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Does the X-ray outflow quasar PDS 456 have a UV outflow at 0.3c?

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ABSTRACT

The quasar PDS 456 (at redshift ~ 0.184) has a prototype ultra-fast outflow (UFO) measured in X-rays. This outflow is highly ionized with relativistic speeds, large total column densities $\log N_{\text{H}}(\text{cm}^{-2}) > 23$, and large kinetic energies that could be important for feedback to the host galaxy. A UV spectrum of PDS 456 obtained with the Hubble Space Telescope in 2000 contains one well-measured broad absorption line (BAL) at $\sim 1346 \text{ \AA}$ (observed) that might be Ly α at $v \approx 0.06c$ or N V $\lambda 1240$ at $v \approx 0.08c$. However, we use photoionisation models and comparisons to other outflow quasars to show that these BAL identifications are problematic because other lines that should accompany them are not detected. We argue that the UV BAL is probably C IV $\lambda 1549$ at $v \approx 0.30c$. This would be the fastest UV outflow ever reported, but its speed is similar to the X-ray outflow and its appearance overall is similar to relativistic UV BALs observed in other quasars. The C IV BAL identification is also supported indirectly by the tentative detection of another broad C IV line at $v \approx 0.19c$. The high speeds suggest that the UV outflow originates with the X-ray UFO crudely 20 to 30 r_g from the central black hole. We speculate that the C IV BAL might form in dense clumps embedded in the X-ray UFO, requiring density enhancements of only $\gtrsim 0.4$ dex compared clumpy structures already inferred for the soft X-ray absorber in PDS 456. The C IV BAL might therefore be the first detection of low-ionisation clumps proposed previously to boost the opacities in UFOs for radiative driving.

Key words: line: formation – quasars: individual: PDS 456 – quasars: absorption lines – quasars: general

1 INTRODUCTION

Accretion disk outflows are an important part of the quasar phenomenon that might drive “feedback” to regulate black hole growth and host galaxy evolution (e.g., Di Matteo et al. 2005; Hopkins et al. 2008; Hopkins & Elvis 2010; Debuhr et al. 2012; Rupke & Veilleux 2013). The outflows are often studied in the rest-frame UV via blueshifted broad absorption lines (BALs) or their narrower cousins called “mini-BALs” (with a nominal boundary near full width at half minimum FWHM $\sim 2000 \text{ km s}^{-1}$, Weymann et al. 1991; Korista et al. 1993; Crenshaw et al. 2003; Hamann & Sabra 2004; Trump et al. 2006; Knigge et al. 2008; Gibson et al. 2009, and refs. therein). These features appear most often at moderate velocity shifts $v < 0.1c$, but relativistic BALs and mini-BALs

at $v \sim 0.1\text{--}0.2c$ have been measured in a small but growing number of quasars (Hamann et al. 1997a, 2013; Rodríguez Hidalgo 2008; Rodríguez Hidalgo et al. 2011; Rogerson et al. 2016).

The basic physical properties for these outflows can be difficult to determine due to limited wavelength coverage and line saturation that is masked by partial line-of-sight covering of the background light source(s). The strongest measured UV lines¹ are typically C IV $\lambda 1549$, N V $\lambda 1240$, and O VI 1034, indicating moderate to high degrees of ionisation. The maximum ionisations are not known, but there might be a wide range. O VI absorption tends to be as strong or stronger than C IV (Section 3.4 below, also Baskin et al.

¹ Throughout this paper we treat unresolved doublets as single lines, such that, for example, C IV $\lambda 1548, 1551$ becomes C IV $\lambda 1549$ with a summed oscillator strength (for column density estimates).

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2013; Moravec et al. 2017, Herbst et al., in prep.) and higher-ionisation lines such as Ne VIII $\lambda 774$ and Mg X $\lambda 615$ have been observed in a few cases with suitable spectral coverage (e.g., Hamann et al. 1997b; Telfer et al. 1998; Arav et al. 2001). Measurements of the low-abundance line P V $\lambda 1121$ help to overcome the saturation issues to reveal generally large total hydrogen column densities, $\log N_{\text{H}}(\text{cm}^{-2}) \gtrsim 22.3$, across a wide range observed BAL strengths (Hamann 1998; Leighly et al. 2009, 2011; Borguet et al. 2012; Capellupo et al. 2014, 2017; Moravec et al. 2017, but see also Arav et al. 2001).

X-ray observations have revealed another variety of ultra-fast outflows (UFOs) that reach relativistic speeds in luminous quasars (Chartas et al. 2002, 2009; Reeves et al. 2009; Gofford et al. 2013, 2015; Tombesi et al. 2010, 2013, and refs. therein). UFOs are also challenging to study because they are highly variable and highly-ionised to the point where the only strong absorption features appear at X-ray wavelengths. They are often characterized by Fe K-shell absorption features with derived total column densities in the range $\log N_{\text{H}}(\text{cm}^{-2}) \sim 22$ to 24 (see refs. above). UFOs also appear to have generally very large kinetic energies, sufficient to drive feedback effects in the quasar host galaxies (Tombesi et al. 2012, 2013; Gofford et al. 2014, 2015; Reeves et al. 2014).

The quasar PDS 456 (at redshift $z_e \approx 0.184$) has the best-studied example of a powerful, relativistic X-ray UFO (e.g., Reeves et al. 2003, 2009, 2014, 2016; Nardini et al. 2015; Gofford et al. 2014; Matzeu et al. 2017). It is the most luminous quasar in the local universe, with bolometric luminosity $L \sim 10^{47}$ ergs s $^{-1}$ and estimated black hole mass $\sim (1-2) \times 10^9 M_{\odot}$ that together indicate an accretion rate relative to Eddington that is $L/L_{\text{Edd}} \gtrsim 0.3$ and perhaps near unity (Nardini et al. 2015). The X-ray absorber is complex and highly variable, with two main components. The main component measured via Fe K-shell absorption has speeds in the range $v \sim 0.25$ – $0.34c$, very high degrees of ionisation featuring Fe XXV and Fe XXVI, and large total column densities $\log N_{\text{H}}(\text{cm}^{-2}) \gtrsim 23$ (Reeves et al. 2014; Gofford et al. 2014, 2015; Nardini et al. 2015). Its radial distance from the black hole is estimated at a few hundred gravitational radii (based on absorber variability, e.g., Nardini et al. 2015). The kinetic power of this outflow is also remarkably large, crudely ~ 20 percent of the bolometric luminosity, which is well above the threshold needed for feedback to the host galaxy (Gofford et al. 2014; Nardini et al. 2015, and refs. therein). The second outflow component measured in soft X-rays has somewhat lower speeds, $v \sim 0.17$ – $0.27c$, lower ionisations, and column densities in the range $\log N_{\text{H}}(\text{cm}^{-2}) \sim 22$ to 23 (Reeves et al. 2016). There is evidence for time-variable covering fractions in this absorber that might be indicative of small dense clumps or substructures embedded in the overall X-ray outflow (Matzeu et al. 2016).

PDS 456 also has signatures of outflow in the UV. A UV spectrum obtained in 2000 by O’Brien et al. (2005) using the Space Telescope Imaging Spectrometer (STIS) on board the Hubble Space Telescope (HST) revealed a highly blueshifted C IV broad emission line at speeds near $v \sim 5000$ km s $^{-1}$ plus a single BAL that might plausibly be Ly α or N V at $v \sim 0.06c$ or $\sim 0.08c$, respectively. These UV features provide further evidence for exotic mass loss from PDS 456 across a wide range of spatial scales, from the X-ray UFO

that is believed to originate near the black hole at radii of ~ 20 – $30 r_g$ (gravitational radii, corresponding to ~ 0.001 – 0.0015 pc Nardini et al. 2015; Matzeu et al. 2017) to an outflow-dominated C IV broad emission-line region that we place crudely at ~ 0.3 pc (based on scaling relations with luminosity, Kaspi et al. 2005, for a $10^9 M_{\odot}$ black hole). The UV BAL is an important component to this outflow picture, but its measured speed and physical nature depend critically on the line identification.

In this paper, we reexamine the UV BAL in PDS 456 with the main result that it is likely to be C IV $\lambda 1549$ at $v \approx 0.30c$. This would be the fastest UV outflow line ever reported but similar in speed to the X-ray UFO in this quasar. Throughout this paper, we adopt a redshift for PDS 456 of $z_e = 0.18375$ based on the emission line [Fe II] $1.6435 \mu\text{m}$ measured by Simpson et al. (1999). We describe two archival HST spectra of PDS 456 in Section 2 below. Section 3 presents our analysis and measurements of the observed BAL, comparisons to photoionisation models, comparisons to UV BALs in other quasars, and a discussion of the plausible BAL identifications. Section 4 presents a summary and discussion of the results.

2 HST SPECTRA

Figure 1 shows the spectrum of PDS 456 obtained by O’Brien et al. (2005) in 2000 using HST-STIS with the G140L and G230L gratings. We obtained this spectrum and another HST spectrum measured in 2014 (described below) from the Mikulski Archive for Space Telescopes (MAST). We present them here without further processing. The STIS G140L grating provided wavelength coverage from $1137 \text{ \AA}$ to $1715 \text{ \AA}$ at resolutions ranging from ~ 310 km s $^{-1}$ to ~ 200 km s $^{-1}$, while STIS G230L covered $1580 \text{ \AA}$ to $3148 \text{ \AA}$ at resolutions from ~ 640 km s $^{-1}$ to ~ 300 km s $^{-1}$ (see O’Brien et al. 2005 for more details). The spectrum plotted in Figure 1 is combined from the two gratings by masking out the extreme ends of the wavelength coverage to avoid excessive noise and then calculating variance-weighted average fluxes at the remaining wavelengths of overlap.

The UV spectrum of PDS 456 is substantially reddened by dust in our Galaxy (due to the quasar’s sky position near the Galactic plane). Previous studies based on visible spectra and photometry indicate that the visual extinction towards PDS 456 is $A_V \sim 1.5$ magnitudes, corresponding to selective extinction $E(B-V) \sim 0.48$ for a standard Galactic reddening curve (Torres et al. 1997; Simpson et al. 1999). We obtain an independent estimate of the reddening by fitting the UV continuum with a single power law modified by the Galactic extinction curve from Cardelli et al. (1989) with $R_V = 3.1$. The fit is constrained by the median flux in wavelength windows between $\sim 1170 \text{ \AA}$ and $\sim 2400 \text{ \AA}$ that avoid strong emission and absorption lines. The dashed red curve in Figure 1 shows our preferred fit using $E(B-V) = 0.45$ and power law index $\alpha_{\lambda} = -1.68$ (for $F_{\lambda} \propto \lambda^{\alpha_{\lambda}}$). An important feature of Galactic extinction at these wavelengths is the “bump” at $\sim 2175 \text{ \AA}$. We experimented with different fit parameters (e.g., the dotted red curves in Figure 1), but the values of $E(B-V)$ and the power law slope are well constrained² by the

² We note, however, that the observed spectrum is poorly fit at

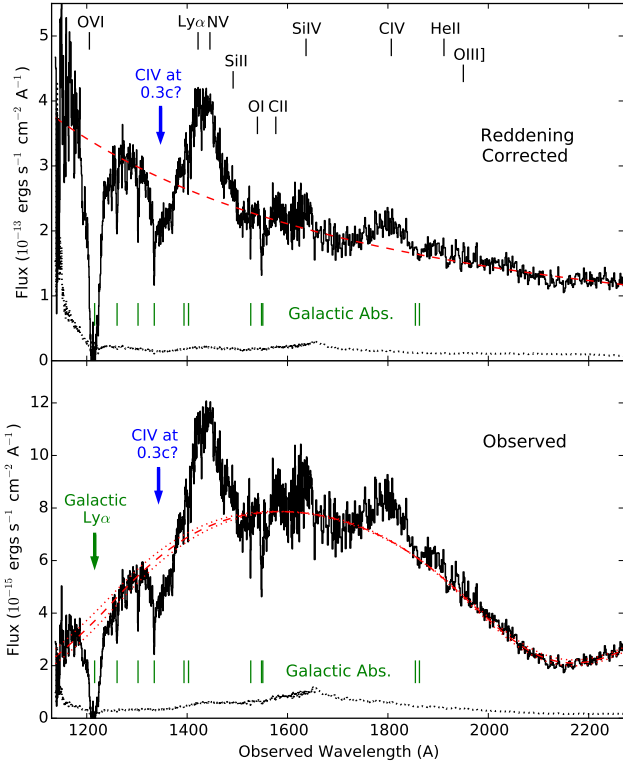


Figure 1. HST STIS spectrum of PDS 456 from 2000 plotted as observed (bottom panel) and reddening-corrected using a Galactic extinction curve with $E(B - V) = 0.45$ (top panel). The grey dotted curves are the corresponding error spectra. The red dashed curves are an approximate fit to the continuum using a power law with index $\alpha_\lambda = -1.68$ shown with (bottom panel) and without (top) Galactic reddening. The red dotted curves in the bottom panel show the same power law reddened alternatively by $E(B - V) = 0.42$ and 0.48 . The BAL we attribute to CIV at $v \approx 0.30c$ is labeled above the spectrum at $1346 \text{ \AA}$. Some broad emission-line wavelengths are marked across the top. Galactic absorption lines (including strong damped Ly α at $1216 \text{ \AA}$) are marked by green vertical lines near the bottom. See Section 2.

data because Galactic reddening curves with a strong $2175 \text{ \AA}$ bump suppress the flux at both ends of the spectral coverage shown in Figure 1. The top panel in this figure shows the reddening-corrected spectrum of PDS 456 together with the best-fit power law.

Figure 2 shows a previously-unpublished spectrum of PDS 456 obtained in 2014 with HST-COS using the G140L grating (PI: O’Brien). This spectrum (shown by the red curve) is plotted on top of the HST-STIS spectrum (black curve) from Figure 1 with no reddening corrections. It is combined from separate exposures totaling 5186 seconds. The spectral resolution ranges from roughly 200 km s^{-1} to 100 km s^{-1} from blue to red across the wavelengths shown

wavelengths $\gtrsim 2400 \text{ \AA}$ (observed). This might be caused by reddening in the quasar environment, a break in the intrinsic/emitted spectral slope, and/or blended Fe II emission lines that are known to be strong in PDS 456 (Simpson et al. 1999). We do not investigate this further because our goal is simply to define a continuum for analysis of the outflow lines at shorter wavelengths (Section 3).

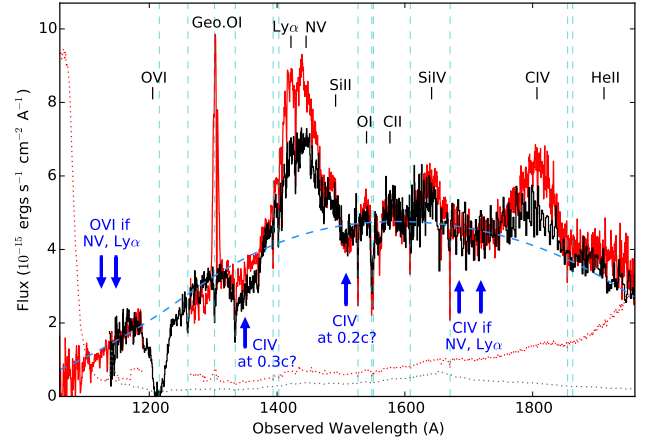


Figure 2. HST spectra of PDS 456 obtained with STIS in 2000 (black curve) and COS in 2014 (red) plotted at observed wavelengths, uncorrected for reddening. The red and grey dotted curves are the corresponding error spectra (1σ uncertainties per pixel). The STIS spectrum is scaled vertically by a factor ~ 0.60 to match the COS spectrum approximately in the continuum. The dashed blue curve shows our fit to the STIS spectrum from Figure 1. The BAL we attribute to CIV $\lambda 1549$ at $0.30c$ is labeled below the spectrum at $\sim 1346 \text{ \AA}$. Also marked are the expected positions of OVI $\lambda 1034$ and CIV BALs if the observed BAL is attributed, instead, to NV $\lambda 1240$ at $0.08c$ or Ly α at $0.06c$ (e.g., “if NV, Ly α ”). We tentatively identify another weak BAL at $\sim 1513 \text{ \AA}$ as CIV at $0.19c$. Dashed blue vertical lines mark the positions of Galactic absorption features. The narrow emission spike at $1304 \text{ \AA}$ is geocoronal OI. See Figure 1 and Section 2 for additional notes.

in Figure 2. A gap in the COS wavelength coverage from $\sim 1170 \text{ \AA}$ to $\sim 1265 \text{ \AA}$ avoids Galactic/geocoronal Ly α absorption/emission. The blue arrows in Figures 1 and 2 mark the positions of observed or expected absorption lines discussed in Section 3 below

3 ANALYSIS

3.1 The UV BAL at $1346 \text{ \AA}$

The HST-STIS spectrum in Figure 1 clearly shows a BAL at $\sim 1346 \text{ \AA}$ (observed), as reported by O’Brien et al. (2005). Figure 3 plots this spectrum again after normalizing by the reddened power law in Figures 1 and 2. We fit the BAL in this normalized spectrum using a simple Gaussian optical depth profile of the form

$$\tau_v = \tau_o e^{-v^2/b^2} \quad (1)$$

where τ_o is the line-center optical depth, v is the velocity shift from line center, and b is the doppler parameter that sets the line width. The BAL fit shown by the magenta curve in Figure 3 yields $\tau_o = 0.35 \pm 0.01$, $b = 5135 \pm 203 \text{ km s}^{-1}$, and observed line-center wavelength $\lambda_o = 1345.7 \pm 0.6 \text{ \AA}$ (where the errors quoted are 1σ uncertainties returned by the line-fitting routine). The full width at half minimum of this fitted profile is $\text{FWHM} \approx 8550 \pm 340 \text{ km s}^{-1}$. In the quasar frame, the fit has rest equivalent width $\text{REW} = 10.8 \pm 0.6 \text{ \AA}$.

Our study concerns identification of the $\sim 1346 \text{ \AA}$ BAL trough. One possibility favored by O’Brien et al. (2005) is

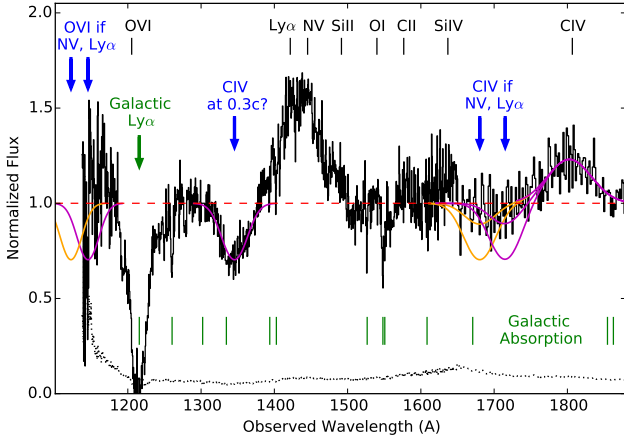


Figure 3. HST-STIS spectrum of PDS 456 in the observed frame, normalized by the continuum fit shown in Figure 1. The smooth magenta and orange curves show our fit to the observed BAL at 1346 Å and then that fit transposed to the wavelengths of O VI and C IV assuming the measured BAL is Ly α (magenta) or N V (orange). The weaker profiles drawn at the C IV positions show the transposed fit at 1/3 its measured strength to illustrate approximate upper limits to the actual absorption there. The absence of absorption at the transposed fit positions argues against the Ly α and N V identifications for the observed BAL (Section 3.5 and Section 3.6). Other labels are the same as Figures 1 and 2. Narrow Galactic absorption lines of O I λ 1304 and C II λ 1335 have been removed by interpolation for clarity.

Ly α 1216 at velocity shift $v \approx 0.06c$ (18,000 km s $^{-1}$). Another is N V λ 1240 at $v \approx 0.08c$. These identifications seem reasonable because the observed BAL is on the blue side of the Ly α –N V emission-line blend and the inferred velocities are in a normal range for BAL outflows in other quasars. However, we will argue below (Section 3.5 and Section 3.6) that the Ly α or N V identifications are problematic because they are not accompanied by C IV and O VI λ 1034 BALs at the same velocity shifts. The absence of these other lines is illustrated by the other orange and magenta curves in Figure 3, which show our fit to the observed BAL transposed to the positions expected for accompanying C IV and O VI BALs. At the C IV positions, the transposed BAL fit is shown twice – once at full strength and once at 1/3 of its measured strength at 1346 Å. We estimate that the 1/3 scalings represent approximate upper limits to C IV BALs that might accompany a N V or Ly α BAL at 1346 Å. We will argue below that the absence O VI and C IV absorption at these wavelengths suggests that the BAL at $\sim$ 1346 Å is actually C IV λ 1549 at $v \approx 0.30c$.

We test these line identifications using photoionisation models (Section 3.3) and comparisons to BALs/mini-BALs observed in other quasars (Section 3.4). First we estimate ionic column densities for each possible identification. These estimates follow from the fitted Gaussian optical depth profile by

$$N_i = \frac{m_e c}{\sqrt{\pi} e^2} \frac{b \tau_o}{f \lambda_r} \quad (2)$$

where N_i is the ionic column density, f and λ_r are the oscillator strength and rest-frame wavelength of the transition, and we assume implicitly that the ground-state column density equals the ionic column density. Thus we find that, if the

observed BAL is Ly α , it corresponds to a neutral hydrogen column density of $\log N_{\text{HI}}(\text{cm}^{-2}) \approx 15.35$. If the BAL is N V or C IV, then the column density is $\log N(\text{N V})(\text{cm}^{-2}) \approx 15.62$ or $\log N(\text{C IV})(\text{cm}^{-2}) \approx 15.44$, respectively. The uncertainties in these estimates depend mainly on the continuum placement used for the BAL profile fit. Experiments with different continua suggest that the maximum uncertainties are roughly 0.1 dex.

It is important to note that the line optical depth and column densities derived above are only lower limits because BAL outflows often exhibit saturation with partial line-of-sight covering of the background light source. This can lead to weak/shallow absorption-line troughs even if the optical depths are large (e.g., Hamann 1998; Hamann & Sabra 2004; Arav et al. 2005; Capellupo et al. 2017; Moravec et al. 2017, and refs. therein). In one recent study, Herbst et al. (in prep.) used median composite spectra of BAL quasars from the Baryon Oscillation Sky Survey (BOSS, Dawson et al. 2013; Ross et al. 2012, part of the Sloan Digital Sky Survey-III, SDSS-III, Eisenstein et al. 2011) to show that the low-abundance doublet P V λ 1118,1128 is often present and that widely-separated doublets like P V, Si IV λ 1393,1402, and O VI λ 1032,1038 have $\sim$ 1:1 doublet ratios across a wide range of BAL strengths. These results indicate that all of the commonly measured BALs of abundant ions in quasar spectra, e.g., C IV λ 1549, N V λ 1240 and O VI λ 1034, are *typically* saturated and the observed depths of the BAL troughs are controlled mainly by the line-of-sight covering fractions.

3.2 Other UV Outflow Features

In addition to the UV BAL at 1346 Å, there are two other features in the 2000 HST-STIS spectrum that point to extreme outflows in PDS 456. First, as noted previously by O’Brien et al. (2005), the broad C IV emission line is extremely blueshifted. Our simple Gaussian fit to this feature shown in Figure 3 indicates that the line centroid is blueshifted by 5200 ± 450 km s $^{-1}$ (consistent with the O’Brien et al. 2005 measurement). This fit also yields a rest equivalent width of $\text{REW} = 14.7 \pm 1.8$ Å and a large velocity width of $\text{FWHM} = 11770 \pm 1060$ km s $^{-1}$ (not corrected for the $\sim$ 500 km s $^{-1}$ doublet separation in C IV). The low-ionisation emission lines O I λ 1304 and C II λ 1335 are poorly measured in this spectrum compared to C IV, but they are clearly less blueshifted. We crudely estimate their centroid blueshifts by visual inspection to be 700 ± 200 km s $^{-1}$.

The large C IV emission-line blueshift identifies an outflow-dominated broad emission-line region with an unusually high outflow speed. At $v \approx 5200 \pm 450$ km s $^{-1}$, the C IV blueshift is outside of the range of values reported by Coatman et al. (2016) for a sample of $\sim$ 31,157 luminous SDSS quasars (see also Richards et al. 2011). However, we can crudely estimate by extrapolation from their Figure 1 that the C IV blueshift in PDS 456 is in the upper $\sim$ 0.1 percent of luminous quasars. To our knowledge, it is matched in the literature only by a few luminous extremely red quasars (ERQs) discovered in SDSS-III/BOSS to have prodigious outflows include BALs in the UV and highly blueshifted [OIII] λ 5007 emission lines (Hamann et al. 2017; Zakamska et al. 2016, Perrotta et al., in prep.). We will present a more thorough analysis of the C IV blueshift in a future paper. Here we note simply that large blueshifts are known

to correlate with small emission-line REWs and intrinsically weak X-ray emissions (compared to other quasars/AGN at similar luminosities, e.g., Leighly 2004; Leighly et al. 2007; Wu et al. 2011, 2012; Richards et al. 2011; Luo et al. 2015). In addition, the work by Coatman et al. (2016) supports speculation in the ERQ studies that large blueshifts and other prominent outflow features are related to high accretion rates (high Eddington ratios) in the quasars.

Another tentative outflow feature in the 2000 HST-STIS spectrum is a weak BAL marginally detected at ~ 1513 Å (observed, see Figures 2 and 3). The reality of this feature depends on the continuum placement, but it is clearly below our best guess at the continuum and might be stronger in the HST-COS 2014 spectrum (Figure 2). There is also tentative absorption near ~ 1560 Å, but this feature is blended with telluric C IV absorption at 1549 Å and it appears to be much narrower than the BAL candidate at ~ 1513 Å. Fits to the BAL candidate at 1513 Å are poorly constrained, but multiple trials indicate a central wavelength of $\lambda_o \approx 1513 \pm 2$ Å, FWHM $\approx 3800 \pm 1000$ km s $^{-1}$, and REW $\approx 1.1 \pm 0.5$ Å in the quasar frame. If this BAL is real, the only plausible identification is C IV at $v \sim 0.19c$ (The leading alternative, S IV $\lambda 1398$ at $v \sim 0.09c$ is ruled out by the absence of C IV absorption at the same velocity shift. See Rodríguez Hidalgo et al. 2011; Hamann et al. 2013).

3.3 Cloudy Simulations

We use the photoionisation and spectral synthesis code Cloudy version 17.00 (Ferland et al. 2013, 2017) to predict absorption-line strengths for different physical conditions that might produce the observed BAL at ~ 1346 Å. The calculations assume twice solar metallicity and a fixed hydrogen density of $n_H = 10^7$ cm $^{-3}$. The specific density has no bearing on our results (e.g., Hamann 1997). The metallicity can affect mainly the metal line strengths relative to the Lyman lines. We choose twice-solar metallicity to be crudely consistent with other studies of the outflows and broad emission-line environments of luminous quasars (Hamann & Ferland 1999; Dietrich et al. 2002; Warner et al. 2004; Nagao et al. 2006; Gabel et al. 2005a, e.g.). The clouds are irradiated by a standard quasar ionising spectrum consistent with measurements of PDS 456 (e.g., Nardini et al. 2015). This spectrum is defined by power laws across optical-UV and X-ray wavelengths with slopes of $\alpha_{uv} = -0.5$ and $\alpha_x = -1.3$, respectively (for $f_\nu \propto \nu^\alpha$). The two power laws are joined smoothly in the far-UV by an exponential Wien function with temperature $T = 250,000$ K. The relative strengths of the UV and X-ray spectral segments are set by a two-point power-law index between 2500 Å and 2 keV equal to $\alpha_{ox} = -1.8$ (Stratova et al. 2005; Steffen et al. 2006). This spectrum is defined in Cloudy by the command: AGN T=250000K, a(ox)=-1.8, a(uv)=-0.5, a(x)=-1.3. The flux of hydrogen ionising radiation incident on the model clouds is set by the ionisation parameter,

$$U \equiv \frac{Q_H}{4\pi c R^2 n_H} \quad (3)$$

where R is the distance from the quasar light source and Q_H is the total emitted luminosity of hydrogen-ionising photons (#/s). Other aspects of the calculations are described below.

3.4 A Comparison Outflow Sample

Here we construct a sample of outflow quasars from SDSS-III/BOSS that can be useful analogues to test the BAL identification in PDS 456. We specifically use data from the BOSS quasar catalog for data release 12 (DR12, Pâris et al. 2017) to select quasars with the following properties: 1) We require that the BOSS spectra have signal-to-noise ratios $\text{SNR} > 5$ at 1700 Å (quasar frame). 2) The emission-line redshifts must be $z_e > 2.4$ to place the outflow lines of C IV, N V and Ly α within the BOSS wavelength coverage at speeds up to at least $\sim 0.1c$. Nearly $2/3$ of the selected quasars have $z_e > 2.6$ such that O VI $\lambda 1034$ is also covered. 3) The quasars must have broad C IV outflow lines (BALs or mini-BALs) characterized by balnicity index $\text{BI} < 3000$ km s $^{-1}$ and absorption index $\text{AI} > 1000$ km s $^{-1}$, both measured at $>4\sigma$ confidence (see Pâris et al. 2017, and refs. therein for definitions of BI and AI). These parameter limits are designed to exclude strong/deep C IV BALs (with large BI) as well as very narrow C IV lines with FWHMs $<$ several hundred km s $^{-1}$ (with small AI) that do not resemble the observed feature in PDS 456. Note that broad outflow lines are recorded in the BOSS quasar catalog only at velocity shifts $v \lesssim 0.1c$. 4) The *minimum* velocity recorded on the red side of the BAL/mini-BAL troughs is $v > 6000$ km s $^{-1}$ (as specified by the AI integration limit $\text{vmin_civ_450} > 6000$). This avoids broad lines near $v \sim 0$ that can have different profiles and other properties compared to higher-velocity BALs and mini-BALs. The specific value of 6000 km s $^{-1}$ is a compromise that requires a substantial velocity shift while still maintaining a large sample size. Finally, 5) we reject a small fraction of the quasars (~ 10 percent) that have damped Ly α absorption (DLAs) or strong sub-DLAs at the wavelengths of interest (based on our own visual inspections of the spectra).

The final sample includes 641 quasars at median redshift $\langle z_e \rangle = 2.84$ and median absolute i -band magnitude $\langle M_i \rangle = -26.8$. We visually inspect all of the quasar spectra to assess the strengths of prominent outflow lines for comparison to PDS 456. At these redshifts and velocity shifts, the N V, Ly α , and O VI outflow lines are inevitably blended with unrelated absorption features in the Ly α forest. However, in most cases, the outflow lines are clearly discernible because they are much broader than the forest features. Figure 4 shows spectra of four quasars we find to have the strongest well-measured N V outflow lines relative to C IV in our BOSS sample. These extreme cases with large N V/C IV are useful for our discussions of the BAL identification in PDS 456 (e.g., Section 3.6).

It is also useful to consider the median outflow line properties in the BOSS sample. Thus we construct a median composite spectrum in the outflow absorber frame for all 641 quasars. This composite “averages out” the Ly α forest contamination to reveal the typical outflow line strengths and profiles across the sample. The result is shown in Figure 5. We derive an absorber redshift for each quasar from the wavelength of minimum flux between the limits of the C IV AI integration (e.g., between vmin_civ_450 and vmax_civ_450) in smoothed versions of the BOSS spectra. The composite is then constructed by normalizing each spectrum to a continuum flux near 1700 Å and then shifting to the absorber frame before calculating the median. We

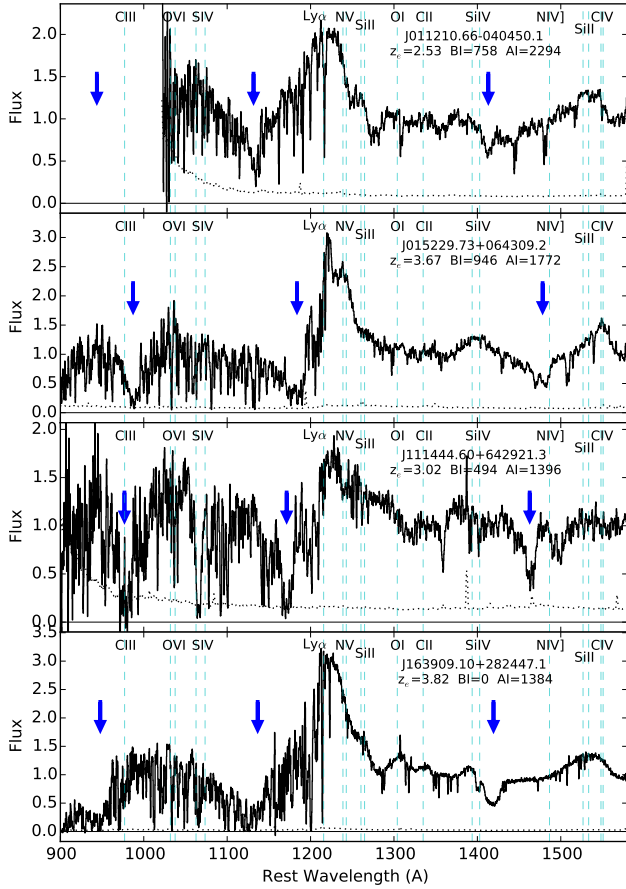


Figure 4. Normalized spectra in the quasar frame of four outflow quasars from our BOSS sample with the strongest well-measured N V absorption lines relative to C IV. Blue arrows mark the broad outflow absorption lines of O VI, NV, and C IV from left to right, respectively. NV and O VI suffer from varying amounts of contamination in the Ly α forest (e.g., in the bottom panel there is Lyman limit absorption at wavelengths $\lesssim 912$ Å and the deepest part of the broad NV trough is enhanced and offset from the absorption minima in C IV and O VI due to Ly α forest contamination). Prominent broad emission lines are labeled across the top. Also listed are the quasar names, emission-line redshifts, z_e , and absorption indices, BI and AI, from the BOSS DR12 quasar catalog. See Section 3.4

remove broad emission lines from the final outflow quasar composite spectrum by dividing by another composite spectrum of ~ 7700 BOSS quasars with similar redshifts and absolute magnitudes (M_i) but *without* broad outflow lines in their spectra. The redshifts of these non-outflow quasars are randomly displaced to match the distribution of shifts used to place the outflow quasars in the absorber frame. A simple division then yields the final composite spectrum shown in Figure 5. See Baskin et al. (2013), Baskin et al. (2015), and Herbst et al. (in prep.) for more discussion of this procedure.

The main result from Figure 5 is that the O VI $\lambda 1034$, NV $\lambda 1240$, and C IV $\lambda 1549$ outflow lines have roughly similar strengths while Ly α (in the blue wing of NV) is substantially weaker. This represents the typical situation for luminous quasars with BALs/mini-BALs crudely similar to the BAL in PDS 456. Another interesting result in Figure 5 is the significant presence of P V $\lambda 1121$ absorption and the resolved

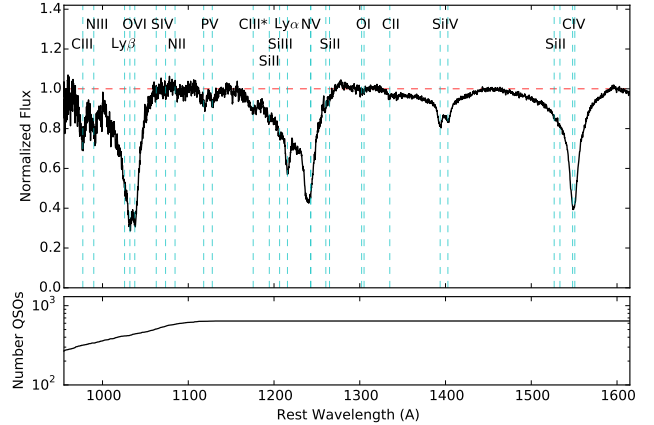


Figure 5. *Top panel:* Normalized composite spectrum in the absorber frame for 641 BOSS quasars selected to have C IV BALs/mini-BALs with moderate strengths and minimum trough velocities > 6000 km s $^{-1}$ (Section 3.4). The dashed blue vertical lines with labels across the top mark absorption lines that are or might be present. The dashed red horizontal line marks the unabsorbed continuum. *Bottom panel:* Number of quasars contributing to the composite at each wavelength.

doublets with $\sim 1:1$ ratios in O VI, P V, and Si IV. This supports the claim by Herbst et al. (in prep.) that BAL/mini-BAL outflows often have large total column densities and saturated absorption in all of the prominent lines.

3.5 Problems with the Ly α Identification

Figure 6 shows theoretical line-center optical depths from our Cloudy simulations (Section 3.2) for lines that should accompany a Ly α BAL at 1346 Å in PDS 456. In particular, the solid curves show the optical depths for lines with $b = 5135$ km s $^{-1}$ formed in clouds with a range of ionisation parameters, U , but with neutral hydrogen column densities held fixed at the derived value of $\log N_{\text{HI}}(\text{cm}^{-2}) = 15.35$ (Section 3.1). These predictions should be considered lower limits because N_{HI} derived from the data is a lower limit. Holding N_{HI} constant in the calculations leads to model clouds that have larger total hydrogen column densities, N_{H} , at larger U . The $\log N_{\text{H}}$ values for each $\log U$ are shown across the top of the figure.

The shaded regions around the curves for C IV, NV, and O VI in Figure 6 illustrate their dependence on the specific values of N_{HI} and the turbulent velocity, v_{turb} , used in the calculations. The solid curves for these lines (and the solid and dashed curves for C III $\lambda 1977$, Si III $\lambda 1206$, and Si IV $\lambda 1398$) correspond to fiducial parameters $\log N_{\text{HI}}(\text{cm}^{-2}) = 15.35$ and $v_{\text{turb}} = 1000$ km s $^{-1}$. In the top panel, the shaded regions show the range of optical depths corresponding to the maximum uncertainty of ± 0.1 dex in our $\log N_{\text{HI}}$ measurement (see Section 3.1, where the smaller/larger N_{HI} value sets the lower/upper envelope to the shaded curves). The optical depth changes in response to these changes in N_{HI} are non-linear at large U (and large N_{H}) because of radiative shielding in the far-UV that affects each ion differently. These shielding effects maintain significant optical depths in the moderate ions C IV, NV, and O VI at a fixed N_{HI} and large U . In other calculations (not shown), we determine that the threshold for important shielding effects at large U is roughly

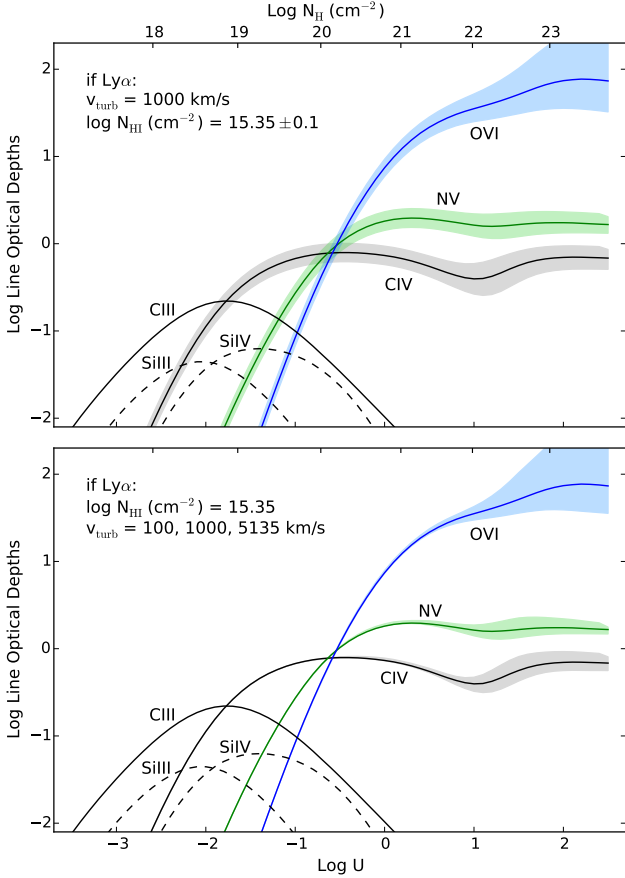


Figure 6. Line-center optical depths versus ionisation parameter, $\log U$ (bottom axis), and total hydrogen column density, $\log N_{\text{H}}$ (top axis) predicted by Cloudy simulations to accompany a Ly α BAL like the observed feature in PDS 456 (with $\tau_o = 0.35$ and $b = 5135 \text{ km s}^{-1}$, Section 3.1). The dashed and dark solid curves are the same in both panels, representing model clouds with $\log N_{\text{H}}(\text{cm}^{-2}) = 15.35$ and $v_{\text{turb}} = 1000 \text{ km s}^{-1}$. Shaded regions around the C IV, NV, and O VI curves show the ranges of optical depths that result from varying $\log N_{\text{H}}$ by ± 0.1 (top panel) or changing the turbulence from v_{turb} from 100 km s^{-1} to its maximum value, 5135 km s^{-1} , set by the observed BAL width (bottom panel). These predictions show that the Ly α BAL identification would require $\log U < -2$ and uniquely low column densities, $\log N_{\text{H}}(\text{cm}^{-2}) < 18.7$, for a BAL outflow (see Section 3.5).

$\log N_{\text{H}}(\text{cm}^{-2}) \gtrsim 14.8$. This is 0.55 dex lower than our lower limit to the H I column density inferred from the observed BAL (Section 3.1), and therefore the predictions in Figure 6 that result from shielding at large U should be applicable.

Another parameter that can affect radiative shielding is the velocity dispersion inside the cloud, which we characterize by the turbulence velocity v_{turb} . Large internal velocity dispersions can enhance the shielding by blending together numerous absorption lines in the far-UV, which in turn affects the ionisation structure deep in the cloud. The shaded regions around the C IV, NV, and O VI curves in the bottom panel of Figure 6 depict a range of results corresponding to $v_{\text{turb}} = 100 \text{ km s}^{-1}$, which yields negligible line shielding (the bottom envelope of the shaded regions), up to $v_{\text{turb}} = 5135 \text{ km s}^{-1}$, which yields the maximum line shielding allowed by the observed BAL width (the upper envelope of the shaded

regions). We adopt an intermediate value $v_{\text{turb}} = 1000 \text{ km s}^{-1}$ as our fiducial case (here and in all of our calculations below) because it yields conservatively small amounts of line shielding and, therefore, conservatively small predicted line optical depths in C IV, NV, and O VI at large U .

The main result from Figure 6 is that a Ly α BAL should be accompanied by significant broad C IV and O VI absorption lines over a wide range of normal BAL outflow physical conditions. It is important to keep in mind that the predicted strengths of the C IV and O VI lines are lower limits based on a lower limit on N_{H} that follows from the assumption of no partial covering. Smaller values of N_{H} are not allowed because they cannot produce the observed BAL attributed to Ly α . Larger N_{H} requiring larger N_{H} (at a given U) would produce stronger C IV and O VI absorption. The absence of these lines in the PDS 456 spectrum (Figures 2 and 3) means that the only conditions consistent with the Ly α BAL identification are low degrees of ionisation, $\log U < -2$, and very low total column densities, $\log N_{\text{H}}(\text{cm}^{-2}) < 18.7$ that are $\gtrsim 3.5$ orders of magnitude lower than expected for BAL outflows. For comparison, analysis of composite BAL quasar spectra by Herbst et al. (in prep.) suggests that *typical* BAL outflows have $\log N_{\text{H}}(\text{cm}^{-2}) \gtrsim 22.2$ (see also Sections 1 and 3.4).

Another problem for the Ly α identification is that a Ly α -only BAL system would be, to our knowledge, unprecedented among observed BAL outflows. Figure 3 shows that the C IV and O VI BALs expected to accompany Ly α in PDS 456 would need to be < 33 percent the strength of Ly α to avoid detection. This is very different from the typical situation illustrated by Figure 5 and by other BAL composite spectra (Baskin et al. 2013, Herbst et al., in prep.) where the C IV and O VI lines that are ~ 2 to ~ 3 times *stronger* than Ly α . BAL surveys based on C IV BAL detections (e.g., Trump et al. 2006; Gibson et al. 2009; Pâris et al. 2017) are naturally biased toward stronger C IV lines and they cannot detect Ly α -only BAL systems if they exist. However, our visual inspections of the 641 outflow quasar spectra in our BOSS sample (Section 3.4) do not find any instances of C IV BALs weaker than, or even comparable to, the corresponding Ly α absorption.

A third argument against the Ly α BAL identification stems from the weakening of the observed BAL between the HST observations in 2000 and 2014 (Figure 2). Variability in BALs and other outflow lines can be attributed to outflow clouds moving across our lines of sight or to changes in the ionisation caused by changes in the incident ionising flux (cf., Filiz Ak et al. 2013; Capellupo et al. 2013, 2014; Misawa et al. 2014a; Grier et al. 2015; Arav et al. 2012, 2015; Rogerson et al. 2016; Moravec et al. 2017, and refs. therein). The evidence for BAL saturation discussed in Section 3.1 and Section 3.4 might favor crossing clouds in many situations. However, if the BAL changes in PDS 456 were caused by ionisation changes and the ionisation was initially low to produce Ly α absorption without accompanying C IV and O VI lines, then the only way to make Ly α weaker at a fixed N_{H} is with larger ionisation parameters that should produce stronger C IV and O VI absorption. The continued absence of these accompanying lines in the 2014 (Figure 2) thus makes the Ly α identification more difficult because the values of N_{H} and U in 2000 would need to be several times lower than the already-low upper limits deduced from Figure 6.

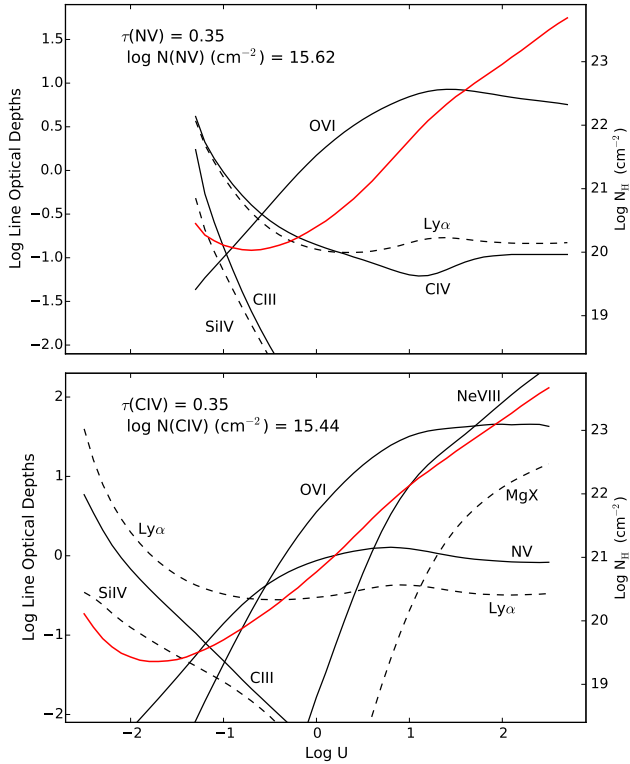


Figure 7. Line-center optical depths (black curves) and total hydrogen column densities, $\log N_{\text{H}}$ (red curve), versus ionisation parameter, $\log U$, predicted by Cloudy to accompany the observed BAL (with $\tau_0 = 0.35$ and $b = 5135 \text{ km s}^{-1}$; Section 3.1) if it is attributed to N v (top panel) or C iv (bottom panel). The ionic column densities held fixed in the model clouds are shown in the upper left of each panel. These predictions and existing spectra (Figures 1–3) rule out low degrees of ionisation for the N v BAL system, and they can guide future observations to test both BAL identifications further (see Section 3.6 and Section 3.7).

3.6 Problems with the N v $\lambda 1240$ Identification

The top panel in Figure 7 shows Cloudy predictions for the line-center optical depths that should accompany the observed BAL if it is N v $\lambda 1240$. As in Figure 6, the calculations assume a cloud velocity dispersion $v_{\text{turb}} = 1000 \text{ km s}^{-1}$ (to include moderate amounts of line shielding) but line optical depths derived for $b = 5135 \text{ km s}^{-1}$. The total column densities in the model clouds are adjusted to yield an N v column density, $\log N(\text{N v})(\text{cm}^{-2}) = 15.62$, fixed to the value inferred from the observed BAL (Section 3.1). The resulting optical depths are again only lower limits because $N(\text{N v})$ is a lower limit.

The observed BAL at $1346 \text{ \AA}$ is compatible with the N v identification only if the accompanying C iv line is undetectable. This requires a ratio of N v / C iv absorption line strengths $\gtrsim 3$ (Section 3.1, Figure 3). Figure 7 shows that moderate-to-low outflow ionisations with $\log U \lesssim 0$, corresponding to total columns $\log N_{\text{H}}(\text{cm}^{-2}) \lesssim 20.2$, are firmly ruled out for the N v BAL identification by the predictions for strong C iv absorption. At higher ionisations, the predicted optical depth ratio hovers around $\tau(\text{N v})/\tau(\text{C iv}) \sim 3\text{--}5$ (Figure 7), which is marginally consistent with the observed spectrum. However, any significant partial covering

effects would push the observed N v / C iv line depth ratio toward unity, such that C iv should be detectable and the N v BAL identification is ruled out.

Figures 4 and 5 show more directly that absorption line ratios N v / C iv $\gtrsim 3$ needed for the N v BAL identification are, at best, extremely rare in observed quasar outflows. In particular, the composite spectrum (Figure 5) shows that typical weak-to-moderate BAL/mini-BAL systems have a median line depth ratio of N v / C iv $\sim 1 \pm 0.1$. Individual outflows can have larger N v / C iv ratios. Measurements in the literature of BALs/mini-BALs at speeds up to nearly $0.2c$ reveal N v / C iv depth ratios in the range ~ 0.5 to 1.5 (Steidel 1990; Jannuzi et al. 1996; Hamann et al. 1997c,a; Barlow & Sargent 1997; Telfer et al. 1998; Sabra et al. 2003; Rodríguez Hidalgo 2008; Rodríguez Hidalgo et al. 2011, 2013; Rogerson et al. 2016; Moravec et al. 2017). Our visual inspections of the 641 outflow quasars in Section 3.4 find crudely ~ 2 percent of cases with N v / C iv $\gtrsim 2$. Figure 4 shows the most extreme well-measured examples. It is difficult to assess the N v / C iv line ratios quantitatively in these spectra due to blending problems in the Ly α forest. However, the most extreme cases³ shown in Figure 4 also appear to be inconsistent with the N v BAL identification in PDS 456.

Another constraint on the N v BAL identification is that it should be accompanied by strong O vi absorption at observed wavelength $\sim 1122 \text{ \AA}$. This is based on both theoretical predictions (Figure 7) and observations of other outflow quasars (Section 3.4 and references immediately above). The wavelength $1122 \text{ \AA}$ is just outside the wavelength coverage of the 2000 STIS spectrum; however, there is no evidence for a broad O vi line wing that could be measurable if this line was present (Figure 3). The 2014 COS spectrum does cover these O vi wavelengths and clearly shows no signs of absorption there (Figure 2), but the BAL at $1346 \text{ \AA}$ was also much weaker in this spectrum. Thus the constraints provided by absence of O vi in existing spectra are weak, but they also do not support the N v BAL identification.

3.7 The Case for C iv $\lambda 1549$ at $0.30c$

One argument favoring the C iv BAL identification in PDS 456 is that it is readily compatible with existing spectra (Figures 1–3). It has none of the problems described above for Ly α and N v because no other lines are expected within the wavelength coverage. Observations of high-velocity BALs/mini-BALs in other quasars indicate that the outflow ionisations are generally high and that the most prominent lines accompanying C iv should be N v and O vi (Section 3.4, Section 3.6, Hamann et al. 2013). The bottom panel in Figure 7 shows specific theoretical predictions for the optical depths in these and other lines that should accompany the observed BAL if it is C iv. The most observationally-accessible lines in this plot are Ly α at low ionisations and N v at high ionisations, both at predicted observer-frame wavelengths near $\sim 1068 \text{ \AA}$. Future observations might test for this absorption, but that will be difficult

³ We need to acknowledge here that hypothetical quasars with strong N v BALs but negligible C iv absorption cannot appear in our BOSS sample because it relies on C iv line detections (via AI and BI).

due to the short wavelength and severe reddening (Section 2). The situation is worse for other corroborating lines. In particular, the corresponding Si IV $\lambda 1398$ line would be at ~ 1216 Å observed, which is directly on top of the Galactic damped Ly α line, and the O VI and Ne VIII $\lambda 774$ lines, which should be strong at high ionisations are at inaccessible observed wavelengths (~ 899 Å and ~ 673 Å, respectively) due to Galactic Lyman limit absorption.

A second circumstantial argument favoring the C IV identification is that its velocity $v \approx 0.30c$ is similar to the speeds measured for X-ray outflow of PDS 456. The X-ray outflow is highly variable with multiple velocity components, but the main component identified by Fe K-shell absorption has measured velocities in the range $v \sim 0.25$ – $0.34c$ with a typical value near $\sim 0.3c$ (Reeves et al. 2014, 2016; Nardini et al. 2015; Matzeu et al. 2017). The C IV BAL might provide evidence for a physical relationship between the UV and X-ray outflows in PDS 456 (see Section 4 below).

A third argument is that relativistic C IV BALs/mini-BALs at speeds approaching $v \sim 0.2c$ have already been measured in a growing number of luminous quasars (Januzzi et al. 1996; Hamann et al. 1997b, 2013; Rodríguez Hidalgo 2008; Rodríguez Hidalgo et al. 2011; Rogerson et al. 2016, Rodríguez Hidalgo et al., in prep.). The C IV BAL at $v \approx 0.30c$ in PDS 456 would set a new speed record for UV outflows, but it is not so dramatic to be a paradigm shift for our understanding of these outflows. The width and depth of the BAL in PDS 456 is roughly similar to these other high-velocity outflow features. We also note that our tentative detection of another weak C IV BAL in PDS 456 at $v \sim 0.19c$ (Figure 2, Section 3.2) provides additional, albeit tentative, evidence that a relativistic UV outflow is present in PDS 456.

4 SUMMARY & DISCUSSION

The UV spectrum of PDS 456 obtained with HST-STIS in 2000 has a distinct BAL at ~ 1346 Å that we identify as C IV $\lambda 1549$ at velocity shift $v \approx 0.30c$, FWHM ≈ 8550 km s $^{-1}$, and minimum optical depth $\tau_o = 0.35$ (Section 3.1). The C IV identification rests on its compatibility with existing spectra and the absence of lines that should accompany the alternative identifications, Ly α $\lambda 1216$ and N V $\lambda 1240$. Ly α is compatible with the observed BAL only if the outflow has a low degree of ionisation, $\log U < -2$, and a very low total column density, $\log N_H(\text{cm}^{-2}) < 18.7$, that would be unprecedented in BAL outflow studies, e.g., several orders of magnitude below recent estimates (Section 3.5). The N V identification might be consistent with absence of accompanying C IV absorption if the gas is highly ionised and it pushes the boundary of observed N V/C IV line strengths beyond what we find in our comparison of 641 outflow quasars in BOSS (Section 3.6). However, this situation should produce strong O VI absorption, which is not well constrained in existing spectra of PDS 456 but it appears to be absent (Figure 3). Thus the N V identification is also strongly disfavored.

The C IV BAL identification has none of these problems (Section 3.7). It would mark the fastest UV outflow line ever reported, but its velocity is consistent with the X-ray outflow in PDS 456 (see Section 1 and below) and not dramatically

different from the relativistic C IV BALs/mini-BALs already measured at speeds approaching $\sim 0.2c$ in other quasars (Januzzi et al. 1996; Hamann et al. 1997b, 2013; Rodríguez Hidalgo 2008; Rodríguez Hidalgo et al. 2011; Rogerson et al. 2016, Rodríguez Hidalgo et al., in prep.). The C IV identification is also weakly supported by our tentative identification of an additional C IV BAL feature at $v \sim 0.19c$ (, Figure 2, Section 3.2). Broad UV outflow lines at speeds near $0.3c$ are surely very rare based on the rarity of such lines at ~ 0.1 – $0.2c$ (see refs. above). However, the incidence of C IV BALs/mini-BALs at $v > 0.2c$ is not known because they have not been searched for in large quasar surveys like SDSS/BOSS and, in any case, they could easily be missed due to blends with unrelated lines in the Ly α forest (e.g., see the search for P V $\lambda 1121$ BALs by Capellupo et al. 2017).

The location of the C IV BAL outflow in PDS 456 is a critical unknown. The range of ionizations consistent with C IV absorption might favor lower ionizations and larger distances from the black hole than the X-ray outflow, perhaps at radii of order ~ 1 pc as inferred from some variability studies of C IV BALs/mini-BALs in other quasars (Hamann et al. 2013; Capellupo et al. 2014; Moravec et al. 2017, McGraw et al., submitted). However, larger distances do not necessarily produce lower ionisations, e.g., if acceleration causes the outflow densities to drop faster than the $1/r^2$ behaviour expected from free expansion at a constant speed. Moreover, the relativistic speed of the C IV BAL at $v \sim 0.30c$ indicates that the UV outflow originated with the X-ray UFO very close to the black hole. If the measured flow speeds are similar to the gravitational escape speed at the launch radius, then the launch point is roughly at $r \sim 20$ – $30 r_g$ (Nardini et al. 2015; Matzeu et al. 2017). Also note that a highly-ionised X-ray outflow launched from this radius and expanding freely outward into a fixed solid angle will not necessarily become less ionised at larger distances (for C IV absorption) because the $1/r^2$ dilution of the radiation field is balanced by a $1/r^2$ decline in the densities to yield a constant ionisation parameter. Lower ionisations will occur if the inner regions of the outflow radiatively shield the material downstream and/or if there are clumps with enhanced densities relative to the ambient flow. Clumping and shielding can occur to the same effect at almost any radius, and there is already evidence for dense clumps in the X-ray outflow of PDS 456 (based on lower ionizations and partial covering in the soft X-ray absorber, Reeves et al. 2016; Matzeu et al. 2016; see also Gofford et al. 2014; Nardini et al. 2015; Hagino et al. 2015).

It is therefore an intriguing possibility that the C IV BAL forms directly within, or in close proximity to, the relativistic X-ray outflow. Our Cloudy simulations (Figure 7) show that the C IV BAL in PDS 456 could form over a wide range of physical conditions, including very high ionisations where C IV is just a trace constituent. This situation is illustrated in Figure 8, which plots the ionisation structure in a single model cloud with $\log U = 1.7$. This cloud reaches the observed minimum BAL optical depth $\tau_o(\text{C IV}) = 0.35$, along with $\tau_o(\text{N V}) = 0.90$ and $\tau_o(\text{O VI}) = 43$, at total column density $\log N_H(\text{cm}^{-2}) = 22.9$ (represented by the unshaded left-hand portion of Figure 8). This front portion of the cloud matches our calculations in the bottom panel of Figure 7. In this environment, the C IV and N V ion fractions are everywhere $\lesssim 4 \times 10^{-4}$ and $\lesssim 3 \times 10^{-3}$, respectively, and the

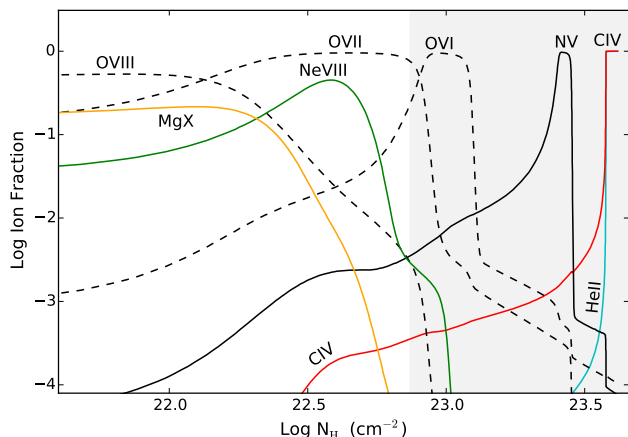


Figure 8. Ionisation fractions versus total hydrogen column density, $\log N_{\text{H}}$, in a single model cloud with ionisation parameter, $\log U = 1.7$. This cloud produces the minimum observed BAL optical depth in C IV $\tau_{\text{o}}(\text{C IV}) = 0.35$, at $\log N_{\text{H}}(\text{cm}^{-2}) \approx 22.9$, such that the unshaded left-hand region represents the model cloud with $\log U = 1.7$ in Figure 7. This environment (at $\log N_{\text{H}}(\text{cm}^{-2}) \leq 22.9$) is highly ionised and optically thick at the bound-free edges of O VII, O VIII, and Ne VIII, with only trace amounts of the lower ions C IV, NV, and O VI. The lower ions dominate at larger column densities (grey shaded region) where there is more radiative shielding, but these lower ionisations are not needed for the C IV BAL in PDS 456.

dominant form of oxygen is O VII. This highly-ionised environment capable of C IV $\lambda 1549$ absorption will also produce saturated absorption in higher ion lines such as Ne VIII $\lambda 774$ and Mg X $\lambda 615$ (see also $\log U = 1.7$ Figure 7) and at the bound-free edges of O VII and O VIII in soft X-rays.

Detailed comparisons between the UV and X-ray outflows in PDS 456 are beyond the scope of the present study. They are subject to uncertainties caused by the outflow variabilities and by unknowns in the shape of the ionising spectrum and the spatial locations of different outflow components relative to the ionizing far-UV and X-ray emission sources. Here we note simply that recent estimates of the X-ray outflow ionisation parameters⁴ are roughly in the range $\log U \sim 4.3$ to 4.8 for the K-shell outflow (Gofford et al. 2014; Nardini et al. 2015) and $\log U \sim 2.8$ for the soft X-ray absorber (Reeves et al. 2016). The ionisation parameters needed for the C IV BAL, with upper limit $\log U \lesssim 2.4$ (Figure 7), might occur in dense clumps embedded in the X-ray outflow if the density enhancements are $\gtrsim 0.4$ dex relative to the soft X-ray absorber or $\gtrsim 2$ dex relative to the Fe K-shell outflow.

The ionization structure of the UV outflow might resemble the quasar SBS 1542+541 (redshift $z \sim 2.36$), which has measured UV BALs ranging in ionization from C IV up to Mg X and Si XII $\lambda 510$ (Telfer et al. 1998). The lines in this quasar appear saturated with shallow troughs that reveal ion-dependent line-of-sight covering fractions from ~ 15 percent in C IV to ~ 50 percent in the higher ions. Telfer et al.

(1998) infer from this a 2-zone outflow that could span a decade or more in ionisation parameter, with the C IV lines forming in small clumps embedded in a more highly-ionised outflow medium. Telfer et al. (1998) also note that this outflow should produce substantial bound-free absorption by O VII and O VIII in soft X-rays.

The general picture of clumpy outflows has become commonplace in quasar outflow studies. High-quality observations of UV outflow lines often provide evidence for clumpy multi-phase outflow structures with a range of covering fractions (Moravec et al. 2017; Misawa et al. 2014b; Hamann et al. 2011; Hamann & Sabra 2004; Gabel et al. 2005a; Leighly et al. 2015, 2011, 2009; Misawa et al. 2007; Arav et al. 2008, 2005; de Kool et al. 2002; Ganguly et al. 1999). Detailed studies of some bright Seyfert 1 galaxies (e.g., Kaspi et al. 2002; Netzer et al. 2003; Gabel et al. 2005b) clearly demonstrate that the UV and X-ray absorption features can form together in complex outflows, with indications that the lower-ionisation UV lines identify clumps or filaments embedded in the X-ray outflow. Clumpy outflow structures are also predicted by recent numerical simulations (Sim et al. 2010; Takeuchi et al. 2013; Waters et al. 2017) and by considerations of the radiative forces that can compress the outflows into small substructures (Stern et al. 2014). There are also theoretical arguments requiring dense clumps to moderate the outflow ionizations in the absence of significant radiative shielding (De Kool 1997; Hamann et al. 2013).

If the C IV-absorbing gas in PDS 456 is indeed embedded in the X-ray outflow, it could have important implications for the outflow energetics. It is a well-known problem for X-ray UFOs (e.g., Tombesi et al. 2011, 2013; Gofford et al. 2015) that their high ionizations lead to low opacities and inefficient radiative acceleration (Gofford et al. 2013, 2014; Higginbottom et al. 2014). Harnessing the full radiative power of the quasar to drive these outflows might require opacities beyond electron scattering at UV/far-UV wavelengths, near the peak of the quasar spectral energy distributions. Matzeu et al. (2017) showed recently that the outflow speeds in PDS 456 correlate with the variable X-ray luminosity, consistent with radiative acceleration (see also Saez & Chartas 2011). If radiative forces are important, then dense clumps with lower ionisations embedded in the X-ray outflows might be important to boost the opacities for radiative driving (Laor & Davis 2014; Hagino et al. 2015). The C IV BAL at $v \approx 0.30c$ in PDS 456 could be the first direct observational evidence for this idea.

PDS 456 is a remarkable object with powerful accretion-disk outflows launched from a range of at least two decades in disk radii, from ~ 0.001 pc for the X-ray outflow to ~ 0.3 pc for the blueshifted C IV broad emission line (Section 1, O’Brien et al. 2005; Gofford et al. 2014). More work is needed to understand how the UV BAL fits into this outflow environment and, specifically, to test the C IV BAL identification. A search for corroborating NV absorption at ~ 1077 Å will be difficult due to Galactic reddening (Section 2) and the poor sensitivity of current instruments at these wavelengths (Section 3.7). However, the alternative BAL identifications, NV and Ly α , might be ruled out or confirmed more easily by searching for O VI absorption near 1122 Å (Section 3.6, although the BAL at 1346 Å needs to be present for a meaningful test). Our team has ongoing programs to extend the

⁴ The ionisation parameter quoted in the X-ray studies is $\xi = L_{\text{ion}}/n_{\text{H}}r^2$ from Tarter et al. (1969), where L_{ion} is the quasar luminosity from 1 to 1000 Rydberg. For ξ in units of ergs cm s^{-1} and the continuum shape used in our Cloudy simulations (Section 3.3), the conversion is $\log U \approx \log \xi - 1.2$.

UV wavelength coverage and monitor PDS 456 in the UV and X-rays that will be described in future papers.

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REFERENCES

- Arav N., et al., 2001, *ApJ*, **561**, 118
- Arav N., Kaastra J., Kriss G. A., Korista K. T., Gabel J., Proga D., 2005, *ApJ*, **620**, 665
- Arav N., Moe M., Costantini E., Korista K. T., Benn C., Ellison S., 2008, *ApJ*, **681**, 954
- Arav N., et al., 2012, *A&A*, **544**, A33
- Arav N., et al., 2015, *A&A*, **577**, A37
- Barlow T. A., Sargent W. L. W., 1997, *AJ*, **113**, 136
- Baskin A., Laor A., Hamann F., 2013, *MNRAS*, **432**, 1525
- Baskin A., Laor A., Hamann F., 2015, *MNRAS*, **449**, 1593
- Borguet B. C. J., Edmonds D., Arav N., Benn C., Chamberlain C., 2012, *ApJ*, **758**, 69
- Capellupo D. M., Hamann F., Shields J. C., Halpern J. P., Barlow T. A., 2013, *MNRAS*, **429**, 1872
- Capellupo D. M., Hamann F., Barlow T. A., 2014, *MNRAS*, **444**, 1893
- Capellupo D. M., et al., 2017, *MNRAS*, **469**, 323
- Cardelli J. A., Clayton G. C., Mathis J. S., 1989, *ApJ*, **345**, 245
- Chartas G., Brandt W. N., Gallagher S. C., Garmire G. P., 2002, *ApJ*, **579**, 169
- Chartas G., et al., 2009, *New Astronomy Review*, **53**, 128
- Coatman L., Hewett P. C., Banerji M., Richards G. T., 2016, *MNRAS*, **461**, 647
- Crenshaw D. M., Kraemer S. B., George I. M., 2003, *ARA&A*, **41**, 117
- Dawson K. S., et al., 2013, *AJ*, **145**, 10
- De Kool M., 1997, in Arav N., Shlosman I., Weymann R. J., eds, *Astronomical Society of the Pacific Conference Series Vol. 128, Mass Ejection from Active Galactic Nuclei*. p. 233 ([arXiv:astro-ph/9704230](https://arxiv.org/abs/astro-ph/9704230))
- Debuhr J., Quataert E., Ma C.-P., 2012, *MNRAS*, **420**, 2221
- Di Matteo T., Springel V., Hernquist L., 2005, *Nature*, **433**, 604
- Dietrich M., Hamann F., Shields J. C., Constantin A., Vestergaard M., Chaffee F., Foltz C. B., Junkkarinen V. T., 2002, *ApJ*, **581**, 912
- Eisenstein D. J., et al., 2011, *AJ*, **142**, 72
- Ferland G. J., et al., 2013, *Rev. Mex. Astron. Astrofis.*, **49**, 137
- Ferland G. J., et al., 2017, preprint, ([arXiv:1705.10877](https://arxiv.org/abs/1705.10877))
- Filiz Ak N., et al., 2013, *ApJ*, **777**, 168
- Gabel J. R., et al., 2005a, *ApJ*, **623**, 85
- Gabel J. R., et al., 2005b, *ApJ*, **631**, 741
- Ganguly R., Eracleous M., Charlton J. C., Churchill C. W., 1999, *AJ*, **117**, 2594
- Gibson R. R., et al., 2009, *ApJ*, **692**, 758
- Gofford J., Reeves J. N., Tombesi F., Braitto V., Turner T. J., Miller L., Cappi M., 2013, *MNRAS*, **430**, 60
- Gofford J., et al., 2014, *ApJ*, **784**, 77
- Gofford J., Reeves J. N., McLaughlin D. E., Braitto V., Turner T. J., Tombesi F., Cappi M., 2015, *MNRAS*, **451**, 4169
- Grier C. J., et al., 2015, *ApJ*, **806**, 111
- Hagino K., Odaka H., Done C., Gandhi P., Watanabe S., Sako M., Takahashi T., 2015, *MNRAS*, **446**, 663
- Hamann F., 1997, *ApJS*, **109**, 279
- Hamann F., 1998, *ApJ*, **500**, 798
- Hamann F., Ferland G., 1999, *ARA&A*, **37**, 487
- Hamann F., Sabra B., 2004, in Richards G. T., Hall P. B., eds, *Astronomical Society of the Pacific Conference Series Vol. 311, AGN Physics with the Sloan Digital Sky Survey*. p. 203
- Hamann F., Barlow T., Cohen R. D., Junkkarinen V., Burbidge E. M., 1997a, in Arav N., Shlosman I., Weymann R. J., eds, *Astronomical Society of the Pacific Conference Series Vol. 128, Mass Ejection from Active Galactic Nuclei*. p. 19
- Hamann F., Barlow T. A., Junkkarinen V., Burbidge E. M., 1997b, *ApJ*, **478**, 80
- Hamann F., Barlow T. A., Junkkarinen V., 1997c, *ApJ*, **478**, 87
- Hamann F., Kanekar N., Prochaska J. X., Murphy M. T., Ellison S., Malec A. L., Milutinovic N., Ubachs W., 2011, *MNRAS*, **410**, 1957
- Hamann F., Chartas G., McGraw S., Rodriguez Hidalgo P., Shields J., Capellupo D., Charlton J., Eracleous M., 2013, *MNRAS*, **435**, 133
- Hamann F., et al., 2017, *MNRAS*, **464**, 3431
- Higginbottom N., Proga D., Knigge C., Long K. S., Matthews J. H., Sim S. A., 2014, *ApJ*, **789**, 19
- Hopkins P. F., Elvis M., 2010, *MNRAS*, **401**, 7
- Hopkins P. F., Hernquist L., Cox T. J., Kereš D., 2008, *ApJS*, **175**, 356
- Jannuzi B. T., et al., 1996, *ApJ*, **470**, L11+
- Kaspi S., et al., 2002, *ApJ*, **574**, 643
- Kaspi S., Maoz D., Netzer H., Peterson B. M., Vestergaard M., Jannuzi B. T., 2005, *ApJ*, **629**, 61
- Knigge C., Scaringi S., Goad M. R., Cottis C. E., 2008, *MNRAS*, **386**, 1426
- Korista K. T., Voit G. M., Morris S. L., Weymann R. J., 1993, *ApJS*, **88**, 357
- Laor A., Davis S. W., 2014, *MNRAS*, **438**, 3024
- Leighly K. M., 2004, *ApJ*, **611**, 125
- Leighly K. M., Halpern J. P., Jenkins E. B., Casebeer D., 2007, *ApJS*, **173**, 1
- Leighly K. M., Hamann F., Casebeer D. A., Grupe D., 2009, *ApJ*, **701**, 176
- Leighly K. M., Dietrich M., Barber S., 2011, *ApJ*, **728**, 94
- Leighly K. M., Cooper E., Grupe D., Terndrup D. M., Komossa S., 2015, *ApJ*, **809**, L13
- Luo B., et al., 2015, *ApJ*, **805**, 122
- Matzeu G. A., Reeves J. N., Nardini E., Braitto V., Costa M. T., Tombesi F., Gofford J., 2016, *MNRAS*, **458**, 1311
- Matzeu G. A., Reeves J. N., Braitto V., Nardini E., McLaughlin D. E., Lobban A. P., Tombesi F., Costa M. T., 2017, preprint, ([arXiv:1708.03546](https://arxiv.org/abs/1708.03546))
- Misawa T., Eracleous M., Charlton J. C., Kashikawa N., 2007, *ApJ*, **660**, 152
- Misawa T., Charlton J. C., Eracleous M., 2014a, *ApJ*, **792**, 77
- Misawa T., Inada N., Oguri M., Gandhi P., Horiuchi T., Koyamada S., Okamoto R., 2014b, *ApJ*, **794**, L20
- Moravec E. A., Hamann F., Capellupo D. M., McGraw S. M., Shields J. C., Rodríguez Hidalgo P., 2017, *MNRAS*, **468**, 4539
- Nagao T., Marconi A., Maiolino R., 2006, *A&A*, **447**, 157
- Nardini E., et al., 2015, *Science*, **347**, 860
- Netzer H., et al., 2003, *ApJ*, **599**, 933
- O'Brien P. T., Reeves J. N., Simpson C., Ward M. J., 2005, *MNRAS*, **360**, L25
- Pâris I., Petitjean P., Ross N. P., Myers A. D., Aubourg É., Streblyanska A., Bailey S. e. a., 2017, *A&A*, **597**, A79

- Reeves J. N., O'Brien P. T., Ward M. J., 2003, *ApJ*, **593**, L65
- Reeves J. N., et al., 2009, *ApJ*, **701**, 493
- Reeves J. N., et al., 2014, *ApJ*, **780**, 45
- Reeves J. N., Braitto V., Nardini E., Behar E., O'Brien P. T., Tombesi F., Turner T. J., Costa M. T., 2016, *ApJ*, **824**, 20
- Richards G. T., et al., 2011, *AJ*, **141**, 167
- Rodríguez Hidalgo P., 2008, PhD thesis, University of Florida, Gainesville, http://etd.fcla.edu/UF/UFE0022674/rodriguez_p.pdf
- Rodríguez Hidalgo P., Hamann F., Hall P., 2011, *MNRAS*, **411**, 247
- Rodríguez Hidalgo P., Eracleous M., Charlton J., Hamann F., Murphy M. T., Nestor D., 2013, *ApJ*, **775**, 14
- Rogerson J. A., Hall P. B., Rodríguez Hidalgo P., Pirkola P., Brandt W. N., Filiz Ak N., 2016, *MNRAS*, **457**, 405
- Ross N. P., Myers A. D., Sheldon E. S., Yèche C., Strauss M. A., Bovy J., Kirkpatrick J. A. e. a., 2012, *ApJS*, **199**, 3
- Rupke D. S. N., Veilleux S., 2013, *ApJ*, **768**, 75
- Sabra B. M., Hamann F., Jannuzi B. T., George I. M., Shields J. C., 2003, *ApJ*, **590**, 66
- Saez C., Chartas G., 2011, *ApJ*, **737**, 91
- Sim S. A., Proga D., Miller L., Long K. S., Turner T. J., 2010, *MNRAS*, **408**, 1396
- Simpson C., Ward M., O'Brien P., Reeves J., 1999, *MNRAS*, **303**, L23
- Steffen A. T., Strateva I., Brandt W. N., Alexander D. M., Koeke-moer A. M., Lehmer B. D., Schneider D. P., Vignali C., 2006, *AJ*, **131**, 2826
- Steidel C. C., 1990, *ApJS*, **74**, 37
- Stern J., Behar E., Laor A., Baskin A., Holczer T., 2014, *MNRAS*, **445**, 3011
- Strateva I. V., Brandt W. N., Schneider D. P., Vanden Berk D. G., Vignali C., 2005, *AJ*, **130**, 387
- Takeuchi S., Ohsuga K., Mineshige S., 2013, *PASJ*, **65**, 88
- Tarter C. B., Tucker W. H., Salpeter E. E., 1969, *ApJ*, **156**, 943
- Telfer R. C., Kriss G. A., Zheng W., Davidsen A. F., Green R. F., 1998, *ApJ*, **509**, 132
- Tombesi F., Cappi M., Reeves J. N., Palumbo G. G. C., Yaqoob T., Braitto V., Dadina M., 2010, *A&A*, **521**, A57
- Tombesi F., Cappi M., Reeves J. N., Palumbo G. G. C., Braitto V., Dadina M., 2011, *ApJ*, **742**, 44
- Tombesi F., Cappi M., Reeves J. N., Braitto V., 2012, *MNRAS*, **422**, L1
- Tombesi F., Cappi M., Reeves J. N., Nemmen R. S., Braitto V., Gaspari M., Reynolds C. S., 2013, *MNRAS*, **430**, 1102
- Torres C. A. O., Quast G. R., Coziol R., Jablonski F., de la Reza R., Lépine J. R. D., Gregório-Hetem J., 1997, *ApJ*, **488**, L19
- Trump J. R., et al., 2006, *ApJS*, **165**, 1
- Warner C., Hamann F., Dietrich M., 2004, *ApJ*, **608**, 136
- Waters T., Proga D., Dannen R., Kallman T. R., 2017, *MNRAS*, **467**, 3160
- Weymann R. J., Morris S. L., Foltz C. B., Hewett P. C., 1991, *ApJ*, **373**, 23
- Wu J., et al., 2011, *ApJ*, **736**, 28
- Wu J., Brandt W. N., Anderson S. F., Diamond-Stanic A. M., Hall P. B., Plotkin R. M., Schneider D. P., Shemmer O., 2012, *ApJ*, **747**, 10
- Zakamska N. L., et al., 2016, *MNRAS*, **459**, 3144
- de Kool M., Korista K. T., Arav N., 2002, *ApJ*, **580**, 54

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