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# Beryllium detection in the very fast nova ASASSN-16kt (V407 Lupi)

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## ABSTRACT

We present high-resolution spectroscopic observations of the fast nova ASASSN-16kt (V407 Lup). A close inspection of spectra obtained at early stages has revealed the presence of low-ionization lines, and among the others we have identified the presence of the ionized <sup>7</sup>Be doublet in a region relatively free from possible contaminants. After studying their intensities, we have inferred that ASASSN-16kt has produced  $(5.9\text{--}7.7) \times 10^{-9} M_{\odot}$  of <sup>7</sup>Be. The identification of bright Ne lines may suggest that the nova progenitor is a massive (1.2  $M_{\odot}$ ) oxygen–neon white dwarf. The high outburst frequency of oxygen–neon novae implies that they likely produce an amount of Be similar, if not larger, to that produced by carbon–oxygen novae, then confirming that classical novae are among the main factories of lithium in the Galaxy.

**Key words:** nuclear reactions, nucleosynthesis, abundances – stars: individual: ASASSN-16kt: V407 Lup – novae, cataclysmic variables – Galaxy: evolution.

## 1 INTRODUCTION

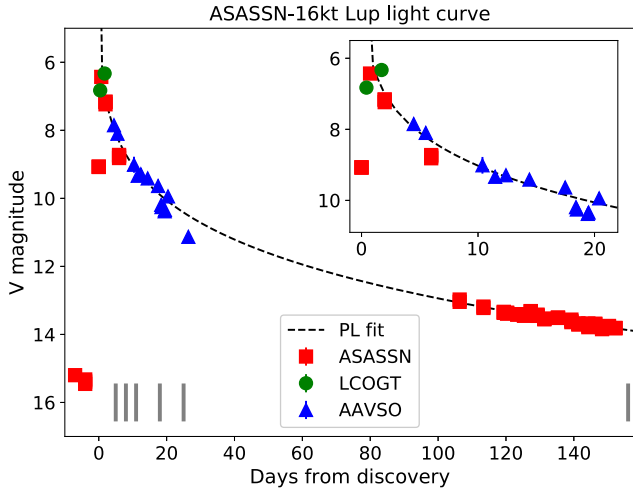
During a nova outburst, the thermo-nuclear runaway (TNR) process produces several beta isotopes of neon, oxygen, and fluorine whose decays release enough energy to eject the accreted layers (Wagoner, Fowler & Hoyle 1967; Fowler, Caughlan & Zimmerman 1975; Hernanz et al. 1996; Starrfield et al. 1996; José & Hernanz 1998; Starrfield et al. 1998; Yaron et al. 2005; Starrfield et al. 2009). At the same time, they furnish the nova ejecta with non-solar CNO isotopes (mainly <sup>13</sup>C, <sup>15</sup>N, and <sup>17</sup>O) that will propagate into the interstellar (IS) medium then contributing to the Galactic chemical enrichment (Gehrz et al. 1998; José et al. 2004). It was proposed by Cameron & Fowler (1971) that during advanced phases of red giant stars the reaction of <sup>3</sup>He with <sup>4</sup>He can give rise to the <sup>7</sup>Be isotope, which decays only through electron capture into Li after a half-time decay

of  $\sim 53$  d (Giraud et al. 2007). This mechanism was later applied to novae by Arnould & Norgaard (1975). Consequently, classical novae have been proposed as one of the possible lithium factories in the Galaxy (Starrfield et al. 1978; D’Antona & Matteucci 1991; Romano et al. 1999).

In recent years, thanks to the use of high-resolution spectrographs, we were for the first time able to observe the presence of <sup>7</sup>Li in the early spectra of novae V1369 Cen (Izzo et al. 2015) and V5668 Sgr (Wagner et al., in preparation, Izzo et al., in preparation). In addition, evidence of large quantities of <sup>7</sup>Be II has also been reported, after  $\sim 60$  d from the outburst, for the carbon–oxygen (CO) novae V339 Del (Tajitsu et al. 2015), V2944 Oph (Tajitsu et al. 2016), and V5668 Sgr Molaro et al. (2016). In the latter case, the observation of unsaturated narrow components allowed an accurate <sup>7</sup>Be abundance determination, which led the authors to suggest that classical novae are one of the main lithium factories in the Galaxy.

In this paper, we present a new detection of <sup>7</sup>Be II in the spectra of ASASSN-16kt (also named V407 Lup), an oxygen–neon (ONe) nova, then providing additional support that ONe novae also

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**Figure 1.**  $V$  light curve of ASASSN-16kt obtained with data from ASASSN, LCOGT, and AAVSO observations. The dashed black line represents the best fit of the data (excluding the discovery observation) with a power-law function, used to estimate the  $t_2$  value. The grey vertical lines correspond to the spectroscopic observation epochs. The gap of data between days 40 and 100 is mainly due to the difficult visibility of the nova when it was too close to the Sun. The inset figure shows a zoom around the first 20 d of the nova outburst.

produce  ${}^7\text{Be II}$ . After presenting our data set (Section 2), we discuss some physical properties of the nova system such as its distance and the extinction, as inferred from the spectral data set (Section 3). Then, we present an analysis of the early phases of the nova, showing the presence of  ${}^7\text{Be II}$  and the absence of  ${}^7\text{Li I}$  in the early spectra (Section 4). Finally, from the unique late nebular spectrum in our data set, we provide an estimate of the mass of the nova ejecta and hence the amount of  ${}^7\text{Be II}$  produced in the nova outburst (Sections 5 and 6), before drawing our conclusions in Section 7.

## 2 OBSERVATIONS

ASASSN-16kt was discovered by the All-Sky Automated Survey for SuperNovae (ASAS-SN; Shappee et al. 2014) as a bright source of  $V = 9.1$  mag on 2016 September 24 (Stanek et al. 2016). The nova brightened in the following two days, reaching a maximum of  $V = 6.3$  on 2016 September 26 (Chen et al. 2016), after which it started to decay rapidly. In Fig. 1, we show the  $V$  mag light curve of ASASSN-16kt obtained with data from the ASAS-SN (Kochanek et al. 2017), LCOGT (Brown et al. 2013) and AAVSO<sup>1</sup> observations. If we exclude the first data points, corresponding to the pre-discovery and when the nova was still rising in luminosity, the light curve (magnitude versus time) is well fitted with a simple power-law function with power-law index  $\gamma = -0.16 \pm 0.01$ . From this, we derive a value for the time the light curve decays of two and three magnitudes from the peak of  $t_2 = 5.9 \pm 0.1$  d and  $t_3 = 13.2 \pm 0.1$  d.

After the discovery, we started a DDT program at ESO-VLT with UVES and X-Shooter [Program ID 297.D-5065(B), PI L. Izzo]. In parallel, we observed the nova with the PUC High Echelle Resolution Optical Spectrograph (PUCHEROS; Vanzi et al. 2012) mounted on the ESO 0.5 m telescope located at the Observatory of Pontificia Universidad Catolica (OUC) in Santiago, Chile. We obtained a total of six epochs, five of these concentrated in the first 25 d

**Table 1.** Log of observations. The values on the first column refer to days from the discovery.

Epoch	Date	Instrument	Range (Å)
Day 5	Sep. 29 2016	PUCHEROS	4250–6900
Day 8	Oct. 2 2016	UVES	3060–9460
Day 11	Oct. 5 2016	PUCHEROS	4250–6900
Day 18	Oct. 12 2016	UVES	3060–9460
Day 25	Oct. 19 2016	X-Shooter	3140–24800
Day 155	Feb. 27 2017	X-Shooter	3140–24800

and the last one after 155 d from its discovery, see also Fig. 1. The detailed log of the observations is shown in Table 1. PUCHEROS data have been reduced with the CERES routine for echelle spectra (Brahm, Jordán & Espinoza 2017). The ESO-VLT UVES data have been reduced and flux-calibrated using the *reflex* environment (Freudling et al. 2013), while for the X-shooter data we used the standard ESO Recipe Execution Tool (*esorex*)<sup>2</sup>. We made extensive use of various python scripts for the analysis and IRAF packages for counter-checking all measurements. In particular, we recognize the use of the NUMPY (Van Der Walt, Colbert & Varoquaux 2011), MATPLOTLIB (Hunter 2007), and the ASTROPY (Astropy Collaboration et al. 2013) packages.

## 3 GENERAL PROPERTIES

According to the classification introduced by Gaposchkin (1957), ASASSN-16kt can be considered a very fast nova ( $t_2 \leq 10$  d). CNe undergo a super-Eddington phase (see della Valle & Livio 1995, and references therein) during their maximum luminosity. This is a parameter which depends on the temperature of the progenitor white dwarf (WD),  $T_{\text{WD}}$ , on the accretion rate,  $\dot{M}$  (Yaron et al. 2005), but mainly on the mass of the WD progenitor (Livio 1992): bright and fast novae are indeed generally related to more massive progenitors (Iben & Tutukov 1985). Moreover, a massive WD needs less accreted matter to initiate the TNR, and consequently the mass ejected in the outburst is less that expelled by less massive nova progenitors. An immediate result of this theoretical deduction is the existence of some correlations between the absolute magnitude at nova maximum and the rate of decline (the maximum-magnitude rate-of-decline, MMRD, see also Downes & Duerbeck 2000). However, we are well aware that the accuracy provided by the MMRD methods used for novae in M31 and LMC is of the order of 30 per cent (van den Bergh 1992; della Valle & Livio 1995), and that not all novae follow these correlations, like the new class of very fast and faint novae with a  $t_2 \sim$  of few hours for which an extension of the MMRD was pointed out (Kasliwal et al. 2011). In what follows, we take into account this additional source of uncertainty.

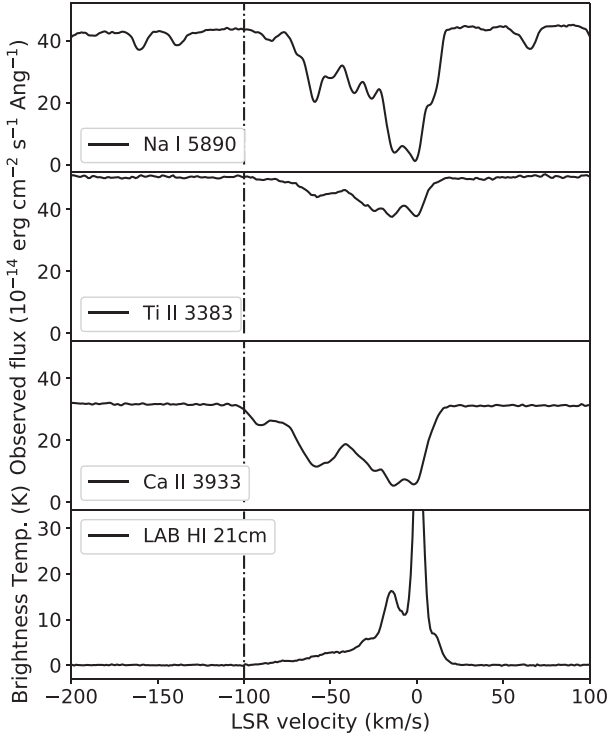
Using the MMRD formulation given in della Valle & Livio (1995):

$$V_{pk} = -7.92 - 0.81 \times \arctan\left(\frac{1.32 - \log t_2}{0.23}\right), \quad (1)$$

with  $V_{pk}$  the absolute  $V$  magnitude at nova maximum and  $t_2$  the time the nova decay of 2 mag from the maximum, we find the distance of ASASSN-16kt to be  $d = 10.9 \pm 1.6$  kpc. Similar results are obtained using the Buscombe & de Vaucouleurs (1955) method, in the formulation given by Downes & Duerbeck (2000), where all novae at 15 d after the maximum have a similar absolute magnitude,

<sup>1</sup> <http://www.aavso.org> (see Kafka 2017)

<sup>2</sup> <http://www.eso.org/sci/software/cpl/esorex.html>



**Figure 2.** The structured line profiles of the IS lines Na I 5890, Ti II 3383, and Ca II 3933 compared with the profile of the H I 21 cm line from the LAB survey (Kalberla et al. 2005), all corrected for the local standard of rest velocity. The black dashed line marks the  $v = -100 \text{ km s}^{-1}$  limit, corresponding to the beginning of the rising line profile, and implying that the nova is located at the border of the Galactic disc.

given by  $M_{V,15} = -6.05 \pm 0.44$ . For ASASSN-16kt, we have that the observed V-band magnitude at 15 d is  $V_{15} = 9.6$  mag, then we obtain a distance of  $d_{V,15} = 11.2 \pm 2.0$  kpc, in agreement with our previous estimate.

At Galactic coordinates of  $l = 330:0938$  and  $b = 9:573$ , we obtain a height above the Galactic plane of  $h = 1800 \pm 300$  pc, which is typical for *bulge* novae (della Valle & Livio 1998). We also report that the position of the nova is 0.5 arcsec distant from a *GAIA* DR1 source (ID 5999691733347769472) with a still unknown parallax and a magnitude of  $G = 18.95$ . Assuming this source is the possible progenitor of ASASSN-16kt, we estimate an amplitude of the outburst of  $A \sim 12.3 \pm 0.2$ . We finally note that, given the above estimate of the distance, the lack of a measured parallax in the DR1 is consistent with our result for the distance of the nova.

At maximum light, novae show an intrinsic colour index of  $(B - V)_{\text{max}} = 0.23 \pm 0.06$ , (van den Bergh & Younger 1987). This is explained by the reprocessing of the incoming radiation from the underlying WD at longer wavelengths through the nova ejecta (Shore 2012). Consequently, we can estimate the intrinsic extinction from the observed colour index: at maximum, ASASSN-16kt has  $V = 6.33 \pm 0.06$  and  $B = 6.86 \pm 0.06$  (Chen et al. 2016), and hence  $(B - V)_{\text{obs}} = 0.53 \pm 0.08$ . Combining this observed data with the results of van den Bergh & Younger (1987), we get an intrinsic extinction of  $E(B - V) = 0.30 \pm 0.09$ .

The high resolution provided by UVES allowed us to identify several IS lines and diffuse interstellar bands (DIBs) in the first UVES spectrum (Day 8). Several IS lines such as Na I D, Ti II, and Ca II H,K, have similar profiles with multiple absorption features at bluer wavelengths, see Fig. 2. The common maximum blue-shifted

**Table 2.** List of observed IS lines and DIBs in the Day 8 UVES spectrum. We give the IS/DIB name (column 1), the observed central wavelength (2), the observed EW (3), and the rest-frame DIB wavelength (4). For the Na I 3302/3 doublet, we report the total EW of both components.

DIB	$\lambda_{\text{obs}}$ (Å)	EW (mÅ)	$\lambda_{\text{rf}}$ (Å)
Ti II 3242	3242.22	$101.0 \pm 5.8$	3241.994
Na I 3302/3	3302.60/3.21	$61.2 \pm 3.8$	3302.369/978
Ti II 3383	3384.00	$152.1 \pm 6.2$	3383.768
Ca II 3933	3933.91	$718.5 \pm 43.7$	3933.663
Ca II 3968	3968.72	$487.1 \pm 14.5$	3968.469
Na I 5890	5890.33	$881.6 \pm 16.9$	5889.950
Na I 5896	5896.31	$695.2 \pm 10.9$	5895.924
CH+ 3958	3957.97	$47.3 \pm 1.0$	3957.7
CH+ 4232	4232.85	$13.7 \pm 0.4$	4232.6
CH 4300	4300.65	$7.8 \pm 0.2$	4300.3
DIB 5780	5780.69	$105.5 \pm 0.9$	5780.5
DIB 5797	5797.36	$50.5 \pm 0.8$	5797.1
DIB 6196	6196.39	$25.1 \pm 0.4$	6196.0
DIB 6614	6613.97	$50.8 \pm 0.3$	6613.6

velocity reported for DIBs is  $v \sim -100 \text{ km s}^{-1}$ . These observations point to the presence of a number of IS clouds along the line of sight to the nova, and the observed maximum blue-shifted components suggests them to be at a large distance from us. In Fig. 2, we also show the neutral hydrogen 21 cm line profile corrected to heliocentric velocities, as observed by the LAB survey (Kalberla et al. 2005), at the position of the nova. We note an increasing brightness temperature in the H I 21 cm profile starting exactly at  $v_{\text{hel}} \approx -100 \text{ km s}^{-1}$ , suggesting that, considering its Galactic coordinates, the nova is located at the opposite border of the Galactic disc, with respect to the Sun location.

Some DIBs can also be used to estimate line-of-sight reddening. To this aim, we used the correlations between the equivalent widths (EW) of these DIBs, the column density of neutral atomic hydrogen  $N(H)$  and the colour excess provided by Friedman et al. (2011). The total list of IS lines and DIBs with secure identification and measurements of their observed central wavelengths as well as of their EWs is reported in Table 2. We obtained an estimate for  $N(H)$  and  $E(B - V)$  using DIBs at  $\lambda\lambda$  5780, 5797, 6196, and 6614. After a first check, we excluded DIB  $\lambda$  6196 since the estimated  $N(H)$  value is larger than that derived from an analysis of the neutral hydrogen line at 21 cm from the LAB survey, i.e.  $N(H)_{\text{HI}} = 1.54 \times 10^{21} \text{ cm}^{-2}$ : we believe indeed that this DIB is blended with some other unidentified component. The final values obtained from the average of the remaining DIB estimates are:  $N(H) = 1.03 \times 10^{21} \text{ cm}^{-2}$  and  $E(B - V) = 0.24 \pm 0.02$ , which is in good agreement with our above estimate obtained from the observed colour index at maximum. We then used these values to correct the spectra for Galactic extinction.

A summary of the main properties of ASASSN-16kt is shown in Table 3.

#### 4 THE EARLY SPECTRA

The first spectroscopic observation of the nova was obtained with PUCHEROS ( $R \sim 20000$ ) five days after the nova discovery, see Fig. 3. The spectrum shows clear Balmer series P-Cygni profiles with broad blue-shifted absorptions centred at  $v = -2030, -3650 \text{ km s}^{-1}$ . The first of these absorption systems coincides in velocity with the observed Na I D P-Cygni profile line. The presence

**Table 3.** Summary of the main physical properties of ASASSN-16kt: (1) the  $V$  magnitude at maximum brightness  $V_{\max}$ ; (2) the amplitude  $A$  in magnitudes from the possible progenitor star; (3–4) the  $t_2$  and  $t_3$  values; (5) the colour index  $E(B - V)$ ; (6) the maximum ejecta velocity  $v_{\text{ej}}$ ; (7) the distance  $d$ , and (8) the height  $h$  above the Galactic plane.

Property	Value
$V_{\max}$	6.3 mag
$A$	$12.3 \pm 0.2$ mag
$t_2$	$5.9 \pm 0.1$ d
$t_3$	$13.2 \pm 0.1$ d
$E(B - V)$	$0.24 \pm 0.02$ mag
$v_{\text{ej}}$	$2370 \text{ km s}^{-1}$
$d$	$10.9 \pm 1.6$ kpc
$h$	$1.8 \pm 0.3$ kpc

of  $\text{N III } \lambda 4640$  and  $\text{He I } \lambda 5016$  in emission suggests a He/N spectral class for the nova (Williams et al. 1991). However, the presence of an emission line coincident with  $\text{Fe II } \lambda 5169$  tells us that the nova could have passed through a very fast iron-curtain phase.

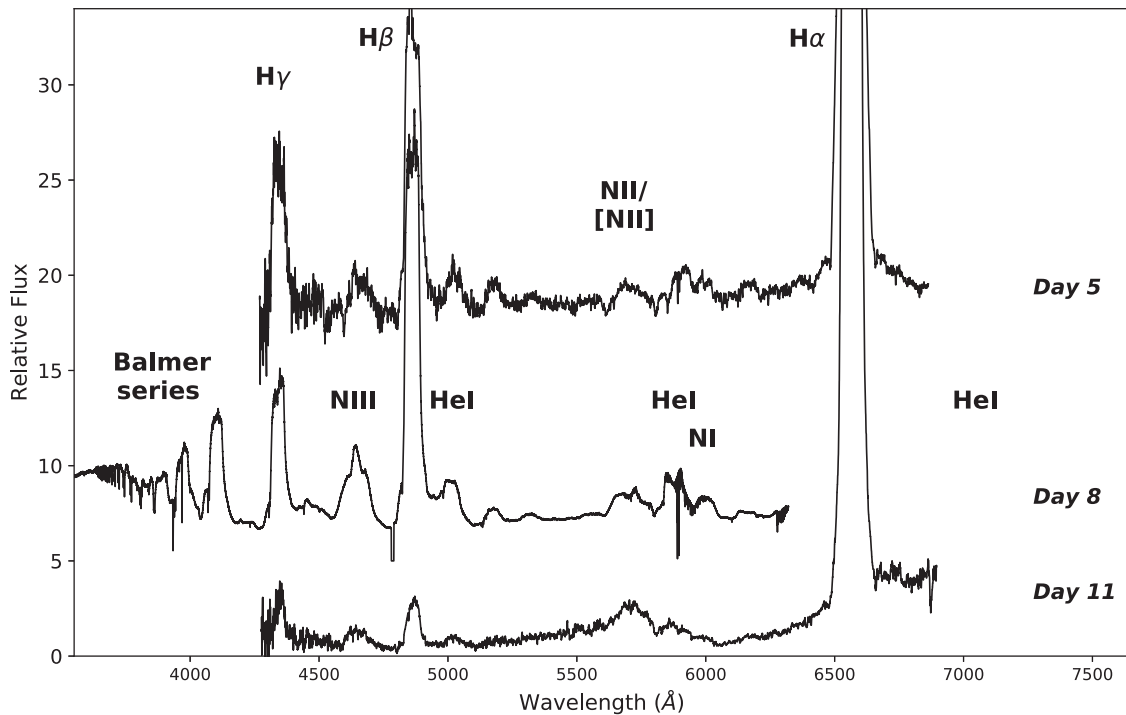
The following spectrum, obtained three days later with UVES at VLT, confirms this hypothesis. We clearly distinguish emission lines of  $\text{He I } \lambda\lambda 5016, 5876$ ,  $\text{N II } \lambda 5678$  and the blend  $\text{N III } \lambda 4640$  as the strongest non-Balmer lines, if we exclude the  $\text{O I } \lambda 8446$ , see Fig. 3. We also identify transitions of  $\text{N I}$  quartet lines  $\lambda\lambda 7454, 8212, 8692$ , and  $\text{N I}$  doublet lines  $\lambda\lambda 7904, 9028/60, 9395$  (partly embedded with telluric lines), as well as  $\text{N I } \lambda 6005$ . All these neutral nitrogen lines are characterized by the absence of P-Cygni profiles and a bell-

shaped emission-line profile, unlike the  $\text{He I}$  lines which show a clear saddle-shaped profile lines, with evidence of P-Cygni absorption at  $v_{\text{ej}} = -2030$  and  $-3830 \text{ km s}^{-1}$ . We also note the presence of forbidden lines like  $[\text{N II}] \lambda 5755$ , which is partially blended with the  $\text{N II } 5678$  line, and faint  $[\text{O II}] \lambda\lambda 7320/30$  doublet, which is embedded in telluric absorption lines. These features are typical of the He/N spectral class.

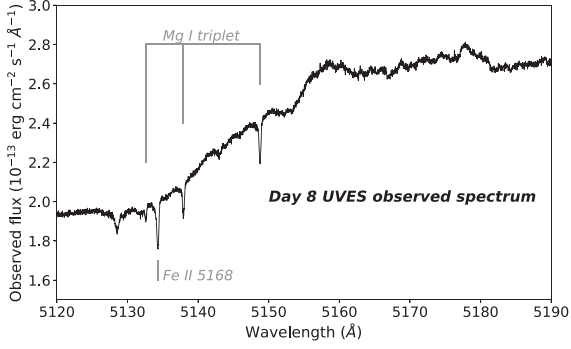
The high resolution, high signal-to-noise data provided by UVES also allow us to identify low-ionization elements in the nova ejecta. The  $\text{Na I D}$  lines serve as reference for the identification of low-ionization transitions (see e.g. the method developed by Williams et al. 2008): we determine an expansion velocity from analysis of  $\text{Na I D}$  lines of  $v_{\text{ej},1} = -2030 \text{ km s}^{-1}$ , similar to what we have observed for  $\text{He I } \lambda 5876$ , and an additional fainter component at  $v_{\text{ej},2} = -2370 \text{ km s}^{-1}$ . Assuming the first value as the ejecta velocity of the low-ionization elements, we are able to identify two lines of  $\text{Fe II}$  (multiplet 42)  $\lambda\lambda 5018, 5169$  (the P-Cygni of the other line at  $4924 \text{ \AA}$  is not clearly identified, since the corresponding absorption is located within the red wing of the bright  $\text{H } \beta$  line) and the  $\text{Mg I } \lambda 5178$  triplet, see Fig. 4. Furthermore, we clearly identify  $\text{Fe II } \lambda 5018$  and other similar transitions at  $\lambda\lambda 3154, 3168, 3210, 3213, 3228$ , the triplet of  $\text{Mg I } \lambda\lambda 3829, 3832, 3838$ ,  $\text{Ca II } \lambda\lambda 3933, 3968$ , and several lines of  $\text{Cr II } \lambda\lambda 3132, 3135, 3137, 3147$ . Concluding, ASASSN-16kt shows traces of low-ionization elements that are typical of the  $\text{Fe II}$  nova spectral class (according to the Cerro-Tololo classification) and confirms the end of a very rapid iron-curtain phase.

#### 4.1 Detection of the $^7\text{Be II } \lambda 3130$ doublet

In addition to the above mentioned low-ionization transitions, we identified  $^7\text{Be II } \lambda\lambda 3130, 3131$ , which is blue-shifted at the same



**Figure 3.** The optical spectra of ASASSN-16kt observed on Day 5 and Day 11 with PUCHEROS and on Day 8 with UVES. We show only a limited wavelength range (3550–7300 Å) of the UVES spectrum for clarity. He/N features are reported and visible in both spectra where coverage with PUCHEROS is available. A faint emission line attributed to  $\text{Fe II } \lambda 5168$  is clearly visible in the first two (Day 5 and Day 8) spectra.



**Figure 4.** Day 8 observed spectrum in the range 5120–5190 Å. The Fe II  $\lambda 5168$  line and the Mg I triplet coincide with their laboratory wavelengths, after correcting for the blue-shifted ejecta velocity as measured from Na I D lines. The additional absorption observed at  $\lambda 5128.5$  corresponds to the blue-shifted component at  $v_{\text{exp},2} = -2370 \text{ km s}^{-1}$  of Fe II  $\lambda 5168$ , see also Fig. 5.

velocity as the Na I D line. The region is relatively free from absorptions of low ionization species and the two lines of the <sup>7</sup>Be II are clearly visible. In addition, the P-Cygni profile extends up to velocities  $\sim -2370 \text{ km s}^{-1}$ , where additional blue-shifted velocities are also observed for Na I D, Fe II  $\lambda 5168$ , and Ca II  $\lambda 3933$ , see Fig. 5. Following the same approach used for V1369 Cen (Izzo et al. 2015) and V5668 Sgr (Molaro et al. 2016), we can estimate the mass abundance of <sup>7</sup>Be II ejected in ASASSN-16kt through comparison with the resonance lines of Ca II. First, we note the absence of the Ca I  $\lambda 4226$  resonance line in the Day 8 spectrum, hence we conclude that the Ca in the nova ejecta is largely ionized. Since the resonance

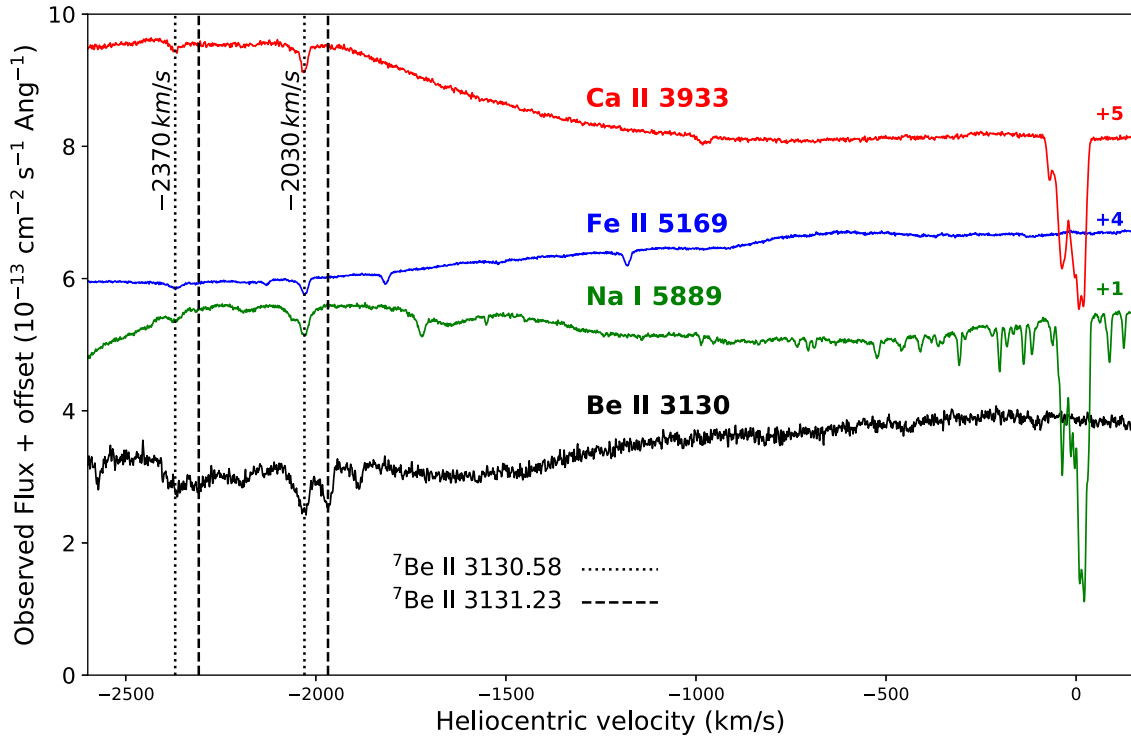
transitions differ by 0.80 eV, we can calculate the mass abundance as follows (Spitzer 1998):

$$\frac{A_m(\text{Be})}{A_m(\text{Ca})} = \left( \frac{\text{EW}_{\text{Be II}}}{3130^2} / \frac{\text{EW}_{\text{Ca II}}}{3933^2} \right) \times \frac{gf_{\text{Ca II}}}{gf_{\text{Be II}}} \times \frac{u_{\text{Be}}}{u_{\text{Ca}}} \quad (2)$$

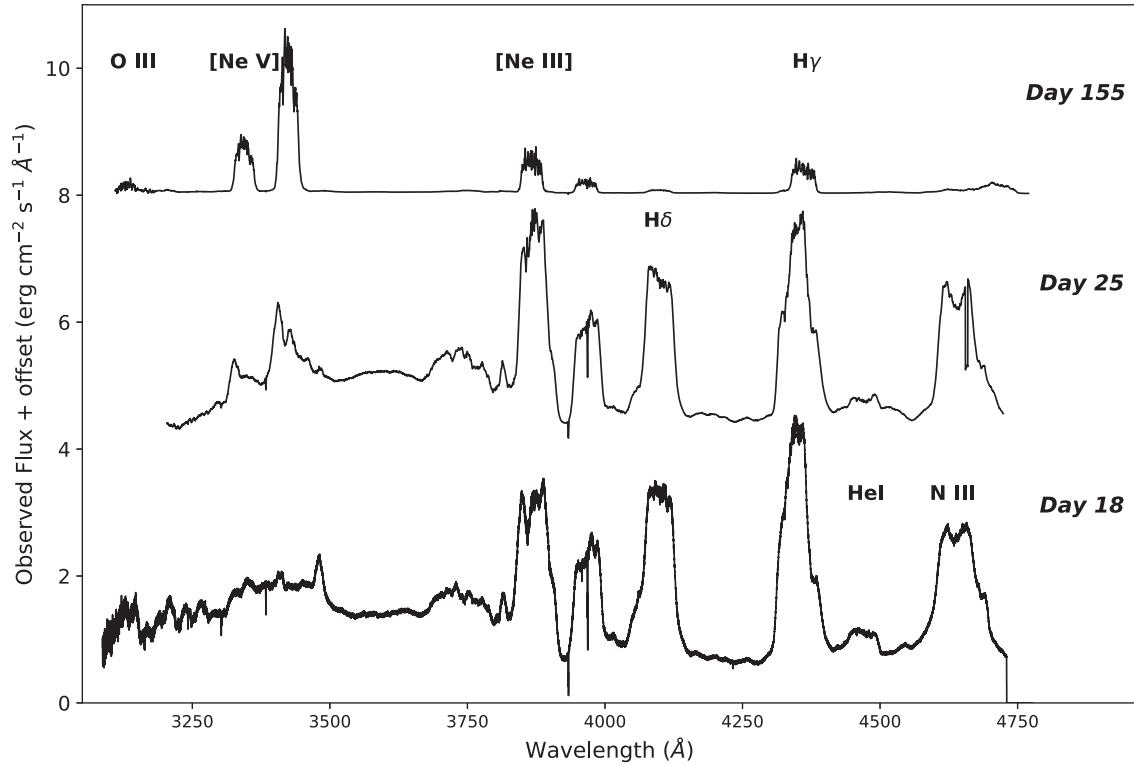
where  $u$  is the atomic mass (40 for Ca and 7 for <sup>7</sup>Be II). The oscillator strengths  $gf$ , are  $gf_{\text{Be}} = 0.68 + 0.34 = 1.02$  (Bièmont 1977) and  $gf_{\text{Ca}} = 1.36$  (Black, Weisheit & Laviana 1972). We then measured the EWs of the blue-shifted absorption lines associated with both elements. The total contribution due to Ca II is provided by  $\text{EW}_{\text{Ca II}} = \text{EW}_{\text{Ca II} \lambda 3933}$ . From the observed spectra, and considering the blue-shifted absorptions at  $v_{\text{ej}} = -2030, -2370 \text{ km s}^{-1}$  for both transitions, we obtain total EWs of  $\text{EW}_{\text{Be II}} = 434.3 \text{ m} \text{ \AA}$ , and  $\text{EW}_{\text{Ca II}} = 41.4 \text{ m} \text{ \AA}$ . Considering the formula in equation (2), we obtain a mass abundance ratio of  $A_m(\text{Be})/A_m(\text{Ca}) = 3.86$ , and considering that after 8 d from the initial outburst  $\sim 10$  percent of <sup>7</sup>Be II has already decayed into <sup>7</sup>Li I, we get an initial <sup>7</sup>Be II mass abundance produced during the TNR of  $A_{m,\text{in}}(\text{Be})/A_m(\text{Ca}) = 4.25$ .

### 5 PHYSICAL PROPERTIES FROM LATE EPOCHS

All low-ionization absorption features were absent in the UVES spectrum obtained 18 d after the nova’s discovery. This spectrum shows the clear presence of He I  $\lambda\lambda 4471, 5876, 7065, 10830$  lines as well as emerging nebular transitions such as [O III]  $\lambda\lambda 4959, 5007$  (the auroral [O III]  $\lambda 4363$  line is blended with H  $\gamma$ ), and high-ionization iron lines of [Fe VI]  $\lambda\lambda 5176, 5677$ . We also note a faint Bowen fluorescence line of O III  $\lambda 3132$  in the spectrum from Day 25, while its detection in the Day 155 spectrum is doubtful. Day 25



**Figure 5.** ASASSN-16kt spectrum at Day 8. The figure displays the spectrum around Be II  $\lambda 3130$  (black line), Na I  $\lambda 5890$  (green line), Fe II  $\lambda 5169$  (blue line) and Ca II  $\lambda 3933$  (red line) plotted on the velocity scale. The fluxes have been scaled to show the common absorption component at  $v_{\text{exp},1} = -2030 \text{ km s}^{-1}$ , which is shown as vertical dashed (for the Be II  $\lambda 3130$  doublet component) and dotted (for the Be II  $\lambda 3131$ ) lines, and at  $v_{\text{exp},2} = -2370 \text{ km s}^{-1}$ , shown as dash-dot lines.

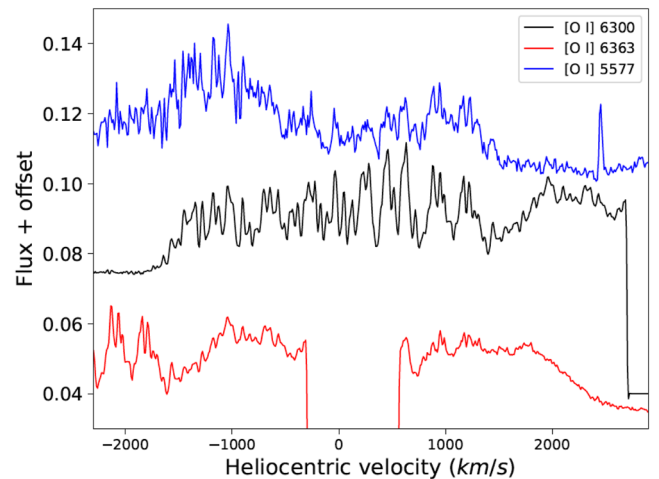


**Figure 6.** The optical spectra of ASASSN-16kt observed on Day 18 with UVES and on Day 25 and Day 155 with X-Shooter. The flux evolution in time of neon lines as well as of Bowen fluorescence lines as O III  $\lambda$ 3132 and the blend N III 4640 is clearly visible. Neon lines are blended with O III and [Ne III] lines in the Day 18 and Day 25 spectra.

spectrum shows also typical forbidden lines of Ne such as [Ne III]  $\lambda\lambda$ 3869,3967 and [Ne V]  $\lambda\lambda$ 3346,3426, the latter ones being fainter than [Ne III] and blended with O III and [Ne III] lines. The situation does not change on the Day 25 spectrum, while the spectrum obtained on Day 155 shows a further increase of excitation, marked by the flux increase of [Ne V] lines, making it even brighter than the H  $\alpha$  line, see Fig. 6, as well as the presence of [Fe VII]  $\lambda$ 6087.

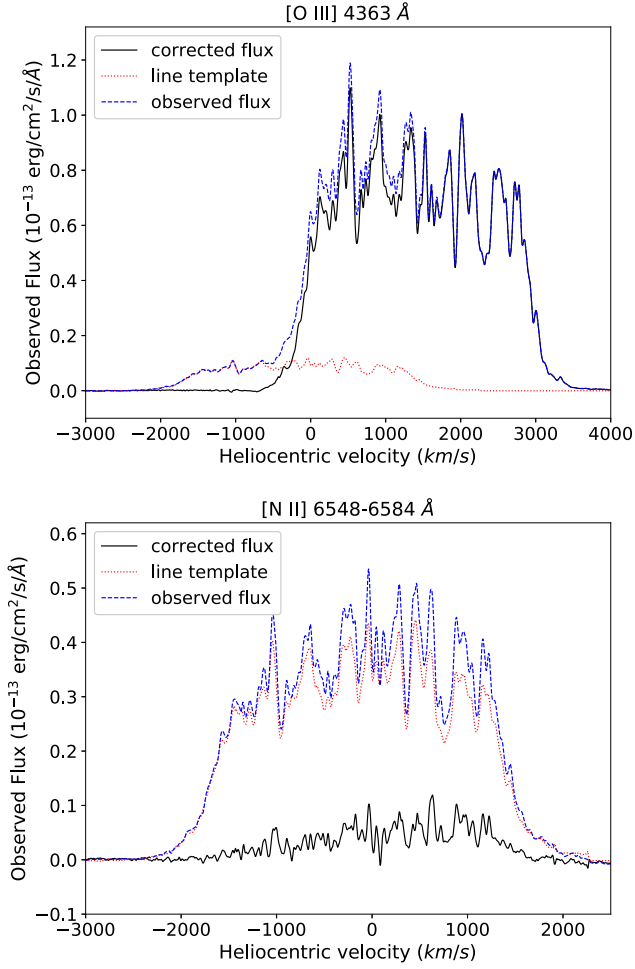
We also detect highly-ionized [Fe X]  $\lambda$ 6375 blended with [O I]  $\lambda$ 6363. The intensity of this line is usually 1/3 of the emission of the nearby [O I]  $\lambda$ 6300, but both lines show a similar observed flux. Williams (1994) explained by existence of cool, high-density blobs of neutral material embedded in the nova shell. In Fig. 7, we show a comparison between the [O I]  $\lambda\lambda$ 5577,6300 and 6363 lines. While the  $\lambda$ 5577 and  $\lambda$ 6300 lines show similar spiky structures at same ejecta velocities on the top of their emission line profiles, in the [O I]  $\lambda$ 6363 line the spiky structures are less pronounced with the red wing of the line extending up to larger velocities. This strongly suggests the presence of an additional emission component, which we associate with the [Fe X]  $\lambda$ 6375 transition. This line is also a good proxy for the presence of a high-energy X-ray photon field, likely originating from a very hot underlying WD (Schwarz et al. 2011). In the Neil Gehrels Swift observatory XRT observations, the maximum of the X-ray supersoft (SSS) emission is around Day 150 from the discovery of the nova (Beardmore et al. 2017). One year after the nova discovery the SSS emission is still active but with a count rate flux in X-rays less than more of one order of magnitude (Orio et al. 2017).

The Day 155 spectrum is the last high-resolution spectrum in our data set, and we can estimate some important physical properties of the nova ejecta from it. We used emission line diagnostics



**Figure 7.** ASASSN-16kt spectrum at Day 155. The figure displays the spectrum centred at O I  $\lambda$ 6363 (blue line), O I  $\lambda$ 6300 (black line), and O I  $\lambda$ 5577 (red line) plotted on the velocity scale. The fluxes have been scaled to show the common structures at the top of their emission line profile, as well as the wider red wing of the [O I]  $\lambda$ 6363 line, suggesting the presence of an underlying [Fe X]  $\lambda$ 6375 line. The gap between 6355 Å and 6375 Å is due to the instrument (X-Shooter in this case).

that are usually developed for the analysis of H II regions as well as planetary nebulae and active Galactic nuclei to determine the electron density and the hydrogen mass ejected in the nova outburst. The lack of typical electron density indicators in the spectrum led us to use the nebular lines of [N II]  $\lambda\lambda$  5755,6548,6584 and



**Figure 8.** The results of the method used to de-blend [O III]  $\lambda$ 4363 (upper panel) and the [N II]  $\lambda\lambda$ 6548,6584 (lower panel) from H  $\gamma$  and H  $\alpha$  emission line, respectively, as described in Section 5.

[O III]  $\lambda\lambda$  4363,4960,5007 to determine the electron density  $n_e$  and temperature  $T_e$ . However, the [N II]  $\lambda\lambda$  6548,6584 and [O III]  $\lambda$ 4363 lines are blended with H  $\alpha$  and H  $\gamma$ , respectively. To disentangle these lines, we used the H  $\beta$  line as a template to rescale their fluxes by taking their blue wing as reference, see Fig. 8. The results are good for the [O III] line while the [N II] lines are not clearly resolved; however, for the estimate of the electron density and temperature, we need their sum, hence we measured the integrated flux of the observed excess as due to both [N II] lines. De-reddened fluxes and full width at half-maximum (FWHM) for the main emission lines observed in the Day 155 spectrum are reported in Table 4.

Using the formulas provided in Osterbrock & Ferland (2006):

$$\frac{I(\lambda 4959) + I(\lambda 5007)}{I(\lambda 4363)} = \frac{7.90 \exp 32900/T_e}{1 + 4.5 \times 10^{-4} \frac{n_e}{T_e^{0.5}}}, \quad (3)$$

$$\frac{I(\lambda 6548) + I(\lambda 6584)}{I(\lambda 5755)} = \frac{8.23 \exp 25000/T_e}{1 + 4.4 \times 10^{-3} \frac{n_e}{T_e^{0.5}}}, \quad (4)$$

which are typically used for H II regions and planetary nebulae but equally applied to the case of novae (see e.g. Della Valle et al. 2002), we obtain an electron density of  $n_e \sim 2 \times 10^6 \text{ cm}^{-3}$  and a temperature of  $T_e \sim 12000 \text{ K}$ .

**Table 4.** De-reddened fluxes and FWHM for bright emission lines measured in the UVES spectrum of Day 155. Fluxes are in units of  $10^{-13} \text{ erg cm}^{-2} \text{ s}^{-1}$  while FWHM are in  $\text{km s}^{-1}$ .

Line	$\lambda_0$	Flux	FWHM
[Ne v]	3346	59.40	2600
[Ne v]	3426	173.33	2600
[Ne III]	3869	42.124	2550
[Ne III]	3967	13.67	2600
[O III]	4363	37.60	3100
H $\beta$	4861	7.43	2550
[O III]	4959	18.27	–
[O III]	5007	58.75	3100
[N II]	5755	3.58	2550
He I	5876	0.51	3000
[N II]	6548+6584	3.06	–
H $\alpha$	6562	22.24	2550
He I	10830	12.29	2500
He I	12528	1.29	3700
Pa $\beta$	12818	1.73	2800

The presence of very bright [Ne v] lines suggests a large abundance of neon in the nova ejecta, pointing to a ONe WD progenitor for ASASSN-16kt. We have computed the ionic abundances for Ne using the observed forbidden lines at [Ne III]  $\lambda\lambda$ 3869,3967 and [Ne v]  $\lambda\lambda$ 3346,3426, respectively. We also note that [Ne IV] lines are likely blended with other lines, including the He I  $\lambda$ 4713, so we have not considered these lines in our computation. In this light, the final results will provide lower limit to the actual ionic abundance of Ne.

We used the NEBULAR.IONIC task (Shaw & Dufour 1995) within the IRAF package to estimate Ne ionic abundances based on the H $\beta$  emissivity (De Robertis, Dufour & Hunt 1987). Using the values for the temperature and density obtained before, we obtain the ionic abundances, with respect to hydrogen, of  $N([\text{Ne v}])/N(\text{H}) = 1.22 \times 10^{-3}$  and  $N([\text{Ne III}])/N(\text{H}) = 3.59 \times 10^{-4}$ . This results in a Ne abundance of  $N(\text{Ne}) \geq 9.2$ , relative to hydrogen,  $N(\text{H}) = 12$ . The solar Ne abundance is  $N(\text{Ne}) = 8.05 \pm 0.10$  (Lodders, Palme & Gail 2009). Consequently, we estimate a Ne abundance in the ejecta of ASASSN-16kt which is  $\sim 14$  times the solar one. This value is comparable to that of the ONe nova V1974 Cyg of  $N(\text{Ne}) \sim 12$  (Gehrz et al. 1994), where the authors used infrared emission lines of Ne to estimate the abundance.

### 5.1 The ejected mass of <sup>7</sup>Be II

Following Mustel & Boyarchuk (1970) we can estimate the mass of the hydrogen envelope ejected in the outburst. To this aim, we first need the knowledge of the geometry of the nova ejecta. It is well known that nova ejecta are not spherical (Cohen & Rosenthal 1983; Slavin, O’Brien & Dunlop 1995; Gill & O’Brien 1998, 2000; Sahman et al. 2015), evidence confirmed also at radio frequencies (Chomiuk et al. 2014). In fact, nova ejecta frequently show equatorial rings and/or polar caps condensations, which make them more similar to oblate or prolate spheroids than spheres. In what follows, we assume a spherical geometry as a first-order approximation for the nova ejecta, but the other driving parameter is given by the filling factor, which quantifies the degree of condensations inside the nova ejecta. Consequently, we first assume that the volume  $V$  of the ejecta can be approximated by a spherical shell with radius  $R = v_{\text{ej}} \Delta t$ , where  $v_{\text{ej}} \approx 2300 \text{ km s}^{-1}$  is the expansion velocity

measured from the Day 155 spectrum and  $\Delta t = 13.4$  Ms is the time since outburst, and with thickness of the shell  $\delta = aR\epsilon$ , where  $a = 0.17$  corresponds to the observed ratio between the width of Balmer lines P-Cygni profiles and their blue-shifted velocity, while  $\epsilon$  is the filling factor. The observed flux of the recombination Balmer lines (in this case we consider H  $\beta$ ) is then given by

$$I_{H\beta} = \left[ \left( \frac{4\pi j_{H\beta}}{n_e n_p} \right) n_e n_p V \right] \frac{1}{4\pi d^2}, \quad (5)$$

with the term  $\frac{4\pi j_{H\beta}}{n_e n_p}$  being the H  $\beta$  emissivity (Osterbrock & Ferland 2006) and  $d$  the distance to the nova. Assuming  $n_e = n_p$ , we can estimate the filling factor from the previous formula, being  $\epsilon = 0.15$  at this epoch, and then the mass of the hydrogen ejected in the outburst  $M_{H,ej} \approx 5.5 \times 10^{28}$  g, corresponding to  $M_{H,ej} = 2.7 \times 10^{-5} M_{\odot}$ .

We can finally compute the mass of Ca following the semi-empirical analysis developed by Izzo et al. (2015), where different nova composition models have been considered to estimate the mass of synthesized elements compared to the ejected H mass. For a generic ONe WD progenitor with a mass varying from 1.0 and 1.35  $M_{\odot}$ , the mass of Ca varies between  $(1.66\text{--}1.82) \times 10^{-5}$  of the total ejecta mass, and between  $(5.1\text{--}6.8) \times 10^{-5}$  the ejected H mass, according to the numerical results of Politano et al. (1995). These values correspond to an ejected Ca mass of  $(1.38\text{--}1.83) \times 10^{-9} M_{\odot}$ . Assuming the ejected H mass computed above and a mass abundance ratio of 4.24 for  ${}^7\text{Be II}$  with respect to Ca as estimated in the previous section, we obtain a total mass of  ${}^7\text{Be II}$  ejected in ASASSN-16kt of  $(5.9\text{--}7.7) \times 10^{-9} M_{\odot}$ , which is of the same order of magnitude of the  ${}^7\text{Be II}$  mass estimated in V5668 Sgr (Molaro et al. 2016) and one order of magnitude larger than the Li mass estimated in V1369 Cen (Izzo et al. 2015), but in this latter case only the observations of the  ${}^7\text{Li } \lambda 6708$  line were considered.

## 6 DISCUSSIONS

All novae with positively identified  ${}^7\text{Be II}$  and/or Li to date have been CO novae. Progenitors of CO novae have generally smaller masses compared to ONe novae (Iben & Tutukov 1985; Straniero, Piersanti & Cristallo 2016). Moreover, a less massive WD progenitor needs a larger amount of accreted matter to reach the critical pressure to hence ignite the TNR (Fujimoto 1982; Prialnik et al. 1982; Truran & Livio 1986). Consequently, CO novae are characterized by more massive ejecta than ONe novae (typically 10 times larger or more). In turn, the mass of  ${}^7\text{Be II}$  produced in ONe novae should be less than that produced in CO novae, as expected from numerical simulations (Hernanz et al. 1996). The  ${}^7\text{Be}$  mass estimated for ASASSN-16kt is indeed less than the  ${}^7\text{Be}$  mass in V5668 Sgr (Molaro et al. 2016), which is in agreement with the above picture.

Another consequence of the above scenario is that the luminosity a nova reaches at maximum correlates with the WD progenitor mass: the larger the WD mass, the smaller the accreted envelope mass, as well as the total binding energy (Truran & Livio 1986). Moreover, for massive WDs, the luminosity also anticorrelates with the accretion rate: the newly-discovered class of very fast and faint novae (Kasliwal et al. 2011) represent the main outcome of this physical scenario for CNe.

An analytical formulation relating the absolute  $B$ -band magnitude at maximum and the WD progenitor mass was proposed by Livio (1992): for a distance of  $d = 10.0$  kpc, we have  $M_B = -9.1$  mag, after correcting the observed value  $B = 6.86$  mag for reddening, and then a WD mass of  $M_{WD} = (1.20 \pm 0.05) M_{\odot}$ . Assuming this value for the WD progenitor mass, according to Truran &

Livio (1986), we obtain a recurrence time for ASASSN-16kt of  $T_{\text{rec}} \sim 3 \times 10^4$  yr. When applying the same formulae to V5668 Sgr, we find a WD progenitor mass of  $0.84 M_{\odot}$ , and a recurrence time of  $T_{\text{rec}} \sim 3 \times 10^5$  yr, which is almost one order of magnitude larger than ASASSN-16kt. If we take into account the estimated  ${}^7\text{Be}$  yield produced by ASASSN-16kt and its recurrence time, and the observed evidence that fast novae represent  $\sim 30$  per cent of the total novae in the Galaxy, we conclude that ONe novae produce a similar quantity of  ${}^7\text{Be II}$  than the one produced by CO novae, like V5668 Sgr.

## 7 CONCLUSIONS

ASASSN-16kt is a very fast nova ( $t_2 = 5.9 \pm 0.1$  d) at a distance of  $d = 10.9 \pm 1.6$  kpc. The first spectra obtained five to eight days from its discovery reveal that ASASSN-16kt is an hybrid nova in the CTIO spectral classification (Williams 1992): the identification of Fe II lines is due to the presence of blue-shifted absorption lines at the same expanding velocities reported for other lines, like Na I, Ca II, and also  ${}^7\text{Be II}$ . The two lines of the  ${}^7\text{Be II}$  doublet are clearly seen and not contaminated by blends for the components at  $v_{ej,1} = -2030$  km s $^{-1}$  and  $v_{ej,2} = -2370$  km s $^{-1}$ . Moreover, the late spectra show it is a ONe nova, due to the presence of very bright [Ne III] and [Ne V] lines and from the Ne abundance value derived from their analysis. Given its height above the Galactic plane ( $h = 1.7$  kpc), it is also a bulge nova, according to della Valle & Livio (1998). However, bulge novae are characterized by distinct properties, such as being slower and fainter than ASASSN-16kt, and are usually characterized by a typical Fe II spectral class. ASASSN-16kt in this sense would represent an interesting outlier of this sub-class.

The detection of  ${}^7\text{Be}$ , which decays completely into  ${}^7\text{Li}$  through electron capture (Giraud et al. 2007), in the early spectra of the ONe nova ASASSN-16kt further confirms that novae represent the main Li-factories in the Galaxy. The estimate of the mass of  ${}^7\text{Be II}$  ejected is in line with other values published in literature (Izzo et al. 2015; Molaro et al. 2016), implying that the contribution of ONe novae to Galactic lithium enrichment is not negligible compared to CO novae, but given also their higher recurrence frequencies, ONe novae produce a similar quantity of  ${}^7\text{Be II}$  as CO novae. The evidence that  ${}^7\text{Be II}$  has been identified only at early epochs in ASASSN-16kt implies that the detection of  ${}^7\text{Be II}$ , is possible only during the epochs of the nova characterized by the presence of low-ionization elements, i.e. during the maximum peak and the diffuse-enhanced phase (McLaughlin 1942; Gaposchkin 1957), or following the CTIO classification (Williams 1992) when Fe II lines are present in the spectrum.

A more precise characterization of the  ${}^7\text{Be II}$  and Li mass produced during a nova outburst as a function of the progenitor WD mass is still needed, in order to precisely quantify the total Galactic yield of lithium provided by novae. The large quantity of  ${}^7\text{Be II}$  measured in the few novae observed so far suggests a very large overabundance if we assume that all novae produce approximately the same quantity of  ${}^7\text{Be}/{}^7\text{Li}$  in each outburst. These results represent an important input for numerical simulations of TNR nucleosynthesis: according to the results presented in Jos e & Hernanz (1998), the quantity of  ${}^7\text{Be II}$ , and hence Li, produced during the TNR varies in between  $\sim 10^{-5}\text{--}10^{-6}$  times the ejected H mass in the outburst, which is a value  $\sim$  an order of magnitude lower than the observed quantity for both CO and ONe novae.

These findings imply that classical novae produced a huge quantity of Li, more than the value of  $150 M_{\odot}$  of Li estimated to be in the

Milky Way (Fields, Molaro & Sarkar 2014). If confirmed by more detections of Be and Li in future classical novae, we could even be faced with a new ‘lithium problem’: since  ${}^7\text{Be}$  decays only via electron-capture producing  ${}^7\text{Li}$  (Giraud et al. 2007), some depletion mechanisms must work in the immediate phases after its formation in order to destroy a sufficient quantity of  ${}^7\text{Li}$ , without claiming the existence of more exotic physics (but see Krasznahorkay et al. 2016). Additional detections of  ${}^7\text{Be II}$  and/or  ${}^7\text{Li}$  combined with a detailed study of the ejecta properties, like its degree of asphericity (Ribeiro et al. 2011; Chesneau et al. 2012; Shore et al. 2013), are needed to quantify the nova Galactic yield, and then to infer the presence of mechanisms acting during the TNR, or in the nova ejecta, that can deplete freshly-formed Li, and finally explain the over-abundances of Li observed in young stellar populations (Spite 1990).

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