

## The XXL Survey

### IX. Optical overdensity and radio continuum analysis of a supercluster at $z = 0.43^{*,**}$

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#### ABSTRACT

We present observations with the *Karl G. Jansky* Very Large Array (VLA) at 3 GHz (10 cm) toward a sub-field of the XXL-North 25 deg<sup>2</sup> field targeting the first supercluster discovered in the XXL Survey. The structure has been found at a spectroscopic redshift of 0.43 and extending over 0°:35 × 0°:1 on the sky. The aim of this paper is twofold. First, we present the 3 GHz VLA radio continuum observations, the final radio mosaic and radio source catalogue, and, second, we perform a detailed analysis of the supercluster in the optical and radio regimes using photometric redshifts from the CFHTLS survey and our new VLA-XXL data. Our final 3 GHz radio mosaic has a resolution of 3'':2 × 1'':9, and encompasses an area of 41' × 41' with rms noise level lower than ~20 μJy beam<sup>-1</sup>. The noise in the central 15' × 15' region is ~11 μJy beam<sup>-1</sup>. From the mosaic we extract a catalogue of 155 radio sources with signal-to-noise ratio (S/N) ≥ 6, eight of which are large, multicomponent sources, and 123 (79%) of which can be associated with optical sources in the CFHTLS W1 catalogue. Applying Voronoi tessellation analysis (VTA) in the area around the X-ray identified supercluster using photometric redshifts from the CFHTLS survey we identify a total of seventeen overdensities at  $z_{\text{phot}} = 0.35\text{--}0.50$ , 7 of which are associated with clusters detected in the *XMM-Newton* XXL data. We find a mean photometric redshift of 0.43 for our overdensities, consistent with the spectroscopic redshifts of the brightest cluster galaxies of seven X-ray detected clusters. The full VTA-identified structure extends over ~0°:6 × 0°:2 on the sky, which corresponds to a physical size of ~12 × 4 Mpc<sup>2</sup> at  $z = 0.43$ . No large radio galaxies are present within the overdensities, and we associate eight (S/N > 7) radio sources with potential group/cluster member galaxies. The spatial distribution of the red and blue VTA-identified potential group member galaxies, selected by their observed  $g - r$  colours, suggests that the clusters are not virialised yet, but are dynamically young, as expected for hierarchical structure growth in a  $\Lambda$ CDM universe. Further spectroscopic data are required to analyse the dynamical state of the groups.

**Key words.** catalogs – galaxies: clusters: general – galaxies: groups: general – radio continuum: galaxies

#### 1. Introduction

Over the last decade, deeper insights into various physical properties of galaxies, their formation and their evolution through cosmic time have been gained by sensitive, multiwavelength

surveys such as GOODS (Dickinson et al. 2003), COSMOS (Scoville et al. 2007), GAMA (Driver et al. 2009, 2011) and CANDELS (Koekemoer et al. 2011; Grogin et al. 2011). The XXL<sup>1</sup> is a panchromatic survey (X-ray to radio) of two regions on the sky, each 25 deg<sup>2</sup> in size. Constraining the time evolution of the Dark Energy equation of state using galaxy clusters is the main goal of the XXL project. In a deep (6 Ms) *XMM-Newton*<sup>2</sup> survey to a depth of ~5 × 10<sup>-15</sup> erg s<sup>-1</sup> cm<sup>-2</sup> in the [0.5–2] keV band several hundred galaxy cluster detections (0 <  $z$  < 1.5) were made over the ~50 deg<sup>2</sup> (Pierre et al. 2016, hereafter Paper I). The XXL South region (XXL-S) is

\* The full catalogue is available as a queryable database table XXL\_VLA\_15 via the XXL Master Catalogue browser <http://cosmosdb.iasf-milano.inaf.it/XXL>. A copy of the catalogue and the mosaic are also available at the CDS via anonymous ftp to [cdsarc.u-strasbg.fr](http://cdsarc.u-strasbg.fr) (130.79.128.5) or via <http://cdsarc.u-strasbg.fr/viz-bin/qcat?IX/49>

\*\* A FITS file of the reduced image is only available at the CDS via anonymous ftp to [cdsarc.u-strasbg.fr](http://cdsarc.u-strasbg.fr) (130.79.128.5) or via <http://cdsarc.u-strasbg.fr/viz-bin/qcat?J/A+A/592/A8>

<sup>1</sup> <http://irfu.cea.fr/xxl>

<sup>2</sup> <http://xmm.esac.esa.int/>

located in the southern hemisphere in the Blanco Cosmology Survey (BCS) field<sup>3</sup> and the XXL North region (XXL-N) near the celestial equator encompasses the smaller XMM Large Scale Structure Survey (XMM-LSS)<sup>4</sup> field. Multiwavelength photometric data (<25 in AB mag) from the UV/optical to the IR drawn from GALEX, Canada-France-Hawaii Telescope Legacy Survey (CFHTLS), BCS, Sloan Digital Sky Survey (SDSS)<sup>5</sup>, the Two Micron All Sky Survey (2MASS)<sup>6</sup>, *Spitzer* Space Telescope<sup>7</sup>, and Wide-field Infrared Survey Explorer (WISE)<sup>8</sup> already exist over almost the full area (Paper I). Even deeper observations are available over smaller areas (Pierre et al. 2016), and will be conducted over the full area with the DECam<sup>9</sup> to magnitude 25–26 in *griz* bands, and 26–27 in *grizY* bands with Hyper-SuprimeCam<sup>10</sup>. Photometric redshifts are expected to reach accuracies better than ~10%, sufficient for large-scale structure studies out to intermediate redshift ( $z \sim 0.5$ ), and for evolutionary studies of galaxies and AGN to  $z \sim 3$ . More than 15 000 optical spectra are already available (see Paper I).

To complement the multiwavelength coverage of the field, we present new radio observations with the *Karl G. Jansky* Very Large Array (VLA) at 3 GHz over a  $\sim 0:7 \times 0:7$  sub-area of the XXL-N field. These observations target the first supercluster discovered in the XXL survey by Pompei et al. (2016, hereafter Paper VII). One cluster within this large structure was previously identified at  $z_{\text{phot}} = 0.48$  using photometric redshifts in the CFHTLS wide field by Durret et al. (2011). The *XMM-Newton* XXL observations revealed six X-ray clusters within  $\sim 0:35 \times 0:1$  as described in Paper I, see also Pacaud et al. (2016, hereafter Paper II). Based on further spectroscopic observations of the brightest cluster galaxies by Koulouridis et al. (2016, hereafter Paper XII), Paper VII reports the redshift of the structure of  $z = 0.43$ . The structure, thus, has a physical extent of  $\sim 10 \times 2.9 \text{ Mpc}^2$  ( $21' \times 6'$ ). Here we present a detailed analysis of this structure based on the CFHTLS optical data and our new VLA radio data.

Throughout this paper we use cosmological parameters in accordance with the *Wilkinson* Microwave Anisotropy Probe satellite final data release (WMAP9) combined with a set of baryon acoustic oscillation measurements and constraints on  $H_0$  from Cepheids and type Ia supernovae (Hinshaw et al. 2013). These parameters are  $\Omega_M = 0.282$ ,  $\Omega_\Lambda = 0.718$  and  $H_0 = 69.7 \text{ km s}^{-1} \text{ Mpc}^{-1}$ . All sizes given in this paper are physical. All magnitudes are in the AB magnitude system, and all coordinates in J2000 epoch.

In Sect. 2 we present the optical and radio data used in this paper. In Sect. 3 we describe optical and radio properties of the supercluster. Discussion of the results is presented in Sect. 4 and the summary is given in Sect. 5.

## 2. Data

### 2.1. Optical data

We use the photometric redshift catalogues from the CFHTLS<sup>11</sup>. The CFHTLS Wide survey is divided into fields W1, W2, W3

<sup>3</sup> <http://www.usm.uni-muenchen.de/BCS/>

<sup>4</sup> <http://wela.astro.ulg.ac.be/themes/spatial/xmm/LSS/>

<sup>5</sup> <http://www.sdss.org/>

<sup>6</sup> <http://www.ipac.caltech.edu/2mass/>

<sup>7</sup> <http://ssc.spitzer.caltech.edu/>

<sup>8</sup> <http://wise.ssl.berkeley.edu/>

<sup>9</sup> <http://www.darkenergysurvey.org/DECam/camera.shtml>

<sup>10</sup> <http://www.naoj.org/Projects/HSC/>

<sup>11</sup> <http://www.cfht.hawaii.edu/Science/CFHTLS/>

**Table 1.** List of VLA pointings used in the observations.

Pointing	RA	Dec	Time on-source
(1)	[h m s]	[° ' '']	[s]
(1)	(2)	(3)	(4)
P1	02 10 14.66	−06 22 20.00	972
P2	02 10 14.66	−06 12 20.00	987
P3	02 10 14.66	−06 02 20.00	975
P4	02 10 54.89	−06 22 20.00	984
P5	02 10 54.89	−06 12 20.00	981
P6	02 10 54.89	−06 02 20.00	975
P7	02 11 35.11	−06 22 20.00	999
P8	02 11 35.11	−06 12 20.00	984
P9	02 11 35.11	−06 02 20.00	978

**Notes.** Pointing names are stated in Col. (1), and their central coordinates in sexagesimal format in Cols. (2) and (3). Total on-source time in seconds is given in Col. (4). All observations were performed in the S-band with 3 GHz central frequency; B configuration has the largest recoverable angular scale of  $58''$ .

and W4 covered by the  $u^*$ ,  $g'$ ,  $r'$ ,  $i'$ ,  $z'$ , and  $y'$  bands with an 80% completeness limit in AB magnitudes for point sources of  $u^* = 25.2$ ,  $g' = 25.5$ ,  $r' = 25.0$ ,  $i' = 24.8$  and  $z' = 23.9$ . The XXL-N field lies within the CFHTLS W1 field. The photometric redshift accuracy in this field is  $\sigma_{\Delta z/(1+z_s)} \sim 0.030\text{--}0.032$  at  $i < 21.5$ ,  $\sigma_{\Delta z/(1+z_s)} \sim 0.066\text{--}0.077$  at  $22.5 < i < 23.5$ , where  $\Delta z = |z_p - z_s|$  is the difference between photometric and spectroscopic redshift as described in Ilbert et al. (2006) and Coupon et al. (2009) (see also catalogue documentation<sup>12</sup>). Since the redshift accuracy rapidly deteriorates for magnitudes above  $i = 23.5$  in the optical analysis of the structure, we have adopted a cut in  $i$ -band magnitude at  $i = 23.5$ .

To obtain stellar mass estimates used in Sect. 4 we used publicly available data from the Canada-France-Hawaii Lensing Survey (CFHTLenS; see Heymans et al. 2012; Hildebrandt et al. 2012; Erben et al. 2013; and Velander 2012). The stellar masses were extracted using “LePhare” (see Ilbert et al. 2006 and Arnouts et al. 1999) assuming a Chabrier initial mass function (see Chabrier 2003).

### 2.2. VLA 3 GHz radio data

#### 2.2.1. Observations, data reduction, and imaging

We observed nine regularly spaced pointings separated by  $10'$  in RA and Dec in the XXL-N field with the VLA in B configuration, as given in the Table 1. The area observed has central coordinates RA =  $02^{\text{h}}10^{\text{m}}54^{\text{s}}.9$ , Dec =  $-06^\circ 12' 20''.2$ . Three hours of observations were taken on August 30, 2012 (7:35–10:35 UTC) in S-band with central frequency of 3 GHz and 2 GHz bandwidth divided into 16 spectral windows 128 MHz in width. Owing to large-amplitude radio frequency interference (RFI) between 2.12–2.38 GHz, the second and third spectral windows were completely flagged. For flux calibration we observed the calibrator 0521+166 (Seyfert 1 galaxy 3C138) for 105 s at the end of the observations. For phase calibration we used J0241-0815 (Seyfert 2 galaxy NGC1052). The phase calibrator was observed every 30 min for 2 min on-source throughout the observations. The observations were taken under a clear sky at a temperature of  $12^\circ\text{C}$ .

<sup>12</sup> [http://cesam.lam.fr/cfhtls-zphotos/files/cfhtls\\_wide\\_T007\\_v1.2\\_Oct2012.pdf](http://cesam.lam.fr/cfhtls-zphotos/files/cfhtls_wide_T007_v1.2_Oct2012.pdf)

Calibration of the data was done using the AIPSLITE pipeline (e.g. Bourke et al. 2014) developed for the Caltech-NRAO Stripe 82 Survey (Mooley et al. 2015). The calibrated data were then imported into the CASA package and data with amplitudes greater than 0.4 Jy were further flagged using the CASA task FLAGMINMAX. Using the task SPLIT the  $uv$  data were then split into separate pointings prior to imaging. Imaging of the  $uv$  data was performed for each pointing separately using multifrequency synthesis within the CASA task CLEAN (Rau & Cornwell 2011; see also Novak et al. 2015). The algorithm models the sky brightness in the entire 2 GHz bandwidth using a linear combination of Gaussian functions with amplitudes following a Taylor polynomial in frequency. A two-term expansion was used with 10 000 iterations, threshold  $3\sigma$  ( $40 \mu\text{Jy}$ ), default gain value of 0.1 and “gridmode” parameter set to wide field with 128  $w$ -projection planes for convolution. We used Briggs weighting with a robust value of 0.5. This value was chosen as it optimally balances rms and sidelobe contamination. After the imaging, a wide-band primary beam correction was applied using the CASA task WIDEBANDPBCOR. This was executed with two Taylor terms (“nterms”) and a threshold of 20% of the primary beam response was set.

Each pointing was convolved to  $3''.2 \times 1''.9$  with position angle of  $20^\circ$  to assure a uniform synthesised beam across the final mosaic. To construct the mosaic we combined the nine individual pointings by performing an rms-weighted sum of the pixels in the overlapping regions. For each pixel the local rms value was obtained from the corresponding pixel in the noise map. These noise maps were constructed using the AIPS<sup>13</sup> task RMSD and are discussed in more detail in Sect. 2.2.2.

In summary, the final  $3''.2 \times 1''.9$  resolution mosaic shown in Fig. 1 covers a total area of  $41' \times 41'$  ( $0.47 \text{ deg}^2$ ) and has an average rms of  $20 \mu\text{Jy beam}^{-1}$ . The average rms of the inner  $15' \times 15'$  region is  $10.8 \mu\text{Jy beam}^{-1}$ . Within this central area the noise is nearly Gaussian.

## 2.2.2. Cataloguing

To catalogue all sources present in the mosaic down to a specific signal-to-noise ratio (S/N) we used the AIPS task Search and Destroy (SAD). The SAD task looks for potential sources in an image by identifying agglomerations of pixels (islands) with values above a certain level and then fits Gaussian components to them (see AIPS Cookbook<sup>14</sup>). We ran SAD using both the mosaic and the rms maps as inputs. The latter was generated via the AIPS task RMSD and is used by SAD to determine the local S/N values (see Fig. 2). The size of the box from which the rms is sampled (parameter “insize”) was set to  $80 \times 80$  pixels and parameters “xinc” and “yinc”, which define the increments for calculating rms, were set to 10 pixels. The task SAD was configured with the following parameters: “ngauss” is the maximum number of components to look for was set to 10 000, “cparm” defines the S/N levels down to which the task will look for potential sources and was set to 50 10 7 4.5, “dparm(9)” was set to 2.1 to set the units of “cparm” levels to S/N, “doall” was set to  $-1$  to fit only one Gaussian per island. The parameter “dparm(7)”, which defines total residual flux in the fitting box, was set to  $85 \mu\text{Jy}$ . All other parameters were kept at default values. This procedure yielded 311 radio source components down to 4.5 times the local rms, all of which were visually inspected and compared to

CFHTLS optical images of field W1 with a  $1''$  position cross-match radius.

A total of 28 components were manually combined into eight multicomponent sources. Their total integrated flux densities were measured using the AIPS task TVSTAT down to  $2\sigma$  contours (while their peak surface brightness values, major and minor axes, position angles, and their respective errors are set to  $-99.99$  values in the catalogue; see below). In all such cases the coordinates of the component that overlapped with the optical detection were taken as the source position.

The fraction of radio and optical matches (within  $1''$  matching radius) is roughly constant at around 80% down to  $6\sigma$ , and then drops to 56% for  $5 < S/N < 6$ . Therefore, in the final catalogue we report only 155 sources above  $6\sigma$ .

To separate resolved from unresolved sources we made use of the ratio of total flux density and peak surface brightness, which is a measure of extendedness of the sources. The ratio  $S_{\text{total}}/S_{\text{peak}}$  as a function of the S/N is shown in Fig. 3, and spreads with decreasing S/N (see, e.g. Bondi et al. 2008). Following the method of Bondi et al. (2008), we fitted a lower envelope containing  $\approx 95\%$  of the sources with  $S_{\text{peak}} > S_{\text{total}}$ . This envelope was then mirrored above the line defined by  $S_{\text{total}}/S_{\text{peak}} = 1$ , giving an equation for the upper envelope

$$\frac{S_{\text{total}}}{S_{\text{peak}}} = 1 + \frac{4}{\left(\frac{S_{\text{peak}}}{\text{rms}}\right)^{1.1}}. \quad (1)$$

Sources in Fig. 3 above the upper envelope are considered resolved. For unresolved sources the values of total flux densities are then set equal to their peak surface brightnesses. Lengths of major and minor axes, and position angles (with their respective errors) are set to 0.0.

A total of 25 (130) sources were considered resolved (unresolved). A sample page of the catalogue is given in Table 2. The Table 2 shows a sample radio catalogue page and the relevant parameters. In summary, 155 sources are catalogued, 25 of which are considered resolved, 8 are flagged as multicomponent, and 11 are positioned on a sidelobe.

## 2.2.3. Multiwavelength properties of the 3 GHz sources

We found 123 optical counterparts of the 155 3 GHz sources using the CFHTLS W1 field catalogue and a  $1''$  matching radius, corresponding to 79% of the sources. All sources with optical counterparts are visible across all CFHTLS bands ( $u^*$ ,  $g$ ,  $r$ ,  $i$ ,  $z$ ). For comparison, a correlation of 155 randomly distributed artificial sources across the field yielded nine matches with the CFHTLS catalogue within  $1''$  matching radius. This sets the false-match probability to  $\leq 6\%$ .

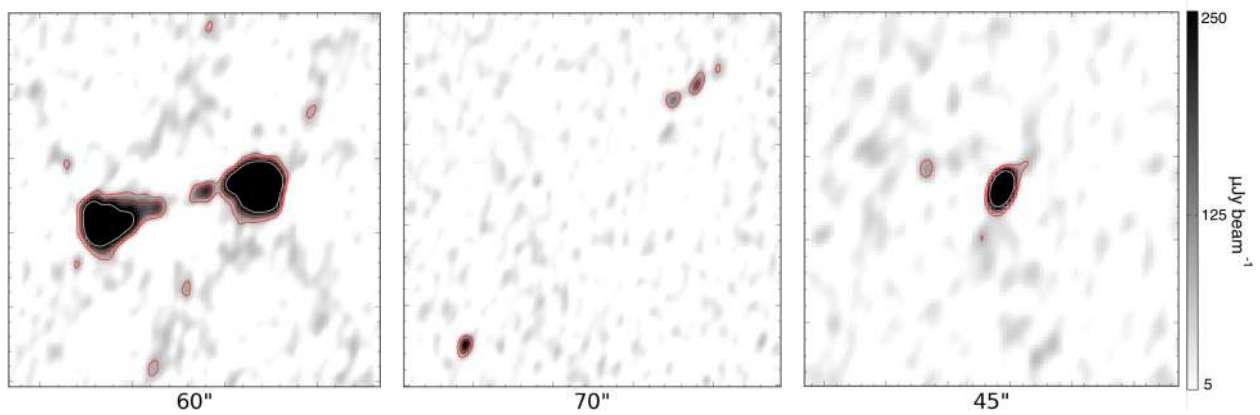
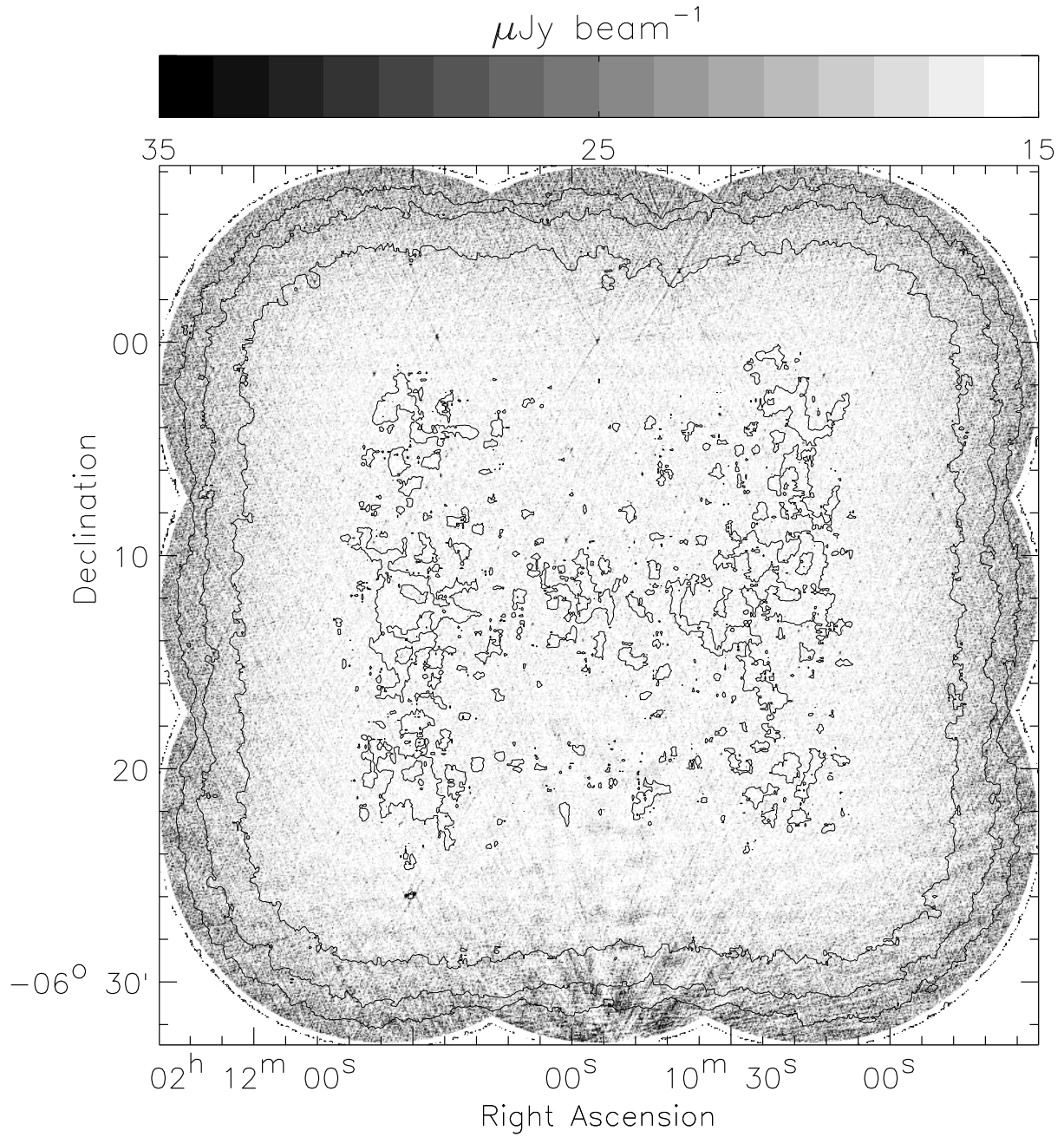
The  $i$ -band magnitudes of the matched sources are in the range  $\sim 15$  to  $\sim 25$  with an average  $\bar{i} \approx 21.3$ . Redshifts range from  $\sim 0.3$  to  $\sim 1.2$  and have an average of  $\bar{z}_{\text{phot}} \approx 0.8$ .

We obtained 1.4 GHz information for our 3 GHz sources using the Faint Images of the Radio Sky at Twenty-cm (FIRST<sup>15</sup>). The FIRST survey is an all-sky survey conducted at 1.4 GHz; it has angular resolution of  $5''$  (Becker et al. 1995) and a typical rms of  $0.15 \text{ mJy}$ . For comparison, the average beam size in the VLA-XXL survey was  $\sim 2''.6$  with an average rms of  $\sim 20 \mu\text{Jy}$ , giving  $\sim 4$  times better sensitivity (for an assumed spectral index of  $\alpha = -0.8$ ) and a factor of  $\sim 2$  better resolution. The FIRST source catalogue lists sources brighter than  $0.75 \text{ mJy}$ .

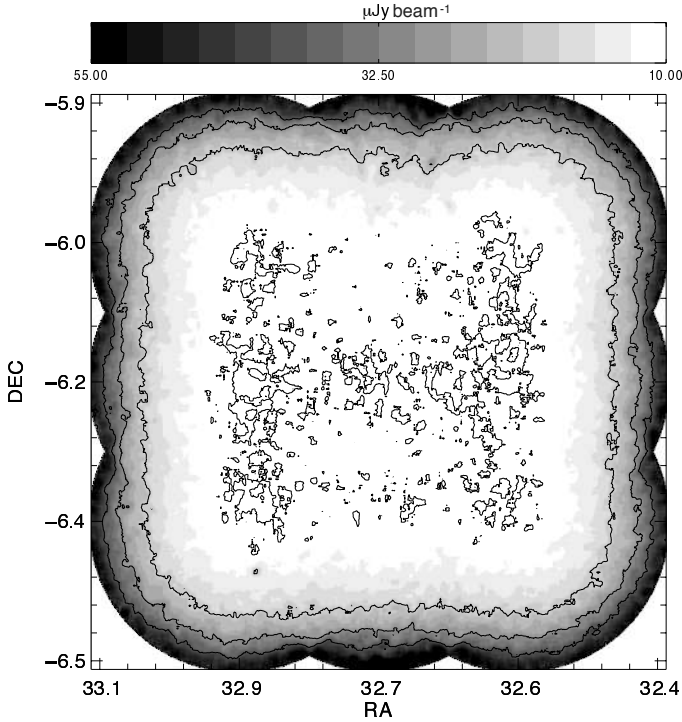
<sup>13</sup> <http://www.aips.nrao.edu/>

<sup>14</sup> <http://www.aips.nrao.edu/cook.html>

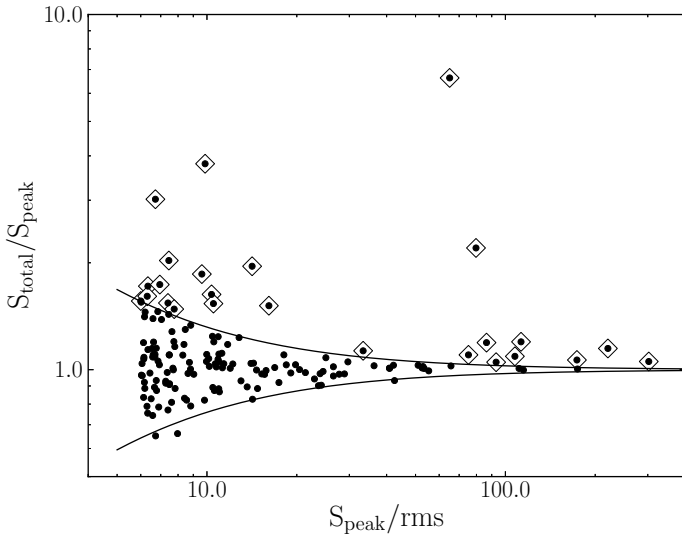
<sup>15</sup> <http://sundog.stsci.edu/>



**Fig. 1.** *Top:* grayscale mosaic of the area of XXL-N field observed with the VLA at 3 GHz with rms contours overlaid (contour levels are 10, 20, 30 and 40  $\mu\text{Jy beam}^{-1}$ ). *Bottom:* three examples of single- and multicomponent sources represented as grayscale maps with contours overlaid. The contour levels are  $2^i\sigma$  ( $i = 2, 3, \text{ and } 4$ ,  $1\sigma = 11 \mu\text{Jy beam}^{-1}$ ).

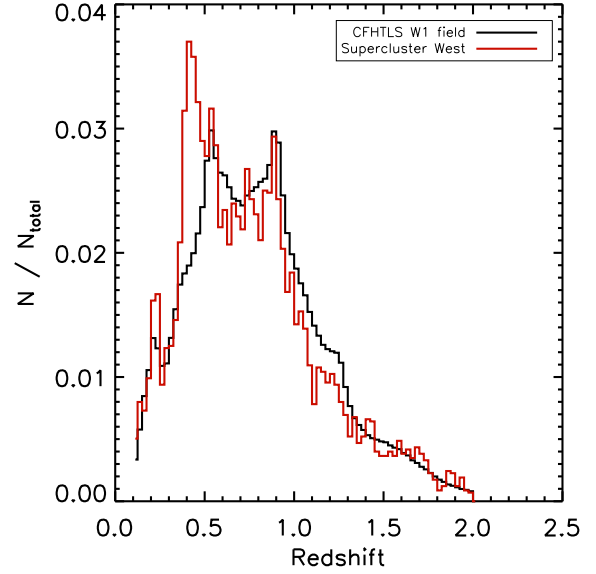


**Fig. 2.** Noise map of VLA – XXL-N mosaic generated using the AIPS task RMSD. The rms contour levels are at 10, 20, 30 and 40  $\mu\text{Jy beam}^{-1}$ . The grey scale (in units of  $\mu\text{Jy beam}^{-1}$ ) is shown at the top.



**Fig. 3.** Ratio of 3 GHz total flux densities and peak surface brightnesses as a function of the signal-to-noise ratio. Sources above (below) the upper envelope are considered resolved (unresolved). Resolved sources are indicated by diamonds. See Sect. 2.2.2 for details.

By positionally matching our 3 GHz source catalogue with the FIRST Survey Catalog (Helfand et al. 2015) using a matching radius of  $2''.5$  we find 21 matches. Of these 21 sources, 16 are singlecomponent at 3 GHz. We find a mean spectral index of  $\alpha = -0.57$  with a standard deviation of  $\sigma = 0.51$  between the 3 and 1.4 GHz fluxes for these singlecomponent sources (assuming  $S_\nu \propto \nu^\alpha$ , where  $S_\nu$  is the radio flux density at frequency  $\nu$ , and  $\alpha$  is the spectral index). These values are expected for radio sources at these flux levels (e.g. Kimball & Ivezić 2008).



**Fig. 4.** Normalised redshift distribution of galaxies in CFHTLS field W1 (black) and the region of the XXL-N field containing the six westernmost X-ray identified clusters (red).

### 3. Analysis of the supercluster

In this section we analyse the first supercluster (XLSSC-e) identified in the XXL-N field (Paper VII). We use Voronoi tessellation analysis (VTA) and the CFHTLS photometric redshifts (Sect. 3.1), as well as the radio properties of the VTA-identified potential supercluster members (Sect. 3.2).

#### 3.1. Optical properties

##### 3.1.1. Identification of overdensities: Voronoi tessellation analysis

To identify the photometric redshift depth that the supercluster occupies, in Fig. 4 we compare the photometric redshift distribution data from the CFHTLS T0007 Data Release (described in Ilbert et al. 2006; Coupon et al. 2009; see also catalogue documentation<sup>16</sup>), between the area encompassing the six westernmost X-ray clusters and the entire field. The comparison indicates an excess in the number of galaxies at  $z \approx 0.43$  in this area. In Fig. 5 we show the spatial distribution of galaxies in this cluster area for six  $\Delta z = 0.15$  wide redshift bins in the range  $z_{\text{phot}} = 0.2-1.1$ . Consistent with the redshift peak at  $z \sim 0.43$  we find the strongest clustering of galaxies in the redshift bin  $z_{\text{phot}} = 0.35-0.50$  (cf. Coupon et al. 2009). This range corresponds to a width of  $3.4\sigma_{\Delta z/(1+z)}(1+z)$  at  $z = 0.43$ , and we consider this redshift range in the analysis.

To identify potential supercluster galaxy members we perform a VTA in the area of interest for CFHTLS W1 galaxies with  $z_{\text{phot}} = 0.35-0.50$ . In this technique (used in van de Weygaert & Icke 1989; and also recently in e.g. Smolčić et al. 2007; Oklopčić et al. 2011; Jelić et al. 2012) we divide the area into polygons and calculate their surface area. Each polygon contains only one galaxy with sides equidistant from the nearest neighbouring galaxies and from the galaxy within. This means the inverse of the surface of these polygons,  $A^{-1}$ , is proportional to the local galaxy density  $\rho_{\text{local}}$ , i.e.

<sup>16</sup> [http://cesam.lam.fr/cfhtls-zphotos/files/cfhtls\\_wide\\_T007\\_v1.2\\_Oct2012.pdf](http://cesam.lam.fr/cfhtls-zphotos/files/cfhtls_wide_T007_v1.2_Oct2012.pdf)

**Table 2.** Sample radio catalogue page.

ID	RA	Dec	$\sigma_{RA}$	$\sigma_{Dec}$	$S_{peak}$	$\sigma_{peak}$	$S_{total}$	$\sigma_{total}$	Res	SL	Multi	Maj	$\sigma_{Maj}$	Min	$\sigma_{Min}$	PA	$\sigma_{PA}$
(1)	[h m s]	[° ' '']	['']	['']	[ $\mu\text{Jy beam}^{-1}$ ]	[ $\mu\text{Jy beam}^{-1}$ ]	[ $\mu\text{Jy}$ ]	[ $\mu\text{Jy}$ ]	(10)	(11)	(12)	[']	[']	[']	[']	[°]	[°]
1	02 09 41.78	-06 19 10.56	0.18	0.43	199	30	601	42	1	0	0	6.7	1.0	2.7	0.4	172.4	6.0
2	02 09 42.38	-05 59 42.66	0.05	0.08	494	29	494	29	0	0	0	0.0	0.0	0.0	0.0	0.0	0.0
3	02 09 45.62	-06 18 43.14	0.02	0.03	1229	23	1229	23	0	0	0	0.0	0.0	0.0	0.0	0.0	0.0
4	02 09 47.75	-06 20 53.02	0.15	0.19	134	20	134	20	0	0	0	0.0	0.0	0.0	0.0	0.0	0.0
5	02 09 48.68	-06 29 41.37	0.09	0.12	453	42	453	42	0	0	0	0.0	0.0	0.0	0.0	0.0	0.0
6	02 09 51.03	-06 06 20.77	0.02	0.03	928	18	928	18	0	0	0	0.0	0.0	0.0	0.0	0.0	0.0
7	02 09 51.90	-06 07 12.60	0.12	0.26	158	16	602	52	1	0	0	6.1	0.6	2.6	0.3	10.3	4.0
8	02 09 54.42	-05 54 54.17	0.13	0.22	184	30	184	30	0	0	0	0.0	0.0	0.0	0.0	0.0	0.0
9	02 09 56.19	-06 09 16.01	0.07	0.09	235	15	356	56	1	0	0	3.4	0.2	2.7	0.2	163.7	11.0
10	02 09 56.43	-06 04 57.98	0.20	0.27	80	13	125	32	1	0	0	4.1	0.7	2.3	0.4	151.7	11.0

**Notes.** Column (1) gives the ID of the source from the VLA radio catalogue, Cols. (2) and (3) represent the coordinates of the source. Position errors in arcseconds, as calculated using the AIPS task SAD, are shown in Cols. (4) and (5). Peak surface brightness and its error are given in Cols. (6) and (7) in units of  $\mu\text{Jy beam}^{-1}$ , followed by total flux density in Col. (8) and its error in Col. (9), both given in  $\mu\text{Jy}$ . A flag denoting a resolved source is stated in column (Res = 1 if the source is resolved, 0 otherwise) (10), while a flag for a source located on a sidelobe is given in Col. (11). Sources found on a sidelobe by visual inspection were flagged with Sidelobe = 1. A flag indicating a source composed of multiple components is stated in Col. (12) (Multi = 1 if the source is multicomponent, 0 if single component). The major axis FWHM of the deconvolved Gaussian component and its error are given in Cols. (13) and (14) in arcseconds, respectively. Column (15) gives the minor axis FWHM of the fitted Gaussian with its error in Col. (16) in arcseconds. The position angle of the major axis of the fit Gaussian and its error are shown in (17) and (18) in degrees, respectively. In total, 25 sources are considered resolved, eight have been flagged as multicomponent and 11 as positioned on a sidelobe. The full catalogue is available as a queryable database table XXL\_VLA\_15 via the XXL Master Catalogue browser <http://cosmosdb.iasf-milano.inaf.it/XXL>. A copy is also available at the CDS.

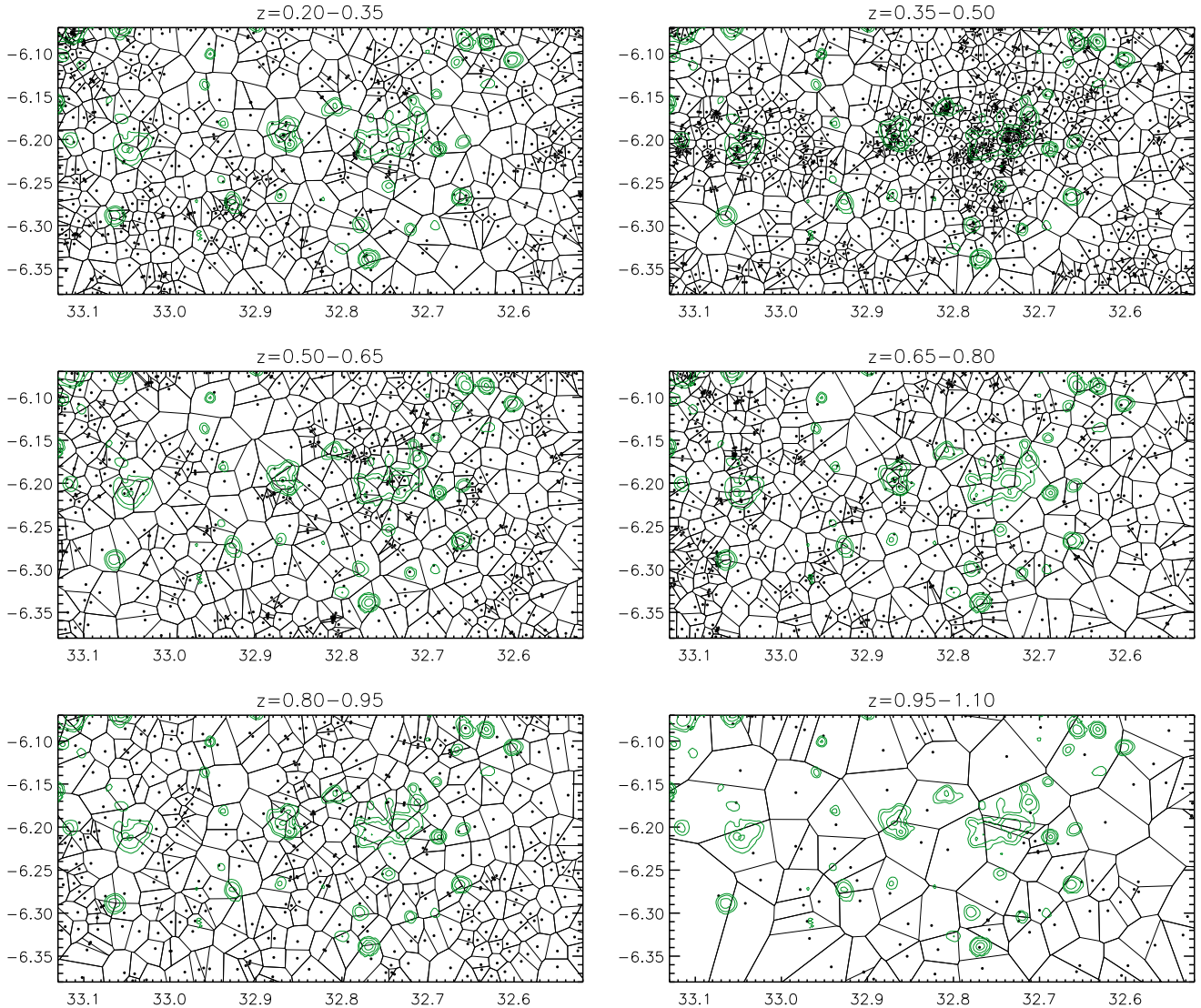
$\rho_{local} \propto A^{-1}$ . A robust “local” galaxy density threshold value then identifies “overdense” regions. To determine this threshold we performed the VTA on an area  $\sim 10$  times larger than the area of interest. This area was populated by  $N$  randomly distributed artificially generated galaxies. The number of galaxies  $N$  is equal to the number of galaxies with  $z_{phot} = 0.35-0.50$  within the CFHTLS catalogue in the same area. Based on the VTA on this field we calculated the average field number density, i.e. the median of the  $A_i^{-1}$  distribution where  $A_i$  is the area of a given Voronoi cell. We note that the CFHTLS images of this region contained masked areas, for example around bright stars or artefacts. The impact of these masks on our density estimates was compensated for by scaling the overdensity criterion threshold by a factor equal to  $(A_{Total} - A_{Masked})/A_{Total}$ , which equates to  $\sim 93\%$ . The threshold was chosen as the 92% quantile of the cumulative distribution of galaxy densities in the mock catalogues (see Fig. 6) and equates to 16 357 galaxies per square degree. We then flagged as overdense those cells for which  $A_i^{-1}$  is greater than this threshold. Galaxies within such flagged cells are considered to be potential members of overdense structures.

The results of the VTA at  $z_{phot} = 0.35-0.50$  are shown in Fig. 7 where we indicate only potential overdense structure galaxy members. The regions from which galaxies were sampled were chosen so as to best encompass the VTA galaxy member candidates located either within the X-ray contours or within the VTA-identified overdensity structures (making sure that each galaxy was counted only in one overdensity). We show an area larger than encompassed by the X-ray emission of the seven identified X-ray clusters to quantify potential additional overdense structures that may be undetected in the X-rays. In addition to the clusters already identified in the X-rays (labelled in Fig. 7; see also Fig. 1 in Paper VII), we identify ten additional galaxy overdense structures, labelled VTA01 to VTA10 in the figure. The overdense structures seem to assemble into two agglomerations (eastern and western), with an extended potential overdense structure in between (VTA07 in Fig. 7). We note,

however, that the area between these two agglomerations contains an optically masked area (a vertical strip at RA  $\sim 32^{\circ}91$ ; cf. Fig. 1 in Paper VII). Thus, no clustering information could be retrieved for this part of the sky, which could contain additional overdense structures connecting the two agglomerations. The VTA may be not effective in detecting highly elongated, filamentary overdense structures. Since the structures in the field studied here are relatively close to each other, we do not expect this feature to influence our results significantly. Our VTA analysis based on photometric redshifts clearly identifies the clusters, independently detected via extended X-ray emission. This assures that the VTA yields robust group/cluster candidates, as further verified in the next section.

### 3.1.2. Verification of the VTA-identified overdensities

We test the probability of finding an overdensity caused by random density fluctuations using our te mock catalogues of randomly distributed galaxies across an area that is  $\sim 10$  times larger than that occupied by the supercluster. By tiling the plane with non-overlapping circular regions of a fixed radius, we identify regions with  $N = 1, 2, 3, \dots$  simulated overdense galaxies. We perform this analysis for three different circular areas with radii  $r = 30'', 60'',$  and  $100''$ . The result is then normalised by the effective area to derive the expected number of false positive detections of overdense regions. In Fig. 8 we show the expected number of false positive detections with  $N$  galaxies within  $r = 30'', 60'',$  and  $100''$  in our supercluster area. Since our overdense regions contain five or more galaxies within an effective radius of  $\gtrsim 30''$ , we expect to find less than one such overdense region that is the result of random density fluctuations in our field. This is in agreement with Ramella et al. (2001) who set the minimum number of closely separated galaxies to five, as the threshold for cluster detection, rendering the effect of Poisson fluctuation negligible.



**Fig. 5.** Voronoi tessellation diagrams of the region encompassing the five westernmost X-ray clusters (RA: 32.52, 33.13, Dec:  $-6.38$ ,  $-6.05$ ) as a function of redshift ( $0.2 < z < 1.1$ ) with Voronoi polygons shown with lines and isophotes of X-ray surface brightness in contours. Each polygon contains only one galaxy, indicated by a black dot. The largest number of sources in the region of X-ray emissions are located in the redshift range  $0.35 < z < 0.50$ . The levels of the X-ray contours are shown in green at 5.28, 7.02, 12.54 and 30  $\text{cts s}^{-1} \text{deg}^{-2}$  in the  $[0.5-2.5]$  keV range.

To assess the significance of Malmquist bias on the detection of overdense structures we calculated absolute magnitudes for galaxies in the studied redshift range of  $0.35-0.5$ . These absolute magnitudes are shown as a function of redshift in the upper panel of Fig. 9. We note a drop in the absolute magnitude limit from  $M_i \sim -18$  at  $z = 0.35$  to  $M_i \sim -19$  at  $z = 0.5$ . Cutting the optical catalogue at  $M_i = -19$  instead of  $m_i = 23.5$ , we find the same overdense regions, although with a slightly lower number of galaxies in each of them (lower panel in Fig. 9). Thus, we conclude that Malmquist bias does not affect the results of the overdensity detections presented here. Finally, we performed an a posteriori verification of the validity of choosing the redshift range  $0.35 < z < 0.5$ . In the first step, we concatenated the CFHTLS photometric redshift and the CFHTLS photometric catalogue for an area  $\sim 10$  times the size of the supercluster. Using this catalogue we performed the VTA (as described in Sect. 3.1.1) for various  $dz = 0.15$  wide redshift ranges centred at  $z = 0.225 + k \times 0.1$ , where  $k = 0, 1, 2, 3, 4, 5$ . In each redshift range the number of overdense galaxies within the  $z = 0.35-0.5$  identified overdensity regions were then identified.

We find that all 17 overdense regions peak in the chosen redshift bin ( $0.35-0.5$ ), suggesting this is the optimal range to describe all the overdense regions. However, to accurately report the mean redshift of each overdense region and the corresponding standard deviation we analyse each region within its own extent in redshift. We consider the region extended in redshift if the number of galaxies in redshift bins adjacent to the peak is within the Poissonian counting error of the peak value. We then use the extended redshift range to determine the mean redshift of overdense galaxies in each VTA region. Redshifts used to calculate the  $\bar{z}_{\text{phot}}$  and its standard deviation for all overdense regions (including XLSSC clusters) are sampled from the range  $0.35-0.5$ , except for VTA02, VTA09, and VTA10. These overdense regions show a broader distribution in redshift, as determined by the criterion described above. Specifically, VTA02 was extended to  $0.25 < z < 0.5$ , VTA09 to  $0.35 < z < 0.6$ , and VTA10 to  $0.25 < z < 0.5$ . The mean photometric redshifts and the corresponding standard deviations of our VTA-identified overdense structure member candidates are listed in Table 3. They are in agreement with the spectroscopic redshifts of all of the brightest

**Table 3.** List of VTA-detected cluster/group candidates.

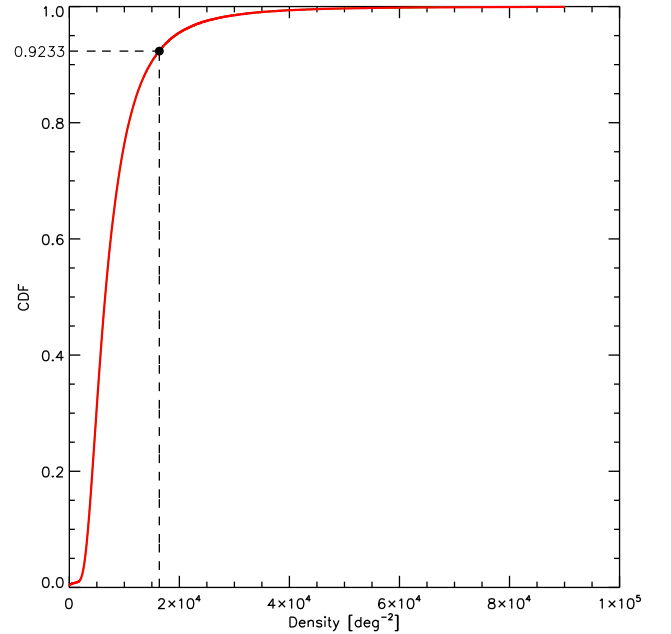
ID	RA	Dec	$\bar{z}_{\text{phot}}$	$\sigma_{\text{phot}}$	$\bar{z}_{\text{spec}}$	$L_{300 \text{ kpc}}^{\text{XXL}}$	$kT_{300 \text{ kpc}}$	$r_{500}$	$M_{500,MT}$
(1)	[h m s] (2)	[d m s] (3)	(4)	(5)	(6)	$[\times 10^{42} \text{ erg s}^{-1}]$ (7)	[keV] (8)	[kpc] (9)	$[\times 10^{13} M_{\odot}]$ (10)
XLSSC 081	02 12 12*	-06 12 39*	0.42	0.04	0.428*	14.9*	1.7*	549*	7*
XLSSC 082	02 10 51*	-06 10 44*	0.42	0.04	0.424*	17.1*	3.9*	878*	29*
XLSSC 083	02 10 56*	-06 11 54*	0.42	0.03	0.430*	31.3*	4.8*	990*	41*
XLSSC 084	02 11 03*	-06 12 47*	0.42	0.04	0.430*	13.8*	4.5*	955*	37*
XLSSC 085	02 11 29*	-06 11 47*	0.43	0.03	0.428*	28.3*	4.8*	985*	41*
XLSSC 086	02 11 14*	-06 09 44*	0.41	0.04	0.424*	11.2*	2.6*	698*	15*
XLSSC 099	02 12 53	-06 12 07	0.43	0.04	0.3911	14.4	5.1	1038	47
VTA01	02 11 10	-06 13 22	0.40	0.04	...	<7.3	<1.8	<560.6	<7.7
VTA02	02 10 58	-06 13 53	0.39	0.05	...	<9.3	<1.9	<590.2	<9.0
VTA03	02 11 01	-06 15 50	0.43	0.04	...	<5.9	<1.6	<535.2	<6.7
VTA04	02 12 18	-06 10 30	0.41	0.03	...	<9.1	<1.9	<586.9	<8.8
VTA05	02 12 20	-06 12 01	0.42	0.05	...	<9.9	<2.0	<597.1	<9.3
VTA06	02 12 30	-06 12 13	0.43	0.03	...	<11.7	<2.1	<618.9	<10.3
VTA07	02 11 50	-06 13 28	0.43	0.03	...	<7.5	<1.8	<563.3	<7.8
VTA08	02 12 01	-06 09 36	0.43	0.04	...	<8.83	<2.04	<606.6	<9.7
VTA09	02 13 06	-06 14 13	0.47	0.06	...	<8.8	<1.9	<582.7	<8.6
VTA10	02 13 01	-06 08 24	0.39	0.06	...	<11.9	<2.30	<649.7	<11.9

**Notes.** For clusters detected in the X-rays the X-ray ID (XLSSC) is adopted. Cluster candidates with unavailable spectroscopic redshift are marked by ellipses (...). Central right ascension and declination of regions shown in Figs. 7, 11 and 12 are given in Cols. (2) and (3). For the galaxies satisfying the VTA overdensity threshold and potentially belonging to the given structure as suggested by the circular or elliptical areas in Figs. 7, 11 and 12, the mean photometric redshift and its standard deviation are given in Cols. (4) and (5), respectively. The spectroscopic redshift is given in Col. (6) (where available). Where available, X-ray luminosity in units of  $\text{erg s}^{-1}$ , temperature in keV,  $r_{500}$  in kpc and  $M_{500,MT}$  in units of solar mass ( $M_{\odot}$ ) are given in Cols. (7)–(10), respectively (as described in Paper VII; Paper III).  $M_{500,MT}$  for XLSSC clusters has been derived in Paper IV using  $M_{\text{WL}} - T$  relation. For X-ray non-detected VTA structures,  $3\sigma$  upper limits are given (see Sect. 4 for details). (\*) Taken from Paper VII.

cluster galaxies determined in Paper VII, also listed in Table 3. Spectroscopic redshift confirmation of the XLSSC 081–086 was performed on spectra of at least three galaxies (up to five) per cluster (Pompei et al. 2016). However, the spectroscopic redshift of the XLSSC 099 was determined from the spectrum of only one galaxy in the cluster (the brightest) using a spectrum from the GAMA<sup>17</sup> survey. The median redshift error is  $\sim 33 \text{ km s}^{-1}$  for the  $z_{\text{spec}}$  of the GAMA objects (Baldry et al. 2014), rendering the  $z_{\text{spec}}$  error negligible. The photometric redshift of this bright galaxy ( $i = 17.88$ ) is estimated to be  $z_{\text{phot}} = 0.3615$ , and is thus consistent with its spectroscopic redshift, since  $\sigma_{\Delta z/(1+z_s)} \sim 0.030\text{--}0.032$  for sources with  $i < 21.5$ . Our analysis suggests that the mean overdensity redshift is slightly shifted relative to the redshift of this galaxy.

### 3.1.3. Galaxy colour properties

We separate blue from red galaxies using a threshold of  $g - r = 1.17$ . This value corresponds to the inflection point of the double-Gaussian function best fit to the  $g - r$  histogram. The  $g - r$  versus  $r$  colour-magnitude diagram for the VTA-identified potential supercluster galaxy members is shown in Fig. 10. In Fig. 11 we show the spatial distribution of the red and blue potential supercluster galaxy members, overlaid with X-ray contours. In virialised clusters, red galaxies occupy the centre, while the blue galaxies occupy the outskirts of the cluster. As can be seen from the top and bottom panels of Fig. 11, this does not appear to be the case here in all identified overdense structures. This will be further discussed in Sect. 4.

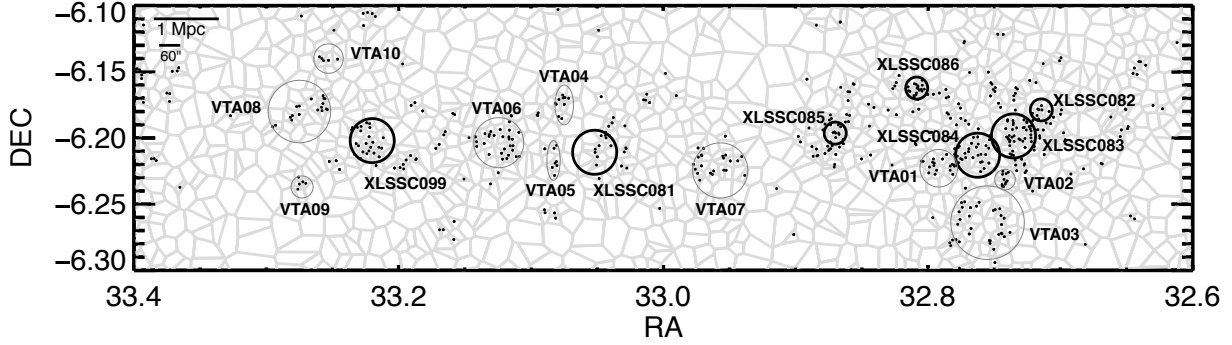


**Fig. 6.** Cumulative distribution function (CDF) of galaxy densities in ten simulated catalogues with randomly distributed artificially generated galaxies. The overdensity threshold used in this paper corresponds to the 92% quantile of the distribution. Galaxies populating Voronoi cells with density above this value are treated as overdense region/cluster member candidates.

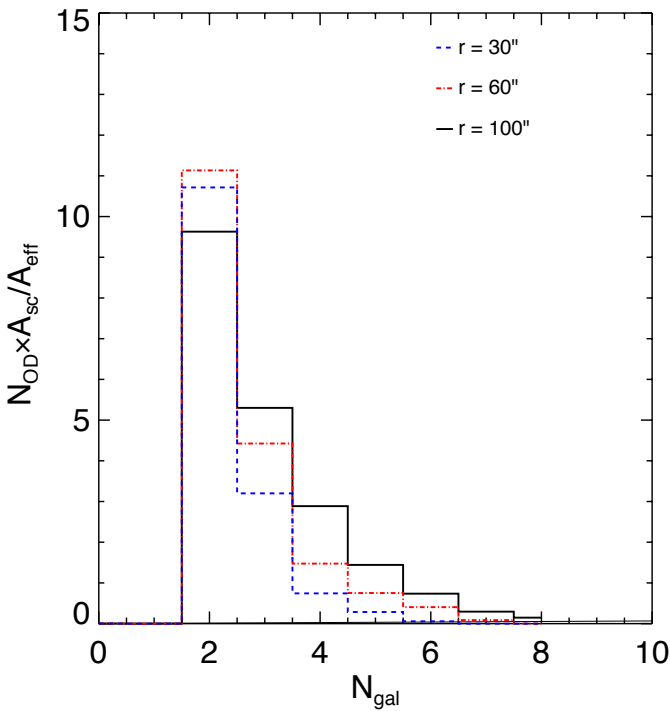
### 3.2. Radio properties

Our VLA-XXL-N mosaic covers only the western part of the supercluster (XLSSC 081–086, VTA01–03, and VTA07). The

<sup>17</sup> <http://www.gama-survey.org/>



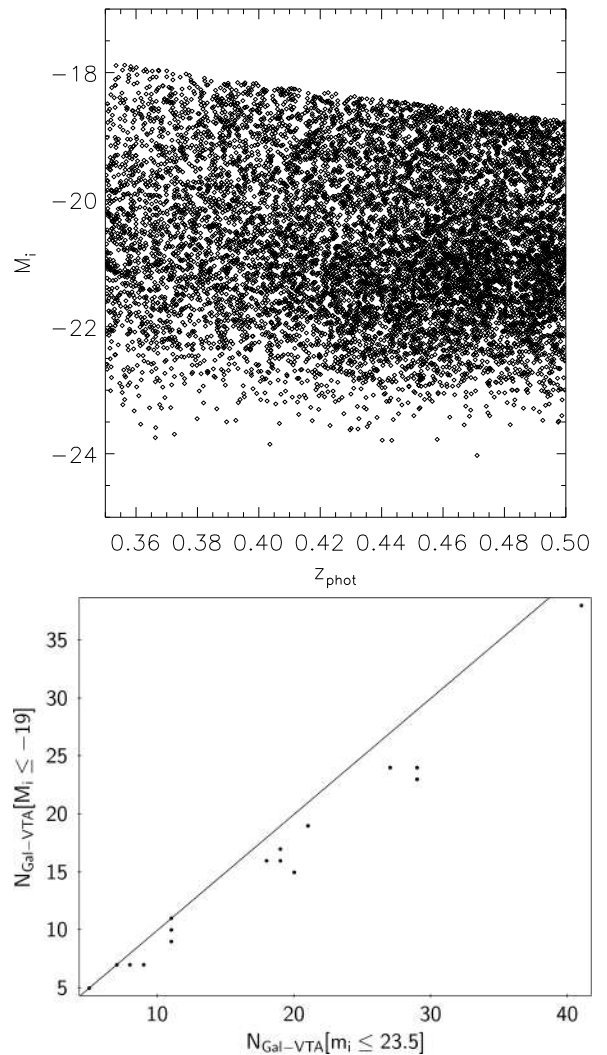
**Fig. 7.** Voronoi tessellation analysis of the area around the X-ray identified cluster structure at  $0.35 < z_{\text{phot}} < 0.50$  (lines). Galaxies associated with overdense structures identified via our VTA analysis (see text for details) are shown as black dots. A projected scale of 1 Mpc is indicated in the top left corner. The X-ray identified clusters are marked with thick black circles, and VTA identified overdense structures are marked with thin grey circles. Galaxies within these regions were sampled to statistically estimate the properties of clusters and overdense structures. (See Table 3 for details.)



**Fig. 8.** Expected number of circular overdense regions detected due to random density fluctuation as a function of the number of galaxies in these regions. The estimates are provided for three different radii, as indicated in the legend.

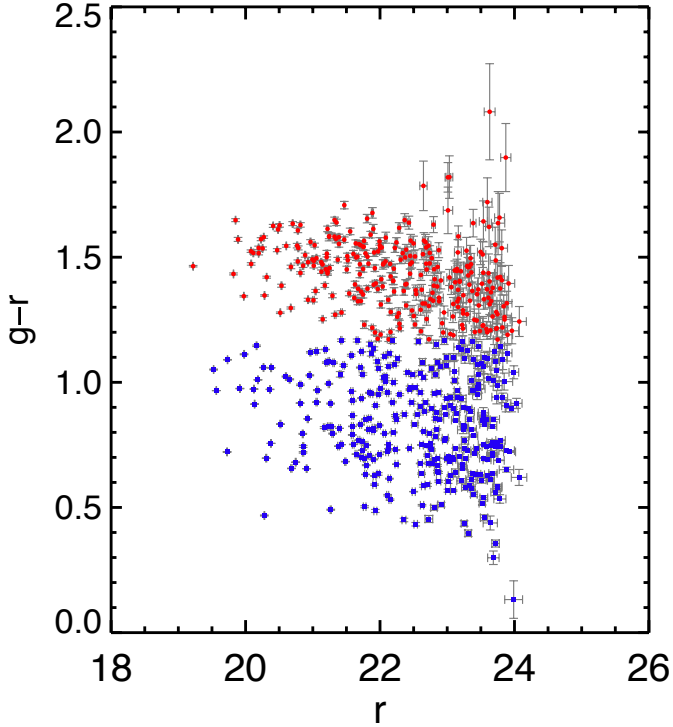
X-ray cluster XLSSC 081 is on the very edge of the mosaic and encompasses one source detected in VLA-XXL-N radio observations (VLA-XXL ID = 154), but without an optical counterpart. VTA04-09 and XLSSC 099 are outside the field-of-view of our 3 GHz VLA-XXL-N mosaic. Thus, for the eastern part of the supercluster we search for radio sources associated with CFHTLS galaxies at  $0.35 < z_{\text{phot}} < 0.50$  using the FIRST and NVSS 1.4 GHz survey catalogues (Becker et al. 1995; Condon et al. 1998). We find no radio sources or radio galaxies associated with galaxies at  $0.35 < z_{\text{phot}} < 0.50$  in the area of interest.

In Fig. 12 we show our 3 GHz radio mosaic of the area hosting the western part of the supercluster. In this area we find eight 3 GHz radio sources with matched CFHTLS photometric redshifts between 0.35 and 0.50. We show zoomed-in images of



**Fig. 9.** *Top:* absolute  $i$  band magnitude as a function of photometric redshift. In the redshift range 0.35–0.5 a slight decrease can be seen. *Bottom:* comparison of the number of galaxies in each VTA overdense region with a cut in absolute  $i$ -band magnitude ( $M_i \leq -19$ ), and a cut in observed  $i$ -band magnitude applied in this paper ( $m_i \leq 23.5$ ).

the host galaxies in the optical with radio contours overlaid in Fig. 13. No large radio galaxies are found within and around the X-ray identified cluster structure.



**Fig. 10.** Colour-magnitude diagram of the sources in overdense structures within the XXL-N supercluster identified via the Voronoi tessellation analysis (see Fig. 7). Sources are classified as red (i.e. early-type) if their  $g - r \geq 1.17$ , and blue (i.e. late-type) if their  $g - r < 1.17$ .

The properties of the radio sources are listed in Table 4. Four out of the eight sources are hosted by red galaxies, one has a colour consistent with blue galaxies, and two have host galaxy colours consistent with the green valley (upper part of the blue cloud; see Fig. 10). A visual inspection of the host galaxies shows that the colours are, as expected, linked to the morphologies of the galaxies (e.g. [Strateva et al. 2001](#)). The red galaxies are elliptical, while the blue galaxies show spiral features.

The 3 GHz luminosity densities of the sources (assuming a spectral index of  $-0.8$ ) are in the range of  $3 \times 10^{22} - 2 \times 10^{23} \text{ W Hz}^{-1}$ . Three out of the eight radio sources can be associated with the central regions of X-ray clusters or VTA identified overdense structures (VLA-XXL ID = 072 in XLSSC 083, VLA-XXL ID = 075 in VTA02 and VLA-XXL ID = 093 in XLSSC 086). The host galaxies of these are among the four brightest galaxies in the cluster/overdense structure. In particular, the host galaxy of VLA-XXL ID = 075 in the overdense structure VTA02 is the optically brightest galaxy in this structure.

## 4. Discussion

### 4.1. Constraining total masses of the overdensities

In Paper VII the six westernmost X-ray detected clusters in the XLSSC-e supercluster have been spectroscopically confirmed in the region of interest to lie at  $z = 0.42 - 0.43$  within an extent of  $\sim 0.35 \times 0.1$  on the sky (or a physical extent of  $\sim 10 \times 2.9 \text{ Mpc}^2$  at  $z = 0.43$ ). Using the XXL X-ray data they have also computed the virial masses within the radius encompassing 500 times the critical density of the six clusters to be in the range of  $M_{500} = (1 - 3) \times 10^{14} M_{\odot}$ . This yields a total mass of all six of

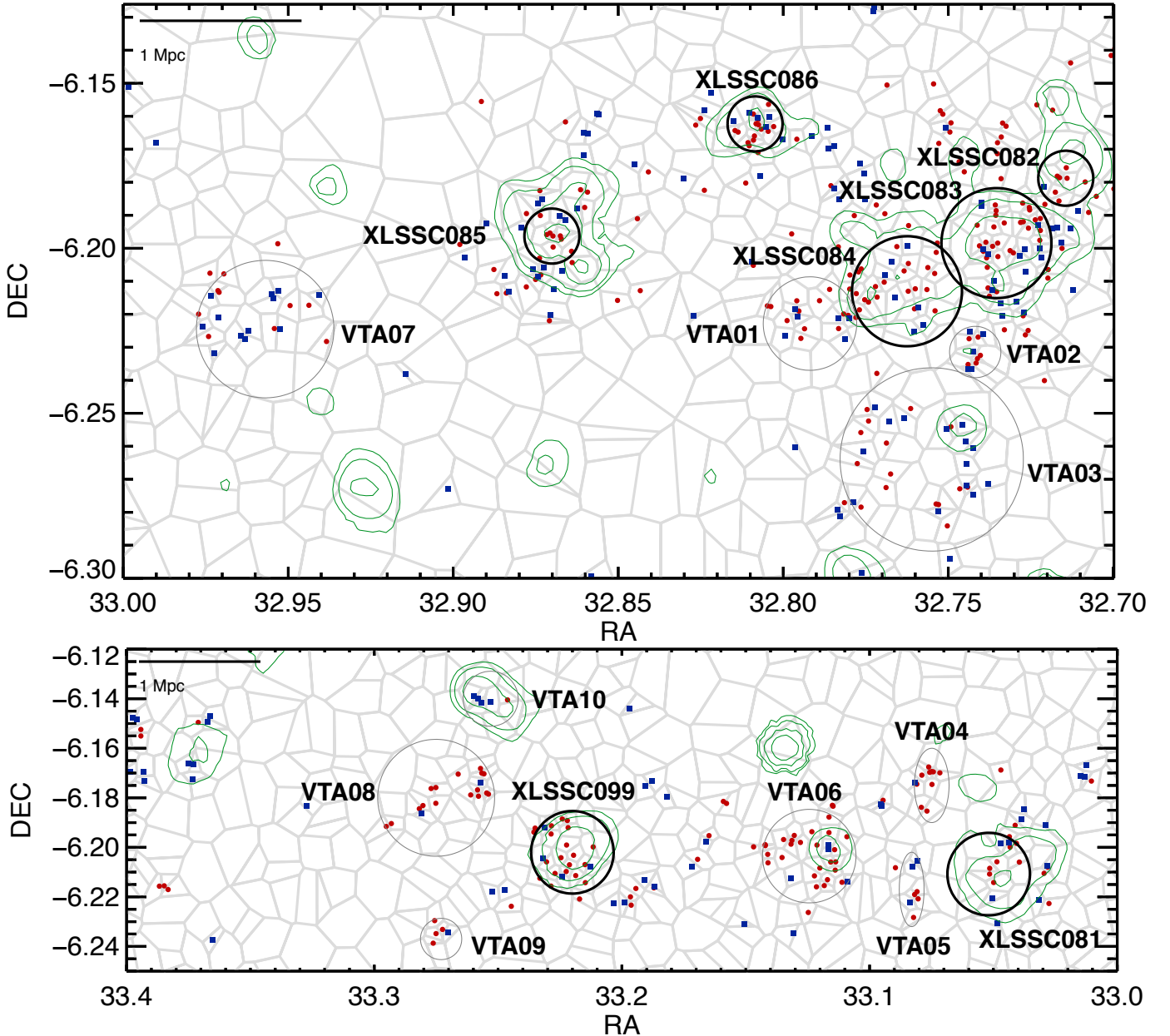
the westernmost X-ray detected clusters of  $M_{500} \sim 10^{15} M_{\odot}$  (see Table 1 in Paper VII; see also Table 3 in Paper II).

Additionally, ten VTA identified overdense structures have not been identified as X-ray clusters, presumably because their X-ray emission falls below the X-ray cluster detection threshold. Nevertheless, for these overdense structures we can put upper limits on their X-ray properties as follows. The upper limits on VTA overdense structures were calculated using upper limits on the count rates in a 300 kpc aperture centred at the middle of the optical structures listed in Table 3. The method consists of the Bayesian approach for calculating X-ray aperture photometry described in detail in Willis et al. (in prep.). It computes values and bounds for the intensity of the source ( $S$ ) using counts and exposure data obtained in source and background apertures. It calculates the background-marginalised posterior probability distribution function of the source intensity ( $S$ ), assuming Poisson likelihoods for the detected number of source counts and background counts in the given exposure time. The mode of this PDF is determined, and the lower and upper bounds of the confidence region are determined by summing values of the PDF alternately above and below the mode until the desired confidence level is attained. When the mode is at  $S = 0$  or the calculation for the lower bound reaches the value  $S = 0$ , only the upper confidence bound is evaluated, which is considered an upper limit. With this upper limit on the count rate in a 300 kpc aperture we used the mass-temperature relation of [Lieu et al. \(2016; hereafter Paper IV\)](#) and the luminosity-temperature relation of [Giles et al. \(2016; hereafter Paper III\)](#) to calculate upper limits on the mass of the VTA overdense structures. Starting with an initial estimate of the temperature of the object, we use the  $M - T$  relation to evaluate the mass and the radius corresponding to the overdense structure of 500 times the critical density. We integrate the given count rate to compute the X-ray luminosity up to  $r_{500}$ , and re-estimate the temperature predicted from the corresponding  $L - T$  relation. We re-iterate until we converge on the value of the temperature. The calculations for the upper limits are given in Table 3. For VTA10 we masked two AGN point sources leaving only 48% of the area of the 300 kpc aperture for the calculation of the upper limit on the flux. As expected, the upper limits on total masses estimated for our VTA-identified overdensities without X-ray detections are consistent with the structures being poor clusters or groups (e.g. [Vajgel et al. 2014](#)). A spectroscopic velocity dispersion of the structures would allow the classification of these structures (e.g. [Dehghan & Johnston-Hollitt 2014](#)). However, no such spectroscopic data is currently available.

### 4.2. Morphology and composition of the supercluster

As is evident from Fig. 11, the X-ray emission of the clusters XLSSC 082, 083, and 084 overlap into a morphologically elongated ensemble in the NW-SE direction. Our VTA analysis reveals multiple overdense structures of optical galaxies in the same field, three of which lie at the same redshift but are not detected in the X-ray (VTA01 – VTA03). Overdense structure VTA01 is located toward the SE from the ensemble revealed in the X-rays, and appears to be a distinct structure. The overdense structures VTA02 and VTA03 are located toward the south of the ensemble. The non-detection of the VTA-identified overdense structures in the X-ray data sets the upper limit to the combined total virial mass of these three overdense structures to  $2.3 \times 10^{14} M_{\odot}$  (see Table 3.).

The spatial distribution of red and blue galaxies in the X-ray ensemble (XLSSC 082, XLSSC 083, XLSSC 084) and VTA01-03 groups appears non-structured, i.e. there is no clear



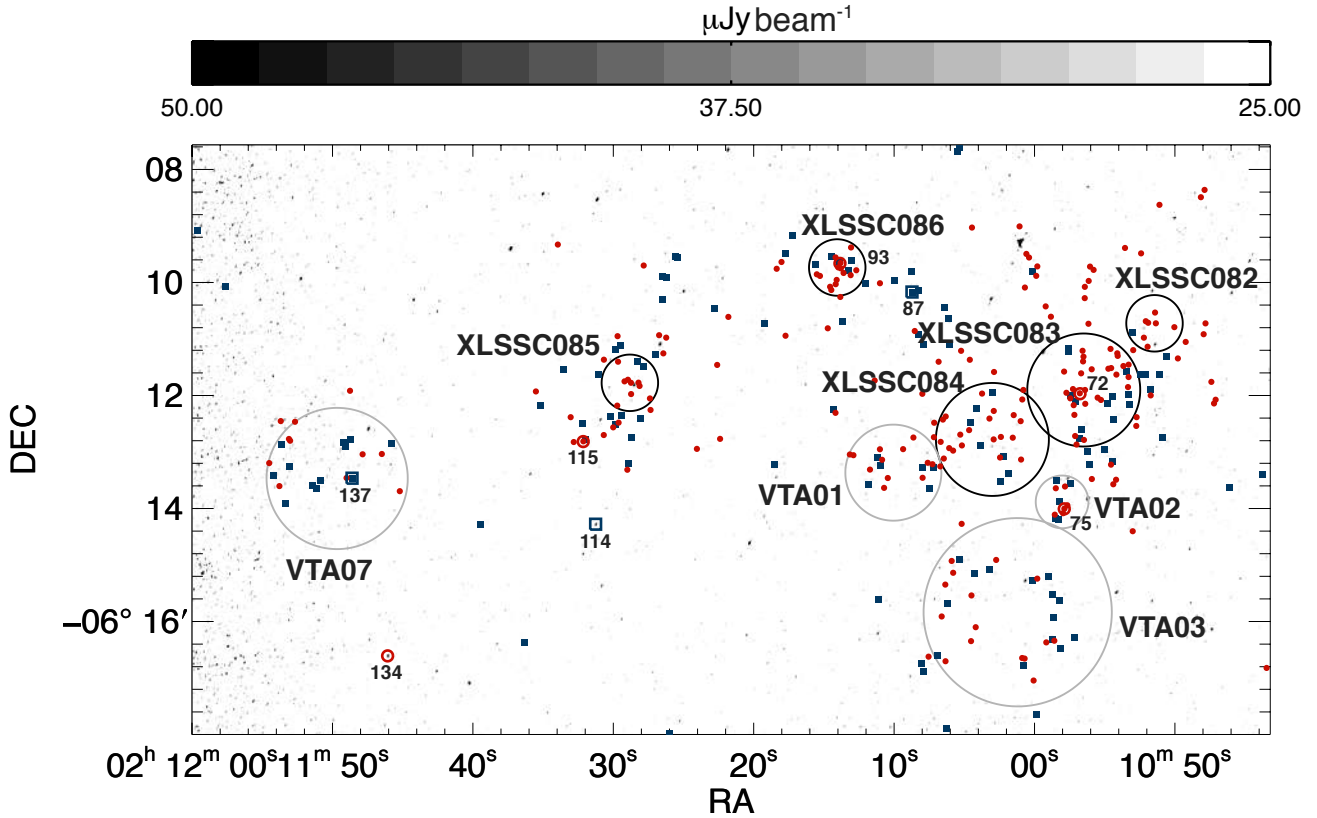
**Fig. 11.** Voronoi tessellation analysis of the western (*top*) and eastern (*bottom*) region of the supercluster at  $0.35 < z < 0.50$  (grey lines). Potential cluster galaxy members, identified via our VTA analysis (see text for details), are separated into red ( $g-r \geq 1.17$ ) and blue ( $g-r < 1.17$ ) galaxies. These are shown by red filled circles and blue filled squares, respectively. The levels of the X-ray contours are shown in green at 5.28, 7.02, 12.54 and 30  $\text{cts s}^{-1} \text{deg}^{-2}$  in the [0.5–2.5] keV range. A scale of 1 Mpc is indicated in the top left corner. The X-ray identified clusters are marked with thick black circles, and VTA identified overdense structures are marked with thin grey circles, and are labelled in the plot. Galaxies within these regions were sampled to statistically estimate the properties of clusters and overdense structures (see text and Table 3 for details).

red-blue galaxy separation with increasing distance from the group centre, which suggests an unvirialised state of the groups. If this is the case, then the calculated virial masses may overestimate the real masses of these cluster candidates.

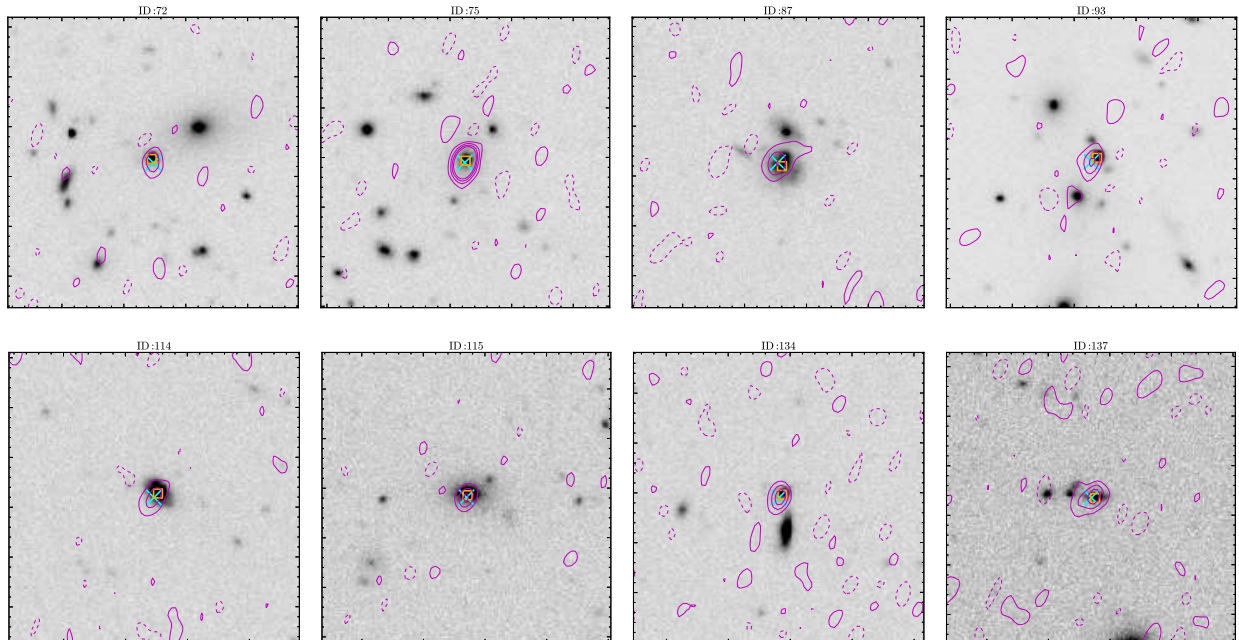
We find a 3 GHz radio source (VLA-XXL ID = 072) associated with the third brightest galaxy within the XLSSC 083 clusters close to the peak of the X-ray emission. Its monochromatic 3 GHz radio luminosity density is  $(4.05 \pm 0.55) \times 10^{22} \text{ W Hz}^{-1}$ , and it is associated with a red-sequence galaxy ( $g-r = 1.6$ ). This suggests that the radio synchrotron emission likely arises from a weak AGN within the galaxy, often found at the bottom of the gravitational potential wells of galaxy clusters (e.g. Smolčić et al. 2008, 2010).

The X-ray, radio and optical data combined, thus suggest that the bottom of the gravitational potential well is within the X-ray peak, and that the six clusters and overdense structures identified in this region (XLSSC 082, 083, 084 and VTA01, 02, 03) are likely in the process of merging and forming a larger structure. Based on the  $M_{500}$  values calculated in Paper VII, and using the upper mass limits given in Table 3, we can estimate that the upper limit to the total mass of the western structure after adding  $M_{500}$  is  $M_{500} \leq 1.3 \times 10^{15} M_{\odot}$ .

The clusters XLSSC 085 and 086 are at a projected distance of 8' and 4/8, respectively, from the XLSSC 083 cluster. This corresponds to a distance of  $\sim 2.7$  and  $\sim 1.6$  Mpc, respectively at  $z = 0.43$ . Our VTA identifies overdense structures associated with both clusters. The red-blue galaxy distribution in



**Fig. 12.** Area of VLA - XXL-N mosaic containing the supercluster. “Overdense” early-type galaxies are shown as red dots and late-type galaxies as blue squares. Galaxies are classified as red if their  $g - r \geq 1.17$ , and blue if  $g - r < 1.17$  (see Fig. 10). Open red circles (blue squares) mark potential red (blue) supercluster member galaxies detected as 3 GHz radio sources (the 3 GHz VLA ID is indicated next to the radio sources). The X-ray identified clusters are marked with thick black circles, and VTA identified overdense structures are marked with thin grey circles, and labelled in the plot. Galaxies within these regions were sampled to statistically estimate the properties of clusters and overdense structures (see Table 3 for details).



**Fig. 13.** Optical images (greyscale; 35'' on the side) of potential supercluster member galaxies also detected at 3 GHz within the VLA-XXL data, overlaid with magenta radio contours. The contours start at  $2\sigma$  ( $1\sigma = 10.8 \mu\text{Jy}$ ) and increase in steps of  $2\sigma$  for all sources except source 75 where they increase in steps of  $4\sigma$ . Dashed negative contours are shown at the same levels as the positive. Cyan crosses show the position of the of the VLA radio sources, orange squares represent the position of the optical counterpart. The VLA-XXL-N source IDs are shown above each panel. North is up, east is left.

**Table 4.** Radio sources with optical counterparts that have redshifts  $0.35 < z_{\text{phot}} < 0.50$  within the area containing the studied overdense structures.

ID	RA	Dec	$S_{\text{peak}}$	$z_{\text{phot}}$	$L_{3 \text{ GHz}}$	$g - r$	Overdense structure ID
(1)	[h m s]	[d m s]	$[\mu\text{Jy beam}^{-1}]$	(5)	$[\times 10^{22} \text{ W/Hz}]$	(7)	(8)
114	02 11 31.27	−06 14 16.84	$59.5 \pm 9.5$	0.41	$3.2 \pm 0.5$	1.09	...
75	02 10 57.88	−06 14 00.62	$287.0 \pm 10.8$	0.41	$16.1 \pm 0.6$	1.56	VTA02
115	02 11 32.15	−06 12 49.11	$74.5 \pm 8.7$	0.40	$3.9 \pm 0.4$	1.62	XLSSC 085
72	02 10 56.76	−06 11 58.37	$72.2 \pm 9.7$	0.41	$4.0 \pm 0.5$	1.59	XLSSC 083
87	02 11 08.73	−06 10 10.38	$66.6 \pm 10.5$	0.49	$5.5 \pm 0.8$	1.02	XLSSC 086
93	02 11 13.86	−06 09 39.96	$74.9 \pm 10.8$	0.40	$4.0 \pm 0.5$	1.28	XLSSC 086
137	02 11 48.62	−06 13 27.91	$80.6 \pm 11.7$	0.42	$4.6 \pm 0.6$	0.71	VTA07
134	02 11 46.09	−06 16 36.59	$91.7 \pm 12.0$	0.50	$7.9 \pm 1.0$	1.50	...

**Notes.** Sources unassociated with a cluster/group candidate are marked by ellipsis. Source identifiers from the VLA-XXL-N 3 GHz radio catalogue are given in Cols. (1), Col. (2) gives the right ascension (RA) of the source in hours, minutes, and seconds, Col. (3) gives the declination (Dec) of the source in degrees, arcminutes and arcseconds. Peak surface brightness is given in Col. (4) in units of  $\mu\text{Jy beam}^{-1}$ . Photometric redshifts of optical counterparts of radio sources are given in Col. (5). Radio luminosities at 3 GHz are given in Col. (6) in units of W/Hz, and are divided by  $10^{22}$ . Column (7) provides  $g - r$  optical colour. If a source is associated with an overdense structure, Col. (8) gives the ID of the source.

XLSSC 086 appears non-structured, and both structures show subclumping and elongated overdense structure features. It is interesting that the brightest galaxy (in the  $r$  band) in XLSSC 086 is associated with a red galaxy toward the south of the X-ray emission, while the next three brightest galaxies (one of them hosting the radio source VLA-XXL ID = 093) are associated with the X-ray centroid. The E-W elongation of the X-ray emission of this cluster, supported by a similar elongation of the distribution of galaxies associated with VTA-identified overdense structures, suggests the existence of two substructures that are in the process of merging.

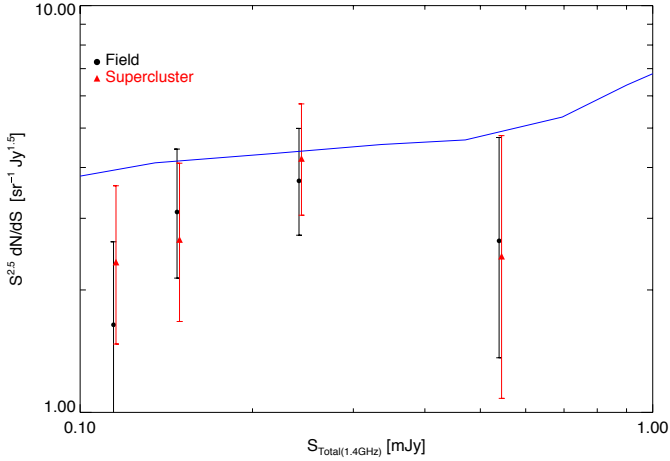
The X-ray detected cluster XLSSC 081 lies  $\sim 0.35$  deg away from XLSSC 083 (see Fig. 11), while the cluster XLSSC 099 is located further to the east,  $\sim 0.43$  deg from XLSSC 083. Furthermore, the VTA associates an elongated overdense structure of galaxies with this cluster. In addition to this cluster, our VTA has revealed seven more overdense structures toward the east of this cluster (VTA04-10), and a potential overdense structure toward the west (VTA07). All overdense structures identified in this region seem to have a non-segregated distribution of red and blue galaxies. This again suggests a dynamically young state for the structures. The X-ray emission coincident with VTA10 originates from two background X-ray sources, spectroscopically found to lie at  $z_{\text{spec}} \sim 0.6$  and  $z_{\text{spec}} \sim 2.1$  by SDSS and reported in Paper XII.

In summary, the VTA resulted in the identification of seventeen overdense structures at  $z_{\text{phot}} = 0.43$ , seven of which are associated with X-ray detected clusters, while the remaining 10 structures are likely groups. The full structure extends over  $0^{\circ}6 \times 0^{\circ}2$  on the sky, which corresponds to a size of  $\sim 12 \times 4 \text{ Mpc}^2$  at  $z = 0.43$ . Two main galaxy cluster agglomerations are located toward the east (XLSSC 081, XLSSC 099, VTA04-10) and west (XLSSC 082-086, VTA01-03) of this structure. The identified overdense structures seem unvirialised and dynamically young, suggesting that these structures are perhaps in the process of forming a larger structure. To shed further light on the dynamical state of the overdensity structures, in the next section we investigate the radio properties of the western cluster/group agglomeration covered by our 3 GHz observations.

#### 4.3. Radio properties of the western cluster/group agglomeration

Previous studies used radio source counts and luminosity functions to investigate enhancement or suppression of radio emission. Several authors have found that in the pre-merging stages the suppression of AGN activity has not yet reached the cluster core (Venturi et al. 1997, 2001) where radio AGN preferably reside (Dressler 1980). A different situation is found in clusters that have already likely undergone the first core-to-core encounter and are now accreting smaller groups. In such cases, an excess in the number of blue galaxies is found at the expected positions of the shock fronts (Miller 2005; Johnston-Hollitt et al. 2008) as the merging process has had enough time to influence the galaxies' optical and radio properties. This implies that cluster mergers in later stages lower the probability of a galaxy developing an AGN (Venturi et al. 2000, 2001), but have also been shown to enhance the number of low-power radio galaxies driven by star formation (Miller 2005; Johnston-Hollitt et al. 2008). In this sense, as can be seen in Fig. 12, the clusters XLSSC 082-084 are close to each other, and we find that the core of XLSSC 083 seems to be preserved. Thus it appears that the agglomeration of XLSSC 082-084 has not yet undergone the first core-to-core encounter. The cluster XLSSC 086 also contains a radio source within a red host galaxy at its centre, and we note the increase in number of blue galaxies between XLSSC 086 and the large cluster structure south-west of it, possibly indicating a shock front present in this area. This picture seems to be consistent with the merger process discussed by Venturi et al. (2001) – most massive structures are the first to undergo a merger (in this case XLSSC 082-084), and they later attract satellite structures of lower mass (in this case XLSSC 086, and possibly VTA02-03).

To further investigate this, in Fig. 14 we show the Euclidean normalised radio source counts for the supercluster area, and the field with good rms excluding the supercluster area, as well as the general field source counts taken from Condon et al. (2012). Given the low number of radio sources in our field the supercluster counts are consistent with those in the field. Our results suggest the radio counts within the supercluster may be slightly enhanced (suppressed) at low (high) fluxes relative



**Fig. 14.** 1.4 GHz radio source counts for the supercluster area (red triangles), and the field with good rms excluding the supercluster area (black). Red dots are slightly offset from the bin centres for better visualisation. Errors are Poissonian, calculated using approximate algebraic expressions from [Gehrels \(1986\)](#). Radio source counts from [Condon et al. \(2012\)](#) are shown by the blue line. The flux densities at 1.4 GHz were calculated from the 3 GHz fluxes assuming a spectral index of  $-0.7$  for consistency with [Condon et al. \(2012\)](#).

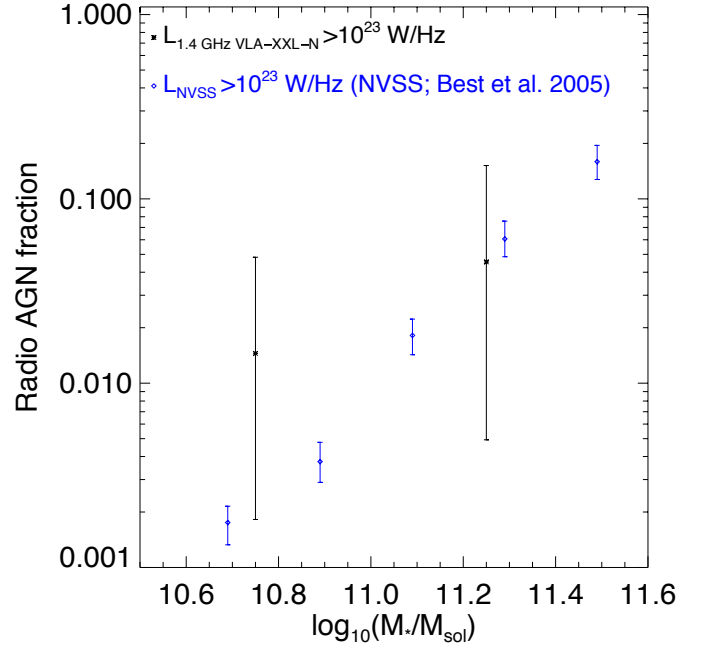
to the field. This is similar to the suppression of radio AGNs and an enhancement of low-power (likely star-forming) radio galaxies found in Shapely Supercluster ([Venturi et al. 2000, 2002](#); [Miller 2005](#)) and the Horologium-Reticulum Supercluster ([Johnston-Hollitt et al. 2008](#)) at  $z \sim 0.07$ . However, this is statistically not significant given the large error bars, and it prevents us from drawing further conclusions.

Our radio data probe  $L_{3 \text{ GHz}} \gtrsim 4 \times 10^{21} \text{ W/Hz}$ . This limit is very close to the division between star-forming and AGN galaxies, thus in the range where volume densities of star-forming and AGN galaxies are comparable. Converting this limit to star formation rate (SFR) using the [Bell \(2003\)](#) relation yields  $SFR \gtrsim 30 M_{\odot}/\text{year}$ . Thus, our data are not deep enough to probe a potential enhancement of star-forming galaxies within the supercluster for  $\lesssim 30 M_{\odot} \text{ yr}^{-1}$ . However, they are sensitive enough to detect powerful radio AGNs usually hosted by red galaxies (e.g. [Best et al. 2006](#); [Smolčić et al. 2008](#)), if present within the supercluster.

In [Fig. 15](#) we show the fraction of red galaxies detected in the radio (radio AGNs, as mentioned above) with  $L_{1.4 \text{ GHz}} > 10^{23} \text{ W/Hz}$  as a function of stellar mass for the region shown in the upper panel of [Fig. 11](#). The fraction was calculated by taking the ratio of the number of red ( $g - r \geq 1.17$ ) galaxies from the supercluster area detected in the radio (see [Tab. 4](#)) with  $L_{1.4 \text{ GHz}} > 10^{23} \text{ W/Hz}$ , and the number of optical red galaxies in the same field. The radio luminosity threshold was chosen to validate the comparison with the red galaxy fraction of NVSS-detected sources hosted by red galaxies from [Best et al. \(2005\)](#), and also shown in [Fig. 15](#).

The fractions derived here for the general field derived by [Best et al. \(2005\)](#) agree well for masses above  $\log(M_*/M_{\odot}) \sim 11$ , however, for  $\log(M_*/M_{\odot}) \lesssim 10.8$  the fraction of radio AGNs is slightly higher than expected, although still within the errors.

If the excess in the fraction of radio detected galaxies of lower stellar mass ( $\log(M_*/M_{\odot}) \sim 10.8$ ) were real, it could be explained by the physical mechanisms in the cluster which favour the AGN activity in red galaxies. [Best et al. \(2005\)](#) find that richer environments preferably host radio AGN. For



**Fig. 15.** The fraction of radio detected red galaxies (radio AGN) as a function of the stellar mass for the region shown in the upper panel of [Fig. 11](#). Black stars show the fractions estimated in this paper for galaxies with  $L_{1.4 \text{ GHz}} > 10^{23} \text{ W/Hz}$ . The errors shown on the black points are calculated from the equations given in [Gehrels \(1986\)](#). Blue diamonds show the fractions obtained for the NVSS sample by [Best et al. \(2005\)](#).

galaxies with stellar masses of  $\sim 10^{11} M_{\odot}$  they estimate a cooling rate of the order of magnitude of  $\sim 1 M_{\odot} \text{ yr}^{-1}$ , which is consistent with a low-power radio AGN. This indicates that the cooling of gas in the hot atmosphere of the host galaxy is a plausible candidate for the fuel source of radio AGN in the cluster structure studied here. This leads to the possibility that this cluster structure is rich in cooling gas, and/or that the cooling gas is likely cospatial with the lower-mass galaxies.

In summary, by analysing the properties of radio sources within the structure, we conclude that the clusters and group candidates in the western agglomeration are likely in the pre-merger state, before the first core-to-core encounter. An increase in the surface density of blue galaxies between the XLSSC 086 and the XLSSC 082-084 agglomeration could indicate a formation of a shock front. Since the XLSSC 082-084 agglomeration is the most massive one in the area, it is likely that in the future the three most massive clusters will be the first to go radio silent, followed by the smaller satellite structures XLSSC 086 and VTA01-03.

## 5. Summary

We have presented observations of a subfield of the XXL-N 25 deg<sup>2</sup> field with the VLA data at 3 GHz (10 cm). The final radio mosaic has a resolution of  $3''.2 \times 1''.9$  and encompasses an area of  $41' \times 41'$  with rms  $\lesssim 20 \mu\text{Jy beam}^{-1}$ . The rms in the central  $15' \times 15'$  area is  $10.8 \mu\text{Jy beam}^{-1}$ . From the mosaic we have extracted 155 sources with  $S/N > 6$ , 8 of which are large, multicomponent sources. Of the 155 sources, we find optical counterparts for 123 sources within the CFHTLS W1 catalogue.

We also analysed the first supercluster discovered in the XXL project, partially located within the radio mosaic. Applying the Voronoi tessellation analysis using photometric redshifts from the CFHTLS survey have identified seventeen overdense

structures at  $z_{\text{phot}} = 0.35\text{--}0.50$ , seven of which are detected in the *XMM-Newton* XXL data. The structure is extended over  $0^{\circ}6 \times 0^{\circ}2$  on the sky, which corresponds to a size of  $\sim 12 \times 4 \text{ Mpc}^2$  at  $z = 0.43$ . No large radio galaxies are present within the overdense structures, and we associate eight ( $S/N > 7$ ) radio sources with potential supercluster members. We find that the spatial distribution of the red and blue potential group member galaxies, selected by their observed  $g-r$  colours, suggests that the groups are not virialised, but are dynamically young, which is consistent with the hierarchical structure growth expected in a  $\Lambda$ CDM universe. Further spectroscopic follow-up is required for a more detailed analysis of the dynamical state of the structure.

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## References

- Arnouts, S., Cristiani, S., Moscardini, L., et al. 1999, *MNRAS*, **310**, 540  
 Baldry, I. K., Alpaslan, M., Bauer, A. E., et al. 2014, *MNRAS*, **441**, 2440  
 Becker, R. H., White, R. L., & Helfand, D. J. 1995, *ApJ*, **450**, 559  
 Bell, E. F. 2003, *ApJ*, **586**, 794  
 Best, P. N., Kauffmann, G., Heckman, T. M., et al. 2005, *MNRAS*, **362**, 25  
 Best, P. N., Kaiser, C. R., Heckman, T. M., & Kauffmann, G. 2006, *MNRAS*, **368**, L67  
 Bondi, M., Ciliegi, P., Schinnerer, E., et al. 2008, *ApJ*, **681**, 1129  
 Bourke, S., Mooley, K., & Hallinan, G. 2014, in *Astronomical Data Analysis Software and Systems XXIII*, eds. N. Manset, & P. Forshay, *ASP Conf. Ser.*, **485**, 367  
 Chabrier, G. 2003, *PASP*, **115**, 763  
 Condon, J. J., Cotton, W. D., Greisen, E. W., et al. 1998, *AJ*, **115**, 1693  
 Condon, J. J., Cotton, W. D., Fomalont, E. B., et al. 2012, *ApJ*, **758**, 23  
 Coupon, J., Ilbert, O., Kilbinger, M., et al. 2009, *A&A*, **500**, 981  
 Dehghan, S., & Johnston-Hollitt, M. 2014, *AJ*, **147**, 52  
 Dickinson, M., Giavalisco, M., & GOODS Team 2003, in *The Mass of Galaxies at Low and High Redshift*, eds. R. Bender, & A. Renzini, 324  
 Dressler, A. 1980, *ApJ*, **236**, 351  
 Driver, S. P., Norberg, P., Baldry, I. K., et al. 2009, *Astron. Geophys.*, **50**, 12  
 Driver, S. P., Hill, D. T., Kelvin, L. S., et al. 2011, *MNRAS*, **413**, 971  
 Durret, F., Adami, C., Cappi, A., et al. 2011, *A&A*, **535**, A65  
 Erben, T., Hildebrandt, H., Miller, L., et al. 2013, *MNRAS*, **433**, 2545  
 Gehrels, N. 1986, *ApJ*, **303**, 336  
 Giles, P. A., Maughan, B. J., Pacaud, F., et al. 2016, *A&A*, **592**, A3 (XXL Survey, III)  
 Grogin, N. A., Kocevski, D. D., Faber, S. M., et al. 2011, *ApJS*, **197**, 35  
 Helfand, D. J., White, R. L., & Becker, R. H. 2015, *ApJ*, **801**, 26  
 Heymans, C., Van Waerbeke, L., Miller, L., et al. 2012, *MNRAS*, **427**, 146  
 Hildebrandt, H., Erben, T., Kuijken, K., et al. 2012, *MNRAS*, **421**, 2355  
 Hinshaw, G., Larson, D., Komatsu, E., et al. 2013, *ApJS*, **208**, 19  
 Ilbert, O., Arnouts, S., McCracken, H. J., et al. 2006, *A&A*, **457**, 841  
 Jelić, V., Smolčić, V., Finoguenov, A., et al. 2012, *MNRAS*, **423**, 2753  
 Johnston-Hollitt, M., Sato, M., Gill, J. A., Fleenor, M. C., & Brick, A.-M. 2008, *MNRAS*, **390**, 289  
 Kimball, A. E., & Ivezić, Ž. 2008, *AJ*, **136**, 684  
 Koekemoer, A. M., Faber, S. M., Ferguson, H. C., et al. 2011, *ApJS*, **197**, 36  
 Koulouridis, E., Poggianti, B., Altieri, B., et al. 2016, *A&A*, **592**, A11 (XXL Survey, XII)  
 Lieu, M., Smith, G. P., Giles, P. A., et al. 2016, *A&A*, **592**, A4 (XXL Survey, IV)  
 Miller, N. A. 2005, *AJ*, **130**, 2541  
 Mooley, K. P., Hallinan, G., Frail, D. A., et al. 2015, in *Am. Astron. Soc. Meet. Abstr.*, **225**, 113.05  
 Novak, M., Smolčić, V., Civano, F., et al. 2015, *MNRAS*, **447**, 1282  
 Oklopčić, A., Smolčić, V., Giodini, S., et al. 2011, *Mem. Soc. Astron. Ital.*, **82**, 161  
 Pacaud, F., Clerc, N., Giles, P. A., et al. 2016, *A&A*, **592**, A2 (XXL Survey, II)  
 Pierre, M., Pacaud, F., Adami, C., et al. 2016, *A&A*, **592**, A1 (XXL Survey, I)  
 Pompei, E., Adami, C., Eckert, D., et al. 2016, *A&A*, **592**, A6 (XXL Survey, VII)  
 Ramella, M., Boschin, W., Fadda, D., & Nonino, M. 2001, *A&A*, **368**, 776  
 Rau, U., & Cornwell, T. J. 2011, *A&A*, **532**, A71  
 Scoville, N., Aussel, H., Brusa, M., et al. 2007, *ApJS*, **172**, 1  
 Smolčić, V., Schinnerer, E., Finoguenov, A., et al. 2007, *ApJS*, **172**, 295  
 Smolčić, V., Schinnerer, E., Scodreggio, M., et al. 2008, *ApJS*, **177**, 14  
 Smolčić, V., Zamorani, G., Schinnerer, E., et al. 2010, *ApJ*, **708**, 909  
 Strateva, I., Ivezić, Ž., Knapp, G. R., et al. 2001, *AJ*, **122**, 1861  
 Vajgel, B., Jones, C., Lopes, P. A. A., et al. 2014, *ApJ*, **794**, 88  
 van de Weygaert, R., & Icke, V. 1989, *A&A*, **213**, 1  
 Velander, M. B. M. 2012, Ph.D. Thesis, Leiden Observatory, Leiden University, The Netherlands  
 Venturi, T., Bardelli, S., Morganti, R., & Hunstead, R. W. 1997, *MNRAS*, **285**, 898  
 Venturi, T., Bardelli, S., Morganti, R., & Hunstead, R. W. 2000, *MNRAS*, **314**, 594  
 Venturi, T., Bardelli, S., Zambelli, G., Morganti, R., & Hunstead, R. W. 2001, *MNRAS*, **324**, 1131  
 Venturi, T., Bardelli, S., Zagaria, M., Prandoni, I., & Morganti, R. 2002, *A&A*, **385**, 39