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Reverberation of pulsar wind nebulae – III. Modelling of the plasma interface empowering a long term radiative evolution

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ABSTRACT

The vast majority of pulsar wind nebulae (PWNe) present in the Galaxy is formed by middle-aged systems characterized by a strong interaction of the PWN itself with the supernova remnant (SNR). Unfortunately, modelling these systems can be quite complex and numerically expensive, due to the non-linearity of the PWN–SNR evolution even in the simple one-dimensional (1D)/one-zone case when the reverse shock of the SNR reaches the PWN, and the two begin to interact (and reverberation starts). Here, we introduce a new numerical technique that couples the numerical efficiency of the one-zone thin shell approach with the reliability of a full ‘Lagrangian’ evolution, able to correctly reproduce the PWN–SNR interaction during the reverberation, and to consistently evolve the particle spectrum beyond. Based on our previous findings, we show that our novel strategy resolves many of the uncertainties present in previous approaches, as the arbitrariness in the SNR structure, and ensure a robust evolution, compatible with results that can be obtained with more complex 1D dynamical approaches. Our approach enable us for the first time to provide reliable spectral models of the later compression phases in the evolution of PWNe. While in general, we found that the compression is less extreme than that obtained without such detailed dynamical considerations, leading to the formation of less structured spectral energy distributions, we still find that a non-negligible fraction of PWNe might experience a super-efficient phase, with the optical and/or X-ray luminosity exceeding the spin-down one.

Key words: radiation mechanisms: non-thermal – methods: numerical – pulsars: general – ISM: supernova remnants.

1 INTRODUCTION

Pulsars (PSRs) are rapidly rotating and strongly magnetized neutron stars (NSs). The combination of rotation and magnetic field induces electric fields strong enough to pull particles out of the NS surface, to accelerate them to ultrarelativistic speeds, and ultimately to drive a pair production cascade that results in the formation of a relativistic pair plasma wind (Michel 1973; Timokhin 2006). The interaction of this wind with the environment (the parent SNR in the case of young PSRs), leads to the formation of relativistic wind bubbles, known as pulsar wind nebulae (PWNe) where the accelerated particles shine through non-thermal synchrotron and inverse Compton (IC) emission, with a very broad-band spectrum extending from radio to TeV and PeV energies (Gaensler & Slane 2006; Slane 2017; Olmi & Bucciantini 2023).

PWNe thus constitute a unique laboratory where high-energy astrophysical processes, non-thermal emission, and relativistic plasma physics can be studied in great detail. For this same reason a

plethora of different approaches have been put forward throughout the years, ranging from analytical to numerical one-zone models (Pacini & Salvati 1973; Reynolds & Chevalier 1984; Gelfand, Slane & Zhang 2009; Bucciantini, Arons & Amato 2011; Torres et al. 2014), from one-dimensional (1D) hydrodynamics analytical (Rees & Gunn 1974; Kennel & Coroniti 1984a, b; Emmering & Chevalier 1987) to 1D relativistic magnetohydrodynamic (MHD) numerical (van der Swaluw et al. 2001; Bucciantini et al. 2003; van der Swaluw, Downes & Keegan 2004; de Jager, Ferreira & Djannati-Ataï 2008; Bandiera et al. 2023), from two-dimensional (2D) MHD (Del Zanna, Amato & Bucciantini 2004; Komissarov & Lyubarsky 2004; Bogovalov et al. 2005; Del Zanna et al. 2006) to the more recent three-dimensional (3D) relativistic MHD (Porth, Komissarov & Keppens 2014; Olmi et al. 2016; Barkov, Lyutikov & Khangulyan 2019; Olmi & Bucciantini 2019). For an exhaustive review on the various methods see Olmi & Bucciantini (2023).

The fact that different approaches keep being developed and used, is a clear testimony to the fact that in the study of PWNe there is still no such thing as an *all-purpose optimal strategy*, but that, depending on the issue at stake, different approaches might prove more suitable. There is now a general consensus that at least 2D, but more appropriately 3D, relativistic MHD is required in order

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to investigate the detail of the dynamics in the immediate post-termination shock region, and how this leads to the observed high energy X-ray jet-torus morphology of many systems (Weisskopf et al. 2000; Gaensler et al. 2002; Lu et al. 2002; Romani & Ng 2003; Camilo et al. 2004; Slane et al. 2004; Romani et al. 2005). However, this approach is so numerically expensive (millions of CPU hours to simulate just a few hundreds of years of evolution), that it is at the moment impossible to use it to model the long term evolution of even a single PWN. Moreover, none of these multidimensional hydrodynamic (HD) or MHD models consistently includes the treatment of radiative losses and the dynamical spectral evolution of the source. Indeed, studies aimed at investigating the long term spectral evolution of many systems, or that requiring the computation of several hundreds of models, as the case of population studies, as in Fiori et al. 2021, or multiple spectral fitting, as in Torres et al. 2014, are carried out using the so called one-zone models, where the PWN is treated as a uniform bubble, without internal structure, expanding within its parent SNR and subject to adiabatic and radiative losses.

All existing one-zone models adopt the same approximation, known as the ‘thin-shell approximation’ (see e.g. Reynolds & Chevalier 1984; Bucciantini et al. 2004; Gelfand et al. 2009; Martin, Torres & Pedalletti 2016; Torres 2017). As the PWN expands inside the parent SNR, it piles up matter into a shell (the swept-up shell). In the thin-shell approximation, such shell is assumed to be infinitely thin, and thus to perfectly trace the radius of the PWN. The shell evolution (a proxy for the PWN) can be obtained by solving the momentum conservation for the shell, which provides the acceleration (or deceleration) of the shell itself, due to the combined action of the PWN pressure from the inside, which acts as a piston pushing out, and the confinement of the outer SNR material. The complete set of equations and boundary conditions can be found e.g. in Bandiera et al. 2020, Paper I hereafter: a summary of them will be presented in the following section. The thin-shell approximation has proven to be highly reliable in the study of young systems, where the PWN evolves within the cold self-similar expanding SNR ejecta, the so called free-expansion phase (Gaensler & Slane 2006). Indeed, it maps very well the 1D dynamics of young PWNe, as found by numerical simulations (van der Swaluw et al. 2001; Bucciantini et al. 2003). The reason is twofold: in this phase the thickness of the shell is indeed much smaller than its radius (hence the correctness of the approximation), but even more important is the fact that the properties of the surrounding SNR ejecta are well-defined (namely they are cold and the expansion is self-similar, Chevalier 1982). This is quite important because all existing one-zone thin-shell models prescribe the outer structure of the SNR with simple analytical recipes.

Unfortunately, while the use of simplified recipes for the SNR is very robust in the early free-expansion phase, for the later reverberation phase, when the PWN interacts directly with the SNR shell, there is no consensus on how to model the confining environment, and this arbitrariness reflects in their increasing unreliability as the evolution proceeds beyond the free-expansion phase. Unfortunately, full 1D HD models, able of capturing the full dynamics of the PWN–SNR system, are still too computationally expensive for population synthesis or multiwavelength broad-band spectral fittings. But, at the same time, the main problem with one-zone thin-shell models is that the internal dynamics in the SNR shell, and the feedback the PWN exerts on it upon interaction, are too complex for any simple recipe to fully capture them.

Recently Bandiera et al. (2023, Paper II hereafter) have presented non-radiative results based on one-zone models that however got rid of the thin-shell approximation, and solved for the correct structure

of both the swept-up shell and the confining SNR using a full Lagrangian approach. This allows for the one-zone evolution of the PWN to be computed taking into account the correct SNR structure. Despite solving completely the ambiguity in the late evolution of all previous models, this approach is still too computationally expensive (it takes about a day on a single CPU to compute the full long term evolution of a single system). While this is undoubtedly the optimal avenue if one is interested in modelling just a few systems, it is not if one needs to compute hundreds or thousands of models, do fittings, or population analysis. The major issue with this full Lagrangian approach is that matter in the swept-up shell undergoes a strong compression, especially in the early phase of evolution, and this imposes a heavy computational burden to model those evolutionary phases. Interestingly, this is the same phase where the thin-shell approach has proved both robust and reliable.

Driven by these considerations, we are presenting here a strategy for PWNe modelling that combines the best of both worlds: the efficiency and speed of the thin-shell approach and the reliability of the full Lagrangian evolution. This is done by combining a thin-shell model for the PWN and swept-up shell (Martin et al. 2016; Martin & Torres 2022), with a Lagrangian model for the surrounding SNR (Bandiera et al. 2021). Our approach is not as fast as a simplified one-zone thin-shell model (where a full PWN evolution can be run in a matter of few seconds), but fast enough (a full evolution can be run in just a few minutes) to perform reliable population studies with multiwavelengths spectral information and fittings. We will show, by comparing with a full Lagrangian approach, that this strategy is able to correctly reproduce the entire evolution of these systems from the very early free-expansion phase, to the late reverberation, capturing with accuracy the dynamics of the complex PWN–SNR interaction.

2 THIN-SHELL LAGRANGIAN CODE: TIDE + L

Here, we briefly illustrate the new computational strategy and how the thin-shell models are coupled with a full Lagrangian evolution of the SNR.

The evolution of the PWN can be divided into two phases: the early free-expansion and the late reverberation. The time at which the latter begins and the former ends, t_{begrev} , occurs when the PWN swept-up shell (whose radius, R_{sh} , is a proxy for the PWN radius R_{pwn} in the thin-shell approximation) equals the radius of the SNR reverse shock, R_{RS} .

During the early free-expansion phase the swept-up shell evolves according to the standard thin-shell equations. One of those describes the conservation of mass:

$$M_{\text{sh}}(t) = 4\pi \int_0^{R_{\text{sh}}(t)} \rho_{\text{ej}}(t, r) r^2 dr = M_{\text{ej}} \left(\frac{R_{\text{sh}}(t)}{V_{\text{ej}} t} \right)^3, \quad (1)$$

where $M_{\text{sh}}(t)$ is the mass in the shell at time t , $\rho_{\text{ej}}(t, r)$ is the density distribution in the ejecta, and the last equality holds for ejecta with a uniform density profile. In the same limit, the ejecta maximum velocity is given as a function of the ejecta mass M_{ej} and the supernova energy E_{sn} as: $V_{\text{ej}}^2 = 10E_{\text{sn}}/(3M_{\text{ej}})$.

Another equation describes the momentum conservation, under the mutual action of the PWN pressure $P_{\text{pwn}}(t)$ from the inside and the confining ejecta from the outside:

$$\frac{d(M_{\text{sh}}(t)\dot{R}_{\text{sh}}(t))}{dt} = 4\pi P_{\text{pwn}}(t) R_{\text{sh}}^2(t) + \frac{dM_{\text{sh}}(t)}{dt} \frac{R_{\text{sh}}(t)}{t}, \quad (2)$$

where $\dot{R}_{\text{sh}}(t)$ is the expansion speed of the shell. The pressure in the PWN is computed taking into account the energy injection from the

pulsar, as well as the adiabatic, radiation, and diffusion losses as:

$$\frac{d}{dt} (4\pi R_{\text{sh}}(t)^3 P_{\text{pwn}}(t)) = L_{\text{psr}}(t) - 4\pi R_{\text{sh}}(t)^2 \dot{R}_{\text{sh}}(t) P_{\text{pwn}}(t) - J(t), \quad (3)$$

where $L_{\text{psr}}(t)$ is the PSR spin-down luminosity, and $J(t)$ factors in radiation and diffusion losses.

This set of equations is solved until $R_{\text{sh}}(t) = R_{\text{RS}}(t)$, where $R_{\text{RS}}(t)$ is given, in the case of uniform ejecta, by equation 22 in Paper II, which is itself based on a fitting of the numerical solution for the full SNR evolution (a more general formula for R_{RS} can be instead found in Bandiera et al. 2021).

The same numerical solution for the full SNR evolution, not only provides the location of the reverse shock, but also the full structure (density, pressure, and velocity) of the SNR, from $R_{\text{RS}}(t)$ outward. This structure is the one used to initialize a Lagrangian code for the evolution of the matter located outside $R_{\text{sh}}(t)$. Moreover, at this same time, we have the size $R_{\text{pwn}}(t) = R_{\text{sh}}(t)$, the expansion rate $\dot{R}_{\text{sh}}(t)$, the mass of the shell $M_{\text{sh}}(t)$, as well as the interior pressure $P_{\text{pwn}}(t)$. These are used to set the initial inner boundary condition of the Lagrangian code, that evolves the matter outside, in the SNR.

For the reader's convenience, we briefly summarize here the equations that are solved (Mezzacappa & Bruenn 1993; Bandiera et al. 2023). The time evolution of the velocity (v) and radius (r) of the interfaces $i + 1/2$ between the shell i and the shell $i + 1$ is given by:

$$H_i^n = \eta_{\text{vnr}} \rho_i^n (v_{i+1/2}^n - v_{i-1/2}^n)^2 \Theta[v_{i-1/2}^n - v_{i+1/2}^n], \quad (4)$$

$$a_{i+1/2}^n = [(r_{i+1/2}^n)^2 (p_{i+1}^n - p_i^n) - (r_{i+1}^n)^2 H_{i+1}^n + (r_i^n)^2 H_i^n] / \Delta m_{i+1/2}, \quad (5)$$

$$v_{i+1/2}^{n+1} = v_{i+1/2}^n - 4\pi \Delta t a_{i+1/2}^n, \quad (6)$$

$$r_{i+1/2}^{n+1} = r_{i+1/2}^n + 4\pi \Delta t v_{i+1/2}^n + 2\pi (\Delta t)^2 a_{i+1/2}^n, \quad (7)$$

where $\Theta[\cdot]$ is the Heavyside function, Δt is the time interval between the steps n and $n + 1$, $\Delta m_{i+1/2}$ is the mass at the interface, defined as a function of the mass of the two bounding shells $\Delta m_{i+1/2} = (\Delta m_i + \Delta m_{i+1})/2$, $\eta_{\text{vnr}} = 2$ (the standard von Neumann–Richtmyer method), and H is the viscous pressure (Schulz 1964). The radius of each shell is defined as the barycenter radius:

$$r_i^n = \left(\frac{(r_{i+1/2}^n)^3 + (r_{i-1/2}^n)^3}{2} \right)^{1/3}, \quad (8)$$

and its density:

$$\rho_i^n = \frac{3\Delta m_i}{4\pi((r_{i+1/2}^n)^3 - (r_{i-1/2}^n)^3)}. \quad (9)$$

Instead, the pressure p_i in the shell is derived by solving (either by successive iterations or by direct analytic solution in the special case of an ideal gas) the following equation for the specific internal energy e_i :

$$e_i^{n+1} = e_i^n - \frac{p_i^{n+1} + p_i^n}{2} \left(\frac{1}{\rho_i^{n+1}} - \frac{1}{\rho_i^n} \right) + 2\pi \Delta t (r_i^{n+1} + r_i^n)^2 H_i^n \frac{v_{i+1/2}^n - v_{i-1/2}^n}{\Delta m_i}, \quad (10)$$

and assuming the following equation of state $p_i = 2\rho_i e_i/3$, appropriate for a perfect gas of adiabatic index $\Gamma = 5/3$. As stated before, the values of the quantities for the first interface are set, at the beginning of the Lagrangian evolution, to those provided by the thin-shell model at t_{begrev} , by identifying $r_{1/2} = R_{\text{pwn}}$, $v_{1/2} = \dot{R}_{\text{pwn}}$, $\Delta m_{1/2} = M_{\text{sh}}$, and $p_{1/2} = P_{\text{pwn}}$.

In practice, according to our treatment of the evolution during the reverberation phase, the PWN acts as a piston inside the SNR, in the same way as it was done for the Lagrangian simulations in Bandiera et al. (2023), except that now the piston includes the swept-up shell, which provides its starting inertia and momentum. The mass of the swept-up shell is supposed not to change during the reverberation, which is equivalent to the assumption that no matter is further accumulated. The evolution of $R_{\text{pwn}}(t)$ is simply that of the inner interface of the first cell of the Lagrangian algorithm.

The thin-shell model we used is the one of the TIDE code (see Martin & Torres 2022 and references therein), a one-zone time-dependent radiative code able to evolve the leptonic population of the PWN. TIDE solves the diffusion-loss equation considering adiabatic and radiative losses (synchrotron, IC, bremsstrahlung), plus diffusion (for a description of the equations see the following sections, and Martin, Torres & Rea 2012 for details of the radiative implementation). We shall call the version of TIDE as developed for this paper, TIDE + L, to easily distinguish it from thin-shell-only incarnation. TIDE + L is based on the modification of TIDE already described in Bandiera et al. (2020), to which we add a number of modules to introduce the Lagrangian treatment during reverberation, following the approach described earlier.

3 MODELS

In order to assess the performance of this new approach, we have selected a few representative cases for the evolution of the PWN–SNR system. We shall compare the results for these cases with those obtained via a full-Lagrangian evolution.

As it was pointed out in our previous works, all quantities can be scaled with the SNR characteristic units defined, following Truelove & McKee (1999), as:

$$R_{\text{ch}} = M_{\text{ej}}^{1/3} \rho_0^{-1/3}, \quad (11)$$

$$t_{\text{ch}} = E_{\text{sn}}^{-1/2} M_{\text{ej}}^{5/6} \rho_0^{-1/3}, \quad (12)$$

where $\rho_0 = 1.4 m_p n_{\text{ism}}$ is the mass–density of the interstellar medium (ISM), with m_p the proton mass, and n_{ism} the number density. As shown in Paper II, making use of this scaling each PWN–SNR system can be represented by a point in the $\log_{10}(\tau_0/t_{\text{ch}}) - \log_{10}(L_0/L_{\text{ch}})$ characteristic plane, where $L_{\text{ch}} = E_{\text{sn}}/t_{\text{ch}}$. The entire PWN–SNR population covers a roughly elliptic area in the same plane, as can be seen from Fig. 1.

Our benchmark models have been selected to range over this region, in order to sample diverse properties of the reverberation phase. They consist of three real sources plus three synthetic ones (see Fig. 1). The three real sources are: the Crab nebula, 3C58 (both well-known young systems, still in their free-expansion phase, with very detailed spectral data), and J1834.9–0846 (with a rather poorly covered spectrum but representative of a highly compressible system). Since the scope of this work is to validate our new approach, and not to recompute spectral fittings, we adopt the same parameters as in Paper I and Martin & Torres (2022) for 3C58. For synthetic sources we assume the same spectral parameters of the Crab nebula, which are needed to compute the dynamical evolution of the system when radiative losses are included.

The relevant parameters defining the properties of the selected sources are given in Table 1.

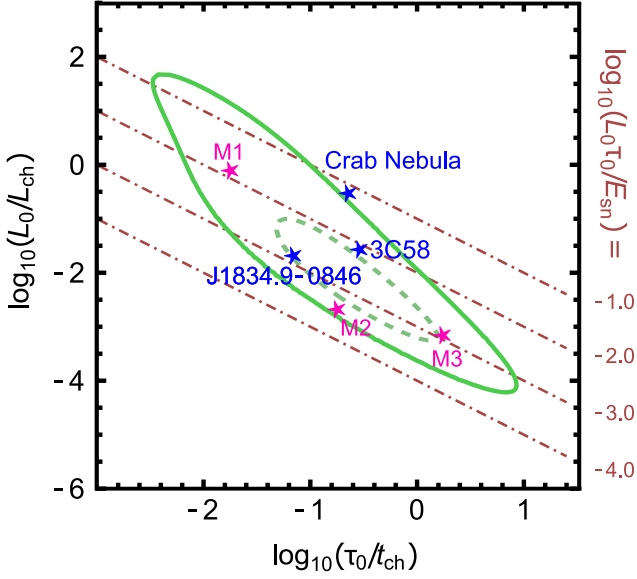


Figure 1. Plot of the PWNe population in the characteristic plane $\log_{10}(\tau_0/t_{\text{ch}}) - \log_{10}(L_0/L_{\text{ch}})$, as defined in Bandiera et al. (2023, Paper II). The green solid line shows the iso-level enclosing the 98 per cent of the population, while the dashed one the 50 per cent. The position of the sources selected as references for the present study is shown with a star symbol: blue stars for real sources, magenta ones for synthetic sources.ch

4 DYNAMICAL EVOLUTION

In this section, we compare the long term evolution (up to $\sim 8t_{\text{ch}}$) of our benchmark sources computed using different approaches and codes, namely:

- (i) the full, albeit non-radiative, Lagrangian code, in its original PWN + SNR version introduced in Paper II;
- (ii) the herein-developed TIDE + L, a fully-radiative, hybrid Lagrangian code;
- (iii) the new code TIDE + L, but used in the non-radiative regime simply obtained by neglecting all radiative and diffusive losses. To distinguish it from the radiative TIDE+L, we will refer to the no-radiative/no-losses version as to TIDE + L^{NL}.

The introduction of TIDE + L^{NL} allows us to make a direct comparison with the results of the pure Lagrangian approach, which is also non-radiative. It serves to cross-calibrate the dynamical part contained in TIDE, and gain confidence when the results are modified by the further appearance of the radiation losses.

The results for the dynamical evolution of all benchmark sources are shown in Fig. 2. For each source, we present the time evolution of the PWN radius and of the SNR characteristic surfaces (the contact discontinuity, CD, the reverse shock, RS, and the forward shock, FS), all in terms of characteristic units. In order to assess the validity of the thin-shell approximation, we also show the swept-up shell thickness, as computed by the full-Lagrangian code, where it is fully resolved.

The strength of the reverberation phase is measured through the compression factor of the PWN, namely:

$$\text{CF} = \frac{R_{\text{pwn}}^{\text{max}}}{R_{\text{pwn}}^{\text{min}}}, \quad (13)$$

always computed at the first compression event. The PWN reaches its maximum radius $R_{\text{pwn}}^{\text{max}}$ close to the time of the beginning of the reverberation, $t_{\text{beg, rev}}$, while $R_{\text{pwn}}^{\text{min}}$ is the PWN radius at the minimum

of the first compression. All these quantities are given in Table 2, for the various approaches.

The difference in the compression factor is very small in all cases when we compare the evolution computed with the Lagrangian code and TIDE + L^{NL}. We found that for those systems who preserve a rather thin-shell after the onset of the reverberation phase (namely Crab, 3C58, and M3), the CF computed with TIDE + L^{NL} is also very similar to the one obtained considering the radius of the barycenter of the shell, computed as:

$$R_{\text{bary}}(t) = \frac{4\pi}{M_{\text{swept}}} \int_{R_{\text{pwn}}(t)}^{R_{\text{sh}}(t)} \rho(r, t) r^3 dr, \quad (14)$$

where M_{swept} is the swept-up mass collected in the shell at $t = t_{\text{beg, rev}}$ and $\rho(r, t)$ the mass-density. This happens because in those cases the thin-shell approximation is particularly good, given the limited thickness of the shell; in fact, we see that at the time the PWN reaches its minimum radius at first compression $R_{\text{pwn}}^{\text{MIN}}$, that we named t_{MIN} , $\Delta R_{\text{sh}}(t_{\text{MIN}})/R_{\text{pwn}}^{\text{MIN}} = 0.04, 0.14, 0.33$, for Crab, 3C58, and M3 respectively, is always less than unity. On the contrary, this is no longer true when we look at the more compressible systems (namely J1834.9–0846, M1, and M2), for which the shell shows an evidently larger extension with a relative thickness with respect to the minimum radius at first compression always larger than unity: $\Delta R_{\text{sh}}(t_{\text{MIN}})/R_{\text{pwn}}^{\text{MIN}} = 2.9, 1.5, 2.8$. Then the thin shell approximation is not valid. Thus, the value of the CF computed with the shell barycenter radius is no longer compatible with that obtained at the PWN boundary.

The SNR characteristic curves appear almost perfectly coincident between the Lagrangian and TIDE + L results, with at most small variations in the CD when there is a larger deviation of the PWN radius after the first re-expansion caused by the presence of reflected shocks in between the PWN and the CD.

The changes introduced by the radiative losses are evident for all systems, except for 3C58 that is characterized by a very low magnetization according to the fitting. Losses reduce the energetics of the PWN since the very early free-expansion phase, leading to a slower expansion rate, and a delay of $t_{\text{beg, rev}}$. This is particularly evident for those systems that, like in the case of J1834.9–0846 or M1, have a short spin-down time $\tau_0 \ll t_{\text{ch}}$, and a large initial magnetic fraction η . As expected, all cases (except 3C58) show a larger compression factor in the radiative regime and, this changes the subsequent evolution.

We thus notice that, except for very compressible systems (e.g. J1834.9–0846 and M2), the variation of the CF induced by the radiative losses is rather modest, although nevertheless non-negligible in general. The ratio between the CF computed from TIDE+L and TIDE + L^{NL} from Table 2 amounts to a factor 1.28 for the Crab Nebula, 1.61 for M1, and 2.5 for M3, the less energetic of the three. Essentially no variation is indeed observed for 3C58. On the other hand, more compressible systems show larger variations as it might be expected, with a factor of 7.5 for J1834.9–0846 and 7.7 for M2, at the boundary of the population.

One might wonder why there is such a big difference in the compression experienced by J1834.9–0846, M2, and M3, all characterized by a rather similar energetics ($L_0\tau_0/E_{\text{sn}}$) and close to the lower boundary of the population. This can be actually understood comparing the initial spin-down time with the characteristic age of the systems: M2 and J1834.9–0846 have $\tau_0/t_{\text{ch}} \simeq 0.18$ and $\tau_0/t_{\text{ch}} \simeq 0.07$ respectively, while M3 has $\tau_0/t_{\text{ch}} \simeq 1.8$. This translates in a huge difference in the power that is still injected in the PWN at

Table 1. Parameters used in the simulations for the Crab Nebula, J1834.9–0846, and 3C58, as taken from Paper I and Martin & Torres (2022). For J1834.9–0846, the unique source which reverberation phase has already started at its actual age, the original best-fit spectral parameters are no more able to fit the spectral data, then we show in parentheses the value of the magnetic fraction needed to get a good one. In the last row, we also report the values of the magnetic field in the PWNe that we obtain at the actual age for the three real sources.

Parameter	Symbol	Crab PWN <i>real</i>	J1834.9–0846 <i>real</i>	3C58 <i>real</i>	M1 <i>synthetic</i>	M2 <i>synthetic</i>	M3 <i>synthetic</i>
Observed properties							
Spin-down period(s)	P	0.0334	2.48	0.0657	–	–	–
Period derivative	\dot{P}	4.2×10^{-13}	7.96×10^{-13}	1.93×10^{-13}	–	–	–
Age (yr)	t_{age}	969	7970	2500	–	–	–
Characteristic age (yr)	τ_c	1296	4900	5398	–	–	–
Distance (kpc)	d	2	4	2	–	–	–
Pulsar injection							
Braking index	n	2.51	2.2	3.0	2.33	2.33	2.33
Initial spin-down time (yr)	τ_0	758	280	2878	220	1000	2500
Initial spin-down luminosity (erg s^{-1})	L_0	3×10^{39}	1.74×10^{38}	9.3×10^{37}	2.13×10^{39}	1.2×10^{37}	1.6×10^{37}
Environment							
SN explosion energy ($\times 10^{51}$ erg)	E_{sn}	1	1	1	1	1	1
ISM density (particles/cm ³)	n_{ism}	0.5	0.5	0.1	0.02	0.5	5.0
SNR ejected mass (M_{\odot})	M_{ej}	9.0	11.3	17.2	12.0	16.9	8.0
Spectral model							
Far infrared temperature (K)	T_{fir}	70	25	25	70	70	70
Far infrared energy density (eV cm^{-3})	w_{fir}	0.1	0.5	0.22	0.1	0.1	0.1
Near infrared temperature (K)	T_{nir}	5000	3000	2900	5000	5000	5000
Near infrared energy density (eV cm^{-3})	w_{nir}	0.3	1	0.41	0.3	0.3	0.3
Energy break	γ_b	9×10^5	10^7	8.8×10^4	5×10^5	5×10^5	5×10^5
Low energy index	α_l	1.5	1	1	1.5	1.5	1.5
High energy index	α_h	2.54	2.1	3.012	2.5	2.5	2.5
Containment factor	ϵ	0.27	0.6	0.5	0.27	0.27	0.27
Magnetic fraction	η	0.02	0.045 (0.15)	0.0106	0.02	0.02	0.02
Characteristic values							
Time (yr)	t_{ch}	3329	4024	9766	12371.5	5623.4	1405.8
Radius (pc)	R_{ch}	8	8.7	17	25.9	9.9	3.6
$\text{Log}_{10}(\tau_0/t_{\text{ch}})$	r^*	–0.64	–1.16	–0.53	–1.75	–0.75	+ 0.25
$\text{Log}_{10}(L_0 t_{\text{ch}}/E_{\text{sn}})$	L^*	–0.5	–1.66	–1.54	–0.08	–2.67	–3.14
Derived quantities							
Magnetic field at t_{age} (μG)	B_{PWN}	100.0	2.2	2.0	–	–	–

$t_{\text{beg,rev}}$, namely: $3 \times 10^{-4} L_0$ for J1834.9–0846, $3 \times 10^{-3} L_0$ for M2, and $0.22 L_0$ for M3. This last system then is still enough powerful to react to the SNR compression when reverberation starts, while the other two can barely contrast the pressure exerted by the SNR. This is reflected in the higher CF we found for J1834.9–0846 and M2.

Systems with large compression $\text{CF} \gg 100$ are possible but rare, limited to the extremes of the known population of PWNe. In these, there will be extreme modifications of the PWN spectrum as time goes by, with significant increase of magnetic field and burn-out of electrons, as well as the appearance of super-efficiency, a phase characterized by an emitted luminosity at a given energy band that exceeds the pulsar spin-down power at the same age (Torres & Lin 2018).

In our non-radiative study of Paper II, we have produced maps of $\log_{10}(\text{CF})$ in the then-considered parameter space, respectively using a thin-shell model with different prescriptions for the pressure outside the PWN or our interpolating formula to our Lagrangian models. We refrain to do so here since the problem is complicated by the existence of radiation, i.e. the values of $\log_{10}(\text{CF})$ not only should account for the dynamical and energetics features, but also for the assumed magnetization. It is not the aim of this work to do a full

population analysis or a phase–space exploration of all combinations we could think are plausible in nature.

Radiative losses induce an extra-compression that is very mild for most of the population. On the other hand, there is limited sector of the population, namely the very low energetic systems, in which losses can lead to a substantial enhancement (up to 10 times) of the compression.

Finally, as for comparison with the present results, we want to recall the CFs we estimated in Paper I for the Crab nebula and J1834.9–0846, using the standard version of the TIDE code without the Lagrangian module that is based on the pure thin-shell approximation (similarly to all the other one-zone models available so far in the literature). What we obtained then was: a CF of 3.530 for the Crab (2.712 with no radiative losses) and a CF of 1054 (80.75 with no losses) for J1834.9–0846. The difference with the present results, which maps very well what is expected from the Lagrangian evolution, is apparent: the pure thin-shell approximation always leads to an overestimation of the CF (by a factor 1.7 for the Crab and 4.3 for J1834.9–0846), as we extensively discussed in Paper II. As a general result, we have shown here that the CF is very sensitive to the details of the model used, so that an accurate treatment is a must,

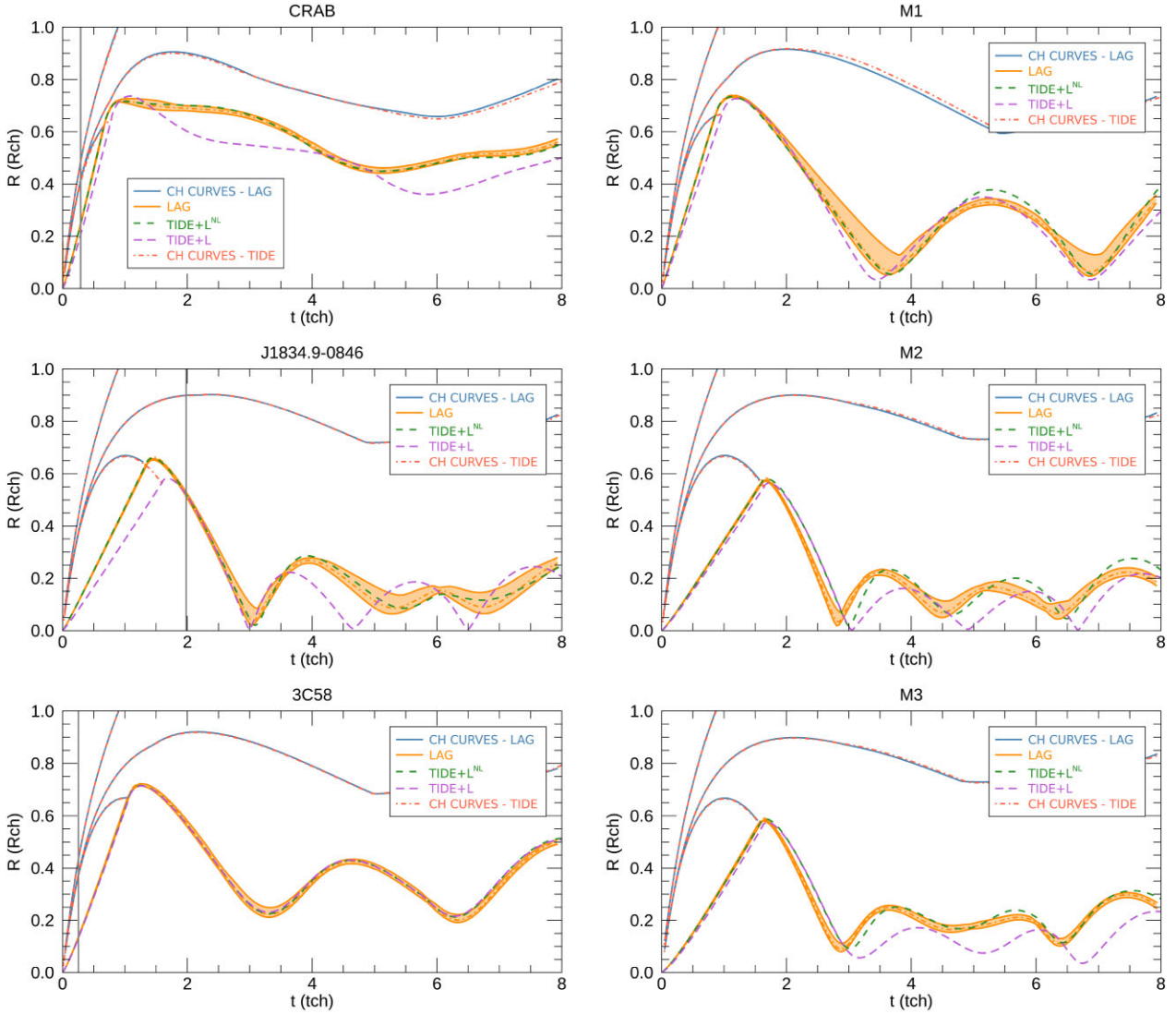


Figure 2. Time evolution of the PWN radius and SNR characteristic curves (named CH CURVES in the legend) for different models, with quantities expressed in characteristic units [see equation (12)]. The SNR FS, CD, and RS are extracted from the Lagrangian simulations (light blue solid lines) and compared with the same curves as obtained with TIDE + L, in the non-radiative case (dot-dashed light red lines). From the outside (left-upper corner of each plot) to the inside, we can recognize the FS going out of the plot at $\sim 1t_{\text{ch}}$, the CD and the RS, impinging on the PWN. It is apparent that, except for minor deviations in the CD, the characteristic curves are perfectly coincident from the Lagrangian code and TIDE + L. From the Lagrangian simulations we do not only plot R_{pwn} , but also the outer edge of the swept-up mass shell (all in orange), to illustrate how much the shell deviates from the thin-shell approximation. The shell barycenter is also shown as an orange dot-dashed line lying in between the PWN radius and the outer edge of the shell. Green and violet dashed lines show the evolution of R_{pwn} as obtained, respectively, with TIDE + L^{NL} and TIDE + L. The presence of early synchrotron losses results in a deviation of the PWN radius already in the free-expansion phase and in the delayed beginning of the reverberation phase (particularly evident in J1834.9–0846). Each source is modelled following the parameters given in Table 1. Real sources are shown on the left column, synthetic ones on the column on the right. Finally, for real sources, we show the positioning of their present age on the evolutionary plot by means of a vertical solid grey line.

to avoid inaccurate results. Despite the corrections introduced by the better dynamical treatment, which by themselves represent an important progress in PWNe simulations, providing the tools to pass through reverberation in a relatively safe approach, the compression factors at the border of the known population of nebulae remain extreme. PWNe can reduce themselves in size by more than two orders of magnitude, at least in the 1D representation, while in 3D this behaviour might change due to the complex interaction of the PWN boundary with the SNR and the onset of border instabilities and mixing. The main spectral evolution effects of such changes are discussed next, with further details left for future discussion elsewhere.

5 SPECTRAL EVOLUTION

Let us discuss here the spectral evolution of our three real sources, as computed by TIDE + L. For all sources we compute the total spectrum at their present age, to be compared with available spectral data. We then perform a multi-epoch analysis to investigate how, especially in relevant moments during the reverberation phase, the spectral properties of the sources may vary.

As we already mentioned before, realizing a good fit to the data, i.e. one that takes into account the modification TIDE + L brings over previous versions regarding the Lagrangian implementation of the evolution, is beyond the scope of this work. We leave this to future

Table 2. CFs and related quantities as obtained, for the different sources considered in this work, with the pure Lagrangian approach and the two versions of TIDE + L (the non-radiative one indicated with TIDE + L^{NL}).

System	Code	$t_{\text{beg, rev}}$ [t_{ch}]	$R_{\text{pwn}}^{\text{MIN}}$ [R_{ch}]	$R_{\text{pwn}}^{\text{MAX}}$ [R_{ch}]	CF
CRAB	LAG	0.6852	0.4414	0.6385	1.447
	LAG(bary)	–	0.4499	0.7205	1.602
	TIDE + L ^{NL}	0.6408	0.4489	0.7153	1.593
	TIDE + L	0.6633	0.3595	0.7357	2.047
J1834.9–0846	LAG	1.323	0.0198	0.6196	31.29
	LAG(bary)	–	0.0368	0.6629	18.01
	TIDE + L ^{NL}	1.315	0.0205	0.6592	32.16
	TIDE + L	1.339	0.0024	0.5819	242.5
3C58	LAG	1.082	0.1915	0.6837	3.570
	LAG(bary)	–	0.2009	0.7232	3.600
	TIDE + L ^{NL}	1.054	0.2128	0.7154	3.362
	TIDE + L	1.062	0.2161	0.7143	3.305
M1	LAG	0.8806	0.0471	0.6746	14.32
	LAG(bary)	–	0.0638	0.7394	11.59
	TIDE + L ^{NL}	0.8466	0.0544	0.7339	13.49
	TIDE + L	0.8574	0.0335	0.7260	21.67
M2	LAG	1.604	0.0181	0.5618	31.04
	LAG(bary)	–	0.0332	0.5829	17.56
	TIDE + L ^{NL}	1.592	0.0182	0.5799	31.86
	TIDE + L	1.606	0.0023	0.5645	245.4
M3	LAG	1.5317	0.0801	0.5483	6.845
	LAG(bary)	–	0.0900	0.5905	6.561
	TIDE + L ^{NL}	1.5393	0.0918	0.5877	6.402
	TIDE + L	1.5838	0.0355	0.5745	16.18

research. Rather, we here give preference to the comparison with the results obtained with the previous – thin-shell – version of TIDE. It is in this comparison where we shall be able to learn the impact of incorporating the new treatment of the dynamics via the Lagrangian modules that represent the passage through reverberation. For this reason, we have not modified the spectral parameters determined in the reference works for the selected sources, not to introduce differences in the spectral properties besides those naturally arising from the different evolution. For a discussion of the fitting procedure and best-fitting values we refer to Paper I (and references therein) and Martin & Torres (2022).

For the Crab and 3C58, even without optimizing the parameters, the fits appear to be still good, given that both are still in free-expansion, and the thin-shell approximation provides a good description of their evolution, such that the difference between TIDE + L and previous versions, are small. Different is the case of J1834.9–0846, which is the only source that is already in reverberation, according to our models (see Torres 2017).

Here, we briefly recall how we compute the spectral model through the source evolution. The energy input, namely the spin-down luminosity of the pulsar $L(t)$, changes with time following:

$$L(t) = L_0 \left(1 + \frac{t}{\tau_0} \right)^{-\beta}, \quad (15)$$

where $\beta = (n + 1)/(n - 1)$ is the fading index. The particle content of the PWN at each time is given by the balance of injection, energy losses, and adiabatic losses (or gains, in compression phases), plus possibly the escape of particles from the source through diffusion. These are taken into account solving the following diffusion-loss equation, as used in Martin et al. (2012) and Torres et al. (2014):

$$\frac{\partial N(\gamma, t)}{\partial t} = -\frac{\partial}{\partial \gamma} [\dot{\gamma}(\gamma, t)N(\gamma, t)] - \frac{N(\gamma, t)}{\tau(\gamma, t)} + Q_{\text{inj}}(\gamma, t), \quad (16)$$

where $N(\gamma, t)$ is the particle number, $\gamma = E/(m_e c^2)$ the particle normalized energy (with m_e the electron mass, considering a fully leptonic pulsar wind), $\dot{\gamma}(\gamma, t)$ the loss (gain) term (containing synchrotron, IC, and bremsstrahlung radiative losses plus adiabatic losses or gains), $\tau(\gamma, t)$ is the escape term (assuming Bohm diffusion), and finally, $Q_{\text{inj}}(\gamma, t)$ is the injection term. Particles are continuously injected in the PWN during the evolution considering a broken power law in energy, namely:

$$Q_{\text{inj}}(\gamma, t) = Q_0(t) \times \begin{cases} \left(\frac{\gamma}{\gamma_b} \right)^{-\alpha_l}, & \text{if } \gamma \leq \gamma_b, \\ \left(\frac{\gamma}{\gamma_b} \right)^{-\alpha_h}, & \text{if } \gamma > \gamma_b. \end{cases} \quad (17)$$

The normalization $Q_0(t)$ is obtained at each time from the requirement that a constant fraction $(1 - \eta)$ of the pulsar input goes into particles, namely:

$$(1 - \eta)L(t) = \int_{\gamma_{\text{min}}}^{\gamma_{\text{max}}} \gamma m_e c^2 Q_{\text{inj}}(\gamma, t) d\gamma. \quad (18)$$

The minimum Lorentz factor at injection is generally assumed to equal the electron energy at rest (then $\gamma_{\text{min}} = 1$), while the maximum Lorentz factor is computed following Martin et al. (2016) (either considering the synchrotron acceleration limit or the maximum energy to maintain acceleration confined into the wind termination shock, see equations 6–7 of that work). The remaining fraction of the injection energy, $\eta L(t)$, is then converted into magnetic field at each time-step. The variation of the magnetic energy ($W_B = B^2 R^3/6$) in the PWN is given by (Martin et al. 2016):

$$\frac{dW_B(t)}{dt} = \eta L(t) - \frac{W_B(t)}{R(t)} \frac{dR(t)}{dt}. \quad (19)$$

The magnetic fraction η , or better called the injection sharing, is kept constant during the evolution, and so is the instantaneous ratio of the energy that goes into particles and into magnetic field. This constant injection sharing along time should not be confused with a constant energy partition, which is in fact not realizing: since particles have dynamical and radiative losses, while the magnetic energy is affected by dynamics only, the ratio of particle energy to magnetic field energy is not constant in time, nor it is equal to $(1 - \eta)$. However, TIDE provides $B(t)$, as well as the particles energy directly, and thus the energy ratio can be computed at each time (see e.g. fig. 4 of Martin et al. 2016) for further discussion. The η parameter can be connected with the pulsar wind magnetization by: $\sigma = \eta/(1 - \eta)$.

Fig. 3 shows the results of the spectral evolution of the selected sources. In the left column, we show the spectrum at their actual age, the available spectral data for reference, and the relevance of the various terms contributing to the total spectrum. In the right column, we show the multi-epoch spectrum, computed at different times for the various sources, to highlight specific moments of their reverberation process (i.e. maximum compressions or expansions).

Let us here discuss in more detail the differences between the present approach and previous ones, with particular reference to Torres & Lin (2018, TL2018 hereafter), where the authors used TIDE (v2.2, following the nomenclature of Paper I) i.e. standard thin-shell model, as normal in literature, without Lagrangian treatment of reverberation and with a different description of the shell dynamics than the one introduced in Paper I.

For the three real sources:

(i) **Crab nebula** (upper row).

A comparison with results obtained with the standard one-zone, pure thin-shell, version of the code TIDE can be made looking at fig. 1 in TL2018. The agreement is excellent at t_{age} , as we discussed

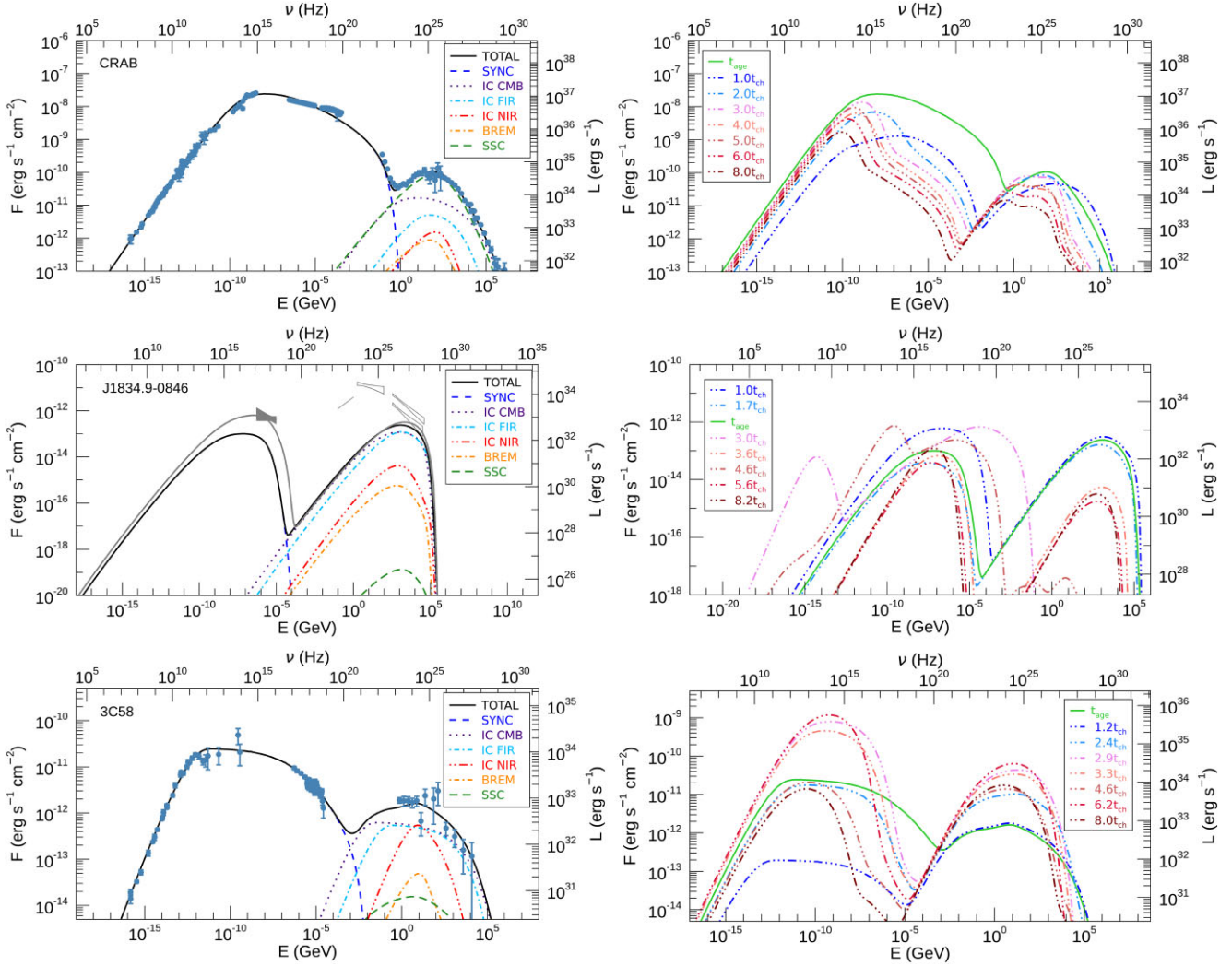


Figure 3. Spectra at t_{age} (panels on the left-hand) and at multiple epochs (panels on the right-hand) for the three real benchmark sources considered in this work (from top to bottom: Crab, J1834.9–0846 and 3C58). The single spectral components contributing to the total emission are plotted separately at t_{age} , where SYNC stands for synchrotron emission from both particles families, SSC for self-synchrotron Compton, and the major photon fields for IC considered are: the cosmic microwave background (CMB), far- and near-infrared photons (FIR and NIR). In the multi-epoch spectra, the spectrum at t_{age} is plotted again for an easier comparison with other ages, in green colour and solid line. The different epochs considered, chosen in order to map salient moments of the evolution of each source (e.g. maximum or minimum of R_{pwn}), are listed in the plot legend, and relative spectra are drawn with the same colour-code. For the Crab Nebula data are taken from Baldwin (1971); Macías-Pérez et al. (2010; radio), Bühler & Blandford (2014; from IR to very high energies), and LHAASO Collaboration (2021) for extremely high energy data. For J1834.9–0846, we consider the same set of data as in Torres (2017): out of them only X-ray data are certainly referring to J1834.9–0846 and coming from the magnetar nebula (Younes et al. 2016), while in the GeV and TeV bands we can only report upper limits. As mentioned in the main text, since this source at $t = t_{\text{age}}$ already entered its reverberation phase, its spectral properties are particularly sensitive to the variation of the PWN radius, pressure, and magnetic field. This is the reason why in order to get a reasonable fit of the data we need to modify substantially the initial magnetization from the original value assumed in the reference work using the pure thin-shell TIDE ($\eta = 0.045$ versus the new value $\eta = 0.15$, much larger). The total spectrum corresponding to the new best-fitting value is shown as a grey solid line. Data for 3C58 come from Green (1986); Morsi & Reich (1987); Salter et al. (1989); Torii et al. (2000; radio and IR); Green (1994); Slane et al. (2008; X-rays), Abdo et al. (2013); Ackermann et al. (2013); Li et al. (2018; gamma-rays), and Aleksić et al. (2014; for the very high energies).

previously, and the same holds at $1 t_{\text{ch}}$, the time at which the PWN reaches its maximum expansion (black dashed line corresponding to t_1 in fig. 1 of TL2018). We can easily recognize in the spectrum the effects of the adiabatic expansion and magnetic field decrease in the lower synchrotron spectrum, which also cause less emission in the IC range, being the major contribution coming from self-synchrotron Compton processes (SSC).

On the contrary, from the onset of reverberation, the difference with the results obtained within the pure thin-shell approximation starts to increase. At the maximum compression ($t_3 = 6841.7$ yr in TL2018,

blue dashed line) the pure thin-shell modelling produces a larger compression, and thus on one side high energy gains (reflecting in the extra-luminosity at radio frequencies in TL2018), while on the other a huge increase of the nebular magnetic field, modifying substantially the high-energy synchrotron spectrum. Depending on how much the magnetic field is increased, most of the particles might be burn off, also causing a strong variation in the IC spectrum, due to the lack of electrons for the scattering events. On the contrary, when we release the pure thin-shell approach and model the evolution with TIDE + L, the maximum compression happens much later (at

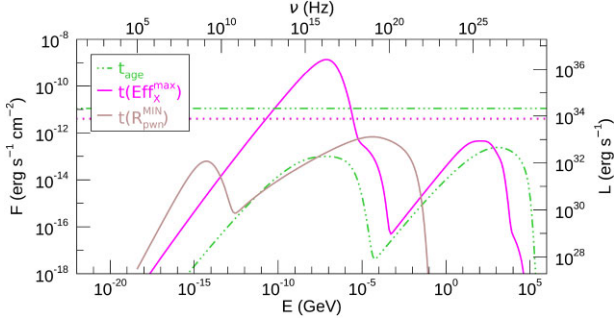


Figure 4. Plot of the spectrum of J1834.9–0846 at three different ages: t_{age} (in dash-dotted green), $t(\text{Eff}_X^{\text{max}}) \simeq 2.9t_{\text{ch}}$, the time at which the efficiency in the X-ray band is maximum (solid magenta line), and $t(R_{\text{pwn}}^{\text{MIN}})$, the maximum compression (at $3.0t_{\text{ch}}$ – solid brown line), to be matched with the same coloured line of Fig. 3. Comparing with the spectral evolution of the same source discusses in TL2018, where TIDE v2.2 was used, we can notice that, despite the difference in times and intensity of the compression, the spectrum at maximum efficiency is very similar (see Fig. 1, second row, orange dashed line in that paper). The horizontal lines show the value of the spin-down luminosity of the pulsar at the three ages considered (with the same colour code). The brown line, referring to $t(R_{\text{pwn}}^{\text{MIN}})$, is hidden in the plot by magenta line.

$t \sim 6t_{\text{ch}} \sim 20\text{kyr}$, corresponding to the medium red line), and we see a different scenario: the luminosity at low energies is slightly diminished while the X-ray component is almost depleted (due to ageing), and the same reflects on the IC component. No effects connected to adiabatic gains or increase of the magnetic field can be easily identified in the full model at the spectral level. From t_{age} on, the magnetic field in the system always remains below $100 \mu\text{G}$; the weak compressions do not cause any important field amplification.

(ii) **J1834.9–0846** (middle row).

In this case, being $t_{\text{age}} > t_{\text{beg, rev}}$, the variation introduced by the coupled Lagrangian and thin-shell evolution of TIDE + L makes the set of spectral parameters recovered from previous works (TL2018 and Paper I) not good for a fit to the spectral data anymore. To improve the representation, we need to substantially increase the magnetic fraction (from 0.045 to 0.15). Of course, this also means that the comparison with previous results (same fig. 1 in TL2018) is no longer straightforward, and not very good even at t_{age} , although the overall properties of the spectrum appear consistent. J1834.9–0846 is one of the two compressible systems we considered and the one with the lowest spin-down luminosity at the maximum compression (at $t = 3t_{\text{ch}} \simeq 12073 \text{ yr}$). If we look at the spectrum at the compression peak, we recognize the effect of the adiabatic compression in the increase of the low-energy emission (the bump peaking at $\sim 10^{10} \text{ Hz}$) combined with the synchrotron cooling of the higher energy synchrotron component, visible with a spectral steepening before the peak. If we compare with the same time from TL2018 (where the maximum of the compression happens at $t_3 = 8373.9 \text{ yr}$), we note the same effect: the extra compression induced by the pure thin-shell approximation reflects in a massive synchrotron cooling at all energies, even in the radio band (no low-energy bump survives). However, in this case, in both approaches the compression is strong enough to increase the magnetic field to the point that most of the particles in the nebula are cooled by synchrotron losses.

In Fig. 4, we plot again the spectrum of J1834.9–0846 to show the detail of the maximum efficiency in X-ray emission (namely $L_X(t)/L(t)$), which the system reaches before the maximum compression (namely at a radius of $\sim 0.03R_{\text{ch}}$ and at $t \simeq 2.9t_{\text{ch}}$). Using the same notation as in TL2018, that time is named $t(\text{Eff}_X^{\text{max}})$ and corresponds

to a spectral peak at $\nu \simeq 2 \times 10^{16} \text{ Hz}$. The spin-down luminosity injected by the pulsar at that time is $L(t(\text{Eff}_X^{\text{max}})) = 7.8 \times 10^{33} \text{ erg s}^{-1}$, while the X-ray luminosity in the 0.5–5 keV (Chandra) band is $L_X(t(\text{Eff}_X^{\text{max}})) = 1.6 \times 10^{36} \text{ erg s}^{-1}$, namely a factor ~ 200 larger than the injected power, meaning that the system will actually be going through a strong super-efficient phase. In this source, even if the magnetic field soon after reverberation is low (see Table 1), during the peak of the compression phases it is amplified to very large values, being $B \propto R_{\text{pwn}}^{-2}$. In the absence of radiative losses, the first compression leads to a magnetic field of 3 mG, while with radiative losses it reaches 100 mG.

(iii) **3C58** (bottom row).

At $t_{\text{beg, rev}} \sim 1.2t_{\text{ch}}$ the spectrum shows the effects of the adiabatic expansion: the synchrotron spectrum lowers due to a combined effect of adiabatic cooling of the particles and decrease of the magnetic field. In this case, differently from what we have earlier discussed for the Crab nebula, being the dominant emission at very high energies the one coming from the scattering between nebular leptons with background photons, instead of the SSC, no variation is observed in the IC spectrum. At the first compression (happening at $t \sim 3.3t_{\text{ch}}$), the spectrum shows an evident modification in morphology: the synchrotron spectrum peaks at higher frequencies due to adiabatic gains. This system has been characterized a very low magnetic field, with little dynamical consequences. This is shown by the coincidence between the non-radiative and radiative evolution (see Fig. 2). It also has a rather modest compression during reverberation, and the increase in magnetic field is not enough to effectively modify the IC emission yield, which instead follows directly the variations of the synchrotron component. For 3C58 we do not have multi-epoch examples to compare with, but only the spectrum at t_{age} as presented in Martin & Torres (2022), with which we find a perfect agreement for the reasons just described. This nebula is weakly magnetized (Slane et al. 2008), and as expected the magnetic field remains low for the entire evolution, with an average up to an age of $8t_{\text{ch}}$ of $6 \mu\text{G}$ ($3 \mu\text{G}$ without losses).

Finally, let us comment upon the possible appearance of super luminous phases in the evolution of our three sources. These can be easily identified in Fig. 5, especially comparing the integrated luminosity in a given energy band with the power injected by the pulsar at the same age (panel on the right-hand). As discussed previously, only J1834.9–0846 shows an X-ray luminosity that exceeds the injection luminosity in some phases, due to the extreme compression the system undergoes in reverberation and the consequent dynamical modifications. Interestingly, Crab and 3C58, which on the contrary do not show important modifications of their high energy emission due to the small compression, both show the presence of phases with an augmented optical emission (at 1 eV).

We believe this is an interesting feature to be further investigated in a dedicated study, extending the analysis of the spectral modifications to a larger population to identify possible observational features for future instruments working in the UV and optical- B bands.

6 CONCLUSIONS

Despite the fact that middle-aged PWNe constitute the vast majority of all PWN–SNR systems, and the fact that perhaps they form one of the main contributor to the diffuse Galactic gamma-ray background, little is known of their evolutionary properties. This is in sharp contrast with younger systems for which, thanks to a plethora of approaches, we have now a well-established canonical picture.

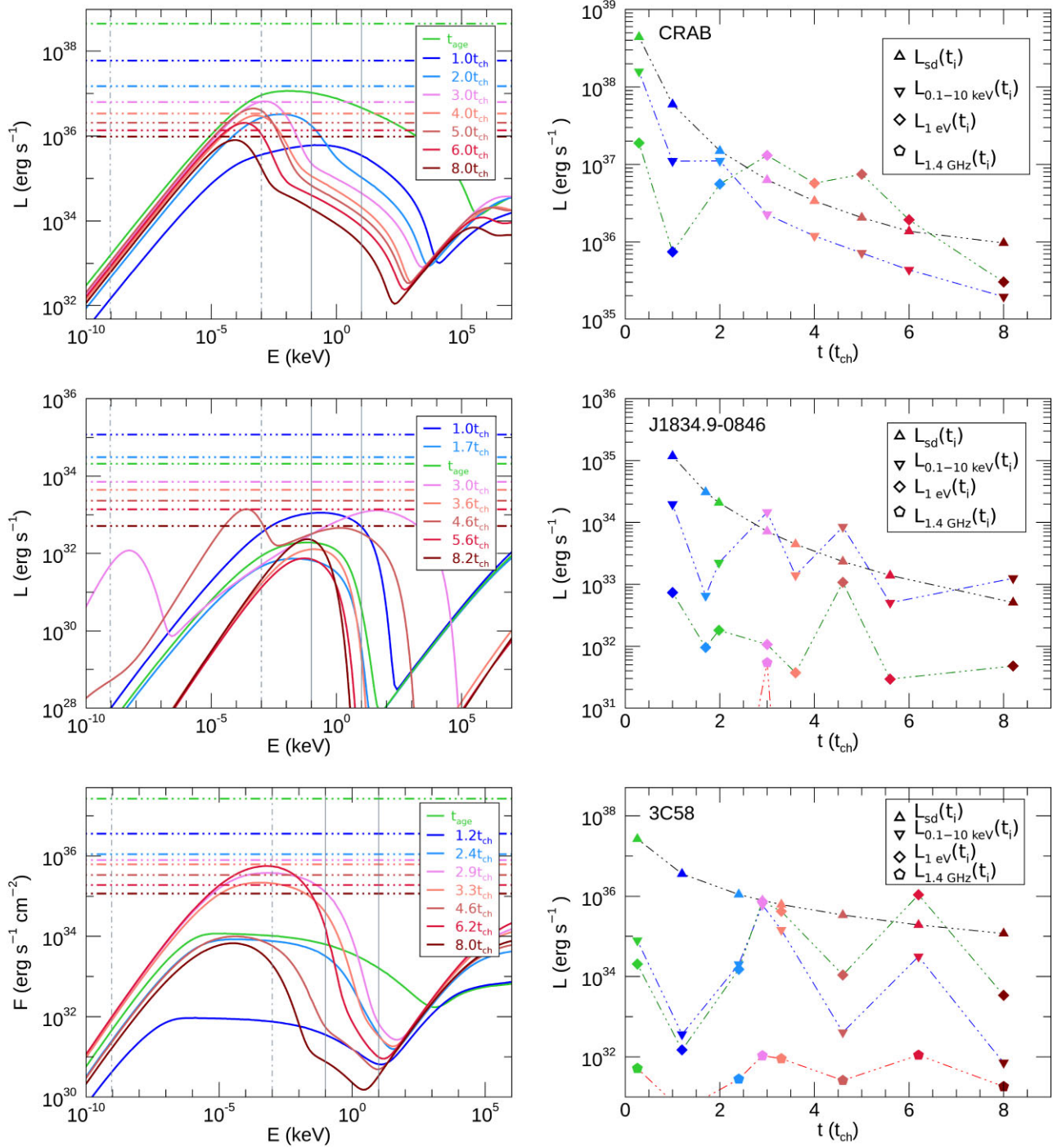


Figure 5. Left-hand panels: Synchrotron spectra of the Crab (upper row), J1834.9–0846 (middle row), and 3C58 (bottom row), at the same ages considered in Fig. 3 (using the same colour coding). In each plot the horizontal dot-dashed lines show the value of the spin-down luminosity of the pulsar at the same age at which the spectral luminosity has been obtained (same colours for the same age). The vertical grey lines identify some energy bands of interest, from left to right: the radio frequency of 1.4 GHz (dot-dashed), the optical at 1 eV (dot-dashed), and the 0.1–10 keV X-ray band (within the solid lines). These are the references to build up the plots of the panels on the right. Right-hand panels: comparison of the integrated PWN luminosity (in radio, optical, and X-rays) with the injection luminosity at the ages considered for the left plots. The colours refer to the same age, while different symbols refer to different observational bands. Up triangles represent the pulsar spin-down at a certain age, and to facilitate reading the symbols are joined with a dot-dashed black line. Down triangles (joined with a dot-dashed blue line) stand for the PWN integrated X-ray luminosity, diamonds for the optical one (joined with a green dot-dashed line), pentagons for the radio luminosity (joined with a red dot-dashed line). Normally the spin-down luminosity is larger than the PWN luminosity in any band, and up triangles are above the other symbols. On the contrary, the system is facing a super-efficient phase when the symbol relative to a certain energy band stays above the up triangle (i.e. $L_i > L_{\text{sd}}$). As it is apparent, in our case this is only happening in J1834.9–0846 at X-rays, for the 0.1–10 keV band. Interestingly, the Crab nebula shows a persistent super-efficient phase in the optical, while 3C58 has a super-efficient optical phase around $6t_{\text{ch}}$. In most cases, the luminosity in the radio band is much lower than other bands and it remains excluded from the plots (this is e.g. the case for the Crab, where the radio luminosity is always below 2×10^{33} erg s $^{-1}$), too low to extend the plot and maintain the visibility of the other points).

Barring a few limited numerical multidimensional simulations of selected systems, currently the most popular approach to investigate the late time evolution of PWNe is within the so called thin-shell approximation. However, a correct description of the dynamics of the compression that a PWN experiences during reverberation requires a correct description of both the shell properties and that of the SNR surrounding it. In Paper II, in fact, we showed how sensitive the PWN evolution is to a rough representation of both *of the latter*. This not only affects the shape and size of the PWN, but also the future evolution of the relativistic pair population, and thus the spectral properties of such systems, together with radiative losses which feedback on the dynamical bouncing properties.

In this work, we have further extended our study of the properties of post-reverberation systems, introducing a new numerical technique for the dynamical-spectral evolution of a PWN interacting with its SNR through and beyond the reverberation phase. Using this hybrid approach, that combines Lagrangian techniques for the dynamical evolution of the SNR, with standard thin-shell evolution for the PWN, together with time dependent radiative model for the spectral properties of the PWN, we are able to model self-consistently the entire evolution, allowing PWNe simulations to extend up to the typical age range where most of the TeV population is found.

The code presented here is based on a modified version of TIDE, and its performances, with respect to a more sophisticated full-Lagrangian approach, has been cross-checked for a variety of input conditions that cover the parameters space of the PWN–SNR population. The new approach correctly recovers the full evolution both during free-expansion and later reverberation.

As it might be expected, we found that radiative losses introduce complexity to the description of the PWN response to the SNR in reverberation, especially for highly-compressible systems: sources at different position in the PWN–SNR parameters space can show similar compression factors, due to the combination of their injection parameters (L_0 , τ_0), the SNR characteristic age (t_{ch} and how it compares with τ_0) and the relevance of synchrotron losses (the magnetization parameter η). A general treatment of the effect of radiation losses and its possible parametrization require a much more detailed investigation of the parameters space, that goes beyond the scope of the present paper and that we leave for future investigations.

The importance of the results presented here is immediately clear when we look at the multiband spectral evolution of the sources, in the later evolutionary phases. While in the free-expansion results are unchanged, substantial differences arises at later times, especially for the low compressive systems. The difference is less critical for highly-compressible cases. For these cases, we confirm the existence of the so called super-efficient phases.

The minor compression reflects in a less structured spectrum of the sources at the different ages, since most of the bumps and variations to the spectral slopes observed with previous models arise from the excessive compression, causing increased synchrotron cooling and appearance of bumps connected with the adiabatic gains in the strong compression phases. We plan to use our new model to produce reliable PWNe population studies, extending the concept presented in Fiori et al. (2021) to multiwavelengths analysis, as well as gather estimates of the extent and the number of super-efficient PWNe existing in the Galaxy, enhancing the work by Torres, Lin & Coti Zelati (2019).

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DATA AVAILABILITY

The data for the models underlying this article will be shared on reasonable request to the corresponding authors.

REFERENCES

- Abdo A. A. et al., 2013, *ApJS*, 208, 17
 Ackermann M. et al., 2013, *ApJS*, 209, 34
 Aleksić J. et al., 2014, *A&A*, 567, L8
 Baldwin J. E., 1971, in Davies R. D., Graham-Smith Feds, Vol. 46, The Crab Nebula, Dordrecht. p. 22
 Bandiera R., Bucciantini N., Martín J., Olmi B., Torres D. F., 2020, *MNRAS*, 499, 2051
 Bandiera R., Bucciantini N., Martín J., Olmi B., Torres D. F., 2021, *MNRAS*, 508, 3194
 Bandiera R., Bucciantini N., Martín J., Olmi B., Torres D. F., 2023, *MNRAS*, 520, 2451
 Barkov M. V., Lyutikov M., Khangulyan D., 2019, *MNRAS*, 484, 4760
 Bogovalov S. V., Chechetkin V. M., Koldoba A. V., Ustyugova G. V., 2005, *MNRAS*, 358, 705
 Bucciantini N., Blondin J. M., Del Zanna L., Amato E., 2003, *A&A*, 405, 617
 Bucciantini N., Amato E., Bandiera R., Blondin J. M., Del Zanna L., 2004, *A&A*, 423, 253
 Bucciantini N., Arons J., Amato E., 2011, *MNRAS*, 410, 381
 Bühler R., Blandford R., 2014, *Rep. Prog. Phys.*, 77, 066901
 Camilo F., Gaensler B. M., Gotthelf E. V., Halpern J. P., Manchester R. N., 2004, *ApJ*, 616, 1118
 Chevalier R. A., 1982, *ApJ*, 258, 790
 Ferreira S. E. S., de Jager O. C., 2008, *A&A*, 478, 17
 Del Zanna L., Amato E., Bucciantini N., 2004, *A&A*, 421, 1063
 Del Zanna L., Volpi D., Amato E., Bucciantini N., 2006, *A&A*, 453, 621
 Emmering R. T., Chevalier R. A., 1987, *ApJ*, 321, 334
 Fiori M., Olmi B., Amato E., Bandiera R., N. Bucciantini, Zampieri L., Burtovoi A., 2022, *MNRAS*, 511, 1439
 Gaensler B. M., Arons J., Kaspi V. M., Pivovarov M. J., Kawai N., Tamura K., 2002, *ApJ*, 569, 878
 Gaensler B. M., Slane P. O., 2006, *ARA&A*, 44, 17
 Gelfand J. D., Slane P. O., Zhang W., 2009, *ApJ*, 703, 2051
 Green D. A., 1986, *MNRAS*, 218, 533
 Green D. A., 1994, *ApJS*, 90, 817
 Kennel C. F., Coroniti F. V., 1984a, *ApJ*, 283, 694
 Kennel C. F., Coroniti F. V., 1984b, *ApJ*, 283, 710
 Komisarov S. S., Lyubarsky Y. E., 2004, *MNRAS*, 349, 779
 LHAASO Collaboration, 2021, *Science*, 373, 425
 Li J., Torres D. F., Lin T. T., Grondin M.-H., Kerr M., Lemoine-Goumard M., de Oña Wilhelmi E., 2018, *ApJ*, 858, 84
 Lu F. J., Wang Q. D., Aschenbach B., Durouchoux P., Song L. M., 2002, *ApJ*, 568, L49
 Macías-Pérez J. F., Mayet F., Aumont J., Désert F. X., 2010, *ApJ*, 711, 417
 Martin J., Torres D. F., Rea N., 2012, *MNRAS*, 427, 415
 Martin J., Torres D. F., Pedalletti G., 2016, *MNRAS*, 459, 3868
 Martin J., Torres D. F., 2022, *JHEAP*, 36, 128
 Mezzacappa A., Bruenn S. W., 1993, *ApJ*, 405, 669
 Michel F. C., 1973, *ApJ*, 180, L133
 Morsi H. W., Reich W., 1987, *A&AS*, 69, 533

- Olmi B., Bucciantini N., 2019, *MNRAS*, 484, 5755
- Olmi B., Bucciantini N., 2023, *PASA*, 40, e007
- Olmi B., Del Zanna L., Amato E., Bucciantini N., Mignone A., 2016, *J. Plasma Phys.*, 82, 635820601
- Pacini F., Salvati M., 1973, *ApJ*, 186, 249
- Porth O., Komissarov S. S., Keppens R., 2014, *MNRAS*, 443, 547
- Rees M. J., Gunn J. E., 1974, *MNRAS*, 167, 1
- Reynolds S. P., Chevalier R. A., 1984, *ApJ*, 278, 630
- Romani R. W., Ng C. Y., 2003, *ApJ*, 585, L41
- Romani R. W., Ng C. Y., Dodson R., Brisken W., 2005, *ApJ*, 631, 480
- Salter C. J., Reynolds S. P., Hogg D. E., Payne J. M., Rhodes P. J., 1989, *ApJ*, 338, 171
- Schulz W. D., 1964, *J. Math. Phys.*, 5, 133
- Slane P., 2017, in Alsabti A. W., Murdin P. eds, *Handbook of Supernovae*, Springer Int. Pub., New York. p. 2159,
- Slane P., Helfand D. J., van der Swaluw E., Murray S. S., 2004, *ApJ*, 616, 403
- Slane P., Helfand D. J., Reynolds S. P., Gaensler B. M., Lemiére A., Wang Z., 2008, *ApJ*, 676, L33
- Timokhin A. N., 2006, *MNRAS*, 368, 1055
- Torii K., Slane P. O., Kinugasa K., Hashimoto-dani K., Tsunemi H., 2000, *PASJ*, 52, 875
- Torres D. F., Cillis A., Martín J., de Oña Wilhelmi E., 2014, *JHEAP*, 1, 31
- Torres D. F., 2017, *ApJ*, 835, 54
- Torres D. F., Lin T., 2018, *ApJ*, 864, L2
- Torres D. F., Lin T., Coti Zelati F., 2019, *MNRAS*, 486, 1019
- Truelove J. K., McKee C. F., 1999, *ApJS*, 120, 299
- van der Swaluw E., Achterberg A., Gallant Y. A., Tóth G., 2001, *A&A*, 380, 309
- van der Swaluw E., Downes T. P., Keegan R., 2004, *A&A*, 420, 937
- Weisskopf M. C. et al., 2000, *ApJ*, 536, L81
- Younes G. et al., 2016, *ApJ*, 824, 138

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