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Observations of the very young Type Ia Supernova 2019np with early-excess emission

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ABSTRACT

Early-time radiative signals from Type Ia supernovae (SNe Ia) can provide important constraints on the explosion mechanism and the progenitor system. We present observations and analysis of SN 2019np, a nearby SN Ia discovered within 1–2 days after the explosion. Follow-up observations were conducted in optical, ultraviolet, and near-infrared bands, covering the phases from ~-16.7 d to $\sim+$ 367.8 d relative to its B-band peak luminosity. The photometric and spectral evolutions of SN 2019np resemble the average behaviour of normal SNe Ia. The absolute B-band peak magnitude and the post-peak decline rate are $M_{\rm max}(B) = -19.52 \pm 0.47$ mag and $\Delta m_{15}(B) = 1.04 \pm 0.04$ mag, respectively. No Hydrogen line has been detected in the nebular-phase spectra of SN 2019np. Assuming that the ^{56}Ni powering the light curve is centrally located, we find that the bolometric light curve of SN 2019np shows a flux excess up to 5.0 per cent in the early phase compared to the radiative diffusion model. Such an extra radiation perhaps suggests the presence of an additional energy source beyond the radioactive decay of central nickel. Comparing the observed colour evolution with that predicted by different models, such as interactions of SN ejecta with circumstellar matter (CSM)/companion star, a double-detonation explosion from a sub-Chandrasekhar mass white dwarf (WD) and surface ^{56}Ni mixing, we propose that the nickel mixing is more favoured for SN 2019np.

Key words: supernovae: general – supernovae: individual: (SN 2019np).

1 INTRODUCTION

Since the late twentieth century, the studies of Type Ia supernovae (SNe Ia) have led to the discovery of the accelerating universe (Riess et al. 1998; Perlmutter et al. 1999; Riess et al. 2016, 2018, 2021). Nowadays, SNe Ia are widely used as standardizable candles in observational cosmology (Riess, Press & Kirshner 1996; Guy et al. 2005; Wang et al. 2005; Howell et al. 2006; Howell 2011; Burns et al. 2018). However, the exact formation channel and explosion physics of SNe Ia are still unknown (Wang & Han 2012; Wang et al. 2013a; Maoz, Mannucci & Nelemans 2014; Jha, Maguire & Sullivan 2019).

Three popular scenarios for the explosion of SNe Ia are: (1) The 'single-degenerate' (SD) channel, in which a carbon-oxygen (CO) WD accretes matter from a non-degenerate companion such as a main-sequence star, a giant, or a He star. A thermonuclear explosion process is expected to occur when the progenitor CO WD reaches the critical Chandrasekhar mass of $M_{\rm Ch} \sim 1.4\,{\rm M}_{\odot}$ (Whelan & Iben 1973; Nomoto, Iwamoto & Kishimoto 1997; Podsiadlowski et al. 2008; Wang et al. 2009a); (2) The 'double-degenerate' (DD) channel, in which the explosions of SNe Ia are triggered by the compressional heat generated by the dynamical merging process of two CO WDs (Iben & Tutukov 1984; Webbink 1984; Benz et al. 1990; Pakmor et al. 2012); (3) The direct head-on collision of two WDs in triple systems (Kushnir et al. 2013); however, the rate of head-on collisions of WDs in triple systems has been suggested to be too low to account for the majority of SNe Ia explosions (Toonen, Perets & Hamers 2018). In the SD channel, an extended CSM profile is expected to form around the progenitor system during its evolution towards a SN Ia explosion. Although the signature of CSM has been revealed for some SNe Ia via high-resolution spectroscopy (Hamuy et al. 2003; Aldering et al. 2006; Patat et al. 2007; Sternberg et al. 2011; Dilday et al. 2012; Maguire et al. 2013; Silverman et al. 2013), the lack of narrow H emission lines in the late-time spectra still challenges the SD model (Mattila et al. 2005; Leonard 2007; Shappee et al. 2013; Maguire et al. 2016; Tucker et al. 2020). Additionally, searches for pre-explosion or surviving companions of the progenitor through deep images (Li et al.

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2011; González Hernández et al. 2012; Schaefer & Pagnotta 2012) have firmly excluded the presence of giant and subgiant companions.

On the other hand, detection of the first light after a SN explosion can also provide important constraints on progenitors of SNe Ia (Kasen 2010). In a SD progenitor system, the SN ejecta could run into the non-degenerate companion or surrounding CSM. Then an excess of ultraviolet (UV)/optical emission could be generated and become detectable within the first few days of the interaction between SN ejecta and a companion star or within a few hours for the interaction between SN ejecta and ambient CSM. In the former case, the flux excess can last for a few days, depending on the size of the companion, the progenitor-companion separation, the ejecta velocity, and the viewing angle. During the *Kepler* mission (Haas et al. 2010), a few SNe were detected in the *Kepler* fields, including two spectroscopically confirmed SNe Ia and three possible ones. There is no signature of ejecta interaction with a stellar companion for the latter three (KSN 2012a, KSN 2011b, and KSN 2011c) (Olling et al. 2015) and SN 2018agk (Wang et al. 2021). Blue excess flux was detected in the early-time light curve of SN 2018oh, however, the explanations are not converged (Dimitriadis et al. 2019; Li et al. 2019; Shappee et al. 2019). For example, both the deep mixing carbon feature and non-detection of hydrogen in the nebular-phase spectrum of SN 2018oh do not favour a SD scenario for its origin (Li et al. 2019; Tucker, Shappee & Wisniewski 2019). A similar 'blue bump' feature was also detected in the nearby SN 2017cbv (Hosseinzadeh et al. 2017), but analysis of the nebular-phase spectroscopy argues against a non-degenerate H- or He-rich companion as its progenitor (Sand et al. 2018). For the latest case, however, the extra emission in the early time usually lasts for a few hours, depending on the density and distance of the CSM. In a recent study, SN 2020hvf is reported to have such an early-time excess emission (Jiang et al. 2021).

Besides interaction with a companion star or CSM disc, the early excess emission may also be produced by other mechanisms (Levanon & Soker 2019), e.g. radioactive decay of ^{56}Ni near the surface of the exploding WD (Piro & Nakar 2013; Magee et al. 2020). Mixing of ^{56}Ni into the outer layers would result in bluer and steeper rising light curves (Piro & Morozova 2016) and this mechanism works for both SD and DD progenitors. Magee et al. (2020) shows that the observed diversity in light curves of SNe Ia can be reproduced by varying the distribution of ^{56}Ni. Another explosion model that can produce such ^{56}Ni distribution is the double-detonation explosion of a sub-Chandrasekhar mass WD, where the initial detonation of the outer-layer He would produce a certain amount of nickel near the surface and send a shockwave to the C/O core to trigger the second detonation (Woosley & Weaver 1994; Wang, Justham & Han 2013b; Noebauer et al. 2017), thus leading to a wide range of absolute magnitudes and colours. However, in order to match the observed peak brightness of SNe Ia, this model requires that the number of iron-group elements (IGEs) synthesized during He shell detonation is very small. The radioactive decay of the surface material would produce photons that diffuse out of the ejecta rapidly, which results in a flux excess in the first few days after explosion.

The detection of SN 2019np within 1-2 days after its explosion provides us a rare opportunity to examine the first-light evolution of a SN Ia. In this paper, we present extensive follow-up observations of SN 2019np within the UV to NIR wavelength ranges. We also analyse its photometric and spectroscopic behaviours and compare the observational properties of SN 2019np with those of other well-studied SNe Ia. Observations and data reduction are outlined in Section 2, the light/colour curves are described in Section 3, and the spectroscopic behaviour of SN 2019np and its temporal evolution are presented in Section 4. We discuss the properties of SN 2019np and its explosion parameters in Section 5. The conclusions are given in Section 6.

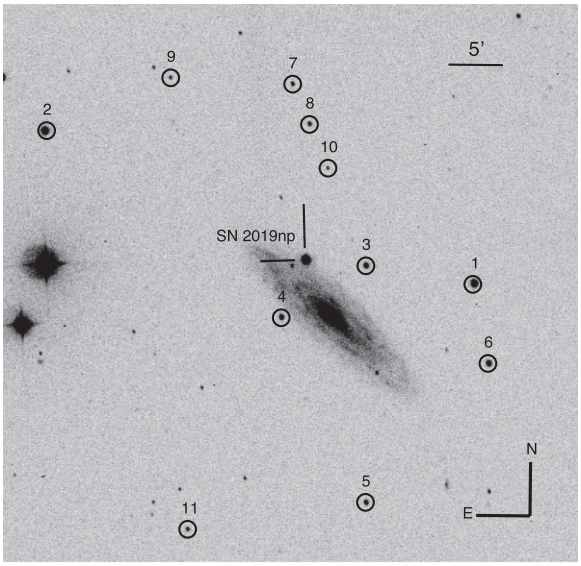

Figure 1. Image of SN 2019np in NGC 3254, taken with Tsinghua-NAOC 0.8 m telescope. Reference stars are marked with black circles and listed in Table A1. North is up and east is left.

2 OBSERVATIONS

SN 2019np was discovered on 2019 January 9.67 (UT) by K. Itagaki in NGC 3254, which is an Sbc-type spiral galaxy (see Fig. 1) with a redshift of only 0.00452. The classification spectrum obtained ∼1 d after the discovery suggests that SN 2019np was a very young, normal SN Ia, at ∼15 d before the maximum luminosity (Kilpatrick & Foley 2019; Wu et al. 2019).

2.1 Photometry

Photometric observations were obtained through various facilities including: (1) the 0.8 m Tsinghua-NAOC telescope (TNT) in China (Huang et al. 2012); (2) the Lijiang 2.4 m telescope (LJT) at Yunnan observatory; (3) the 0.8 m Joan Oró Telescope+MEIA3 at Montsec astronomical observatory (Spain); (4) the 0.67/0.92 m Schmidt telescope at Mount Ekar ebservatory (Italy); (5) the 1.8 m Copernico telescope+AFOSC at Mount Ekar observatory (Italy); (6) the 2.6 m Nordic optical telescope (NOT)+ALFOSC at the Roque de Los Muchachos observatory (Spain); (7) the 0.6 m reflector telescope+BITRAN-CCD at Itagaki astronomical observatory (Japan). All images were pre-processed, including bias subtraction, flat-field correction, and cosmic ray rejection with the standard IRAF[1] packages. We use the pipeline `Zuruphot` developed for automatic photometry of the TNT iamges (Mo et al. in prep) to perform point spread function (PSF) photometry for both SN 2019np and local reference stars.

[1]IRAF is distributed by the National Optical Astronomy Observatories, which are operated by the Association of Universities for Research in Astronomy, Inc., under cooperative agreement with the National Science Foundation (NSF).

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The instrumental magnitudes were then converted to the Johnsons *BV* (Johnson et al. 1966) and Sloan digital sky survey (SDSS) *gri*-band photometry (Fukugita et al. 1996), based on the magnitude of 10 relatively bright local comparison stars from the SDSS data release nine catalogue. SN 2019np was also observed by the Ultraviolet/Optical telescope (UVOT; Gehrels et al. 2004; Roming et al. 2005) on the Neil Gehrels *Swift* observatory (Gehrels et al. 2004) in three UV (*uvw*2, *uvm*2, *uvw*1) and three optical filters (*u*, *b*, *v*). Photometry was extracted using the HEASOFT[2] with the latest *Swift* calibration database.[3] The final optical and UV magnitudes are reported in Tables A5 and A2, respectively.

2.2 Spectroscopy

Optical spectra of SN 2019np were obtained by the Lijiang 2.4 m telescope (LJT+YFOSC); Xinglong 2.16 m telescope (XLT+BFOSC) (Zhang et al. 2016a); the 1.8 m Copernico telescope+AFOSC at the Mount Ekar observatory (Italy); the Alhambra faint object spectrograph camera (ALFOSC) on the 2.56 m Nordic optical telescope (NOT[4]); the 10.4 m Gran telescopio CANARIAS (GTC)+OSIRIS at the Roque de Los Muchachos observatory (Spain); and the 2.0 m Faulkes telescope north (FTN)+FLOYDS at Haleakala observatory (USA), spanning from −16.7 d to 367.8 d relative to the *B*-band maximum light. A journal of spectroscopic observations of SN 2019np is presented in Table A3. We reduced all the spectra using the standard IRAF routine. Flux calibration of the spectra was performed with spectrophotometric standard stars taken on the same nights. We correct the atmospheric extinction using the extinction curves of local observatories. The telluric correction was derived from the spectrophotometric standard star spectral observations and applied to the SN spectra.

Two near-infrared (NIR) spectra of SN 2019np were obtained using the Folded port InfraRed Echellette (FIRE, Simcoe et al. 2013) spectrograph mounted on the 6.5-m Magellan Baade telescope at Las campanas observatory, Chile. The spectra cover a wavelength range of 0.8–2.5 μm, with the high-throughput prism mode coupled to a 0.″6 slit. For telluric correction, we observed an A0V star close in time and at approximately similar airmass to SN 2019np (Hsiao et al. 2015, 2019). The spectra were reduced using the IDL pipeline `firehose` (Simcoe et al. 2013). We also present two NIR spectra obtained with the SpeX spectrograph (Rayner et al. 2003) mounted on the NASA infrared telescope facility (IRTF) in Hawaii. A short cross-dispersed mode has been used and a 0.″5 slit was placed on the target, providing a wavelength coverage of 0.8–2.5 μm. The SpeX data were reduced using the IDL code `Spextool` (Cushing, Vacca & Rayner 2004). A journal of spectroscopic observations of SN 2019np is presented in Table A4.

3 PHOTOMETRY

3.1 Optical and UV light curves

Fig. 2 shows the UV and optical light curves of SN 2019np sampled during the period from −16.5 to +32.8 d and that from −17.7 to 419.5 d relative to the *B*-band maximum light, respectively. Similar to other normal SNe Ia, the light curves of SN 2019np show shoulders in the *R*/*r*-bands. Secondary peaks can be identified in *I*/*i*-band light curves, for which the first peak appears about two days earlier than the *B*-band peak.

We applied high-order polynomial fits to the light curves of SN 2019np around the maximum light and estimated that it reached a *B*-band peak magnitude of 13.55 ± 0.02 mag on MJD 58510.2 ± 0.8 and a *V*-band peak magnitude of 13.62 ± 0.15 mag on 58512.5 ± 0.7, respectively. The post-peak magnitude decline in 15 d from the *B*-band peak, $\Delta m_{15}(B)$ (Phillips et al. 1999), as determined using the light-curve fitting tools SALT2 (Guy et al. 2010) and SuperNovae in object-oriented Python (SNooPY2, Burns et al. 2011), is estimated as 1.038 ± 0.004 and 1.040 ± 0.069 mag, respectively. The best-fitting light-curve models and the associated parameters are presented in Fig. 3.

In Figs 4 and 5, we compare the optical and UV-band light curves, respectively, between SN 2019np and several other well-observed normal SNe Ia with similar values of Δm_{15}. The comparison sample include SN 2018oh (Li et al. 2019), SN 2017cbv (Hosseinzadeh et al. 2017; Wang et al. 2020), and SN 2011fe (Silverman et al. 2012; Tsvetkov et al. 2013; Zhang et al. 2016b; Stahl et al. 2019). The optical light curves of SN 2019np exhibit high similarities to those of SN 2017cbv ($\Delta m_{15}(B) = 1.06 \pm 0.3$ mag). Closer inspection of the UV light curves reveals that SN 2019np shows exceptionally blue UV radiation in the very early phase, which is also similar to SN 2017cbv.

3.2 Reddening

The Galactic extinction toward SN 2019np is estimated as $A_V(\mathrm{Gal}) = 0.055$ mag according to the dust map derived by Schlafly & Finkbeiner (2011). Adopting the Cardelli, Clayton & Mathis (1989) extinction law with a total-to-selective extinction ratio of 3.1, the reddening due to Milky Way is $E(B-V)_{\mathrm{Gal}} = 0.018$ mag. This is consistent with the weak Na ID absorption due to the Milky Way.

Furthermore, we examined the absorption of the interstellar Na ID doublet (5895.92, 5889.95 Å) due to the host galaxy and derived an equivalent width (EW) of 0.68Å for SN 2019np from its near-maximum-light spectra. This corresponds to a host-galaxy reddening of 0.10 ± 0.02 mag according to the empirical relation proposed by Poznanski, Prochaska & Bloom (2012) (i.e. $\log_{10}(E(B-V)) = 1.17 \times \mathrm{EW}(\mathrm{D}_1+\mathrm{D}_2) - 1.85$), and a similar result with larger dispersion has been updated by Phillips et al. (2013). We also employed SNooPy2 (Burns et al. 2011) to fit the multi-band light curves of SN 2019np to determine the host-galaxy reddening, which gives $E(B-V)_{\mathrm{host}} = 0.110 \pm 0.066$ mag. The best-fitting results are also shown in Fig. 3. We averaged the reddening estimated by the above methods and adopted $E(B-V)_{\mathrm{host}} = 0.10 \pm 0.04$ mag as the final value. Moreover, an extinction law with $R_V = 3.1$ (Cardelli et al. 1989) is adopted throughout this paper.

After correcting the reddening due to the Milky Way and host galaxy, the $B-V$ colour of SN 2019np is found to be -0.06 ± 0.03 mag around the *B*-band maximum light, in consistent with that of a normal SN Ia (see, e.g. Wang et al. 2009b).

3.3 Colour curves

In Fig. 6, we compare the colour evolution of SN 2019np with that of some well-observed SNe Ia with similar $\Delta m_{15}(B)$. At early times, SN 2019np showed blue colours similar to SN 2011fe, SN 2017cbv, and SN 2018oh. After that, the $B-V$ and $g-r$ colour curves evolve bluewards and reach the minimum value at about 10 d prior to the *B*-band maximum. Then they evolve redwards again and reach the red peak at t ∼ 30 d; and they gradually become blue after the red peak.

[2]HEASOFT, the High Energy Astrophysics Software, https://www.swift.ac.uk/analysis/software.php
[3]https://heasarc.gsfc.nasa.gov/docs/heasarc/caldb/swift/
[4]http://www.not.iac.es/instruments/alfosc/

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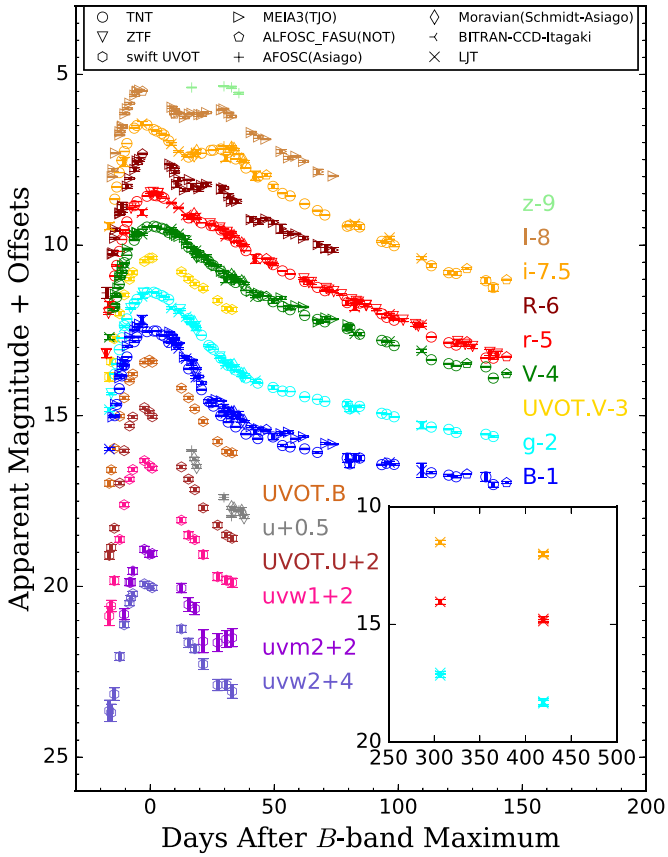

Figure 2. Optical and UV light curves of SN 2019np. Different colours represent different bands, including the *Swift uvw2, uvm2, uvw1,* UVOT *U, u*, UVOT *B, B, g*, UVOT *V, V, r, R, i, I*, and *z*. The insert panel shows the late-time light curves.

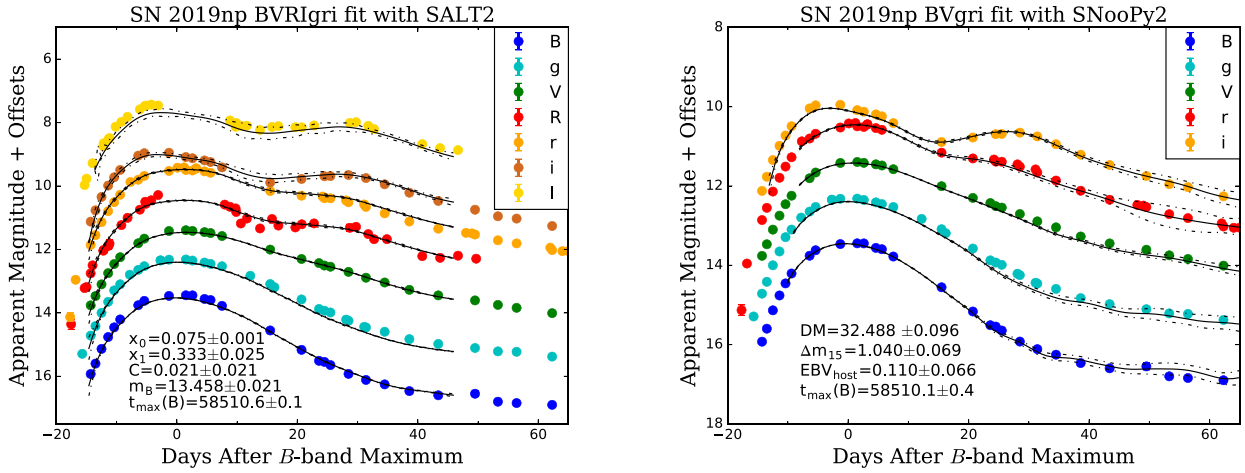

Figure 3. Best-fitting light-curve models from SALT2 (left-hand panel) and SNooPy2 (right-hand panel). The light curves of different bands have been shifted vertically for better display. The dashed lines represent the 1-σ uncertainty of the light-curve templates. The parameter x_0 represents the normalization of the SED sequence, x_1 is the light curve shape parameter, and C stands for the colour. m_B and $t_{\max}(B)$ denote the maximum magnitude of the B-band and the corresponding time. DM is the distance modulus.

4 SPECTROSCOPY

4.1 Temporal evolution of the optical spectra

Fig. 7 shows the spectral evolution of SN 2019np in the optical wavelength range. The spectral evolution follows that of normal SNe Ia. For example, within the first weeks after the explosion, the spectra of SN 2019np are characterized by prominent absorption lines of intermediate-mass elements (IMEs) and ionized IGEs. By the time SN 2019np stepped into the early nebular phase, absorption features of IGEs start to dominate the spectra.

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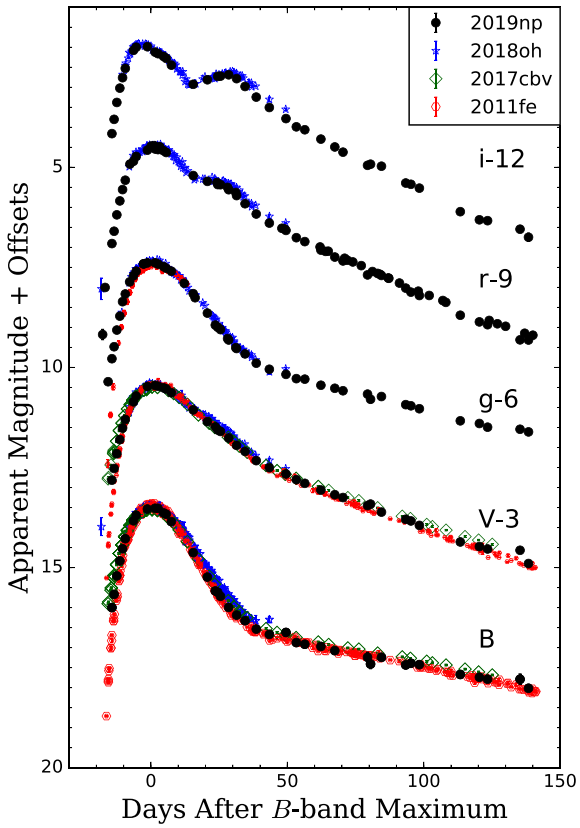

Figure 4. Comparison of the optical light curves of SN 2019np with other well-observed SNe Ia with similar decline rates. The light curves of the comparison SNe Ia have been normalized to match the peaks of SN 2019np.

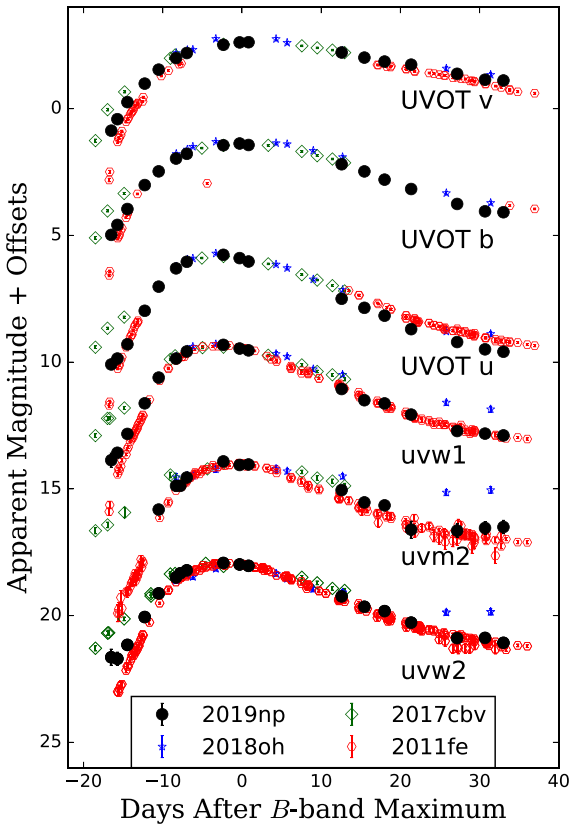

Figure 5. The UV light curves of SN 2019np compared with other well-sampled SNe Ia with similar decline rates. The light curves of the comparison SNe Ia are normalized to match the peaks of SN 2019np.

Fig. 8 shows the spectral comparison between SN 2019np and some well-observed normal SNe Ia with similar decline rates at different epochs. The earliest spectrum obtained at t ~ -16.6 d is dominated by prominent absorption lines of IMEs and IGEs. Note that weak absorption features of C II λ6580 and C II λ7234 can clearly be identified in its earliest spectra. The absorption feature near 4300 Å could be due to the Mg II λ4481 line blended with the Fe II λ4404 line, while the blended lines of Fe III λ5129, Fe II λλλ4924, 5018, 5169, and Si II λ5051 are responsible for the broad absorption near 4800 Å. The Ca II H&K and Ca II NIR triplet contribute to the prominent absorption features near 3700 Å and 8000 Å, respectively. In the spectrum at t ~ -10.5 d, the 'W'-shaped S II absorption features near 5400 Å and the Si II λ5972 feature near 5800 Å start to show up in the spectra and gain their strengths. By this phase, the C II λ6580 absorption is still visible in SN 2019np, SN 2011fe, and SN 2018oh, but it disappears in SN 2012cg and SN 2013dy. Detached high-velocity feature (HVF) can be clearly seen in Ca II NIR triplet of SN 2019np, which is similar to the comparison SNe Ia. The spectra around the maximum light are characterized by the prominent absorption features of Ca II H&K, S II, Si II λ6355, and Ca II NIR triplet. Note that the C II λ6580 absorption feature still seems visible in SN 2019np, SN 2011fe, and SN 2018oh. The HVF of the Ca II NIR triplet becomes weak while the photospheric component gains its strength. At t $\sim$ 1 month, the characteristic Fe II features within the wavelength range of 4600–5300Å start to develop and dominate in the spectra. At this phase, the overall spectral features of SN 2019np and the comparison SNe Ia become quite similar. Spectroscopically, SN 2019np is found to be more similar to SN 2011fe.

Fig. 9 shows the spectra taken at t $\sim$ 144 d and t $\sim$ 303 d after the maximum light, when forbidden lines of ionized IGEs, such as Fe and Co, dominate in the spectra. One can see that the late-time behaviour of SN 2019np also exhibits considerable similarities to that of SNe 2011fe, 2013dy, and 2018oh. Combining the spectra at t $\sim$ 144 d and $\sim$303 d, we notice that the [Fe III] feature at $\sim$4700Å [Co II] at $\sim$6000Å and [Fe II]/[Ni II] at $\sim$7200Å tend to become relatively weak with time. From the t $\sim$ 303 d spectrum, we measured the velocity shift of forbidden emission lines of [Fe II]7155Å and [Ni II]7378Å as 1550 ± 140 km s^{-1}, which is consistent with that of the comparison normal SNe Ia (i.e. SN 2011fe and SN 2018oh).

4.2 Ejecta velocity

We measure the ejecta velocities of SN 2019np using the absorption minima of a series of spectral features, including Si II λ6355, S II λ5468, C II λ6580, Ca II NIR triplet, and the result is shown in Fig. 10. All velocities have been corrected to the rest-frame of the host galaxy based on its redshift. The photospheric velocity of Si II λ6355 at t ~ -16.6 d is $\sim$16 100 km s^{-1}, which is higher compared to that measured from the C II λ6580 velocity of $\sim$15 000 km s^{-1}.

The velocity of Si II λ6355 is measured as 10 200 km s^{-1} at the time of B-band maximum, which is comparable to the typical value (i.e. $\sim$10 500 km s^{-1}) of normal SNe Ia and can be clearly put into the normal velocity group according to the classification scheme proposed by Wang et al. (2009c). The velocity of Si II λ6355 exhibits a similar evolution as that measured from the C II λ6580 and S II λ5468.

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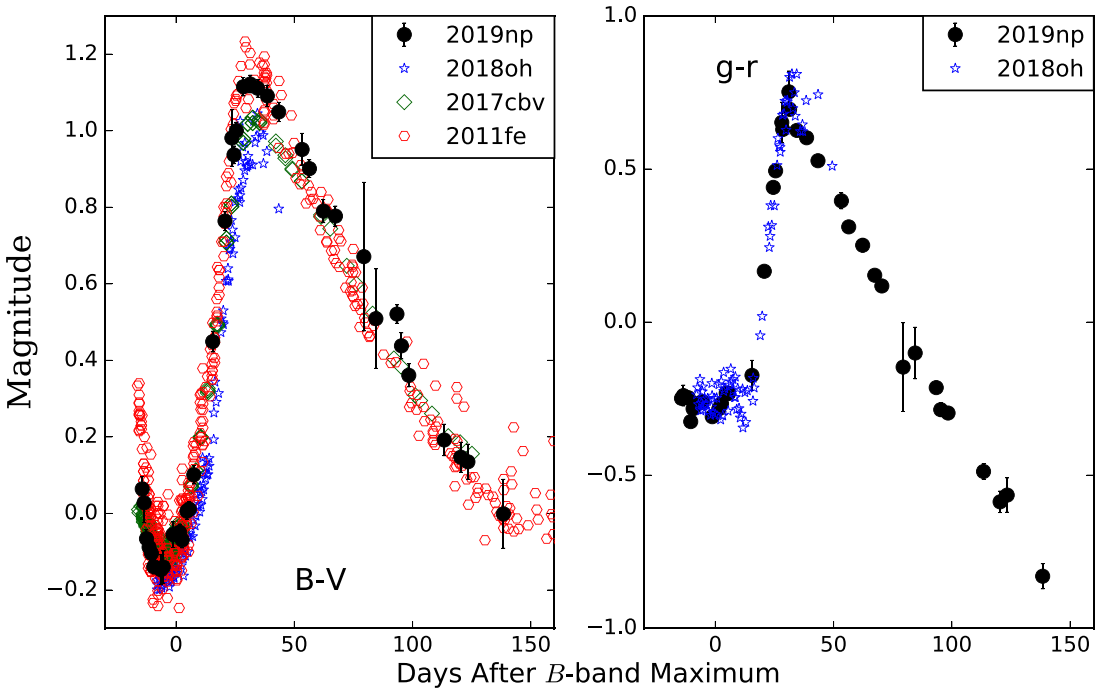

Figure 6. Reddening-corrected $B-V$ (left-hand panel) and $g-r$ (right-hand panel) colour curves of SN 2019np, compared with those of SNe 2018oh, 2017cbv, and 2011fe.

Basic photometric and spectroscopic parameters of SN 2019np are listed in Table 1. The velocity evolution of SN 2019np as derived from the Si II λ6355 line is presented in Fig. 11. The velocity evolution of SNe 2018oh, 2013dy, 2005cf (Wang et al. 2009b), and 2011fe are also shown for comparison. The velocity gradient of Si II λ6355 measured 10 days past maximum is $\dot{v}_{\rm Si} = 21 \pm 5\,{\rm km\,s^{-1}\,d^{-1}}$, indicating that SN 2019np belongs to the low-velocity gradient (LVG) subclass according to the classification scheme proposed by Benetti et al. (2005).

4.3 NIR spectra

Fig. 12 shows the NIR spectra of SN 2019np taken with IRTF and FIRE. We also compared the $t = +23.0$ d and $+56.0$ d spectra with SN 2011fe and SN 2017cbv. The non-detection of the Paβ feature indicates that the mass of hydrogen contained in companion star is less than 0.1 $M_\odot$ (Wang et al. 2020). The lack of hydrogen features in the nebular spectra is also in accord with a low hydrogen mass limit (Sand et al. 2018).

Moreover, the H-band break is observed around 1.5 μm in the NIR spectra of SN 2019np, which is related to the dramatic shift in the amount of line blanketing from IGEs (Wheeler et al. 1998). Similar to other SNe Ia (Sand et al. 2018), a decrease in the strength of the H-band break can be identified in SN 2019np during the phases from $t = +23.0$ to $+108.8$ d. Such a time evolution of the H-band break is related to the quantity of intermediate-mass elements, depending on different explosion scenarios, which can be further used to determine mass of ^{56}Ni (Ashall et al. 2019; Hsiao et al. 2019).

5 DISCUSSION

5.1 Quasi-bolometric light curve

The distance of the host galaxy NGC 3254 is 32.80 ± 6.8 Mpc according to Tully–Fisher relation (Tully, Courtois & Sorce 2016). Assuming $R_{\rm V} = 3.1$, the absolute magnitude of B-band peak is estimated as $M_{\rm max}(B) = -19.52 \pm 0.47$ mag after correcting for the Galactic and host-galaxy extinctions. We derived the quasi-bolometric luminosity of SN 2019np based on *Swift* UV and optical photometry. Assuming that the NIR-band emission of SN 2019np has a similar contribution as in SN 2011fe (Zhang et al. 2016b), the quasi-bolometric light curve of SN 2019np is shown in Fig. 13. The maximum-light luminosity is $L_{\rm peak} = 1.70 \pm 0.63 \times 10^{43}\,{\rm erg\,s^{-1}}$, which is slightly larger than SN 2011fe ($L_{\rm peak}^{\rm 11fe} = 1.13 \times 10^{43}$ erg s^{-1}, Zhang et al. 2016b). We use the `Minim` Code (Chatzopoulos et al. 2013), based on the radiation diffusion model from Arnett (Arnett 1982; Chatzopoulos, Wheeler & Vinko 2012; Li et al. 2019; Zeng et al. 2021), to estimate the nickel mass and other parameters. The relevant result is shown in Fig. 14. The fitting parameters include first-light time $t_0 = 58494.09 \pm 0.03$, the mass of the radioactive nickel $M_{\rm Ni} = 0.66 \pm 0.05\,M_\odot$, the light-curve time-scale $t_{\rm lc} = 11.19 \pm 0.01$ d, and the gamma-ray leaking time-scale $t_\gamma = 37.42 \pm 0.70$ d (Chatzopoulos et al. 2012). Then the inferred ejecta mass ($M_{\rm ej}$) and expansion velocity ($v_{\rm exp}$) can be calculated from $t_{\rm lc}$ and t_γ:

$$t_{\rm lc}^2 = \frac{2\kappa M_{\rm ej}}{\beta c v_{\rm exp}} \quad \text{and} \quad t_\gamma^2 = \frac{3\kappa_\gamma M_{\rm ej}}{4\pi v_{\rm exp}^2} \tag{1}$$

(Arnett 1982; Clocchiatti & Wheeler 1997; Valenti et al. 2008; Chatzopoulos et al. 2012; Wheeler, Johnson & Clocchiatti 2015; Li et al. 2019), where κ is the effective opacity in optical, κ_γ is the opacity for γ rays (assuming that positrons released during cobalt decay are completely captured) and $\beta \sim 13.8$ is the light curve parameter related to the density profile of the ejecta (Arnett 1982). The κ_γ can be estimated as $\kappa_\gamma \sim 0.03$ cm^2 g^{-1} (Wheeler et al. 2015). So the $M_{\rm ej}$ and $v_{\rm exp}$ only depend on the value of the optical opacity κ. The $M_{\rm ej}$ should be smaller than the Chandrasekhar mass (i.e. $M_{\rm ej} < M_{\rm ch}$) and the $v_{\rm exp}$ should be at least as large as

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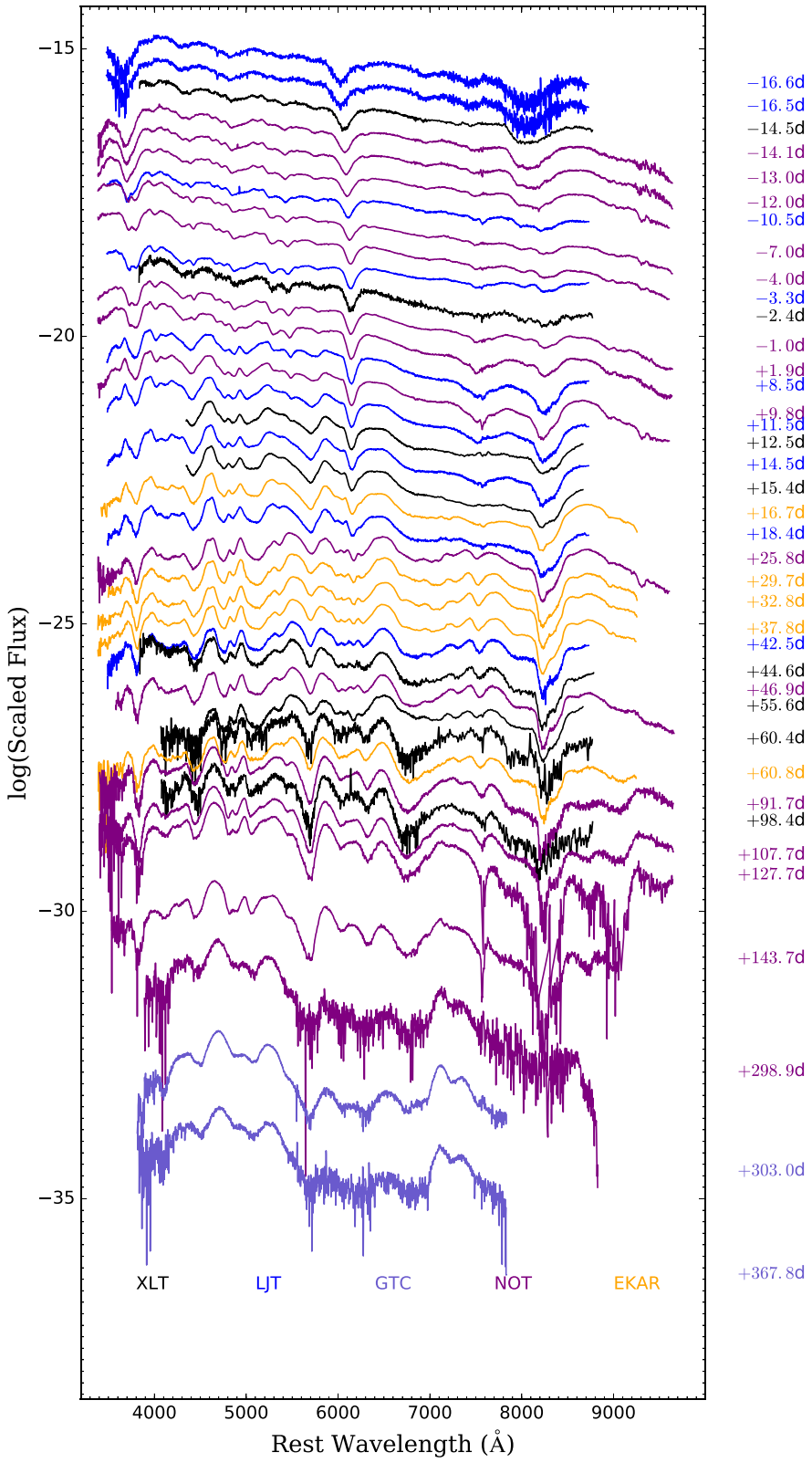

Figure 7. Optical spectral evolution of SN 2019np from −16.6 d to +367.8 d relative to the *B*-band maximum light. The spectra have been corrected for host-galaxy redshift ($z = 0.00452$) and reddening. The text on the right-hand side of each spectrum denotes the phase in days due to the *B*-band maximum light. Different colours of the spectra represent that they were taken with different spectroscopic instruments (i.e. XLT, LJT, GTC, NOT, or EKAR), which are indicated at the bottom of the plot.

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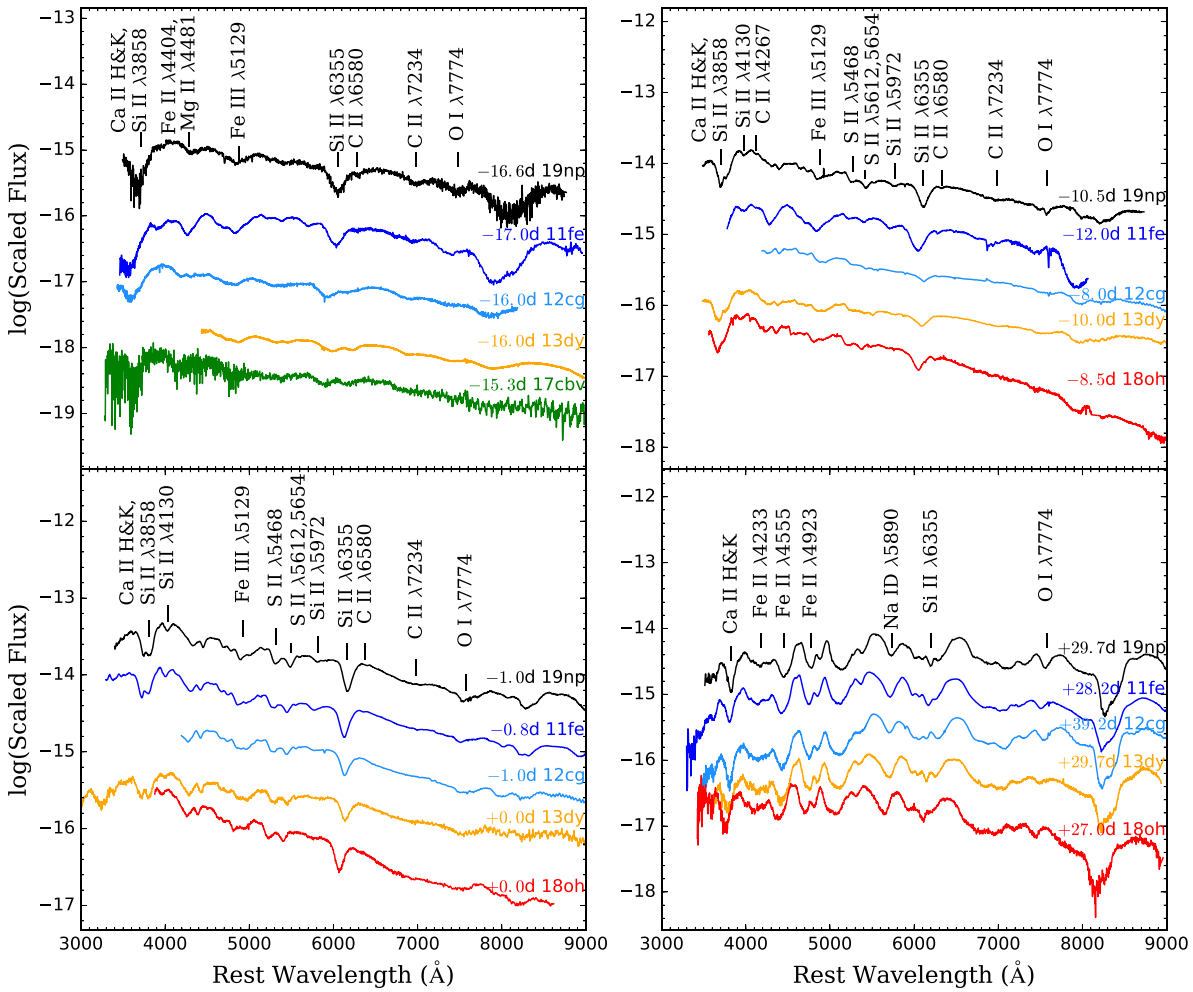

Figure 8. Spectral comparison between SN 2019np and some well-studied SNe Ia at several selected epochs (i.e. t ∼ −16.0 d, −10.0 d, 0 d, and +30.0 d). The comparison objects include SNe 2011fe (Mazzali et al. 2014; Zhang et al. 2016b), 2012cg (Marion et al. 2016), 2013dy (Zheng et al. 2013; Pan et al. 2015; Zhai et al. 2016), 2017cbv (Hosseinzadeh et al. 2017), and 2018oh (Li et al. 2019). All spectra have been corrected for reddening and host-galaxy redshifts. For better display, all spectra were shifted arbitrarily along the vertical direction.

the observed expansion velocity ($v_{\rm exp} > 11\,000$ km s^{-1}). With these constraints, we derive $0.08 \lesssim \kappa \lesssim 0.09$ cm^2 g^{-1}. Thus, we adopt $\kappa \sim 0.085$ cm^2 g^{-1} for SN 2019np (see also Li et al. (2019) for a similar analysis of SN 2018oh). Then we get the ejecta mass $M_{\rm ej} = 1.34 \pm 0.12 M_\odot$ and the kinetic energy of the supernova as $E_{\rm sn} = 0.3 M_{\rm ej} v_{\rm exp}^2 = 1.65^{+0.17}_{-0.15} \times 10^{51}$ erg.

5.2 The early-excess flux

In Figs 14 and 5, one can see that excess flux exists in the early-time light curves. Such an excess emission is hard to be explained by diffusion of centrally located ^{56}Ni, but it could be related to the collision between the SN-ejecta and the companion star. Such a strong interaction would produce an optical/UV excess in the first week after the explosion. Thus, we use a hybrid model to fit the early-time light curves of SN 2019np. Assuming that the density distribution of the ejecta follows a broken power law, i.e. $\rho \propto r^{-1}$ for the inner region and $\rho \propto r^{-10}$ for the outer one, the collision-powered SN luminosity ($L_{\rm col}$) and effective temperature ($T_{\rm col}$) can be expressed as follows (Kasen 2010):

$$L_{\rm col} \propto a M_{\rm ej}^{-5/8} E_{\rm sn}^{7/8} \kappa^{-3/4} t^{-1/2}, \quad (2)$$

$$T_{\rm col} \propto a^{1/4} \kappa^{-35/144} t^{-37/72}, \quad (3)$$

where a is the distance between the mass center of the WD and its companion, and t represents the dynamical time. We can then derive the time-dependent collision-powered flux ($F_{\nu,\,\rm col}$) under the assumption of blackbody radiation. For simplicity, the contribution from ^{56}Ni-powered radiation is described by the fireball model ($F_{\nu,\,\rm Ni} \propto t^2$; Nugent et al. 2011). So the total flux can be given by $F_\nu = F_{\nu,\,\rm col} + F_{\nu,\,\rm Ni}$.

In this circumstance, we still set $\kappa \sim 0.085$ cm^2 g^{-1} and adopt the above fitting results ($M_{\rm ej}$, $E_{\rm sn}$) from Arnett model. Then we fit the data observed earlier than MJD 58497 and derive the best-fitting parameters, which suggests that the explosion date is MJD 58495.2 ± 0.1 and the separation distance between the WD and its companion star (a) is 5.74 ± 0.38 × 10^{11} cm. In most cases, the separation distance is comparable to the Roche lobe radius of the companion star, for typical mass ratios, $a/R_{\rm companion} = 2$–3 (Kasen 2010). So the companion star could have a radius of (1.91–2.87) × 10^{11} cm. Thus, the corresponding companion star is likely to be a main-sequence star of $\sim 1 M_\odot$ (see table 1 in Kasen 2010). The best-fitting result is presented in Fig. 15.

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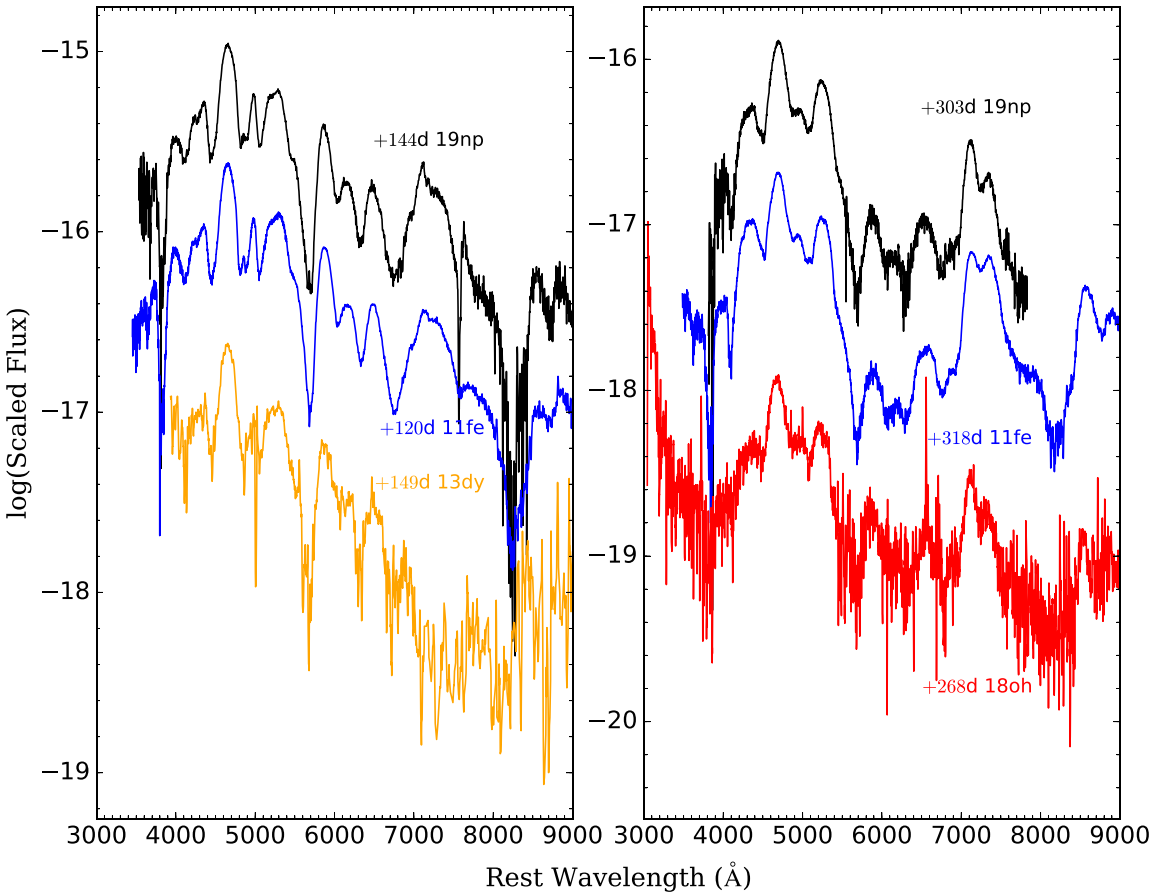

Figure 9. Nebular-phase spectra of SN 2019np obtained at t $\sim$ +144 d (left-hand panel) and +303 d (right-hand panel) compared with the spectra of a few well-observed SNe Ia at similar phases, including SNe 2011fe (Silverman et al. 2012; Yaron & Gal-Yam 2012), 2013dy (Zheng et al. 2013; Zhai et al. 2016), and 2018oh (Tucker et al. 2019). All spectra were corrected for host-galaxy redshifts. The yellow and red spectra have been smoothed by a Savitsky–Golay filter.

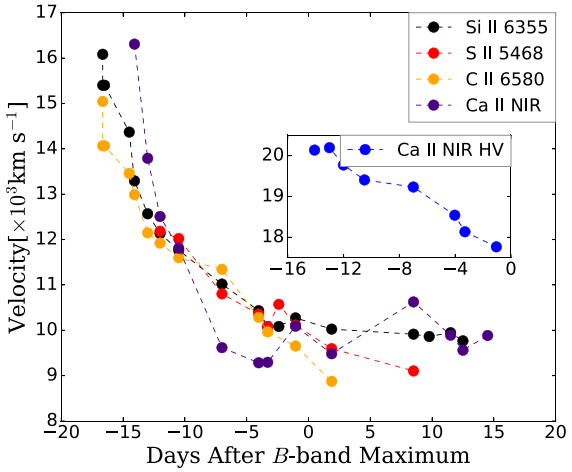

Figure 10. Evolution of the expansion velocity of SN 2019np as measured from the absorption minimum of Si II λ6355, S II λ5468, C II λ6580, and Ca II NIR triplet.

Table 1. Parameters of SN 2019np.

Parameter	Value
	Photometric
$B_{\rm max}$	13.55 ± 0.02 mag
$B_{\rm max} - V_{\rm max}$	0.005 ± 0.007 mag
$M_{\rm max}(B)$	-19.52 ± 0.47 mag
$E(B-V)_{\rm host}$	0.10 ± 0.04 mag
$\Delta m_{15}(B)$	1.04 ± 0.04 mag
s_{BV}	0.98 ± 0.01
$t_{\rm max}(B)$	58510.2 ± 0.8 d
t_0	58494.09 ± 0.03 d
$\tau_{\rm rise}$	15.1 ± 0.8 d
$L_{\rm bol}^{\rm max}$	$1.70 \pm 0.63 \times 10^{43}$ erg s^{-1}
$M_{^{56}\rm Ni}$	0.66 ± 0.05 M$_\odot$
μ	32.58 ± 0.45 mag
	Spectroscopic
v_0(Si II)	$10,200 \pm 141$ km s^{-1}
$\dot{v}$(Si II)	21 ± 5 km s^{-1} d^{-1}
R(Si II)	0.17 ± 0.01

In addition, there are other possible explanations for the early bump in the light curves and one particular mechanism is the ^{56}Ni mixing with the outermost layers (Piro & Nakar 2013; Dimitriadis et al. 2019). In this scenario, the mass fraction of ^{56}Ni near the surface exceeds that in the 'expanding fireball' model, which can produce extra flux, and it can happen in both SD and DD progenitor channels. A specific explosion model for such a configuration is the double-detonation sub-Chandrasekhar explosion, where the surface ^{56}Ni can

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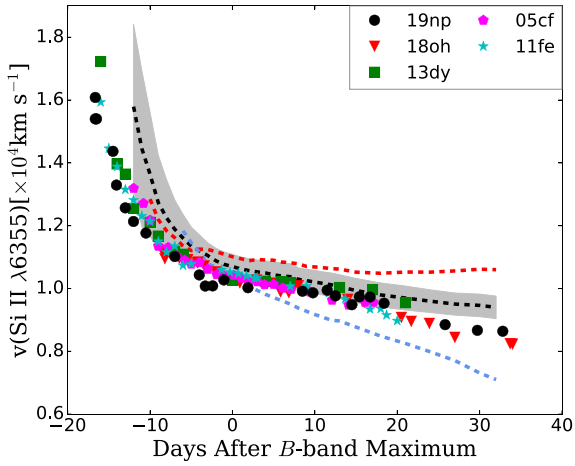

Figure 11. Velocity evolution of SN 2019np as measured from the absorption minimum of Si II λ6355, compared to that of SNe 2018oh, 2013dy, 2005cf, and 2011fe. Overplotted curves present the velocity evolution obtained for SN 1991T-like (red-dashed), SN 1991bg-like (blue-dotted), and normal (solid black) subclasses of SNe Ia (Wang et al. 2009c). The grey-shaded region indicates the 1-σ standard deviation of the mean velocity of normal SNe Ia.

be produced through the detonation of surface helium (Noebauer et al. 2017).

For the double detonation model (DDM), the explosion is thought to start with a shell of helium-rich material accreted by a sub-Chandrasekhar WD from its companion. The shell donation drives a shock wave that propagates towards the core of the CO WD, which could cause a secondary explosion in the core. The core explosion completely disrupts the WD (Woosley & Weaver 1994). The DDMs 1–6, with parameters spanning from 0.810–1.385$M_{\odot}$ WDs and the helium shells from 0.0035–0.126$M_{\odot}$ (Fink et al. 2010; Kromer et al. 2010), are available from the Heidelberg supernova model archive.[5] We compared the bolometric light curves of SN 2019np with those predicted by the above models in Fig. 15. The explosion time given by Arnett model is $t_0 = 58494.09 \pm 0.03$ d. Note that the peak bolometric luminosity predicted by DDMs 4–6 is higher than the observations. In addition, the DDMs 4–6 predict a shoulder in the bolometric light curve at around ∼25.0 d, which is not seen in SN 2019np.

Furthermore, it is possible that ^{56}Ni can also be mixed to the outer layer in some other scenarios (Piro & Nakar 2013; Piro & Morozova 2016). We consider a more general model in which the amount of surface ^{56}Ni is tuned to explore the influence on the rising light curves (Piro & Morozova 2016). In Fig. 15, we present some typical light curves given by theoretical models (Piro & Morozova 2016), which adopted a fixed amount of ^{56}Ni (i.e. 0.5 $M_{\odot}$) and vary the distribution with a boxcar and a width from 0.05 to 0.25 $M_{\odot}$. As mentioned in Stritzinger et al. (2018) and Dimitriadis et al. (2019), the early-time colour is an important parameter to distinguish the above models. In Fig. 16, we show the colour evolution of SN 2019np and the expected colours for different models. SN 2019np shows broader red–blue–red (in $B - V$ colour) evolution than the DDMs. Moreover, the $g - r$ colour evolution of SN 2019np is quite flat and does not conform with that predicted by the DDM. For the interaction model, the early-time colour was initially blue and then became red, which is inconsistent with the observed colour evolution of SN 2019np. Based on the discussions above, we conclude that the early flux excess in SN 2019np is most consistent with the ^{56}Ni mixing model.

[5]https://hesma.h-its.org/

6 CONCLUSION

We present extensive UV-optical photometry and optical-NIR spectroscopy of the Type Ia SN 2019np. The following conclusions can be made:

(1) The light curves suggest that SN 2019np resembles that of other well-sampled normal SNe Ia. A B-band peak magnitude of 13.55 ± 0.02 mag was reached, corresponding to an absolute magnitude of -19.52 ± 0.47 mag. A B-band luminosity decline rate is found to be $\Delta m_{15}(B) = 1.04 \pm 0.04$ mag. Both values are consistent with the photometric behaviour of normal SNe Ia;

(2) We construct the quasi-bolometric light curve of SN 2019np based on the UV-optical photometry and an adoption of a NIR flux correction. The estimated peak bolometric luminosity gives $L_{\rm peak} = 1.70 \pm 0.63 \times 10^{43}$ erg s^{-1}, indicating a synthesized nickel mass of 0.66 ± 0.05 M$_{\odot}$;

(3) Photometry of SN 2019np started within the first two days of the explosion reveals an early blue bump in its light curves. The early bolometric flux evolution is up to ∼5.0 per cent higher compared to that predicted by the diffusion of internal radiation through a homogeneous expanding ejecta (Arnett 1982), suggesting the presence of energy sources in addition to the radioactive decay of nickel;

(4) We also compare the early colour evolution of SN 2019np to various models, including i) interaction between the SN ejecta and the CSM/companion star, ii) double-detonation explosion of a sub-Chandrasekhar mass WD, and iii) nickel mixed into the surface layers of the SN ejecta. We suggest that model iii) provides the most satisfactory fit;

(5) NIR and nebular-phase spectra of SN 2019np show no evidence of a significant amount of Hydrogen, which is incompatible with the picture of the ejecta interacting with a non-degenerate companion or H-rich CSM.

Early observations of SNe Ia, especially those with conspicuous flux excess that deviates from radiation diffusion through a homogeneously expanding ejecta within the first few days ('early-excess SNe Ia'), play an important role in understanding the SNe Ia explosion mechanism and the nature of their progenitors (Jiang et al. 2018; Stritzinger et al. 2018). To our knowledge, only a handful of SNe Ia have been reported to show such an early excess, e.g. SNe 2012cg (Marion et al. 2016), 2013dy (Zheng et al. 2013; Zhai et al. 2016), 2017cbv (Hosseinzadeh et al. 2017; Wang et al. 2020), 2018oh (Dimitriadis et al. 2019), iPTF13dge (Ferretti et al. 2016). The ejecta of the SN Ia expand rapidly and wipe out almost all traces of the pre-explosion configuration within days after the explosion of the progenitor WD. Therefore, photometry and spectroscopy starting within days after the explosion provides a unique view of the kinematics and mean chemical composition of the outermost layers of the SN ejecta and their circumstellar environments. High-cadence monitoring is also essential to probe these properties from the outer to inner parts of the ejecta as the photosphere continuously recedes into the exploding WD. Facilitated by the alert streams of modern wide-field, high-cadence transient searches, such as the Zwicky Transient Facility (Bellm et al. 2019a, b), ATLAS (Tonry 2011), ASAS-SN (Shappee et al. 2014), DLT40 (Tartaglia et al. 2018) and TMTS (Lin et al. 2022), rapid photometric and spectroscopic follow-up observations will provide a more comprehensive characterization of the early behaviours of SNe Ia. An extensive sample of early SNe Ia

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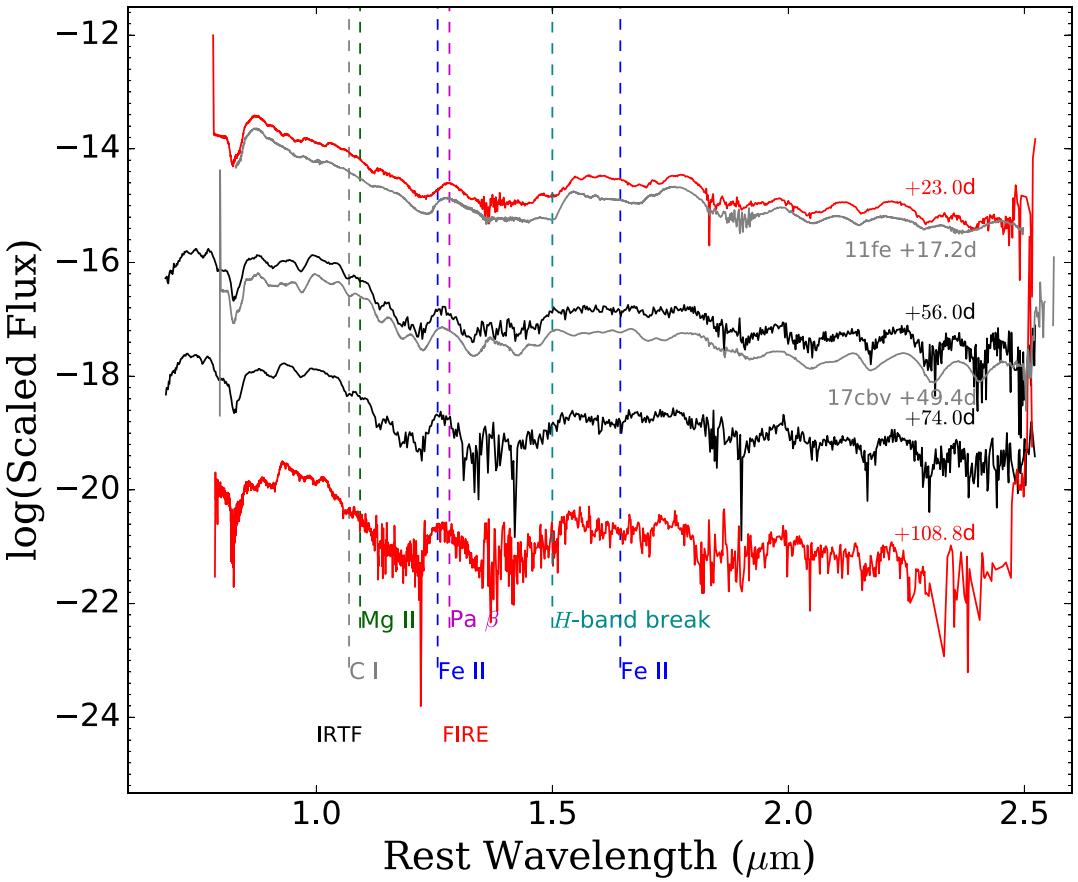

Figure 12. NIR spectra of SN 2019np from FIRE and IRTF, along with the comparable-phase spectra of SN 2011fe (Hsiao et al. 2013) and SN 2017cbv (Wang et al. 2020).

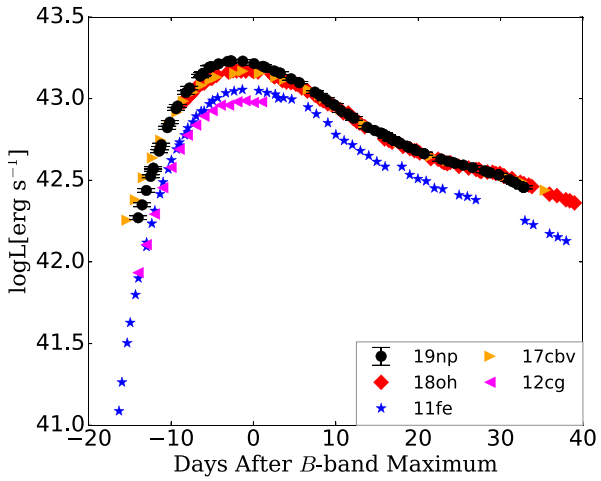

Figure 13. Quasi-bolometric light curve of SN 2019np compared to that of SNe 2018oh, 2011fe, 2012cg, and 2017cbv.

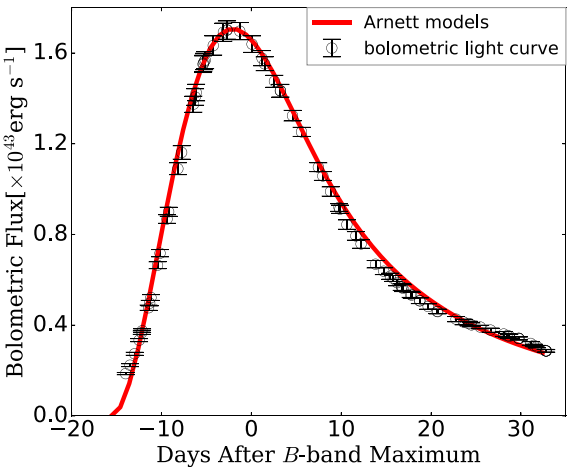

Figure 14. The quasi-bolometric light curve of SN 2019np (open black circles) compared to the best-fitting radiation diffusion model (red line; Arnett 1982).

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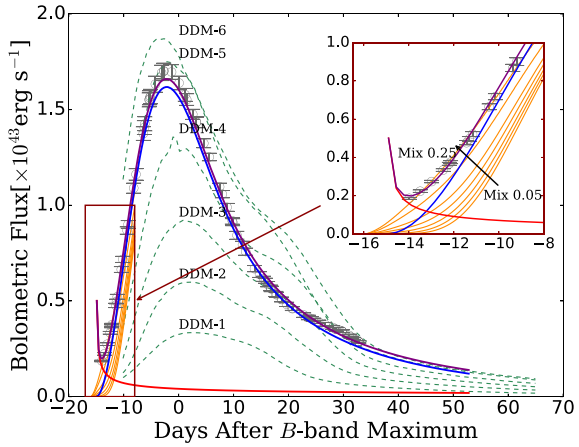

Figure 15. Bolometric light curve of SN 2019np (grey open dots) compared with various models as indicated by different colour-coded curves. The solid blue line represents the radioactive decay model, and the red solid line denotes the interaction model, while the purple solid line shows the sum of the above two. DDMs are shown in green dashed lines. Six models (DDM-1 to 6) are computed for an initial CO core of the WD and He shell masses of ($0.810M_{\odot}$, $0.126M_{\odot}$), ($0.920M_{\odot}$, $0.084M_{\odot}$), ($1.025M_{\odot}$, $0.055M_{\odot}$), ($1.125M_{\odot}$, $0.039M_{\odot}$), ($1.280M_{\odot}$, $0.013M_{\odot}$), and ($1.385M_{\odot}$, $0.0035M_{\odot}$), respectively. The inset portrays the data and the models within the first week as the region outlined by the red square. The orange lines denote the ^{56}Ni mixing models; and the series of lines following the direction of the black arrow are generated using a boxcar with widths of 0.05, 0.075, 0.1, 0.125, 0.15, 0.2, and 0.25 $M_{\odot}$.

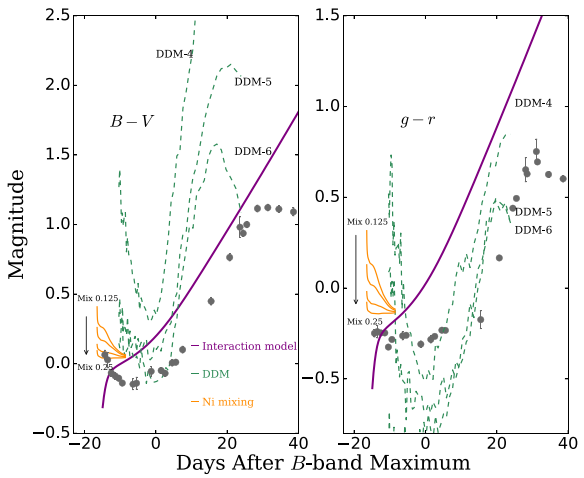

Figure 16. The $B-V$ bands (left-hand panel) and $g-r$ bands (right-hand panel) colour of SN 2019np (grey-dots) compared with various explosion models as indicated by different colour-coded curves. The models are shown in the same colour with Fig. 15.

will enable stringent tests among various models that may account for the flux excess in their early luminosity evolution. It will also provide chances to test the aspect-angle dependency and uniformities of SNe Ia that belong to different subtypes.

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DATA AVAILABILITY

The data underlying this article are available in the article. Our photometric data has been attached in Tables A1, A2 and A5 in the appendix. The log tables of optical and NIR spectral data are also shown in Table A3 and Table A4. The spectra presented in this article will be available on the WISeREP repository (Yaron & Gal-Yam 2012).

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APPENDIX A: PHOTOMETRIC AND SPECTROSCOPIC DATA

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Table A1. Photometric standards in the SN 2019np field for *BVgri* bands.

Num.	α(J2000)	δ(J2000)	B (mag)	V (mag)	g (mag)	r (mag)	i (mag)
1	$10^h29^m06\overset{s}{.}86$	$29°30'01\overset{''}{.}81$	14.826(035)	14.167(012)	14.375(003)	13.969(002)	13.845(004)
2	$10^h29^m45\overset{s}{.}12$	$29°33'23\overset{''}{.}79$	14.446(036)	13.824(013)	14.014(007)	13.641(001)	13.527(013)
3	$10^h29^m16\overset{s}{.}58$	$29°30'28\overset{''}{.}17$	16.456(035)	15.745(013)	15.977(005)	15.524(002)	15.348(003)
4	$10^h29^m24\overset{s}{.}40$	$29°29'31\overset{''}{.}28$	16.825(035)	15.862(013)	16.213(005)	15.534(003)	15.272(002)
5	$10^h29^m17\overset{s}{.}43$	$29°25'47\overset{''}{.}52$	17.187(035)	16.279(014)	16.604(006)	15.974(004)	15.756(003)
6	$10^h29^m05\overset{s}{.}83$	$29°28'26\overset{''}{.}37$	16.281(035)	15.674(013)	15.857(004)	15.498(002)	15.390(003)
7	$10^h29^m22\overset{s}{.}53$	$29°34'07\overset{''}{.}31$	18.035(035)	16.958(014)	17.363(006)	16.581(004)	16.283(004)
8	$10^h29^m21\overset{s}{.}14$	$29°33'19\overset{''}{.}29$	17.820(035)	16.962(013)	17.264(005)	16.679(003)	16.456(002)
9	$10^h29^m33\overset{s}{.}57$	$29°34'21\overset{''}{.}00$	18.793(038)	17.996(021)	18.269(010)	17.739(011)	17.451(005)
10	$10^h29^m19\overset{s}{.}68$	$29°32'25\overset{''}{.}88$	19.127(035)	17.578(014)	18.206(006)	17.000(004)	15.672(002)
11	$10^h29^m33\overset{s}{.}65$	$29°25'24\overset{''}{.}47$	18.281(035)	17.325(014)	17.673(006)	17.000(004)	16.748(005)

Note. Uncertainties, in units of 0.001 mag, are 1σ.

Table A2. *Swift* photometry of SN 2019np.

MJD	*uvw2* (mag)	*uvm2* (mag)	*uvw1* (mag)	*UVOT u* (mag)	*UVOT b* (mag)	*UVOT v* (mag)
58493.7	19.644(315)	...	18.868(281)	17.091(107)	16.979(084)	16.868(123)
58494.5	19.703(255)	...	18.576(170)	16.859(085)	16.586(067)	16.415(086)
58495.7	19.156(182)	...	17.832(115)	16.295(073)	15.959(060)	15.742(069)
58497.9	18.057(114)	...	16.625(078)	14.974(055)	15.014(048)	15.012(066)
58499.7	17.125(084)	18.817(161)	15.610(063)	14.022(044)	14.473(043)	14.460(054)
58501.9	16.496(072)	17.880(104)	14.860(057)	13.300(041)	13.968(041)	14.003(048)
58502.4	16.334(068)	17.881(100)	...	...	...	...
58503.2	16.216(069)	17.552(093)	14.579(055)	13.041(040)	13.781(041)	13.805(046)
58507.8	15.925(075)	16.916(091)	14.318(055)	12.766(040)	13.443(040)	13.479(045)
58509.9	15.983(071)	17.058(086)	14.464(056)	12.896(040)	13.374(040)	13.385(044)
58511.0	16.037(072)	17.038(086)	14.536(057)	13.030(040)	13.433(040)	13.378(043)
58522.8	17.244(104)	18.048(138)	16.061(079)	14.500(050)	14.193(043)	13.783(047)
58525.7	17.653(137)	18.534(201)	16.503(101)	14.857(058)	14.475(045)	13.985(051)
58528.2	17.824(124)	18.651(179)	16.629(090)	15.168(062)	14.800(047)	14.141(051)
58531.6	18.279(156)	19.610(335)	17.070(109)	15.699(076)	15.170(051)	14.263(054)
58537.4	18.889(166)	19.647(263)	17.718(115)	16.207(073)	15.759(060)	14.624(057)
58540.9	18.882(166)	19.544(247)	17.820(121)	16.495(080)	16.050(062)	14.858(059)
58543.2	19.075(200)	19.505(263)	17.891(134)	16.590(087)	16.086(065)	14.888(062)

Note. Uncertainties, in units of 0.001 mag, are 1σ.

Table A3. Log of optical spectroscopic observations of SN 2019np.

MJD	Phase[a]	Range (Å)	Resolution. (Å)	Telescope + Inst.
58493.6	−16.6	3477–8727	25	LJT + YFOSC
58493.7	−16.5	3477-8727	25	LJT + YFOSC
58495.7	−14.5	3838–8770	15	XLT + BFOSC
58496.1	−14.1	3385–9639	18	NOT + ALFOSC
58497.2	−13.0	3385–9640	18	NOT + ALFOSC
58498.2	−12.0	3385–9606	18	NOT + ALFOSC
58499.7	−10.5	3487–8730	25	LJT + YFOSC
58503.2	−7.0	3385–9639	18	NOT + ALFOSC
58506.2	−4.0	3386–9604	18	NOT + ALFOSC
58506.9	−3.3	3483–8727	25	LJT + YFOSC
58507.9	−2.4	3833–8770	15	XLT + BFOSC
58509.2	−1.1	3385–9640	18	NOT + ALFOSC
58512.1	1.9	3385–9635	18	NOT + ALFOSC
58518.7	8.5	3484–8727	25	LJT + YFOSC
58520.0	9.8	3385–9605	18	NOT + ALFOSC
58521.7	11.5	3488–8727	25	LJT + YFOSC
58522.7	12.5	4347–8665	15	XLT + BFOSC
58524.7	14.5	3485–8728	25	LJT + YFOSC
58525.6	15.4	4348–8666	15	XLT + BFOSC
58526.9	16.7	3420–9255	24	EKAR + AFOSC

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Table A3 – *continued*

MJD	Phase[a]	Range (Å)	Resolution. (Å)	Telescope + Inst.
58528.7	18.5	3489–8727	25	LJT + YFOSC
58536.0	25.8	3385–9606	18	NOT + ALFOSC
58539.9	29.7	3498–9256	24	EKAR + AFOSC
58543.0	32.8	3385–9253	24	EKAR + AFOSC
58548.0	37.8	3385–9242	24	EKAR + AFOSC
58552.7	42.5	3489–8729	25	LJT + YFOSC
58554.8	44.6	3841–8783	15	XLT + BFOSC
58557.1	46.9	3580–9658	18	NOT + ALFOSC
58565.8	55.6	4373–8666	15	XLT + BFOSC
58570.6	60.4	4071–8769	15	XLT + BFOSC
58571.0	60.8	3389–9246	24	EKAR + AFOSC
58601.9	91.7	3403–9650	18	NOT + ALFOSC
58608.6	98.4	4077–8772	15	XLT + BFOSC
58617.9	107.7	3402–9649	18	NOT + ALFOSC
58637.9	127.7	3484–9644	18	NOT + ALFOSC
58653.9	143.7	3534–9053	18	NOT + ALFOSC
58809.2	298.9	3891–8826	18	NOT + ALFOSC
58813.2	303.0	3813–7831	7	GTC + OSIRIS
58878.0	367.8	3833–7831	7	GTC + OSIRIS

Note.[a] Days relative to the B −band maximum on MJD 58510.2.

Table A4. Log of NIR spectroscopic observations of SN 2019np.

MJD	Phase[a]	Range (Å)	R	Telescope + Inst.
58533.2	23.0	7839–25345	450	FIRE
58566.2	56.0	6835–25343	1200	IRTF
58584.2	74.0	6844–25336	1200	IRTF
58619.0	108.8	7887–25276	450	FIRE

Note.[a] Days relative to the B-band maximum on MJD 58510.2.

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Table A5. Ground-based optical photometry of SN 2019np.

MJD	*B* (mag)	*V* (mag)	*R* (mag)	*I* (mag)	*u* (mag)	*g* (mag)	*r* (mag)	*i* (mag)	*z* (mag)	data source
58492.5	...	...	...	...	...	...	18.177(131)	...	...	ZTF
58492.7	...	...	17.401(161)	...	...	...	...	...	...	BITRAN-CCD
58493.4	...	...	...	...	...	...	17.001(033)	...	...	ZTF
58493.7	16.977(043)	16.701(078)	...	...	...	16.820(060)	16.738(080)	16.950(100)	...	LJT
58494.5	...	...	...	...	...	16.356(036)	...	...	...	ZTF
58494.9	...	...	16.263(011)	16.000(022)	...	...	...	...	...	MEIA3
58495.2	16.034(027)	15.890(039)	16.229(080)	15.796(036)	...	...	...	...	...	MEIA3
58495.9	...	15.735(013)	15.797(027)	...	...	...	...	...	...	MEIA3
58495.9	15.999(021)	15.816(011)	...	...	...	15.781(010)	15.904(010)	16.157(013)	...	TNT
58496.2	15.668(023)	15.444(025)	15.545(044)	15.308(020)	...	...	...	...	...	MEIA3
58496.7	15.671(038)	15.524(014)	...	...	...	15.482(016)	15.596(018)	15.800(012)	...	TNT
58497.2	...	15.258(016)	...	...	...	...	...	...	...	MEI3A
58497.7	15.208(016)	15.155(007)	...	...	...	15.065(005)	15.186(004)	15.378(005)	...	TNT
58498.0	14.888(028)	14.977(022)	15.083(037)	14.776(024)	...	...	...	...	...	MEIAA
58498.1	14.812(019)	14.922(022)	15.060(053)	14.682(028)	...	...	...	...	...	MEIA3
58498.8	14.833(018)	14.802(006)	...	...	...	14.717(007)	14.838(005)	15.020(007)	...	TNT
58498.9	14.600(042)	14.785(044)	14.929(111)	14.547(043)	...	...	...	...	...	MEIA3
58499.0	14.409(051)	14.610(040)	14.855(111)	14.518(042)	...	...	...	...	...	MEIA3
58499.7	14.444(051)	14.544(011)	...	...	...	14.394(029)	14.644(017)	15.069(133)	...	LJT
58499.8	14.529(014)	14.514(006)	...	...	...	14.356(007)	14.555(005)	14.759(007)	...	TNT
58500.0	...	...	...	14.159(029)	...	...	...	...	...	MEIA3
58500.8	14.279(013)	14.299(006)	...	...	...	14.164(006)	14.321(006)	14.526(005)	...	TNT
58501.0	13.962(026)	14.117(028)	14.277(056)	14.009(024)	...	...	...	...	...	MEIA3
58502.0	13.695(027)	...	14.046(024)	13.846(028)	...	...	...	...	...	MEIA3
58502.4	...	...	...	...	...	13.862(020)	...	...	...	ZTF
58502.5	...	...	...	...	...	...	13.919(027)	...	...	ZTF
58503.8	13.831(026)	13.859(013)	...	...	...	13.710(014)	13.847(008)	14.076(008)	...	TNT
58504.0	...	13.726(023)	...	...	...	...	...	...	...	MEIA3
58504.1	13.491(024)	13.723(030)	13.789(057)	13.612(029)	...	...	...	...	...	MEIA3
58504.8	13.693(026)	13.714(016)	...	...	...	13.593(005)	13.727(006)	13.998(006)	...	TNT
58504.9	...	13.627(012)	...	13.511(019)	...	...	...	...	...	MEIA3
58505.0	13.403(016)	13.646(014)	13.539(022)	...	...	...	...	...	...	MEIA3
58505.1	13.395(015)	13.648(017)	13.604(038)	13.494(020)	...	...	...	...	...	MEIA3
58506.0	13.321(019)	13.684(029)	13.493(059)	13.463(024)	...	...	...	...	...	MEIA3
58506.9	13.19(0130)	13.683(035)	...	...	...	13.505(064)	14.049(085)	13.944(060)	...	LJT
58507.1	13.407(004)	13.534(003)	13.328(023)	13.491(004)	...	...	...	...	...	ALFOSC_FASU
58507.5	...	...	...	...	...	13.412(016)	...	...	...	ZTF
58508.9	13.537(022)	13.473(012)	...	...	...	13.387(009)	13.570(008)	13.985(011)	...	TNT
58510.3	...	...	...	...	...	...	13.468(026)	...	...	ZTF
58511.4	...	...	...	...	...	...	13.457(031)	...	...	ZTF
58511.7	13.518(015)	13.448(006)	...	...	...	13.383(009)	13.540(010)	14.120(007)	...	TNT
58512.8	13.520(015)	13.470(006)	...	...	...	13.416(005)	13.556(003)	14.159(006)	...	TNT
58513.4	...	...	...	...	...	...	13.491(026)	...	...	ZTF
58514.8	13.633(015)	13.508(006)	...	...	...	13.478(008)	13.584(006)	14.227(007)	...	TNT
58515.8	13.667(013)	13.537(006)	...	...	...	13.521(004)	13.628(003)	14.276(006)	...	TNT
58517.7	13.851(016)	13.631(008)	...	...	...	13.595(009)	...	14.446(006)	...	TNT
58518.2	13.552(015)	13.576(011)	13.630(020)	...	...	...	...	...	...	MEIA3
58518.7	13.852(018)	13.531(043)	...	...	...	13.844(034)	13.769(026)	14.538(033)	...	LJT
58519.0	13.625(017)	13.572(016)	13.698(041)	13.963(036)	...	...	...	...	...	MEIA3
58520.0	13.791(012)	13.637(021)	13.751(004)	14.160(004)	...	...	...	...	...	ALFOSC_FASU
58520.1	13.722(017)	13.679(018)	13.862(044)	14.041(028)	...	...	...	...	...	MEIA3
58520.8	13.930(026)	13.717(028)	14.225(077)	14.131(046)	...	...	...	...	...	MEIA3
58521.7	14.030(026)	13.674(024)	...	...	...	13.963(033)	13.912(025)	14.758(024)	...	LJT
58521.8	13.870(017)	13.746(019)	14.070(048)	14.129(020)	...	...	...	...	...	MEIA3
58522.4	...	...	...	...	...	13.895(029)	...	...	...	ZTF
58524.0	14.282(036)	13.773(051)	14.382(040)	14.262(040)	...	...	...	...	...	MEIA3
58524.7	14.329(052)	13.874(038)	...	...	...	14.102(026)	14.120(015)	14.881(027)	...	LJT
58524.9	...	13.958(020)	...	...	...	...	...	...	...	MEIA3
58525.7	14.627(020)	14.059(007)	...	...	...	14.156(014)	14.204(035)	14.924(006)	...	TNT
58526.0	...	...	14.083(014)	14.168(022)	...	...	...	...	...	MEIA3
58526.4	...	...	...	...	...	14.257(020)	...	...	...	ZTF
58526.9	14.461(067)	14.072(021)	...	...	15.519(024)	14.321(018)	14.208(014)	14.791(012)	14.385(014)	AFOSC
58527.1	14.371(025)	14.071(027)	14.290(095)	14.141(028)	...	...	...	...	...	MEIA3
58527.9	...	...	...	...	15.765(055)	14.272(028)	14.114(024)	14.852(054)	...	Moravian
58528.0	14.621(028)	14.219(039)	...	...	...	...	...	...	...	Moravian
58528.6	14.873(042)	14.165(030)	...	...	...	14.369(025)	14.206(021)	14.812(023)	...	LJT
58528.9	14.626(025)	14.311(040)	...	...	15.982(055)	14.464(048)	14.218(027)	14.763(048)	...	Moravian
58529.9	14.852(039)	14.197(030)	14.350(079)	14.177(055)	...	...	...	...	...	MEIA3
58530.9	15.238(018)	14.355(007)	...	...	...	14.642(005)	14.350(04)	14.811(006)	...	TNT
58530.9	15.039(044)	14.344(035)	14.251(066)	14.179(053)	...	...	...	...	...	MEIA3
58533.0	...	14.354(064)	14.223(067)	14.120(033)	...	...	...	...	...	MEIA3
58533.8	15.584(047)	14.484(027)	...	...	...	14.941(022)	...	...	...	TNT

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Table A5 – *continued*

MJD	*B* (mag)	*V* (mag)	*R* (mag)	*I* (mag)	*u* (mag)	*g* (mag)	*r* (mag)	*i* (mag)	*z* (mag)	data source
58534.4	...	...	...	...	...	...	14.378(019)	...	...	ZTF
58534.7	15.618(016)	14.562(007)	...	...	...	14.996(005)	14.430(006)	14.726(007)	...	TNT
58535.7	15.718(015)	14.599(007)	...	...	...	15.053(006)	14.433(005)	14.720(007)	...	TNT
58536.8	15.527(031)	14.607(032)	...	...	...	...	...	...	...	MEIA3
58538.3	...	...	...	...	...	15.278(031)	14.500(035)	...	...	ZTF
58538.7	15.996(017)	14.762(007)	...	...	...	15.313(005)	14.558(008)	14.687(007)	...	TNT
58538.9	15.567(016)	14.767(021)	14.362(115)	14.039(023)	...	...	...	...	...	MEIA3
58540.0	15.892(027)	14.801(034)	14.329(045)	14.015(022)	...	...	...	...	...	MEIA3
58540.0	...	...	...	...	16.891(073)	15.405(025)	14.521(015)	14.604(018)	14.350(016)	AFOSC
58540.0	15.725(057)	14.753(010)	...	...	...	...	...	...	...	AFOSC
58540.7	15.929(087)	15.035(070)	...	...	...	15.483(056)	14.744(072)	14.950(095)	...	LJT
58541.3	...	...	...	...	...	15.505(031)	14.627(037)	...	...	ZTF
58541.6	16.187(016)	14.946(007)	...	...	...	15.526(006)	14.705(007)	14.783(007)	...	TNT
58542.0	15.808(035)	14.931(020)	14.559(059)	14.131(028)	...	...	...	...	...	MEIA3
58542.9	16.053(040)	14.905(027)	14.720(074)	14.235(042)	...	...	...	...	...	MEIA3
58543.0	...	...	...	...	17.444(024)	15.562(036)	14.628(029)	14.631(034)	14.380(037)	AFOSC
58543.0	15.926(034)	14.895(027)	...	...	17.197(043)	15.566(019)	14.744(024)	14.668(030)	...	Moravian
58544.0	15.913(042)	14.963(032)	...	...	17.238(043)	15.532(026)	14.756(027)	14.702(039)	...	Moravian
58544.7	16.327(016)	15.097(007)	...	...	...	15.660(010)	14.908(006)	14.979(007)	...	TNT
58545.1	15.984(023)	15.023(031)	14.719(060)	...	...	...	...	...	...	MEIA3
58546.0	15.976(102)	15.073(030)	...	...	17.276(043)	15.696(017)	14.880(021)	14.834(018)	14.551(031)	AFOSC
58547.1	16.152(037)	15.080(026)	...	...	17.307(117)	15.795(066)	14.973(026)	15.012(031)	...	Moravian
58548.0	16.171(025)	15.123(033)	...	...	17.455(060)	15.795(031)	15.018(033)	15.055(035)	...	Moravian
58548.8	16.538(020)	15.328(008)	...	...	...	15.893(008)	15.165(009)	15.244(007)	...	TNT
58550.9	16.321(032)	15.361(027)	15.254(083)	14.707(040)	...	...	...	...	...	MEIA3
58552.7	16.478(064)	15.511(038)	...	...	...	15.975(034)	15.407(130)	15.464(133)	...	LJT
58553.6	16.671(018)	15.503(007)	...	...	...	16.045(004)	15.392(003)	15.504(009)	...	TNT
58553.9	16.383(032)	15.413(022)	15.308(079)	14.833(038)	...	...	...	...	...	MEIA3
58556.9	16.449(026)	15.487(010)	15.240(050)	14.897(018)	...	...	...	...	...	MEIA3
58557.1	16.400(005)	15.500(004)	...	...	...	...	15.395(015)	15.443(041)	...	ALFOSC_FASU
58558.2	...	...	...	...	...	...	15.522(025)	...	...	ZTF
58559.2	...	...	...	...	...	...	15.539(034)	...	...	ZTF
58559.7	16.622(089)	15.661(045)	...	...	...	16.172(048)	15.575(032)	15.781(064)	...	TNT
58559.9	...	15.594(022)	15.335(051)	...	...	...	...	...	...	MEIA3
58563.0	16.527(051)	15.549(044)	15.517(131)	15.259(053)	...	...	...	...	...	MEIA3
58563.5	16.873(034)	15.803(008)	...	...	...	16.279(015)	15.757(011)	15.987(012)	...	TNT
58566.0	16.485(041)	15.812(042)	15.665(052)	15.446(043)	...	...	...	...	...	MEIA3
58566.6	16.915(016)	15.895(007)	...	...	...	16.292(006)	15.855(007)	16.060(011)	...	TNT
58569.0	16.670(014)	15.776(011)	15.676(026)	15.399(023)	...	...	...	...	...	MEIA3
58572.0	16.615(021)	15.815(021)	15.797(053)	15.549(032)	...	...	...	...	...	MEIA3
58572.2	...	...	...	...	...	...	15.989(027)	...	...	ZTF
58572.5	16.972(021)	16.063(009)	...	...	...	16.444(010)	16.067(012)	16.294(014)	...	TNT
58574.2	...	...	...	...	...	...	16.099(057)	...	...	ZTF
58575.2	...	...	...	...	...	...	16.095(062)	...	...	ZTF
58577.6	17.075(019)	16.179(007)	...	...	...	16.520(007)	16.241(009)	16.490(016)	...	TNT
58578.0	...	16.250(042)	15.985(085)	15.857(080)	...	...	...	...	...	MEIA3
58580.6	...	16.250(018)	...	...	...	16.582(009)	16.338(009)	16.620(009)	...	TNT
58581.1	16.812(029)	16.135(043)	16.095(109)	...	...	...	...	...	...	MEIA3
58581.2	...	...	...	...	...	...	16.271(037)	...	...	ZTF
58582.2	...	...	...	...	...	...	16.320(030)	...	...	ZTF
58584.0	16.836(021)	16.175(013)	16.145(073)	15.977(019)	...	...	...	...	...	MEIA3
58584.2	...	...	...	...	...	...	16.363(032)	...	...	ZTF
58587.3	...	...	...	...	...	...	16.455(024)	...	...	ZTF
58589.6	17.24(126)	16.450(068)	...	...	...	16.666(073)	16.687(072)	16.949(022)	...	TNT
58590.0	...	16.335(018)	...	...	...	...	...	...	...	MEIA3
58590.7	17.411(115)	16.414(056)	...	...	...	16.795(092)	...	16.921(026)	...	TNT
58591.4	...	...	...	...	...	...	16.603(028)	...	...	ZTF
58592.7	17.231(122)	16.601(042)	...	...	...	16.811(068)	16.819(099)	16.883(060)	...	LJT
58593.3	...	...	...	...	...	...	16.650(041)	...	...	ZTF
58594.2	...	...	...	...	...	...	16.673(042)	...	...	ZTF
58594.7	17.242(082)	16.614(048)	...	...	...	16.726(039)	16.701(045)	16.971(075)	...	TNT
58597.2	...	...	...	...	...	...	16.764(044)	...	...	ZTF
58597.4	...	...	...	...	...	...	16.772(040)	...	...	ZTF
58601.3	...	...	...	...	...	...	16.891(037)	...	...	ZTF
58603.6	17.433(016)	16.793(008)	...	...	...	16.929(009)	17.017(007)	17.391(026)	...	TNT
58605.5	17.392(024)	16.835(010)	...	...	...	16.958(014)	17.118(008)	17.423(030)	...	TNT
58606.7	17.368(041)	16.821(027)	...	...	...	...	17.026(032)	17.282(041)	...	LJT
58608.2	...	...	...	...	...	...	17.094(044)	...	...	ZTF
58608.6	17.426(022)	16.946(008)	...	...	...	17.031(009)	17.202(009)	17.519(033)	...	TNT
58612.2	...	...	...	...	...	...	17.201(045)	...	...	ZTF
58617.2	...	...	...	...	...	...	17.327(055)	...	...	ZTF
58618.3	...	...	...	...	...	...	17.373(042)	...	...	ZTF

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Table A5 – *continued*

MJD	*B* (mag)	*V* (mag)	*R* (mag)	*I* (mag)	*u* (mag)	*g* (mag)	*r* (mag)	*i* (mag)	*z* (mag)	data source
58619.7	17.598(211)	17.098(052)	...	...	...	17.281(105)	17.334(047)	17.888(046)	...	LJT
58623.6	17.671(031)	17.360(010)	...	...	...	17.333(017)	17.695(009)	18.107(052)	...	TNT
58630.6	17.741(028)	17.475(011)	...	...	...	17.400(017)	17.861(018)	18.310(066)	...	TNT
58633.6	17.789(034)	17.535(012)	...	...	...	17.492(022)	17.931(034)	18.332(057)	...	TNT
58634.2	...	...	...	...	...	...	17.825(067)	...	...	ZTF
58637.2	...	...	...	...	...	...	17.908(050)	...	...	ZTF
58637.9	17.704(007)	17.469(016)	...	...	...	...	17.831(017)	18.188(024)	...	ALFOSC_FASU
58640.2	...	...	...	...	...	...	17.971(065)	...	...	ZTF
58645.5	17.791(131)	17.568(016)	...	...	...	17.544(018)	18.310(015)	18.545(086)	...	TNT
58647.2	...	...	...	...	...	...	18.145(091)	...	...	ZTF
58648.6	18.017(072)	17.899(018)	...	...	...	17.610(016)	18.314(025)	18.746(103)	...	TNT
58650.2	...	...	...	...	...	...	18.194(089)	...	...	ZTF
58653.9	17.957(031)	17.767(020)	...	...	...	...	18.270(014)	18.514(017)	...	ALFOSC_FASU
58816.8	...	...	...	...	...	19.115(073)	19.045(104)	19.008(080)	...	LJT
58929.7	...	...	...	...	...	20.307(123)	19.816(085)	19.525(086)	...	LJT

Note. Uncertainties, in units of 0.001 mag, are 1σ.

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