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Multi-epoch VLBI of a double maser super burst

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In a rare and spectacular display, two well-known massive star forming regions, W49N and G25.65+1.05, recently underwent maser 'super burst' - their fluxes suddenly increasing above 30,000 and 18,000 Jy, respectively, reaching several orders of magnitude above their usual values. In quick-response, ToO observations with the EVN, VLBA and KaVA were obtained constituting a 4 week campaign - producing a high-cadence multi-epoch VLBI investigation of the maser emission. The combination of high-resolution, polarisation and flux monitoring during the burst provides one of the best accounts, to date, of the maser super burst phenomenon, aiding their use as astrophysical tools. These proceedings contain the preliminary results of our campaign.

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8-11 October 2018

Granada, Spain

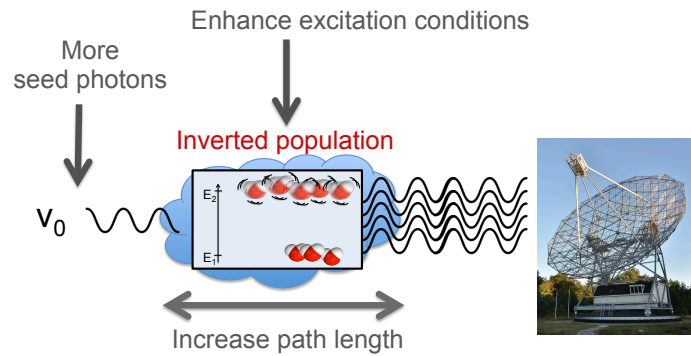


Figure 1: Illustration of the main factors influencing the observed maser flux. Photo credit: Paul Boven

1. Introduction

The maser ‘super burst’ describes an event where a maser emitting region exhibits a sudden increase in flux density of several orders of magnitude in the space of months, weeks, days, or even hours (*ex.* [1, 2, 3, 4, 5]). Currently there are 5 recognised super burst star forming regions (SFRs): Orion KL, NGC 6334I, S255, W49N and G25.65+1.05.

Super bursts constitute rare, transient events and consequently their mechanism of action remains enigmatic. The widely cited Deguchi & Watson model [6] generally allows for three avenues of flux enhancement: an increase in the incident continuum photon flux entering the maser; mechanical/radiative induction of more favourable maser pumping conditions; or an increase in the path length of the maser along the line of sight to the observer. These three scenarios are illustrated in the schematic diagram in Figure 1. Using observations of the super burst phenomenon we hope to distinguish which of these processes is taking place in the observed maser sources.

During the IAU symposium 336 in September 2017 two well-known massive star forming regions, W49N and G25.65+1.05, simultaneously underwent maser super bursts - their fluxes increasing several orders of magnitude above their usual values. These proceedings present early results from an observing campaign investigating the mechanism of these maser super bursts.

2. Observations

Burst activity of $> 18,000$ Jy was reported by the single-dish Maser Monitoring Organisation (M2O) on the 7-8th of September, after which the maser rose to 65,000 Jy in late September [7, 5]. Target of opportunity observations of the 22 GHz water maser in G25.65+1.05 were requested to the European VLBI network (EVN), the KVN and VERA Array (KaVA) and the Very Long Baseline Array (VLBA) and observations were conducted on the 2nd, 11th and 28th of October 2017, respectively. Data were correlated by the representative correlators for each array and processing was conducted with AIPS (Astronomical Image Processing System) details of which will be presented in a forthcoming series of papers. Maser self-calibration was performed without phase referencing to a continuum source, thus the maps center on a common reference maser.

*Speaker.

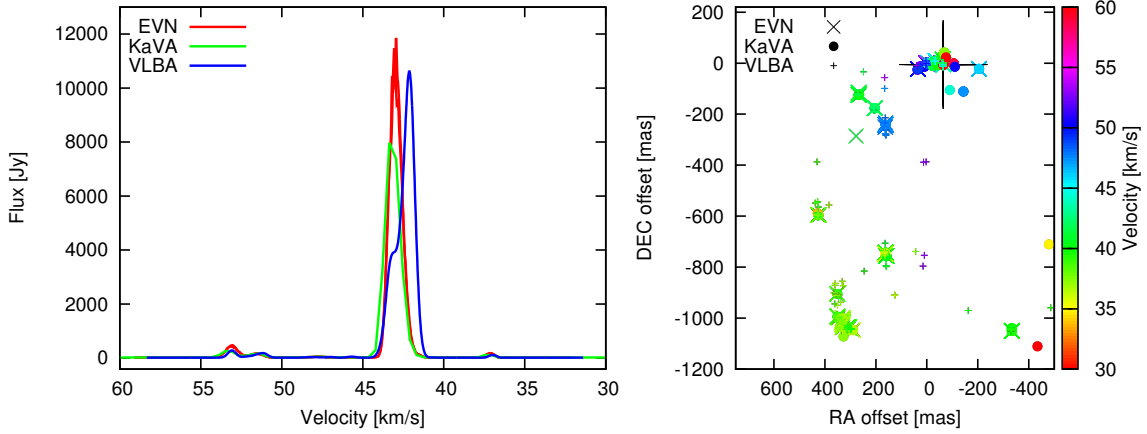


Figure 2: *Left* shows the scalar averaged spectrum of maser emission in G25.65+1.05 for each of the arrays for observing dates: 2nd October (EVN), 11th October (KaVA), 28th October (VLBA). *Right* shows the distribution of maser spots in G25.65+1.05 where colour scale indicates LSR velocity. A black cross indicates the position of the bursting maser feature.

3. Discussion

A preliminary maser spotmap is shown in Figure 2. Maser emission is seen to form two arc structures at $(RA, DEC) = (0, 0)$ and $(RA, DEC) = (350, -1000)$ mas offset. These are both sites of continuum emission revealed by JVLA observations by Bayandina et al. (in prep) and are thought to be regions of ionised gas either of a stellar origin or relating to feedback-driven shocks. Such shocks likely provide the collisional pumping energy required to induce maser action. The super burst maser resides in the center of an arc of masers that extends E-W, resembling a shock surface.

The flux density of the bursting feature is seen to vary by several hundreds of Jy during the observation campaign - indeed G25.65+1.05 is known to be extremely variable [5]. Thus the change is too rapid to be explained as a change in physical conditions such as incident photons and shock formation - masers tracing such conditions are generally more persistent [8]. It can also be noted from the spectrum in Figure 2 that significant flux variation is seen exclusively in the bursting maser spectral feature; the low-flux spectral features remain constant during the campaign. Thus, the exclusivity of the flux variations to one spectral (and spatial) component requires whatever mechanisms to be driving the super burst to be strictly localised to a sub-milliarcsecond region.

A comparison of the simultaneous total flux measured by single dish instruments, with the VLBI cross correlation flux indicates the compactness of the maser emission. During the VLBI campaign, VLBI flux recovery of up to 100% was observed (during the VLBA measurement) indicating highly compact emission. Furthermore, the masers in G25 were shown to be unsaturated [5], which is a fundamental requirement for non-linear flux enhancement in the maser overlap scenario.

Taken together, the observational data strongly favor the explanation of the super burst in G25.65+1.05 to have been caused by the overlapping of maser features along the line of sight to the observer. Such a scenario, in the context of extreme daily flux variability, leads to a method to

constrain the angular size of the maser overlap region despite it being unresolved by dividing the relative proper motions of maser features involved in the overlap by the burst duration. In the case of $\sim \text{mas yr}^{-1}$ proper motions typical to VLBI studies [9], daily viability would imply an overlap region of $1/365$ mas. True proper motion analyses, in addition to continued flux monitoring, are required to make a true estimate.

Further analysis of the observational data including structural analyses, proper motions and observations with various other instruments will be able to further clarify the origin of the 22 GHz water maser super burst in G25.65+1.05.

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