as described in Section 2.1.1, during the simulation these particles can decrease their mass by spawning up to four stellar particles, to simulate star formation. On the other hand, they can also gain mass due to gas ejecta produced by neighbourhood stellar particles. Thus, the mass of each gas particle evolve with time. At  $z = 0$ , the 5 per cent and 95 per cent percentiles of their mass distribution turns out to be 0.6 and 1.06 the initial values, respectively.

<sup>2</sup> $M_{200}$  ( $M_{500}$ ) is the mass enclosed by a sphere whose mean density is 200 (500) times the critical density at the considered redshift.

<sup>1</sup>Usually, dust grains are simply approximated by spheres.

For the gravitational force, a Plummer-equivalent softening length of  $\epsilon = 5.6 h^{-1}$  kpc is used for DM and gas particles, whereas  $\epsilon = 3 h^{-1}$  kpc for black hole and star particles. The DM softening length is kept fixed in comoving coordinates for  $z > 2$  and in physical coordinates at lower redshift. To treat hydrodynamical forces, we adopt the SPH formulation by Beck et al. (2016), that includes artificial thermal diffusion and a higher order interpolation kernel, which improves the standard SPH performance in its capability of treating discontinuities and following the development of gas-dynamical instabilities.

### 2.1.1 Unresolved physics

Our simulations include a treatment of all the unresolved baryonic processes usually taken into account in galaxy formation simulations. For details on the adopted implementation of cooling, star formation (SF), and associated feedback, we refer the reader to Planelles et al. (2014). In brief, the model of SF is an updated version of the implementation by Springel & Hernquist (2003), in which gas particles with a density above  $0.1 \text{ cm}^{-3}$  and a temperature below  $T_{\text{MPh}} = 5 \times 10^5$  K are classified as multiphase. Multiphase particles consist of a cold and a hot phase, in pressure equilibrium. The cold phase is the star formation reservoir. At each time-step, multiphase particles can stochastically spawn non-collisional star particles, each representing a simple stellar population, with expectation value consistent with the star formation rate. The latter is computed on the basis of the physical state of the particle. Each star-forming gas particle can generate up to  $N_g$  ( $=4$  in our simulations) star particles. The  $N_g$ -th event, when it occurs, consists in the complete conversion into a star particle. We implement kinetic feedback following the prescription by Springel & Hernquist (2003), in which a multiphase star-forming particle has a probability to be loaded into galactic outflows (we assume  $v_w = 350 \text{ km s}^{-1}$  for the outflow velocity). This probability is set to generate on average an outflow rate proportional to the star formation rate, with a proportionality factor of 2.

A full account of the AGN feedback model can be found in appendix A of Ragone-Figueroa et al. (2013), with some modifications described in Ragone-Figueroa et al. (2018). We refer the reader to those papers for the numerical details. In short, the BHs are represented by collisionless particles, seeded at the centre of a DM haloes when they become more massive than  $M_{\text{th}} = 2.5 \times 10^{11} h^{-1} M_{\odot}$  and do not already contain a SMBH. The initial BH mass at seeding is  $M_{\text{seed}} = 5 \cdot 10^6 h^{-1} M_{\odot}$ . We reposition at each time-step the SMBHs particle at the position of the nearby particle, of whatever type, having the minimum value of the gravitational potential within the gravitational softening. The SMBH grows with an accretion rate given by the minimum between the  $\alpha$ -modified Bondi accretion rate and the Eddington limit. The formula for the  $\alpha$ -modified Bondi accretion rate:

$$\dot{M}_{\text{Bondi},\alpha} = \alpha \frac{4\pi G^2 M_{\text{BH}}^2 \rho}{(c_s^2 + v_{\text{BH}}^2)^{2/3}} \quad (1)$$

is applied separately to the hot and cold gas components. The threshold between the two components is set at  $T = 5 \times 10^5$  K, and the adopted values of the fudge factor  $\alpha$  are 100 and 10, respectively. This distinction between cold and hot accretion modes has been inspired by the result of high-resolution AMR simulations of the gas flowing to SMBHs (Gaspari, Ruszkowski & Oh 2013). BHs particles merge when they fall within the gravitational softening. The accretion of gas onto the SMBH produces an energy deter-

mined by the radiative efficiency parameter  $\epsilon_r$ . Another parameter  $\epsilon_f$  defines the fraction of this energy that is thermally coupled to the surrounding gas. We calibrated these parameters to reproduce the observed scaling relations of SMBH mass in spheroids. Here, we set  $\epsilon_r = 0.07$  and  $\epsilon_f = 0.1$ . Finally, we assume a transition from a *quasar mode* to a *radio mode* AGN feedback when the accretion rate becomes smaller than a given limit,  $\dot{M}_{\text{BH}}/\dot{M}_{\text{Edd}} = 10^{-2}$ . In this case, we increase the feedback efficiency  $\epsilon_f$  to 0.7.

## 2.2 The model for chemical enrichment

Stellar evolution and metal enrichment follow Tornatore et al. (2007). The production of heavy elements considers separately contributions from Asymptotic Giant Branch (AGB) stellar winds, Type Ia Supernovae (SNIa) and Type II Supernovae (SNII), which is required also when implementing the production of dust from the different stellar channels (Section 2.4.1). Whereas all three types of stars contribute to the chemical enrichment, only SNIa and SNII provide thermal feedback. In addition, as described in Springel & Hernquist (2003), kinetic feedback from SNII is implemented as galactic outflows with a wind velocity of  $350 \text{ km s}^{-1}$ . The initial mass distribution for the star population is described by the initial mass function of Chabrier (2003). We assume the mass-dependent lifetimes of Padovani & Matteucci (1993) to account for the different time-scales of stars of different masses.

To estimate the production of metals due to the evolution of stellar particles, we consider different sets of stellar yields, as detailed in Biffi et al. (2017). We trace the production and evolution of 15 chemical elements: H, He, C, Ca, O, N, Ne, Mg, S, Si, Fe, Na, Al, Ar, and Ni. Consistently, metallicity-dependent radiative cooling of gas is calculated by taking into account the contribution of these 15 chemical species.

The code do not include any special treatment of metal diffusion, although it accounts for the spreading of metals (both in gaseous and dust form, see below) from star particles to the surrounding gas particles by using the same kernel of the SPH interpolation. This is mainly done in order to avoid a noisy estimation of metal-dependent cooling rates. Therefore, heavy elements and dust can be spatially distributed after that only via dynamical processes involving the metal-rich gas.

## 2.3 The test subsample for this work

To test our implementation of dust evolution, we select 4 of the 29 zoomed-in regions. Two of them, dubbed D2 and D3, belong to the small mass sample of five regions discussed above, while the other two, D1 and D6, are in the large mass sample. Their respective  $M_{200}$  at  $z = 0$  are 5.4, 6.8, 11.8, and  $17.5 \times 10^{14} M_{\odot}$ . Most of the following analysis refers to the main cluster of region D2.

## 2.4 Dust formation and evolution

We implemented within our version of GADGET-3 a treatment of dust evolution similar to that introduced by Hirashita (2015) and Aoyama et al. (2017), with some modifications. We anticipate them here for the sake of clarity, before detailing our approach in the rest of this section: (i) We take advantage of a detailed treatment of chemical evolution (Section 2.2), which tracks individually several heavy elements instead of just the global metallicity. Therefore, we can follow separately the fate of the two distinct dust chemical phases which are believed to populate the ISM, namely carbonaceous and silicate grains; (ii) rather than considering just the instantaneous

ISM dust pollution due to SNII explosions, we are able to properly compute also the delayed effect of SNIa explosions and of AGB stars winds; (iii) we take into account also the dust erosion caused by thermal sputtering, which becomes dominant in the hot intra-cluster medium (ICM).

As for the chemical composition of dust, we adopt the standard view, supported by depletion and radiative studies (e.g. Draine 2003; Jenkins 2009, and references therein), that dust composition is dominated by C, O, Si, Mg, and Fe. These elements are believed to be organized in two major chemical classes, one composed mostly by carbon and another one composed by 'astronomical silicates', comprising the four latter elements. More specifically, we accept that the former one is reasonably well approximated for our purposes by pure C grains, and the latter one by Olivine  $\text{MgFeSiO}_4$  grains (Draine 2003). We explicitly note that the majority of radiative transfer computations in astrophysical dusty media are done adopting optical properties of dust grains calculated according to these assumptions. However, it is also worth pointing out that our treatment can be immediately adapted to different dust composition models, as long as the relevant elements are included in the chemical evolution treatment described in Section 2.2.

As in Aoyama et al. (2017), we account for the grain size using the two-size approximation proposed by Hirashita (2015), which has been derived from the full treatment of grain size distribution put forward by Asano et al. (2013). In the two size approach, just two discrete size populations are considered to represent the whole continuum range of grain radii  $a$ . We refer to the two populations as large and small grains. This is justified by the fact that the various processes affecting dust population in the ISM act differently in these two representative populations. The boundary between them is set by Hirashita (2015) at  $a \simeq 0.03 \mu\text{m}$ . We thus track four types of dust grains: large and small carbonaceous grains, and large and small silicates grains.

It is worth pointing out that specifying the grain size distribution, albeit minimal, has a two-fold effect. It allows a more realistic treatment of the evolution of dust population, and it is also useful for post-processing computations involving radiative transfer that we plan to include in the next future. Indeed, the optical properties of grains strongly depends on their size. The two-size approximation allows for such a grain size treatment without affecting too heavily computation times and memory requirements. In the present application, the overhead of computing time arising from the inclusion of dust evolution is limited to less than 20 per cent.

This two-size approximation has been already adopted in previous simulations works (Hou, Hirashita & Michałowski 2016; Hou et al. 2017; Chen et al. 2018), which also tracked separately carbon and silicate dust. Note, however, that their simulations used a simplified treatment of chemical evolution, the Instantaneous Recycling Approximation (IRA). As a consequence, they included only SNII for the primary production of dust grains from stars, neglecting both AGB and SNIa. According to our results, based on a full treatment of chemical evolution, the SNII channel turns out to be insufficient to properly predict the relative abundance of silicate and carbon grains, and thus the radiative effect of dust (see Section 3.2.3).

Previous studies have identified the main processes that should be taken into account to understand the dust content of galaxies. The life cycle of grains begins with *dust production* in the ejecta of stars, including the relatively quiescent winds of AGB stars, but also SNII and SNIa explosion. It can be assumed that dust production affects directly only large grains. Indeed, SNe shocks are more efficient at the destruction of small grains than of large grains, therefore over chemical evolution time-scales only the latter

will survive (Nozawa et al. 2007; Bianchi & Schneider 2007). As for AGB winds, their infrared SEDs suggest that the typical grain size is skewed toward large grains (Groenewegen 1997; Gauger et al. 1999). These findings are supported by expectations of models including dust formation in AGB winds (e.g. Winters et al. 1997).

Once injected into the ISM, the grain population is subject to several evolutionary processes. Metals from the surrounding gas can deposit on the grain surface, causing the grains to gain mass, in a process named *accretion* (Dwek 1998; Hirashita 1999). An important role of this process has been invoked to explain the remarkable amount of dust in high- $z$  quasars and starbursts (e.g. Michałowski et al. 2010; Rowlands et al. 2014; Nozawa et al. 2015, and references therein), which cannot be accounted by stellar production only. Being a surface process, it acts more efficiently on small grains, which have larger total surface per unit mass. However, accretion is not efficient enough to increase significantly a small grain's radius. Its effect is just to increase the amount of small grains, while the evolution of small grains into large grains by accretion is safely negligible. Grain-grain collision can either result in grain *coagulation* or *shattering*. The former process dominates in dense ISM, and can be regarded as a destruction mechanism for small grains and a formation mechanism for large ones. The latter mechanism wins in the diffuse ISM, leading to the formation of small grains and destruction of large ones.

Sputtering, also a surface process, consists of grain erosion due to collisions with energetic ions. The eroded atoms are given back to the gas phase. It occurs both when grains are swept by SNe shocks, and when they are subject to the harsh ion collisions in the hot ICM. On top of SNe dust destruction, we include thermal sputtering in the ICM, which is important only when grains are surrounded by hot gas  $T \gtrsim 10^6$  K. Hirashita (2015) and Aoyama et al. (2017) did not consider ICM thermal sputtering, because they were interested only in galactic ISM. We refer to this process simply as *sputtering*, whether we call the former one *SNe destruction*.

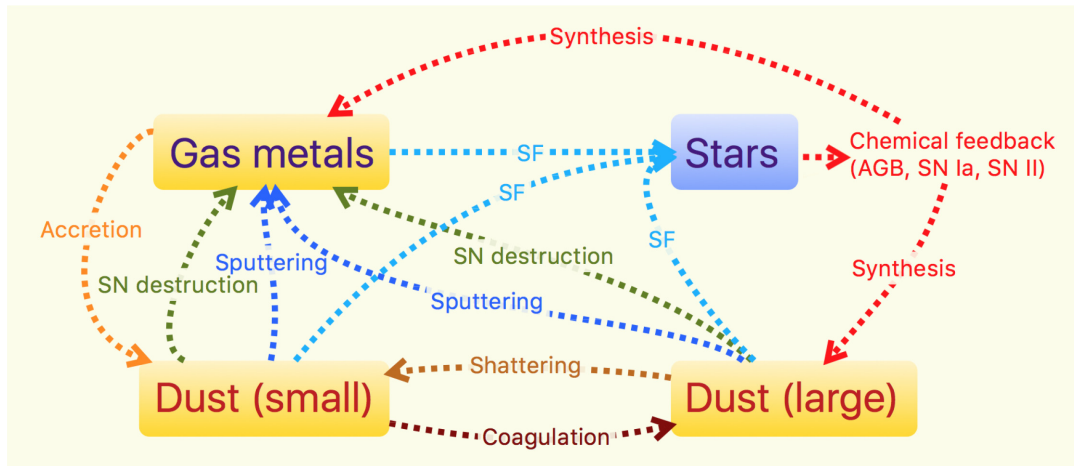
Needless to say, when gas is turned into stars, its dust content is proportionally subtracted from the ISM. This event is often referred to as *astration*. A visual summary of how the various processes act on the components of a galaxy is provided in Fig. 1.

We implement this sequence of processes at the subgrid level. A variable fraction of every metal produced by star particles is given to neighbour gas particles in the form of solid state dust rather than gas. Accordingly, we modified the gas particle structure to trace two extra vectors along with the gas metals: one for large and one for small grains. The three vectors have a length equal to the number of chemical species we trace, which are 15 in this paper. Thus, a certain fraction of the mass of each SPH 'gas' particle, never exceeding a few per cent, is actually used to represent metals which are locked up in solid grains. We still refer to them simply as SPH gas particles.

For a SPH gas particle of total mass  $M_{\text{gas}}$ , the time derivatives of the dust mass content in the two size populations of small grains  $M_{\text{d,S}}$  and large grains  $M_{\text{d,L}}$  can be expressed as follows:

$$\begin{aligned} \frac{dM_{\text{d,L}}}{dt} &= \frac{dM_{P*}}{dt} - \frac{M_{\text{d,L}}}{\tau_{\text{sh}}} + \frac{M_{\text{d,S}}}{\tau_{\text{co}}} - \frac{M_{\text{d,L}}}{\tau_{\text{SN,L}}} - \frac{M_{\text{d,L}}}{\tau_{\text{sp,L}}} - \frac{M_{\text{d,L}}}{M_{\text{gas}}} \psi \\ \frac{dM_{\text{d,S}}}{dt} &= \frac{M_{\text{d,S}}}{\tau_{\text{acc}}} + \frac{M_{\text{d,L}}}{\tau_{\text{sh}}} - \frac{M_{\text{d,S}}}{\tau_{\text{co}}} - \frac{M_{\text{d,S}}}{\tau_{\text{SN,S}}} - \frac{M_{\text{d,S}}}{\tau_{\text{sp,S}}} - \frac{M_{\text{d,S}}}{M_{\text{gas}}} \psi, \end{aligned} \quad (2)$$

where  $\frac{dM_{P*}}{dt}$  is the dust enrichment arising from production by stars, and local time-scales for the various dust processes occurring in the



**Figure 1.** Diagram of mass flows due to dust processes. Boxed in yellow are gas metals, large and small dust, which belong to the gas particle structure. Boxed in blue is the star particle, responsible for the production of metals in the subgrid chemical evolution through AGB winds, SN Ia, and SN II. Star particles spread metals to the surrounding gas particles by enriching gas metals and large dust alike. When a gas particle is ready to form a star particle, all three of the gas particle metal channels are depleted in proportion, in favour of the star particle. SN destruction and sputtering subtract metals from both dust channels and enrich the gas metals.

ISM are introduced:  $\tau_{\text{sh}}$  for shattering,  $\tau_{\text{co}}$  for coagulation,  $\tau_{\text{SN,L}}$  and  $\tau_{\text{SN,S}}$  for SN shock destruction,  $\tau_{\text{sp}}$  for sputtering, and  $\tau_{\text{acc}}$  for accretion. The last term of each equation represents the dust mass loss due to star formation (astration), specifically  $\psi$  is the star formation rate. These equations are applied separately for each element entering into the dust grains. This implies that if a gas particle has gas-phase metals for accretion of silicates but not gas phase C (a quite unlikely possibility, in particular at our relatively coarse resolution), only the mass of small silicates is affected by accretion, but not that of C grains. In the next subsections, we give details on how each of these terms is computed in our simulations.

#### 2.4.1 Dust production by stars

We adapt to each simulation star particle the formalism for dust production by stars introduced by Dwek (1998), in the context of *monolithic* galaxy formation models. However, we introduce an important modification related to our specific assumption for the chemical composition of silicate grains.

Dust is produced by the same stellar enrichment channels which are responsible of metals production, namely AGB winds, SNIa, and SNI. We assume that a certain fraction of the elements concurring to dust formation is originally given to the ISM locked to solid dust particles rather than in gaseous form. These elements are C, Si, O, Mg, and Fe. This is because, as motivated at the beginning of Section 2.4, we consider carbon and silicate dust, and we approximate silicates with olivine  $\text{MgFeSiO}_4$ . In this exploratory work, we maintain for simplicity the assumption that, for each stellar channel the fraction dust elements that condense to dust does not depend on stellar mass nor on metallicity. This approach, introduced in one-zone models by Dwek (1998), has been adopted by many later computations (e.g. Calura et al. 2008; Hirashita 2015; Aoyama et al. 2017). However, we point out that there are several works that studied how dust condensation efficiencies actually depend on stellar mass, metallicity, and ambient gas density (see e.g. Nozawa et al. 2007; Schneider et al. 2014, and references therein). In the future, it will be relatively straightforward to incorporate the results of such computations into our formalism.

#### 2.4.2 AGB stars

As for AGB winds, following Dwek (1998) we assume that the formation of carbon or silicate dust is mutually exclusive, and depends on the C/O number ratio in the ejecta. When  $\text{C/O} > 1$ , all the oxygen is tied up in CO molecules and only carbon dust is formed. If instead  $\text{C/O} < 1$ , all the carbon is consumed to form CO molecules, while the excess oxygen will combine with Si, Mg, and Fe to produce mainly silicate grains. This view is broadly supported by many observations (e.g. Kastner et al. 1993), and implies the assumption that AGB ejecta are mixed at microscopic level, so that the maximum possible amount of CO is formed.

Let  $M_{\text{ej,C}}^{\text{AGB}}$  and  $M_{\text{ej,O}}^{\text{AGB}}$  be the C and O masses ejected by a star particle through the AGB enrichment channel during a given time-step. Carbon (silicate) grains form when  $M_{\text{ej,C}}^{\text{AGB}}$  is greater (smaller) than  $0.75 M_{\text{ej,O}}^{\text{AGB}}$ , where 0.75 is the ratio between O and C atomic weights. Due to the C mass lost to CO, the mass of carbon dust produced during the time-step is given by

$$M_{\text{dust,C}}^{\text{AGB}} = \max \left[ \delta_{\text{AGB,C}} \left( M_{\text{ej,C}}^{\text{AGB}} - 0.75 M_{\text{ej,O}}^{\text{AGB}} \right), 0 \right], \quad (3)$$

where  $\delta_{\text{AGB,C}}$  is a *condensation efficiency* of carbon grains in AGB winds, which we set equal to 1, as in Dwek (1998) (see also Calura et al. 2008).

If  $M_{\text{ej,C}}^{\text{AGB}} < 0.75 M_{\text{ej,O}}^{\text{AGB}}$ , silicate grains are instead produced. Dwek (1998) estimated the masses of the elements going into silicate dust under the simple assumption that for each atom of them, a single O atom will go into dust as well (see his equation 23). While we have implemented in our code also his formulation, we prefer a more specific assumption on the chemical composition of silicates produced by stars, which leads us to different formulae. As expected, we have verified with our code that Dwek (1998) approach leads to the production of ‘silicate grains’ featuring very variable mass ratios between the various elements, and typically very different from those implicitly assumed by most radiative transfer computations. This is particularly true for the SNIa channel, since these stellar events produce mostly iron, much in excess to the quantity that can be bound directly with O in the SNe ejecta. When we use the formalism by Dwek (1998), dust grains reach a mass fraction for Fe which is double that of olivine. Mg is underabundant by a factor

of 3. Our method on the other hand, described below, preserves the olivine-like mass fractions in the four metal components at all redshifts.

Generalizing the notation introduced above, we indicate as  $M_{ej,X}^{AGB}$  the masses ejected by a star particle, through the AGB enrichment channel and during a given time-step, in form of the X element which enters into the olivine compound. X stands then for Mg, Fe, Si, and O. We indicate  $N_{mol,sil}^{AGB}$  as the number of 'molecules' of  $MgFeSiO_4$  that can be formed over the time-step. This is set by the less abundant element, taking into account how many atoms of each element  $N_{ato}^X$  enter into the compound (1 for Mg, Fe, Si, and 4 for O). Then

$$N_{mol,sil}^{AGB} = \delta_{AGB,sil} \min_{X \in \{Mg, Fe, Si, O\}} \left( \frac{M_{ej,X}^{AGB}}{\mu_X N_{ato}^X} \right), \quad (4)$$

where  $\mu_X$  is the atomic weight of the X element, and we have introduced an efficiency factor of condensation for silicate grains  $\delta_{AGB,sil}$ , set to 1 in our reference model. Therefore, the mass of the X element locked into silicate grains is given by

$$M_{dust,X}^{AGB} = \begin{cases} N_{mol,sil}^{AGB} \mu_X N_{ato}^X & \text{for } \frac{M_{ej,X}^{AGB}}{M_{ej,O}^{AGB}} < 0.75. \\ 0 & \text{otherwise} \end{cases} \quad (5)$$

It is straightforward to adapt the above treatment to chemical compounds different from olivine.

#### 2.4.3 SNe II and Ia

Unlike AGB winds, it is reasonable to assume that the explosive outflows produced by SNe are mixed only at a macroscopic level (e.g. Dwek 1998, and references therein), which implies that these channels can produce both carbon and silicate dust at the same time. Hence:

$$M_{dust,C}^{SNx} = \delta_{SNx} M_{ej,C}^{SNx} \quad (6)$$

$$M_{dust,X}^{SNx} = N_{mol,sil}^{SNx} \mu_X N_{ato}^X \quad (7)$$

$$N_{mol,sil}^{SNx} = \delta_{SNx,sil} \min_{X \in \{Mg, Fe, Si, O\}} \left( \frac{M_{ej,X}^{SNx}}{\mu_X N_{ato}^X} \right), \quad (8)$$

where SNx stands either for SNIi or SNIa, and  $M_{ej,X}^{SNx}$  is the mass of the generic element X ejected by a star particle through the SNx enrichment channel during a given time-step. In the case of SNe, we decrease the dust condensation efficiency to  $\delta_{SNIi,C} = \delta_{SNIa,C} = 0.5$  and  $\delta_{SNIi,sil} = \delta_{SNIa,sil} = 0.8$ . These values are the same as those adopted in Dwek (1998) and are meant to account for the grains destroyed by the SN shock or incomplete condensation of available elements.

#### 2.4.4 Shattering

In the diffuse gas, large grains tend to have high velocities (typically  $v = 10 \text{ km s}^{-1}$ , Yan, Lazarian & Draine 2004) due to decoupling from small-scale turbulent motions (Hirashita & Yan 2009). Therefore, in such environments, large grains are most prone to shattering due to collision forming small grains. In our framework, shattering is the process which originates small grains, without directly affecting the total dust mass. The growth of small grain mass occurs by accretion (Section 2.4.5). The shattering time-scale is derived following the

prescription in appendix B of Aoyama et al. (2017). The gas environment diffuse enough to promote efficient shattering is identified by the condition  $n_{gas} < 1 \text{ cm}^{-3}$  (see Hirashita & Yan 2009), in which case:

$$\tau_{sh} = \tau_{sh,0} \left( \frac{0.01}{D_L} \right) \left( \frac{1 \text{ cm}^{-3}}{n_{gas}} \right) \quad \text{if } n_{gas} < 1 \text{ cm}^{-3}, \quad (9)$$

where  $D_L = M_{dust,L}/M_{gas}$  is the dust to gas ratio for large grains. The value of the proportionality constant suggested by Aoyama et al. (2017) is  $\tau_{sh,0} = 5.41 \times 10^7 \text{ yr}$ , obtained assuming the typical grain velocity dispersion quoted above  $v = 10 \text{ km s}^{-1}$ , a grain size of  $0.1 \mu\text{m}$ , and a material density of grains  $3 \text{ g cm}^{-3}$ .

#### 2.4.5 Accretion and coagulation in dense molecular gas

Accretion of gas metals onto dust grains as well as coagulation of grains are processes whose time-scale is inversely proportional to the gas density, and become significant only in the densest regions of molecular clouds, at  $n_H \gtrsim 10^2 - 10^3 \text{ cm}^{-3}$  (see Hirashita & Voshchinnikov 2014, equations 7 and 8). These high densities are unresolved in cosmological simulations, therefore we have to resort to a sub-resolution prescription to estimate the fraction of gas  $f_{dense}$  that is locally in this condition. Aoyama et al. (2017) simply assumed a fixed fraction. Although for their simulated galaxy the results are weakly sensitive to the adopted value, this solution could be in general unsatisfactory. In fact,  $f_{dense}$  should depend on the local condition of the gas and on the numerical resolution reached in a simulation. We thus introduce a more flexible approach. According to the results of high resolution ( $\lesssim 10 \text{ pc}$ ) simulations (e.g. Wada & Norman 2007; Tasker & Tan 2009), the probability distribution functions of ISM density is well described by a log-normal function, characterized by a dispersion  $\sigma$ , found to lie in the range 2–3, and a number density normalization parameter  $n_0$ .

$$f_{pd}(n) dn = \frac{1}{\sqrt{2\pi}\sigma} \exp \left[ -\frac{\ln(n/n_0)^2}{2\sigma^2} \right] d \ln n. \quad (10)$$

Therefore the mass fraction  $f(>n_{th})$  of ISM gas above a given density threshold  $n_{th}$  is (Wada & Norman 2007, equation 19)

$$f(>n_{th}) = \frac{1}{2} \left[ 1 - \text{erf} \left( \left[ \ln \left( \frac{n_{th}}{n_0} \right) - \sigma^2 \right] (\sqrt{2}\sigma)^{-1} \right) \right]. \quad (11)$$

We use this equation in order to estimate  $f_{dense} \equiv f(>n_{th})$  for each SPH gas particle, we adopt  $\sigma = 2.5$  and  $n_{th} = 10^3 \text{ cm}^{-3}$  in our fiducial model. In order to set the local value of the density normalization  $n_0$ , we take into account that for the distribution given by equation (10), the mass averaged density is  $n_M = n_0 \exp(2\sigma^2)$  and we identify  $n_M$  with a suitable number density associated to the SPH particle. Since conditions suitable for accretion and coagulation can only be achieved in star-forming multiphase particles (see Section 2.1), we set  $n_M$  equal to the density of the cold phase  $n_{cold}$ . In the multiphase models of SF by Springel & Hernquist (2003),  $n_{cold}$  is computed from the condition of pressure equilibrium between the two phases, and assuming  $T_{cold} = 1000 \text{ K}$ , while the temperature of the ionized hot phase  $T_{hot}$  is the SPH gas temperature. In our simulations, both the median and the standard deviation of  $n_{cold}$  distribution increases with redshift, and range (in  $\text{cm}^{-3}$ ) from about  $\sim 1$  to  $\sim 20$  and from a few tens to a few hundreds, respectively.

As already remarked, in the two size approximations, only small grains increase their mass when gas metals deposit on their surface. Following Aoyama et al. (2017), we use the following expression

to estimate the corresponding time-scale:

$$\tau_{\text{acc}} = \tau_{\text{acc},0} \left( \frac{Z_{\odot}}{Z_{\text{gas}}} \right) f_{\text{dense}}^{-1} f_{\text{clid}}^{-1}, \quad (12)$$

where  $f_{\text{clid}}$  is the fraction of cold gas in multiphase particles (MP). Indeed this equation does not apply to single-phase particles, nor to the hot phase of multiphase particles, since their density is far too low for the process to occur. Aoyama et al. (2017) set  $\tau_{\text{acc},0} = 1.2 \times 10^6 \text{ yr}$ .

In MP particles, small grains can also coagulate to form large grains. The time-scale is given by (see Aoyama et al. 2017)

$$\tau_{\text{co}} = \tau_{\text{co},0} \left( \frac{0.01}{D_S} \right) \left( \frac{0.1 \text{ km s}^{-1}}{v_{\text{co}}} \right) f_{\text{dense}}^{-1} f_{\text{clid}}^{-1}, \quad (13)$$

where  $D_S = M_{\text{dust},S}/M_{\text{gas}}$  is the dust-to-gas ratio for small grains and  $v_{\text{co}}$  is the velocity dispersion of small grains, for which a value of  $0.1 \text{ km s}^{-1}$  is adopted based on calculations from Yan et al. (2004). Just like with  $\tau_{\text{acc}}$ , the equation has been modified to consider only the cold phase of multiphase particles. Aoyama et al. (2017) used  $\tau_{\text{co},0} = 2.71 \times 10^5 \text{ yr}$ , a value derived assuming a typical size of small grains of  $0.005 \mu\text{m}$ , a material density of  $3 \text{ g cm}^{-3}$  and  $n_{\text{gas}} = 10^3 \text{ cm}^{-3}$ .

#### 2.4.6 SNe destruction

$\text{SNIa/SNII}$  Dust grains are efficiently destroyed by thermal and non-thermal sputtering in SNe shocks (for a review, see McKee 1989). To avoid confusion, hereafter we refer to destruction by SNe events as SNe destruction, and we reserve the term sputtering to what we describe in Section 2.4.7. To treat this process, we follow Aoyama et al. (2017). Let  $N_{\text{SN}}$  be the number of supernovae exploding over a given integration time-step  $\Delta t$ . For simplicity, we do not distinguish by now between the effects of SNII and SNIa. However, the code can be trivially adapted to take into account different values of the relevant parameters for the two classes, since the respective  $N$  and  $N$  are already computed separately. The time-scale for the process in the time-step can be written as

$$\tau_{\text{SN}} = \frac{\Delta t}{1 - (1 - \eta)^{N_{\text{SN}}}}, \quad (14)$$

where

$$\eta = \epsilon_{\text{SN}} \min \left( \frac{m_{\text{SW}}}{m_g}, 1 \right). \quad (15)$$

In this expression,  $m_g$  is the gas mass of the SPH particle,  $m_{\text{SW}}$  is the gas mass swept by a single SN event,<sup>3</sup> and  $\epsilon_{\text{SN}}$  is the efficiency of grain destruction in the shock, that we set to the same reference value 0.1 used by Aoyama et al. (2017). The shocked gas mass has been estimated in McKee (1989) as:

$$m_{\text{SW}} = 6800 M_{\odot} \left( \frac{E_{\text{SN}}}{10^{51} \text{ erg}} \right) \left( \frac{v_s}{100 \text{ km s}^{-1}} \right)^{-2}, \quad (16)$$

where  $E_{\text{SN}}$  is the energy from a single SN explosion and  $v_s$  the shock velocity. We adopt the fiducial value of  $E_{\text{SN}} = 10^{51} \text{ erg}$ . McKee et al. (1987) give the for  $v_s$  the following expression:

$$v_s = 200 \text{ km s}^{-1} (n_0/1 \text{ cm}^{-3})^{1/7} (E_{\text{SN}}/10^{51} \text{ erg})^{1/14}. \quad (17)$$

However, since the number density of ambient gas for single SN blasts is not resolved, we simply set  $v_s = 200 \text{ km s}^{-1}$ . As a result,  $\eta = 170 M_{\odot}/M_g$  in our simulations.

<sup>3</sup>In our relative low resolution simulation  $m_s \ll m_g$ .

#### 2.4.7 Thermal sputtering

Dust grains can be efficiently eroded by collisions with energetic ions. This process becomes important whenever the ambient gas is hot enough,  $T_g \gtrsim 10^6 \text{ K}$ , in which case it is dubbed *thermal sputtering*. To describe it, we employ the analytical formula given by Tsai & Mathews (1995). This is an accurate enough approximation of detailed calculations for both carbonaceous and silicate grains, at least for gas temperature smaller than a few times  $10^7 \text{ K}$ . Above this temperature, the process efficiency tends to stall (Tielens et al. 1994). Taking this into account and combining equations 14 and 15 of Tsai & Mathews (1995), we get:

$$\tau_{\text{sp}} = \tau_{\text{sp},0} \left( \frac{a}{0.1 \mu\text{m}} \right) \left( \frac{0.01 \text{ cm}^{-3}}{n_g} \right) \times \left[ \left( \frac{T_{\text{sp},0}}{\max(T_g, 3 \times 10^7 \text{ K})} \right)^{\omega} + 1 \right] \quad (18)$$

with  $T_{\text{sp},0} = 2 \times 10^6 \text{ K}$ ,  $\omega = 2.5$  and  $\tau_{\text{sp},0} = 5.5 \times 10^6 \text{ yr}$ . In this equation,  $n_g = \rho/\mu m_p$  is the number density including both ions and electrons. To derive the normalization constant  $\tau_{\text{sp},0}$ , we adopted the mean molecular weight  $\mu = 0.59$  of a mixture of 75 per cent H and 25 per cent He, both fully ionized. Moreover, we have taken into account that for spherical grains the mass variations time-scale is related to that for radius variations by  $m/\dot{m} = a/3\dot{a}$ . We adopt  $a = 0.05 \mu\text{m}$  and a 10 times smaller value for the typical radii of large and small grains, respectively. The former is the average radius for a power law size distribution with index  $-3.5$ , ranging from  $0.03 \mu\text{m}$  (the adopted boundary between small and large grains) to  $0.25 \mu\text{m}$  (e.g. Silva et al. 1998).

#### 2.5 Fiducial run and its variations

In Table 1, we list a selection of interesting runs discussed in the paper. Most of the analysis refers to the 'fiducial' run `fid`, for which we essentially adopted the same parameters values as Aoyama et al. (2017). These authors simulated an isolated spiral galaxy, and found results broadly consistent with spatially resolved observations of nearby dusty galaxies. It is however worth pointing out that we are working at resolutions lower by about 3 orders of magnitude in gas particle mass. Moreover, our model has two significant differences with respect to their work: we treat separately the production and evolution of the two distinct chemical species of carbon and silicate grains, and we include a treatment of thermal sputtering in the hot ICM. The latter process is crucial for galaxy cluster environments. Nevertheless, we believe that the parameter values defined by Aoyama et al. (2017) are a good starting point for our work. We will not explore the full dependence of dust properties on the parameters. Such work is postponed to future papers. As for sputtering, we adopt in the fiducial run parameter value derived from the classic work by Tielens et al. (1994).

In order to separate the other dust evolutionary processes from dust production and sputtering, we performed runs `f-crsp`, where only the latter processes are turned on, and `f-nosp` in which all processes but sputtering and SN destruction are active. To compare our dust production method, conserving the element fraction of olivine, to that by Dwek (1998) we run `f-dw` in which his method is instead employed. We also performed a run `f-snII` where only SNII contributes to dust production, but not SNIa nor AGBs. With this run we test, and to some extent disprove, the claim by Aoyama et al. (2017) that dust production by SNII is sufficient to predict the

**Table 1.** List of the test runs discussed in this paper with their respective parameters. Column 1 represents the run name used in the text. Column 2 is the chosen initial condition region. Column 3 is the  $z = 0$   $M_{200}$  of the main cluster of the region. Column 4 is the dust production method from stars: with ‘new’ we refer to the prescription presented in this paper, assuring that dust grains at production respect a given proportion of O, Mg, Fe, and Si, namely that of olivine  $\text{MgFeSiO}_4$  in this work, ‘dwek’ is the prescription proposed by Dwek (1998) (see Section 2.4.1 for details). Finally, with ‘SNII-only’ we refer to runs in which stellar dust production is active only for the SNII channel, as done by Aoyama et al. (2017). This is to test the relative importance of the other two channels SNIa and AGB used in our full implementation. Columns 5–9 are the normalization time-scales defined in the text for each ISM evolution process (2.4).

Run name	IC	$M_{z=0}^{200}$ ( $10^{14} M_{\odot}$ )	Production	$\tau_{\text{sh},0}$ (Myr)	$\tau_{\text{co},0}$ /Myr)	$\tau_{\text{acc},0}$ (Myr)	$\tau_{\text{sp},0}$ (Myr)
fid	D2	5.41	New	54.1	0.271	1.20	5.5
f-crsp	..	..	..	Off	Off	Off	5.5
f-nosp	..	..	..	54.1	0.271	1.20	Off
f-sp.2	..	..	..	..	..	..	27.5
f-snII	..	..	SNII-only	..	..	..	5.5
f-dw	..	..	Dwek	..	..	..	..
f-D3	D3	6.80	New	..	..	1.20	..
f-D6	D6	15.5	..	..	..	..	..
f-D1	D1	17.5	..	..	..	..	..

dust properties. In *f-sp.2* we increase the sputtering time-scale by a factor 5, for illustrative reasons discussed in Section 3.3.

These different runs were performed on the small mass region D2 of our set of initial conditions. The fiducial model was also applied to the other three regions D3, D1, and D6 (See Section 2.3). Unless otherwise specified, the analysis that follows applies to the fiducial run *fid*.

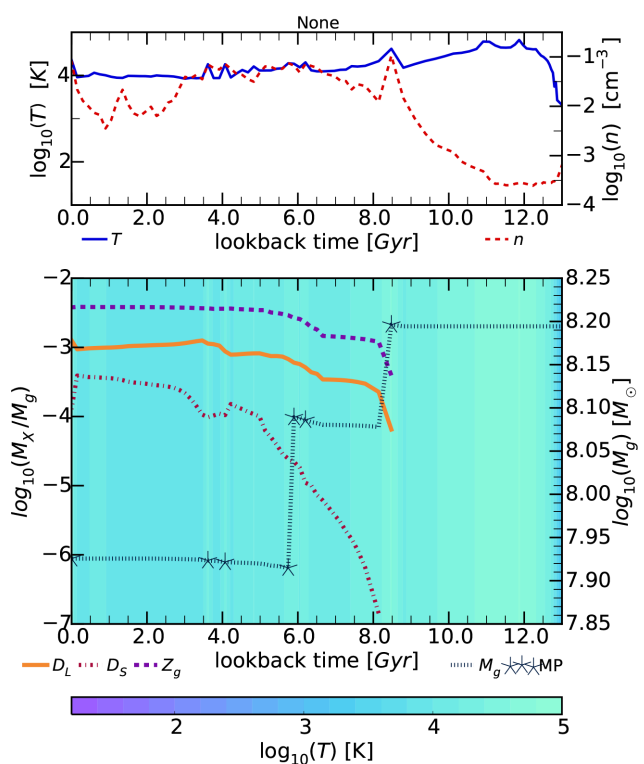
### 3 RESULTS

In this section, we study the behaviour of the dust model within the simulation. First, we follow the evolution of a couple token SPH gas particles representative of two interesting environments. A single gas particle could be thought of as the ISM of a one zone-model, such as those put forward by Dwek (1998) and Hirashita (2015). Then, we investigate the evolution of global properties of dust in the main cluster region. We conclude with a preliminary comparison to observations.

#### 3.1 Inside individual gas particles

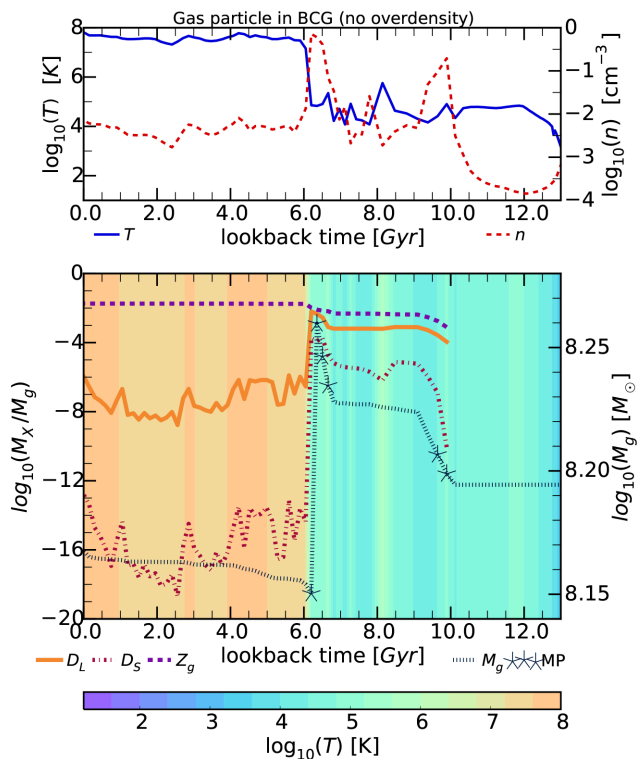
In Fig. 2, we follow the dust, gas metal, total mass, temperature, and density of a simulated gas particle which by  $z = 0$  is located within an isolated galaxy at the periphery of a cluster in our fiducial run. This particle features a relatively quiet evolution. It spawns two star particles, around a lookback time ( $t_{\text{fb}}$ ) of 8.5 and 6 Gyr, as a consequence of its star-forming state. The particle’s mass drops to about 75 per cent and 50 per cent of its original value during each episode (see Section 2.1; dotted dark blue line of the bottom panel). It begins to be enriched in gas metals and large grains by the neighbouring stellar particles at a  $t_{\text{fb}}$  of 8.5 Gyr, around the time of star formation. Soon after that, shattering begins to produce a population of small grains. Its growth rate is steady for the first 4 Gyr, thanks to a fairly constant gas density (red dashed line in the top panel). The gas temperature does not deviate significantly from  $10^4$  K throughout its history, whether the gas density starts low at  $3 \times 10^{-3} \text{ cm}^{-3}$  and then fluctuates by 1 dex below  $10^{-1} \text{ cm}^{-3}$ . The second star particle formation episode at 6 Gyr does not affect the DiG and metallicity, as these quantities decrease proportionally to the mass loss.

While the particle is multiphase, accretion sticks gas metals onto small grains and small grain coagulation in turn convert small grains to large ones. The seven multiphase snapshots are highlighted by a star symbol on the mass line of the bottom panel. The last three



**Figure 2.** For a gas particle in the fiducial run residing in a small, quiet, peripheral galaxy by  $z = 0$ : Top and bottom plots represent the time evolution ( $x$ -axis  $t_{\text{fb}}$  time in Gyr, with 0 being today) of various gas properties: (top plot) temperature (left  $y$ -axis, blue solid line) and number density (right  $y$ -axis, red dashed line). (Bottom plot) Gas metallicity (left  $y$ -axis, dashed purple line), as well as the DiG for large dust grains (left  $y$ -axis, solid orange line) and small dust grains (left  $y$ -axis, dot-dashed red line), total mass (right  $y$ -axis, dotted dark blue line) of the gas particle. The stars mark the snapshots in which the gas particle was captured in multiphase state (MP).

ones, two around 4 Gyr in  $t_{\text{fb}}$  and the third very recent, occurs while the small grain abundance is within 1 dex to that of large grains. In this case, coagulation prevails over accretion, and therefore a depletion of small grains is visible, whilst a tiny bump in the large grains mass ratio can be recognized. Note that it is possible that the particle underwent other multiphase periods between two subsequent snapshots. Similar periods are not highlighted by star symbols in the plot.



**Figure 3.** Same as Fig. 2 but for a gas particle residing in the BCG by  $z = 0$ . The colourbar gradient for temperature is fixed for both of the lower plots in the two figures.

Fig. 3 represents the evolution of another particle which instead ends up in the brightest cluster galaxy (BCG) by  $z = 0$ . This gas particle is first enriched with gas metals and large grains as soon as it becomes multiphase and enters a cold overdensity of star-forming gas. This period lasts a few hundred million years around 10 Gyr in  $t_{\text{fb}}$ . The particle density spikes, and with it shattering increases and produces small grains. At 8 Gyr, the particle temperature increases above  $10^5$  K for a while, increasing the efficiency of sputtering enough to cause a small dip of DtG, visible (almost) only for small grains. This is because the sputtering time-scale for small grains is 10 times smaller than that for large grains. At a  $t_{\text{fb}}$  of about 6.5 Gyr, the particle enters another cold overdensity. At this time, the large grains abundance approaches values, similar to the expected ISM values in galaxies (Li & Draine 2001), close to the gas metal abundance via coagulation, and also small grains manage to rise via accretion. Shortly after, at 6 Gyr in  $t_{\text{fb}}$ , this gas particle spawns a star particle. Soon after it enters in a hot state  $T \gtrsim 10^7$  K, where it remains until the present time. In the hot state, sputtering becomes a strong source of dust destruction. Note again the steeper decay of small grains caused by the size dependence of sputtering. The fluctuations in dust abundances after 8 Gyr are due to slight enrichments from the surrounding star particles. While these contributions are negligible in the relatively high gas metals mass fraction, they are evident in that of large grains, and even more so in that of small grains.

### 3.2 Evolution of the global properties of dust

#### 3.2.1 Dust distribution in the cluster region

In this section, we analyse the global evolution of dust distribution around the central cluster. Fig. 4 follows from top to bottom the

redshifts  $z = 4, 3, 2, 1, 0$ , and from left to right the column densities of total gas mass, gas-phase metals, large dust and small dust grains. At each redshift, we project a physical cube of size 1 Mpc around the main cluster’s progenitor. Regions in which the column density is lower than  $10^{-8} \sigma_{\text{peak}}$ , where  $\sigma_{\text{peak}}$  is the peak column density, are omitted.

The maps shows that gas-phase metals approximately mirror the total gas distribution at all redshifts. Large grains follow loosely a similar pattern until  $z \sim 2$ . At lower  $z$ , this correlation breaks down since many particles reach  $T \gtrsim T_{\text{sp},0} = 2 \times 10^6$  K, above which sputtering erodes grains effectively, particularly small ones. Moreover, small grains are less abundant than large ones at all times and do not grow as rapidly, except in cold ( $T < T_{\text{MPh}} = 5 \times 10^5$  K) star-forming gas overdensities, where shattering and accretion work efficiently.

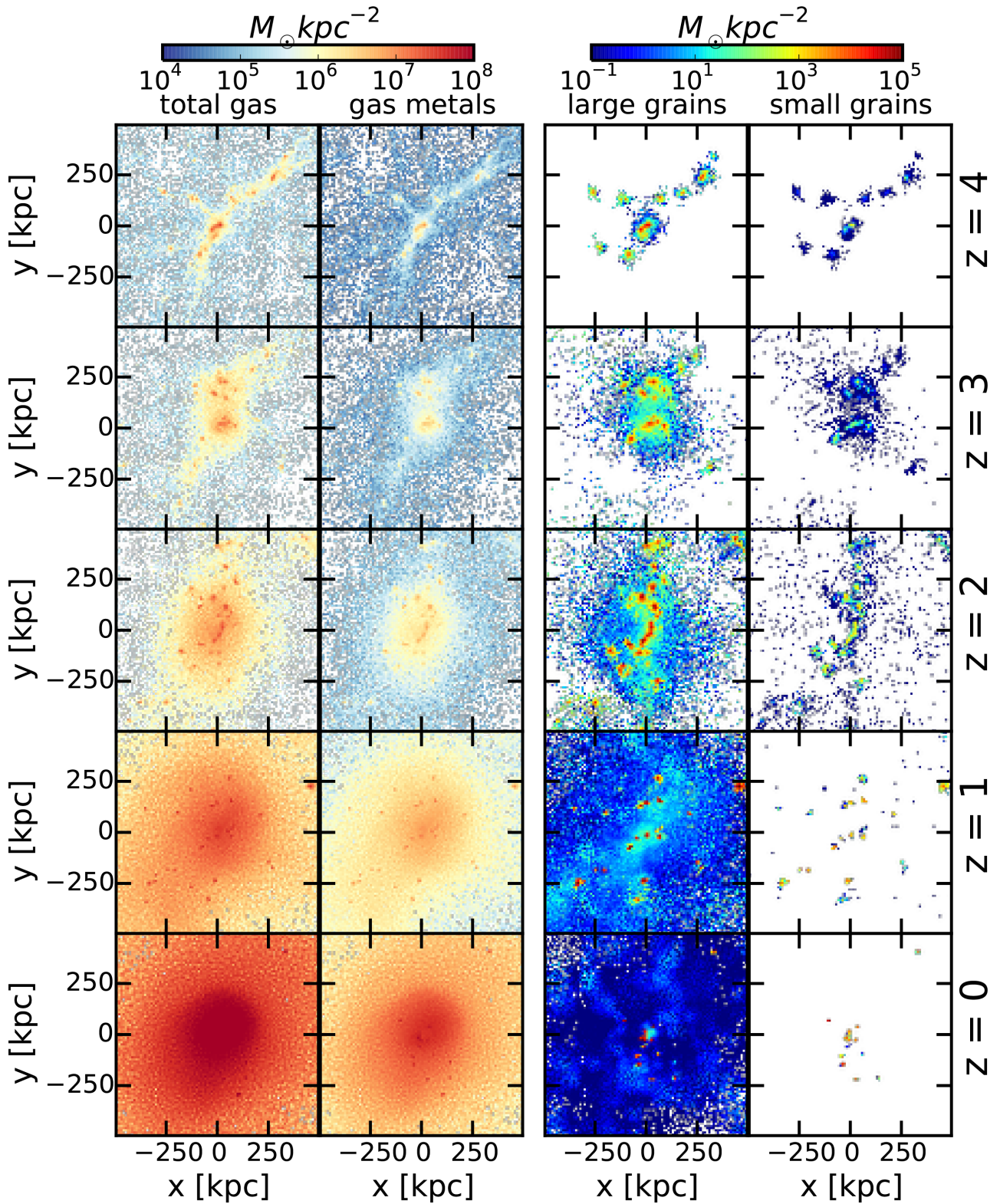
Fig. 5 refers to the same 1 Mpc physical cube around the main progenitor. It represents from left to right the star particle column density, the star formation rate (SFR) column density, the mass-weighted mean temperature and mean number density of gas particles. These maps help in interpreting the former ones. For instance, it can be appreciated that the structures wherein dust survives and evolves undisturbed by sputtering are characterized by low temperature. They feature SFR between 0.01 and 0.1  $M_{\odot} \text{yr}^{-1} \text{kpc}^{-2}$ . As the cluster gains mass its average temperature rises, and sputtering begins to destroy efficiently dust in the ICM.

Fig. 6 maps the average dust-to-gas-metal ratio (left), small-to-large grain ratio (centre) and silicates versus carbonaceous dust (right). At  $z > 3$ , when the SF activity in the proto-cluster region is at its maximum and dust reprocessing is expected to be important, dust properties are predicted to differ significantly from those derived for the MW dust. The latter are almost always adopted in computations to account for dust reprocessing (e.g. Domínguez-Tenreiro et al. 2014, and references therein). For instance, compared to the dust models proposed for the MW by Weingartner & Draine (2001), in the central 100 kpc at  $z = 4$  the mass ratio of small/large carbon grains is more than 2 dex smaller, while the silicate/carbon mass ratio is about twice smaller.<sup>4</sup> We illustrate the possible consequences of these differences on the predicted SEDs in Fig. 7. Here, we show two SEDs computed with the public radiative transport code SKIRT<sup>5</sup> (Camps & Baes 2015). In one case, we have simply used a dust mixture similar to that adopted so far in most GRASIL (Silva et al. 1998) and all GRASIL3D (Domínguez-Tenreiro et al. 2014) applications.<sup>6</sup> Here, the only information on dust derived from the simulation is the total dust content of each SPH particle. In another computation, we have instead exploited the full information concerning the relative partition in the four categories of dust grains (graphite and silicate, small and large), by locally adjusting the adopted mixture at the position of each SPH particle, in order to

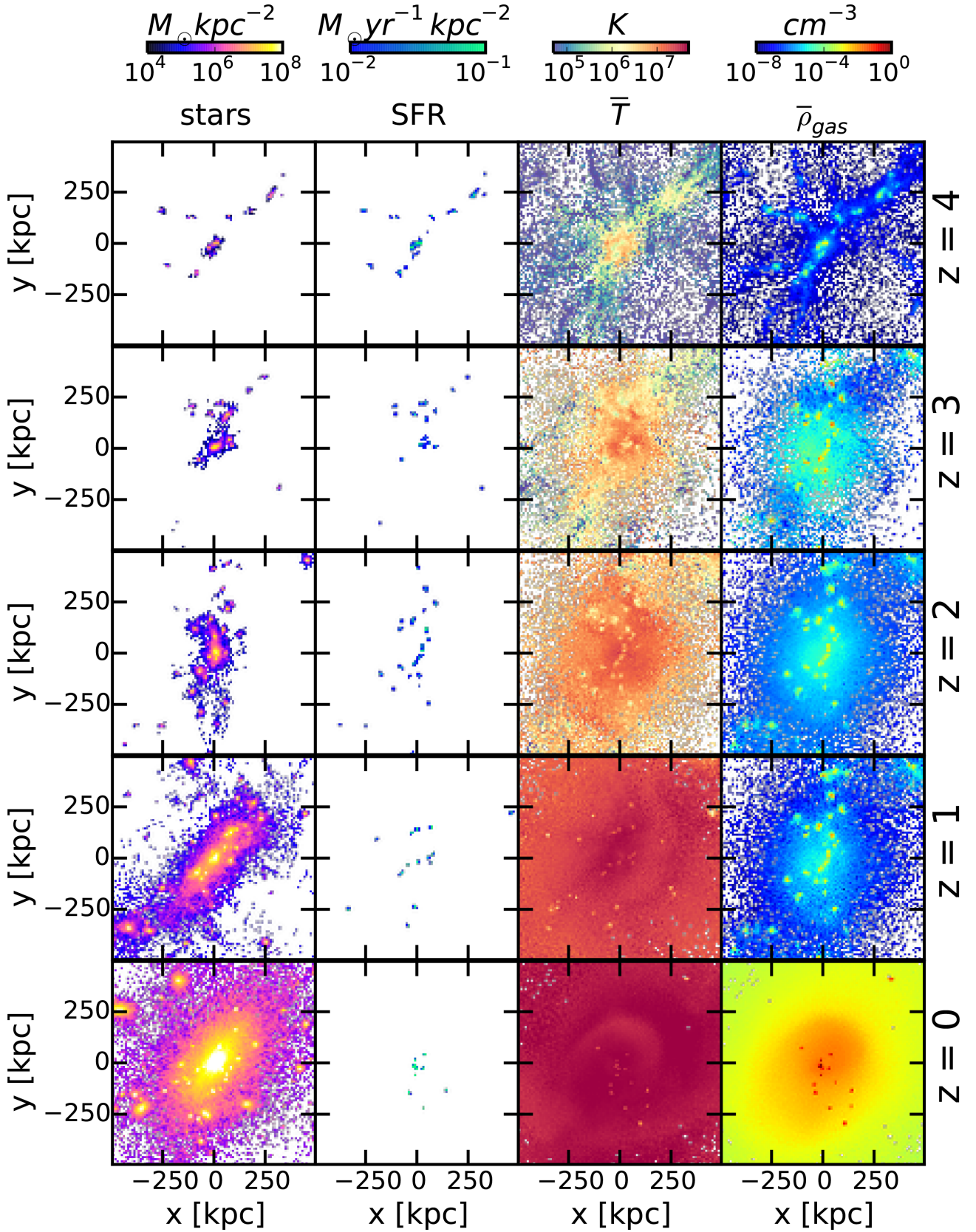
<sup>4</sup>For the most commonly adopted model 4 by Weingartner & Draine (2001), the mass ratio of small-to-large C grains is 0.34, while the mass ratio of Silicate to Carbon grains is 2.5. The model proposed by Silva et al. (1998) for the MW dust, usually adopted in GRASIL code features values not significantly different.

<sup>5</sup><http://www.skirt.ugent.be>

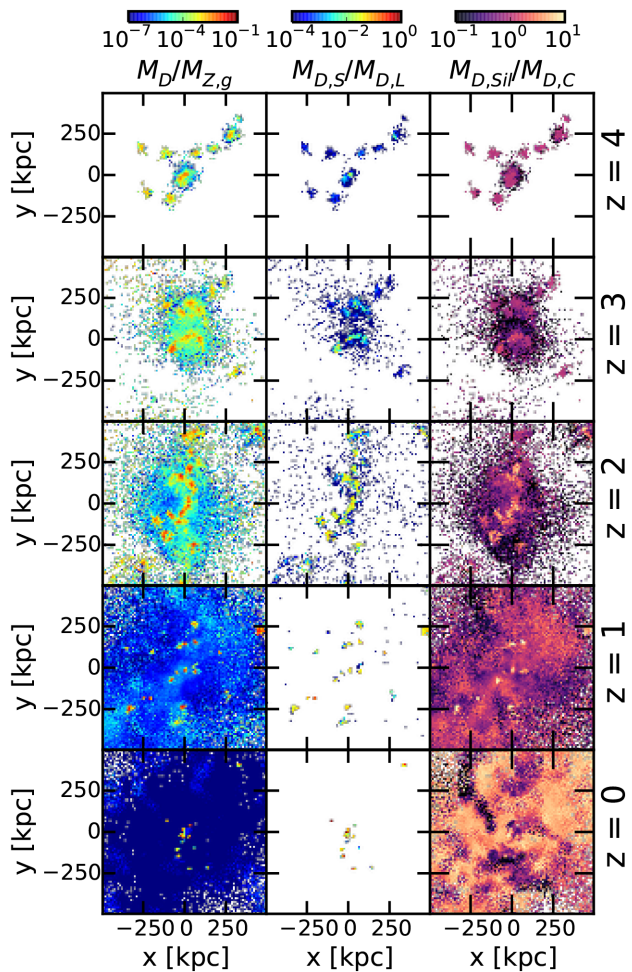
<sup>6</sup>These codes are often employed to compute dust reprocessing in semi-analytic models and galaxy formation simulations, respectively. They both allow for custom variations of the mixture, albeit this feature is seldom used due to lack of information. A spacial dependence on these properties is not implemented in the present versions.



**Figure 4.** For the fid run, column density maps for total gas mass, gas-phase metals, large dust grains, and small dust grains, in a box of 1 Mpc in physical size over five redshifts (from top to bottom,  $z = 4, 3, 2, 1$ , and 0). The two colourbars are fixed at all redshifts from total gas and gas metals and for large and small grains, respectively. Dust abundances trace gas mass distributions until about  $z = 2$ . After  $z = 2$  sputtering destroys dust. Small grains evolve to the point of reaching large grains abundances only in cold overdensities, but they are destroyed more efficiently than large grains in hot gas particles.



**Figure 5.** Similarly to Fig. 4 for the `fid` run, columns represent from left to right: maps of stellar mass and SFR column densities, and of mean mass-weighted temperature and mean number density in a box of 1 Mpc in physical size over 5 redshifts ( $z = 4, 3, 2, 1, 0$ ). Both means are weighted with the gas particle masses. Star formation occurs in cold overdensities mostly occupied by multiphase gas particles.

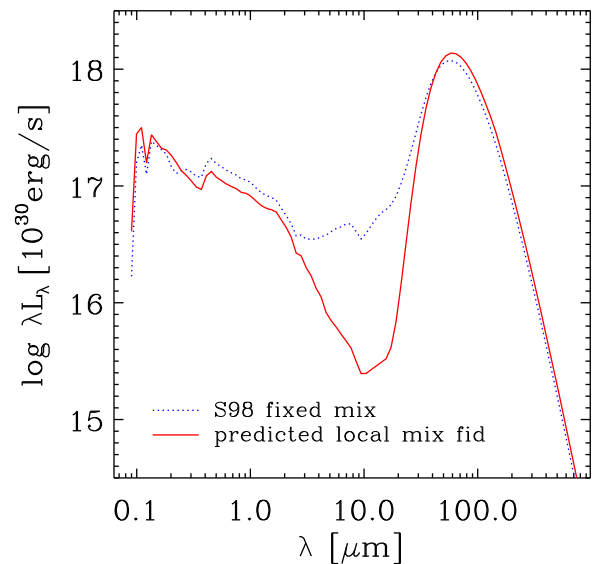


**Figure 6.** Similarly to Fig. 4, columns represent maps of the dust-to-gas-metal ratio (left), small-to-large grain ratio (centre) and silicates versus carbonaceous dust (right) in a box of 1 Mpc in physical size over five redshifts ( $z = 4, 3, 2, 1,$  and  $0$ ).

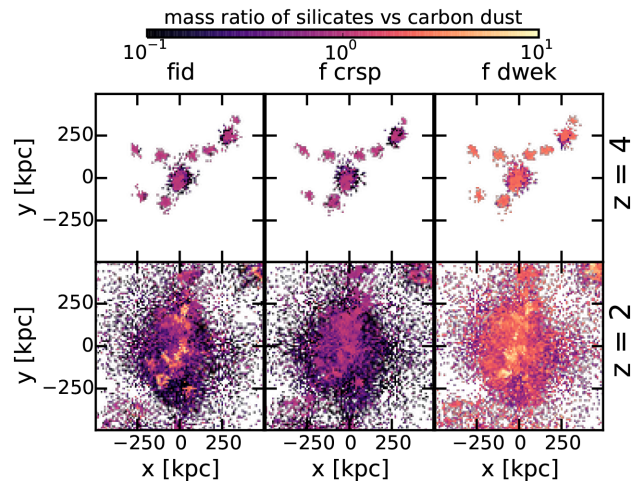
reflect this partition.<sup>7</sup> As it can be seen, the differences are important, particularly in the optical to mid-IR regime. We plan to explore in detail the observational consequences of these differences in the near future.

The ratio silicate-to-carbon dust takes some time to reach values close to those adopted in ‘standard’ dust mixtures. Actually, the ratio arising from dust production by stars is significantly lower in the model, and its increase occurs via dust evolution in the ISM. The behaviour can be appreciated from the maps shown in Fig. 8. From left to right, the three columns compare maps of the ratio for the fiducial run (*fid*), the run including only dust production from stars and sputtering (*f-crsp*), and the run in which the silicate dust is produced by stars adopting the prescription by Dwek (1998) (*f-dw*, see Section 2.4.1), respectively. At early time  $z = 4$ ,

<sup>7</sup>We fed *SKIRT* with a superposition of four spatial distribution of dust densities, each one to represent one of the four grain types followed by our model. The size distributions for each of them, provided by suitable input parameters of *SKIRT*, were the same as that proposed by Silva et al. (1998), but having  $0.03 \mu\text{m}$  as the limiting size to distinguish between small and large grains, respectively. Moreover, the normalization of each of the four distributions has been computed at each point according to the local density of the corresponding grain type, as predicted by the simulation.



**Figure 7.** Spectral energy distributions predicted by *SKIRT* (Camps & Baes 2015) for the 100 kpc central box of the region D2 at  $z = 4$ . The dotted blue line has been computed adopting for the whole volume a ‘standard’ dust mixture reproducing the properties of dust in the diffuse ISM of the MW. The solid red line instead takes into account the point-to-point variations predicted by the simulation for the relative abundances of small and large grains, as well as those for silicate and graphite grains. See Section 3.2.1 for more details.



**Figure 8.** Similarly to Fig. 6 right column, The three columns compare maps of the ratio for the fiducial run (*fid*), the run including only dust production from stars and sputtering (*f-crsp*), and the run in which the silicate dust is produced by stars adopting the prescription by Dwek (1998).

the first two runs are virtually indistinguishable. It is also worth noticing that the last one is characterized by substantially higher silicate-to-carbon (Sil-to-C) ratios, already closer to, or even higher than, the standard one. Indeed the prescription put forward by Dwek (1998) is substantially more liberal in using the ejecta to produce silicate dust. At lower redshift  $z = 2$ , when evolution in the ISM has been important (if included in the computation), the fiducial run has increased the ratio in most of the region. This is not the case for the production and sputtering only run (*f-crsp*, lacking all the processes causing dust evolution in the cold ISM). Also the *f-dw* run does not show a sizable increment with respect to  $z = 4$ , but in that

case it was already high. This is because it leaves by construction less silicate elements in the gas phase at stellar production, meaning that gas accretion onto grains has less to add. The time evolution of the sil-to-C ratio, integrated over  $r_{200}$  is also shown in a later figure.

### 3.2.2 Temperature dependence of dust contents

Additional insights on how and when the various processes affect the dust content of the SPH gas particles can be obtained from inspection of Fig. 9. Here, we plot 2D histograms for DtG (y-axis) versus temperature (x-axis) over the usual five redshifts. DtGs are separately shown for large and small grains. The figure on the right refers to the entire simulation. That on the left includes instead only particles that have spawned at least one stellar particle before. This is meant to select preferentially gas particles which have spent a significant fraction of their life in a star-forming environment.

At  $z \geq 3$  large grains show a peak in the two dimensional distributions at  $T \lesssim 10^4$  and  $\text{DtG} \sim 10^{-4}$ . This last value is about 2 orders of magnitude below the standard Galactic DtG of  $\simeq 0.01$ , marked by horizontal magenta lines in the panels. Indeed, at these early epochs the evolutionary processes in the ISM have not had sufficient time to affect much the dust content of most SPH particles. We have verified this by comparing with the run `f-crsp`. The major effect of ISM evolution before  $z \simeq 3$  is the production of a certain amount of small grains by sputtering. ISM evolution effects manifest appreciably at  $z = 2$ . A population of gas particles featuring DtG close to  $10^{-2}$  for large grains and  $10^{-3}$  for small ones shows up at  $T \lesssim 10^5$ . The particles belonging to this population have undergone multiphase periods, during which accretion onto small grains and their coagulation to form large ones have raised the DtG up to values similar to the Galactic one. Indeed these local peaks appear very similar in the right and in the left figures. As pointed out, the latter is meant to select gas particles characterized by multiphase star-forming periods in their past. Moreover, in run `f-crsp`, gas particles featuring  $\text{DtG} \geq 10^{-2}$  are not produced, as it can be seen in Fig. 10. In this case, the peak of the DtG distribution remains 1–2 orders of magnitude lower even at  $z \leq 2$ , only becoming more and more populated.

At still lower redshift, the high temperatures reached by most SPH gas particles promote efficient thermal sputtering. As a result, a well-defined maximum in the 2D histograms of the entire simulation plots develops at  $T \gtrsim 10^7$  and at very low  $\text{DtG} \lesssim 10^{-5}$  for large grains. The peak for small grains occurs at an even smaller DtG, since they are more strongly disrupted by sputtering.

### 3.2.3 Evolution history of run variations

Fig. 12 illustrates the evolutionary history of various masses computed within  $R_{200}$ , in the main progenitor of the  $z = 0$  cluster. We include results for a selection of run variations on the D2 region. The four top plots show the gas metal and dust masses normalized to the evolving  $M_{200}$ , while the six plots in the bottom show various interesting ratios. In order to facilitate the interpretation, the evolution of  $M_{200}$  and  $R_{200}$  can be seen in Fig. 11.

From the total dust panel, we can appreciate that, excluding the run without sputtering `f-nosp`, the dust production is faster than or comparable to the increase of the cluster mass only in the  $t_{\text{lb}}$  lapse between 11 and 9 Gyr and at  $t_{\text{lb}} \lesssim 4$  Gyr. In the former interval, this is achieved thanks to the combined effect of shattering, accretion,

and coagulation occurring in the the gas, which enhance the dust content by a factor  $\sim 2$ –3 after  $t_{\text{lb}} \sim 10$  Gyr, as can be understood by comparing the runs `fid` with `f-crsp`. In the latter run, the three above mentioned processes are switched off, and the ratio  $M_{\text{dust}}/M_{200}$  is monotonically decreasing down to low  $z$ . The late  $t_{\text{lb}} \lesssim 4$  Gyr flattening of the evolution is instead related to slow down of the  $M_{200}$  increase and to the fact that sputtering has already destroyed most of the dust in the hot ICM.

By adopting the (Dwek 1998) recipe to compute the production rate of silicate grains from stars (run `f-dw`), more liberal than our fiducial method imposing the chemical composition of olivine, we get about 50 per cent more dust at early time  $t_{\text{lb}} \gtrsim 9$  Gyr, and a higher ratio  $M_{\text{sil}}/M_{\text{C}}$  over the whole evolution.

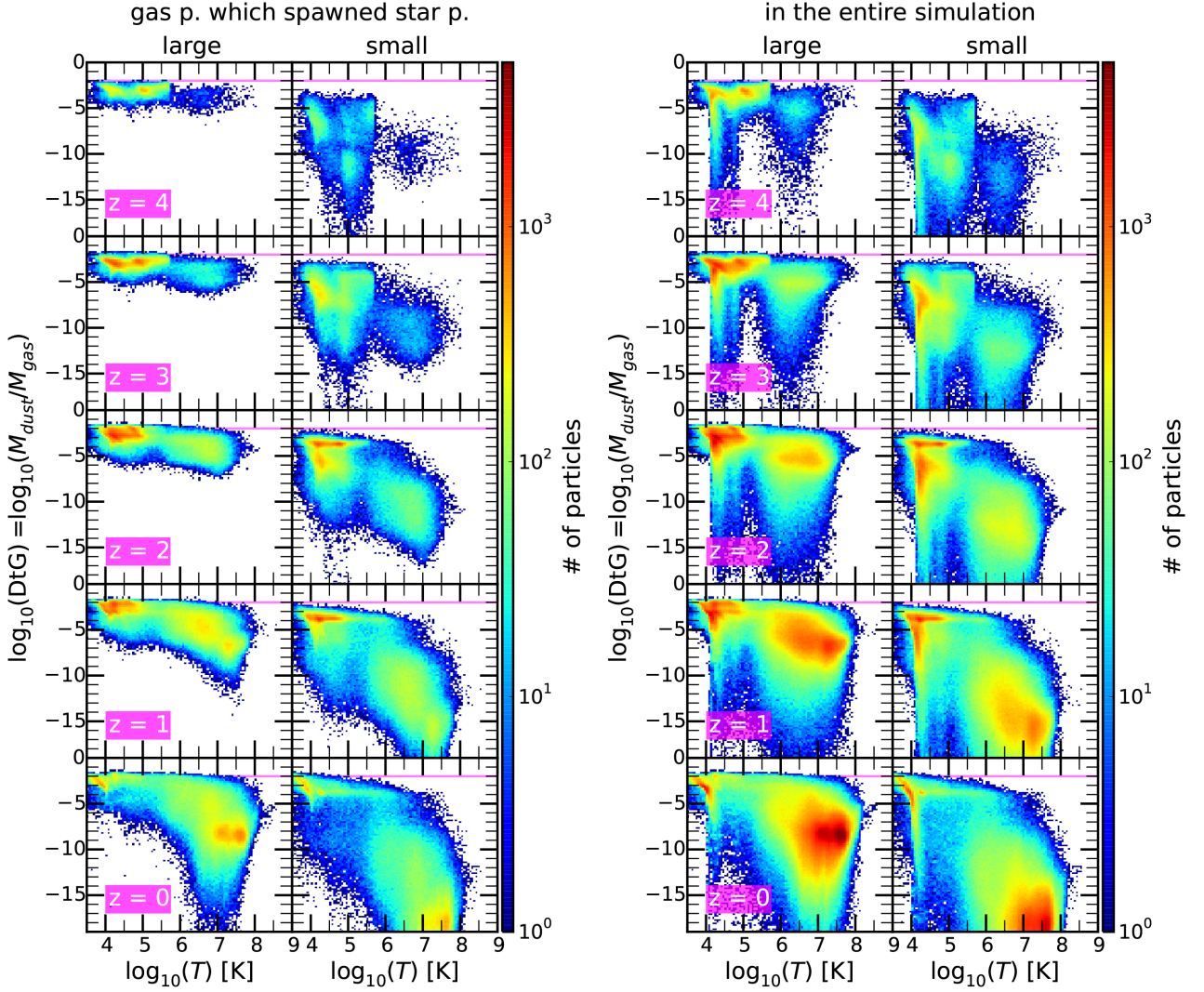
It is interesting to note that by considering only the SNII channel for the production of dust (run `f-snii`), the total dust content is somewhat underpredicted with respect to the fiducial run by up to  $\sim 50$  per cent. But what is more important is that the ratio  $M_{\text{sil}}/M_{\text{C}}$  is significantly overpredicted. Thus, this approximation, sometimes adopted in other works (e.g. Hou et al. 2016; Aoyama et al. 2017; Hou et al. 2017; Chen et al. 2018), seems to be insufficient at least for certain purposes, such as computing the radiative effects of dust.

As we already pointed out, in the fiducial run the mass fraction of small grains at high redshift, when SF is most active in the simulated region is much lower than that of standard mixtures ( $\sim 0.2$ ). These mixtures are calibrated on the average properties of MW dust. This result is quite robust, in the sense that occurs in all the run variations we considered. On the other hand, the ratio  $M_{\text{sil}}/M_{\text{C}}$  is not very different from the standard value  $\sim 2$  at early time, while it ends up significantly higher in most runs. The only exception are those runs with no or simply reduced sputtering (runs `f-nosp` and `f-sp.2`), in which the ratio increases less. The latter run could be also more consistent than the fiducial one with recent estimates of dust content in clusters at low redshift (see Section 3.3). On average, the total dust surviving in `f-sp.2` is a factor of 3 greater than in `fid`. This dust excess is almost entirely due to the increased survival of large grains in the hot ICM, while small grains do not deviate significantly in the two cases. This could seem at first sight unexpected, since sputtering affects more promptly small grains than large ones (equation 18). What happens is that increasing its time-scale by a factor of 5 impacts on the survival of large grains, but the process remains still sufficiently effective to obliterate almost entirely small ones in the hot ICM. Consequently, in the hot environment we predict that the small-to-large grain ratio is affected considerably. In both `fid` and `f-sp.2`, the small grain content shown in Fig. 12 comes from regions characterized by gas temperature  $\lesssim 10^6$  K.

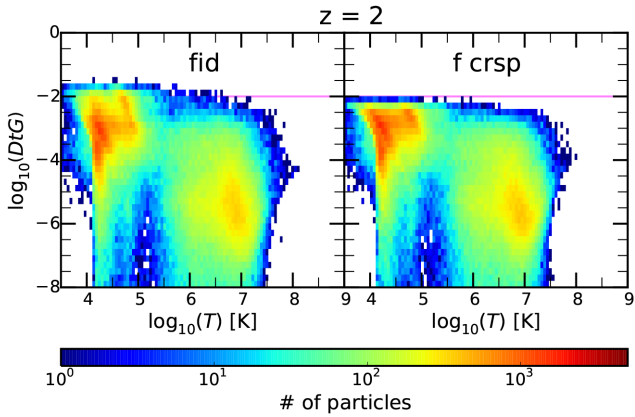
### 3.2.4 DtG and metallicity

Fig. 13 presents 2D histograms of the DtG versus gas metallicity  $Z_{\text{gas}}$  at  $z = 4, 2$ , and 0 for gas particles which are either multiphase in the entire simulated region (left) or within  $R_{200}$  of the main cluster (right). In both cases, we used the `fid` run.

Looking at the multiphase star-forming particles, we notice that at an early time of  $z = 4$ , the gas particles cluster along a linear relationship between DtG of large grains and  $Z$ . Large grains dominate the dust content so this also represents the total behaviour. In this initial phase, the often adopted approximation of a constant ratio between dust and gas metal content is qualitatively good enough over the entire metallicity range of star-forming particles. Note that



**Figure 9.** 2D Histograms of the total DtG versus temperature for gas particles which have spawned star particles in their past (left) and in the entire simulation (right) for the fiducial run and over five redshifts ( $z = 4, 3, 2, 1,$  and  $0$ ). The magenta line represents  $\text{DtG} = 10^{-2}$ , close to the commonly accepted ISM dust abundance in the MW.

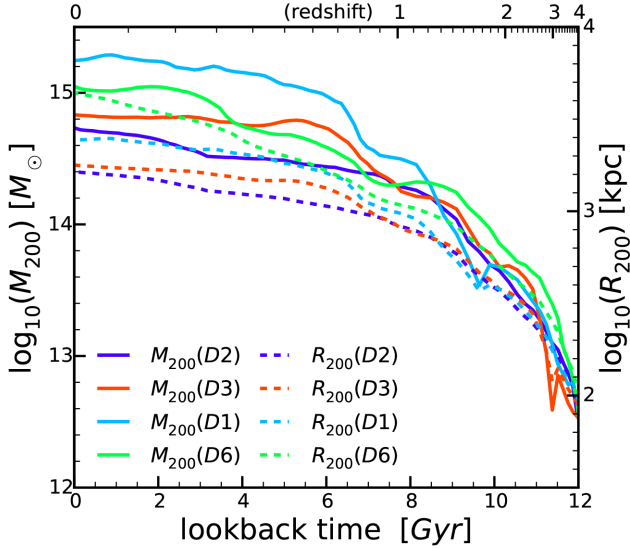


**Figure 10.** 2D histogram of DtG versus temperature within the *fid* and *f-crsp* at  $z = 2$  including all gas particles.

the normalization of this linear relationship between DtG and  $Z$  is lower by a factor  $\sim 3$  than that generally adopted (i.e.  $\text{DtG} \simeq 0.01$  for  $Z \simeq Z_{\odot}$ ). The linearity arises because initially the dominating

mechanism responsible for the presence of dust in the system is production by stars, for which the basic assumption is that a certain fraction of produced metals goes to the solid state rather than gaseous form. As time goes on, the ISM evolution processes boost the dust content for a given metallicity. Accretion of gas metals onto small grains provides sufficient mass to enhance the coagulation of small grains onto large ones. Hence, coagulation dominates over shattering in dense regions. This occurs only above a critical  $\log Z \simeq -2, 5$  for the fiducial parameters, producing a more than linear increase of DtG with  $Z$ . However, at still 2–3 times higher  $Z$ , the DtG slope slows down again toward a linear relationship, now featuring a normalization close to that given by standard MW values.

As noticed by Aoyama et al. (2017, see their figure 7), the overall result of this evolution is that at late time  $z \leq 2$  multiphase particles tend to concentrate in a region whose shape resembles the line along which the one-zone model by Hirashita (2015) evolves in time, shown in the figure with a black line. However, when evaluating this result one should keep in mind the very different nature of one-zone computations and simulations. In the latter, the gas and star density fields are sampled with a certain resolution. In the one-zone



**Figure 11.**  $M_{200}$  and  $R_{200}$  of the fiducial run for the D2, D3, D1, and D6 regions.

model, both quantities reported in the 2D histograms of Fig. 13 are a function of time. In other words, the system moves with time along the line. On the contrary, in the simulation at any given time the gas particles have a broad distribution in the plane, albeit with well defined concentrations.

The right-hand panels of Fig. 13 shows that when the 2D distributions are computed for the main cluster of the region and including also non multiphase particles, a population of particles characterized by very low values of DtG shows up. This population becomes more and more dominant over time, because it is produced by sputtering which efficiently destroy dust grains in the hot ICM at  $T \gtrsim 10^6$  K.

### 3.3 Observational consistency

#### 3.3.1 Dust abundance versus metallicity

In Fig. 14, we show the 2D distribution of DtG versus (gas) Oxygen abundances of star-forming multiphase particles at  $z = 0$ , for the four regions on which we run our fiducial model. This information is similar to that already reported in Fig. 13, but now for all the regions considered in this work, and in a form more directly comparable with observations. We include in the figure data on nearby galaxy samples from Kennicutt et al. (2011) and Madden et al. (2012). It is apparent the broad similarity between the distribution of simulated and the data points. Note, however, that the former refers to individual particles residing in star-forming regions rather than entire unresolved galaxies, which could explain the larger dispersion. In any case, it is reassuring that the model reproduces reasonably well the observed trend of dust abundance with metallicity.

#### 3.3.2 Dust content at low redshift

In this section, we briefly compare with the still limited reliable detections of global dust content in galaxy clusters. Gutiérrez & López-Corredoira (2017) analysed 327 clusters of galaxies in the redshift range 0.06–0.70, using maps and catalogues from the Herschel MerMES project. They reported average integrated fluxes of 118.2, 82.3, and 38.0 mJy within 5 arcmin of the cluster centres at

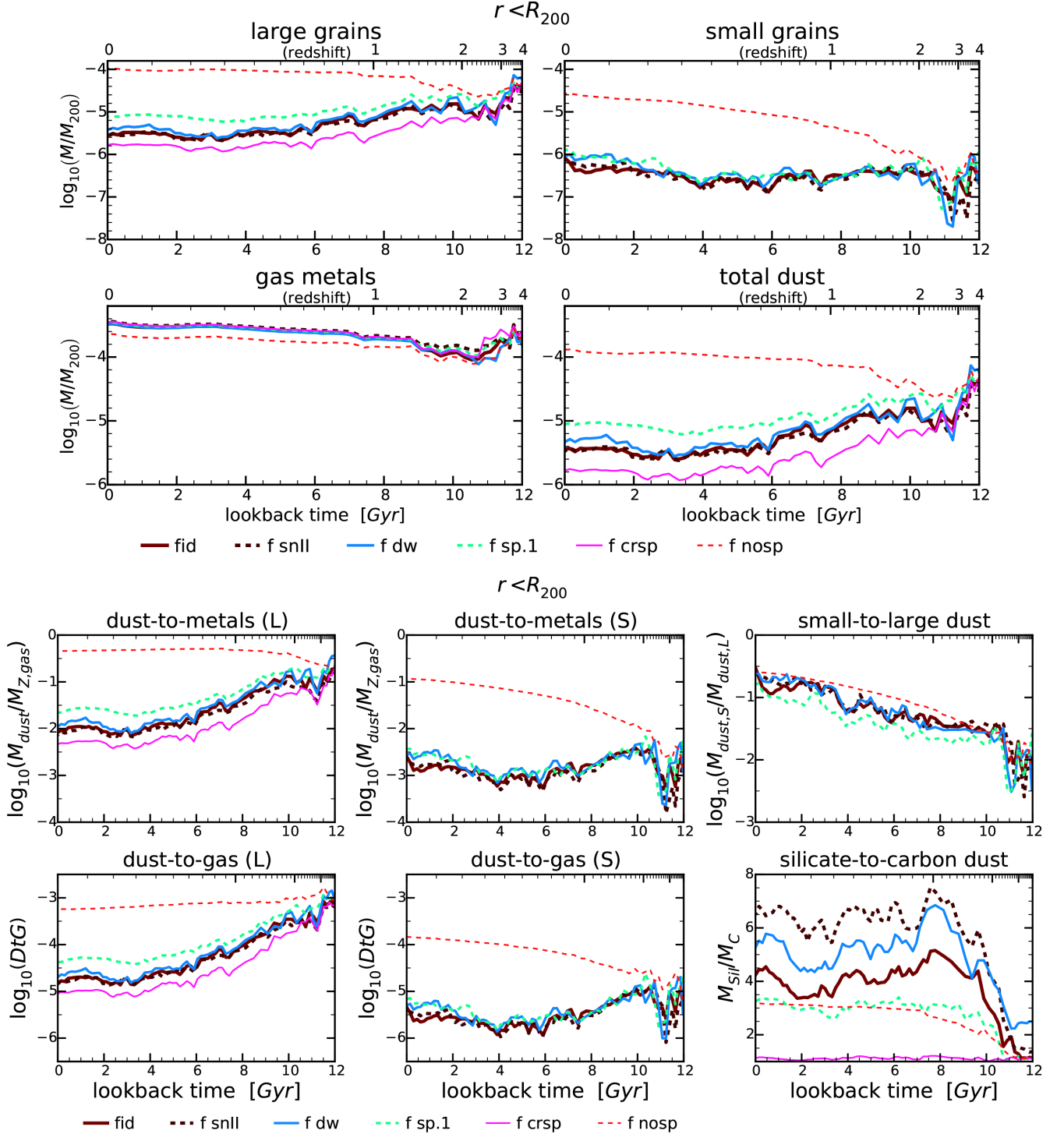
250, 350, and 500  $\mu\text{m}$ , respectively. Adopting their same assumptions on the dust optical properties and temperature (essentially the same as the ‘average’ MW dust), these fluxes translate to a total dust mass of  $1.7 \times 10^9 M_{\odot}$ . Given that their average cluster mass is  $1.1 \times 10^{14} M_{\odot}$  within the same radius, the fraction of dust mass turns out to be about  $1.5 \times 10^{-5}$ . For the main clusters in our four regions, we predict an average fraction in the same redshift range of  $0.25 \times 10^{-5}$ . Thus, we are underpredicting the dust content by a factor  $\sim 6$ . On one hand, this discrepancy could be due at least in part to the quite strong and uncertain assumptions entering into the masses estimated from observations, such as the dust optical properties and (constant) temperature. In particular, they adopted a long wavelength emissivity that scales as  $\lambda^{-\beta}$ , with  $\beta = 2$ , which is the standard result of first principle computations of dust opacities by Draine & Lee (1984). There are several indications that dust emission is better represented at long  $\lambda$  by a shallower index  $\beta \sim 1.5$ , when it is described by single  $T$  blackbody (e.g. Planck Collaboration XXII 2015). From a physical point of view, this could be actually the result of a spread in dust temperature. In any case, adopting a smaller  $\beta$  in interpreting observed fluxes would lead to smaller dust masses in better agreement with our result. On the other hand, the dust content at low redshift in our model clusters is strongly dictated by the sputtering efficiency. For instance, assuming a time-scale longer than that favoured by existing literature (e.g. Tsai & Mathews 1995, and references therein) by a factor 5 (equation 18), the final dust content increases by a factor  $\sim 3$  (See Fig. 12).

Planck Collaboration XLIII (2016) performed a stacking analysis of several hundreds of clusters, wherein *IRAS* and *Planck* data are combined to provide a well-sampled FIR-submm average SED of clusters. Then they fit this SED with modified black body models to derive dust masses and temperature simultaneously. For the whole sample, whose mean redshift is 0.26 and mean total mass is  $M_{200} = 5.6 \times 10^{14} M_{\odot}$ , the estimated dust mass is  $1.1 \times 10^{10} M_{\odot}$ , adopting their preferred emissivity index  $\beta = 1.5$ . Strictly speaking their dust mass refers to the radius of 15 arcmin within which the fluxes have been integrated.

In Fig. 15, the time evolution of the ratio  $M_{\text{dust}}(r < 15 \text{ arcmin})/M_{00}$  is shown individually for the four clusters simulated with the fiducial set of parameters. This is compared with the same ratio as derived from these *Planck* data. The comparison is shown not only with the full sample, but also for the two subsample at  $z \leq 0.25$  and  $z > 0.25$ , comprising 307 and 254 clusters, respectively. This figure confirms that our fiducial models underpredict the dust content of the clusters, albeit by a smaller factor  $\sim 3$ . As such, the model with increased sputtering time-scale turns out to be in reasonable agreement with the data.

## 4 SUMMARY AND FUTURE PROSPECTS

In this work, we have introduced a state of the art treatment of dust production and evolution in our version of the simulation code GADGET-3. We take advantage of the detailed description of chemical evolution already included in our version of GADGET-3 to trace separately the two dust species that are believed to populate the ISM, namely carbonaceous and silicate grains. We also trace at a basic level the continuum size distribution of grains by means of the two-grain size approximation introduced and tested by Hirashita (2015). In our code, large (nominally  $0.1 \mu\text{m}$ ) dust grains are originated by simulation stellar particles from three stellar channels, AGB winds, core collapse SNe, and SNIa. These grains are spread along with metals to the surrounding SPH gas particles, where we

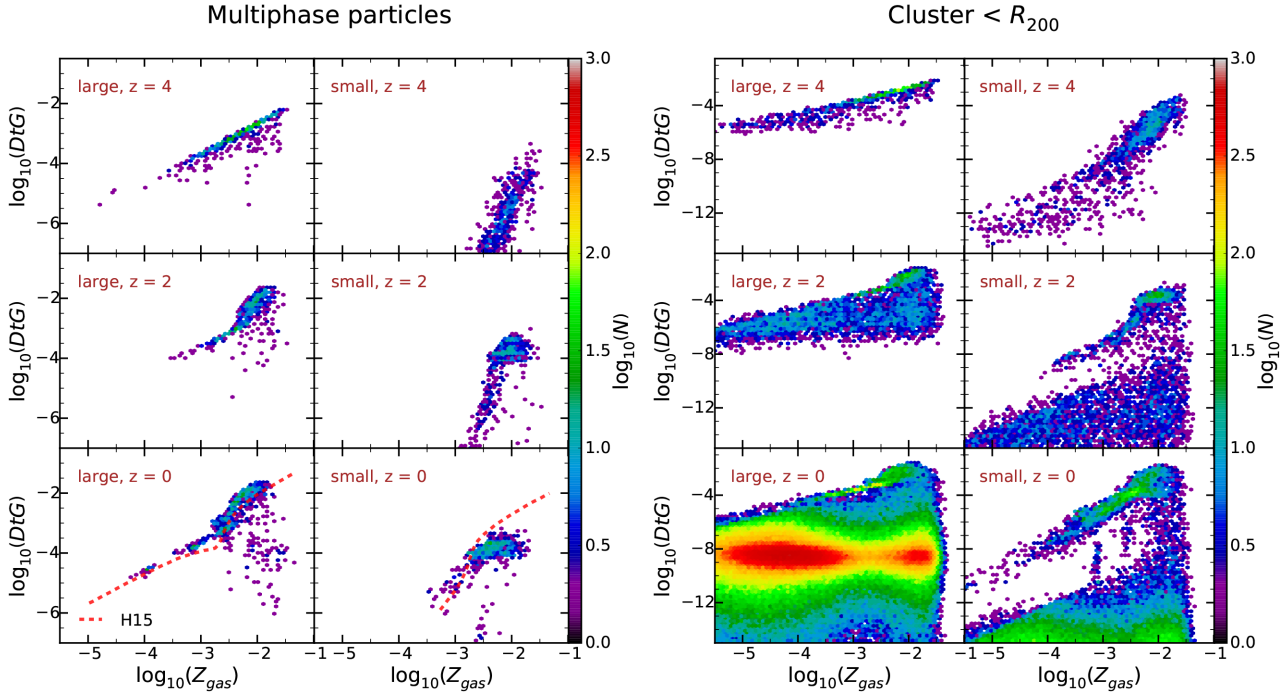


**Figure 12.** Time evolution within  $R_{200}$  of the main progenitor for a selection of runs. (Top 4 plots) masses of large grains, small grains, gas metals, total dust mass and (bottom 6 plots) large-dust-to-gas-metal, small-dust-to-gas-metal, small-to-large ratios, DtG for large and small grains, and lastly the mass ratio between carbonaceous dust and silicates. Small grains are affected more strongly than large grains by the time-scales of the evolution processes.

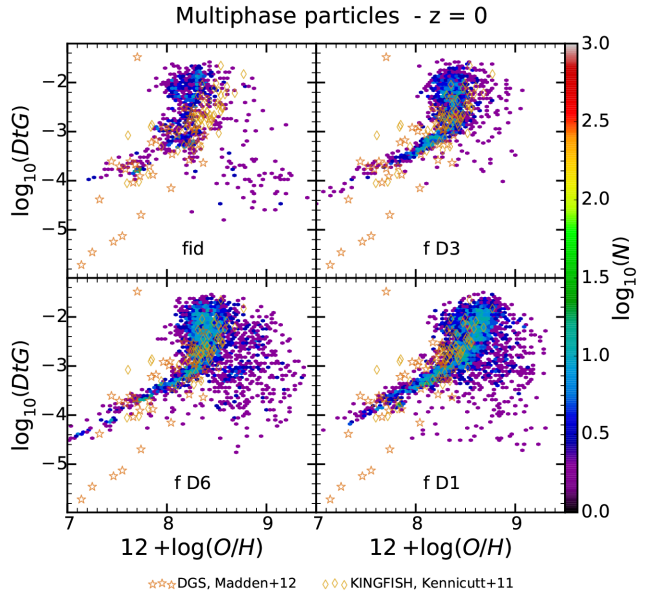
allow for various ISM evolution processes affecting dust properties to occur. In brief, large grains are shattered onto small (nominally  $0.01 \mu\text{m}$ ) grains if the gas density is low enough. On the other hand, in dense star-forming SPH gas particle metals accrete onto small grains and small grains coagulate onto large grains. We also take into account dust destruction by SN shocks and by sputtering in the hot ( $T \gtrsim 10^6$  K) ICM. We evaluate time-scales for the above

mentioned ISM processes as a function of the physical conditions of the SPH particles, and within each we evolve the dust and gas metal contents. As remarked above, the former is divided into four types, namely large carbon, small carbon, large silicates, and small silicates.

As a first test, we apply the method to cosmological zoom-in simulations of four massive ( $M_{200} \geq 3 \times 10^{14} M_{\odot}$ ) galaxy clusters.

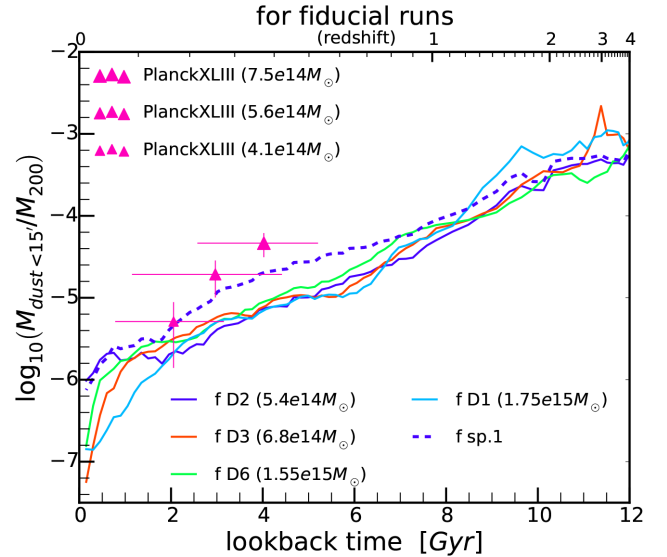


**Figure 13.** 2D histogram of the DtG versus gas metallicity for multiphase star-forming gas particles (left) and for all gas particles (right) in the standard run fid within  $R_{200}$ . The black lines represent the relationship predicted by the one-zone model by Hirashita (2015).



**Figure 14.** 2D histogram of the DtG versus oxygen abundance at  $z = 0$  compared with galactic data from Kennicutt et al. (2011) (red diamonds) and Madden et al. (2012) (black stars).

During the early stages of assembly of the cluster at  $z \gtrsim 3$ , where the star formation activity is at its maximum in our simulations, the proto-cluster regions are rich of dusty gas (Fig. 4). Compared to runs in which only dust production in stellar ejecta is active, runs including processes occurring in the cold ISM enhance the dust content by a factor 2–3. However, the dust properties in this stage turn out to be significantly different than those observationally derived for the *average* Milky Way dust, and commonly adopted in calculations of dust reprocessing (Fig. 8). We stress that this results



**Figure 15.** For the the four different zoom-in cluster simulations considered in this work [fid (D2), f-D3, f-D1, f-D6], we show the evolution of the ratio between the dust mass within a 15 arcmin aperture and  $M_{200}$ . For the region D2, we show also the run with reduced sputtering, f-sp.2. These are compared with the same quantity as estimated in Planck Collaboration XLIII (2016), for the whole sample (central point) and for the subsamples at  $z < 0.25$  and  $z > 0.25$ . The error-bars represent dispersions.

is not unexpected. We show that these differences may have a strong impact on the predicted spectral energy distributions (Fig. 7). At low redshift, our model reproduces reasonably well the trend of dust abundances over metallicity as observed in local galaxies (Fig. 14). However, we underproduce by a factor of 2–3 the total dust content of clusters estimated observationally at low  $z \lesssim 0.5$  using

IRAS, Planck, and Herschel satellites data. This discrepancy can be solved by decreasing the efficiency of sputtering which erodes dust grains in the hot ICM (Fig. 15).

The most immediate purpose of our effort is to have enough information on the simulated ISM/ICM to compute observational properties of simulated objects by means of radiative transfer post processing. This is in line with what we already achieved in Granato et al. (2015), but with less assumptions required on dust properties. However, this work can be also regarded as a first step in the direction of a more sophisticated prescriptions for the sub-resolution physics. For instance, it will be possible to estimate the contribution of dust to the formation of molecular gas. This estimate is a primary ingredient in more advanced implementations of star formation in simulations (e.g. MUPPI, Murante et al. 2015).

## ACKNOWLEDGEMENTS

We are indebted to the anonymous referee for the very careful reading of our manuscript and several important suggestions that improved its quality. We thank Volker Springel for providing the non-public version of the GADGET-3 code, and Peter Camps for prompt assistance and advice in using the radiative transfer code SKIRT. EG is supported by the INAF 2015 PhD grant ‘Galaxy clusters in their infancy: confronting simulated and observed IR/sub-mm properties’. This project has received funding from the European Union’s Horizon 2020 Research and Innovation Programme under the Marie Skłodowska-Curie grant agreement no. 73437. The simulations were carried out at the following facilities (i) Mendieta Cluster from CCAD-UNC, which is part of SNCAD-MinCyT, Argentina; (ii) European Exascale System Interconnect and Storage; and (iii) CINECA (Italy), with CPU time assigned through IS CRA proposals and an agreement with the University of Trieste. The post-processing and analysis has been performed using the PICO HPC cluster at CINECA through our expression of interest. This publication has received funding from the European Union’s Horizon 2020 research and innovation program under grant agreement no. 730562 [RadioNet].

The analysis was conducted using IPYTHON2 (Pérez & Granger 2007), SCIPY3 (Jones et al. 2001), NUMPY4 (Van Der Walt, Colbert & Varoquaux 2011), and MATPLOTLIB5 (Hunter 2007).

## REFERENCES

Aoyama S., Hou K.-C., Shimizu I., Hirashita H., Todoroki K., Choi J.-H., Nagamine K., 2017, *MNRAS*, 466, 105  
 Asano R. S., Takeuchi T. T., Hirashita H., Nozawa T., 2013, *MNRAS*, 432, 637  
 Barlow M. J., Silk J., 1976, *ApJ*, 207, 131  
 Beck A. M. et al., 2016, *MNRAS*, 455, 2110  
 Bekki K., 2013, *MNRAS*, 432, 2298  
 Bianchi S., Schneider R., 2007, *MNRAS*, 378, 973  
 Biffi V. et al., 2017, *MNRAS*, 468, 531  
 Calura F., Pipino A., Matteucci F., 2008, *A&A*, 479, 669  
 Camps P., Baes M., 2015, *Astron. Comput.*, 9, 20  
 Chabrier G., 2003, *PASP*, 115, 763  
 Charlot S., Fall S. M., 2000, *ApJ*, 539, 718  
 Chen L.-H., Hirashita H., Hou K.-C., Aoyama S., Shimizu I., Nagamine K., 2018, *MNRAS*, 474, 1545  
 Domínguez-Tenreiro R., Obreja A., Granato G. L., Schurer A., Alpresa P., Silva L., Brook C. B., Serna A., 2014, *MNRAS*, 439, 3868  
 Draine B. T., 2003, *ARA&A*, 41, 241  
 Draine B. T., Lee H. M., 1984, *ApJ*, 285, 89  
 Dwek E., 1998, *ApJ*, 501, 643

Galametz M., Madden S. C., Galliano F., Hony S., Bendo G. J., Sauvage M., 2011, *A&A*, 532, A56  
 Gaspari M., Ruszkowski M., Oh S. P., 2013, *MNRAS*, 432, 3401  
 Gauger A., Balega Y. Y., Irrgang P., Osterbart R., Weigelt G., 1999, *A&A*, 346, 505  
 Granato G. L., Lacey C. G., Silva L., Bressan A., Baugh C. M., Cole S., Frenk C. S., 2000, *ApJ*, 542, 710  
 Granato G. L., Ragone-Figueroa C., Domínguez-Tenreiro R., Obreja A., Borgani S., De Lucia G., Murante G., 2015, *MNRAS*, 450, 1320  
 Groenewegen M. A. T., 1997, *A&A*, 317, 503  
 Gutiérrez C. M., López-Corredoira M., 2017, *ApJ*, 835, 111  
 Hirashita H., 1999, *ApJ*, 510, L99  
 Hirashita H., 2015, *MNRAS*, 447, 2937  
 Hirashita H., Voshchinnikov N. V., 2014, *MNRAS*, 437, 1636  
 Hirashita H., Yan H., 2009, *MNRAS*, 394, 1061  
 Hou K.-C., Hirashita H., Michałowski M. J., 2016, *PASJ*, 68, 94  
 Hou K.-C., Hirashita H., Nagamine K., Aoyama S., Shimizu I., 2017, *MNRAS*, 469, 870  
 Hunter J. D., 2007, *Comput. Sci. Eng.*, 9, 90  
 Jenkins E. B., 2009, *ApJ*, 700, 1299  
 Jones A., 2013, *Proc. Sci.*  
 Jones E., Oliphant T., Peterson P. et al., 2001, SciPy: Open source scientific tools for Python  
 Kastner J. H., Forveille T., Zuckerman B., Omont A., 1993, *A&A*, 275, 163  
 Kennicutt R. C. et al., 2011, *PASP*, 123, 1347  
 Kennicutt R. C., Evans N. J., 2012, *ARA&A*, 50, 531  
 Li A., Draine B. T., 2001, *ApJ*, 554, 778  
 Madden S. C., Rémy A., Galliano F., Galametz M., Bendo G., Cormier D., Lebouteiller V., Hony S., 2012, in Tuffs R. J., Popescu C. C., eds, *Proc. IAU Symp.*, Vol. 284, The Spectral Energy Distribution of Galaxies - SED 2011, Preston, UK, p. 141  
 McKee C., 1989, in Allamandola L. J., Tielens A. G. G. M., eds, *Proc. IAU Symp.*, Vol. 135, Interstellar Dust, Santa Clara, CA, p. 431  
 McKee C. F., Hollenbach D. J., Seab G. C., Tielens A. G. G. M., 1987, *ApJ*, 318, 674  
 McKinnon R., Torrey P., Vogelsberger M., 2016, *MNRAS*, 457, 3775  
 McKinnon R., Torrey P., Vogelsberger M., Hayward C. C., Marinacci F., 2017, *MNRAS*, 468, 1505  
 Michałowski M. J., Murphy E. J., Hjorth J., Watson D., Gall C., Dunlop J. S., 2010, *A&A*, 522, A15  
 Montier L. A., Giard M., 2004, *A&A*, 417, 401  
 Murante G., Monaco P., Borgani S., Tornatore L., Dolag K., Goz D., 2015, *MNRAS*, 447, 178  
 Nozawa T., Kozasa T., Habe A., Dwek E., Umeda H., Tominaga N., Maeda K., Nomoto K., 2007, *ApJ*, 666, 955  
 Nozawa T., Asano R. S., Hirashita H., Takeuchi T. T., 2015, *MNRAS*, 447, L16  
 Padovani P., Matteucci F., 1993, *ApJ*, 416, 26  
 Pérez F., Granger B. E., 2007, *Comput. Sci. Eng.*, 9, 21  
 Planck Collaboration XLIII, 2016, *A&A*, 596, A104  
 Planck Collaboration XXII, 2015, *A&A*, 576, A107  
 Planelles S., Borgani S., Fabjan D., Killeddar M., Murante G., Granato G. L., Ragone-Figueroa C., Dolag K., 2014, *MNRAS*, 438, 195  
 Popping G., Somerville R. S., Galametz M., 2017, *MNRAS*, 471, 3152  
 Ragone-Figueroa C., Granato G. L., Murante G., Borgani S., Cui W., 2013, *MNRAS*, 436, 1750  
 Ragone-Figueroa C., Granato G. L., Ferraro M. E., Murante G., Biffi V., Borgani S., Planelles S., Rasia E., 2018, *MNRAS*, 479, 1175  
 Rémy-Ruyer A. et al., 2014, *A&A*, 563, A31  
 Rowlands K., Gomez H. L., Dunne L., Aragón-Salamanca A., Dye S., Maddox S., da Cunha E., van der Werf P., 2014, *MNRAS*, 441, 1040  
 Sanders D. B., Mirabel I. F., 1996, *ARA&A*, 34, 749  
 Schneider R., Valiante R., Ventura P., dell’Aglì F., Di Criscienzo M., Hirashita H., Kemper F., 2014, *MNRAS*, 442, 1440  
 Silva L., Granato G. L., Bressan A., Danese L., 1998, *ApJ*, 509, 103  
 Soifer B. T., Neugebauer G., 1991, *AJ*, 101, 354  
 Springel V., 2005, *MNRAS*, 364, 1105  
 Springel V., Hernquist L., 2003, *MNRAS*, 339, 289

- Tasker E. J., Tan J. C., 2009, *ApJ*, 700, 358
- Tielens A. G. G. M., McKee C. F., Seab C. G., Hollenbach D. J., 1994, *ApJ*, 431, 321
- Tornatore L., Borgani S., Dolag K., Matteucci F., 2007, *MNRAS*, 382, 1050
- Tsai J. C., Mathews W. G., 1995, *ApJ*, 448, 84
- Van Der Walt S., Colbert S. C., Varoquaux G., 2011, preprint ([arXiv:1102.1523](https://arxiv.org/abs/1102.1523))
- Wada K., Norman C. A., 2007, *ApJ*, 660, 276
- Weingartner J. C., Draine B. T., 2001, *ApJ*, 548, 296
- Winters J. M., Fleischer A. J., Le Bertre T., Sedlmayr E., 1997, *A&A*, 326, 305
- Yan H., Lazarian A., Draine B. T., 2004, *ApJ*, 616, 895
- Zhukovska S., Dobbs C., Jenkins E. B., Klessen R. S., 2016, *ApJ*, 831, 147

This paper has been typeset from a  $\text{\TeX}/\text{\LaTeX}$  file prepared by the author.