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Authors	SPINIELLO, CHIARA, NAPOLITANO, NICOLA ROSARIO, Arnaboldi, M., TORTORA, CRESCENZO, Coccato, L., Capaccioli, M., Gerhard, O., IODICE, ENRICHETTA, SPAVONE, MARILENA, CANTIELLO, Michele, Peletier, R., Paolillo, Maurizio, SCHIPANI, Pietro
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The Fornax Cluster VLT Spectroscopic Survey II – Planetary Nebulae kinematics within 200 kpc of the cluster core

C. Spiniello,^{1,2} N. R. Napolitano,^{1*} M. Arnaboldi,² C. Tortora,³ L. Coccato,²
M. Capaccioli,¹ O. Gerhard,⁴ E. Iodice,¹ M. Spavone,¹ M. Cantiello,⁵ R. Peletier,³
M. Paolillo⁶ and P. Schipani¹

¹INAF-Osservatorio Astronomico di Capodimonte, Salita Moiarriello, 16, I-80131 Napoli, Italy

²European Southern Observatory, Karl-Schwarzschild-Str. 2, D-85748 Garching, Germany

³Kapteyn Astronomical Institute, University of Groningen, PO Box 800, NL-9700 AV Groningen, the Netherlands

⁴Max Planck Institute for Extraterrestrial Physics, Karl-Schwarzschild-Str. 1, D-85741 Garching, Germany

⁵INAF-Osservatorio Astronomico di Teramo, via Maggini, I-64100 Teramo, Italy

⁶Dip. di Fisica, Università di Napoli Federico II, C.U. di Monte Sant'Angelo, Via Cintia, I-80126 Naples, Italy

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ABSTRACT

We present the largest and most spatially extended planetary nebulae (PNe) catalogue ever obtained for the Fornax cluster. We measured velocities of 1452 PNe out to 200 kpc in the cluster core using a counter-dispersed slitless spectroscopic technique with data from FORS2 on the Very Large Telescope (VLT). With such an extended spatial coverage, we can study separately the stellar haloes of some of the cluster main galaxies and the intracluster light. In this second paper of the Fornax Cluster VLT Spectroscopic Survey, we identify and classify the emission-line sources, describe the method to select PNe, and calculate their coordinates and velocities from the dispersed slitless images. From the PN 2D velocity map, we identify stellar streams that are possibly tracing the gravitational interaction of NGC 1399 with NGC 1404 and NGC 1387. We also present the velocity dispersion profile out to ~ 200 kpc radii, which shows signatures of a superposition of the bright central galaxy and the cluster potential, with the latter clearly dominating the regions outside $R \sim 1000$ arcsec (~ 100 kpc).

Key words: Astronomical data bases – catalogues – galaxies: formation – galaxies: kinematics and dynamics – dark matter.

1 INTRODUCTION

In the context of hierarchical structure formation and evolution, galaxy clusters represent the final stage of the growth of large-scale cosmic structures (Blumenthal et al. 1984; Davis et al. 1985; Bahcall 1988; Navarro, Frenk & White 1996). The Cold-Dark Matter scenario predicts that as the Universe cooled, clumps of dark matter (DM) began to condense and within them gas began to collapse, forming the first galaxies which then experienced structure growth via the so-called ‘merging tree’ (Kauffmann, White & Guiderdoni 1993; Cole et al. 2000; de Lucia & Blaizot 2007). During this growth, various physical processes like tidal interactions, ram pressure stripping, mergers, dynamical instabilities, secular evolution, and finally gas accretion occurred followed by cooling and star formation (e.g. Mihos 2003).

All these processes contributed to shape galaxies and the DM surrounding them and they are expected to have left their imprints in the outskirts of galaxies and beyond, out to the intracluster regions (see e.g. Napolitano et al. 2003; Arnaboldi et al. 2012). Here, dynamical times are longer and galaxy formation mechanisms leave signatures of gravitational interactions such as shells and tidal tails in the kinematics of their components or in the chemical composition of the stars (Murante et al. 2004, 2007; Bullock & Johnston 2005; Rudick, Mihos & McBride 2006; Duc et al. 2011; Longobardi et al. 2015a,b; Pulsoni et al. 2017).

Often the light profiles in the external region of central galaxies in group/clusters change slope with respect to the inner stellar light, passing from a Sérsic to an exponential profile (Seigar, Graham & Jerjen 2007; Donzelli, Muriel & Madrid 2011). This is the case for the Fornax cluster and its central galaxy NGC 1399 (Iodice et al. 2016; Spavone et al. 2017). NGC 1399 surface brightness profile has strong variations of the radial slope and ellipticity with radius, as demonstrated by deep photometry of the central regions of the cluster from VST (Iodice et al. 2016, hereafter **FDS-I**) and

* E-mail: chiara.spiniello@gmail.com

kinematical studies of the cluster core with globular clusters (out to 200 kpc, see Pota et al. 2018, in preparation, and Schuberth et al. 2010 for earlier results). In particular, FDS-I find that the averaged light profile of NGC 1399 in g band can be fitted with two components that contribute in different percentages to the total light at different radii: a Sérsic profile with $n = 4.5$, which dominates for $R < 10$ arcmin and an exponential outer component that contributes about 60 per cent to the total light and extends out to large radii (see figs 3 and 12 in FDS-I). Moreover, surface density maps of GCs, derived from multiband wide-field photometry, reveal the presence of a complex system of structures and substructures connecting NGC 1399 and its intergalactic environment (D’Abrusco et al. 2016; Cantiello et al. 2018). Photometry alone is however not sufficient to clarify the nature of this exponential profile nor to clearly separate the contribution to the light of the extended halo of NGC 1399 (stars gravitationally bound to the galaxy) from that of the intracluster light (ICL, stars that have been tidally stripped from the outer regions of the galaxy, mixing over time to form a diffuse light orbiting in the cluster potential; Dolag, Murante & Borgani 2010). Kinematical information is crucial for understanding if the ‘excess of light’ is made of diffuse halo stars that are still bound to the surrounding galaxies or by stars that are freely flying in the cluster potential.

A first attempt to distinguish dynamically the contribution of the central galaxy from the cluster environment was made by Napolitano, Arnaboldi & Capaccioli (2002), who argued that the velocity dispersion of the outer regions of NGC 1399 had their kinematics modified by the interaction with the potential, concluding that, realistically, the kinematics of the PNe bound to the galaxy potential only might flatten to $\sigma \sim 260 \text{ km s}^{-1}$.

Measuring kinematics in the outer parts of early-type galaxies (ETGs) remains, however, very challenging since the signal-to-noise ratio of the integrated light spectra drops off before the mass profile flattens. Long-slit, multi-object or integral field stellar spectroscopy provides measurements only up to few effective radii (R_{eff} , e.g. Cappellari et al. 2006; Tortora et al. 2009); only discrete tracers can push these limits beyond and help us to understand the formation history of the outer regions of elliptical galaxies (Hui et al. 1995; Arnaboldi et al. 1996; Méndez et al. 2001; Napolitano, Arnaboldi & Capaccioli 2002; Romanowsky et al. 2003; Douglas et al. 2007; de Lorenzi et al. 2008, 2009; Coccato et al. 2008, 2009; Napolitano et al. 2009, 2011; Richtler et al. 2011; Forbes et al. 2011; Pota et al. 2013; Longobardi et al. 2015a).

As a matter of fact, hints of the presence of kinematically distinct components in NGC 1399 have been found by Schuberth et al. (2010) in a dynamical study of 700 globular clusters (GCs) out to 80 kpc. They showed that the red (metal-rich) GCs might trace the spheroidal galaxy component and can be used to constrain the central DM halo, whereas the blue (metal-poor) GCs show evidence for kinematical substructures from accretion episodes during the assembly of the Fornax cluster. These results have been recently confirmed and expanded to larger radii in the companion paper of the Fornax Cluster Very Large Telescope (VLT) Spectroscopic Survey (FVSS), namely FVSS-I, by Pota et al. (2018, to be submitted), where GC velocities are measured out to 200 kpc.

However, GCs usually do not follow the same spatial distribution of the stars and they show most of the time a bi-modal colour distribution (Harris 1991; Ashman & Zepf 1998; Brodie & Strader 2006). Planetary nebulae (PNe) instead provide kinematics that are most of the time directly linked to integrated light measurements in ETGs (Douglas et al. 2007; Coccato et al. 2009). PNe represent part of the post-main-sequence evolution of most stars with masses in the range $0.8\text{--}8 M_{\odot}$, which means that they are drawn from the

same old/intermediate population that composes most of the light in ETGs.¹ PNe are sufficiently bright to be detected also in external galaxies in the local universe, and they are easier to detect at a large galactocentric radius where the background continuum is fainter. Thus, they represent a unique tool to measure the kinematics of elusive stars in low-surface brightness regions where they are easily observable through their [O III] emission at 5007 \AA . Measuring kinematics of Intracluster PNe (ICPNe) is a very effective way to measure the dynamical stage of the intracluster stellar population, even commonly called ICL, and to assess how and when its light originated (Napolitano et al. 2003; Murante et al. 2004, 2007; Gerhard et al. 2007).

Traditionally, PNe have been detected using an on-band/off-band technique (Ciardullo et al. 1989; Arnaboldi et al. 1998) where two images of the same portion of the sky are taken, one with a narrow-band filter centred at the redshifted [O III] $\lambda 5007$ line and the second with a broad-band filter. Sources in O III will be then be detected in the on-band image and too faint to be detected in the second one, whereas foreground stars will appear in both images with similar brightness. With this technique, however, a spectroscopic follow up is then necessary to measure PN velocities (Méndez et al. 2001; Arnaboldi et al. 2004; Teodorescu et al. 2005; Doherty et al. 2009).

In parallel, a number of techniques have been developed that allow detection and velocity measurements to be taken in a single step. One of the most successful is the so-called ‘counter-dispersed imaging’ (CDI), first developed by Douglas & Taylor (1999). With this technique, two images are obtained using a slitless spectrograph with a dispersive element rotated by 180 deg between the two exposures. In this way, the emission-line objects appear as unresolved dots in each dispersed image while stars appear as elongated streaks in the direction of dispersion with a length determined by the spectral resolution and the filter full width at half-maximum (FWHM). Due to the fact that at the two position angles (PAs), the grating disperses the monochromatic emissions in two opposite directions, the PNe will be shifted in the two rotated images by an amount proportional to their velocities. By registering the two images using dispersed spectra of foreground stars, matching the pairs of unresolved emitters and measuring the distance between them, one can simultaneously identify PNe and measure their velocities.

The use of CDI was demonstrated to be so successful and efficient that Douglas et al. (2002) decided to build a dedicated slitless spectrograph to study the kinematics of extragalactic PNe: the Planetary Nebulae Spectrograph (PN.S) mounted at the 4.2-m William Herschel Telescope. In the last 15 yr, the PN.S has produced a conspicuous number of referred publications and has substantially contributed to measure kinematics of galaxy haloes (Romanowsky et al. 2003; Merrett et al. 2006; Douglas et al. 2007; Noordermeer et al. 2008; Coccato et al. 2009; Napolitano et al. 2009; Cortesi et al. 2013), and of galaxy disc (Aniyan et al. 2018). In order to extend the study of diffuse outer haloes of ETGs using PNe as tracers for galaxies situated in the Southern hemisphere, we obtained counter-dispersed images with the FORS2 at the VLT. To carry out a complete mapping of an area of $50 \text{ arcmin} \times 30 \text{ arcmin}$ centred in the core of the Fornax cluster, for which we have already deep imaging coverage with VST (ugri) and VISTA (J and K), we obtained a total of 20 FORS2 pointings each covering an area of $\sim 6.8 \text{ arcmin} \times 6.8 \text{ arcmin}$. Complementing this data with previous

¹ We note that recent studies (e.g. Miller Bertolami 2016) point to the fact that the progenitors of PNe could be older than previously expected, up to 9–10 Gyr (especially for the bright ones).

CDI FORS1 observations of ~ 180 PNe around NCG 1399 and NCG 1404 published in McNeil et al. (2010), hereafter **MN10**, we are able to obtain velocity and velocity dispersion profiles up to ~ 30 arcmin from the central galaxy NCG 1399 (~ 6 times the effective radius, $R_{\text{eff}} \sim 5$ arcmin, as reported in **FDS-I**).

With the spectroscopic maps of the central 200 kpc of the Fornax cluster, PNe (and GCs) can be traced out to intracluster regions where they allow us to probe the cluster potential together with other satellite systems like ultra-compact dwarfs (UCDs) or dwarf galaxies (see e.g. Prada et al. 2003). The simultaneous constraints from independent tracers will allow us to study in great detail the transition region where the cluster potential starts to dominate the galaxy subhaloes. Here, one can expect to isolate unmixed substructures in the phase space, as the relics of interactions of satellite disruptions (Napolitano et al. 2003; Arnaboldi et al. 2004; Bullock & Johnston 2005; McNeil et al. 2010; Arnaboldi et al. 2012; Romanowsky et al. 2012; Coccato, Arnaboldi & Gerhard 2013; Longobardi et al. 2015a).

In this paper, we present positions and velocities of 1635 PNe (1452 of which are new detections), describe the data reduction and calibration (Section 2), explain our techniques and methods to identify PNe (Section 3), and infer their velocities directly from the calibrated and registered images (Section 4). A more complete and detailed description of the techniques is given in **MN10**, we therefore refer the reader to that paper for detailed information on the CDI calibration and data reduction. In Section 5, we present the final catalogue and the histogram of velocities, as well as histograms of velocities of the PNe associated with the three main galaxies and the line-of-sight (LOS) velocity and velocity dispersion distributions as a function of radial distance from NGC 1399. We extend by a factor of ~ 8 the number of detected PNe and by a factor of ~ 4 the radial spatial coverage of the previous PNe sample. Finally, conclusions and future perspectives are presented in Section 6. Throughout the paper, we assume a distance to the core of the Fornax Cluster of 20.9 ± 0.9 Mpc from Blakeslee et al. (2009), and therefore 1 arcsec corresponds to ~ 100 pc.

2 OBSERVATIONS AND DATA REDUCTION

2.1 Instrumental set-up and calibration

The observations have been acquired in P96 (096.B-0412(A), PI: M. Capaccioli) from 2015 November to 2016 December. 20 pointings, for a total of 50 h of observing time, were carried out with FORS2 on the 8-m large ESO VLT, with seeing generally below 0.8 arcsec and always below 1 arcsec, covering a total final area of ~ 50 arcmin \times 33 arcmin, centred around $\alpha = 3:37:51.8$ and $\delta = -35:26:13.6$. We show the final coverage of the FORS2 pointings (numbered from 1 to 20) on top of a Digital Sky Survey (DSS) image² centred on NGC 1399 in Fig. 1.

For each pointing, we acquired three scientific exposures, for a total of ~ 2050 s on target, on each PA (labelled respectively W0 and W180, where ‘W’ stands for West and the numbers refer to the angle aligned to this direction). The final spatial coverage complements³ the region already observed by **MN10** and extends from the centre

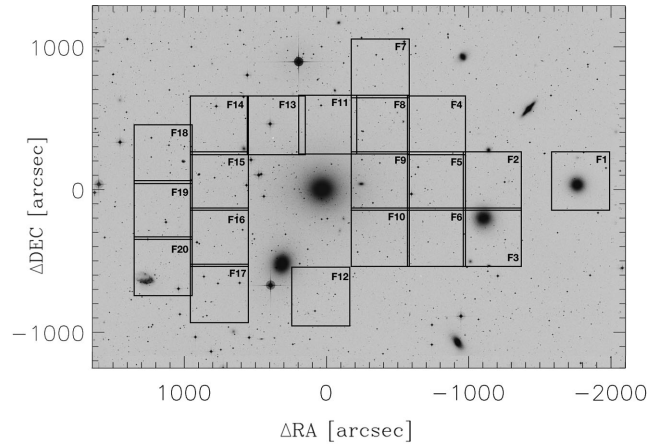


Figure 1. DSS Image of the Fornax Cluster Core (60 arcmin \times 40 arcmin) centred around NGC 1399. Black numbered boxes show the 20 FORS2 pointings obtained in P96.

of NGC 1399 out to ~ 1000 arcsec along the North–South direction and out to ~ 1300 arcsec along East–West, plus a separate pointing centred on NGC1379 (~ 1800 arcsec west of NGC1 399). We used the 1400V grating with a mean dispersion of $0.64 \text{ \AA pix}^{-1}$ and the O III/3000+51 filter, centred on the redshifted [O III] 5007 line (with the SR collimator, $\lambda_{\text{central}} = 5054 \text{ \AA}$, FWHM=59 \AA). In the focal plane, the presence of the grism causes an anamorphic distortion resulting in a contraction in the direction of dispersion which must be corrected before PN velocities can be calculated (see below). With this instrumental configuration, the direction of dispersion is along the row axis (x -axis) but because of the slitless technique, both axes retain spatial information.

We perform all the standard calibration tasks with IRAF and IDL. The scientific frames are bias subtracted and flat fielded individually. Cosmic rays are removed using the routine by van Dokkum (2001). The three scientific exposures of each pointing (for the same dispersion direction) are then combined using the IRAF task *imcombine*. Finally, the background is computed by smoothing with *fmedian* (30 pixels) each combined frame and then subtracting the smoothed image from the original one.

In addition to this, we also require special calibrations to take into account the effect of the dispersing optical element on our images. We use an MXU mask with a uniform array of slits. We requested to illuminate the mask with white light, then we requested to add a grism to the light-path and finally to disperse an arc lamp through the same filter ([O III]3000+51). These three calibration frames, shown in Fig. 2, allow us to solve for the local dispersion, the bandpass filter shift, and the anamorphic distortion introduced by the grism.

In order to calibrate and correct our images, we follow the recipes of **MN10**, using the same IRAF and MATHEMATICA tasks and complementing them with self-written IDL scripts. We refer the reader to that paper for an extensive data-reduction overview. Here, we briefly list the three main corrections we applied and the procedures we used to register the images, which is necessary to identify the PNe candidates.

(i) *Mapping the anamorphic distortion:* We use the x , y -positions from the image of the slitlets and the positions of the brightest line ($\lambda 5015.67$) in the dispersed arc lamp calibration frame to map the anamorphic distortion introduced by the grism. We tabulate the difference in x and y from the slit x_s and y_s to the monochromatic x_{λ_1} , y_{λ_1} , and use it to make a map of the anamorphic distortion with

²The image has been taken from the IRSA at <http://irsa.ipac.caltech.edu/data/DSS/>.

³We did not re-observe the central region already presented in **MN10** and we only have a minimal overlap between the pointings, necessary to calibrate our PN velocities, as explained in Section 4.2.

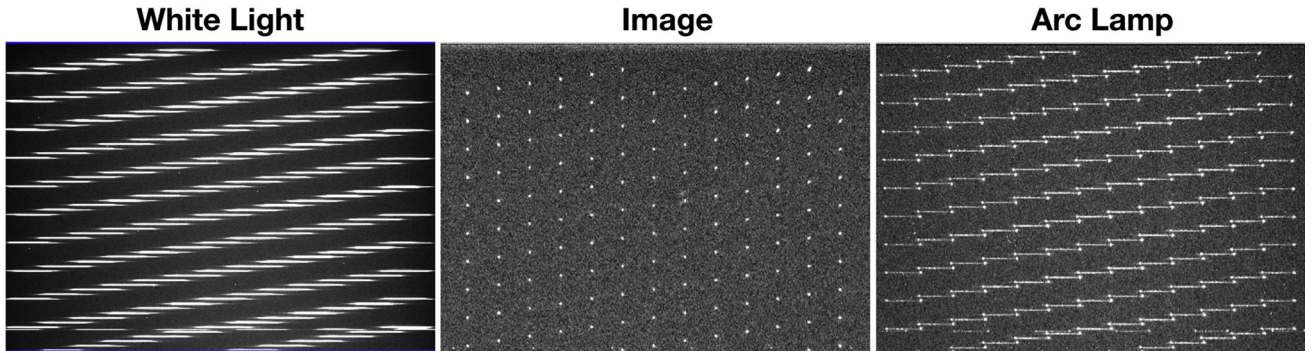


Figure 2. Calibration MXU frames. The left-hand panel shows the white light lamp used to calculate the local dispersion and to measure the bandpass shift. Middle and right-hand panels show the image and the arc lamp used to map the anamorphic distortion. We use the standard HgCdHeAr arcs template and, in particular, two He lines (4921.93\AA , and 5015.67\AA) are transmitted through the narrow filter bandpass and are visible in the right-most panel.

the IRAF task *geomap*. Using *geotran*, we then applied the solution from *geomap* to all of our images to remove such a distortion.

(ii) *Calculating the local dispersion*: We measure the distance between two lines in the dispersed arc lamp. Assuming that the dispersion is linear on this scale, we then convert our dispersions from the measured spectral-plane dispersions to image-plane dispersions using IRAF tasks and IDL routines.

(iii) *Measuring the bandpass shift*: We map the bandpass filter shift as a function of position. We do this by comparing the Gaussian centre of the dispersed white light to the position of the $\lambda 5015.67$ line. As in MN10, we find the underlying smooth 2D function that describes the bandpass shift caused by the presence of the interference filter into a converging beam and we model it using MATHEMATICA.

(iv) *Registration of the images*: Before calculating velocities, we must register the W0 and W180 corrected and calibrated frames in order to have all the strips and extended objects in the same relative positions in each pair. The registration is done by applying a rigid shift calculated by mapping the position of few bright stars, after adjusting them for the bandpass shift (see MN10). The location of a star in our images is always measured with a 2D Gaussian fit using the *n2gaussfit* task in IRAF. For each pointing, we measured the positions of ~ 15 profiles of stars spanning the whole field of view of the pointing in the spectral plane; we apply the described bandpass shift, correct for anamorphic distortion, and rotate the corrected positions of the W180 frame.

Spectral and image planes are defined in Arnaboldi et al. (2007): the spectral plane is the dispersed image as observed, and the image plane has been corrected for the distortion introduced by the grism thereby making an undispersed, distortion-corrected image for one wavelength. In general, to switch from the spectral-plane, where the bandpass shift is measured, to the image-plane, which we need to use to identify PNe, we use the IRAF task *geoxytran*.

3 CANDIDATES IDENTIFICATION

The CDI is a slitless technique that makes use of two counter-dispersed frames (taken by rotating the PA of the field by 180 deg) where the light is selected by a [O III] filter: here, oxygen emission-line objects (like PNe) appear as spatially unresolved monochromatic sources, while continuum sources (stars, background galaxies) show up as strikes or star trails. The velocity of the PNe is obtained by measuring the displacement of the [O III] emission with respect to a calibrated frame (see MN10, for further details). Thus,

CDI allows both the detection and measurement of the Doppler shift with a single observation. Fig. 3 shows an example of the best-quality single FORS2 field (FIELD3) taken in the two PAs, W0 and W180. In these images, which have already been corrected, calibrated, and registered, monochromatic spatially unresolved objects are easy to spot because they appear as unresolved sources. Emitters will appear in the two images at the same y -position, with similar intensity (not necessarily identical, given that the images quality might be slightly different) but shifted in x -direction of an amount which is indeed proportional to their LOS velocity.

Standard criteria to identify PNe have been established and successfully tested by the PN.S community and previous authors (Arnaboldi et al. 2002; Romanowsky et al. 2003; Douglas et al. 2007; McNeil et al. 2010; McNeil-Moylan et al. 2012; Pulsoni et al. 2017). In particular, they are based on the identification of emission-line objects that are then classified into three main classes:

- 1) *PNe candidates*: They are monochromatic sources spatially unresolved in both wavelength and space.
- 2) *Lyman- α galaxies*: These are Lyman- α galaxies at $z \sim 3$ that are most of the time associated with a continuum and therefore can be easily excluded from our final PNe catalogue (see Fig. 3). To estimate the contamination of Lyman- α emitters not associated with a continuum, we use a very simple argument: there is no reason to believe that the number of background galaxies changes for different velocity bins. Directly from the final histogram of velocity that we present in the next section, we see that we find only few objects (4) in the velocity range $2500\text{--}3000 \text{ km s}^{-1}$. This implies that on the full velocity range (0–3000), we expect to find $\lesssim 30$ Lyman- α contaminants, which represent $\lesssim 2$ per cent of our total sample.
- 3) *[O II] emitters at $z \sim 0.347$* : At this redshift, the [O II] doublet will fall in our narrow-band filter and therefore objects will appear bright in the images. However, thanks to spectral resolution of the grism (GRIS_1400V), we are able to resolve the doublet (see Fig. 3) ensuring in this way a minimal amount of contamination from this class of objects in the final PNe catalogue.

In Fig. 3, we highlight with green circles our PNe candidates and with orange ones the other background emission-line objects.

For the identification of the PNe, we rely on visual inspection: we blink the two counter-rotated, registered and shifted images and select point-like sources with same y -position and different x -positions and not associated with a continuum. Moreover, to identify PNe at the very centre of galaxies, where the integrated stellar light dominates, we subtract the W180 image from the W0 frame

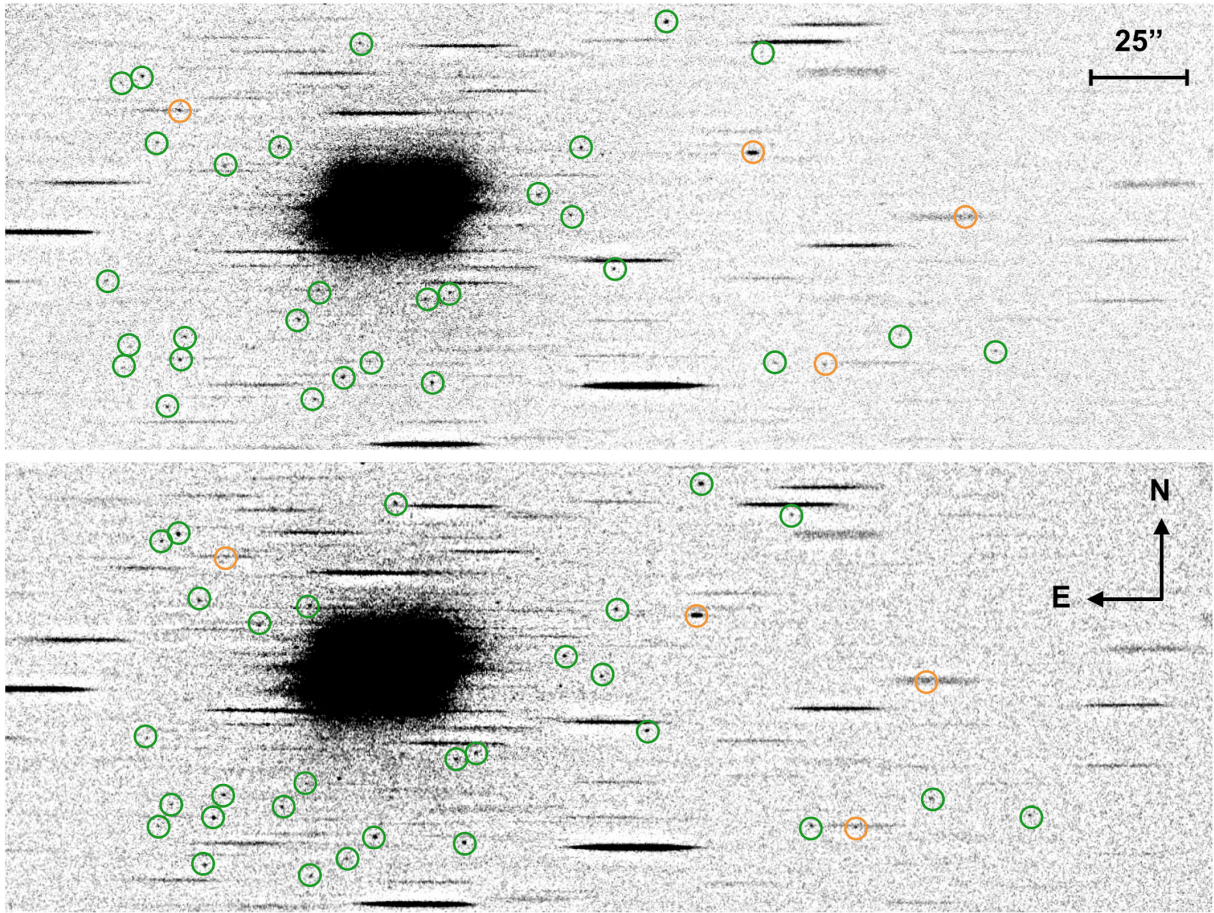


Figure 3. Part of a single pointing (Field 3). The *top frame* is the W0 exposure while the *bottom frame* is the counter-dispersed W180 exposure rotated back to be compared to the first. The big diffuse light on the left centre of each frame is NGC 1387 while the other horizontal strikes are stars. The two images have been corrected, registered and calibrated (see the text for more details) in a way that the galaxy and the stars appear in the same position. Only PNe and point-like emitters have different x -positions in the two panels. However, the [O II] doublet (the emission from galaxies at $z \sim 0.347$) is resolved, thanks to the adopted grism, and most of the background Lyman- α galaxies are associated with an elongated continuum. Some examples (not all of them, for clarity of the image) of emitters are circled (PNe in green, and background objects in orange).

and we search for emission-line objects as positive/negative residuals.

In order to make the procedure as objective as possible and to limit the spurious detections, sources are identified by two different members of our team, who have checked the registered images in a completely independent way. A PN is added to the final catalogue only if it is confirmed by both researchers independently.

4 PLANETARY NEBULAE LINE-OF-SIGHT VELOCITIES

As already discussed, the counter-dispersion shifts the x -position of the PNe that appear in different places in the W0 and the W180 images. The LOS velocity of a PN is thus a function of the separation between the two positions:

$$\lambda = \lambda_0 + \frac{d\lambda}{dx} \frac{\Delta x}{2}, \quad (1)$$

where λ is the measured wavelength of the planetary, λ_0 is the central wavelength of the passband narrow filter (used in the registration procedure), $d\lambda/dx$ is the local dispersion, and finally Δx is the separation, in pixels, between the position of the same PN in the W0 and W180 frames.

4.1 Errors on the velocity measurements

The uncertainties on velocity measurements of single PNe come from three different sources: i) the uncertainties on the PN position in each image (W0 and W180) – roughly corresponding to half of a pixel, ii) the uncertainties on transformation from spectral-plane to image-plane necessary to correct for anamorphic distortion and local dispersion, and iii) The uncertainties on the bandpass shift. We used the 1400V grating with a measured mean dispersion of $0.6396 \text{ \AA px}^{-1}$. Thus, propagating these errors from the formula in equation (1), we obtain uncertainties on the single velocity measurements. These are of the order of $30\text{--}45 \text{ km s}^{-1}$, perfectly consistent with the errors given in MN10 and in Méndez et al. (2001) and slightly larger than the 20 km s^{-1} errors in the Planetary Nebula Spectrograph data of Douglas et al. (2002). Detailed information on the errors on each PN velocity will be provided with the full catalogue (positions, v_{LOS} and m_{5007}) in a publication in preparation.

4.2 Absolute velocity calibration

The LOS PN velocities cause a shift of the monochromatic emission with respect to the LOS velocity correspondent to the central wavelength of the narrow-band filter. The mean velocity of the full

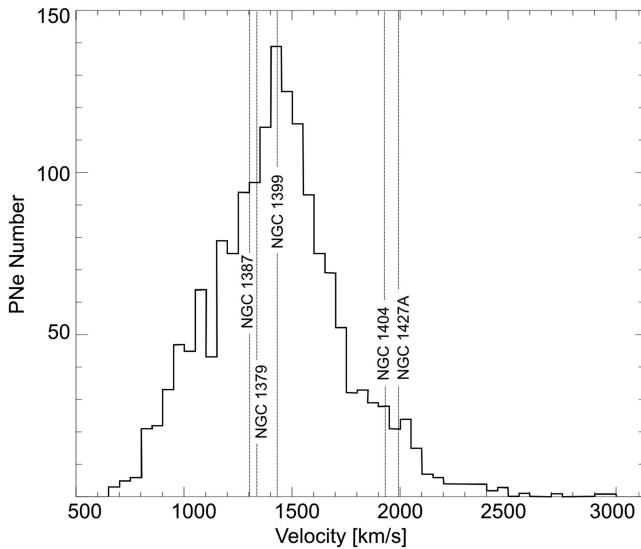


Figure 4. Histogram of velocities of the full catalogue of 1635 PNe, including **MN10** data. Vertical dotted lines show the systemic velocities of the Fornax Galaxy members observed within our pointings.

catalogue is 1433 km s^{-1} with a standard deviation of 312 km s^{-1} , calculated after applying heliocentric correction to each field separately. In Fig. 4, we show the total, global histogram of velocities. In the plot, we also include the PNe presented in **MN10** to have a homogeneous spatial coverage from NGC 1399 outwards. Vertical arrows highlight the systemic velocities of the Fornax galaxy members observed within our pointings. We cover NGC 1379 and NGC 1387. We also complete the coverage on NGC 1404 already partially observed by **MN10** and finally, we also observe the irregular galaxy NGC 1427A.⁴

We calibrate our line-of-sight velocities (LOS_V) measurements on the basis of those by **MN10** by comparing the 14 PNe in common between the two samples. **MN10** objects were in turn calibrated against the 1994 NTT multi-object-spectroscopy measurements of Arnaboldi et al. (1994).⁵ **MN10** found a systematic offset of $+166 \text{ km s}^{-1}$ with respect to Arnaboldi et al. (1994), which was independent of the position on the CCD. For the few (14) PNe in common between us and **MN10**, we find a negative systematic offset of 46 km s^{-1} between our velocities and those of **MN10**, which we apply as a zero velocity offset to our catalogue. The objects in common with **MN10** are found in three different fields, pointing at three different sky positions. Moreover, we detect few (3) objects in an overlapping region between two of our different fields. For them we obtain two independent measurements of velocity, which resulted to be in perfect agreement. We are therefore confident that the offset does not change from pointing to pointing.

To further confirm our velocity calibration, we focus on fields where we observe galaxy members of the Fornax cluster: NGC

1379, with a systemic velocity of $V_{\text{sys}} = 1324 \text{ km s}^{-1}$ and an effective radius of $R_{\text{eff}} = 23.3 \text{ arcsec}$ (Caon, Capaccioli & D’Onofrio 1994), observed in FIELD1; NGC 1387, with a systemic velocity of $V_{\text{sys}} = 1302 \text{ km s}^{-1}$ and $R_{\text{eff}} = 42 \text{ arcsec}$ (de Vaucouleurs 1991), observed in FIELD2 and FIELD3 (see Fig. 1) and finally the central galaxy NGC 1399 ($V_{\text{sys}} = 1425 \text{ km s}^{-1}$ and $R_{\text{eff}} = 303 \text{ arcsec}$; **FDS-I**⁶), mainly covered by **MN10**. For NGC 1399 and NGC 1379, systemic velocities were taken from the NASA/IPAC Extragalactic Database (NED⁷). For NGC 1387, we calculate the velocity directly from a spectrum obtained with The Wide Field Spectrograph at the Australian National University 2.3 Telescope.⁸ We take all the PNe within one effective radius of each galaxy, assuming that the probability that they are bound to such galaxy is high, and calculate their median velocity, after applying the absolute velocity calibration described above. Fig. 5 shows the histograms of velocities for the PNe associated with these three galaxies. In each panel, the black Gaussian line is centred on the median velocity inferred by the PNe distribution and broadened by the standard deviation of the PN velocities, and the magenta Gaussian line is centred on the systemic velocity from the NED data base and has a standard deviation equal to the tabulated central velocity dispersion of the galaxy, further aperture corrected to the R_{eff} using the formula in equation (1) of Cappellari et al. (2006). We caution the reader on the fact that not all the PNe spatially close to a given galaxy must in fact be bound to it. This could cause the overestimation of the standard deviation of the PNe, which is larger than the stellar velocity dispersion. A more rigorous bounding criterion will be defined in a forthcoming paper of the current series. For two out of the three systems, we found a very good agreement. Only for NGC 1387, we infer a velocity which is $\sim 100 \text{ km s}^{-1}$ larger than that tabulated in the literature. However, we note that the two measurements are within 1σ of each other. We believe that part of this difference could be caused by the different image quality between FIELD2 and FIELD3 combined with the fact that this galaxy (and the PNe bound to it) shows a clear sign of rotation. In fact, FIELD3 points to the receding part of the galaxy and is the field with the best image quality we have (in this pointing, we found a total of 143 PNe). On the contrary, FIELD2, on the approaching side of the galaxy rotation, has lower image quality (in total 95 PNe). Consequently, the number of PNe with receding velocity with respect to the systemic velocity of NGC 1387 is larger than the number of PNe with approaching velocity. We finally note that the PNe number within $1 R_{\text{eff}}$ of NGC 1387 is lower than the number found for the other two galaxies (most probably because NGC 1387 might have a lower PN specific density and a smaller effective radius⁹), making the statistics in the histogram poorer. Finally, we will show in the following sections that hints for gravitational interaction between NGC 1399 and NGC 1387 are suggested by the presence of a stream of high-velocity PNe connecting the two galaxies. Some of these high-velocity PNe might overlap spatially with the PNe gravitationally bound to NGC 1387, thus broadening the PN LOSVD at the location of this galaxy.

⁴ Rather than being disrupted during its first passage through the cluster as reported in (Chanamé, Infante & Reisenegger 2000), a recent paper by Lee-Waddell et al. (2018) claims that the irregular optical appearance of NGC 1427A might have tidal origins.

⁵ Since our field pointings strategy is complementary to the observations of **MN10**, who covered the central region of NGC 1399, we do not have objects in common with Arnaboldi et al. (1994), which are restricted to the innermost 200 arcsec from NGC 1399. This is the reason why the two-steps calibration is adopted.

⁶ We note that **FDS-I** recently recomputed the effective radius for NGC 1399, originally given in Caon, Capaccioli & D’Onofrio (1994), using deep photometry in *g* and *i* band and obtaining larger values of $R_{\text{eff}, g} = 5.87 \pm 0.10 \text{ arcmin}$ and $R_{\text{eff}, i} = 5.05 \pm 0.12 \text{ arcmin}$.

⁷ <https://ned.ipac.caltech.edu>

⁸ The spectra have been kindly shared with us by Kenneth Freeman and Mike Bessell, to whom we are grateful.

⁹ For a definition of the PN luminosity-specific number and its variation as a function of galaxy type and colour, we refer to Buzzoni, Arnaboldi & Corradi (2006).

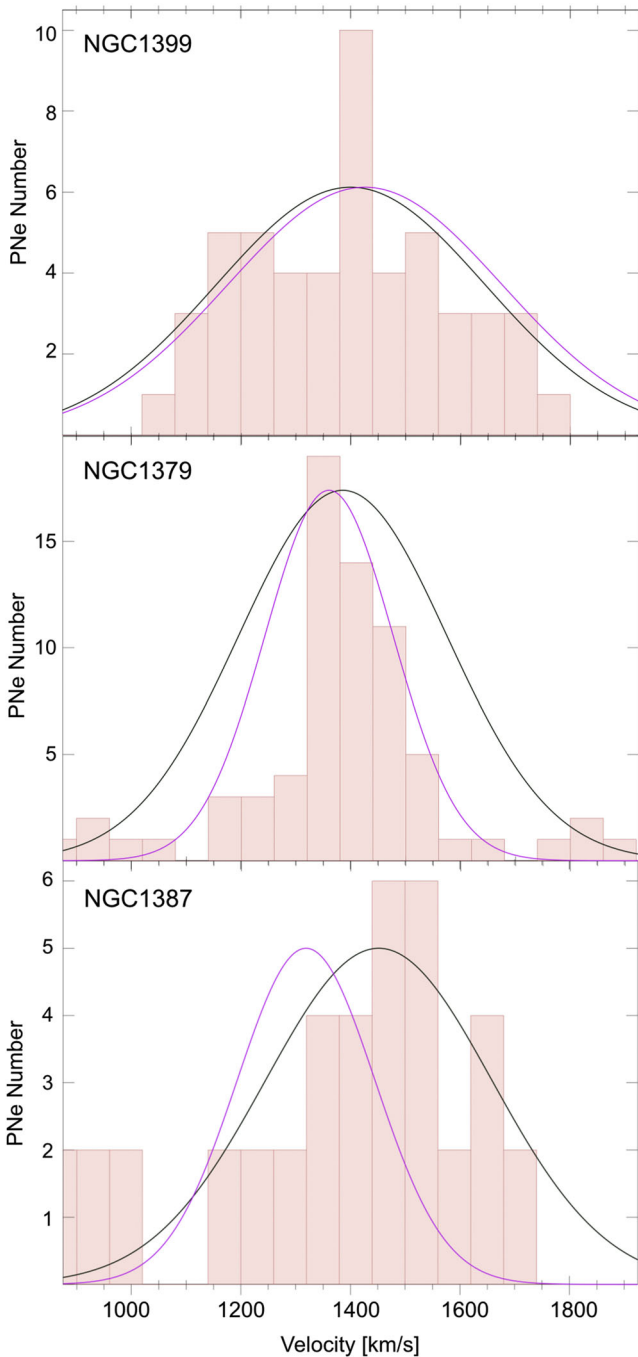


Figure 5. Histograms of velocities for the PNe within the effective radius of NGC 1399 (upper panel, from MN10), NGC 1379 (in FIELD 1), and NGC 1387 (between FIELD 2 and 3). In each panel, the black Gaussian line is centred on the median velocity inferred by the PNe distribution with its standard deviation, and the magenta Gaussian line is centred on the systemic velocity and has a standard deviation equal to the tabulated central velocity dispersion of the galaxy, corrected for aperture, using the formula of Cappellari et al. (2006). An overall good agreement is found, although the median velocity calculated for NGC 1387 is $\sim 100 \text{ km s}^{-1}$ larger than the one inferred from a spectrum centred on core of the galaxy (see the text for a possible explanation for this partial disagreement).

It is worth noticing that this kinematical stream is in the same spatial region where FDS-I found an ~ 5 arcmin long faint stellar bridge ($\mu_g \sim 29\text{--}30 \text{ mag arcsec}^{-2}$) and D’Abrusco et al. (2016) and Cantiello et al. (2018) reported an overdensity of the distribution of blue GCs. We argue that the stream is the result of the stripping of the outer envelope of NGC 1387 on its east side. We will further speculate on this point in the following section.

We thus conclude about the robustness of our velocity calibration. In fact, we reproduce convincingly the systemic velocities of the three galaxies that fall within our pointings and we find a good agreement (after the shifting) with the results of MN10. Further confirmation can be also found in Fig. 4 where we show the final histogram of velocities and overplot the systemic velocities of the observed Fornax galaxy members.

5 RESULTS

5.1 The final PN velocity sample

The final catalogue of PNe with measured velocities comprises 1635 objects (of which 1452 are new detections) and extends conspicuously in both number and spatial coverage in all previous overlapping PNe catalogues of the Fornax Cluster. The phase space distribution diagram is shown in Fig. 6. PN velocities are plotted as a function of radius (calculated as distance from NGC 1399, in circular radii) whereas NGC 1404, NGC 1379, NGC 1387, and NGC 1427A are plotted as coloured circles. NGC 1399 is plotted as a black triangle and its systemic velocity is shown as a dotted black line through the figure. Note that the absence of data at radii $1450 \text{ arcsec} < R < 1650 \text{ arcsec}$ is due to the geometrical coverage of our pointings.

Fig. 7 shows the spatial location of the PNe identified in this work (circles) and those presented in MN10 (squares) colour-coded by their velocity. The PNe are overplotted on a DSS image of $\sim 70 \times 40$ arcmin. An interesting result that one can infer from this figure is the presence of streams that might be directly related to the history of the cluster as a whole, tracing recent streams falling into the cluster (highlighted in the figure with dotted lines). Already MN10 reported the presence of a low-velocity PNe subpopulation moving at about 700 km s^{-1} slower than NGC 1399. We confirm this detection finding ~ 50 PNe with $700 \text{ km s}^{-1} < V < 1000 \text{ km s}^{-1}$ distributed between $+300 \text{ arcsec} \leq \Delta\delta \leq +700 \text{ arcsec}$ north of NGC 1399 (dotted big circle in the figure). Thanks to our more extended coverage, we also highlight a high-velocity PNe stream extending for $\sim 700 \text{ arcsec}$ in Dec. to the west side of NGC 1399, connecting the central galaxy to NGC 1387, which confirms the findings of FDS-I. A more detailed analysis of PN velocities streams and a direct comparison with 2D kinematical maps obtained with GCs covering a similar region (presented in FVSS-I) and with deep photometry (from the FDS) will be presented in a forthcoming paper of the series. Here, we wish to stress the similarity and good spatial agreement with FDS-I and Iodice et al. (2017), who detected this previously unknown region of ICL and also related it to an overdensity in the population of blue GCs (D’Abrusco et al. 2016; Cantiello et al. 2018). These authors found that the ICL in this region contribute ~ 5 per cent to the total light of the brightest cluster member. They also compared their findings to theoretical predictions for the ICL formation and support a scenario in which the intracluster population is built up by the tidal stripping of material from galaxy outskirts in close passages of the central galaxy.

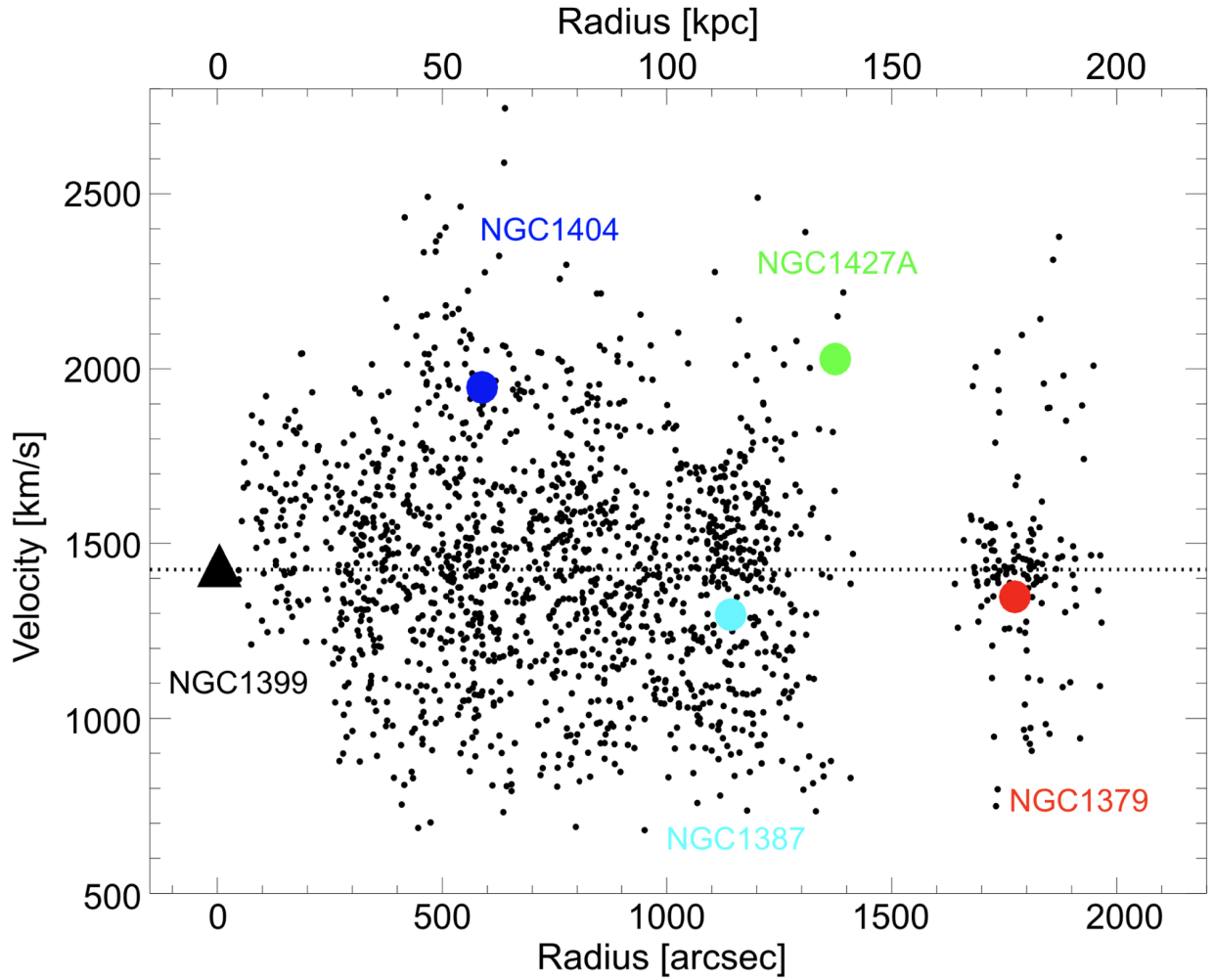


Figure 6. Phase space distribution diagram. NGC 1399 is plotted as a black triangle and its systemic velocity is shown as a dotted black line. The coloured circles represent the other cluster members, with names plotted in the figure with the same colour as the symbol. PN velocities are shown as a function of radius from NGC 1399 in black (black small circles).

5.2 The radial velocity dispersion profile

Thanks to the large spatial coverage of our PNe catalogue, we can trace the velocity dispersion of NGC 1399 out to large radii, well beyond the integrated-stellar-light measurements (Saglia et al. 2000). This is crucial if one aims at covering the transition from the extended galaxy halo to the cluster potential (Gerhard et al. 2007; Dolag, Murante & Borgani 2010). To estimate the LOS velocity dispersion as a function of projected radius, we define circular apertures (annuli) at different radial distances from NGC 1399 and calculate the standard deviation of all the PNe in each of them. Given the irregular final spatial coverage of our sample, which extends more in RA than in Dec., we also follow a different approach and simulate a slit that extends 2600 arcsec in the EW direction (± 1300 arcsec from the centre of NGC 1399) and 600 arcsec in the NS direction (± 300 arcsec from the centre of NGC 1399). We choose this slit-width and orientation to maximize the number of PNe used and at the same time to have a homogeneous and complete spatial coverage. Moreover, we did not orient the slit along the major axis of NGC 1399 because we are not interested in features belonging to the central galaxies but we want to map the kinematics of the Fornax Cluster core. We then divide the slit in distance bins. The radial limits of both the circular annuli and the slit-bins are set

such that all the bins have about the same number of objects (90–91 in the first case, 69–70 in the second). Moreover, in both cases, we add an additional point considering the PNe around NGC 1379. A visualization of the slit and the additional point around NGC 1379 is shown in Fig. 7 as magenta dashed boxes. In Table 1, we report the computed values for the velocity dispersion in all radial bins for the two approaches we undertook. Errors on the velocity dispersion values were estimated using fig. 7 in Napolitano et al. (2001), considering that we have 90 (70) PNe for each annulus (slit-bin). Finally, we added two additional data points (empty symbols) corresponding to the radii where there is an overlap with cluster galaxies in the field (namely NGC 1387 and NGC 1379) where we have recomputed the velocity dispersion excluding the PNe that in first approximation could be bound to the corresponding galaxies. We use the same ‘spatial’ criterion defined in Section 4.1 and used in Fig. 5. We caution the reader that, however, this is only a first-order approximation and that a more robust classification will be performed in forthcoming papers.

Fig. 8 shows our LOS velocity dispersion profile (filled black points for the annuli and magenta points for the slit) compared to different collections of data from the literature. In particular, we overplot the velocity distribution values for the red and blue GCs presented in Schuberth et al. (2010) (red and blue squared,

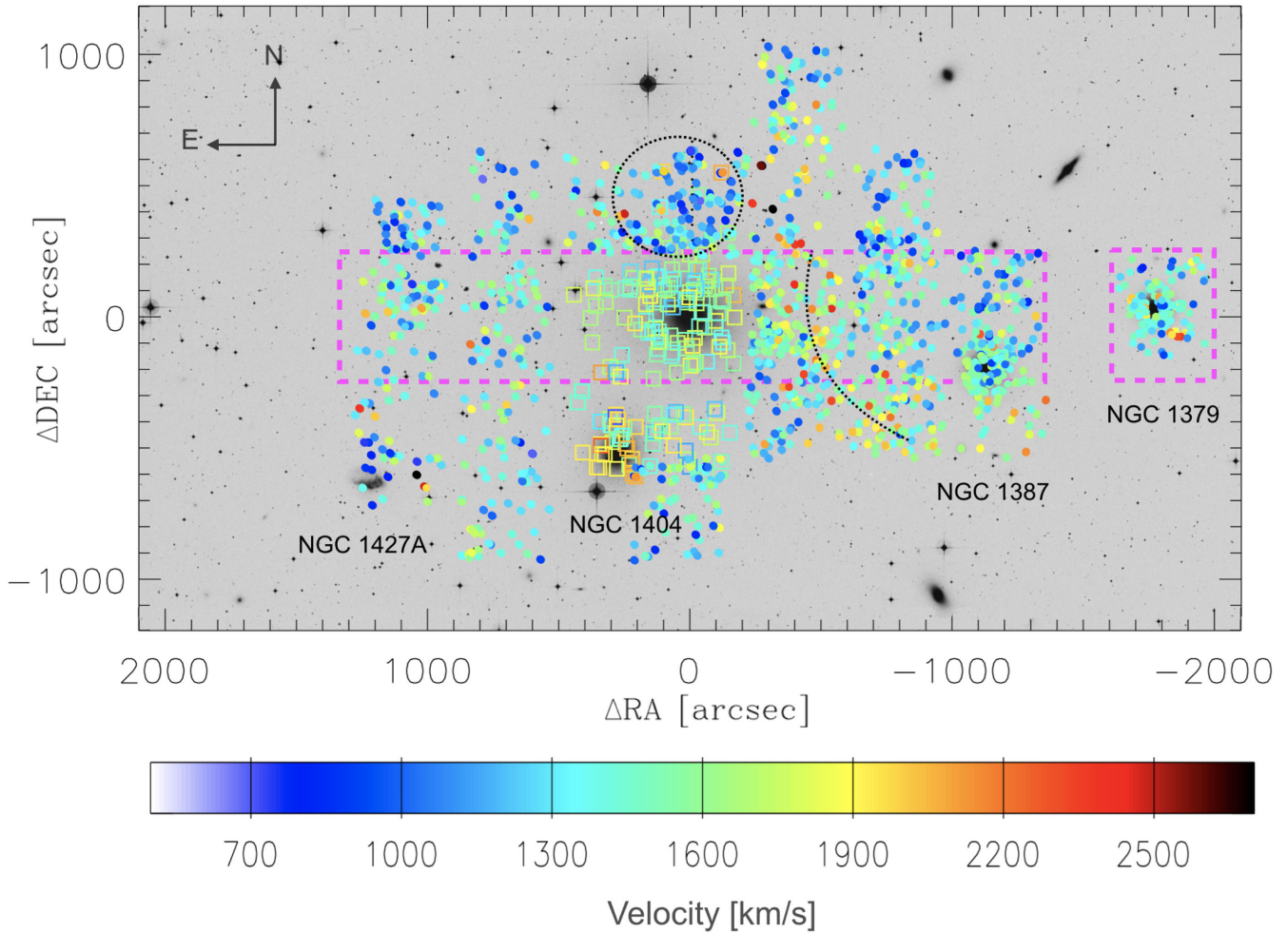


Figure 7. The final PNe catalogue overlaid on a DSS image of the Fornax Cluster. Circles are objects identified in this work, and squares are the PNe already presented in MN10. The PNe are colour-coded by their velocity. Two streams of intracluster PNe can be highlighted from the image: a low-velocity PNe patch north of NGC 1399 (identified by a dotted big circle) and a high-velocity PNe bridge connecting the central galaxy with NGC 1387 (identified by the dotted curved line). See the text for more details. The magenta slit overlaid on the PNe is the slit we define to obtain the velocity dispersion radial profile (corresponding to magenta points in Fig. 8).

respectively)¹⁰ and the integrated-stellar-light profile along the major axis of NGC 1399 from Saglia et al. (2000) (yellow triangles). Finally, the grey shaded horizontal region in the figure represents the 1σ contours of the velocity dispersion of the Fornax Cluster ($\sigma_{cl} = 374 \pm 26 \text{ km s}^{-1}$) reported in Drinkwater, Gregg & Colless (2001), hereafter D01. We go twice as far from NGC 1399 compared to all the previously published velocity dispersion profiles, covering a radial distance of ~ 2000 arcsec (~ 200 kpc) from NGC 1399.

In the plot, we identify three regions delimited by two characteristic radii, separated in the figure with dotted cyan lines:

1) $R < 250$ arcsec. Here, the PN velocity dispersions match the gradient found by Saglia et al. (2000) from integrated-light and agree well with the dispersions reported in Schuberth et al. (2010) for the red, metal-rich population of GCs that trace the halo of

NGC 1399. PNe trace the kinematics of the central galaxy and in this region the population of *virialized* stars dominates, based on results from cosmic simulations (e.g. Cooper et al. 2013, 2015). Indeed, Napolitano, Arnaboldi & Capaccioli (2002) modelled the regions within ~ 400 arcsec and found a dynamical signature at around ~ 200 arcsec of non-equilibrium, consistent with the presence of a non-mixed population possibly in relation with the interaction with the cluster potential or with companion galaxies (e.g. NGC 1404).

2) $250 \text{ arcsec} < R < 1000 \text{ arcsec}$. The PNe velocity dispersion rises steeply around $300\text{--}400$ arcsec ($\sim 1\text{--}1.5R_{\text{eff}}$) from NGC 1399, reaching a peak around ~ 500 arcsec, similar to the blue, metal-poor GCs. As already stressed, a similar ‘transition’ region has been identified by FDS-I, who also performed a cross-analysis joining light, GCs, and X-ray (Paolillo et al. 2002) and highlighting that all these tracers show a different slope with respect to the innermost regions (see fig. 13 in FDS-I). In their section 4.4, FDS-I highlight that the shape of the external light profile can be used to study the accretion mechanisms responsible for the mass assembly of galaxy clusters and that only kinematics can disentangle between ‘broken’ light profiles which are associated with short and recent accretion of massive satellites (Deason et al. 2013) and ‘shallower profiles’

¹⁰ We plot the full data set from Schuberth et al. (2010) without interloper removal and including the GCs within 3 arcsec of NGC1404 and those with extreme velocities. We trust this is an unbiased approach to carry out the comparison of GCs and PNe LOSVDs in the surveyed regions.

Table 1. Velocity dispersion measurements for the two binning approaches. By constructions, the bins have the same number of objects (90–91 in the first case, 69–70 in the second).

Radius (annuli) arcsec	Vel. Disp. (annuli) km s ⁻¹	Radius (slit) arcsec	Vel. Disp. (slit) km s ⁻¹
144 ± 98	197 ± 20	116 ± 70	191 ± 23
285 ± 42	240 ± 24	243 ± 56	253 ± 27
359 ± 32	265 ± 26	335 ± 35	258 ± 28
428 ± 37	376 ± 37	426 ± 56	295 ± 33
499 ± 33	391 ± 39	537 ± 52	341 ± 39
555 ± 23	354 ± 36	660 ± 70	303 ± 38
609 ± 31	364 ± 36	786 ± 51	328 ± 41
682 ± 42	322 ± 32	916 ± 78	296 ± 33
756 ± 31	319 ± 32	1051 ± 56	271 ± 30
817 ± 30	312 ± 31	1131 ± 24	275 ± 30
879 ± 31	329 ± 33	1185 ± 30	257 ± 28
956 ± 46	269 ± 27	1310 ± 95	285 ± 32
1041 ± 38	293 ± 29	1860 ± 83	303 ± 38
1107 ± 27	254 ± 25		
1157 ± 32	288 ± 28		
1227 ± 37	330 ± 33		
1860 ± 150	284 ± 28		

Notes. Col. 1, 3: Distances from NGC 1399 of the annuli and slit-bins, respectively. Col. 2, 4: Velocity dispersions calculated as standard deviation of all the PNe belonging to the correspondent bin for the circular annuli (col. 2) and slit (col. 4) approach. Uncertainties are assumed to be equal to 10 per cent of the measurement for the first case and 12 per cent of the measurement (extrapolated from fig. 7 of Napolitano et al. 2001, given the number of PNe in each bin).

(with excess of light) which indicate minor mergers events that took over long time-scales where stars have undergone more mixing. In fact, in the first case (accretion of a massive satellite), a peak in velocity dispersion can be observed for radii larger than the break radius. Adding the missing piece (kinematics information up to large radius) to the study of FDS-I, we show here that this is indeed the case for Fornax. In this region, of course, also PNe bound to NGC 1404 might contribute in the rise in velocity dispersion. However, in the slit approach (magenta points), where NGC 1404 and the associated PNe are not present by construction, the rise is still present, although less pronounced.

3) $R > 1000$ arcsec. Finally, at these distances, our measurements flatten out, at a value of $\sigma_{\text{ICPNe}} \sim 300$ km s⁻¹, higher than the velocity dispersions measured for the individual galaxies and closer (but slightly lower) to the value reported by D01 for the Fornax Cluster (σ_{cl}). Here, the PNe trace the cluster potential, measuring the kinematics of the ICL. We note also that our σ_{ICPNe} is in perfect agreement with the velocity dispersion of the giant galaxies in the Fornax Cluster (308 ± 30 km s⁻¹), which according to D01 are virialized.¹¹ We also note that the value of the velocity dispersion that we infer when removing the PNe within $1 R_{\text{eff}}$ of NGC 1379 in the last bin is larger (330 ± 38 km s⁻¹) and almost reaches the σ_{cl} (empty black circle in Fig. 8).

5.3 PNe as tracers of the ICL

As a prelude to a more rigorous dynamical analysis, here we try to assess whether the difference in σ is caused by the fact that

¹¹ D01 find a difference of ~ 130 km s⁻¹ in velocity dispersion between the dwarf population and the giant population in the Fornax Cluster, which is consistent with the expected ratio of 2:1 for infalling and virialized, as predicted from Colless & Dunn (1996).

the PNe and the cluster galaxies have different density profiles ($\alpha_{\text{PNe}} = 3.0$ and $\alpha_{\text{cl}} = 2.0$) but still live in the same potential. We use equations (2), (3), and (4) in Napolitano et al. (2014), where they use: $\alpha \equiv -d\ln j/d\ln r$, $\gamma \equiv -d\ln \sigma^2/d\ln r$, and $\beta = 1 - \sigma_\theta^2/\sigma_r^2$ with j being the deprojection of the surface brightness and σ_θ and σ_r the azimuthal and radial components of the velocity dispersion in spherical coordinates. Following their prescriptions, we assume that α , β , and γ are constant with radius and we use the projected dispersion written as in Dekel et al. (2005)

$$\sigma_p(R) = A(\alpha, \gamma)B(\alpha, \gamma, \beta)V_0^2 R^{-\gamma}, \quad (2)$$

where

$$A(\alpha, \gamma) = \frac{1}{(\alpha + \gamma)} \frac{\Gamma[(\alpha + \gamma - 1)/2]}{\Gamma[(\alpha + \gamma)/2]} \frac{\Gamma[\alpha/2]}{\Gamma[(\alpha - 1)/2]} \quad (3)$$

and

$$B(\alpha, \gamma, \beta) = \frac{(\alpha + \gamma) - (\alpha + \gamma - 1)\beta}{(\alpha + \gamma) - 2\beta}. \quad (4)$$

Thus, at a given radius and under the working assumption that PNe and galaxies live in the same potential (i.e. V_0 is the same for both), the projected velocity dispersions of the two components depend only on the factors $A(\alpha, \gamma)$ and $B(\alpha, \gamma, \beta)$, which are functions of the anisotropy parameter (β), the dispersion slope (γ), and the 3D density slope (α). At a fixed radius $R = 1200$ arcsec, where the DM dominates, we assume that the intrinsic σ profiles are both flat ($\gamma_{\text{cl}} = \gamma_{\text{PNe}} = 0$), we take an isothermal profile for the cluster ($\alpha_{\text{cl}} = 2$) and measure the stellar density slope from FDS-I ($\alpha_{\text{PNe}} = 3$). As for the anisotropy parameter, we consider the fig. 2 in Mamon & Łokas (2005) where they found that cosmological simulations suggest that moving towards the outskirts, the galaxies prefer a radial anisotropy ($\beta = 0.3$ at r_{200}). Under these hypotheses, we infer $\sigma_{\text{cl}} = 1.3 \times \sigma_{\text{PNe}} = 380$ km s⁻¹, which is perfectly consistent with the value reported in D01 and indeed confirms that the PNe at that distance do not trace the spheroidal galaxy component but share the dynamics and the potential of the cluster galaxies. This demonstrates that indeed we are mapping the region of transition between the extended halo of NGC 1399 and the ICL, and we go well beyond it, clearly separating the two components. We will investigate this point further when performing a detailed dynamical modelling. In particular, we plan to clearly and more rigorously distinguish PNe that are bound to the haloes of the different galaxies to the ones that are, instead, tracing the ICL (ICPNe) and directly compare our results to those obtained with GCs of FVSS-I. We will focus on the PNe and GCs streams, linking their velocity dispersions with possible mechanisms able to generate the ICL (Gerhard et al. 2007; Murante et al. 2007; Rudick et al. 2009; Cui et al. 2014).

6 SUMMARY AND CONCLUSIONS

We have presented *the largest and most extended PNe kinematic catalogue ever obtained for the Fornax cluster*. We have obtained velocities of 1635 (1452 of which are newly measured) PNe in a region of about 50 arcmin \times 30 arcmin in the core of the Fornax cluster using a counter-dispersed slitless spectroscopic technique with data from FORS2 on the VLT. The catalogue also completes and spatially extends the sample of ~ 180 PNe in the halo of NGC 1399 and around NGC 1404 published by McNeil et al. (2010).

In this paper, we described our techniques and methods to identify and select PNe directly from the calibrated and registered images and we calculated their velocities. We use the PNe LOS velocity and velocity dispersion distributions to trace the ICL within 200 kpc of the Fornax Cluster core.

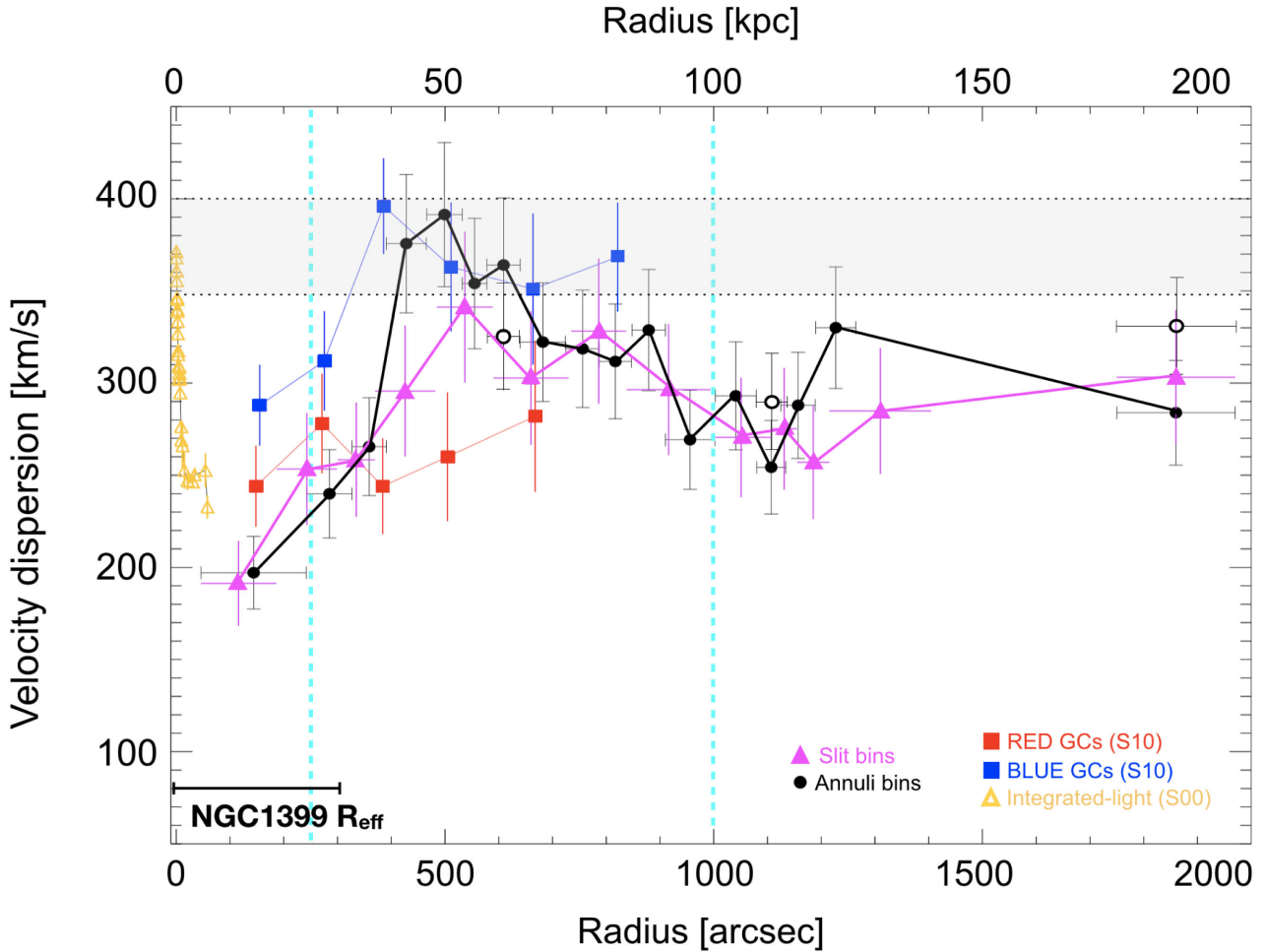


Figure 8. LOS velocity dispersion as a function of projected radius from the centre of NGC 1399. The velocity dispersion obtained from our final catalogue of PNe is plotted as filled black circles for the circular annuli calculation and as filled magenta triangles for the slit-bins. The three empty symbols correspond to the velocity dispersion inferred at a given radius removing the contribution of the PNe that are more likely associated with companion galaxies (all PNe within the effective radius of the corresponding galaxy). Red and blue GCs presented in Schuberth et al. (2010, S10 in the figure) are overplotted as red and blue squares, respectively, while the integrated-stellar-light profile from Saglia et al. (2000, S00 in the figure) is shown as yellow triangles (major axis). Finally, the grey shaded horizontal region represents the velocity dispersion of the Fornax cluster with its 1σ error.

In particular:

(i) We presented the final LOS velocity distribution of the full catalogue (this work + MN10) that comprises 1635 PNe.

(ii) From the 2D spatial distribution of velocities, we identified two streams of PNe. One low-velocity ‘group’, north of NGC 1399, which was already observed in MN10, and a high-velocity PNe ‘bridge’, connecting NGC 1399 and NGC 1387. Interestingly, a faint bridge in the same region was reported in Iodice et al. (2017) from deep photometry and in D’Abrusco et al. (2016) and Cantiello et al. (2018) from an overdensity in the GCs. From the FDS-I photometry results, we argue that this stream between NGC 1399 and NGC 1387 might indicate gravitational interaction that, however, did not happen in a very recent epoch. Indeed, the bridge is rather faint ($\mu_r \sim 28\text{--}29$ mag arcsec $^{-2}$) and above all does not look like a long tail but rather like a ‘diffuse light’ (e.g. FDS-I). Different is, for example, the case of the Virgo cluster, where many bright and much more extended streams and tidal tails have been observed (Mihos et al. 2005; Janowiecki et al. 2010; Capaccioli et al. 2015; Longobardi et al. 2015a).

(iii) We obtained the LOS velocity dispersion as a function of projected radius from the centre of NGC 1399 and we compared it to other velocity dispersion profiles from different kinematical tracers. We reach a distance of ~ 2000 arcsec, corresponding to roughly 200 kpc, from NGC 1399, extending by far all the previous velocity dispersion profiles ever presented.

(iv) In the overlap region, we found a good agreement with the LOS σ profiles obtained from GCs in Schuberth et al. (2010), in particular the velocity dispersions of the PNe within $R < 250$ arcsec from NGC 1399 agree well with the dispersion of the red GC population, which shares the dynamical history of the central galaxy itself. For radii $R \geq 400$ arcsec, instead, the PN velocity dispersion better fits with that measured for the blue cluster population and it is closer to the value of the velocity dispersion of the main Fornax Cluster reported in D01 (σ_{c1}), but it does not reach it. We noted that the measured $\sigma_{\text{PNe}} = 300$ km s $^{-1}$ is perfectly consistent with the velocity dispersion measured by Drinkwater, Gregg & Colless (2001) for the giant galaxies in the Fornax Cluster (308 ± 30 km s $^{-1}$), which are virialized. We therefore concluded that PNe at all radii probed in this study are virialized.

(v) We concluded, consequently, that after $R \sim 400$ arcsec the PNe are not bound to the NGC 1399 central halo. We also note that for intermediate radii ($400 \text{ arcsec} < R < 1000 \text{ arcsec}$) the presence of subhaloes causes a rise in the PNe velocity dispersion measurements.

(vi) Finally, at very large radii ($R \geq 1000$ arcsec), where the velocity dispersion profile is roughly flat, we found that the difference in velocity dispersion between σ_{cl} and the σ_{ICPNe} , where ICPNe have a σ which is $\sim 80 \text{ km s}^{-1}$ lower, is consistent with a dynamical scenario where the PNe share the same potential of the Cluster but have different density profiles from the galaxies ($\alpha_{PNe} = 3.0$ at $R = 1200$ arcsec, assuming $\alpha_{cl} = 2.0$).

This is the second paper of Fornax Cluster VLT Spectroscopic Survey (FVSS), which aims at studying in depth the assembly history of one of the nearby dense cluster environments. With the biggest PNe catalogue ever collected, spectra of GCs (Pota et al. 2018, Paper I, to be submitted) and UCDs with VIMOS@VLT, spectroscopic data on the three large ETGs (NGC 1399, NGC 1404, and NGC 1387) and seven dwarf galaxies (NGC 1396, FCC188, FCC211, FCC215, FCC222, FCC223, and FCC227; Spiniello et al. 2018) with MUSE@VLT, we have *the most complete and uniform collection of dynamical tracers to characterize the stellar population of the different systems and the DM profile deriving in this way the baryonic and dark mass distribution in the core of Fornax ($\sim 200 \text{ kpc}$) with a precision never reached until now.*

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